XLIV Int. Symposium on Multiparticle Dynamics,
Bologna, 8 – 12 Sept., 2014

Overcoming the
POST-HIGGS DEPRESSION
through the
ASTROPARTICLE ALLIANCE

or WHY it ain't to be so in the next few years !!!

Antonio Masiero
INFN and Univ. of Padova

2012: the conquest of a new energy scale in physics

- ~1900 ATOMIC SCALE 10^{-8} cm. $1/(\alpha m_e)$
- ~1970 STRONG SCALE 10^{-13} cm. $Me^{-2\Pi/\alpha_s b}$
- ~2010 WEAK SCALE 10^{-17} cm. TeV^{-1}

FUNDAMENTAL OR DERIVED SCALE?

EX. EXTRA-DIMENSIONS
or
TeV STRING THEORY

EX.: TECHNICOLOR or
SUSY with ELW RAD. BREAKING

NEW PARTICLES AT THE TEV SCALE?

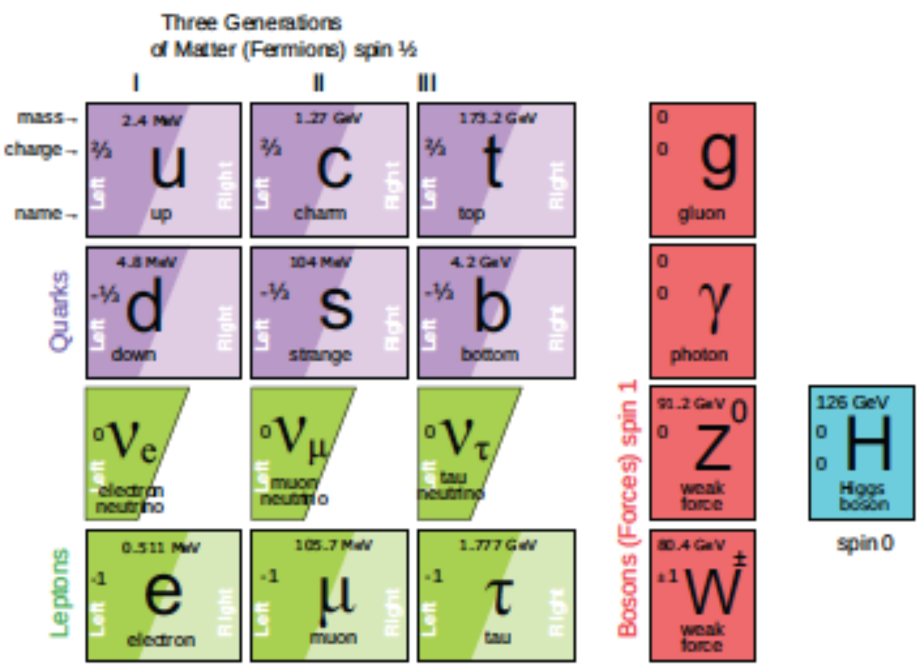
SIMPLE (STANDARD)

IS

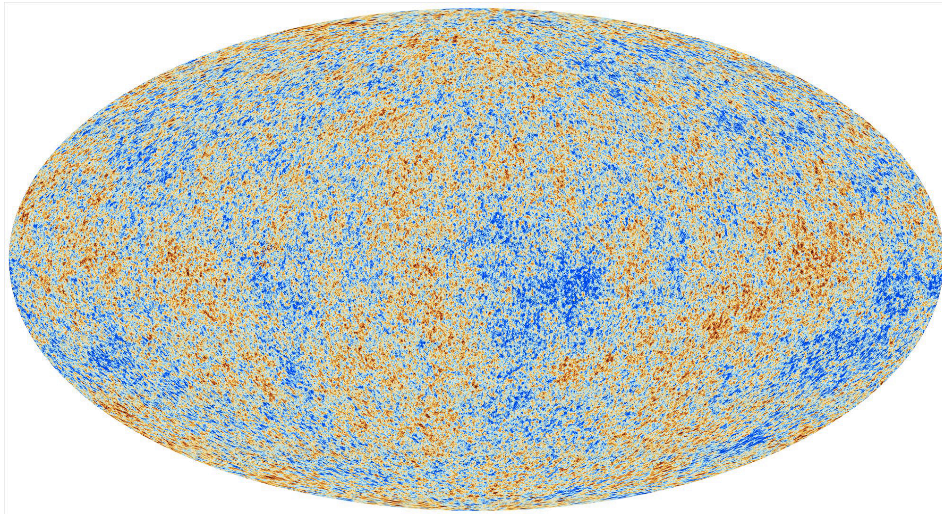
BEAUTIFUL

2013: the thiumph of the **STANDARD**

- **PARTICLE STANDARD MODEL**



- **COSMOLOGY STANDARD MODEL**



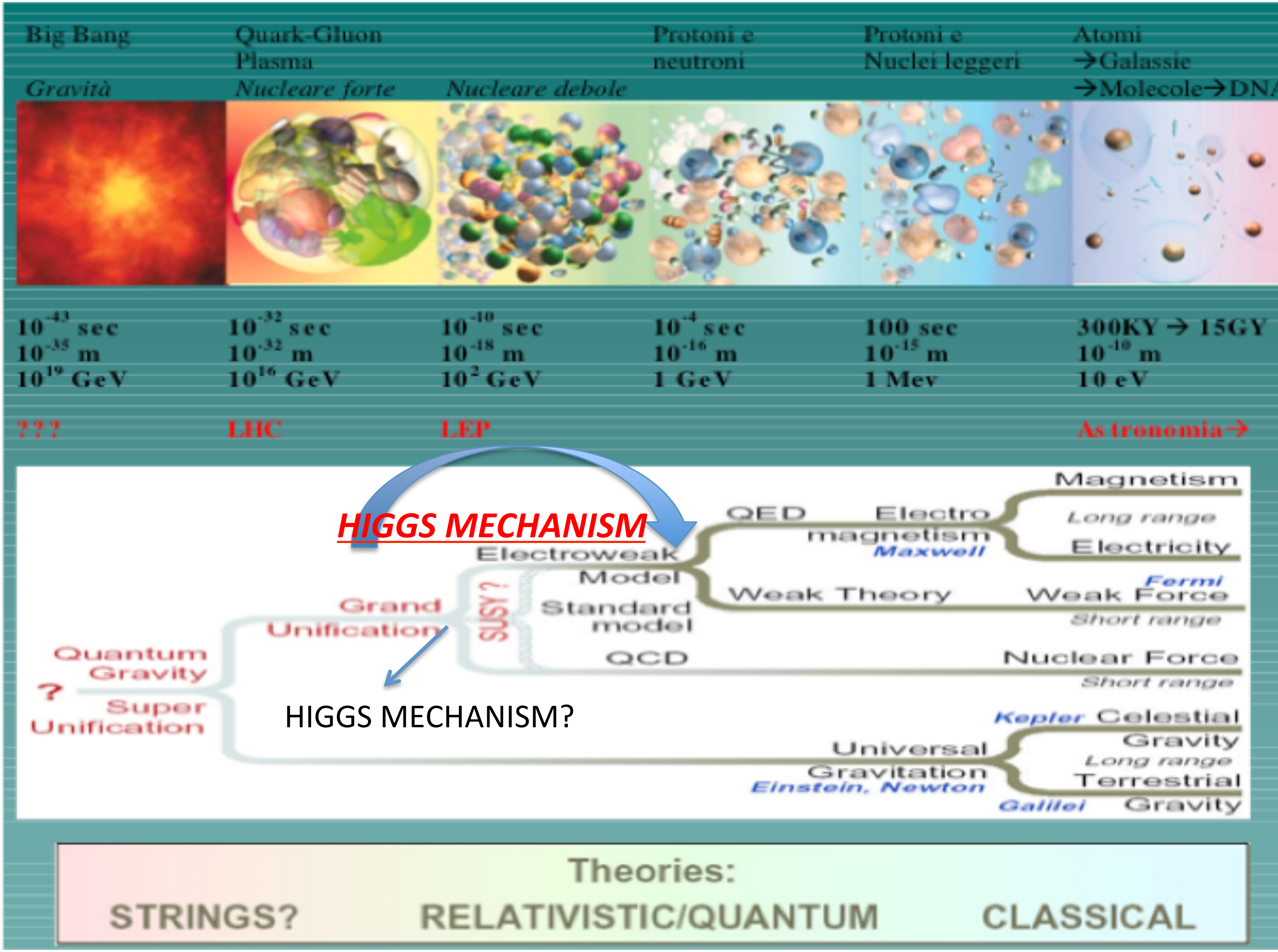
ΛCDM + "SIMPLE" INFLATION

$\Omega_\Lambda = 0.686 \pm 0.020$

$\Omega_m = 0.314 \pm 0.020$

$\Omega_b h^2 = 0.02207 \pm 0.00033$

$h = 0.674 \pm 0.014$

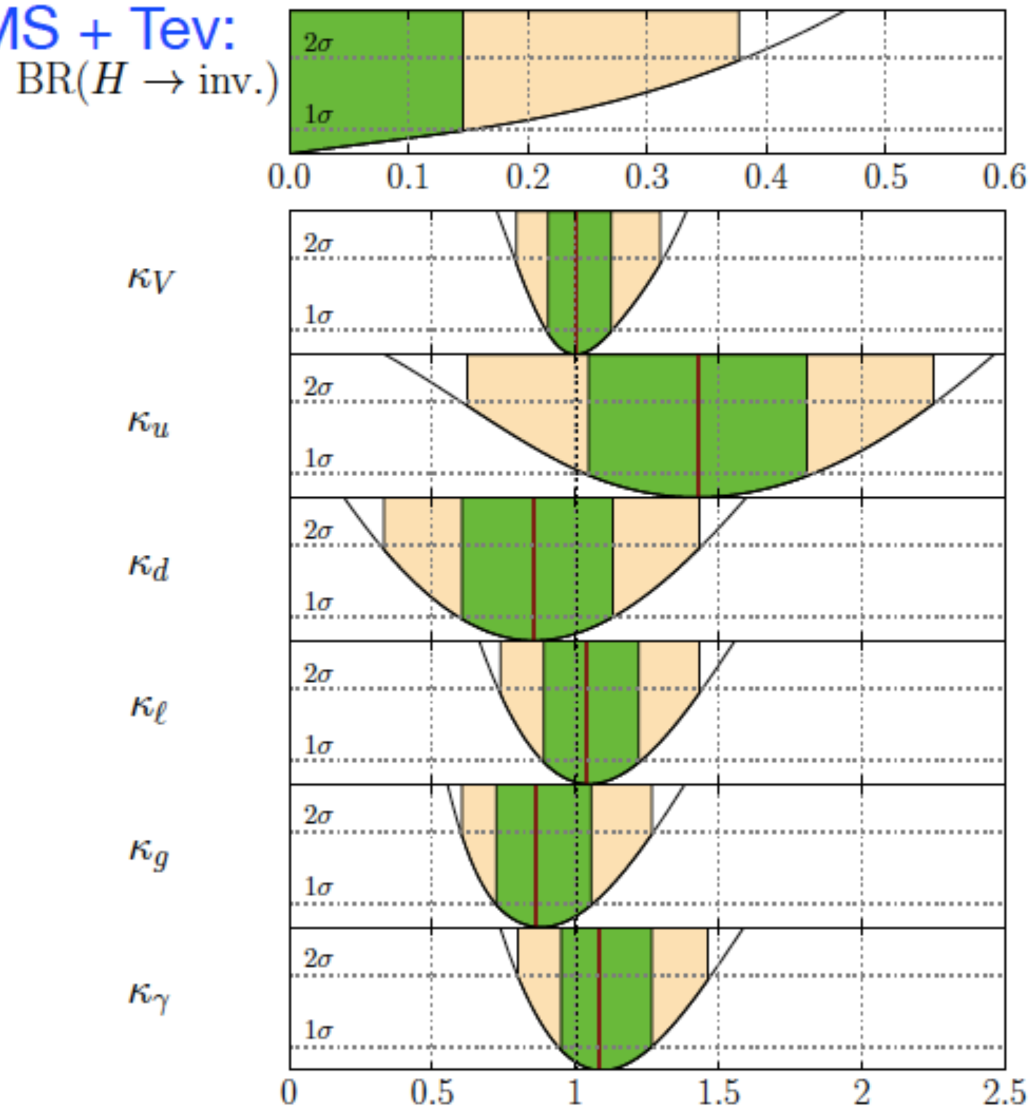


Big Bang
Gravità
Quark-Gluon Plasma
Nucleare forte
Nucleare debole
Protoni e neutroni
Protoni e Nuclei leggeri
Atomi →Galassie →Molecole→DNA
10^-43 sec 10^-35 m 10^19 GeV
10^-32 sec 10^-32 m 10^16 GeV
10^-10 sec 10^-18 m 10^2 GeV
10^-4 sec 10^-16 m 1 GeV
100 sec 10^-15 m 1 Mev
300KY → 15GY 10^-10 m 10 eV
???
LHC
LEP
Astronomia→
HIGGS MECHANISM
Magnetism
QED
Electro magnetism
Maxwell
Long range
Electricity
Electroweak Model
Fermi
Weak Theory
Weak Force
Short range
Grand Unification
SUSY ?
Standard model
Quantum Gravity
QCD
Nuclear Force
Short range
?
Super Unification
HIGGS MECHANISM?
Kepler
Celestial Gravity
Universal Gravitation
Long range
Einstein, Newton
Terrestrial Gravity
Galilei
Theories:
STRINGS?
RELATIVISTIC/QUANTUM
CLASSICAL

Constraints on coupling scale factors from ATLAS + CMS + Tevatron data

ATLAS + CMS + Tev:

Seven fit parameters

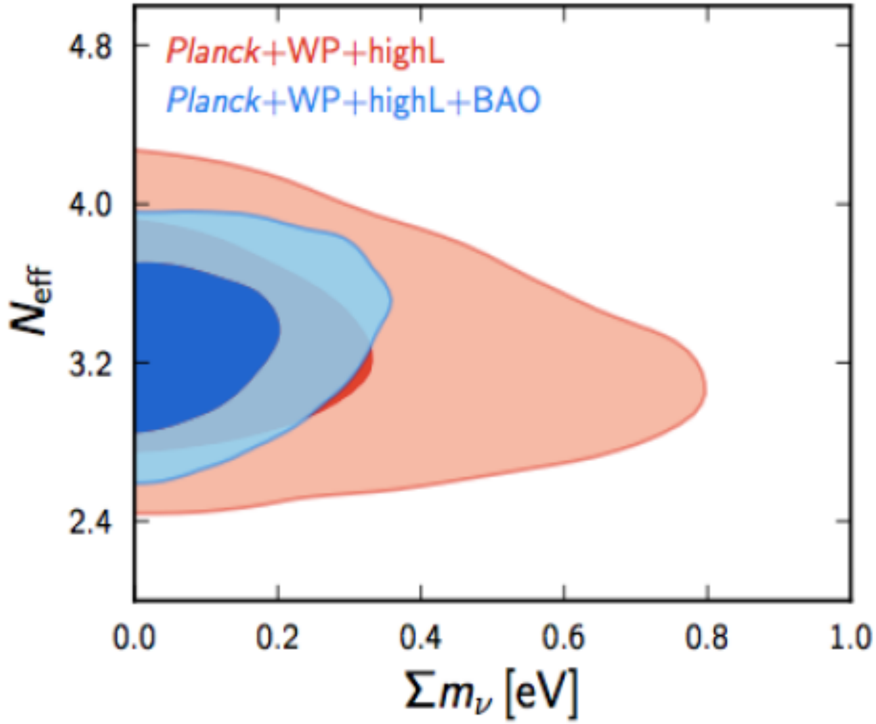


HiggsSignals

[P. Bechtle, S. Heinemeyer, O. Stål, T. Stefaniak, G. W. '14]

G. WEIGLEIN, IPA2014

⇒ Significantly improved precision compared to ATLAS or CMS results alone



N_{eff}=3.36±0.34

The extracted value of N_{eff} depends whether one makes use of the value of the Hubble parameter from the Planck data or from independent observations

Σm_ν< 0.23- 0.8 eV

Recent (and controversial!) **BICEP2** results:

from the measurement of the B-mode polarization of the CMB photons

→ initial **inflationary epoch** at energies ~ **$V^{1/4}$ = 1.94 x 10^{16} GeV $(r/0.12)^{1/4}$;** r= ratio of the CMB tensorial/scalar components – from BICEP2 r ~0.2 , r≠0 at ~6 σ

INFLATON at ~ 10^{16} GeV, not standard Higgs inflation (see, however, Bezrukov and Shaposhnikov)

MICRO

GWS STANDARD MODEL

MACRO

HOT BIG BANG
STANDARD MODEL

UNIVERSE EXPANSION +
WEAK INTERACTIONS
1 sec. after BB

NUCLEOYINTHESIS

NUMBER OF BARYONS and **OF NEUTRINO SPECIES** →
CONFIRMED FROM CMB **350000 YEARS** AFTER BB

BUT ALSO

FRICTION POINTS

-COSMIC MATTER-ANTIMATTER ASYMMETRY

-INFLATION ???

- DARK MATTER + DARK ENERGY

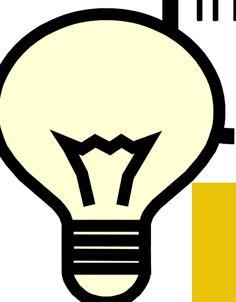
OBSERVATIONAL EVIDENCE OF NEW PHYSICS

BEYOND THE STANDARD

The Energy Scale from the "Observational" New Physics

- neutrino masses
- dark matter
- baryogenesis
- inflation

→ NO NEED FOR THE NP SCALE TO BE CLOSE TO THE ELW. SCALE

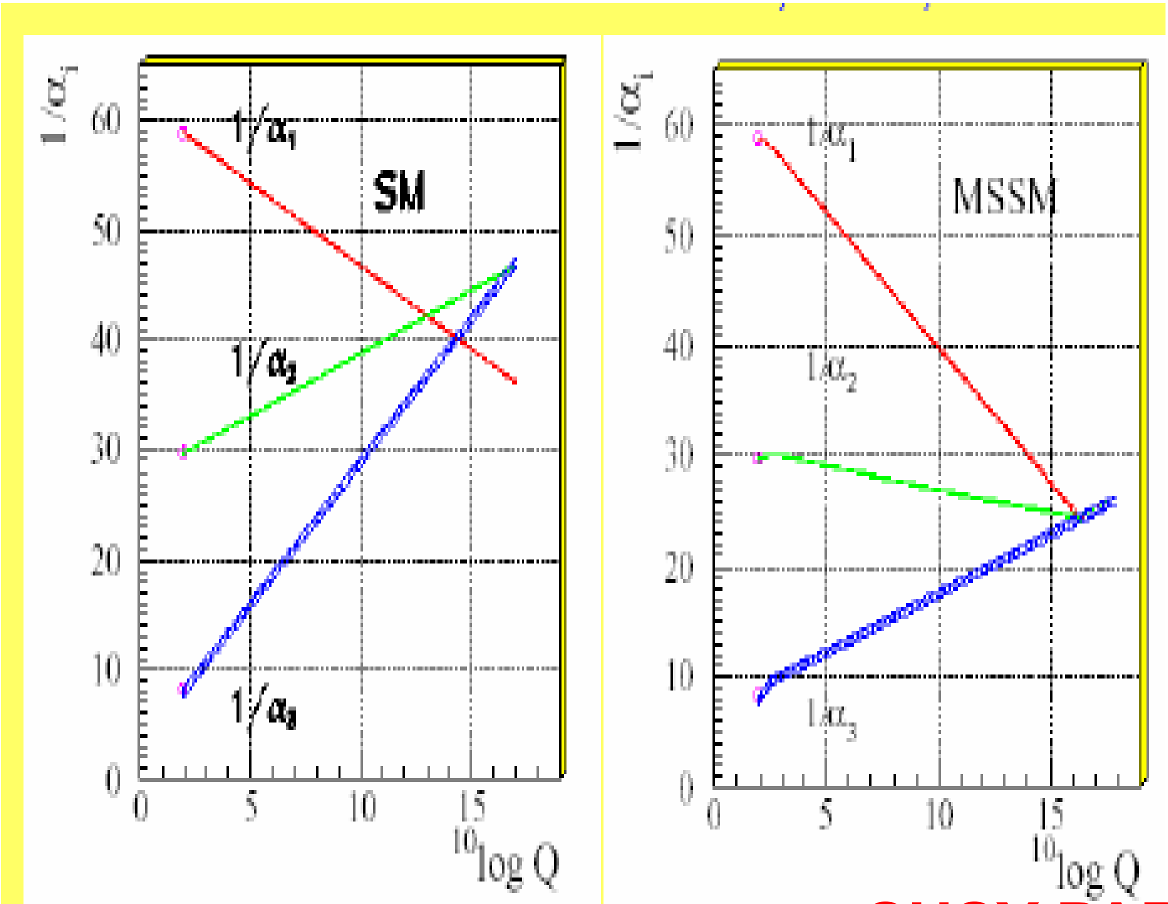


The Energy Scale from the "Theoretical" New Physics

★ ★ ★ **Stabilization of the electroweak symmetry breaking at M_W calls for an ULTRAVIOLET COMPLETION of the SM already at the TeV scale** **+**

★ **CORRECT GRAND UNIFICATION "CALLS" FOR NEW PARTICLES AT THE ELW. SCALE**

LOW-ENERGY SUSY AND UNIFICATION



Input

$$\alpha^{-1}(M_Z) = 128.978 \pm 0.027$$

$$\sin^2 \theta_{\overline{MS}} = 0.23146 \pm 0.00017$$

$$\alpha_s(M_Z) = 0.1184 \pm 0.0031$$

Output

$$M_{SUSY} = 10^{3.4 \pm 0.9 \pm 0.4}\ \text{GeV}$$

$$M_{GUT} = 10^{15.8 \pm 0.3 \pm 0.1}\ \text{GeV}$$

$$\alpha_{GUT}^{-1} = 26.3 \pm 1.9 \pm 1.0$$

e.g. **SUSY PARTICLES AT THE TEV SCALE !**

THE "COMPREHENSION" OF THE ELECTROWEAK SCALE

$$V = \mu^2|H|^2 + \lambda|H|^4$$ **μ ~ 10^2 GeV**

- M = O(10^{16} GeV)

	SU(3)	SU(2)	U(1)
L	1	2	-1/2
e	1	1	1
Q	3	2	1/6
u	3*	1	-2/3
d	3*	1	1/3

→ SO(10): 16

$$m_H^2 \sim -2\mu^2 + \frac{g^2}{(4\pi)^2} M^2$$

ONLY FOR SCALARS; SM FERMIONS AND GAUGE BOSON MASSES ARE PROTECTED BY THE SU(2) × U(1) SYMMETRY !

To comprehend (i.e. stabilize) the elw. scale need NEW PHYSICS (NP) to be operative at a scale $m_{NP} \ll M$

3 comments on m_{NP}

ROMANINO at WHAT NEXT 2014

- Any upper bound on m_{NP} is subjective: any value of m_{NP} acceptable provided one accepts a cancellation

$$\Delta \gtrsim \left(\frac{m_{\rm NP}}{0.5\,{\rm TeV}}\right)^2$$

$m_{NP} > 1.5$ TeV	$\leftrightarrow$	$\Delta > 10$
$m_{NP} > 5$ TeV	$\leftrightarrow$	$\Delta > 100$

m_{NP} x 2 → Δ x 4

- The bound on Δ is model-dependent:

"supersoft" $$\Delta \sim \left(\frac{m_{\rm NP}}{0.5\,{\rm TeV}}\right)^2$$

"soft" $$\Delta \sim \left(\frac{m_{\rm NP}}{0.5\,{\rm TeV}}\right)^2 \times \log\left(\frac{M^2}{m_{\rm NP}^2}\right)$$

- The argument assumes that the electroweak scale can be understood in terms of physics at a scale ~

$$M \gg m_h$$

- **Alternative 1** : it could be that there is nothing at scal $M \gg m_h$ **FINITE NATURALNESS**

- **Alternative 2**: it could be that there is indeed new physics at M, but **"REDUCTIONISM" DOES NOT HOLD** (**anthropic selection**) – i.e. **physics at 10^2 Gev depends on specific choices of parameters made at 10^{16} GeV !** (**unprecedented in physics**)

$$m_H^2 \sim -2\mu^2 + \frac{g^2}{(4\pi)^2} M^2$$

- **UNNATURAL or FINE-TUNING SOLUTION** tuning of parameters at the scale M with precision O($m_H/M)^2$

- **NATURAL SOLUTION**
  Dynamics or symmetries or space-time modifications giving rise to a UV cut-off ~ $(m_H)^2$

- **SYMMETRY vs. MULTIVERSE**

The BIG and the SMALL- dim[m] ≠ 0

- $V = \mu^2|H|^2 + \lambda|H|^4$ what is the value of the energy of its vacuum, i.e. the SM **vacuum energy**?
  → $V_o = \mu^2 \langle H\rangle^2 + \lambda \langle H\rangle^4 \sim (100\ \mathrm{GeV})^2$

observed vacuum energy, i.e. dark energy accelerating the expansion of the Universe O(10^{-3} eV)

- V defined up to a constant → choose such constant to **cancel** the O $(100\ \mathrm{GeV})^2$ contribution
- **10^{-3} eV 10^{2} GeV 10^{16} GeV 10^{19} GeV**

- **Why** so different mass scales ?

- **How** to guarantee their separation → symmetry vs. multiverse

The BIG and the SMALL – dim[m]=0

- $h_t - h_e$ **flavour** issue
- L_{SM} no symmetry prevents to add a term violating **CP in the strong interactions** whose size depends on a **dimensionless** parameter θ → the bound on the neutron EDM → $\theta < 10^{-10}$
- **The θ – problem** : the symmetry solution

Axion from breaking of global chiral symmetry; axion field acts as dynamical theta para-meter, [Peccei,Quinn 77; Weinberg 78; Wilczek 78]

$$\mathcal{L} \supset -\frac{\alpha_s}{8\pi} \underbrace{\frac{A}{f_A}}_{\bar{\theta}} G^a_{\mu\nu}\tilde{G}^{a,\mu\nu}$$

spontaneously relaxing to zero, $\langle A \rangle = 0$ (thus CP conserved)

- mass due to chiral symmetry breaking $m_A \sim m_\pi f_\pi / f_A$
- has universal coupling to photons, $\mathcal{L} \supset -\frac{\alpha}{8\pi} C_0 \frac{A}{f_A} F_{\mu\nu}\tilde{F}^{\mu\nu}$

Ringwald

but θ very small TODAY at $T=0$ $\not\Rightarrow$

θ very small at $T \neq 0$ in the early Universe

Strong CP violation from the QCD axion can be responsible for the matter antimatter asymmetry of the universe in the context of cold baryogenesis

if the EW phase transition is delayed down to the QCD scale

G. SERVANT IPA2014

LOW-ENERGY SIGNATURES OF UNIFICATION AT 10^{16} GeV

- **PROTON DECAY** mediated by new particles (scalars or gauge bosons) related to the unified physics at 10^{16} GeV which DOES NOT respect the BARYON and LEPTON NUMBER SYMMETRIES → for a mediator of mass ~ 10^{16} GeV we expect a proton lifetime in the ballpark of ~ 10^{34} years → exp. accessible

- **NEUTRON-ANTINEUTRON OSCILLATION** if the unified symmetry (ex. SO(10)) breaks down to an intermediate symmetry subsequently spontaneously broken at ~ 10^6 GeV with the breaking of Baryon number of two units (ex. $SO(10) \rightarrow SU(4)_{PS} \times SU(2)_L \times SU(2)_R \rightarrow SU(3) \times SU(2)_L \times U(1)_Y$) → exp. accessible (for instance , at the ESS)

Ways to implement the BEH mechanism

- **STANDARD HIGGS:**

- **FINE-TUNED (unnatural Higgs – anthropic road, high-scale fundamental theory taking care of it, …)**
- **NATURAL (protection mechanism: low-energy SUSY; inexistence of the scale hierarchy problem: extra dimensions, warped space, …)**

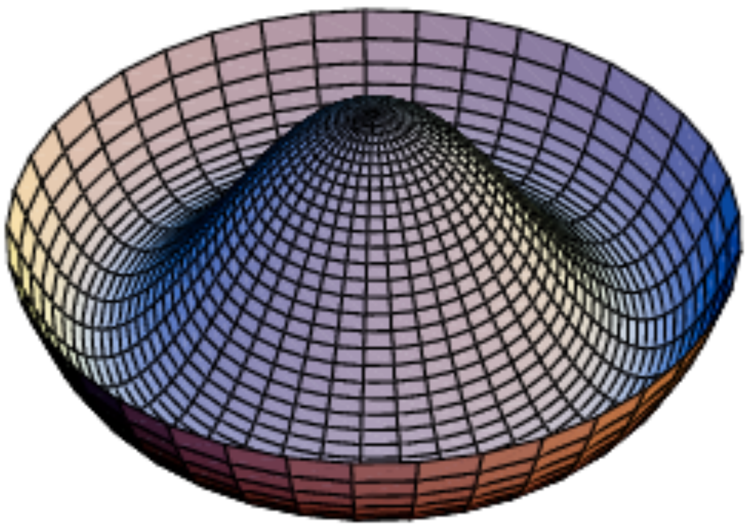
- **COMPOSITE HIGGS: Pseudo-Goldstone boson**

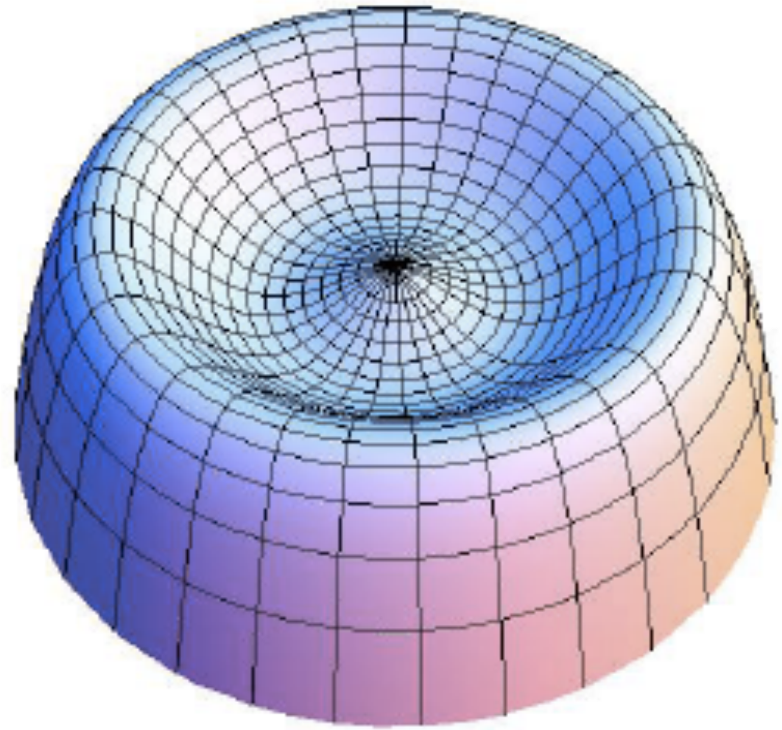
- **More EXOTIC HIGGS (SEMI-IMPOSTOR HIGGS – completely impostor higgs like in the truly higgs-less models by now indefensible) ex.: HIGGS-RADION FIELD, just a single higgs doublet accomplishing a two-fold task, namely breaking the EW symmetry AND stablizing the 5th extra dimension (Geller, Bar-Shalom, Soni)**

On the peculiar value of M_H

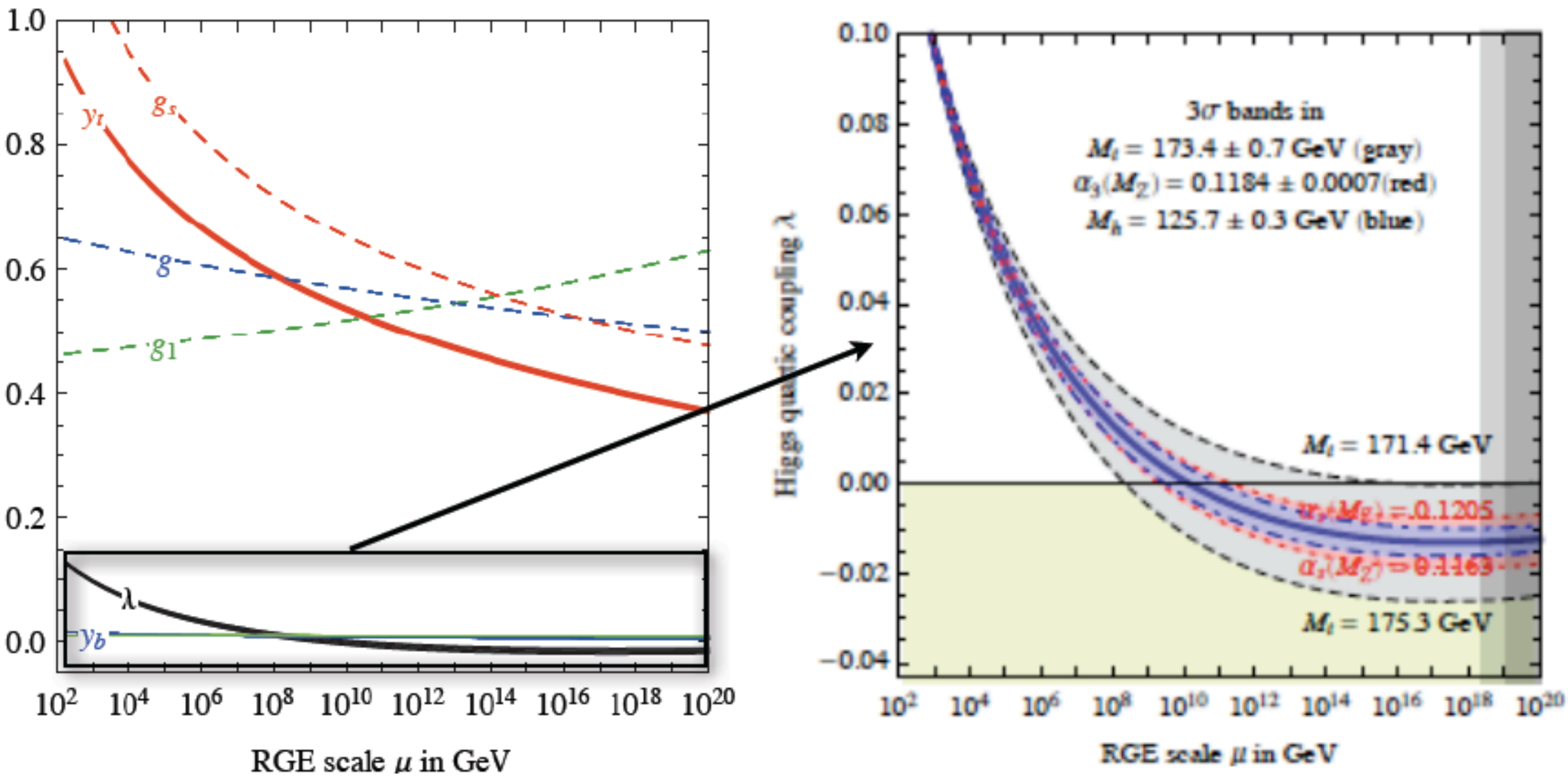
- **For the SM to survive up to a very large scale, M_{GUT} or M_{Planck} : M_H in the fork 125 – 180 GeV, with ~ 125 GeV just on the verge between stability and instability of the vacuum state where the SM sits**
- For the existence of a **(minimal) supersymmetric extension of the SM** at the elw. scale, the lightest SUSY Higgs must have **$M_h < 130$ GeV** (for $M_h > 120$ GeV, the radiative correction to M_h is ~ 50% of the tree-level value)

STABILITY ⟷ INSTABILITY





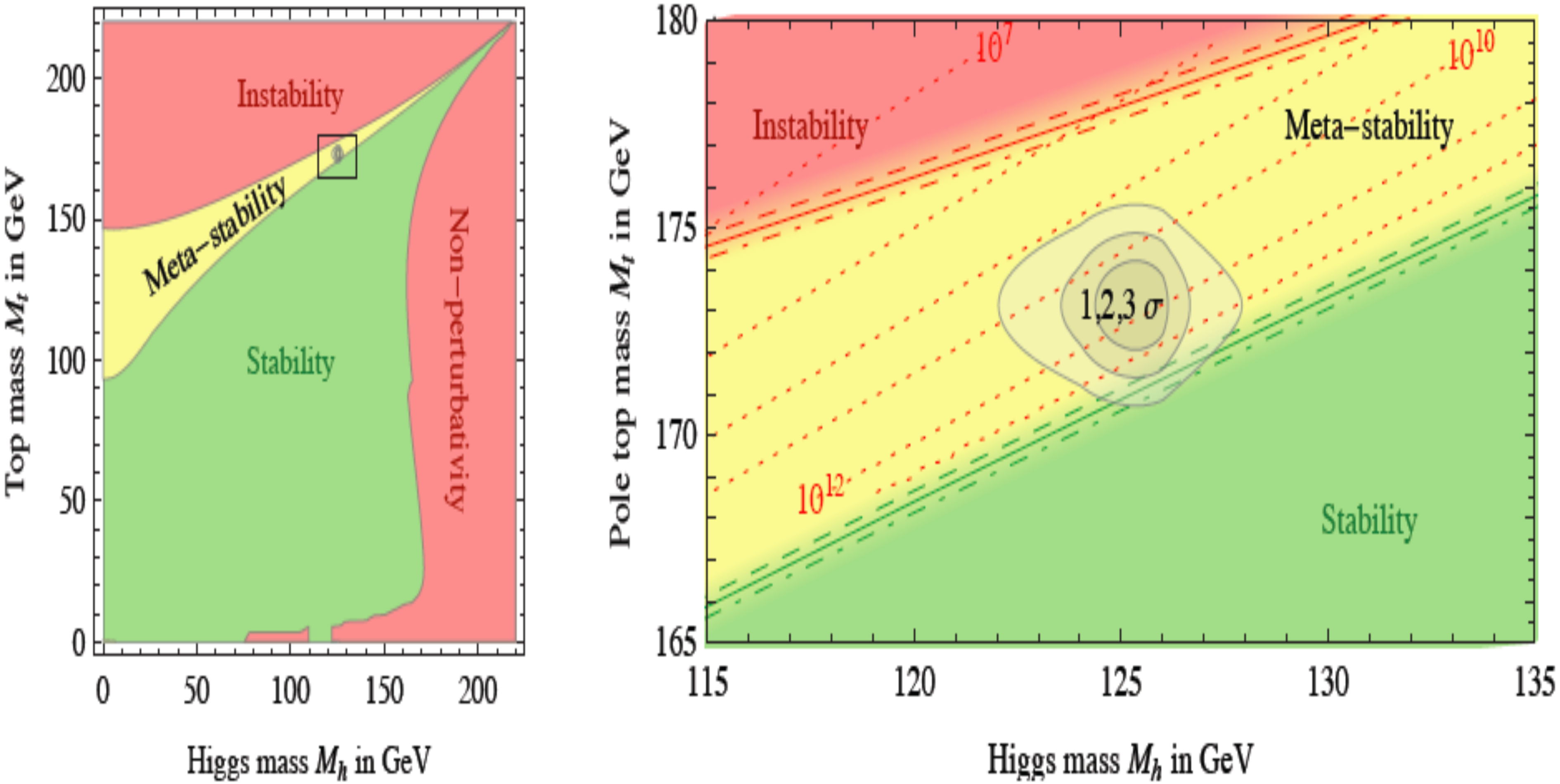
ON THE IMPORTANCE OF PRECISELY MEASURING HIGGS and TOP MASSES



Buttazzo, Degrassi, Giardino, Giudice, Sala, Salvio, Strumia 2013

For previous works: Krive, Linde ‘76; Krasnikov ‘78; Maiani, Parisi, Petronzio ‘78; Cabibbo et al ‘79; Lindner ‘86; Altarelli, Isisdori ’96; Ellis et al 2009; Shaposhnikov et al ’12; Elias-Miro’ ‘et a ‘’12; …..
Degrassi, Di Vita, Elias-Miro, Espinosa, Giudice, Isidori, Strumia ‘12

LIVING DANGEREOUSLY IN A "PROBABLE" METASTABLE UNIVERSE



BEZRUKOV, KALMIKOV, KNIEHL, SHAPOSHNIKOV 2012;
DEGRASSI, DI VITA, ELIAS-MIRO', ESPINOSA, GIUDICE, ISIDORI, STRUMIA 2012

FIRST COMPLETE ANALYSIS NNLO OF THE SM HIGGS POTENTIAL

ON THE IMPORTANCE OF PRECISELY MEASURING HIGGS and TOP MASSES

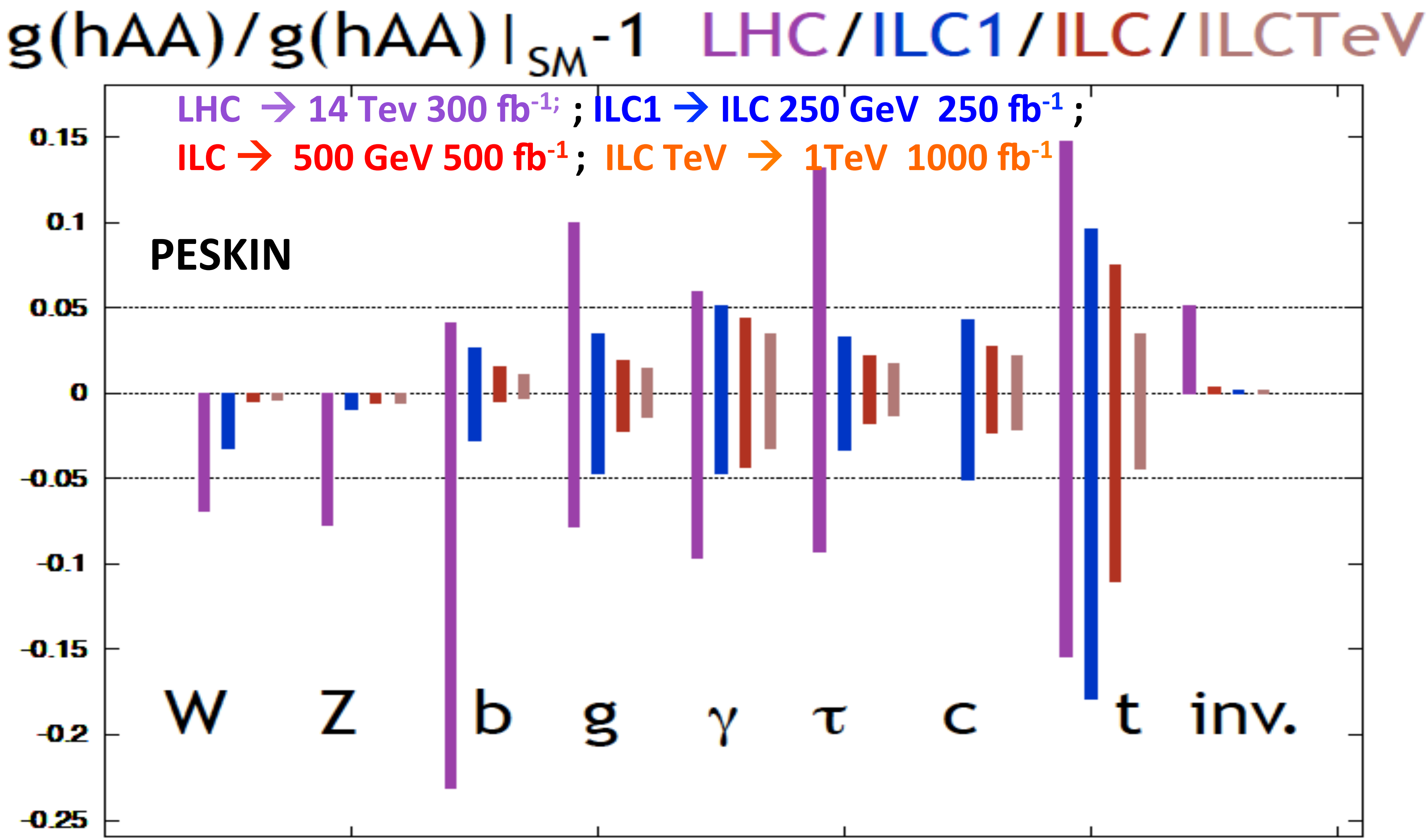
DEGRASSI ET AL

Type of error	Estimate of the error	Impact on M_h
M_t	experimental uncertainty in M_t	± 1.4 GeV
α_s	experimental uncertainty in α_s	± 0.5 GeV
Experiment	Total combined in quadrature	± 1.5 GeV
λ	scale variation in λ	± 0.7 GeV
y_t	$\mathcal{O}(\Lambda_{QCD})$ correction to M_t	± 0.6 GeV
y_t	QCD threshold at 4 loops	± 0.3 GeV
RGE	EW at 3 loops + QCD at 4 loops	± 0.2 GeV
Theory	Total combined in quadrature	± 1.0 GeV

INTRINSIC DIFFICULTY TO "DEFINE" WHAT THE TOP MASS IS AT A **HADRON COLLIDER** WITH UNCERTAINTY ≤ 1 GeV

HIGGS Couplings Sensitivity at LHC and ILC

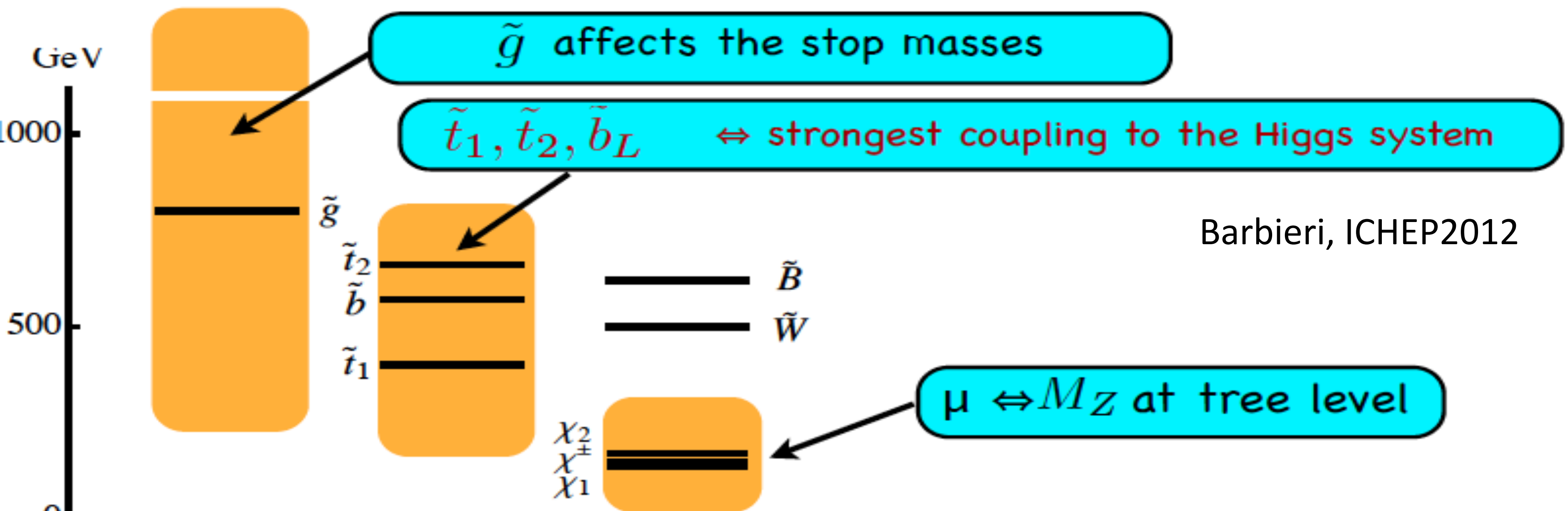
TANABE at this meeting



NATURAL SUSY

LOW-ENERGY SUSY to cope with the gauge hierarchy problem: only the SUSY particles involved in the cancellation of the quadratic div. to the Higgs mass have to remain "light"

"s-particles at their naturalness limit"



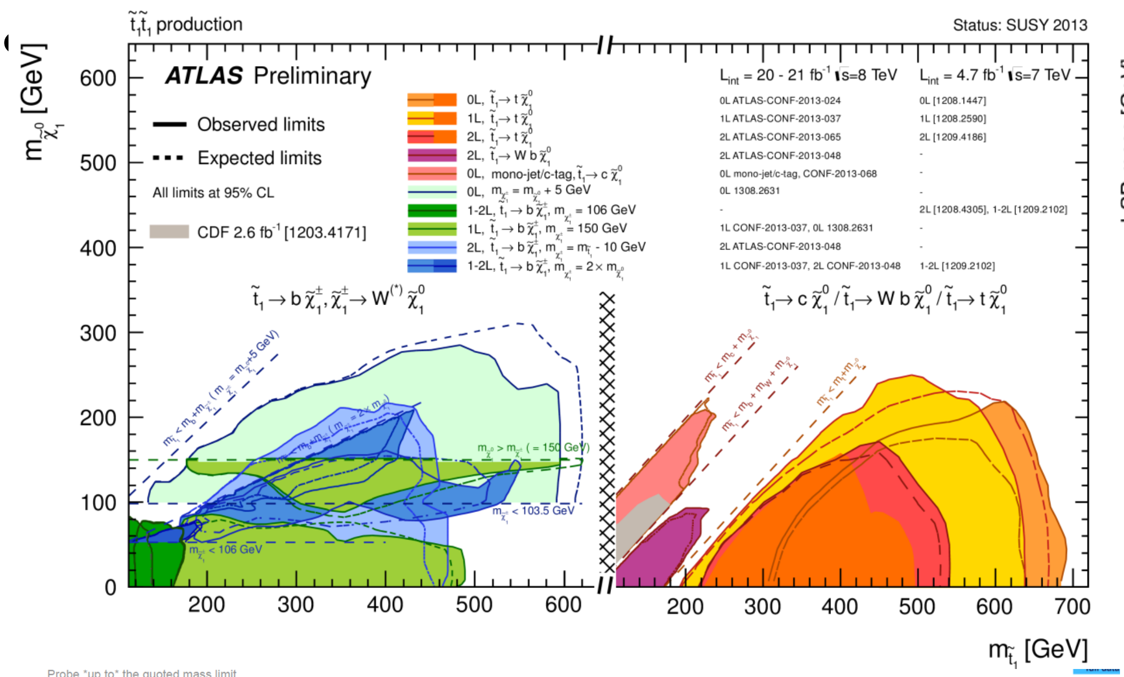
Barbieri, ICHEP2012

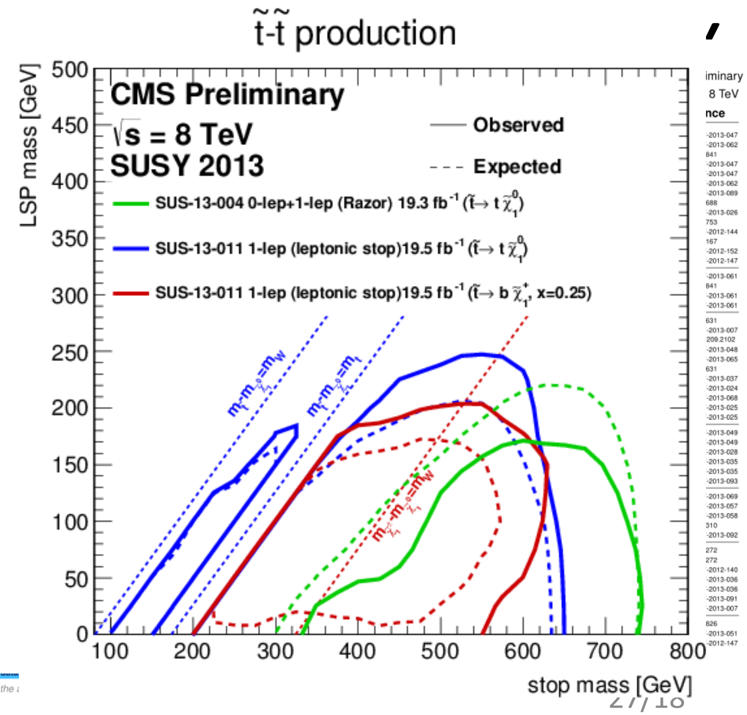
orange areas indicative and dependent on how the Higgs boson gets its mass

$\tilde{B}, \tilde{W}$ not much constrained but expected below $m_{\tilde{g}}$

SUSY searches

- Higgs mass value reduces SUSY parameter space
- Direct searches unsuccessful → gluinos well above 1 TeV
- Working on stop expected lighter





REALITY vs. VIRTUALITY RACE:

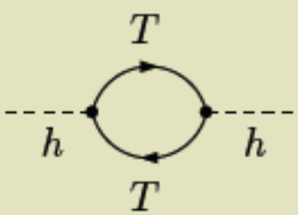
SEARCHING FOR NEW PARTICLES THROUGH THEIR VIRTUAL EFFECTS

Higgs and flavor physics as indirect BSM probes

NEUBERT SUSY2012

$$\mathcal{L}_{\mathrm{EFT}} = \Lambda_{\mathrm{UV}}^2\,\Phi^\dagger\Phi - \lambda(\Phi^\dagger\Phi)^2 + \mathcal{L}_{\mathrm{SM}}^{\mathrm{gauge}} + \mathcal{L}_{\mathrm{SM}}^{\mathrm{Yukawa}} + \frac{\mathcal{L}^{(5)}}{\Lambda_{\mathrm{UV}}} + \frac{\mathcal{L}^{(6)}}{\Lambda_{\mathrm{UV}}^2} + \dots$$

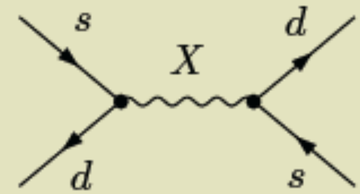
electroweak symmetry breaking $\Downarrow$ Higgs mass



$$\sim \frac{g_T^2}{16\pi^2}\Lambda_{\mathrm{UV}}^2$$

no fine-tuning $\Downarrow$

$$\Downarrow$$



$$\sim \frac{g_X^2}{\Lambda_{\mathrm{UV}}^2}$$

bounds on flavor mixing $\Downarrow$ assuming *generic* flavor structure

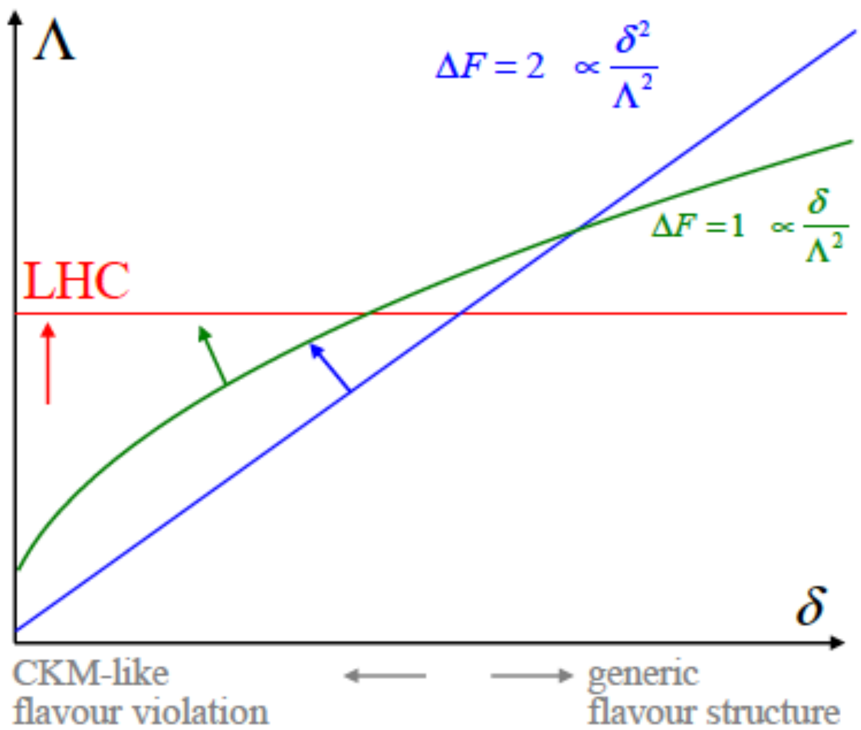
$\Lambda_{\mathrm{Higgs}} \lesssim 1$ TeV $\qquad$ $\Lambda_{\mathrm{flavor}} \gtrsim 10^3$ TeV

Possible solutions to flavor problem explaining $\Lambda_{\mathrm{Higgs}} << \Lambda_{\mathrm{flavor}}$:

(i) $\Lambda_{\mathrm{UV}} >> 1$ TeV: **Higgs fine tuned**, new particles too heavy for LHC

(ii) $\Lambda_{\mathrm{UV}} \approx 1$ TeV: quark flavor-mixing protected by a **flavor symmetry**

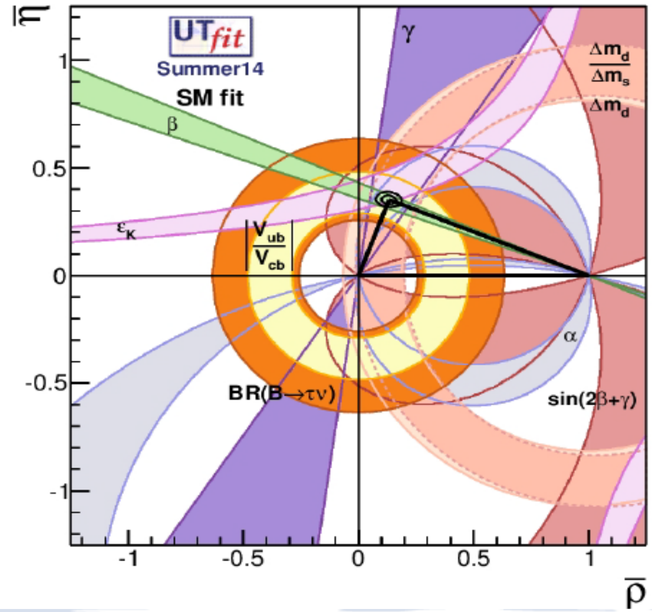
Flavour vs. collider bounds on the new physics scale Λ



- Rare decays: much more observables, more sensitive to "small" flavour violation at "natural" scales

Unitarity Triangle analysis in the SM:

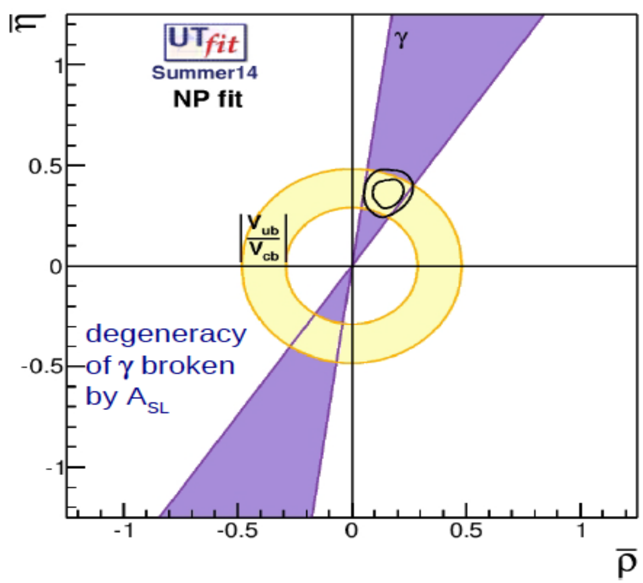
Observables	Accuracy
$\lvert V_{ub}/V_{cb}\rvert$	~ 13%
ε_K	~ 0.5%
Δm_d	~ 1%
$\lvert \Delta m_d/\Delta m_s\rvert$	~ 1%
$\sin 2\beta$	~ 3%
α	~ 8%
γ	~ 10%
$BR(B \to \tau\nu)$	~ 19%



levels @ 95% Prob

$\bar{\rho} = 0.132 \pm 0.023$
$\bar{\eta} = 0.351 \pm 0.013$

NP analysis results



$\bar{\rho} = 0.149 \pm 0.045$
$\bar{\eta} = 0.363 \pm 0.049$

SM is

$\bar{\rho} = 0.132 \pm 0.023$
$\bar{\eta} = 0.351 \pm 0.013$

M. Bona 2014

Unitarity Triangle analysis in the SM:

obtained excluding the given constraint from the fit

Observables	Measurement	Prediction	Pull (#σ)
sin2β	0.679 ± 0.024	0.753 ± 0.042	~ 1.5
γ	68.3 ± 7.5	69.5 ± 3.9	< 1
α	90.8 ± 7.0	87.2 ± 3.9	< 1
$\|V_{ub}\| \cdot 10^3$	3.75 ± 0.46	3.62 ± 0.12	< 1
$\|V_{ub}\| \cdot 10^3$ (incl)	4.40 ± 0.31	–	~ 2.3
$\|V_{ub}\| \cdot 10^3$ (excl)	3.42 ± 0.22	–	< 1
$\|V_{cb}\| \cdot 10^3$	40.9 ± 1.0	42.1 ± 0.7	< 1
B_K	0.766 ± 0.010	0.841 ± 0.078	< 1
$BR(B \to \tau\nu)[10^{-4}]$	1.14 ± 0.22	0.81 ± 0.07	~ 1.4
$BR(B_s \to ll)[10^{-9}]$	2.9 ± 0.7	3.92 ± 0.16	~ 1.3
$BR(B_d \to ll)[10^{-9}]$	0.37 ± 0.15	0.114 ± 0.007	~ 1.7
$A_{SL}^d \cdot 10^3$	-4.8 ± 5.2	0.012 ± 0.002	< 1
$A_{\mu\mu} \cdot 10^3$	-7.9 ± 2.0	-0.12 ± 0.02	~ 3.9

THE FLAVOUR PROBLEMS

FERMION MASSES

What is the rationale hiding behind the spectrum of fermion masses and mixing angles (our "**Balmer lines**" problem)

→ **LACK OF A FLAVOUR "THEORY"**

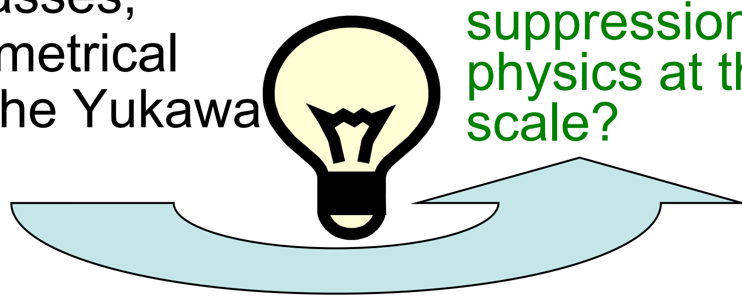
(new flavour – horizontal symmetry, radiatively induced lighter fermion masses, dynamical or geometrical determination of the Yukawa couplings, …?)

FCNC

Flavour changing neutral current (FCNC) processes are suppressed.

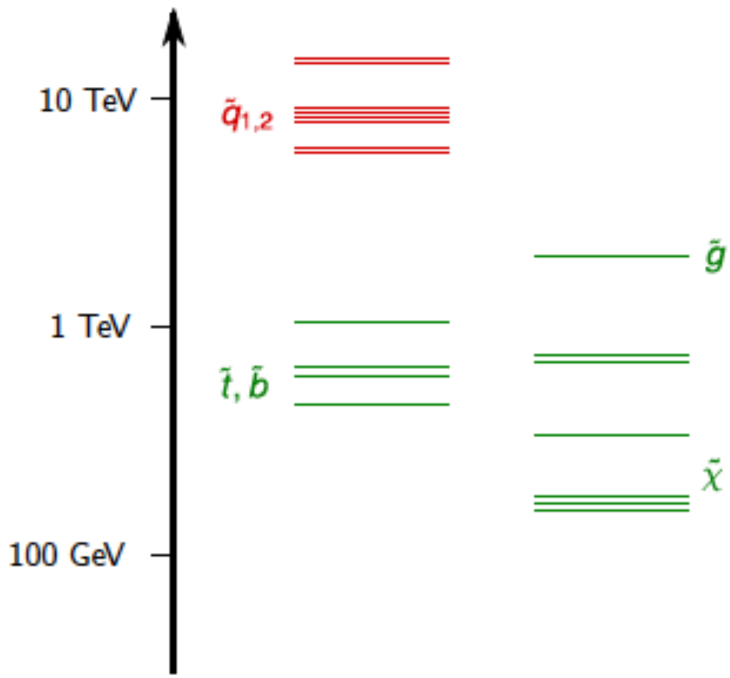
In the SM two nice mechanisms are at work: the **GIM mechanism** and the structure of the **CKM mixing matrix.**

How to cope with such delicate suppression if the there is new physics at the electroweak scale?



Natural SUSY & $U(2)^3$

Straub 2014



- Strong LHC bounds on 1st/2nd generation squarks
- Light 3rd generation squarks to solve hierarchy problem

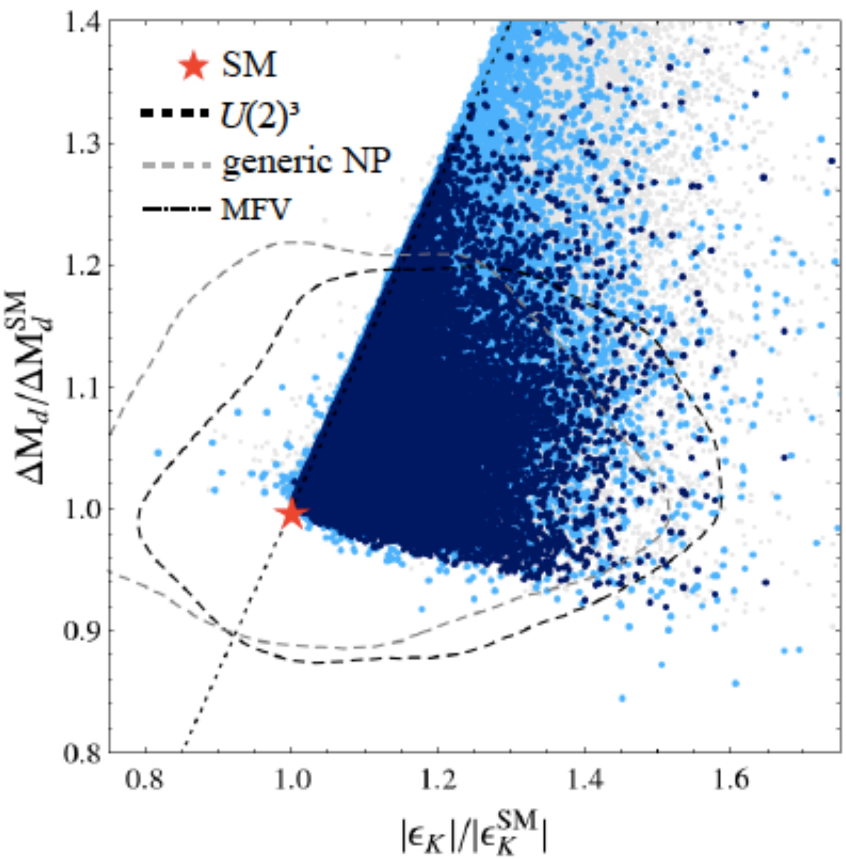
Natural SUSY with split generations + CKM-like flavour violation: approximate, minimally broken $U(2)^3$ flavour symmetry ("split MFV")

[Barbieri et al. 1105.2296]

LHC (real NP particles) – FCNC (virtual NP particles) CONSTRANING New Physics

D. STRAUB, IPA2014

Numerical results for $\Delta F = 2$ observables



- All blue points fullfil collider bounds
- Dashed lines: $\Delta F = 2$ constraints (black: $U(2)^3$, gray: generic)
- Direct bounds almost as constraining as flavour, except for compressed spectra

[Barbieri et al. 1402.6677]

B_s mixing phase measurement

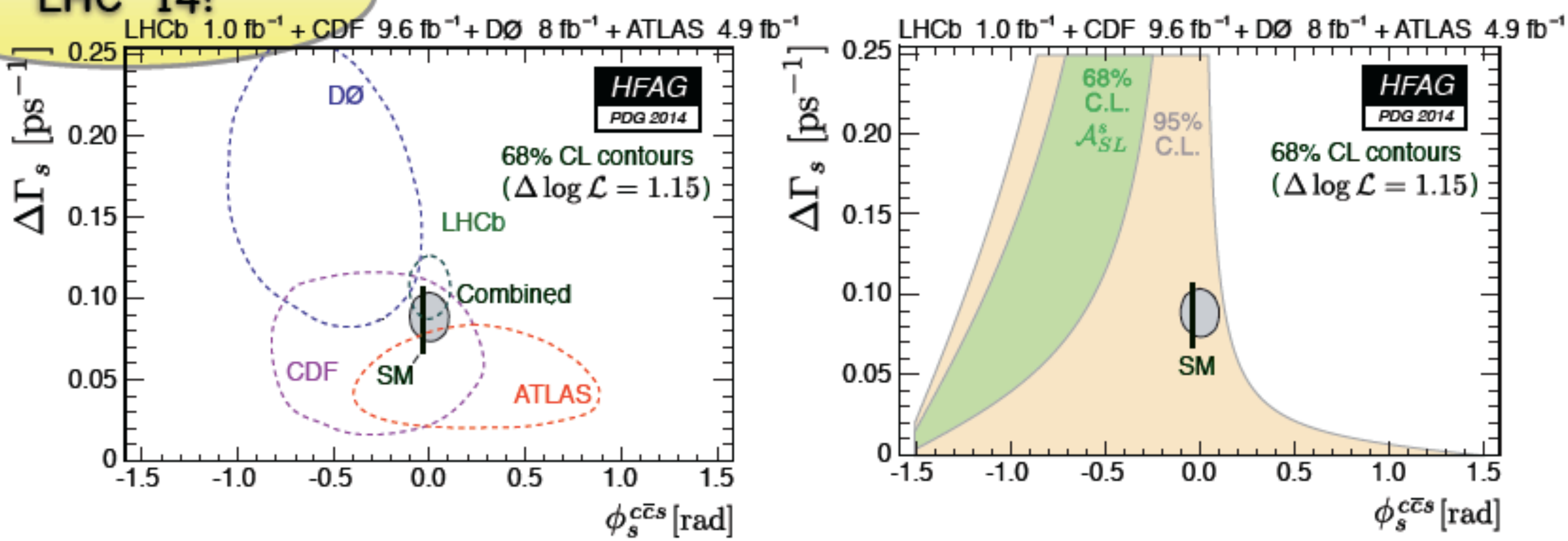
- The experimental errors are significantly reduced by the new LHCb measurement.
- The current measurement of Φs is consistent to SM (=-0.0363±0.0016):

  Φs=0.070±0.055

- LHCb has an ability to reach to the error down to δΦs=±0.025 by 2018 and the upgrade can reach to the precision of δΦs=±0.009.
- So it is too early to conclude! New physics may appear here!

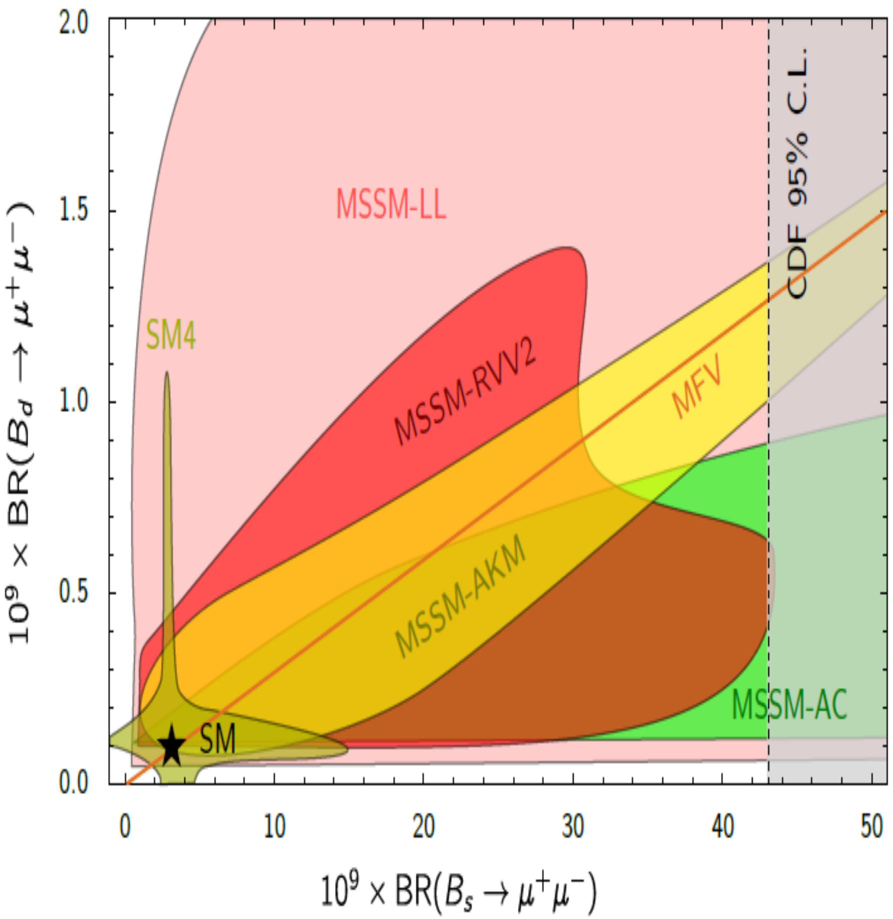
LHC '14!

E. Kou, IPA2014



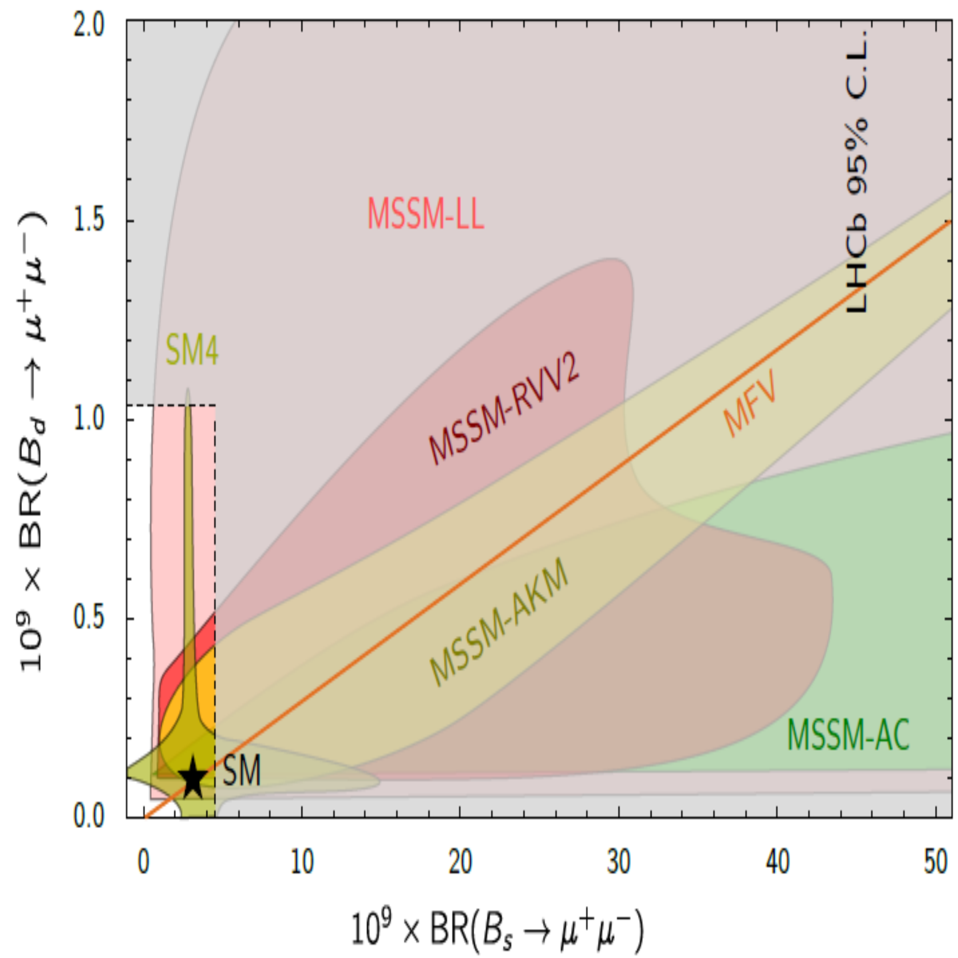
David Straub: arXiv:1205.6094

2011

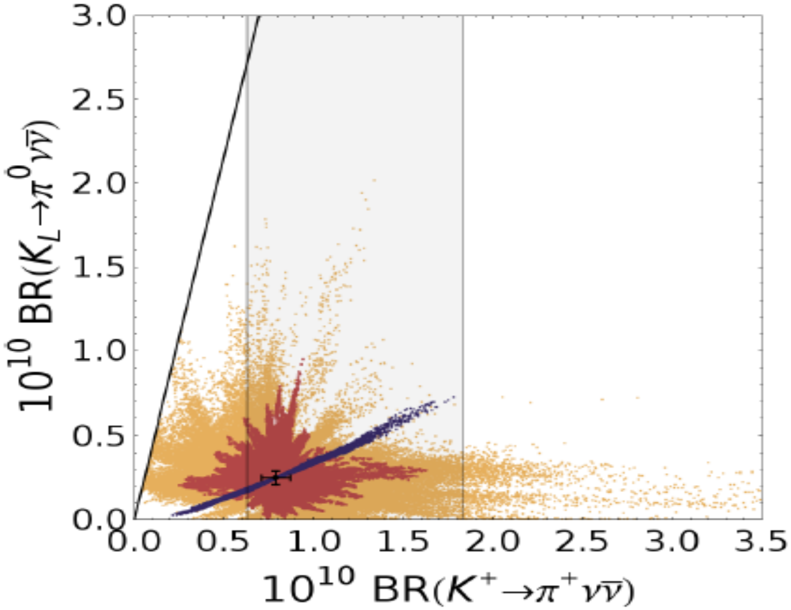


2012

ATLAS, CMS and **LHCb** results combined:
BPH-12-009, ATLAS-CONF-2012-061, LHCb-CONF-2012-017



$K^+ \to \pi^+ \nu \bar{\nu}$ vs. $K_L \to \pi^0 \nu \bar{\nu}$



triplet + anarchy

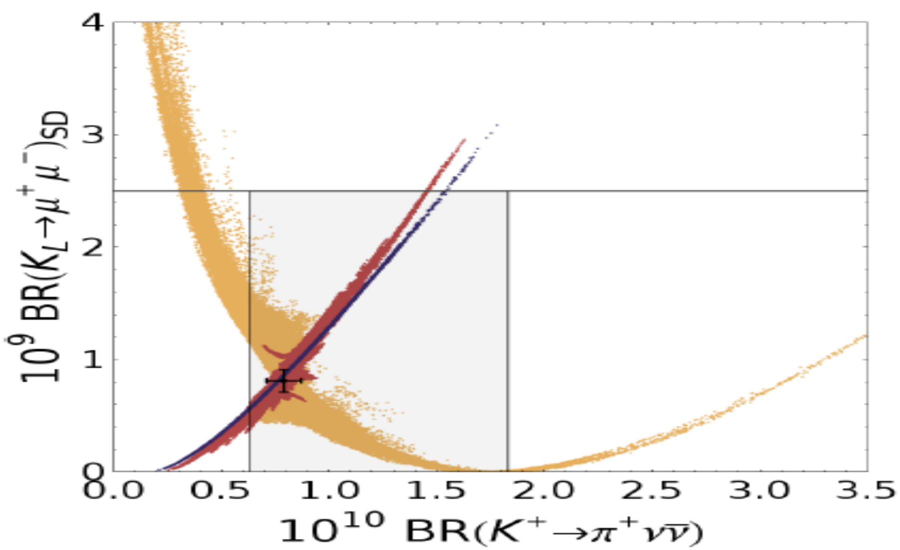
bidoublet + anarchy

bidoublet + $U(2)^3$

- Visible effects in both modes
- $U(2)^3$: aligned in phase with the SM

[Straub 1302.4651]

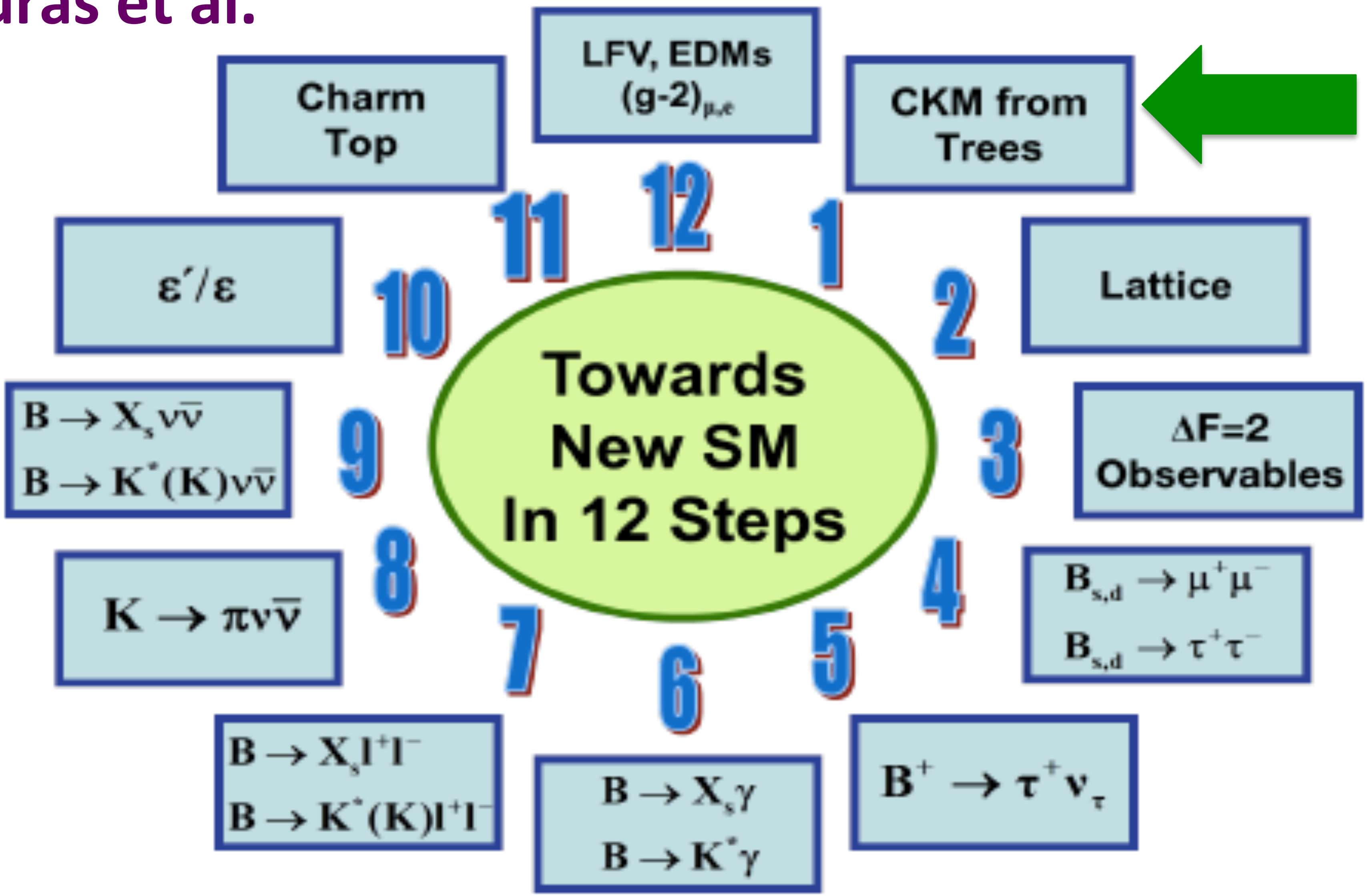
$K^+ \to \pi^+ \nu \bar{\nu}$ vs. $K_L \to \mu^+ \mu^-$



STILL ROOM FOR VISIBLE NEW PHYSICS EFFECTS FOR INSTANCE IN FCNC KAON PHYSICS IN SPITE OF ALL THE STRINGENT HIGH-ENERGY AND HIGH-INTENSITY CONSTRAINTS

The flavour clock cannot stop !!!

Buras et al.



THE FATE OF LEPTON NUMBER

L VIOLATED → ν Majorana ferm.

L CONSERVED → ν Dirac ferm. (dull option?)

SMALLNESS of m_ν

$h\, \bar{\nu}_L H\, \nu_R \rightarrow m_\nu = h\, \langle H \rangle$ $\quad M_\nu < 5$ eV $\rightarrow h < 10^{-11}$

EXTRA-DIM. ν_R in the bulk: small overlap?

↓

PRESENCE OF A NEW PHYSICAL MASS SCALE

NEW HIGH SCALE | NEW LOW SCALE

SEE - SAW MECHAN.

Minkowski; Gell-Mann, Ramond, SlansKy, Vanagida

↓

ν_R ENLARGEMENT OF THE FERMIONIC SPECTRUM

$M \nu_R \nu_R + h\, \bar{\nu}_L\, \phi\, \nu_R$

	ν_L	ν_R
ν_L	~0	$h \langle \phi \rangle$
ν_R	$h \langle \phi \rangle$	M

LR Models?

MAJORON MODELS

Gelmini, Roncadelli

↓

Δ ENLARGEMENT OF THE HIGGS SCALAR SECTOR

$h\, \nu_L \nu_L\, \Delta$

$m_\nu = h \langle \Delta \rangle$

N.B.: EXCLUDED BY LEP!

LFV ⟷ PHYSICS BSM

- LFV ⟷ NEUTRINO MASSES
- LFV ⟷ MATTER-ANTIMATTER ASYMMETRY
- LFV ⟷ GAUGE UNIFICATION
- LFV ⟷ GAUGE HIERARCHY PROBLEM

Going beyond the SM:
the NEUTRINO MASS

A. GIULIANI, SAC APPEC 2013

Cosmology, **single** and **double β decay** measure different combinations of the neutrino mass eigenvalues, constraining the **neutrino mass scale**

In a standard three active neutrino scenario:

$$\Sigma \equiv \sum_{i=1}^{3} M_i$$

cosmology
simple sum
pure kinematical effect

$$\langle M_\beta \rangle \equiv \left(\sum_{i=1}^{3} M_i^2 \, |U_{ei}|^2 \right)^{1/2}$$

β decay
incoherent sum
real neutrino

$$\langle M_{\beta\beta} \rangle \equiv \left| \sum_{i=1}^{3} M_i \, |U_{ei}|^2 \, e^{i\alpha_i} \right|$$

double β decay
coherent sum
virtual neutrino
Majorana phases

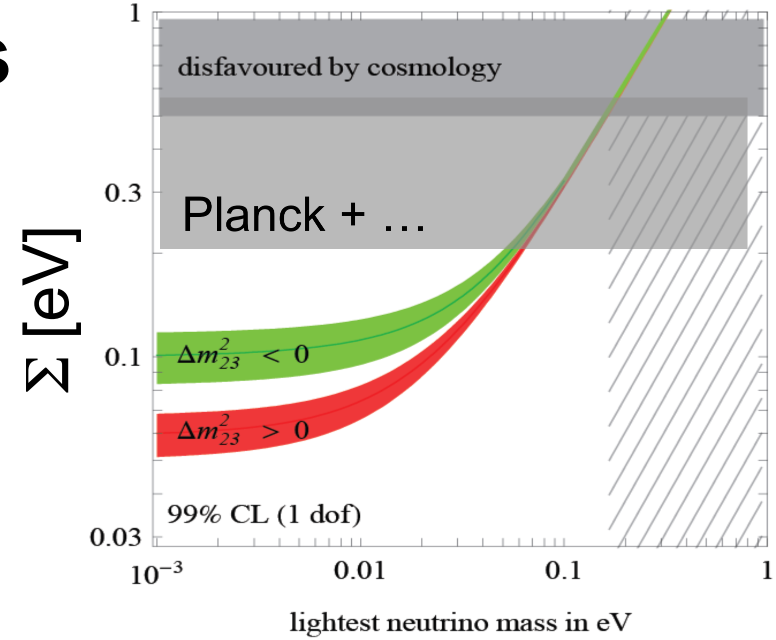
A. GIULIANI, SAC APPEC 2013

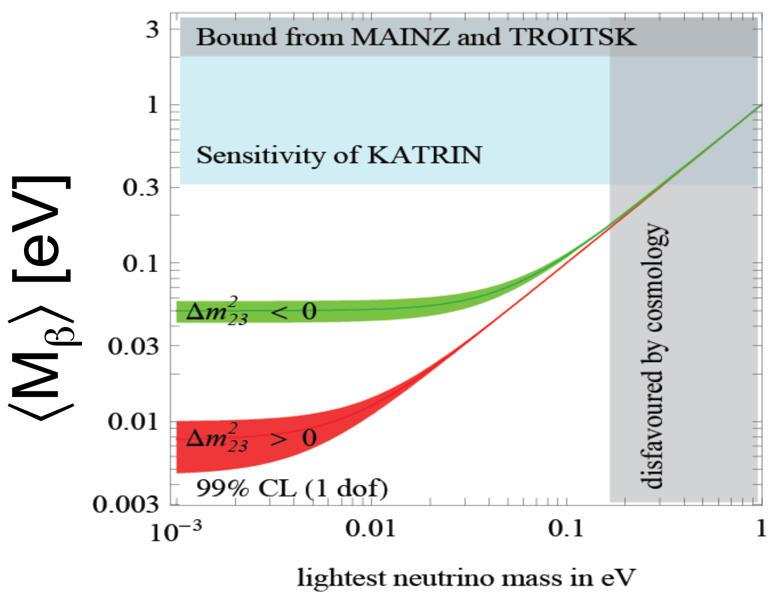
Status on neutrino mass

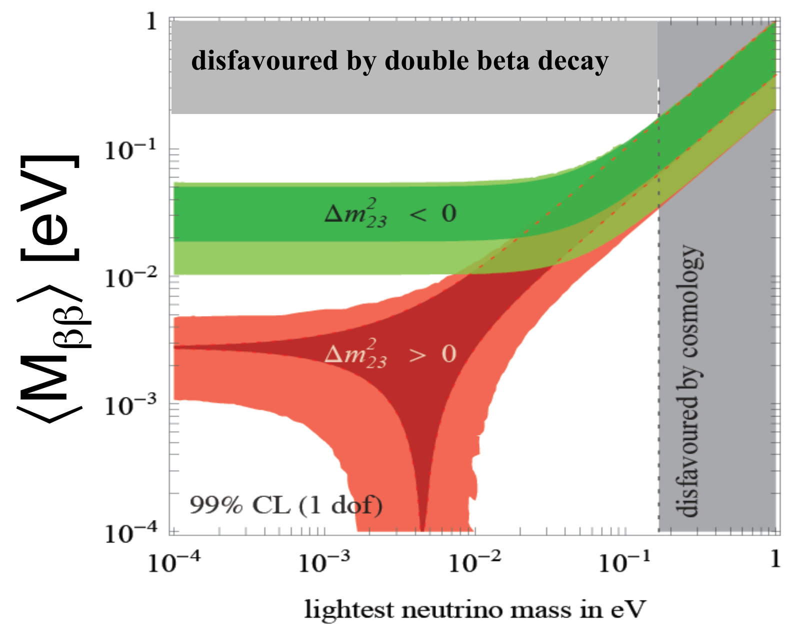
The three constrained parameters can be plotted as a function of the lightest neutrino mass

Two bands appear in each plot, corresponding to **inverted** and **direct hierarchy**

The two bands merge in the **degenerate case** (the only one presently probed)



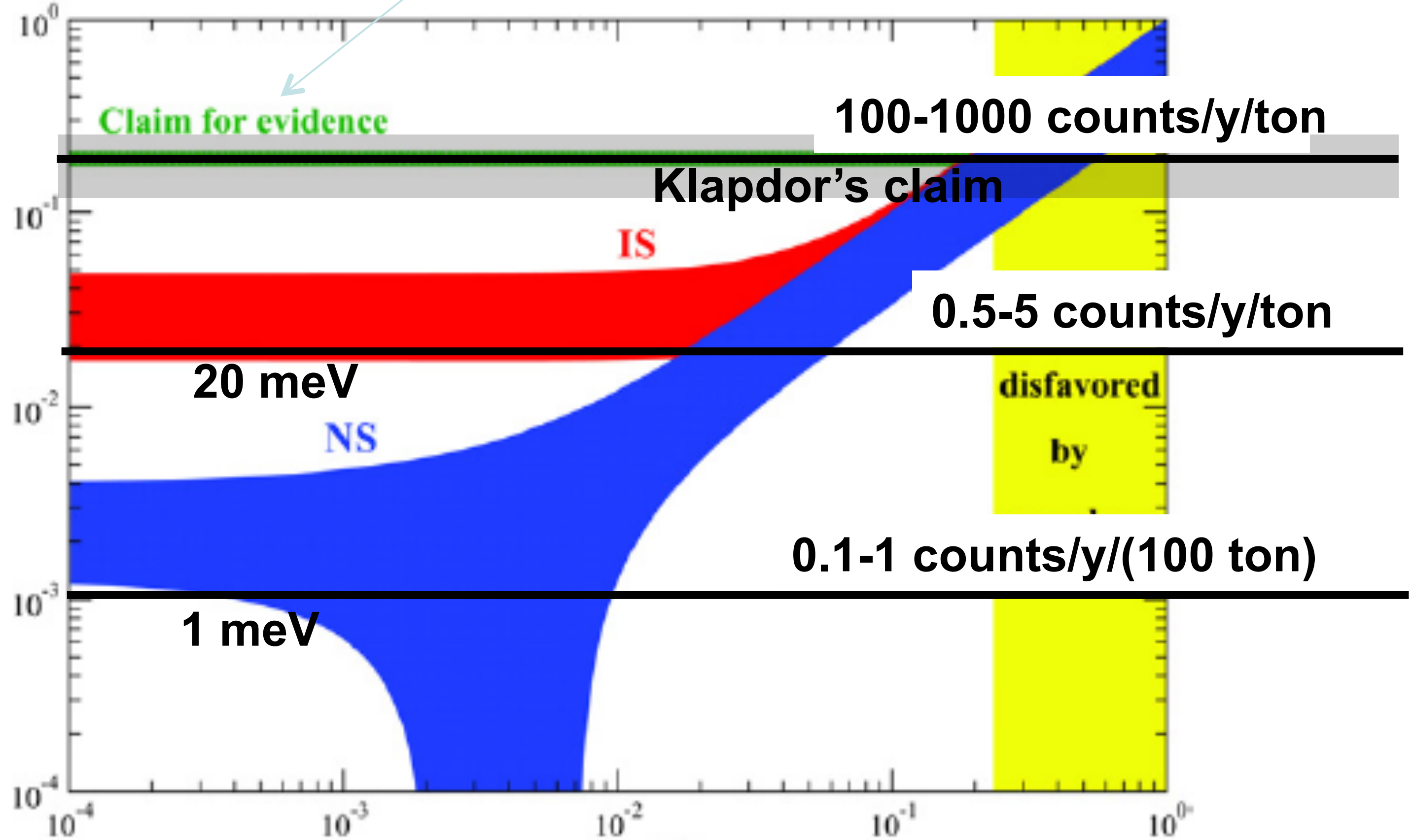




Three challenges for 0ν-DBD search

$\langle M_{\beta\beta} \rangle$ [eV]

Klapdor Krivosheina Modern Physics Letters A 21, No. 20 (2006) 1547



MATTER-ANTIMATTER ASYMMETRY ⟷ NEUTRINO MASSES CONNECTION: BARYOGENESIS THROUGH LEPTOGENESIS

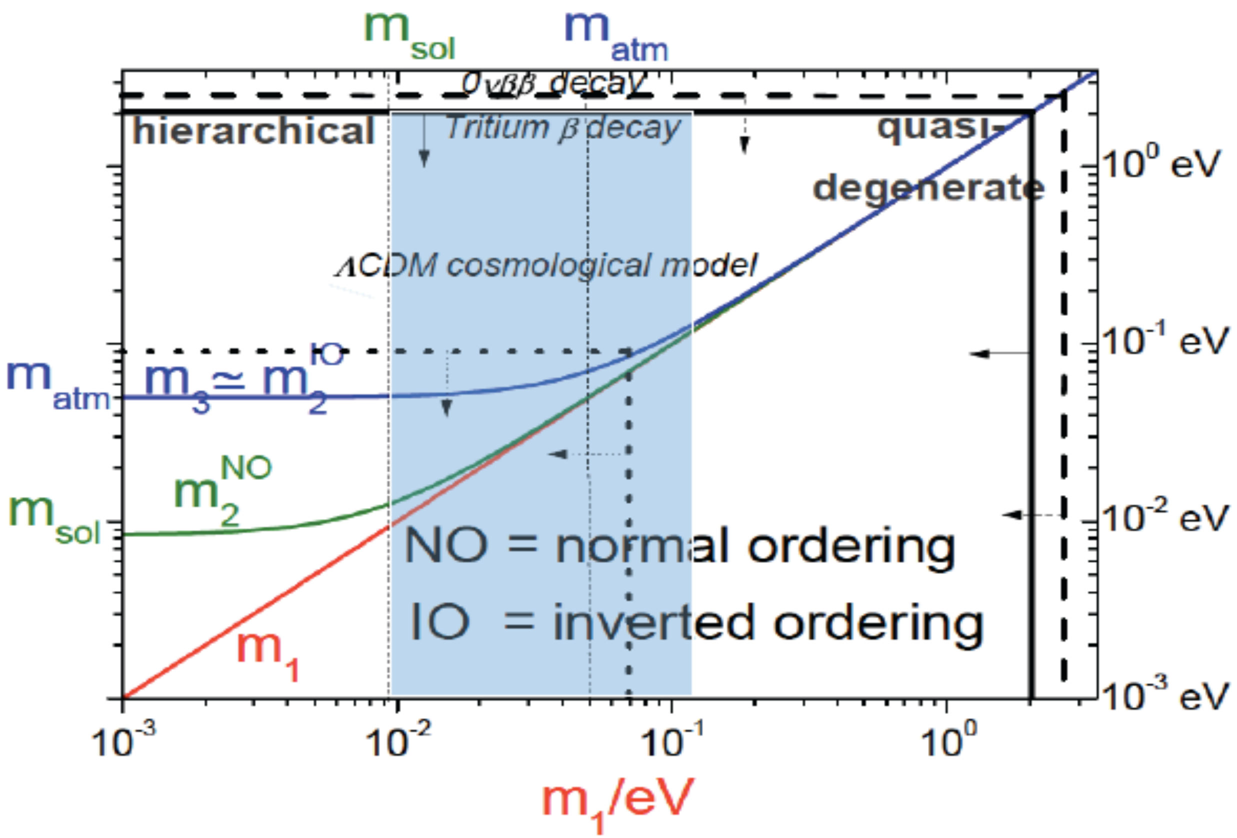
- Key-ingredient of the SEE-SAW mechanism for neutrino masses: **large Majorana mass for RIGHT-HANDED neutrino**
- In the early Universe the heavy RH neutrino decays with Lepton Number violatiion; if these decays are accompanied by a new source of CP violation in the leptonic sector, then

VANILLA LEPTOGENESISIS !

→ it is possible to create a lepton-antilepton asymmetry at the moment RH neutrinos decay. Since SM interactions preserve Baryon and Lepton numbers at all orders in perturbation theory, but violate them at the quantum level, such **LEPTON ASYMMETRY** can be converted by these purely quantum effects into a BARYON-ANTIBARYON ASYMMETRY (**Fukugita-Yanagida mechanism for leptogenesis**)

A new neutrino mass window for leptogenesis

FLAVOURED LEPTOGENESIS !



$$0.01 \text{ eV} \lesssim m_1 \lesssim 0.1 \text{ eV}$$

P. Di Bari, Sophie King, M. Re Fiorentin 2014

LFV IN SUSY SEE-SAW

SEE- SAW (type 1) **LOW-ENERGY SUSY**

New source of (leptonic) flavor:

YUKAWA COUPLINGS OF THE NEUTRINO DIRAC MASS CONTRIBUTIONS, i.e. **THE YUKAWAs** of the **HIGGS couplings to the LETF- and RIGHT – HANDED NEUTRINOS**

The scalar lepton masses through their running bring memory of those new sources of leptonic flavor at the TeV scale, i.e. at energies much below the (Majorana) mass of the RH neutrinos

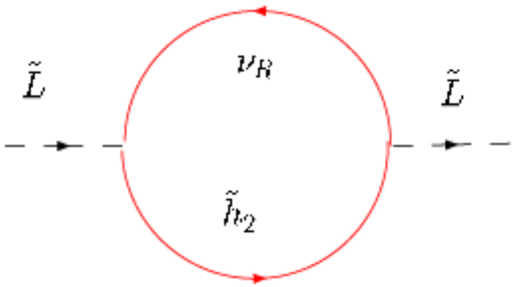
THE STRONG ENHANCEMENT OF LFV IN SUSY SEESAW MODELS CAN OCCUR

EVEN IF THE MECHANISM RESPONSIBLE FOR SUSY BREAKING IS ABSOLUTELY FLAVOR BLIND

SUSY SEESAW: *Flavor universal SUSY breaking and yet large lepton flavor violation*

Borzumati, A. M.

$$L = f_l \, \bar{e}_R L h_1 + f_\nu \, \bar{\nu}_R L h_2 + M \, \nu_R \nu_R$$



$$\left(m^2_{\tilde{L}}\right)_{ij} \simeq \frac{1}{8\pi^2}\left(3m_0^2 + A_0^2\right)\left(f_\nu^\dagger f_\nu\right)_{ij} \log\frac{M}{M_G}$$

Non-diagonality of the slepton mass matrix in the basis of diagonal lepton mass matrix depends on the **unitary matrix U which diagonalizes ($f_\nu^+ f_\nu$)**

How Large LFV in SUSY SEESAW?

- 1) Size of the **Dirac neutrino couplings f_ν**
- 2) Size of the **diagonalizing matrix U**

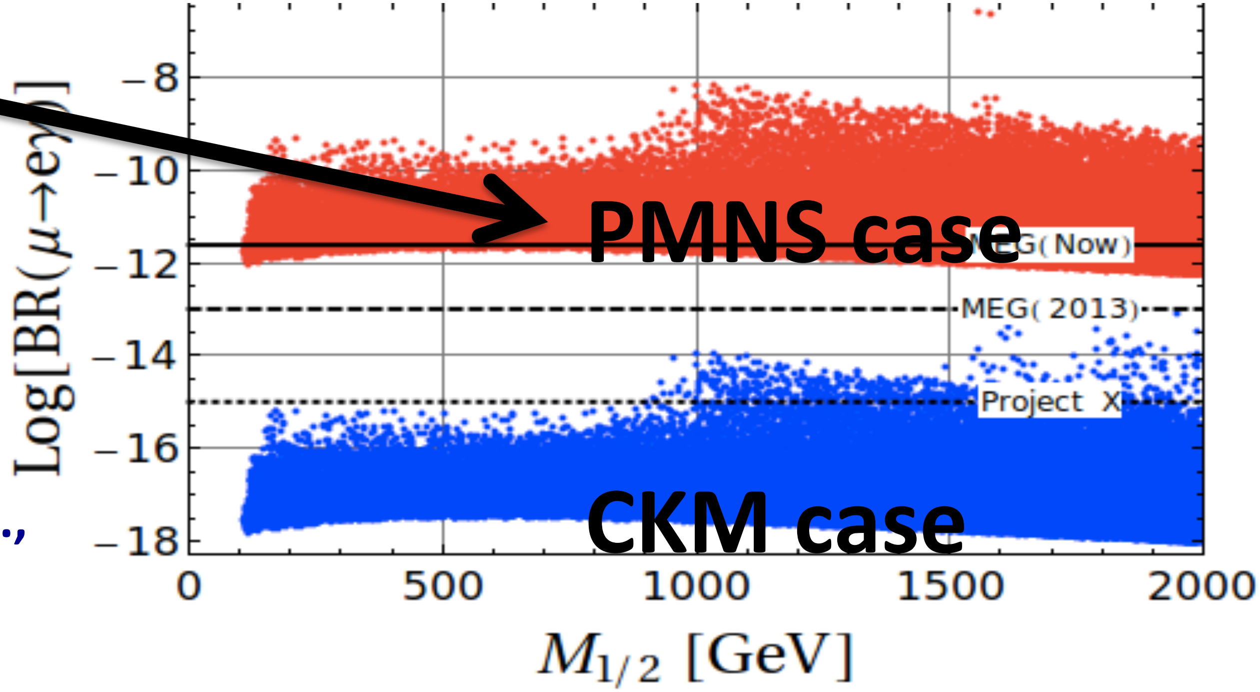
In **MSSM seesaw** or in **SUSY SU(5)** (Moroi): not possible to correlate the neutrino Yukawa couplings to know Yukawas;

In **SUSY SO(10)** (A.M., Vempati, Vives) at least one neutrino Dirac Yukawa coupling has to be of the **order of the top Yukawa coupling** one large of O(1) f_ν

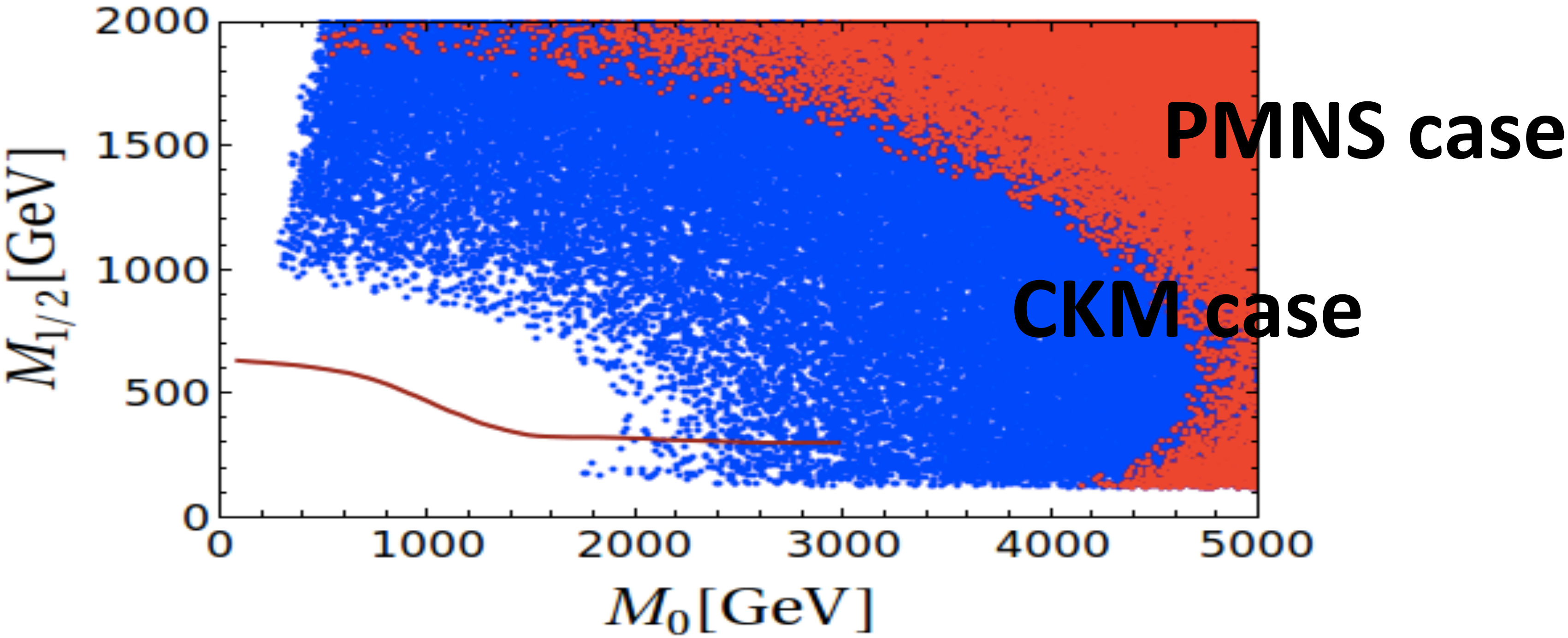
U → two "extreme" cases:

a) **U with "small" entries** → ***U = CKM***;
b) **U with "large" entries** with the exception of the 13 entry → ***U = PMNS*** matrix responsible for the diagonalization of the neutrino mass matrix

PMNS case in mSUGRA with $\tan\beta = 10$

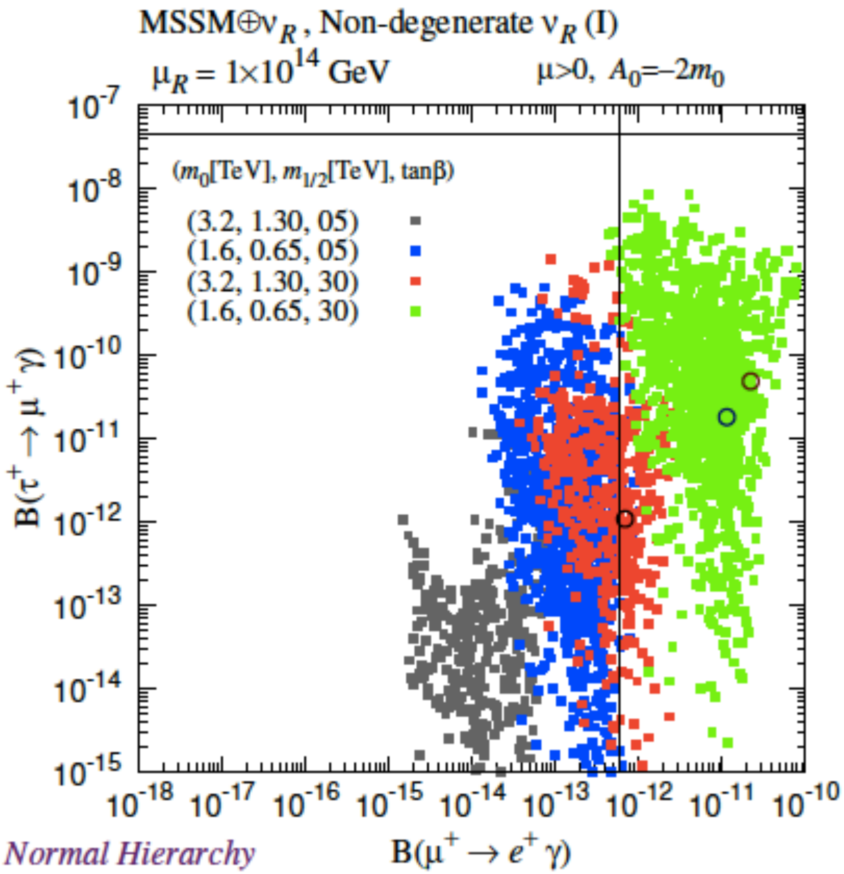


Calibbi, Chowdhuri, A. M., Patel, Vempati 2012

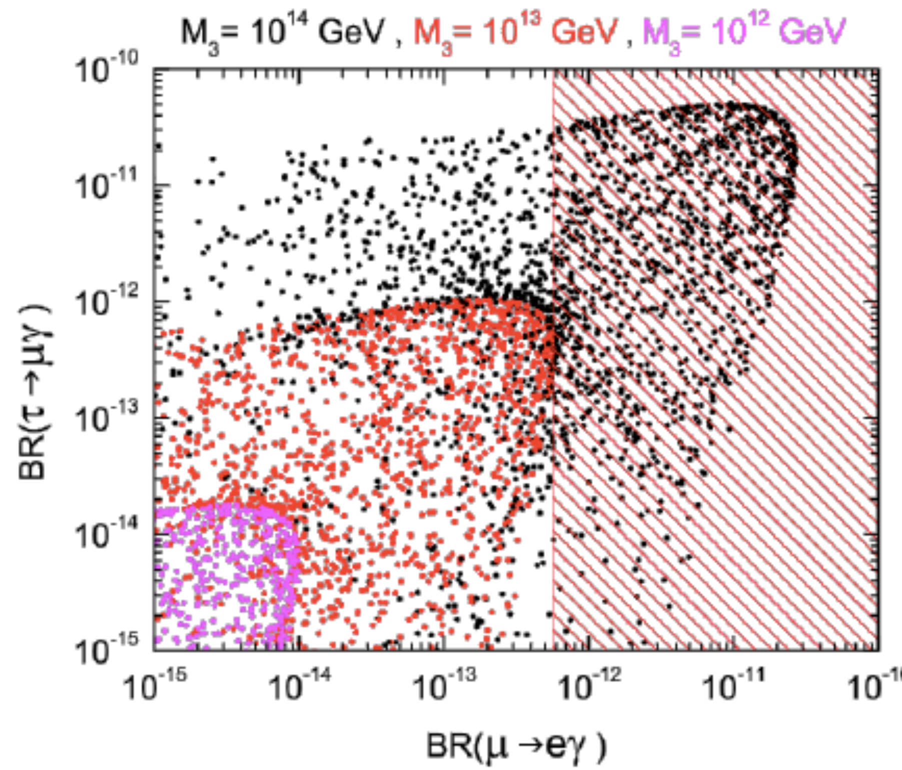


SUSY Type I seesaw model

SUSY parameters fixed, ν_R parameters varied:

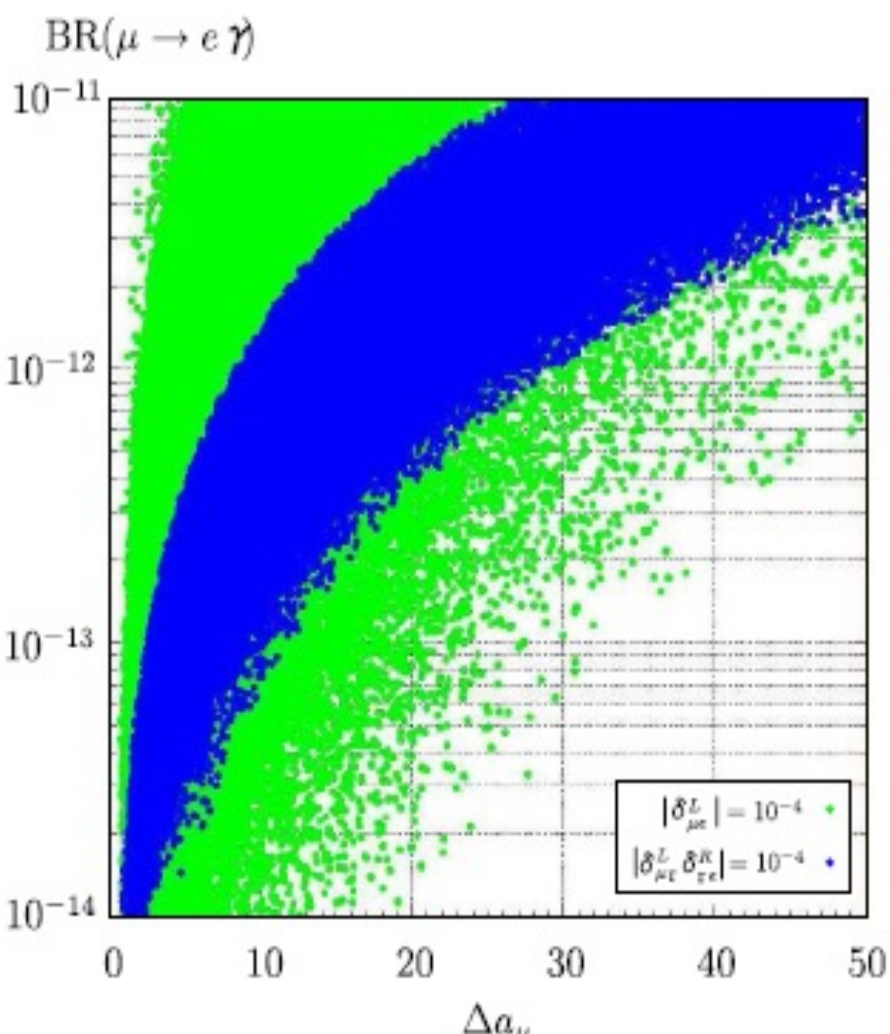


[Goto *et al.*, FLASY2013]

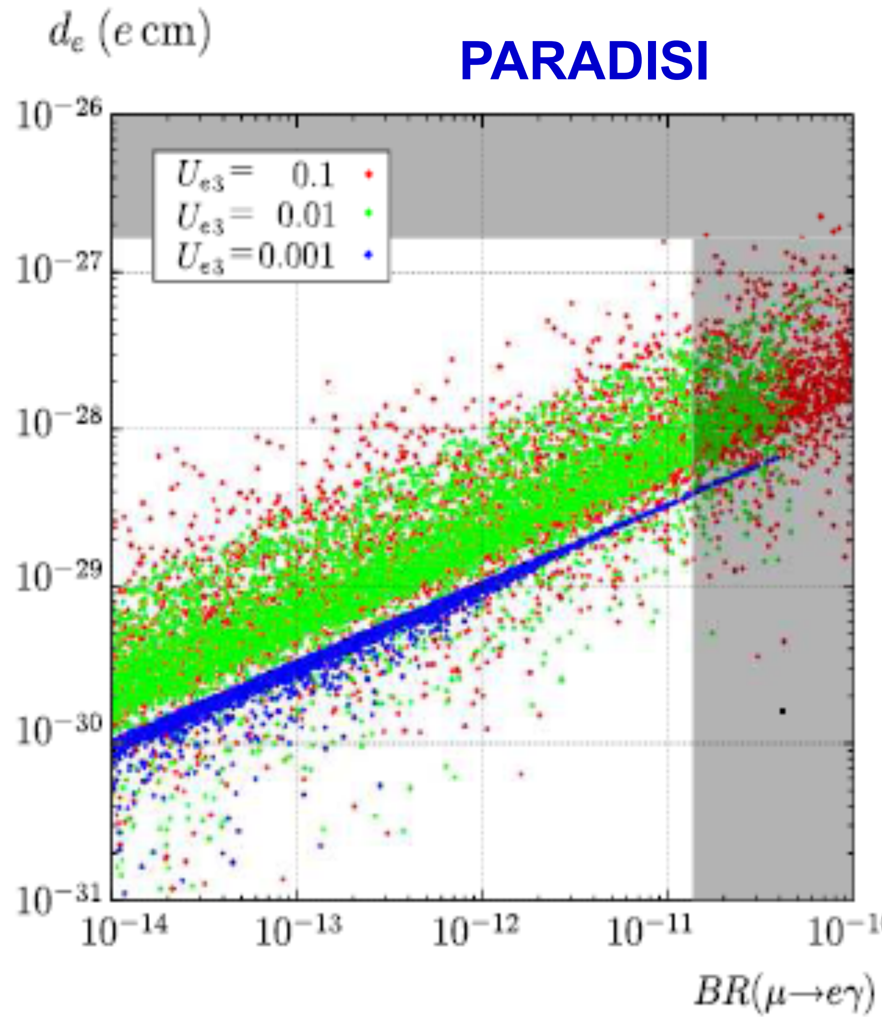


[Joaquim, ICHEP2014]

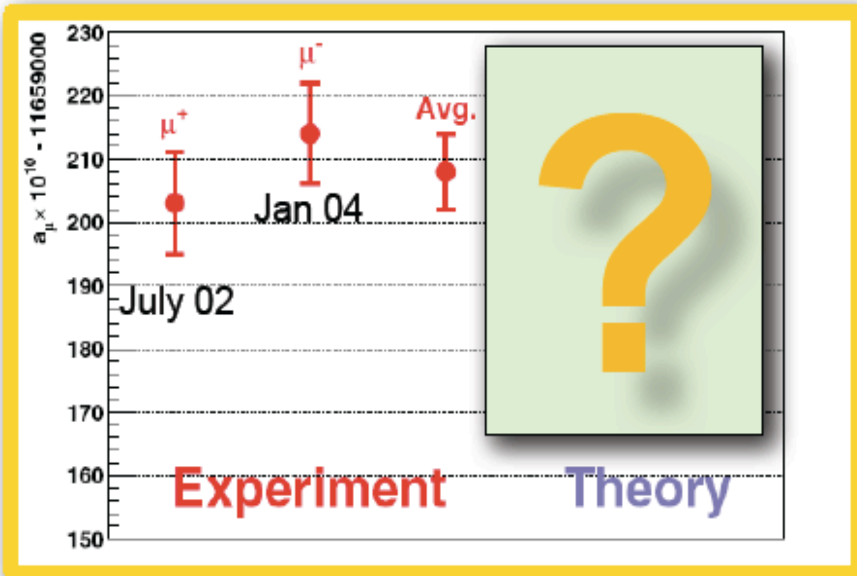
LFV, g – 2, EDM: a promising correlation in SUSY SEESAW



PARADISI



The muon g-2: the experimental result



- **Today:** $a_\mu^{EXP} = (116592089 \pm 54_{stat} \pm 33_{sys}) \times 10^{-11}$ [0.5ppm].
- **Future:** new muon g-2 experiments proposed at:
  - Fermilab E989, aiming at $\pm 16 \times 10^{-11}$, ie 0.14ppm
  - J-PARC aiming at 0.1 ppm

  See B. Lee Roberts & T. Mibe @ Tau2012, September 2012

Sep 2012:
CD0 approval!
Data in (late)
2016?

- Are theorists ready for this (amazing) precision? No(t yet)

The muon g-2: SM vs. Experiment

Adding up all contributions, we get the following SM predictions and comparisons with the measured value:

$a_\mu^{EXP} = 116592089\ (63) \times 10^{-11}$

E821 – Final Report: PRD73 (2006) 072 with latest value of $\lambda=\mu_\mu/\mu_p$ from CODATA'06

$a_\mu^{SM} \times 10^{11}$	$\Delta a_\mu = a_\mu^{EXP} - a_\mu^{SM}$	σ
116 591 794 (66)	$295\ (91) \times 10^{-11}$	3.2 [1]
116 591 814 (57)	$275\ (85) \times 10^{-11}$	3.2 [2]
116 591 840 (58)	$249\ (86) \times 10^{-11}$	2.9 [3]

with the "conservative" a_μ^{HHO}(lbl) = 116 (39) x 10^{-11} and the LO hadronic from:

[1] Jegerlehner & Nyffeler, Phys. Rept. 477 (2009) 1
[2] Davier et al, EPJ C71 (2011) 1515 (includes BaBar & KLOE10 2π)
[3] Hagiwara et al, JPG38 (2011) 085003 (includes BaBar & KLOE10 2π)

Note that the th. error is now about the same as the exp. one

THE EDM CHALLENGE

FOR **ANY NEW PHYSICS AT THE TEV SCALE** WITH **NEW SOURCES OF CP VIOLATION** → NEED FOR **FINE-TUNING** TO PASS THE EDM TESTS OR SOME **DYNAMICS TO SUPPRESS THE CPV** IN FLAVOR CONSERVING EDMS

Current and projected sensitivities

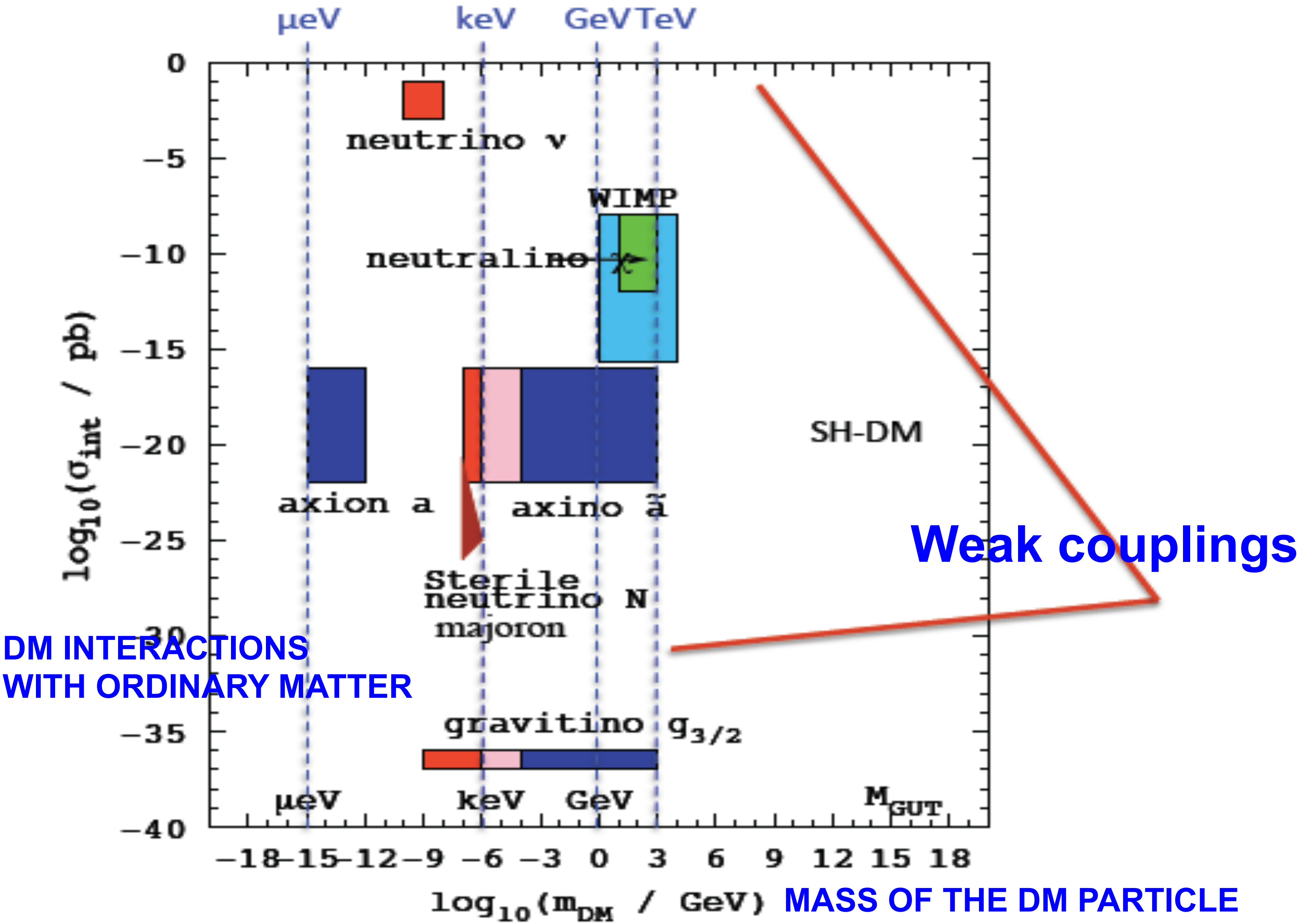
C. GRIFFITH, IPA2014

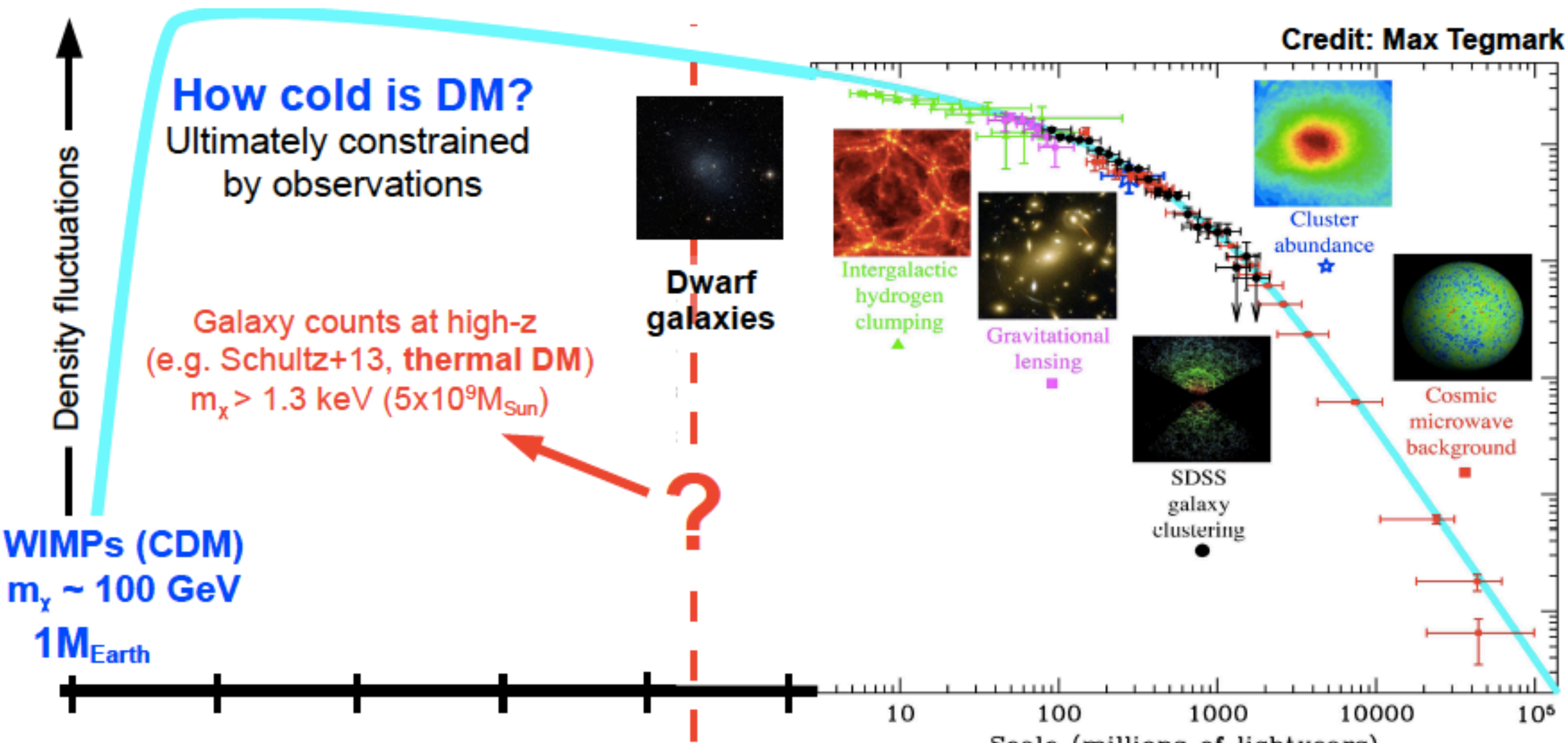
	current limit	projected sens. from planned exp.	standard model CKM prediction
n	3×10^{-26}	10^{-28}	$10^{-31} - 10^{-33}$
e	9×10^{-29}	10^{-30}	$\sim10^{-38}$
Hg	3×10^{-29}	10^{-30}	$<10^{-35}$

(units are in e.cm)

***THE DM ROAD TO NEW PHYSICS BEYOND THE SM*:**

IS **DM** A PARTICLE OF THE ***NEW PHYSICS AT THE ELECTROWEAK ENERGY SCALE*** **?**


μeV
keV
GeV
TeV
neutrino ν
WIMP
neutralino χ
SH-DM
axion a
axino ã
Sterile
neutrino N
majoron
Weak couplings
DM INTERACTIONS
WITH ORDINARY MATTER
gravitino $g_{3/2}$
μeV
keV
GeV
M_{GUT}
$\log_{10}(\sigma_{int}$ / pb)
$\log_{10}(m_{DM}$ / GeV)
MASS OF THE DM PARTICLE
0
−5
−10
−15
−20
−25
−30
−35
−40
−18
−15
−12
−9
−6
−3
0
3
6
9
12
15
18



Credit: Max Tegmark
How cold is DM?
Ultimately constrained by observations
Density fluctuations
Dwarf galaxies
Galaxy counts at high-z
(e.g. Schultz+13, thermal DM)
$m_\chi > 1.3$ keV ($5x10^9 M_{Sun}$)
?
WIMPs (CDM)
$m_\chi \sim 100$ GeV
$1 M_{Earth}$
Intergalactic hydrogen clumping
Gravitational lensing
Cluster abundance
SDSS galaxy clustering
Cosmic microwave background
10
100
1000
10000
10^5
Scale (millions of lightyears)

CDM Controversies?

- Cusp-vs-Core problem
- Missing satellites problem
- To-big-to-fail problem

Possible solutions

- Baryonic physics:
  gas cooling, star formation,
  supernova feedback,…
- Dark Matter:
  warm dark matter
  Decaying DM
  Self-Interacting DM

Spergel et al, Sigurdson et al,
Boehm et al, Kaplinghat et al,
Loeb et al, Tulin et al,
van de Aarseen et al,
….

Y. TANG at this meeting

Example:$\nu\Lambda$MDM

P. Ko and Y. Tang 2014

two right-handed gauge singlets,
a dark sector with an extra $U(1)_X$ gauge
symmetry,

CONNECTION DM – ELW. SCALE

THE WIMP MIRACLE :STABLE ELW. SCALE WIMPs

	SUSY	EXTRA DIM.	LITTLE HIGGS.
1) **ENLARGEMENT OF THE SM**	(x^{μ}, θ)	(x^{μ}, j^{i})	SM part + new part
	Anticomm. Coord.	New bosonic Coord.	to cancel Λ^2 at 1-Loop
2) **SELECTION RULE**	***R-PARITY LSP***	***KK-PARITY LKP***	***T-PARITY LTP***
→ **DISCRETE SYMM.**	Neutralino spin 1/2	spin1	spin0
→ **STABLE NEW PART.**	↓	↓	↓
3) FIND REGION (S) PARAM. SPACE WHERE THE “L” NEW PART. IS NEUTRAL + $\Omega_L h^2$ OK	m_{LSP} ~100 - 200 GeV	m_{LKP} ~600 - 800 GeV	m_{LTP} ~400 - 800 GeV

Info to extract from the **direct searches**

Y. Kahn, IPA2014
Kahn, McCullough, Fox 2014

$$= g(v_{min})$$

$$\frac{dR}{dE_R} = \frac{\rho_\chi \sigma_n}{2 m_\chi \mu_{n\chi}^2} N_A m_n C_T^2(A,Z) \int dE'_R G(E_R, E'_R) \epsilon(E'_R) F^2(E'_R) \int_{v_{min}(E'_R)}^{\infty} \frac{f(\mathbf{v}+\mathbf{v}_E)}{v} d^3 v$$

DM model | detector properties | nuclear physics | DM halo model

$v_{min}(E'_R)$: min. DM velocity required for nuclear recoil E'_R

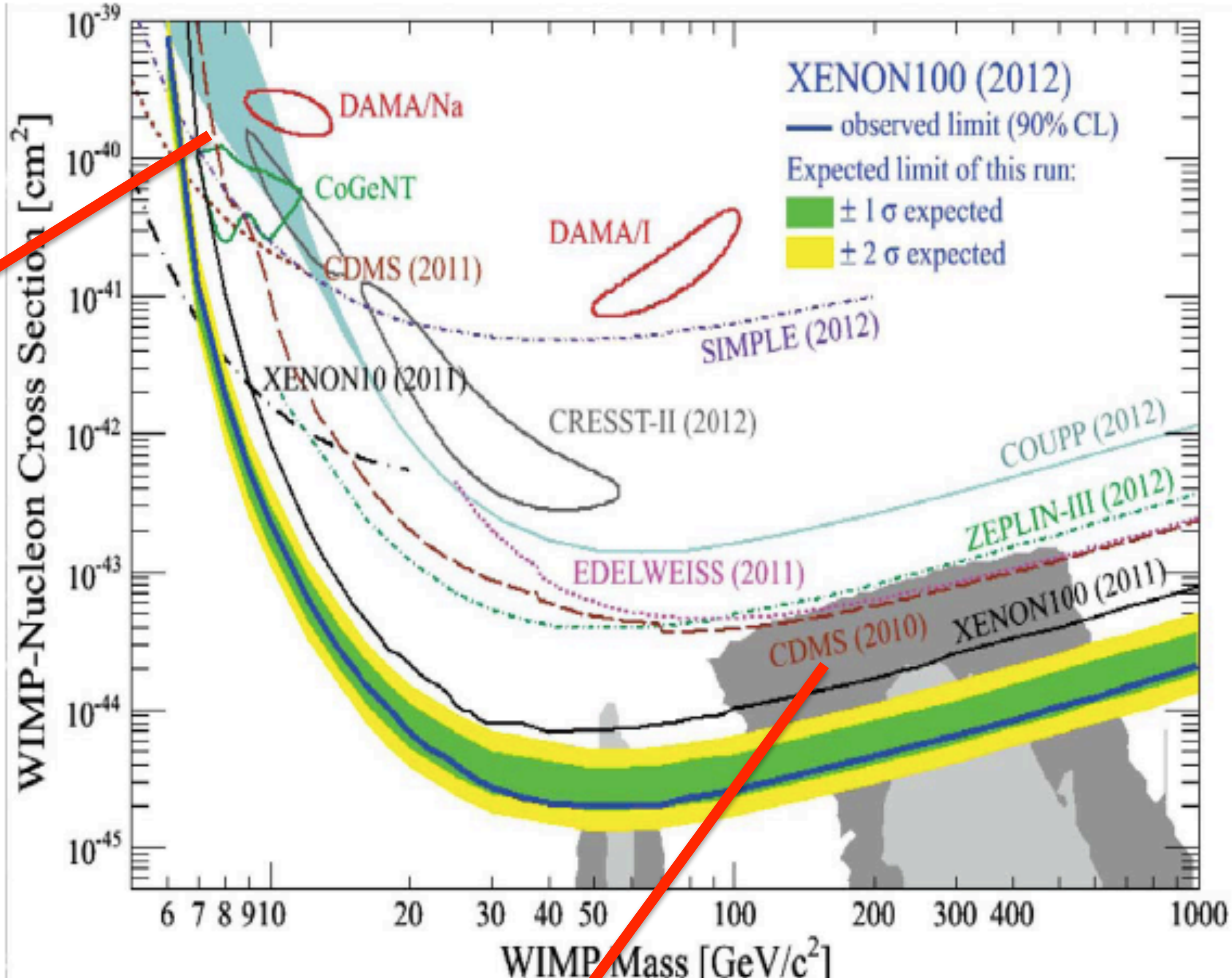
Usual method: DM model + halo model → limits/preferred values in $m_\chi - \sigma_n$ space

Halo-independent: DM model → limits/preferred values in $v_{min} - g(v_{min})$ space

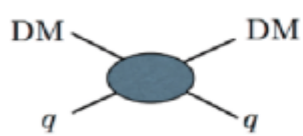
No assumptions about DM halo, easy to compare multiple experiments (esp. signal vs. exclusion)

Low-mass region:
either unexplained backgrounds in DAMA, CoGeNT, and CRESST-II, …
or
… other experiments do not understand low recoil energy calibration, …
or
… can't compare different experiments

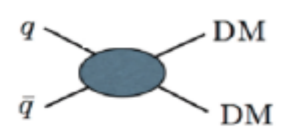
Kolb SUSY2012



Relevant to intensify the efforts here: ex. **asymmetric DM** with **DM particles of mass~ baryon mass given that ρ_{DM} not much different from ρ_B**

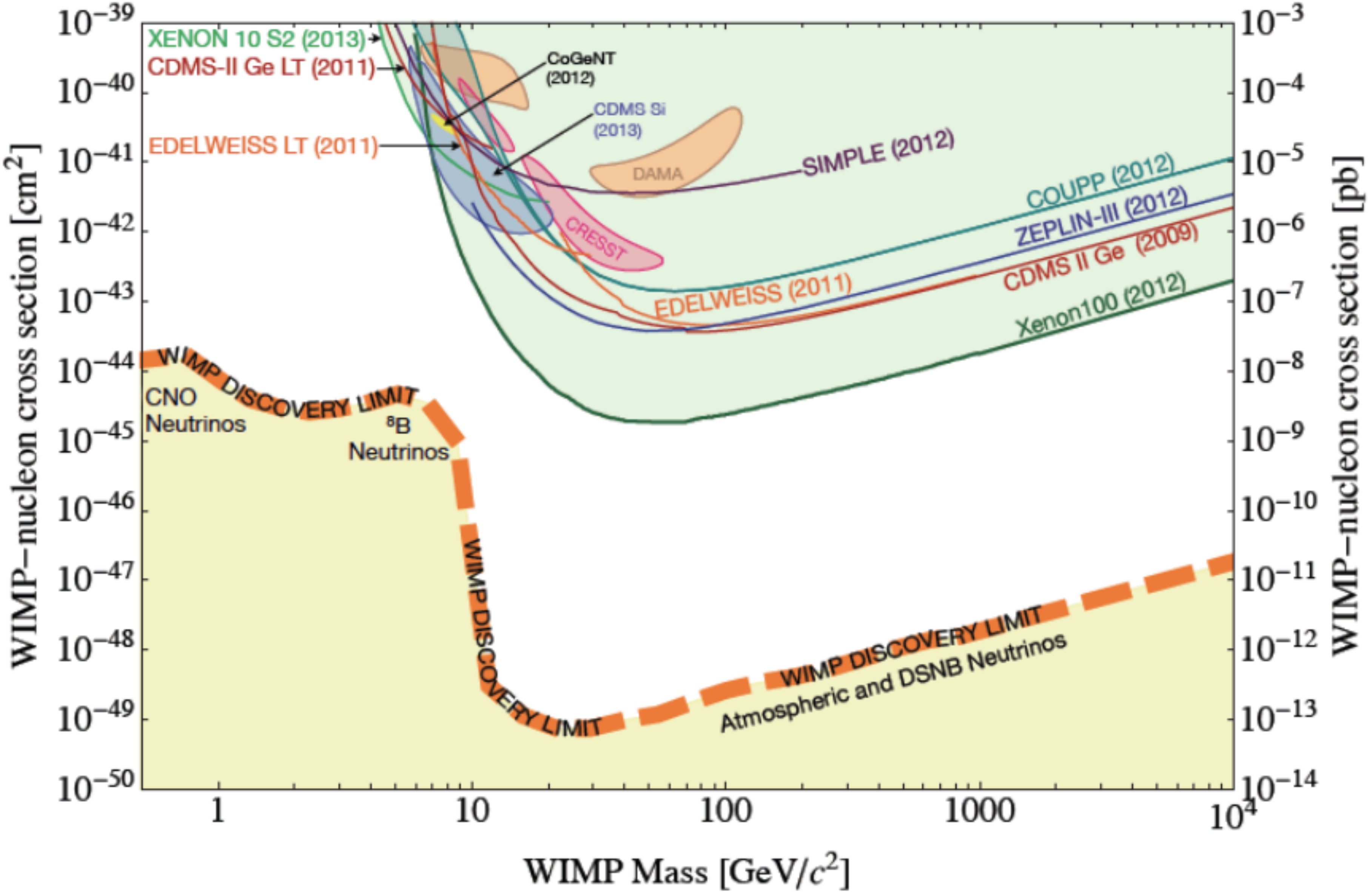


Direct Detection (t-channel)

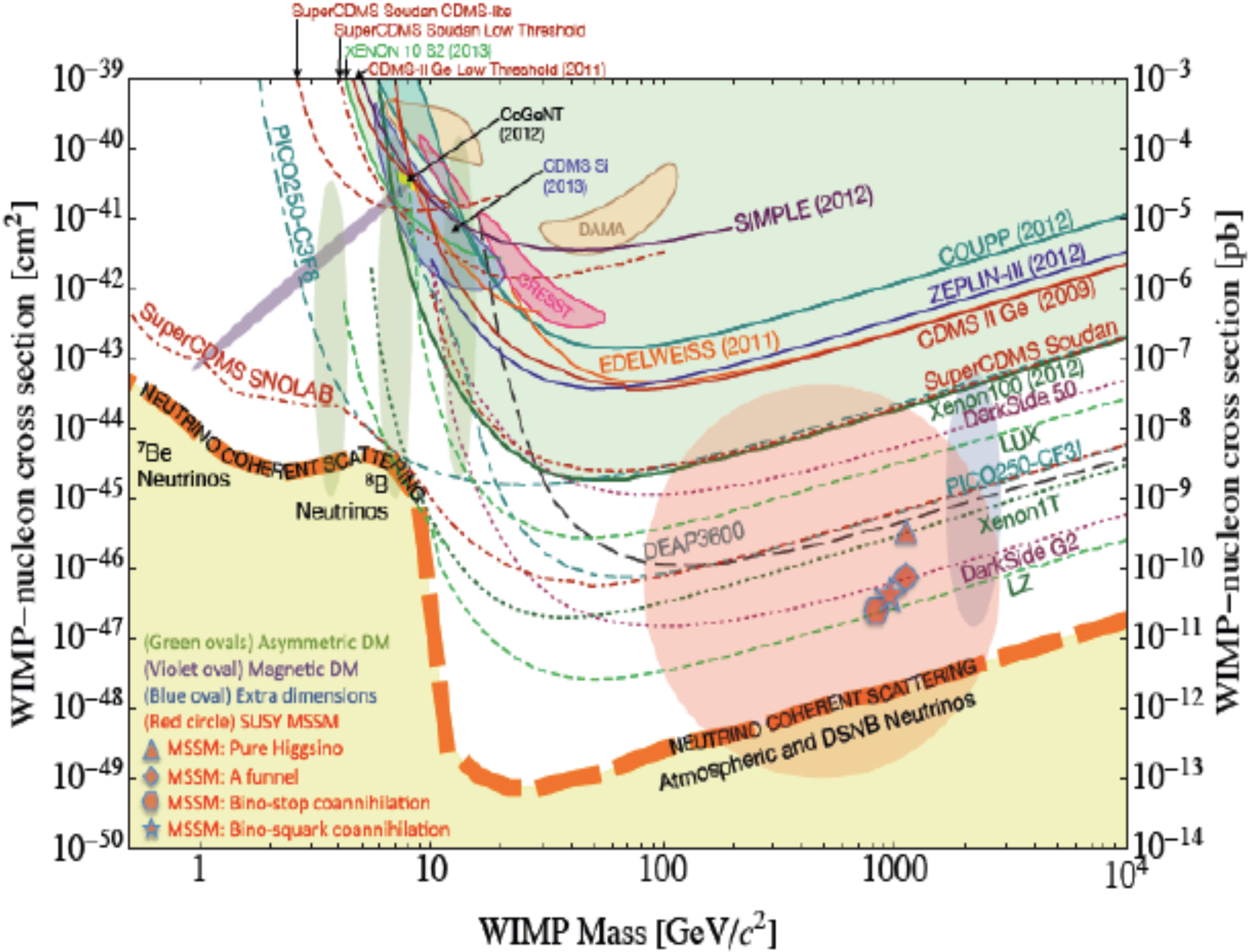


Collider Searches (s-channel)

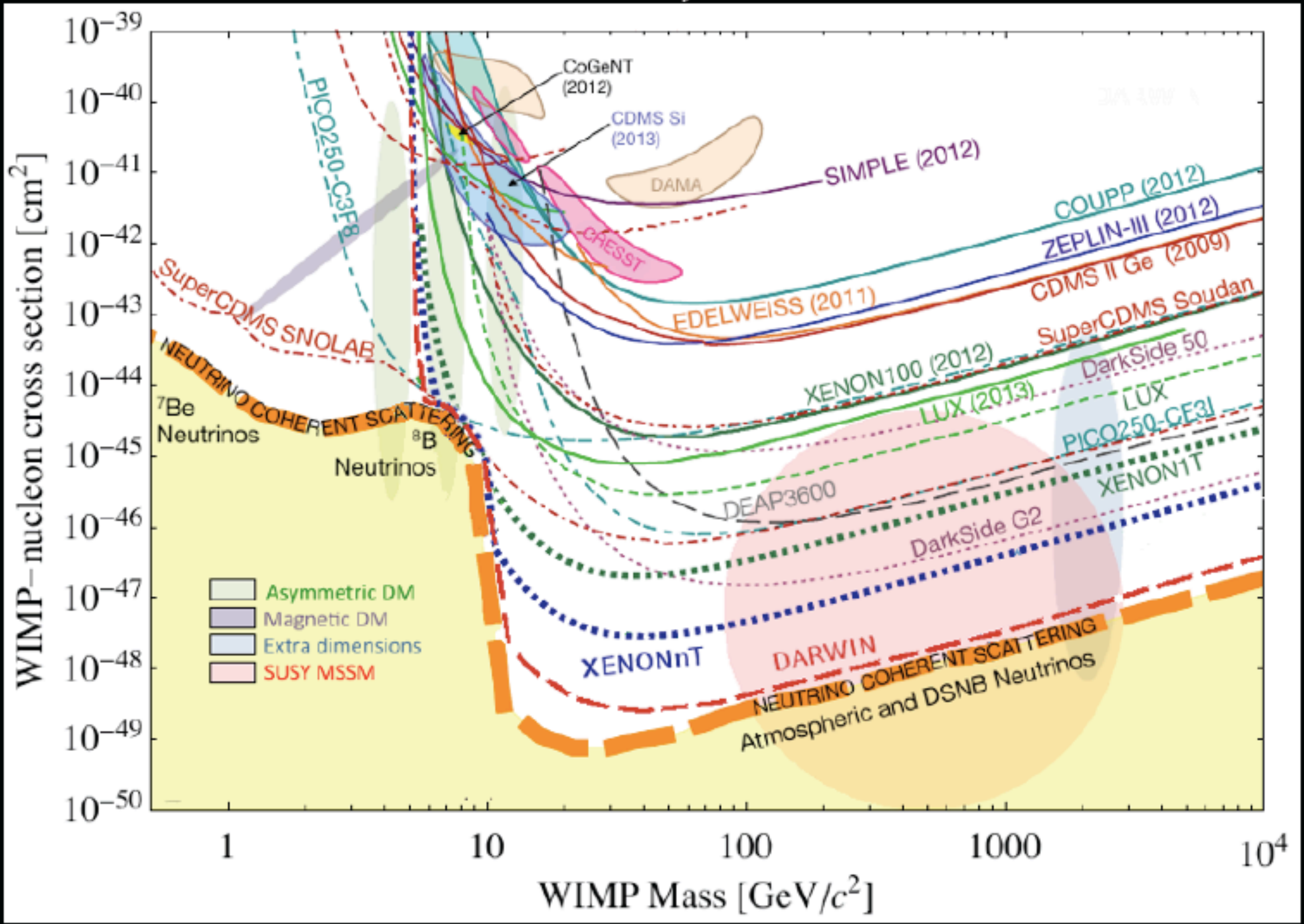
Spin-Independent Cross Section: Current Experiment Results



so far: ~3 years / order of magnitude

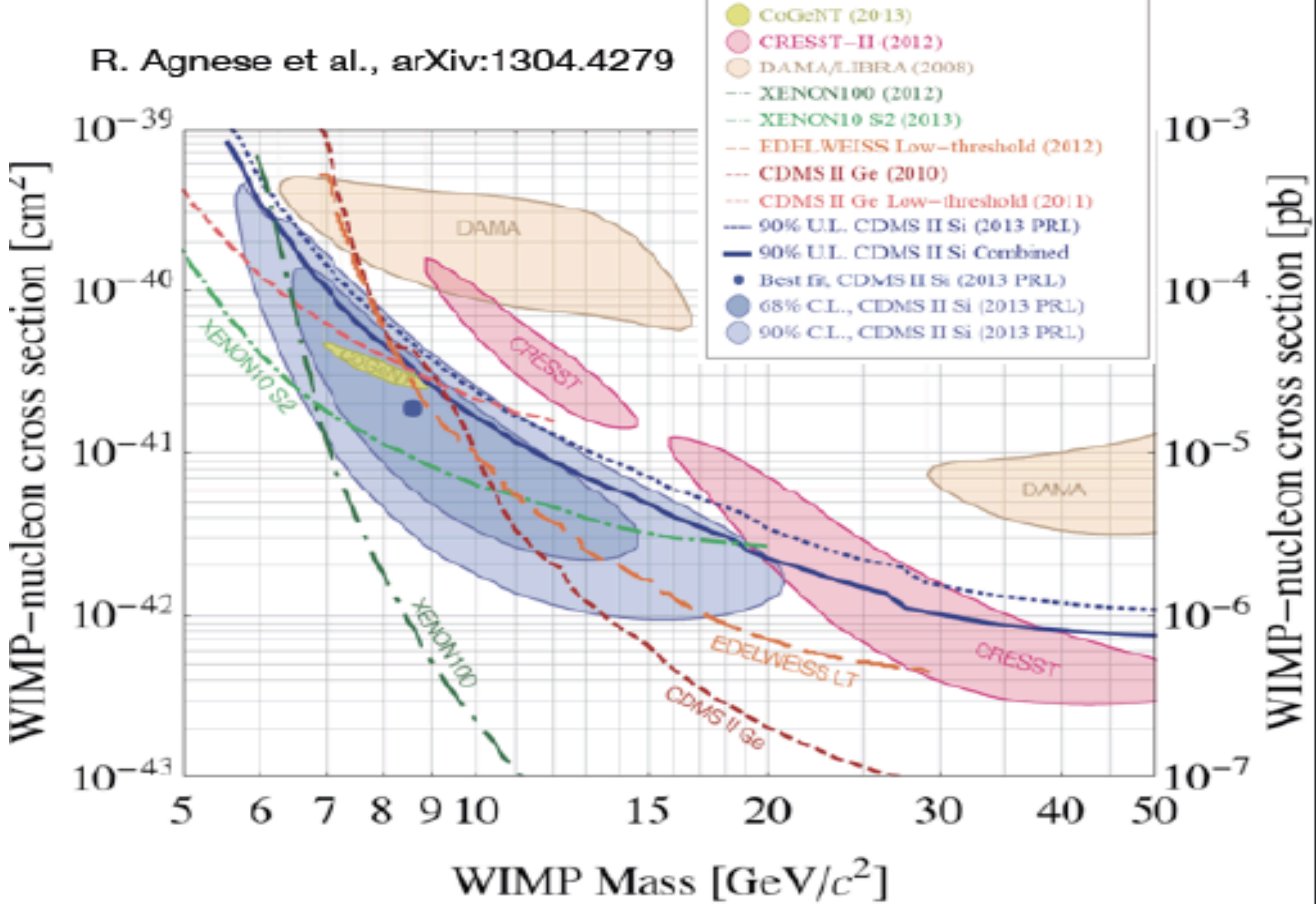

SuperCDMS Soudan CDMS-lite
SuperCDMS Soudan Low Threshold
XENON 10 S2 (2013)
CDMS-II Ge Low Threshold (2011)
CoGeNT (2012)
CDMS Si (2013)
PICO250-C3F8
DAMA
SIMPLE (2012)
COUPP (2012)
ZEPLIN-III (2012)
CDMS II Ge (2009)
CRESST
EDELWEISS (2011)
SuperCDMS Soudan
Xenon100 (2012)
DarkSide 50
LUX
PICO250-CF3I
Xenon1T
DarkSide G2
LZ
DEAP3600
SuperCDMS SNOLAB
NEUTRINO COHERENT SCATTERING
^{7}Be Neutrinos
^{8}B Neutrinos
Atmospheric and DSNB Neutrinos
(Green ovals) Asymmetric DM
(Violet oval) Magnetic DM
(Blue oval) Extra dimensions
(Red circle) SUSY MSSM
MSSM: Pure Higgsino
MSSM: A funnel
MSSM: Bino-stop coannihilation
MSSM: Bino-squark coannihilation
WIMP–nucleon cross section [cm^2]
WIMP–nucleon cross section [pb]
WIMP Mass [GeV/c^2]

1) Science Goals: Dark Matter Projected Sensitivities

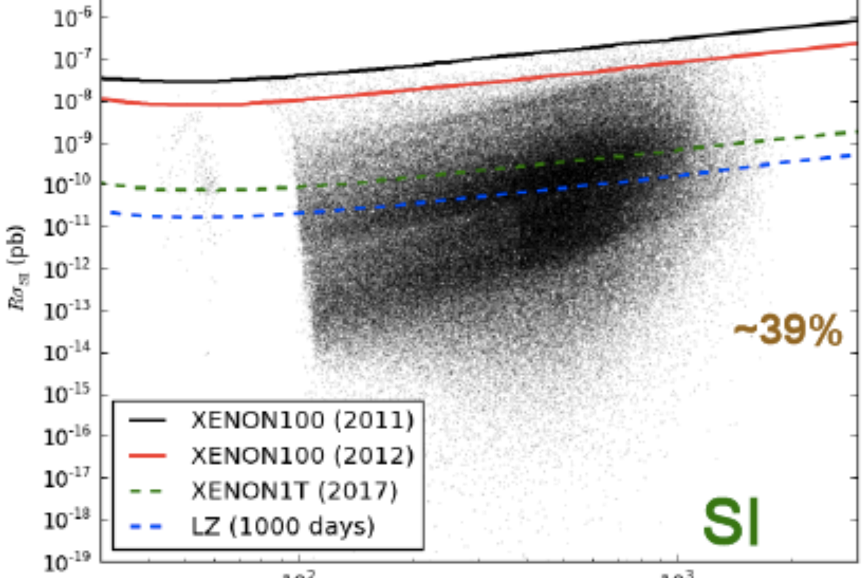


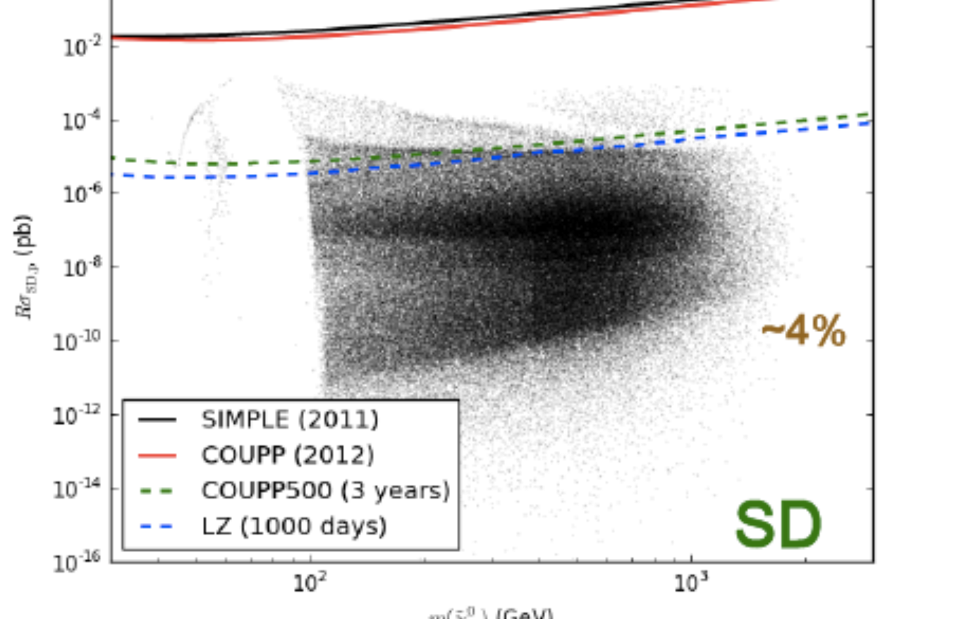
What if 2+ of these experiments observe strong candidate dark matter signals?

Build a directional detector to establish astrophysical origin.



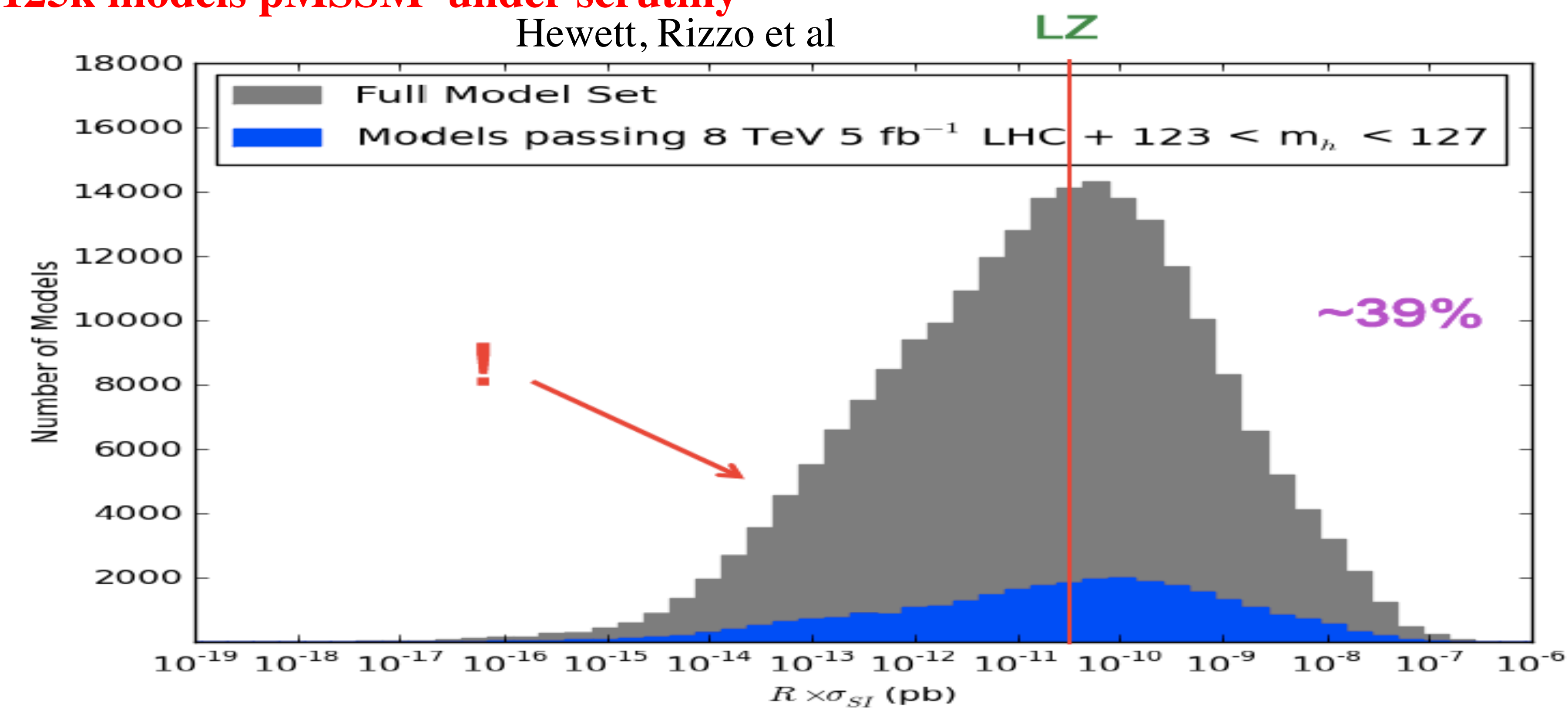
RELEVANCE OF THE DAMA-LIBRA RESULT– IMPORTANCE OF AN INDEPENDENT VERIFICATION (hard to reach the same level of sensitivity)

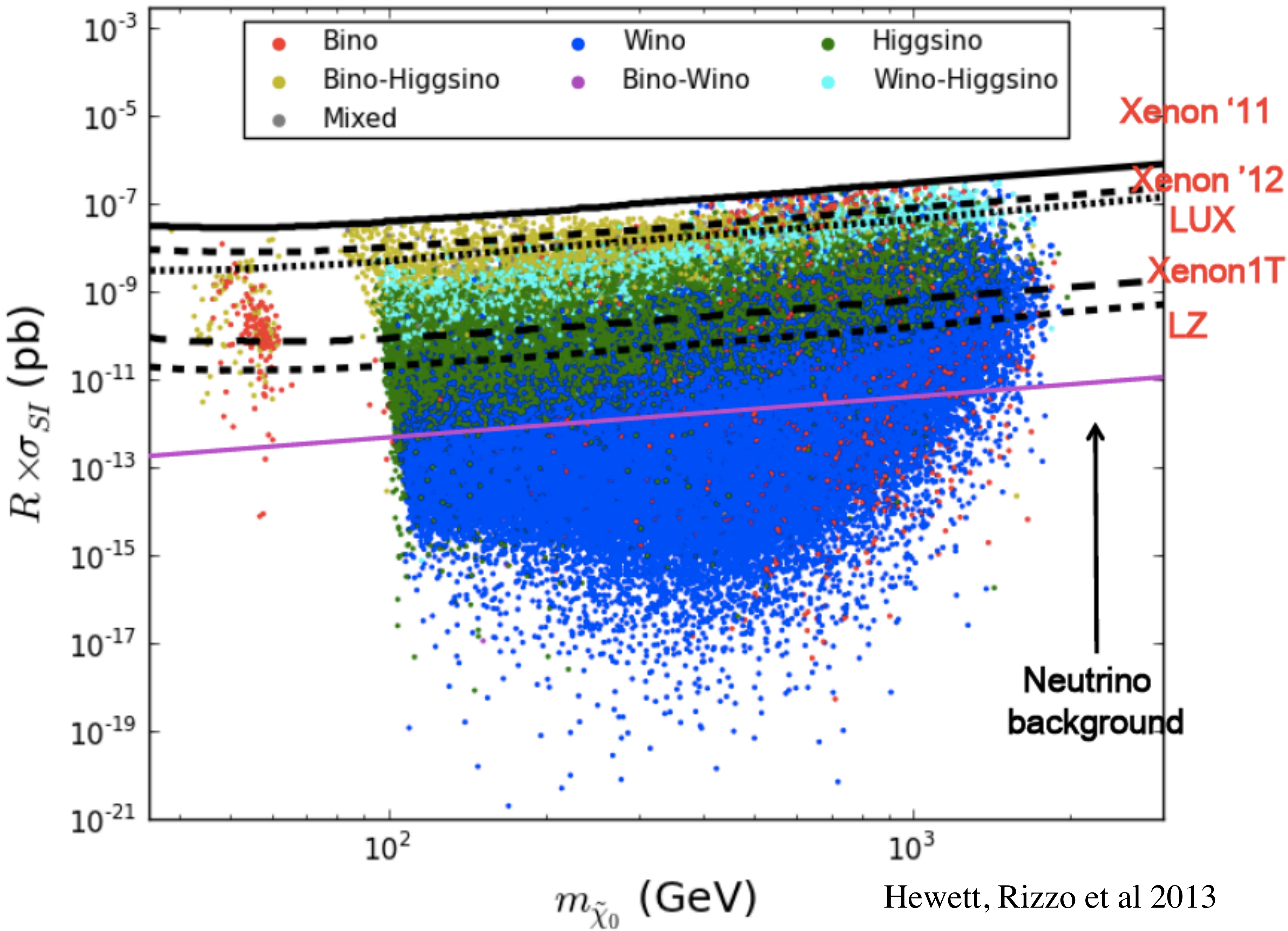




125k models pMSSM under scrutiny

Hewett, Rizzo et al

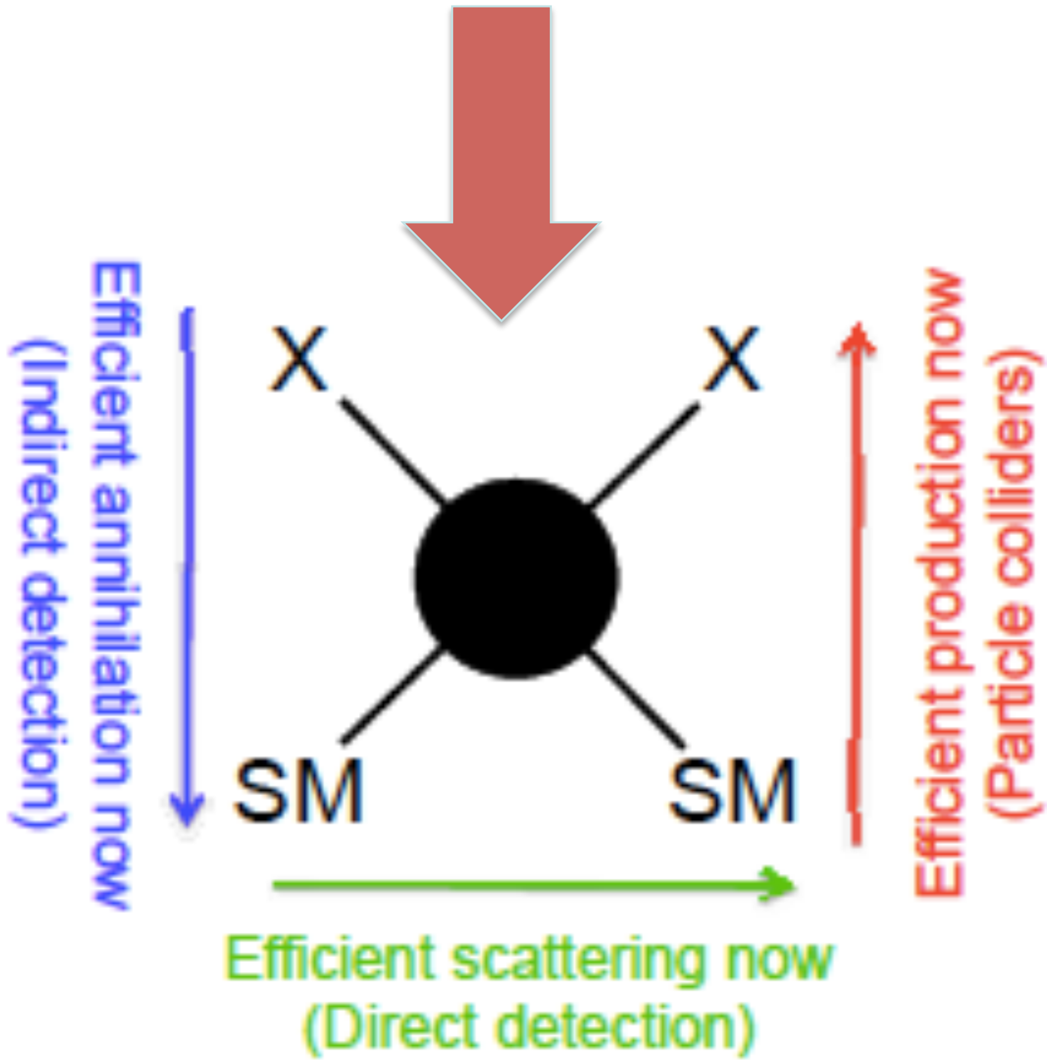


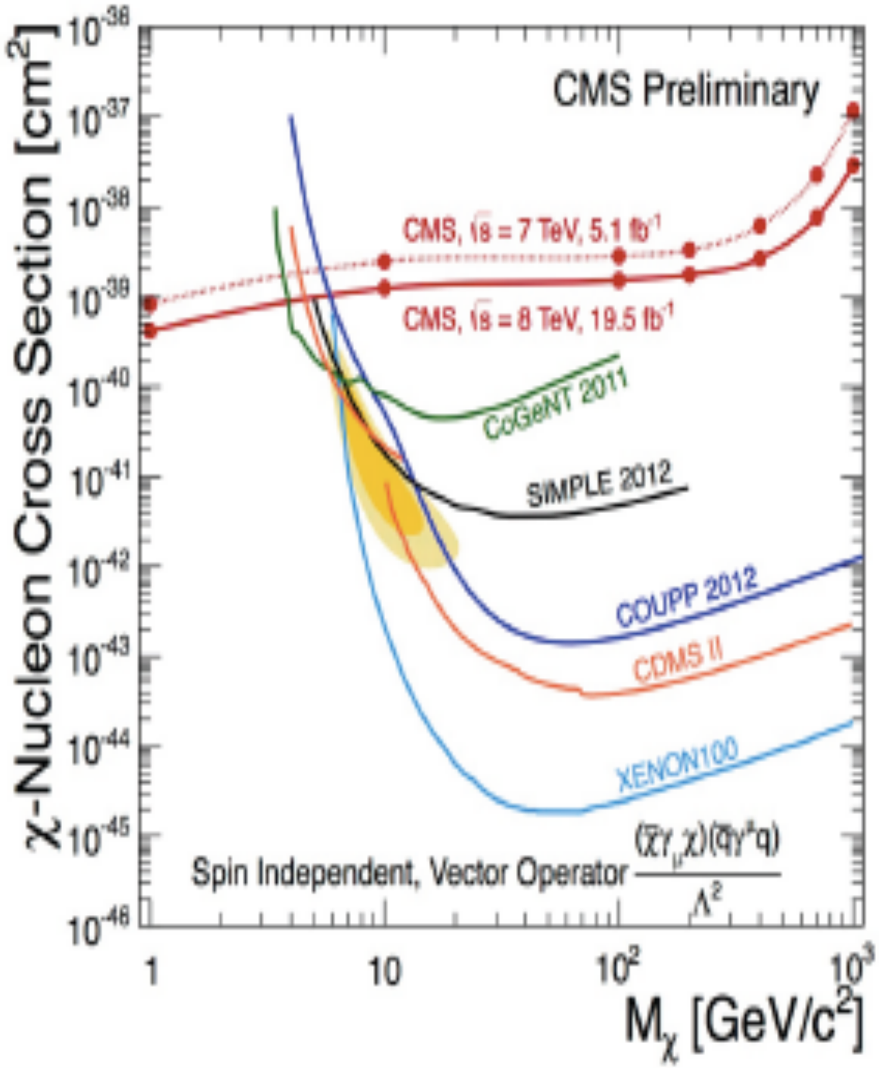


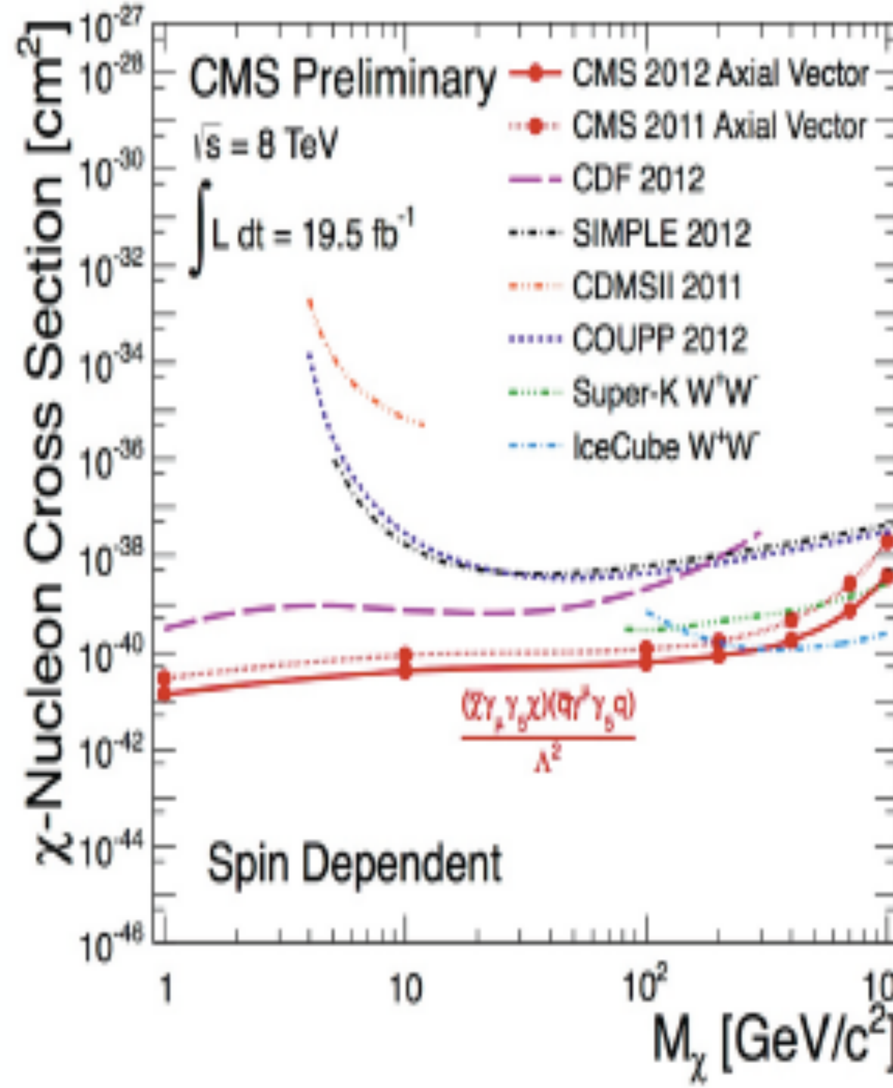
Bino
Wino
Higgsino
Bino-Higgsino
Bino-Wino
Wino-Higgsino
Mixed
Xenon '11
Xenon '12
LUX
Xenon1T
LZ
Neutrino background
$R \times \sigma_{SI}$ (pb)
$m_{\tilde{\chi}_0}$ (GeV)
Hewett, Rizzo et al 2013

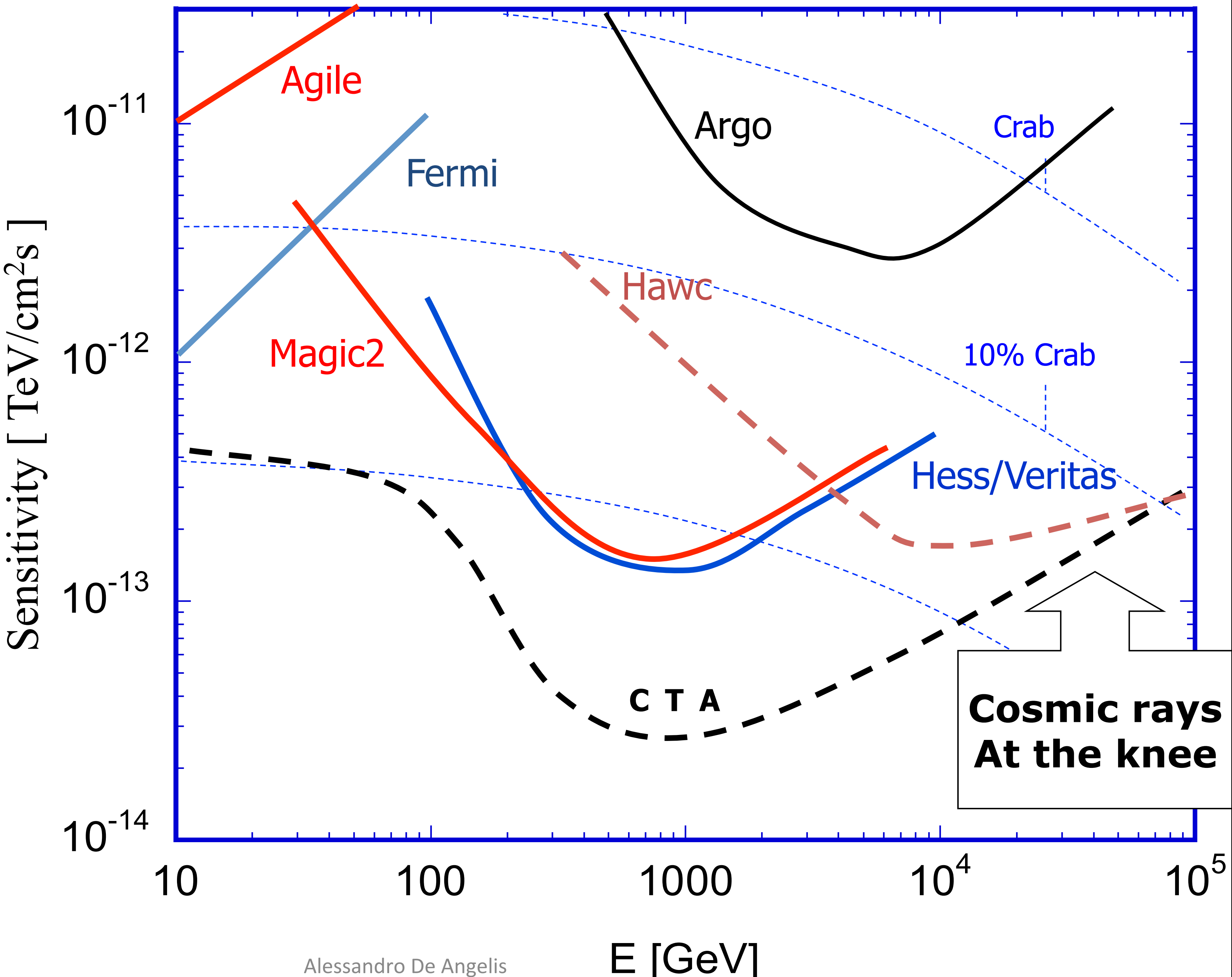
DM COMPLEMENTARITY: efficient annihilation in the early Universe implies today

R. Poettgen (ATLAS), E. Torassa (CMS)



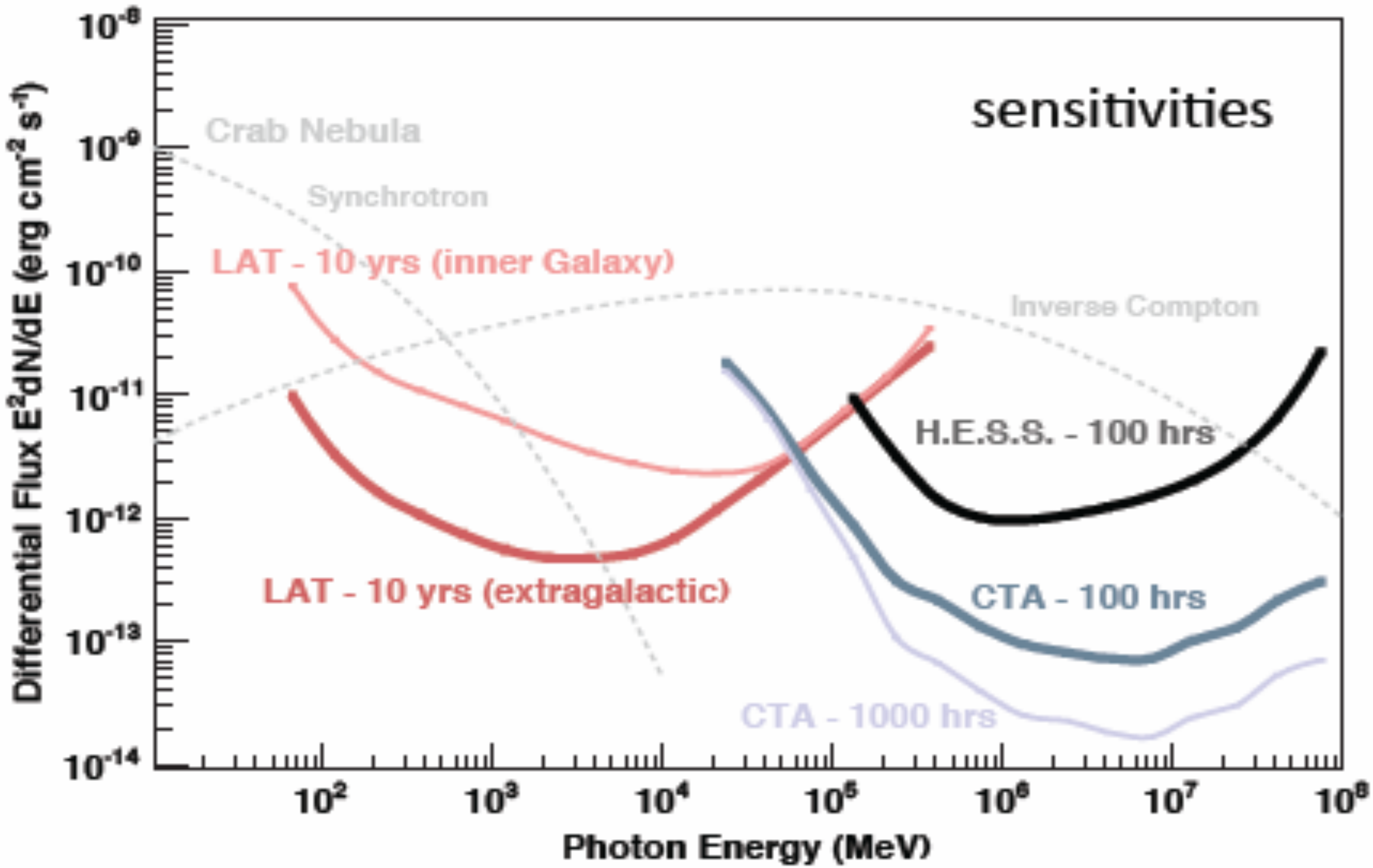

CMS Preliminary
CMS, √s = 7 TeV, 5.1 fb⁻¹
CMS, √s = 8 TeV, 19.5 fb⁻¹
CoGeNT 2011
SIMPLE 2012
COUPP 2012
CDMS II
XENON100
Spin Independent, Vector Operator
χ-Nucleon Cross Section [cm²]
M_χ [GeV/c²]


CMS Preliminary
√s = 8 TeV
∫L dt = 19.5 fb⁻¹
CMS 2012 Axial Vector
CMS 2011 Axial Vector
CDF 2012
SIMPLE 2012
CDMSII 2011
COUPP 2012
Super-K W⁺W⁻
IceCube W⁺W⁻
Spin Dependent
χ-Nucleon Cross Section [cm²]
M_χ [GeV/c²]

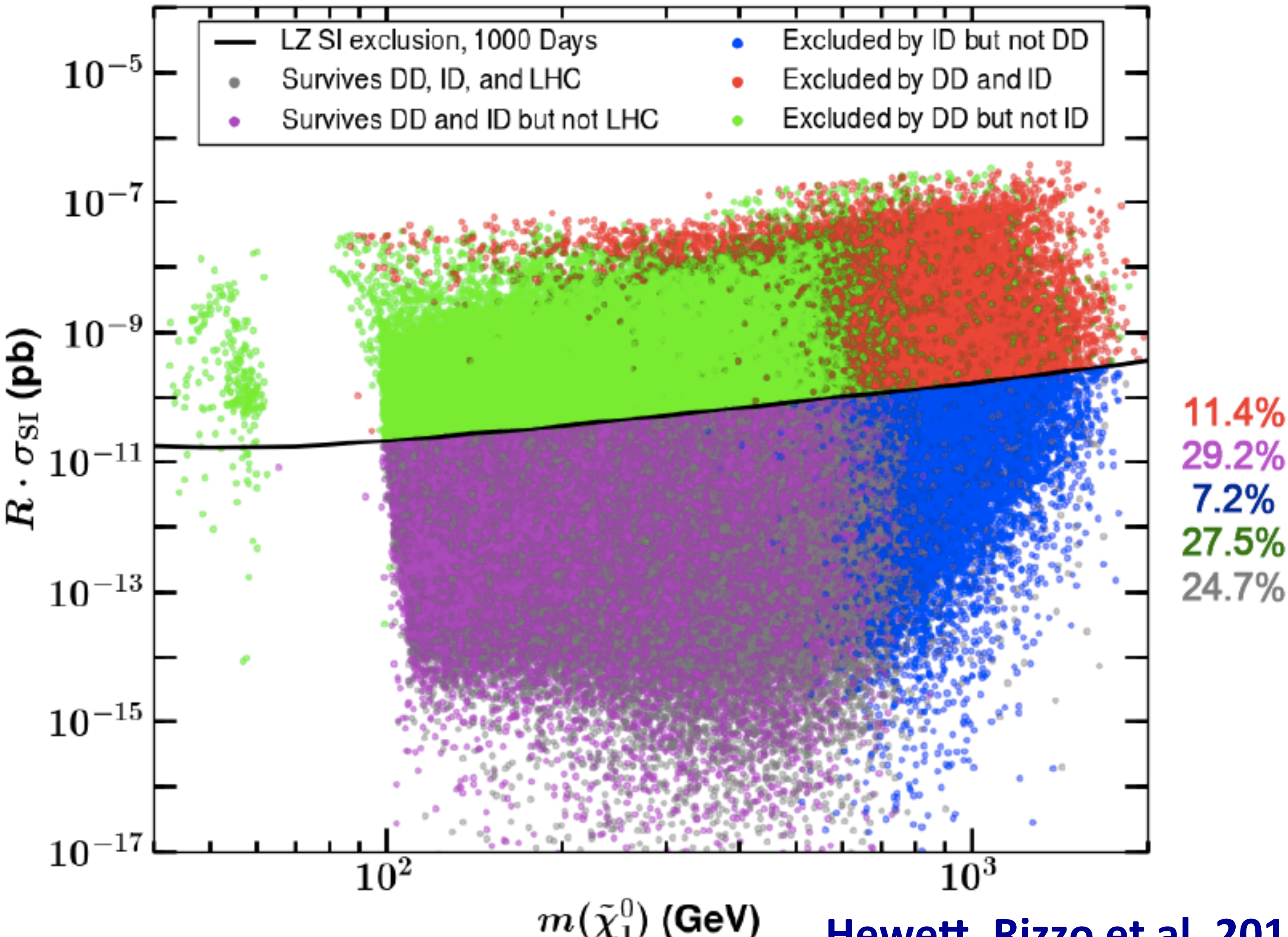


Alessandro De Angelis

GAMMA – ASTRONOMY FROM EARTH AND SPACE



pMSSM models DD = LZ both SI + SD ID = FERMI + CTA



Hewett, Rizzo et al. 2013

AMS Positron Fraction 2014 @ Low Energies

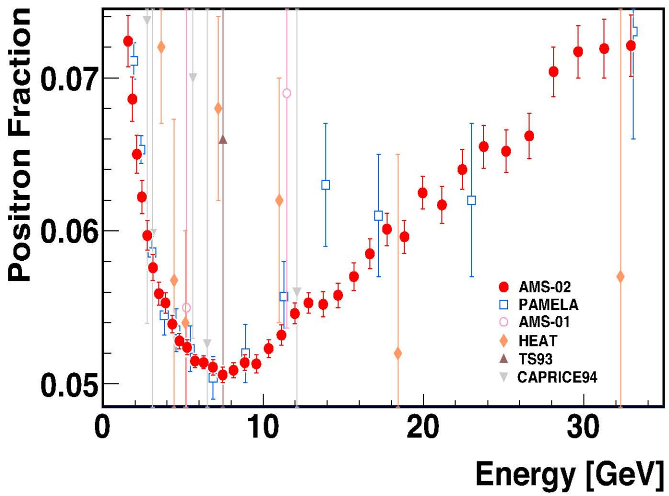
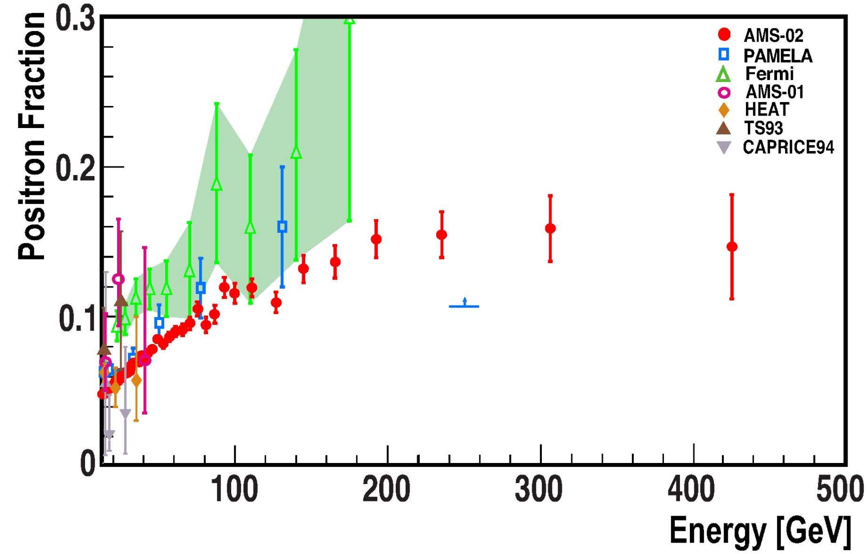


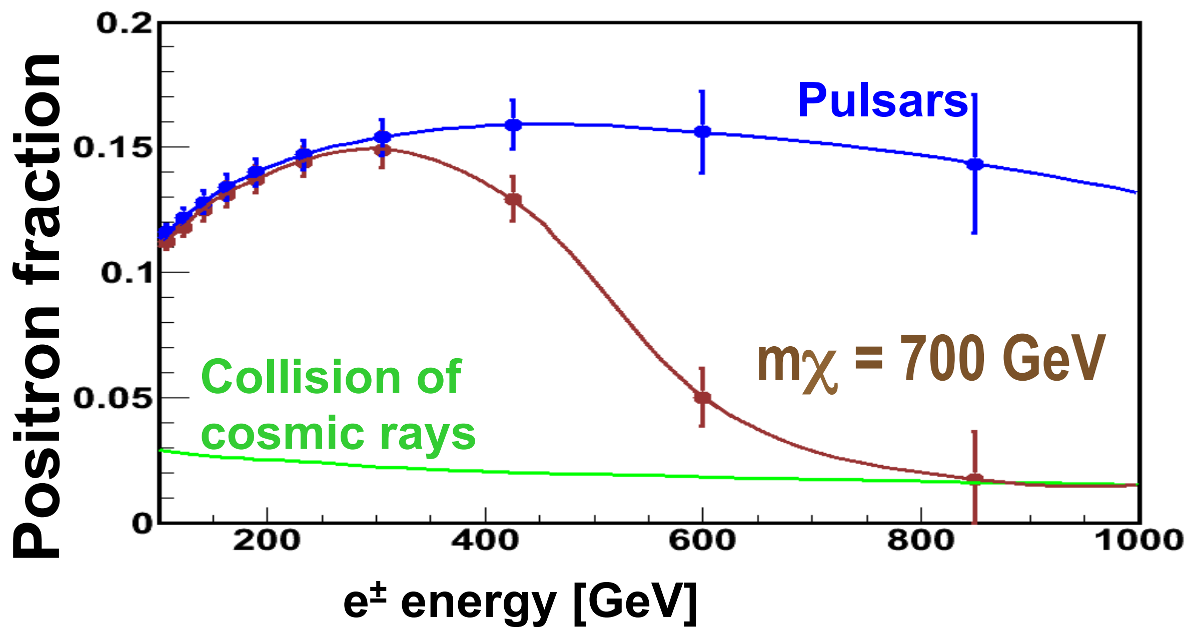
FIG. 2. The positron fraction from 1 to 35 GeV. It shows
provide accurate information on the minimum of the posi

AMS Positron Fraction 2014 @ High Energies



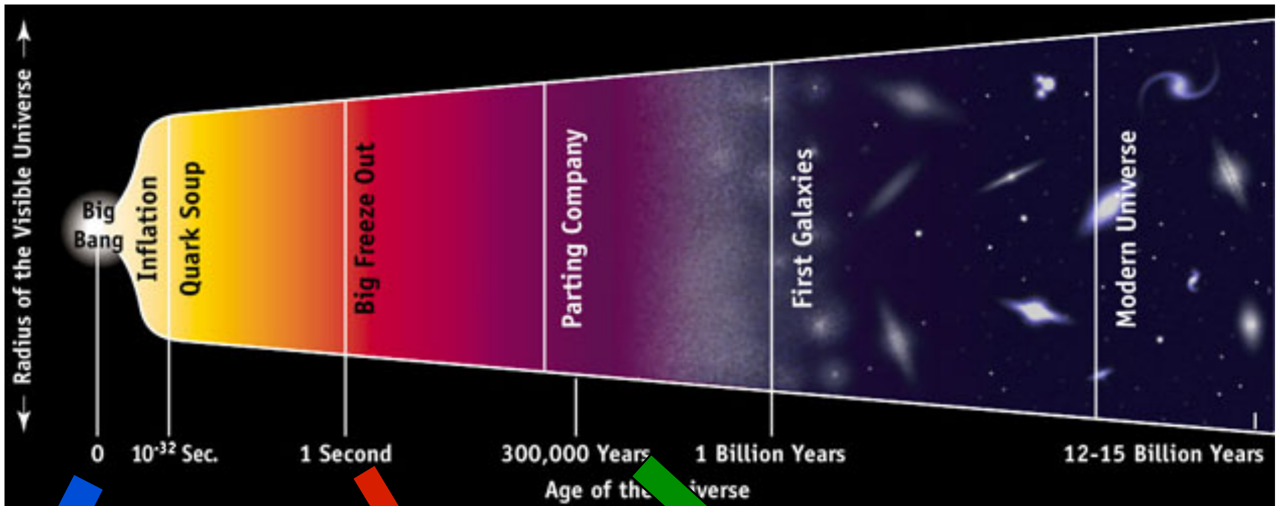
nds the energy range to 500 GeV
m PAMELA [21] (the horizontal

Origin of Positron Fraction: Particle Physics or Astrophysics ?



M. Pohl, IPA2014

Relic Stochastic Background



Relic gravitons

Relic neutrinos

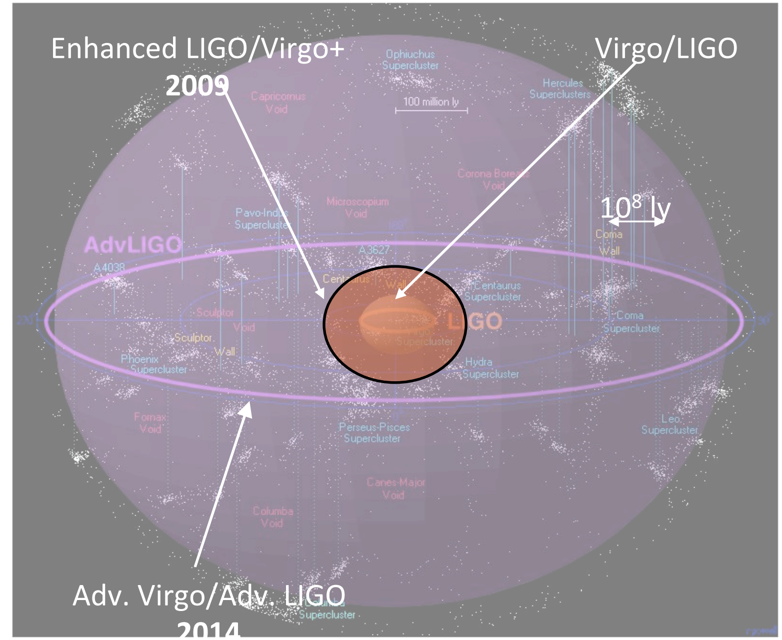
CMBR

- Imprinting of the early expansion of the universe
- Correlation of at least two detectors needed

2nd GENERATION: DISCOVERY AND ASTRONOMY

2nd generation detectors:
Advanced Virgo, Advanced LIGO

GOAL:
sensitivity 10x better →
look 10x further →
Detection rate 1000x larger



B. BARISH at this meeting

Credit: R.Powell, B.Berger

Epoch	Estimated Run Duration	$E_{GW} = 10^{-2} M_\odot c^2$ Burst Range (Mpc) LIGO	$E_{GW} = 10^{-2} M_\odot c^2$ Burst Range (Mpc) Virgo	BNS Range (Mpc) LIGO	BNS Range (Mpc) Virgo	Number of BNS Detections
2015	3 months	40 – 60	–	40 – 80	–	0.0004 – 3
2016–17	6 months	60 – 75	20 – 40	80 – 120	20 – 60	0.006 – 20
2017–18	9 months	75 – 90	40 – 50	120 – 170	60 – 85	0.04 – 100

PHD Syndrome

(Post Higgs Depression)?

(Savas Dimopoulos, GGI, July 2013)

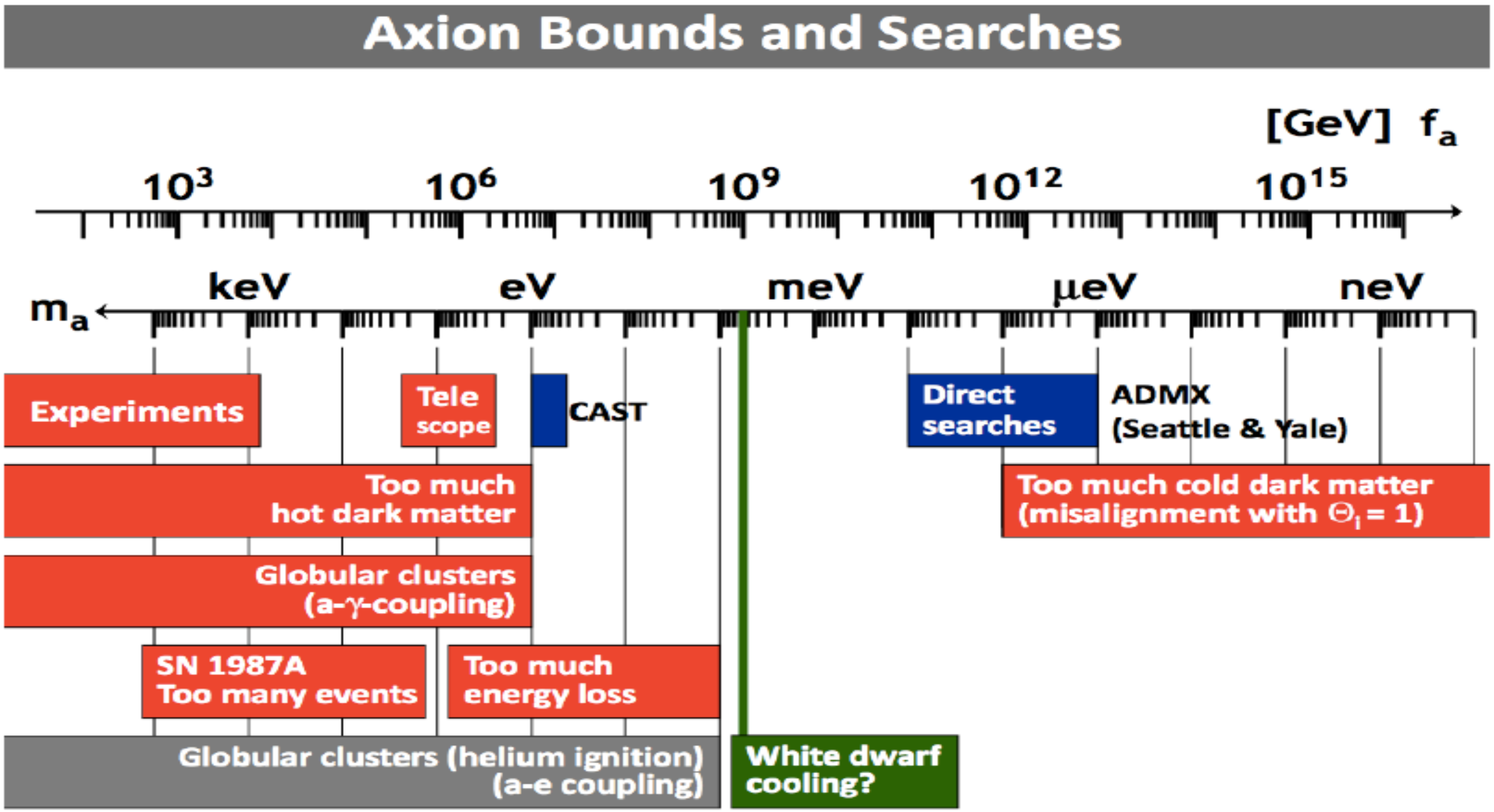
Post-Higgs Depression? No, thanks just the opposite....

- **If the naturalness issue is indeed a relevant issue, the fact that we discovered a light higgs means that there MUST EXIST some mechanism stabilizing its mass and this mechanism NECESSARILY ENTAILS THE PRESENCE OF SOME FORM OF NEW PHYSICS AT THE ELECTROWEAK SCALE**
- Time to get ready (joint exp.-theor. effort) for the new results **in high energy, high intensity, neutrino physics, gravitational waves, cosmic radiation, dark matter and dark energy searches**

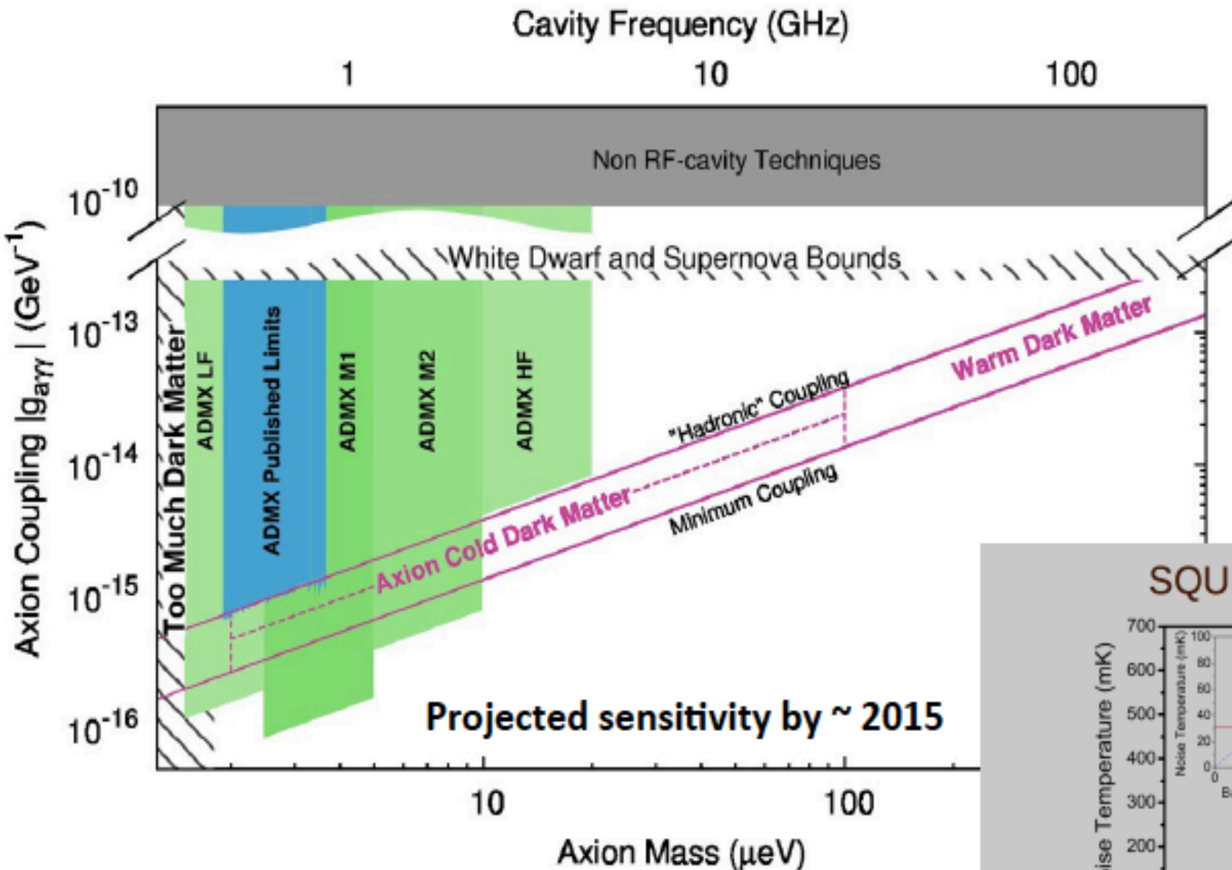

THEORIES
PROOFS
S. Harris

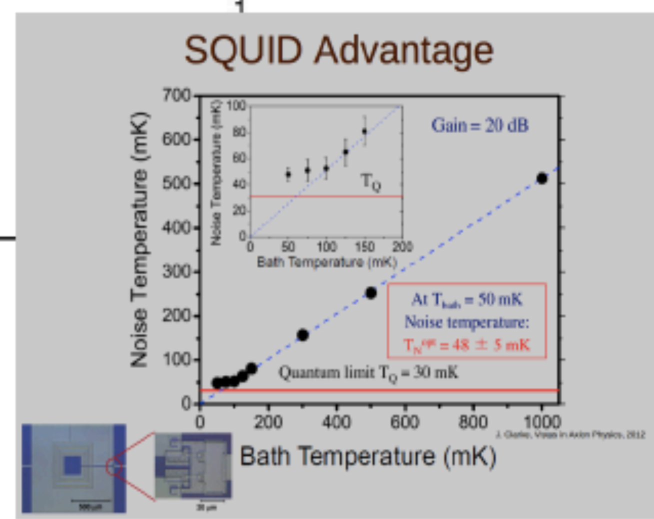
BACK-UP SLIDES

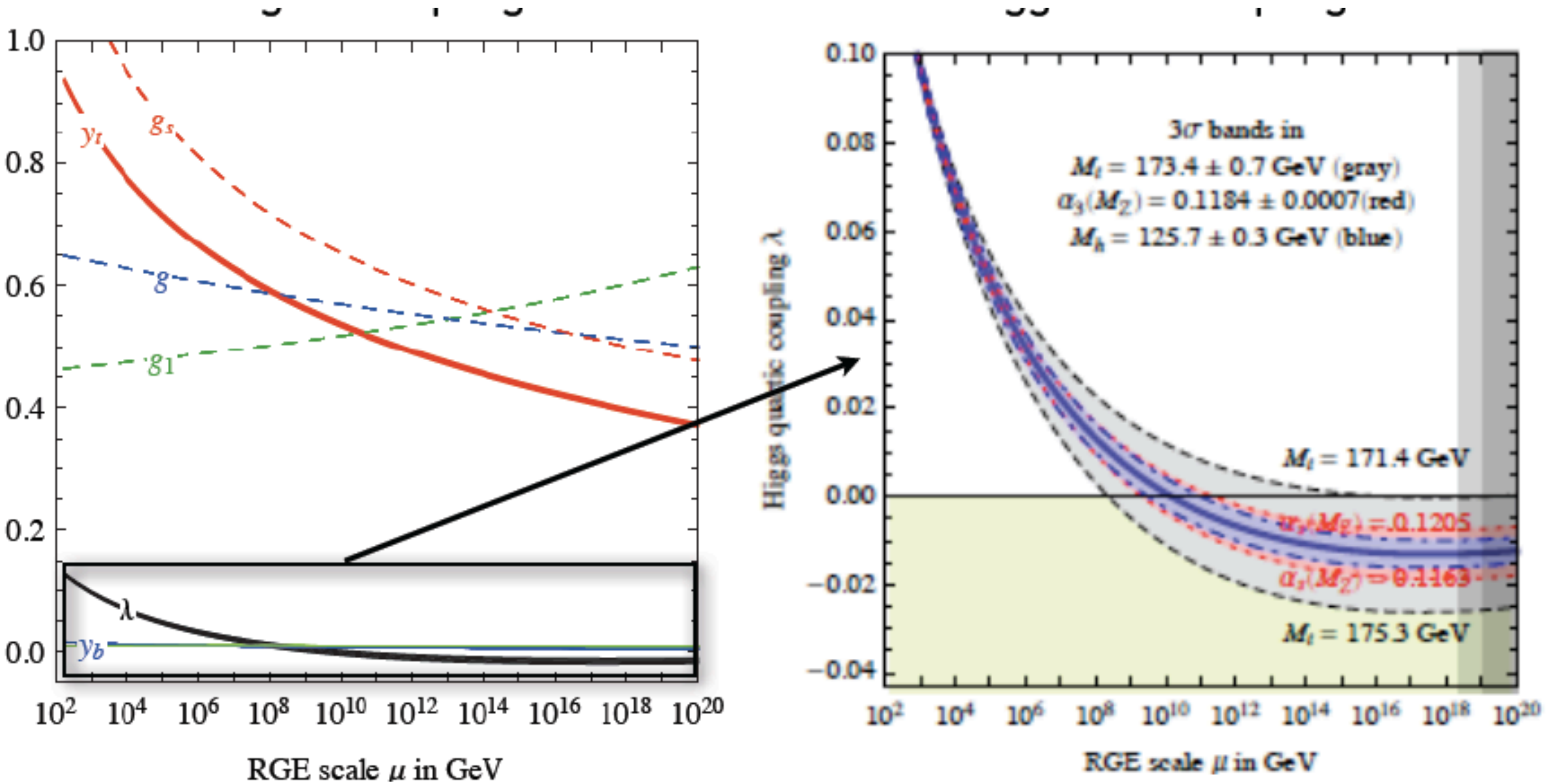
Keep in mind: we don't know at all what DM is made of ! Alternatives to WIMPs – for instance, AXIONS



ADMX achieved and projected sensitivity





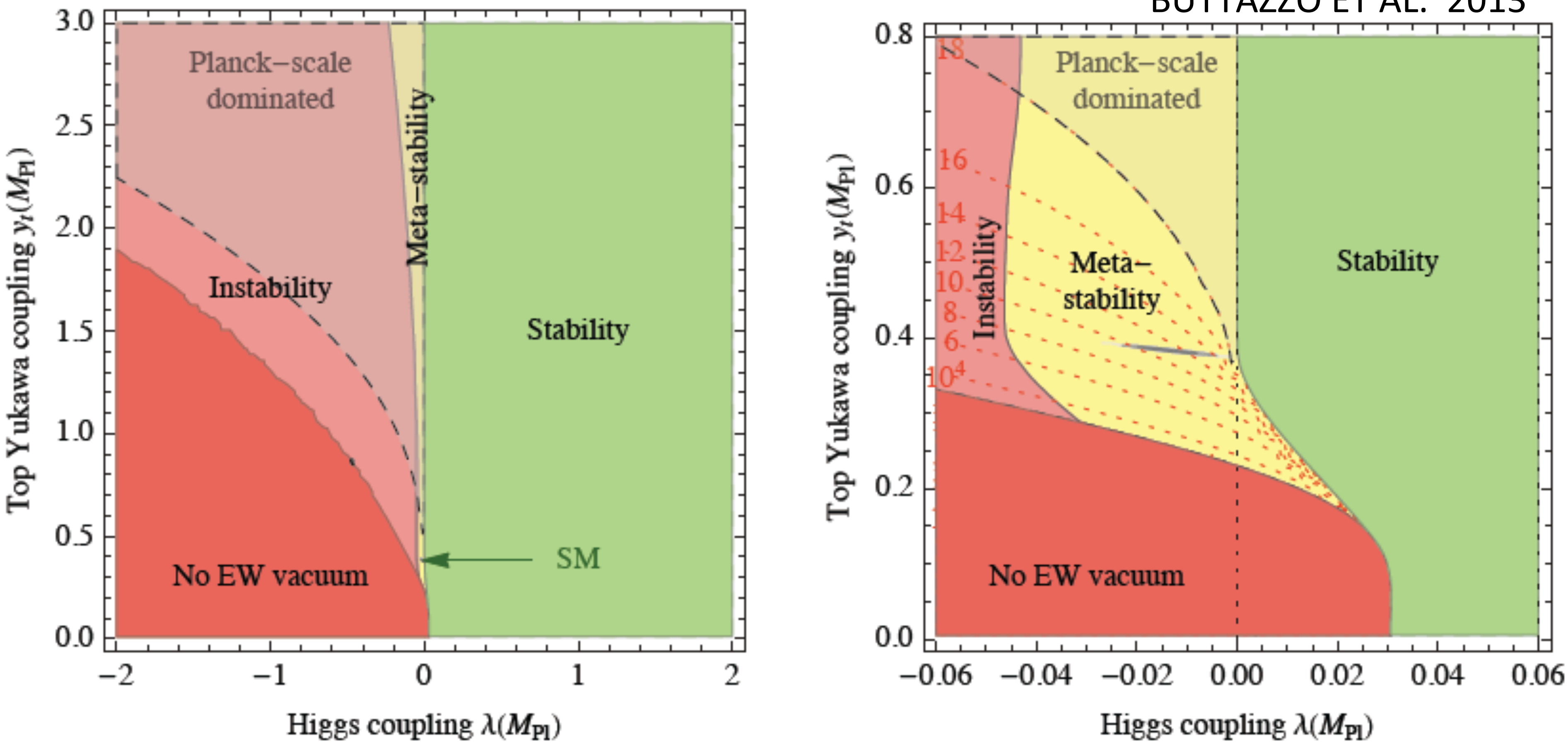


Buttazzo, Degrassi, Giardino, Giudice, Sala, Salvio, Strumia 2013

For previous works: Krive, Linde '76; Krasnikov '78; Maiani, Parisi, Petronzio '78; Cabibbo et al '79; Lindner '86; Altarelli, Isisdori '96; Ellis et al 2009; Shaposhnikov et al '12; Elias-Miro' 'et a ''12; …..
Degrassi, Di Vita, Elias-Miro, Espinosa, Giudice, Isidori, Strumia '12

IF SM VALID UP TO M_{PLANCK} → M_H formidable telescope to sneak into unexplorable energies…

BUTTAZZO ET AL. 2013



The Universe looks very close to **CRITICALITY**

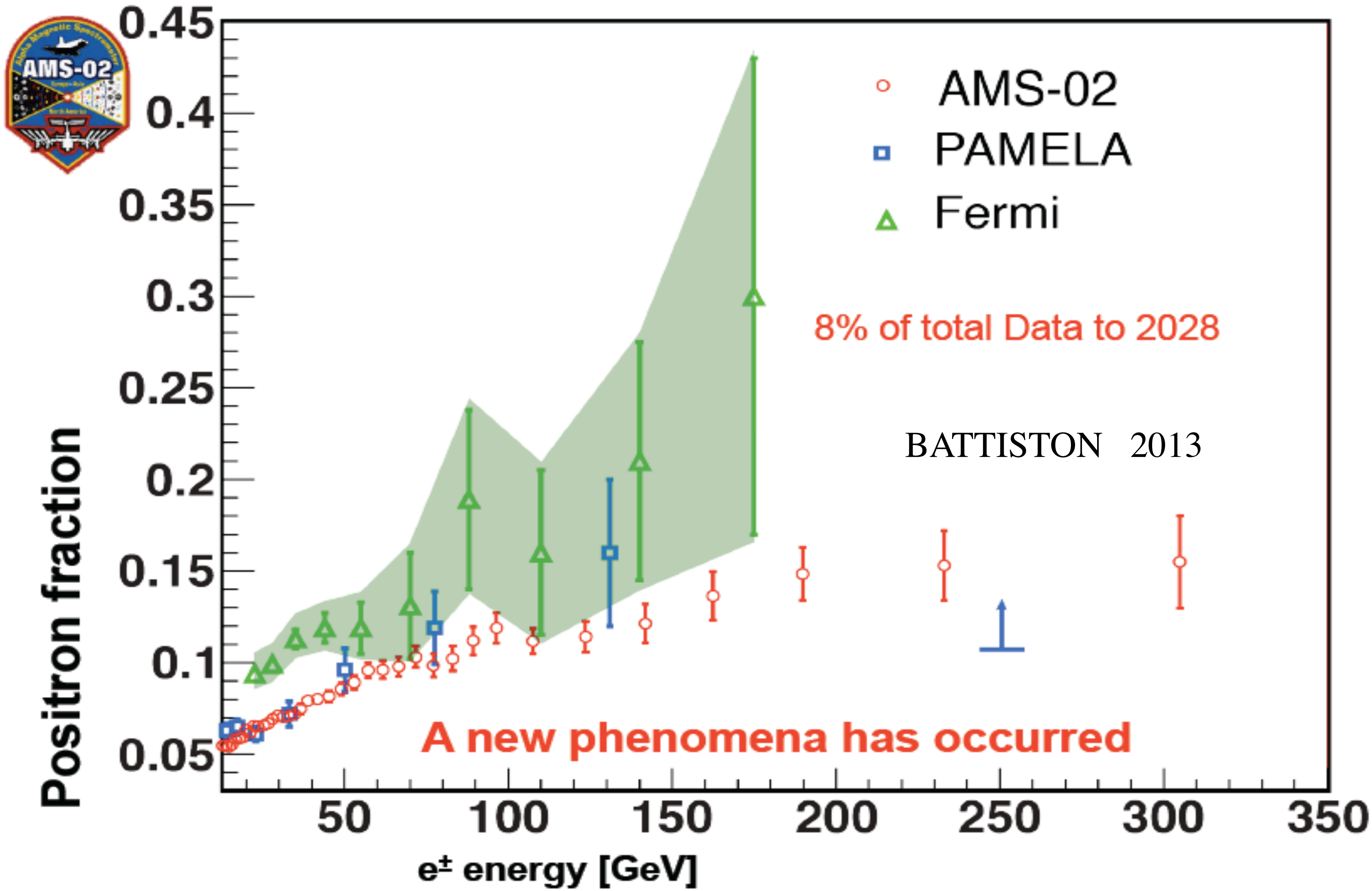
ON THE IMPORTANCE OF PRECISELY MEASURING HIGGS and TOP MASSES

DEGRASSI ET AL

Type of error	Estimate of the error	Impact on M_h
M_t	experimental uncertainty in M_t	± 1.4 GeV
α_s	experimental uncertainty in α_s	± 0.5 GeV
Experiment	Total combined in quadrature	± 1.5 GeV
λ	scale variation in λ	± 0.7 GeV
y_t	$\mathcal{O}(\Lambda_{QCD})$ correction to M_t	± 0.6 GeV
y_t	QCD threshold at 4 loops	± 0.3 GeV
RGE	EW at 3 loops + QCD at 4 loops	± 0.2 GeV
Theory	Total combined in quadrature	± 1.0 GeV

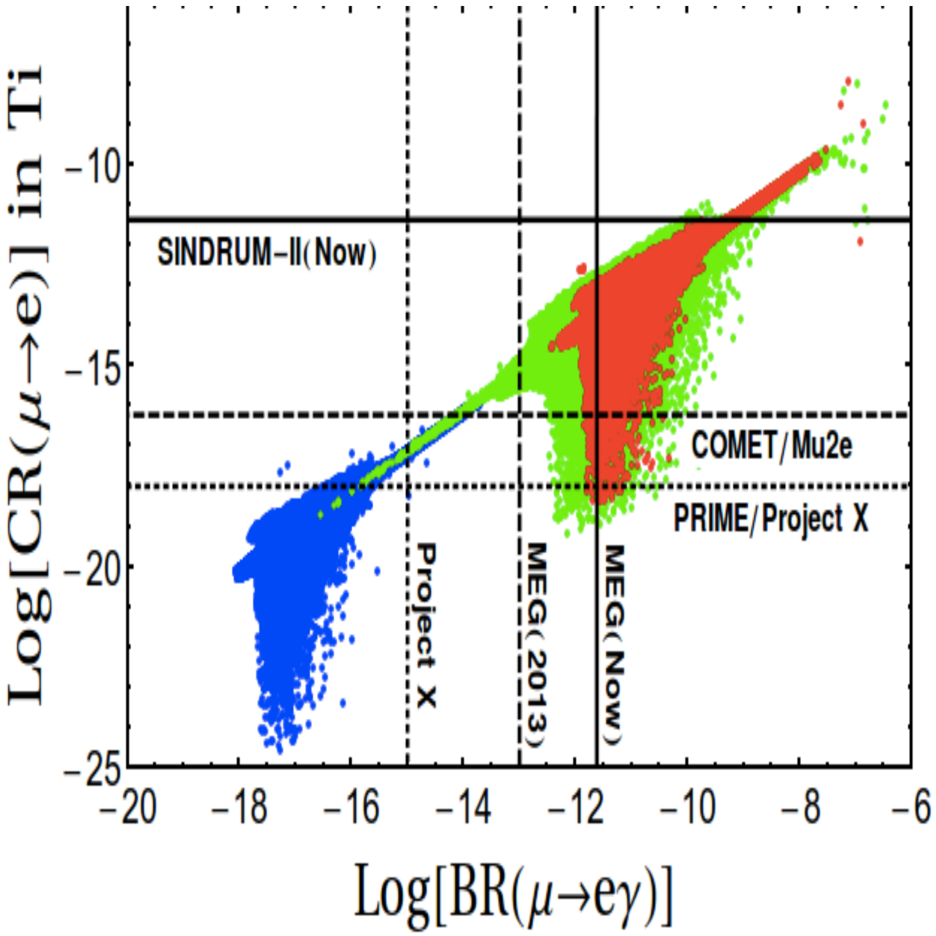
INTRINSIC DIFFICULTY TO “DEFINE” WHAT THE TOP MASS IS AT A **HADRON COLLIDER** WITH UNCERTAINTY ≤ 1 GeV

INDIRECT SEARCHES FOR DM

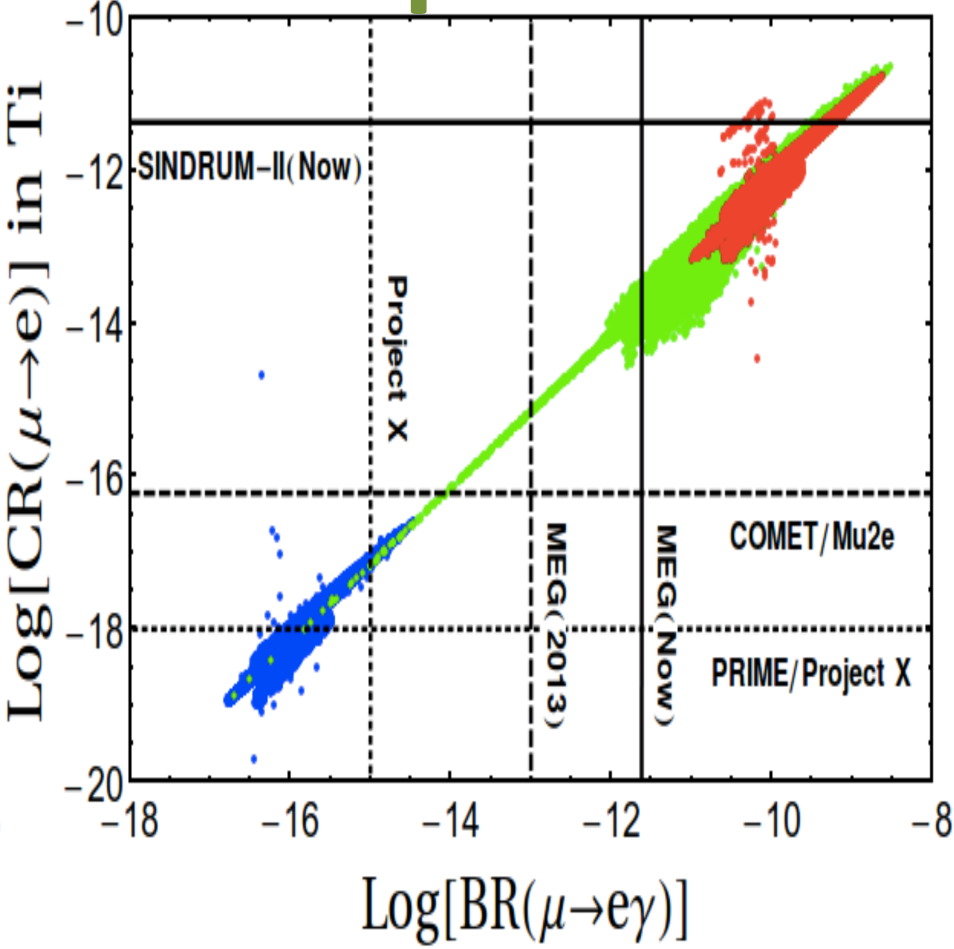


μ – e conversion vs μ → eγ

Tanβ = 10

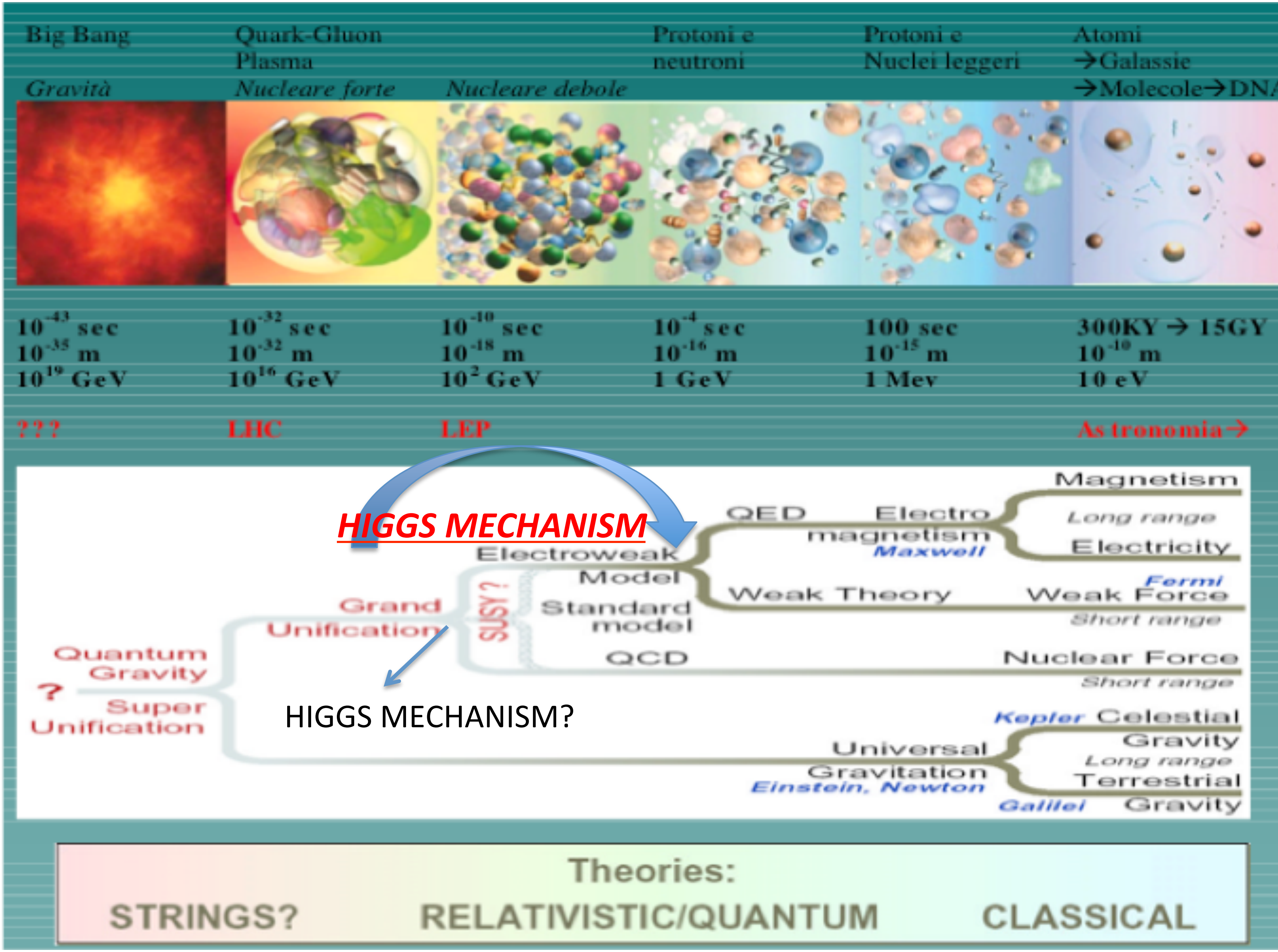


Tanβ = 40



Is the DM a portal to new physics beyond the SM? (I)

- DM: most of the gravitationally clusterized form of energy of the Universe that we call MATTER is of non-baryonic nature, i.e. **non-baryonic DM exists**, and **it is by itself new physics**, i.e. it is made of particle(s) which are not present in the SM particle spectrum
- **Is (are) the mass(es) of the DM particle(s) at the electroweak scale**, i.e. of O(1TeV), or is the DM scale not correlated at all with the elw. scale?



Big Bang
Quark-Gluon Plasma
Protoni e neutroni
Protoni e Nuclei leggeri
Atomi →Galassie →Molecole→DNA
Gravità
Nucleare forte
Nucleare debole
10^-43 sec 10^-35 m 10^19 GeV
10^-32 sec 10^-32 m 10^16 GeV
10^-10 sec 10^-18 m 10^2 GeV
10^-4 sec 10^-16 m 1 GeV
100 sec 10^-15 m 1 Mev
300KY → 15GY 10^-10 m 10 eV
? ? ?
LHC
LEP
As tronomia→
HIGGS MECHANISM
Magnetism
QED
Electro magnetism
Maxwell
Long range
Electricity
Electroweak Model
Fermi
Weak Theory
Weak Force
Short range
Grand Unification
SUSY ?
Standard model
Quantum Gravity
?
Super Unification
QCD
Nuclear Force
Short range
HIGGS MECHANISM?
Kepler
Celestial Gravity
Universal Gravitation
Einstein, Newton
Long range
Terrestrial Gravity
Galilei
Theories:
STRINGS?
RELATIVISTIC/QUANTUM
CLASSICAL

Higgs (and top?) physics: the importance of being precise

Higgs mass measurement: the need for high precision

Measuring the mass of the discovered signal with high precision is of interest in its own right

But a high-precision measurement has also direct implications for probing Higgs physics

M_H: crucial input parameter for Higgs physics

BR($H \to ZZ^*$), BR($H \to WW^*$): highly sensitive to precise numerical value of M_H

A change in M_H of 0.2 GeV shifts BR($H \to ZZ^*$) by 2.5%!

⇒ Need high-precision determination of M_H to exploit the sensitivity of BR($H \to ZZ^*$), ... to test BSM physics

WHAT NEXT

In view of the complex landscape we have to confront, INFN has recently started a process to identify the most important research themes that we should focus on amongst those that in this moment do not receive enough attention (people, funding). FERRONI

HIGH ENERGY, HIGH-INTENSITY, ASTROPARTICLE PHYSICS COMPLEMENTARY ATTACK TO THE NEW PHYSICS FORTRESS

The next step

It is clear that at least **1 ton isotope** is required to explore the inverted hierarchy region

- Impact of enrichement cost

Isotope	Abundance	Price/kg [k$]	Price/(10 t) [M$]
^{76}Ge	7.61	~ 80	800 (640)$^{*)}$
^{82}Se	8.73	~ 120	1200 (1000)$^{*)}$
^{100}Mo	9.63	~ 80	800 (640)$^{*)}$
^{116}Cd	7.49	~ 180	1800 (1440)$^{*)}$
^{130}Te	34.08	~ 20	200 (160)$^{*)}$
^{136}Xe	8.87	$\sim$ 5-10	50-100 (40-80)$^{*)}$
^{150}Nd (?)	5.6	> 200	> 2000

Barabash, 2013

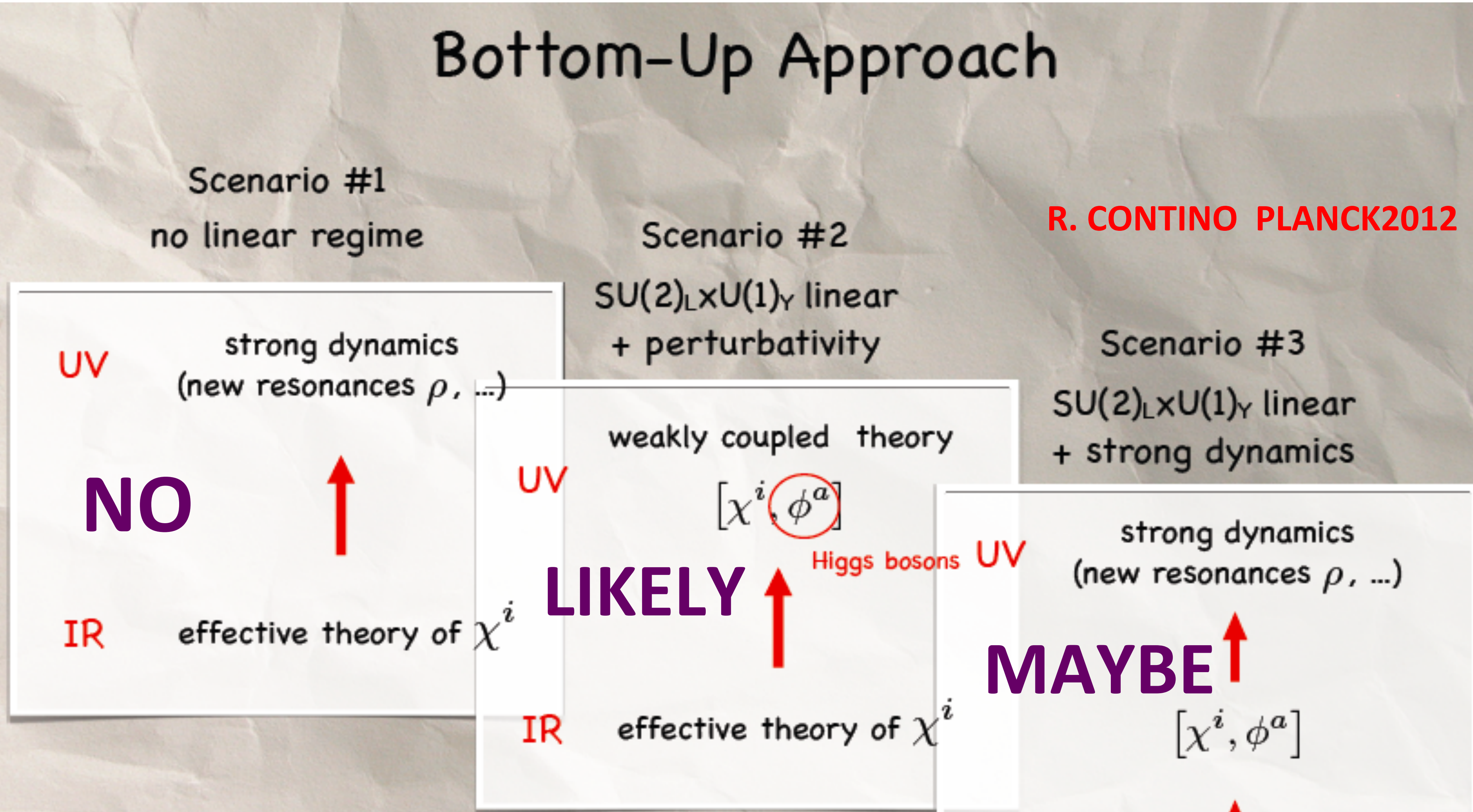
- How many technological approaches and which ones?

- How many isotopes and which ones?

- **Which infrastructures?**
  Which chance do we have to get them in Europe?

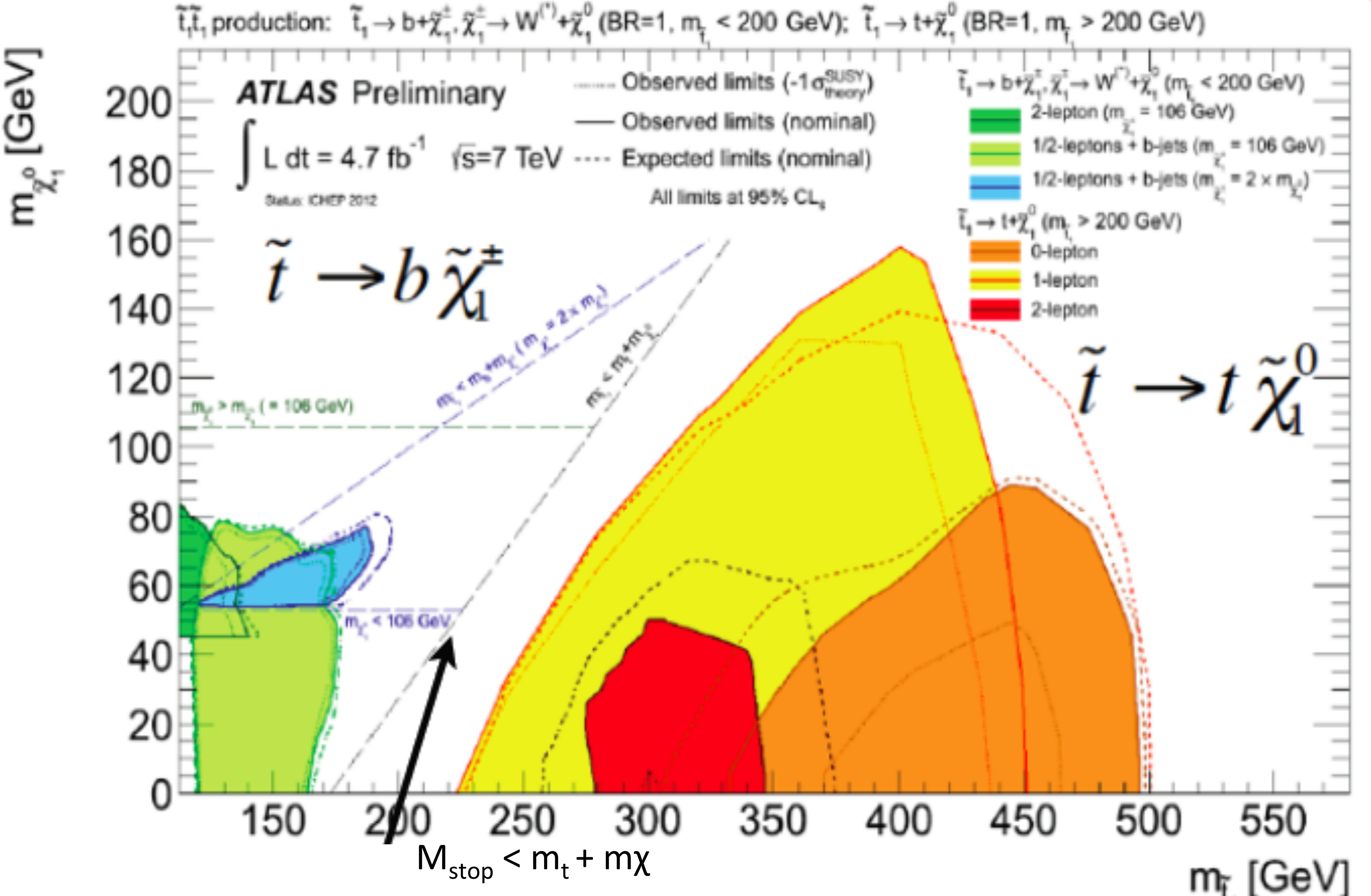
A. Giuliani SAC APPEC 2013

EWSB: WITH OR WITHOUT A HIGGS BOSON



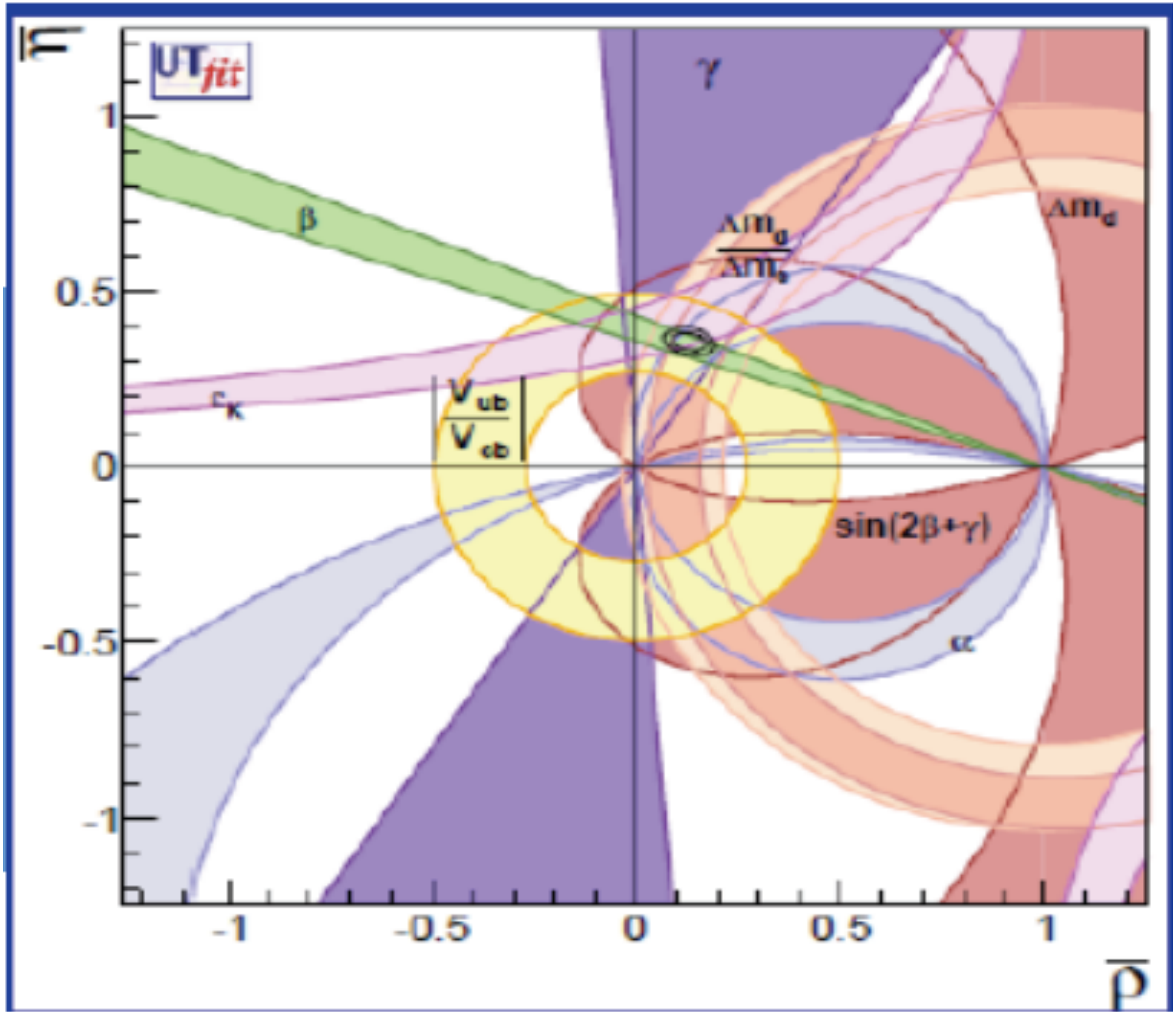
CAN LHC TELL US WHAT NATURE HAS CHOSEN TO BREAK THE ELW SYMMETRY?

Hunting for a light s-top



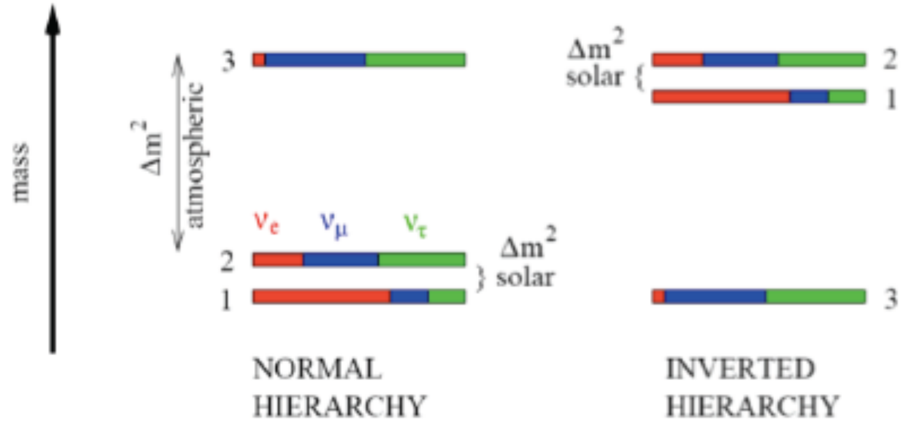
$M_{stop} < m_t + m\chi$

the (almost complete) CKM triumph



Neutrino oscillations & sterile neutrinos

- **Atmospheric neutrinos(θ_{23})**
  - SuperK, HyperK/UNO, INO, TITAND,...
- **Solar neutrinos(θ_{12}):**
  - GALLEX/SAGE, SuperK, SNO, Borexino, XMASS, ...
- **Reactor neutrinos(θ_{12},θ_{13} → mass hierarchy):**
  – KamLAND, Daya Bay→ JUNO, Double CHOOZ, Reno,...
- **Accelerator neutrinos(θ_{23},θ_{13} → mass hierarchy, δ, ...):**
  – MINOS, OPERA, MiniBooNe, T2K, NOVA, ICARUS...



CPV

+ A number of anomalies:

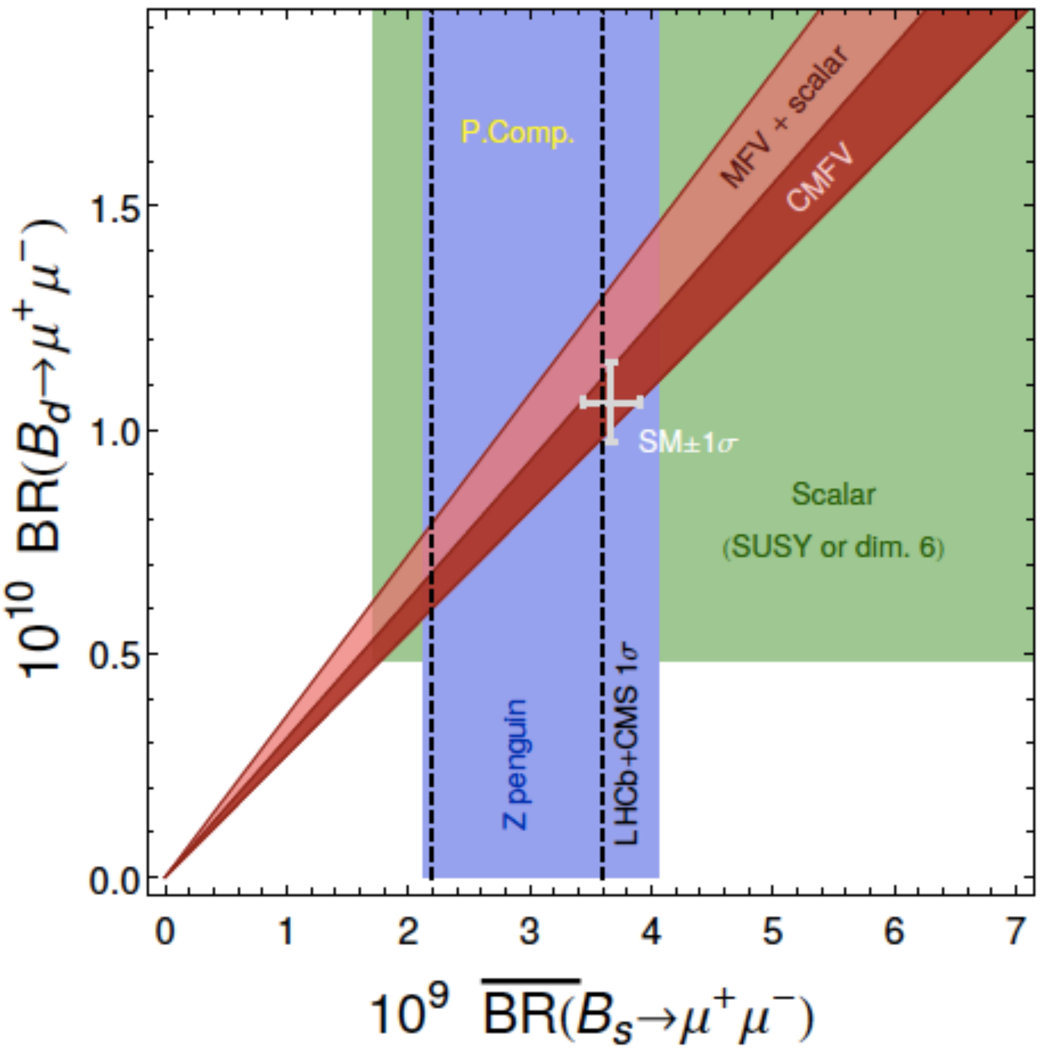
LSND ?

Reactor neutrino flux ?

Sterile neutrinos ? MiniBoone

D. STRAUB

Model-independently: $B_s \to \mu\mu$ vs. $B_d \to \mu\mu$



- $B_s \to \mu\mu$ vs. $B_d \to \mu\mu$: powerful test of Minimal Flavour Violation
- Beyond MFV: scalar operators (e.g. SUSY) unconstrained by other processes
- Z penguins beyond MFV: size constrained by $b \to s\ell^+\ell^-$ processes (notably $B \to K^*\mu^+\mu^-$)