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AGB star nucleosynthesis: when new data from nuclear physics help to solve puzzles

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Low mass stars contribute to the chemical evolution of the Galaxy as well as more massive supernova progenitors. Indeed the limited amount of processed matter released into the interstellar medium by small objects is compensated by the large number of them.

At the late stages of their evolution, stars with mass smaller than $3M_{\odot}$ undergo the Asymptotic Giant Branch (AGB) phase, which has been found to be a unique site for synthesis of some nuclei heavier than Fe through slow neutron capture reactions. AGB nucleosynthesis is also characterized by H-burning coupled with mixing phenomena, which have been proved to account for anomalies in the Li abundance and C, N and O isotopic ratio observed in stellar spectra and meteorite grains.

Nowadays the improvements in stellar spectroscopy and the growing number of geochemical analysis of meteorite grains offer new challenges and stronger constraints for the model of stellar nucleosynthesis and a high precision of the nuclear physics data employed in calculations is required.

We present how updated measurements of nuclear cross sections contribute to solve puzzles of oxygen isotopic mix in presolar grains of AGB origins. While despite accurate measurement of the reaction rate the high abundances of ^{26}Al found in the same grains require a specific mixing model to be reproduced, similar to the one needed to trigger the formation of an extended neutron source in low mass AGB stars.

Finally, to show the crucial role of a proper estimation of electron density at the nucleus in the stellar plasma we present an evaluation of the complex problem of Li abundances in AGB stars in the light of an ad-hoc estimate of the ^7Be time-of-life.

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