Contribution of non-AGN sources to diffuse high energy astrophysical neutrino

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Cosmic-ray and Dark Matter

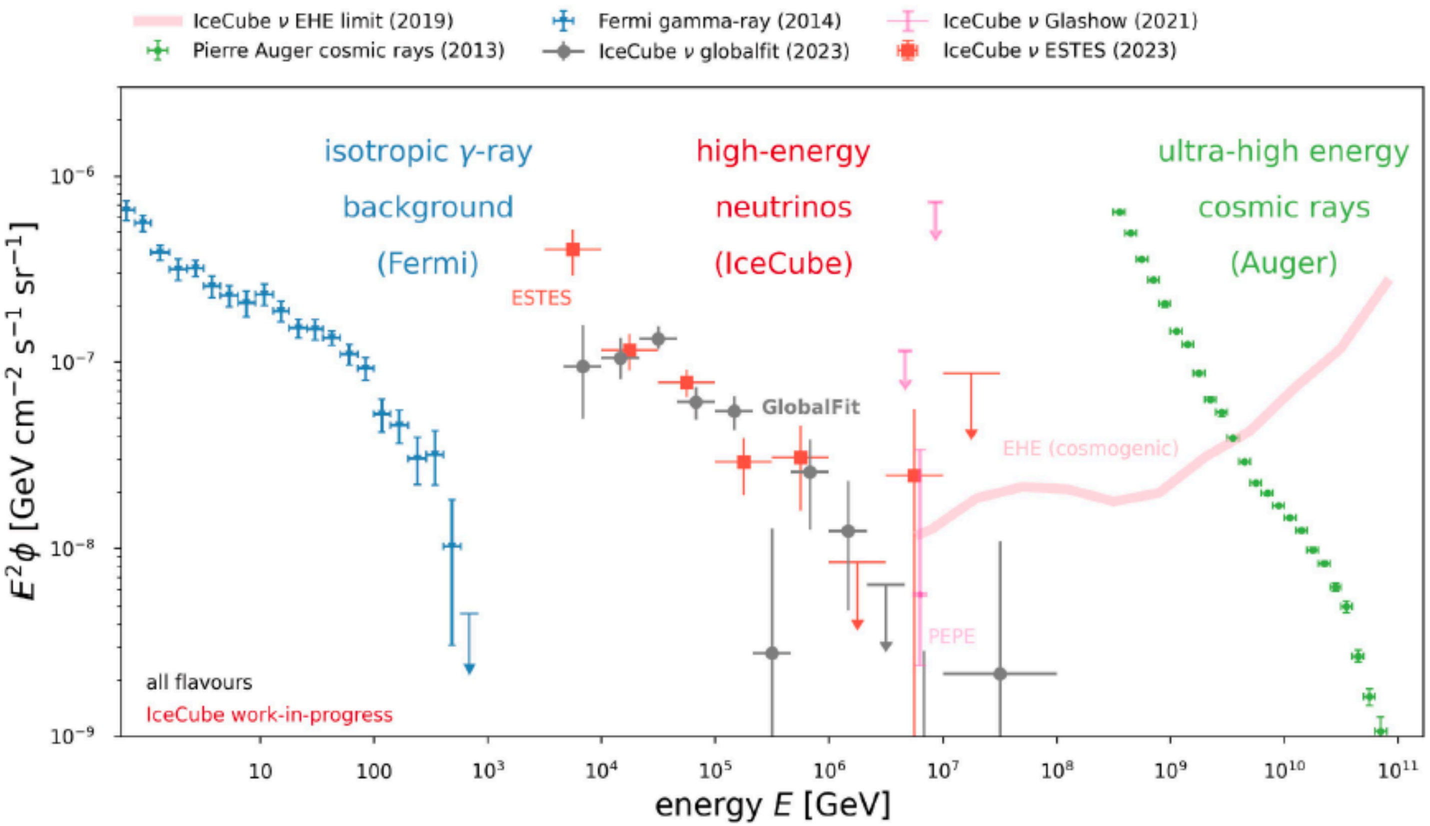
10-11/11/2024

TAsP Meeting And brainstorming

Multimessenger connections

From J.A Aguilar RICAP24

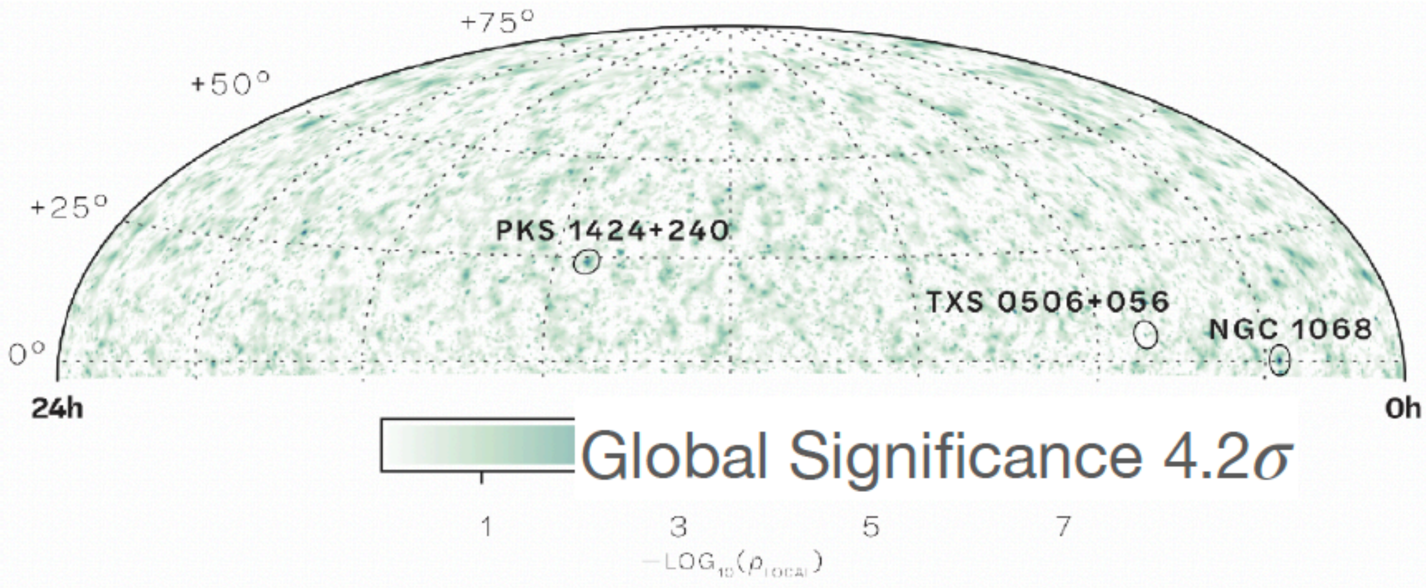
Gamma-rays, neutrinos, and cosmic rays connection

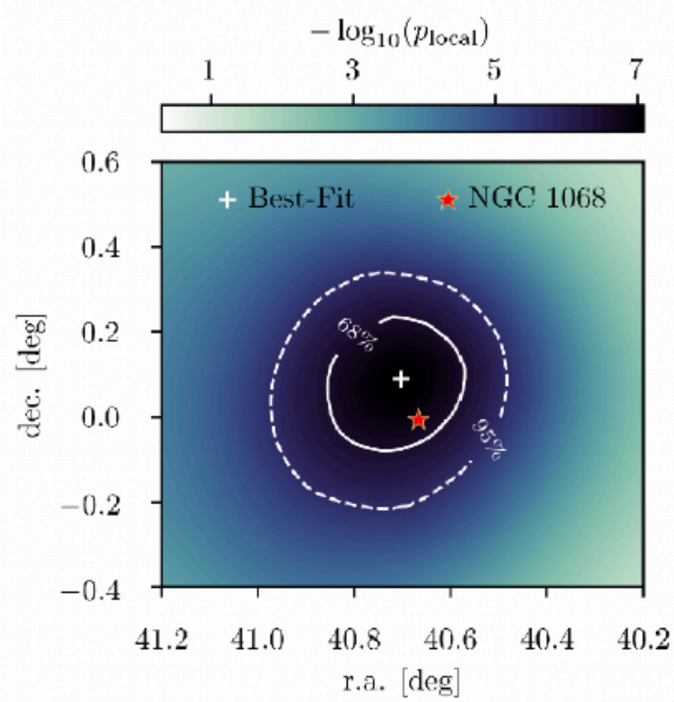


- Diffuse background with 3 different messengers:
  - **Similar energy densities…**
  - …but also **evidence of different origin**
- Interesting **interfaces between messengers**

The observed AGNs with ν

The Neutrino Sky





- Strongest neutrino emission (best-fit):
  - Located at R.A. 40.69° and Dec. 0.09°.
  - $\hat{n} = 81$
  - $\hat{\gamma} = 3.2$
- It also appears in the the list of 110 pre-define sources

RESEARCH

RESEARCH ARTICLE

NEUTRINO ASTROPHYSICS

Evidence for neutrino emission from the nearby active galaxy NGC 1068

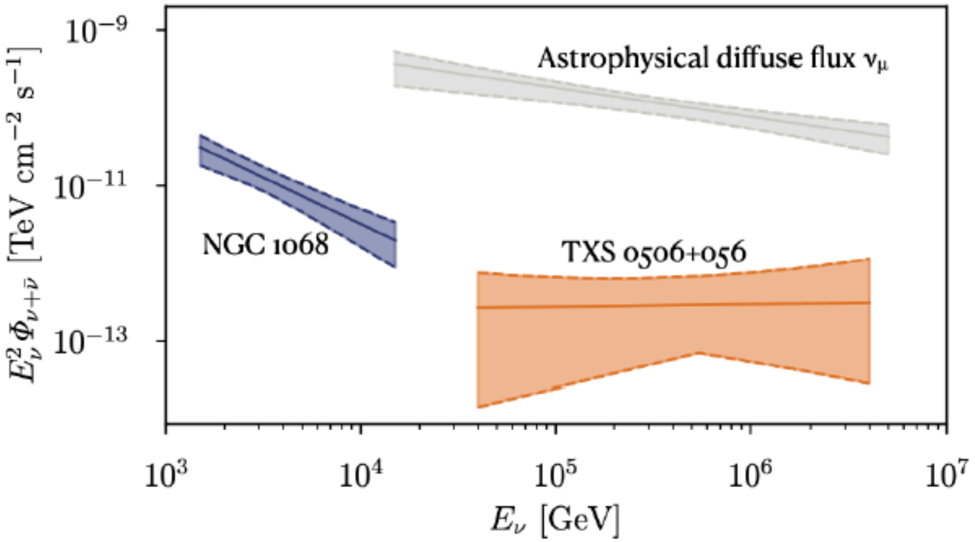
IceCube Collaboration*†

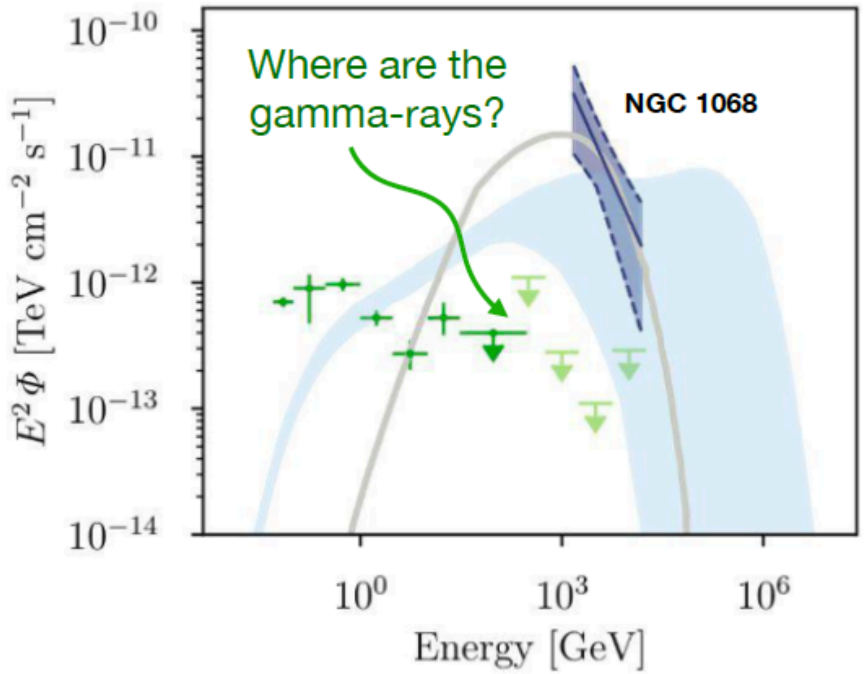
The observed AGNs with ν

NGC 1068

Science 378 (2022) 538-543

Neutrino Flux





- TXS 0506+056 and NGC 1068 contribute each ~1% of the total astrophysical diffuse neutrino
- Measured **neutrino flux** exceeds TeV **gamma-ray upper limits**

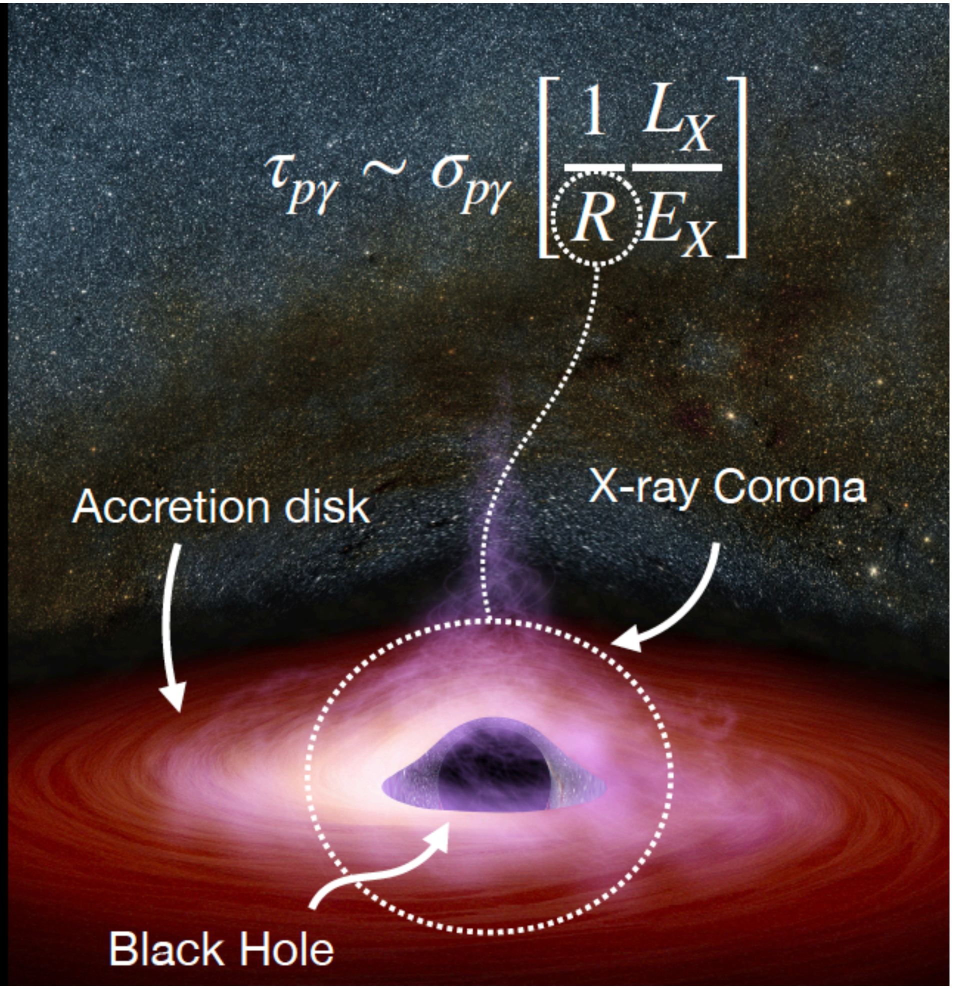
The observed AGNs with ν

The peculiarities of NGC1068

The observed AGNs with ν

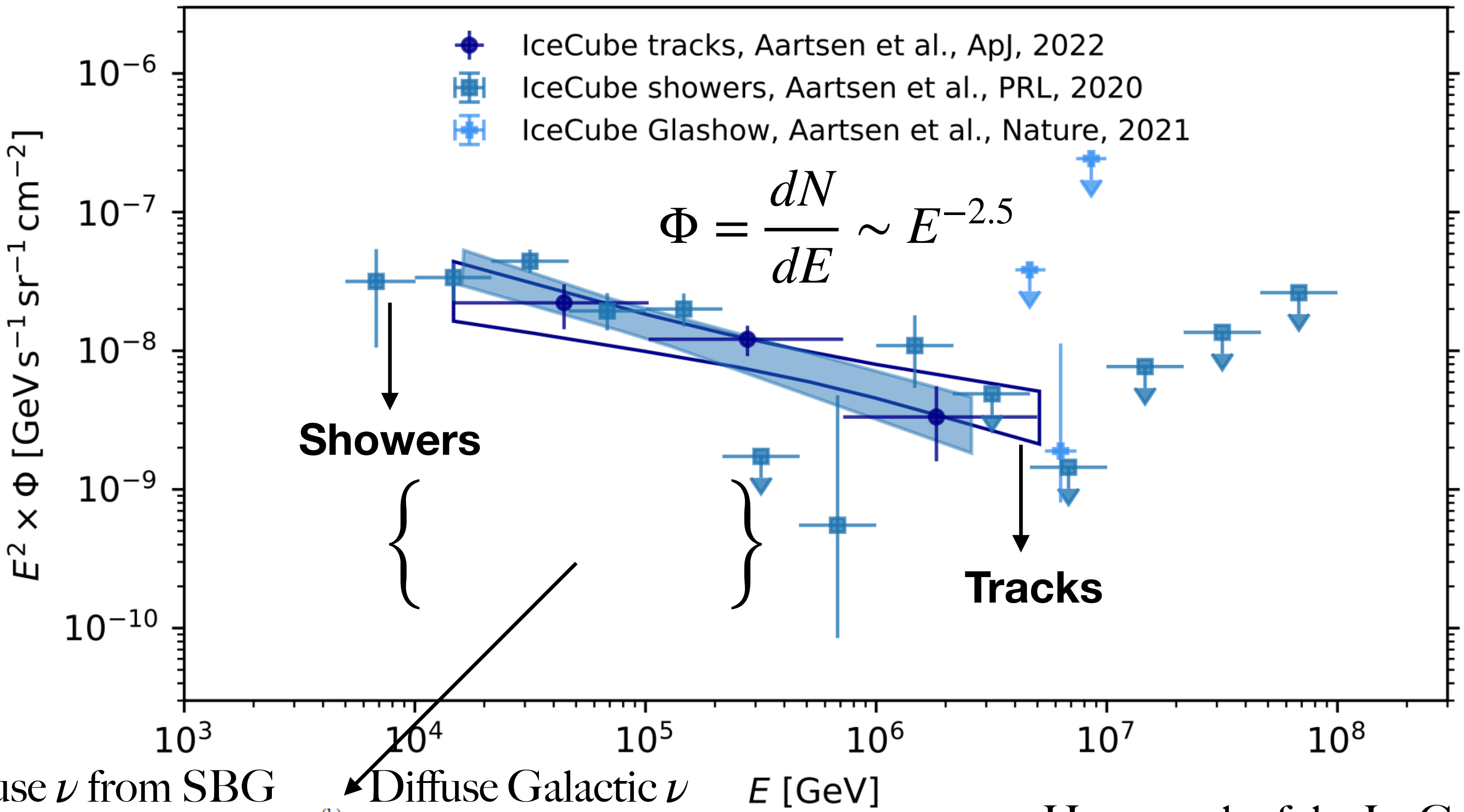
The Disk-Corona Model

- Electron and protons are accelerated in the high field regions associated with the black hole and the accretion disk
- They produce neutrinos in the optical thick corona
  - **Gamma-rays are absorbed**



In several IceCube presentations used as a universal description of neutrino emission from the entire AGN class and responsible for most of the astrophysical neutrinos observed by IceCube.

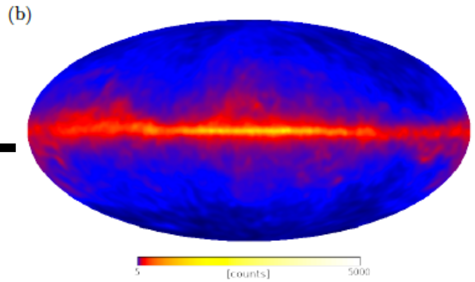
Setting the multicomponent energy ranges



Diffuse ν from SBG

\+

Diffuse Galactic ν



How much of the IceCube neutrinos can be related to non-AGN-galaxy emission.
Is the Milky Way a real neutrino desert?

Hadronic production in the SBGs

p-p interaction is likely to occur when
density of gas higher than density of radiation
(for example in Starburst Galaxies)

The Starburst Galaxy M82

Credit:
NASA, ESA and the Hubble Heritage Team (STScI/ AURA). Acknowledgment: J. Gallagher (University of Wisconsin), M. Mountain (STScI) and P. Puxley (NSF).

Properties of SBGs

- **~100 Myr phase** in the life of a Galaxy
- High Star Formation Rate **(10-100 times higher than Milky Way),** mainly concentrated into the nucleus
- They are abundant **(~10^4 — 10^5 Gpc^{-3})**
- Strong Magnetic field 10^2-$10^3 \mu$**G** and high degree of magnetic turbulence which traps protons for ~ 10^5 years (**Reservoirs**)
- Not very brilliant in gamma-rays **(only a small portion currently observed)**

Starburst Galaxies properties

Interstellar gas as the target

$$p + p \rightarrow \pi^+ \pi^- \pi^0 \ldots$$

Leaky-box-like model for CR transport

$$f(p)\left(\frac{1}{\tau_{\rm loss}(p)} + \frac{1}{\tau_{\rm adv}(p)} + \frac{1}{\tau_{\rm diff}(p)}\right) = Q(p)$$

injected CR from SN explosion

E.Peretti et al. MNRAS 2019

$$Q(p) \propto \left(\frac{p}{m_p}\right)^{-\alpha} \cdot e^{-p/p_{\rm max}}$$

CR acceleration up to the knee in Supernovae Remnants

- $$\tau_{\rm loss} \simeq \tau_{pp} \propto \frac{1}{n_{\rm ISM}}$$

Tloss > Tesc in the core

- The denser the SBN, the more the energy losses affects the CR transport

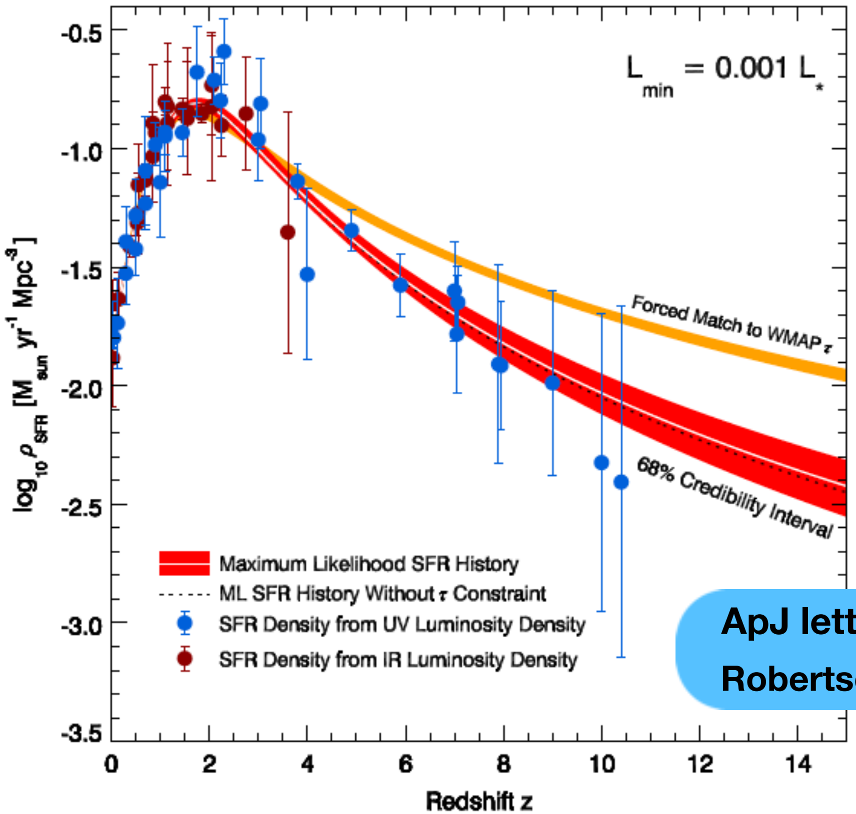
- $\tau_{\rm adv} = R/v_{\rm wind}$
- $\tau_{\rm diff} = R^2/D$

$R_{\rm SBN} = 200\,{\rm pc}$ $\quad v_{\rm wind} = 500\,{\rm km/s}$

✦ **Neutrinos** and **γ-rays** from pions decays:

$$\pi^{\pm} \rightarrow e^{\pm}\,\nu_e\,\nu_\mu\,\overline{\nu}_\mu$$
$$\pi^0 \rightarrow \gamma\,\gamma$$

Star formation rate can trace the ν emission



ApJ letter 2015 Robertson et al.

Important to estimate the diffuse ν emission of this galaxies population, we can follow the star formation density along the distance considered.

A proposed multimessenger fit

The Gamma-Ray Contributions:
1. SBGs
2. Blazar + Electromagnetic Cascades
3. Radio Galaxies

For Blazars and Radio Galaxies, we used the estimations given by Ajello et al. 2015 (ArXiv: 1501.05301)

The Neutrino Contributions:
1. SBGs
2. Blazars

For Blazars, we used the estimations given by Palladino et. Al 2019 (ArXiv:1806.04769)

MNRAS 503 4032 (2021) Ambrosone, Chianese, Fiorillo, A.M., Miele, Pisanti

Observational Samples Used

Extragalactic gamma-ray Background (EGB)

1. **7.5 years HESE**
2. **6 years Cascades**

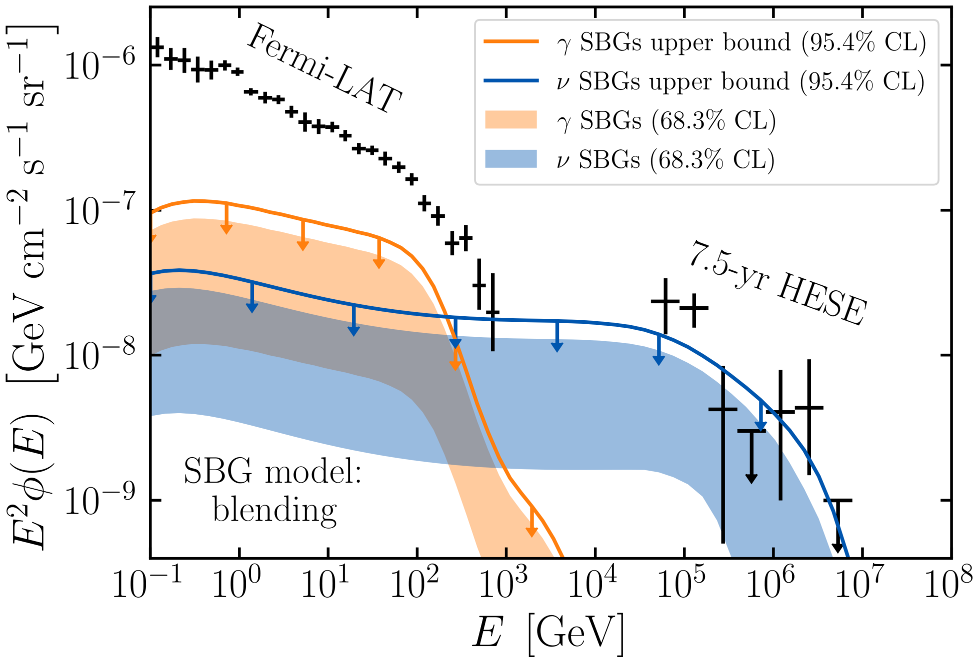
$$\chi^2_{\nu+\gamma}(N_{SBG}, N_{RG}, N_{Blazars}, p^{max}) = \chi^2_\nu + \chi^2_\gamma + \left(\frac{N_{Blazars}-1}{0.26}\right)^2 + \left(\frac{N_{RG}-1}{0.65}\right)^2 + \left(\frac{N_{Blazars}-0.80}{0.11}\right)^2$$

They come from uncertainties of the Non-SBG components

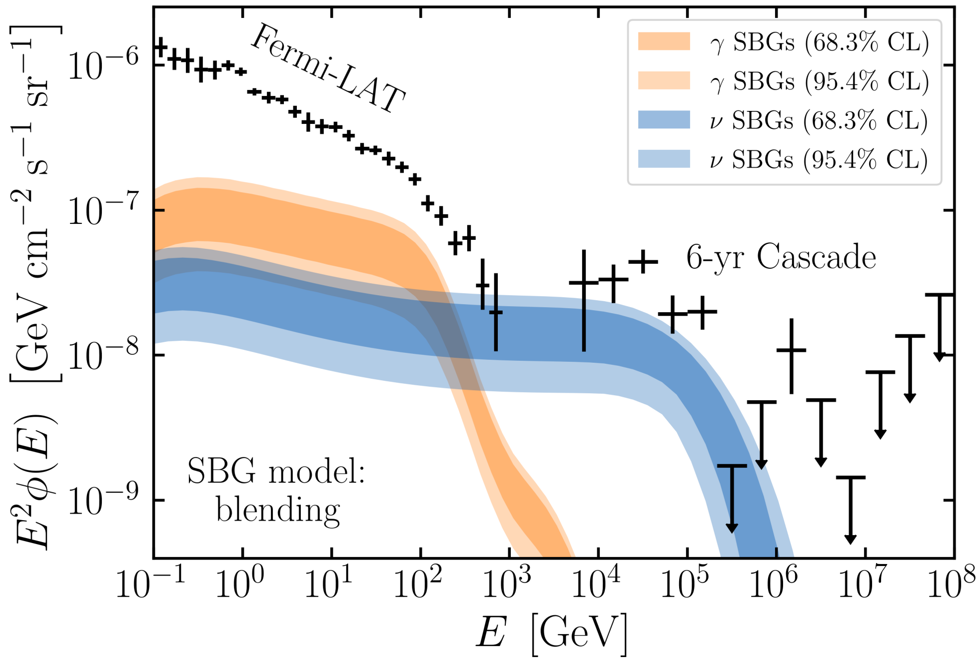
It comes from the positional limit of Point Sources above 50 GeV (Lisanti et al. 2016)

A proposed multimessenger fit

MNRAS 503 4032 (2021) Ambrosone, Chianese, Fiorillo, A.M., Miele, Pisanti



2 sigmas allowed SED considering Fermi-LAT EGB and IceCube HESE data

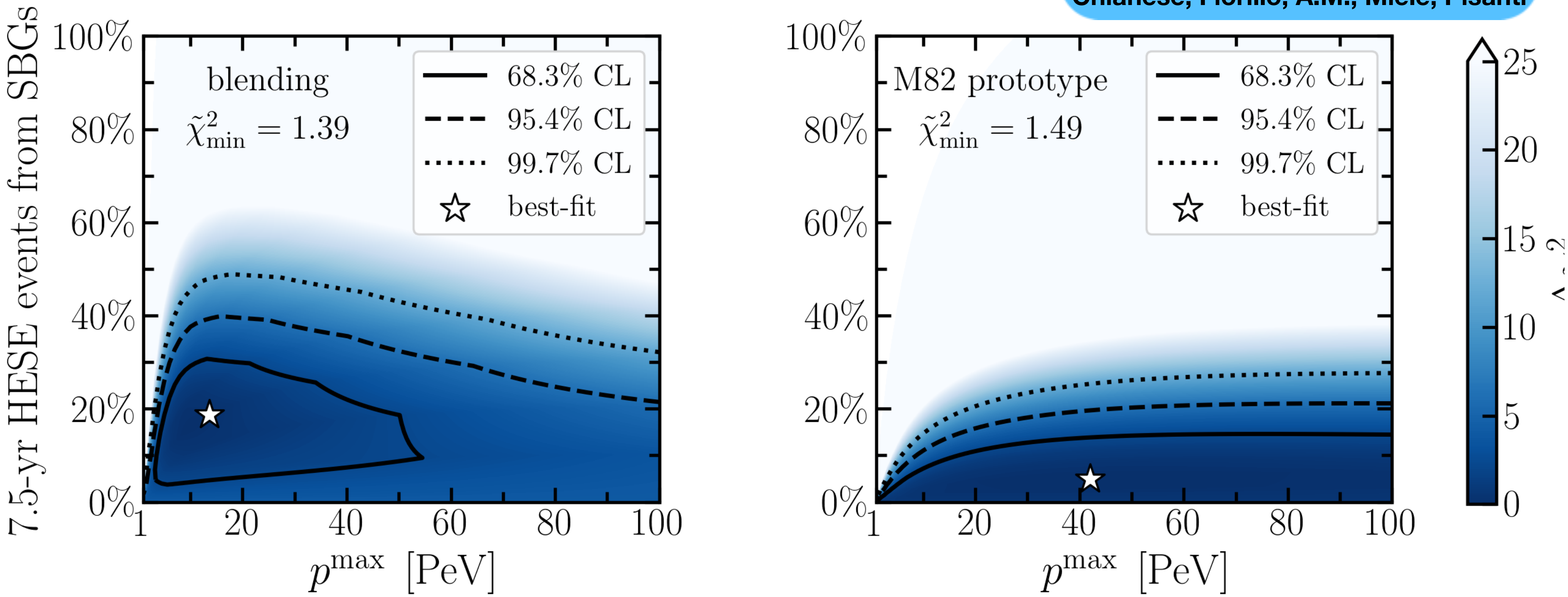


2 sigmas allowed SED considering Fermi-LAT EGB and IceCube CASCADE data

A proposed multimessenger fit

MNRAS 503 4032 (2021) Ambrosone, Chianese, Fiorillo, A.M., Miele, Pisanti

Redshift interval: $0 \leq z \leq 4.2$



At 2 sigma level the "blending" scenario can account up to 40% of IceCube HESE measured flux, moreover at 1 sigma a Pmax up to 50 PeV is permitted, however a cutoff ~ 10 PeV is favored.

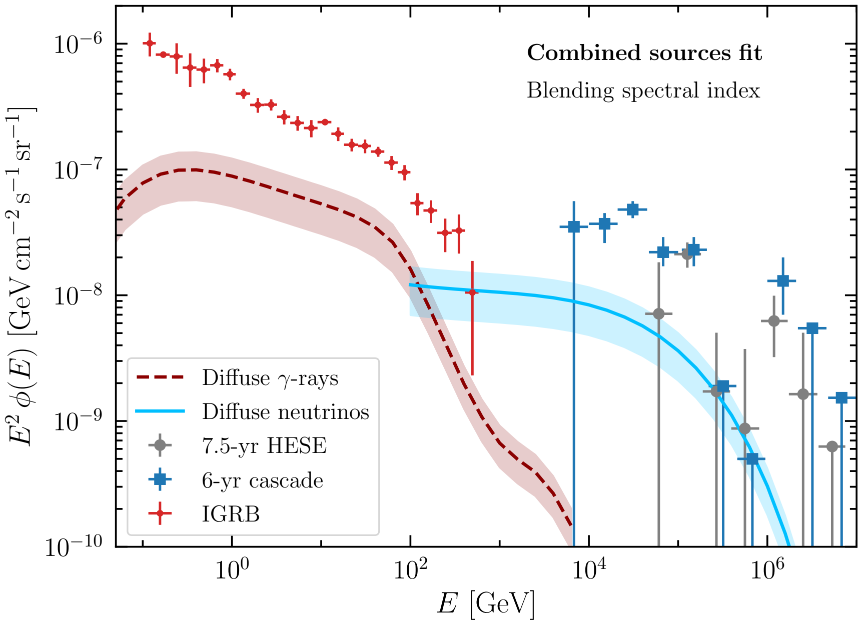
Diffuse Fluxes of SBGs using the entire Fermi-LAT data

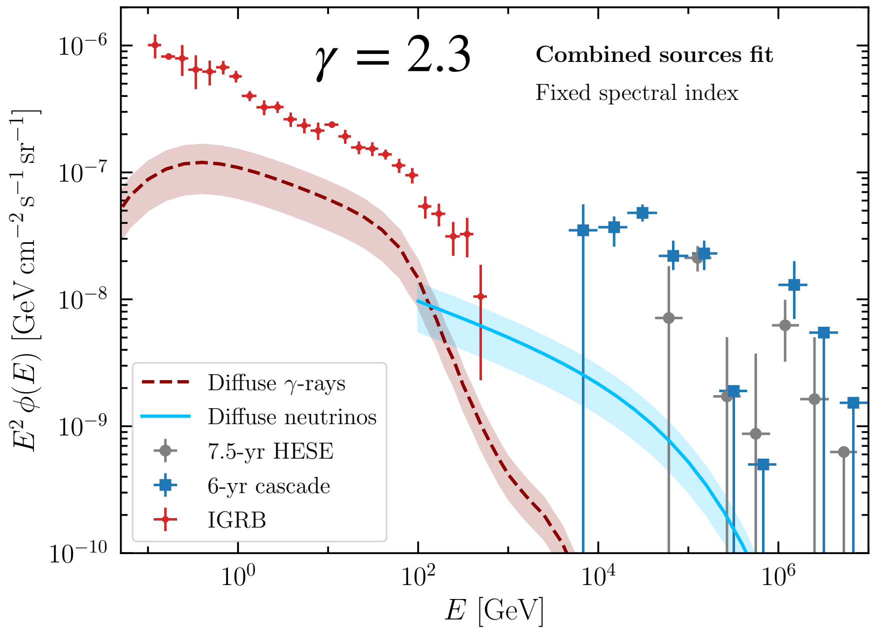
The emission from all SBGs in the Universe

$$\phi^{\rm diff}_{\gamma,\nu} = \frac{c}{4\pi H_0}\int_0^{z_{\max}} \frac{{\rm d}z}{E(z)}\int_{10^6 L_\odot}^{\infty} \frac{{\rm d}L_{\rm IR}}{\ln(10)\, L_{\rm IR}} \mathcal{S}_{\rm SFR}(L_{\rm IR}, z) \times Q_{\gamma,\nu}\left(E(1+z), {\rm R_{SN}}({\rm L_{IR}}), {\rm F_{cal}}({\rm R_{SN}}({\rm L_{IR}}))\right)\, e^{-\tau_{\gamma,\nu}(E,z,L_{\rm IR})}$$

$E_{\max}$ assumed is 10 PeV

JCAP08(2024)040 Ambrosone, Chianese, + A.M.





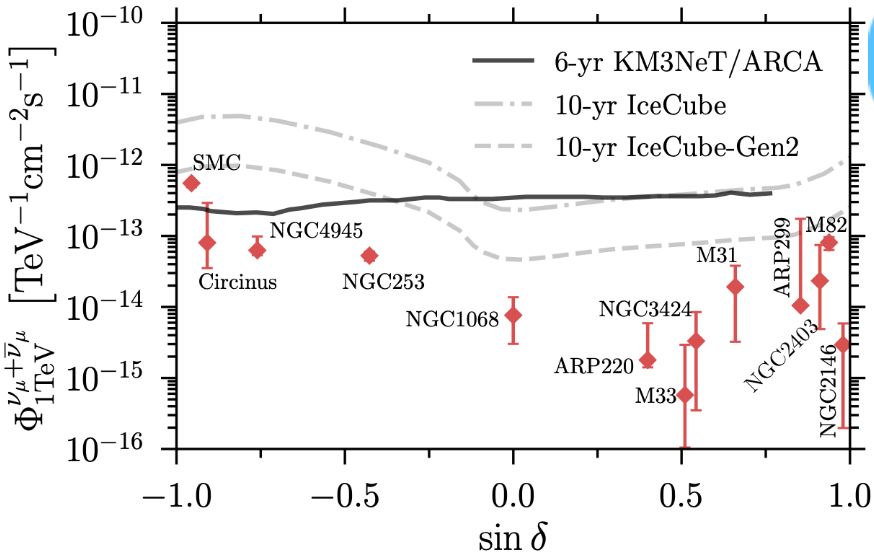
- The blending scenario increases the neutrino flux ($\sim 20\,\%$ of the IceCube measurements)
- The neutrino flux is constrained to be $\sim$ % of the IceCube diffuse Fluxes

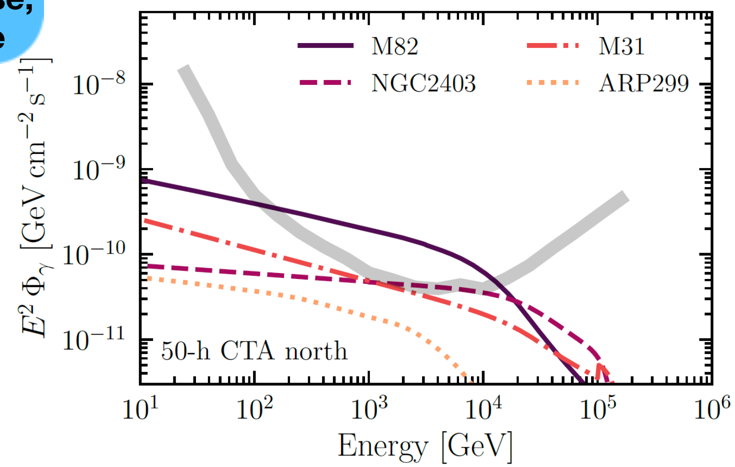
Probing the SBG Calorimetric Scenario

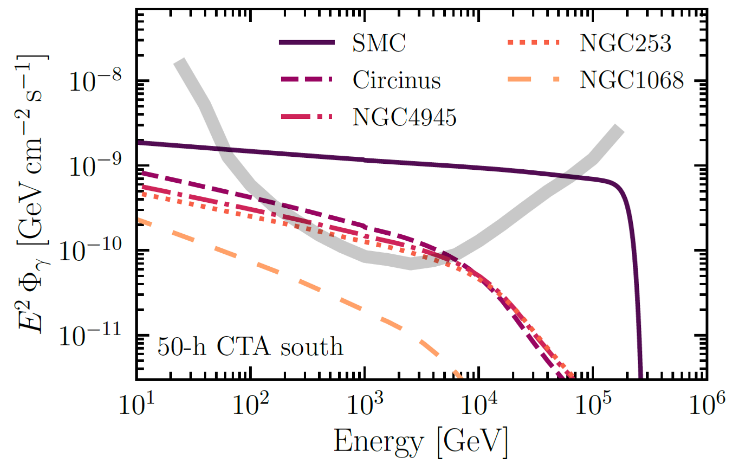
Tracing Neutrino Emission from Local resolved SFGs

Neutrino Expectations: KM3NeT Forecast Gamma-Rays Expectations. CTA Forecast

ApJL 919 2021
Ambrosone, Chianese, Fiorillo, A.M., Miele





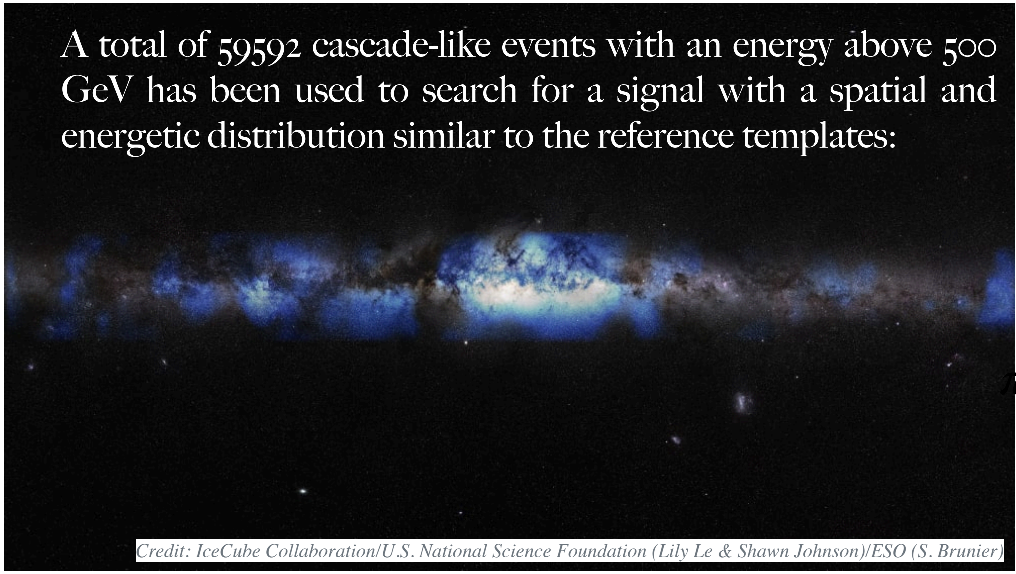


Future γ/ν observations will be fundamental to:

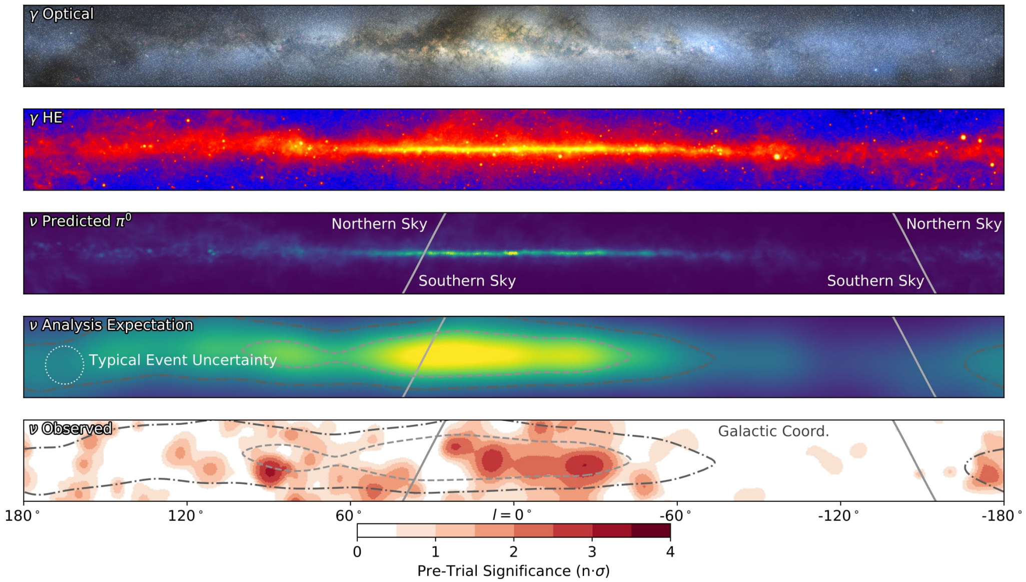
- ✦ Discover if Neutrino Astronomy is a tracer for star-forming activity
- ✦ Probe the calorimetric fraction inside SBG: If there will be no detection, nearby SBGs are dominated by diffusion and not by either p-p collisions or advection.

High energy emission from Milky Way: a new piece of diffuse flux puzzle : VHE ν

IceCube Coll.
Science 2023



π_0

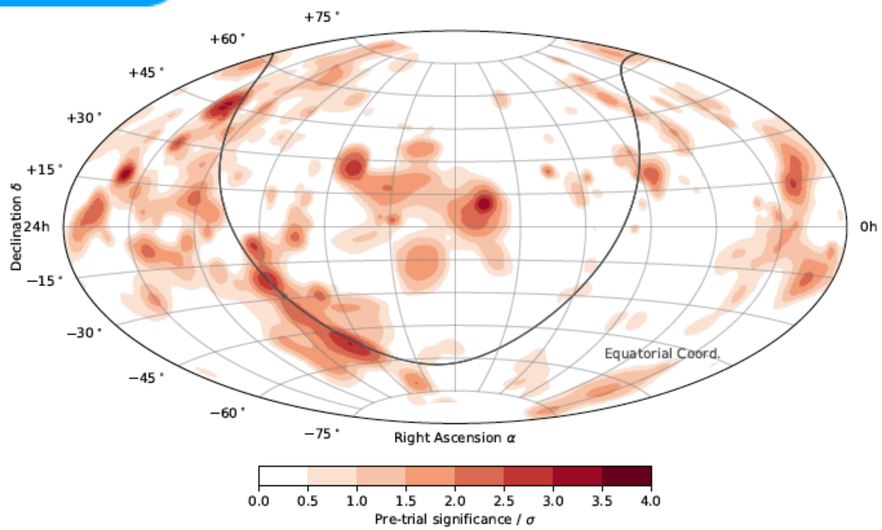


Π_0 → A Fermi-LAT coll. template based on a homogeneous diffusion coefficient along the Milky Way longitude and a 2012 molecular gas map.

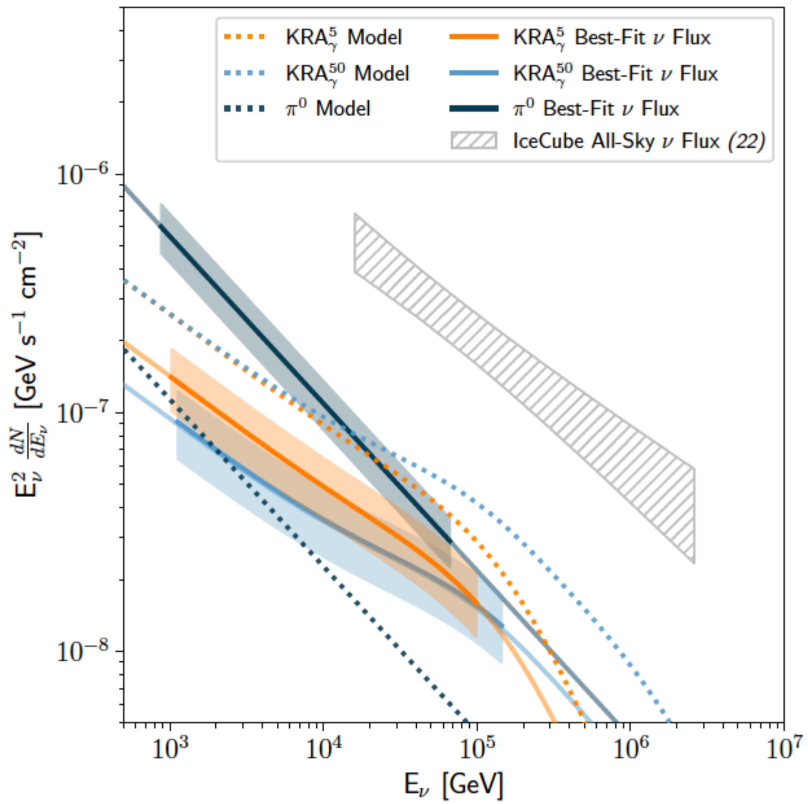
KRA-γ5 and KRA-γ50 → A template obtained with DRAGON and Gamma-sky codes based on a inhomogeneous diffusion coefficient and a CR spectral hardening toward the Milky Way center (radial dependent) and two different CR cutoffs at 5 and 50 PeV

IceCube observation of Galactic neutrinos

IceCube coll. Science 2023



IceCube CASCADE best fit pre-trial significance distribution map



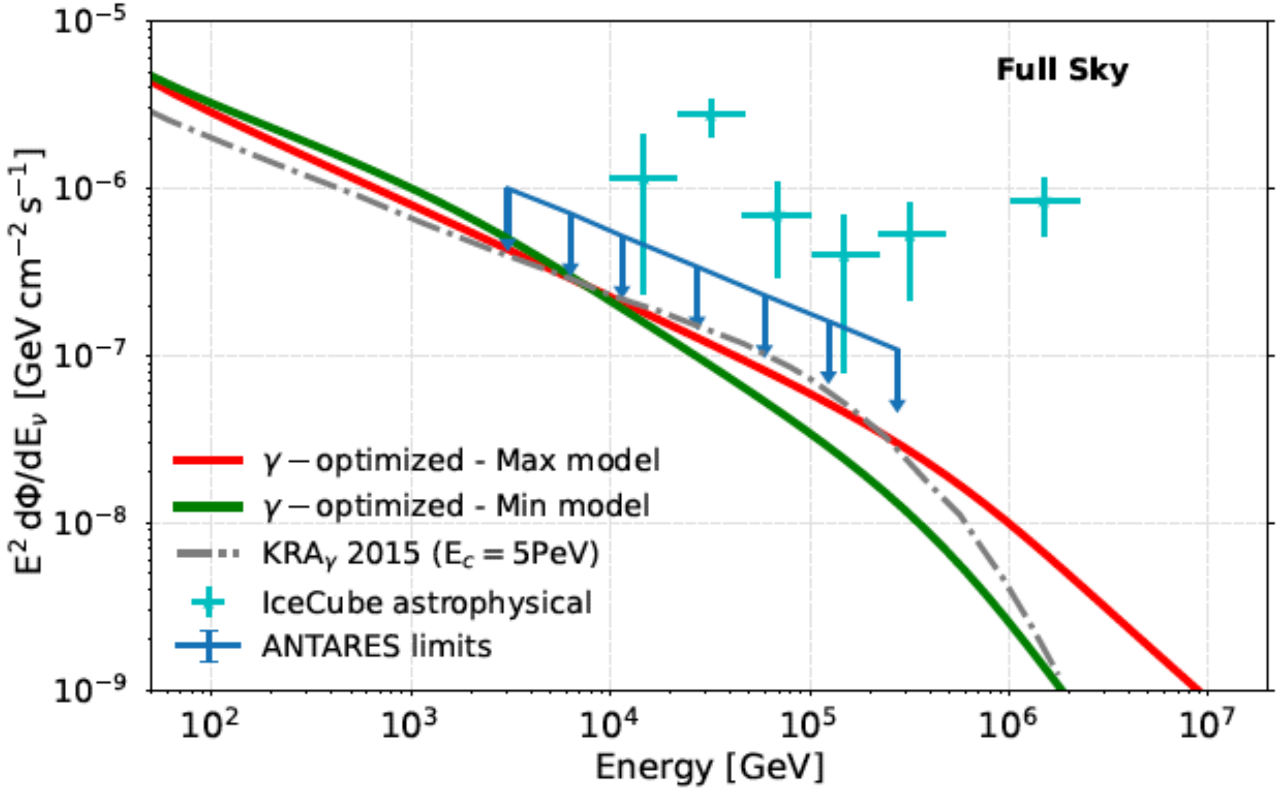
Diffuse Galactic plane analyses	Flux sensitivity Φ	p-value	Best-fitting flux Φ
π^0	5.98	1.26×10^{-6} (4.71σ)	$21.8^{+5.3}_{-4.9}$
KRA$^5_\gamma$	0.16×MF	6.13×10^{-6} (4.37σ)	$0.55^{+0.18}_{-0.15}$×MF
KRA$^{50}_\gamma$	0.11×MF	3.72×10^{-5} (3.96σ)	$0.37^{+0.13}_{-0.11}$×MF

→ Considering the obtained best fit normalizations seems the more motivated case.

ν expectations from the new KRA-γ models

De La Torre, Gaggero, Grasso, A.M.
Frontiers 2022

The expected new full sky ν SED in comparison with IceCube and ANTARES

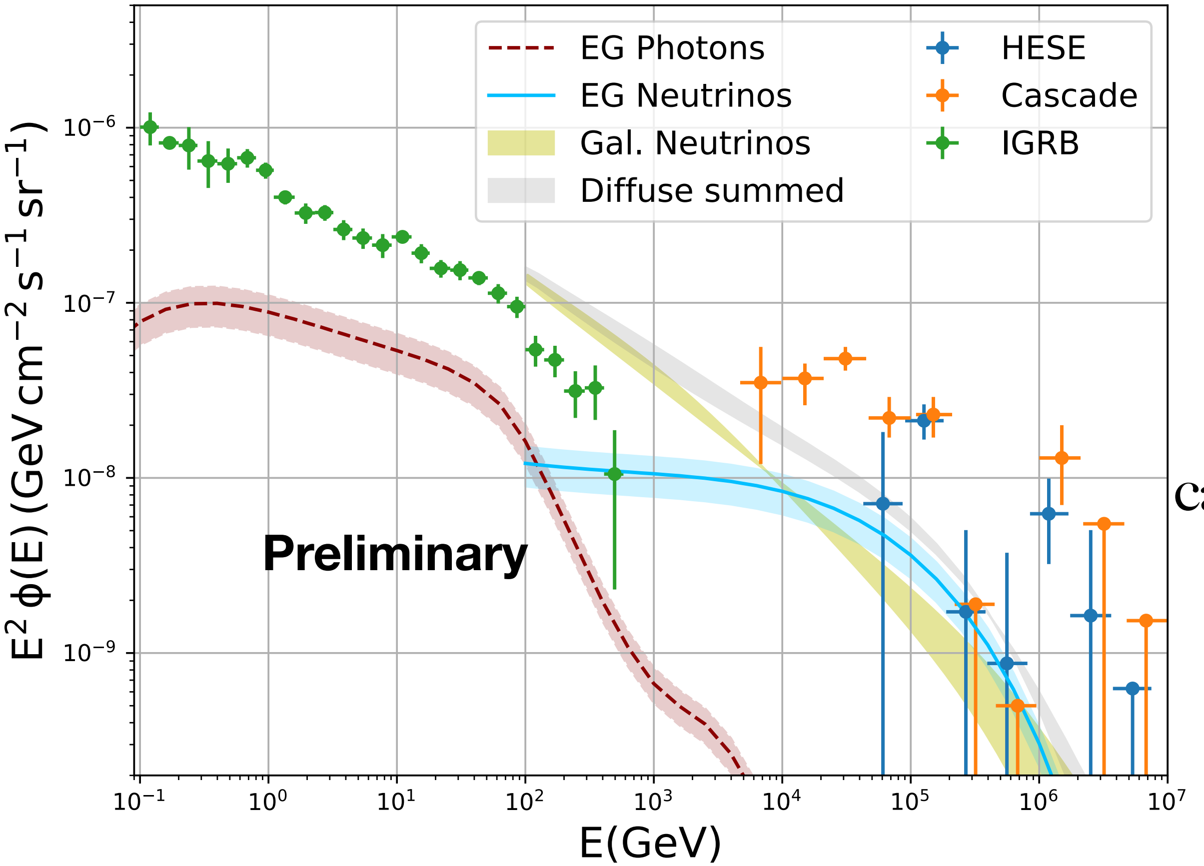


Still at the level of 10% of IceCube full sky observations and compatible with ANTARES model independent old upper limits

The updated KRA-gammas remain consistent with the previous KRA-gamma with CR cutoff at 5 PeV.

A diffuse emission associated to non AGNs

Milky Way diffuse ν emission spatially related to Galactic disk while the SBG diffuse emission is a isotropic component



EG neutrino = Diffuse SFG component

Gal Neutrino = Diffuse Galactic from the updated KRA_γ model

The sum of these two components can account for a non negligible part of the flux measured by IceCube

The absence of a AGN for the SFG population does not prevent a accountable diffuse neutrino emission when integrating over high redshifts.

SUMMARY

- Galaxies who don't host a AGN can still be a important astrophysical component to explain the IceCube diffuse observations.
- Star forming abundance can be a tracer of neutrino emission.
- Single Starburst galaxies are hardly resolved as a point like ν emitters however the ν diffuse component becomes important when integrating up to high redshifts.
- IceCube and Antares start to characterize the Galactic neutrino emission (still nice questions to go through).
- The low energy part of the IceCube observed astrophysical neutrinos can be explained by reservoirs emission: starforming galaxies+ Milky Way

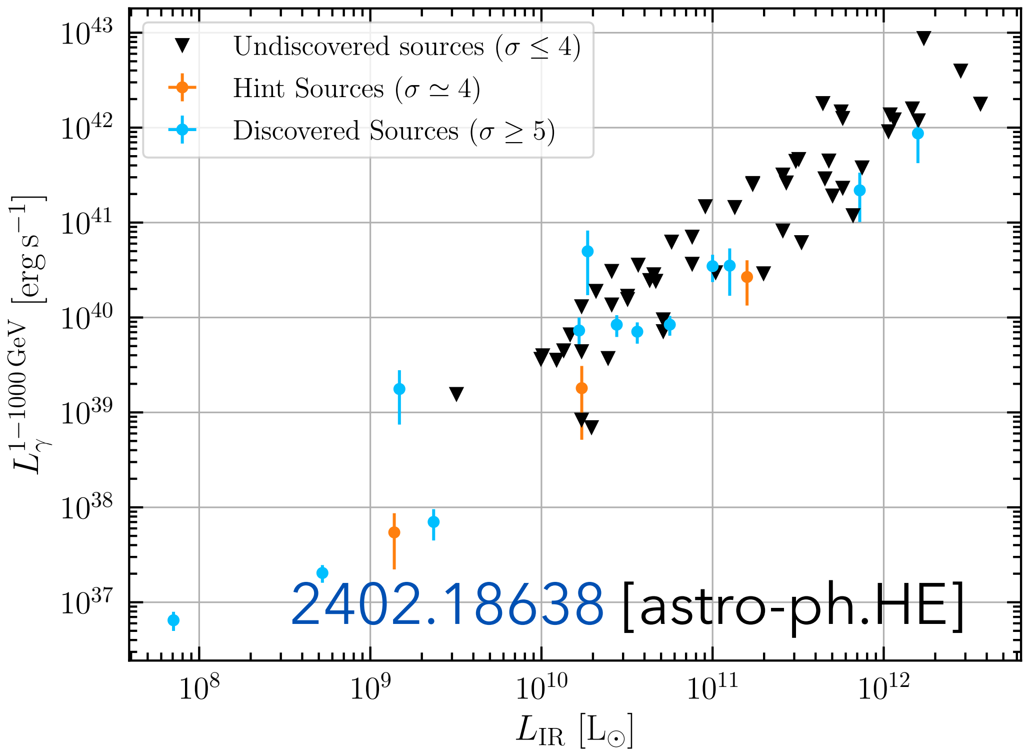
BACK UP SLIDES

Relation star-formation and γ-rays luminosity

There is a tight correlation between the γ-ray luminosity and the Infrared (IR) Luminosity

Redshift interval: $0 \leq z \leq 8$

Analysing a catalogue of 70 sources with 15 years of Fermi-LAT data



2402.18638 [astro-ph.HE]

A new analysis

The IR Luminosity is strictly connected to the Star Formation Rate (SFR)

$$\mathrm{SFR} = 1.36 \cdot 10^{-10} \left(\frac{L_{\mathrm{IR}}}{L_\odot}\right)\left(1 + \sqrt{\frac{10^9 L_\odot}{L_{\mathrm{IR}}}}\right) [M_\odot\, \mathrm{yr}^{-1}]$$

- The higher the SFR, the more CRs get injected in SBG disk
- The higher the SFR, the more dense the system is and the CRs are trapped into the system **(Complete CR calorimetry)**

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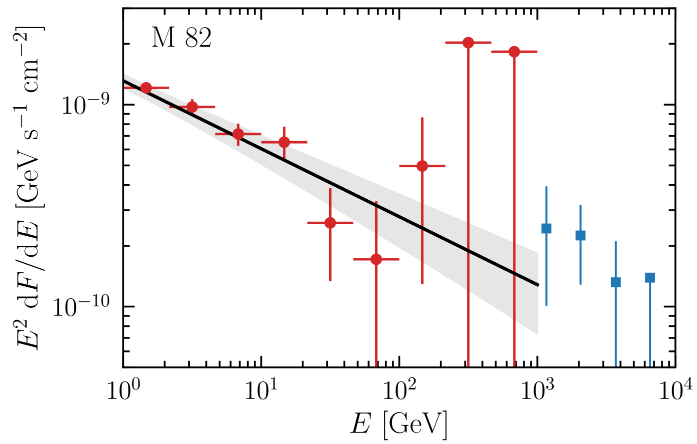
Properties of Discovered Sources

14 sources are discovered with more than $5\sigma\,(TS > 25)$

$E_\gamma \in [1 - 1000]$ GeV

Take a look at 2402.18638 [astro-ph.HE]

Source	D_L [Mpc]	L_{IR} [10^{10} $L_\odot$]	$F_{1-1000\,GeV}$ [10^{-10} ph cm^{-2} s^{-1}]	ϕ_0 [10^{-12} MeV^{-1} cm^{-2} s^{-1}]	γ	TS (σ)	TS_{SM}
M 82	3.53	5.6	9.8 ± 0.5	1.31 ± 0.10	2.34 ± 0.06	1104 (33)	0.35
NGC 253	3.56	3.6	8.1 ± 0.9	1.08 ± 0.10	2.33 ± 0.08	730 (27)	1.03
ARP 220	84.3	$1.7 \cdot 10^2$	1.6 ± 0.6	$(2.0 \pm 0.7) \cdot 10^{-1}$	2.2 ± 0.2	50 (7.1)	–
NGC 1068	10.1	10.0	4.5 ± 0.5	$(5.8 \pm 0.9) \cdot 10^{-1}$	2.28 ± 0.15	238 (15)	–
Circinus	4.21	1.7	5.1 ± 1.3	$(6.2 \pm 1.7) \cdot 10^{-1}$	2.23 ± 0.14	78 (8.8)	–
SMC	0.06	$7.1 \cdot 10^{-3}$	$(3.0 \pm 0.3) \cdot 10^1$	4.4 ± 0.3	2.44 ± 0.06	801 (28)	4.13
M 31	0.77	$2.3 \cdot 10^{-1}$	3.1 ± 0.8	$(6.3 \pm 1.3) \cdot 10^{-1}$	3.0 ± 0.3	74.6 (8.6)	0.22
NGC 2146	17.2	12.6	1.3 ± 0.5	$(1.5 \pm 0.5) \cdot 10^{-1}$	2.16 ± 0.18	41.5 (6.4)	–
ARP 299	48.6	72.6	1.3 ± 0.5	$(1.7 \cdot 0.6) \cdot 10^{-1}$	2.3 ± 0.2	46.4 (6.8)	–
NGC 4945	3.72	2.8	9.6 ± 1.3	1.34 ± 0.15	2.40 ± 0.08	412 (20)	–
NGC 2403	3.18	0.15	1.5 ± 0.5	$(10 \pm 4) \cdot 10^{-2}$	1.92 ± 0.17	52.8 (7.3)	–
NGC 3424	27.2	2.1	10 ± 5	$(1.3 \pm 0.5) \cdot 10^{-1}$	2.3 ± 0.3	28 (5.3)	–
LMC	0.05	$5.2 \cdot 10^{-2}$	$(1.38 \pm 0.07) \cdot 10^2$	$(1.85 \pm 0.08) \cdot 10^1$	2.41 ± 0.04	1493 (38)	0.24
M 33	0.91	0.14	$1.2 \pm 0.6^\dagger$	$(1.8 \pm 0.7) \cdot 10^{-1}$	2.5 ± 0.3	16 (4)	–



Two sources have a strong hint of γ-ray emissions($\sim 4\sigma$)

M83, NGC 1365

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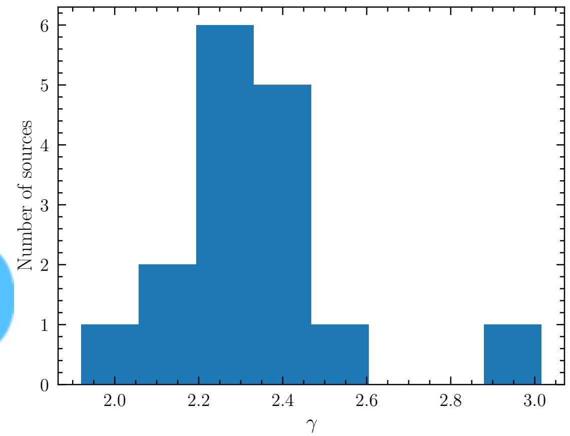
- All spectra are consistent with simple power-laws

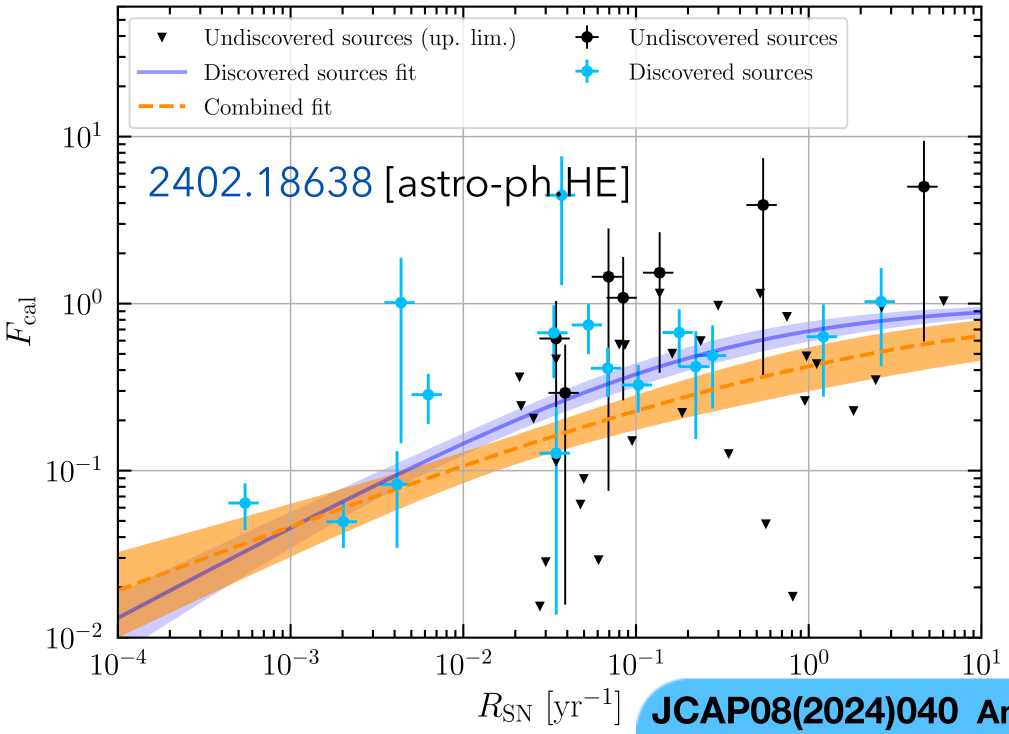


Calorimetric Fraction and Star Formation Rate

$$f(p)\left(\frac{1}{\tau_{\rm loss}(p)}+\frac{1}{\tau_{\rm adv}(p)}+\frac{1}{\tau_{\rm diff}(p)}\right)=Q(p) \quad \Rightarrow \quad f(p) \simeq Q(p)\tau_{pp} F_{\rm cal}$$

$F_{\rm cal}$: Fraction of CRs which actually interact and produce γ and ν

Average $F_{\rm cal}$ between $10 - 10^4\,\mathrm{GeV}$ for CRs



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$$R_{\rm SN}[{\rm yr}^{-1}] \simeq \frac{1}{83}{\rm SFR}[{\rm M}_\odot\,{\rm yr}^{-1}]$$

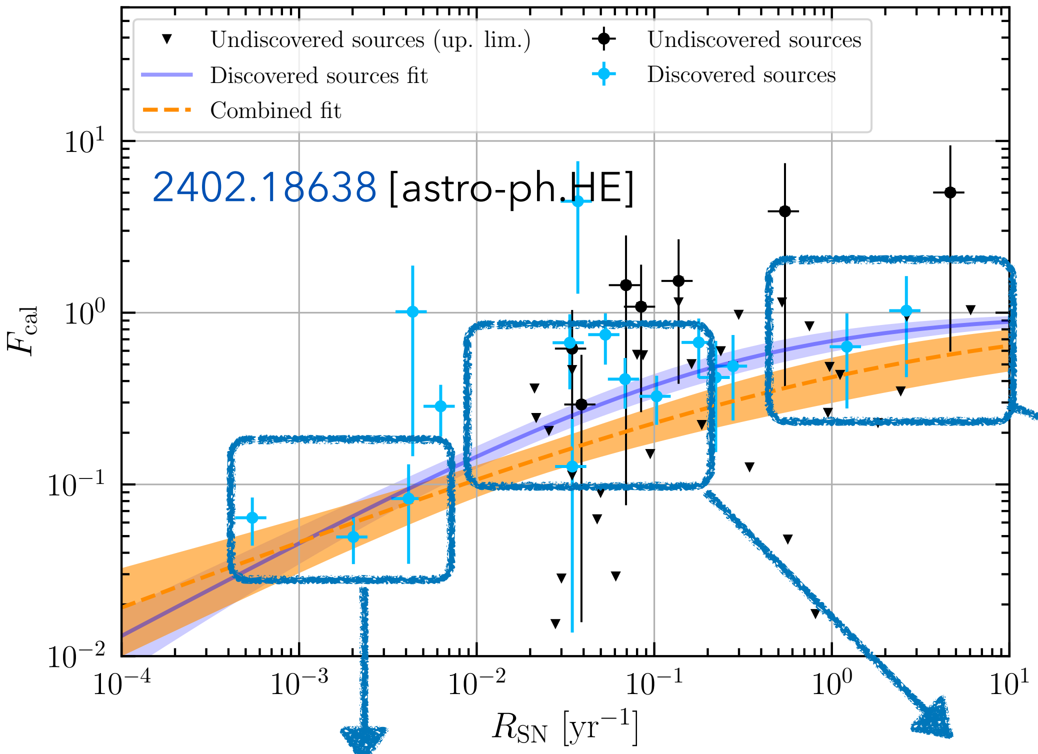
$F_{\rm cal}$ correlates with the SFR and the Supernovae explosion rate

$$F_{\rm cal} = A\left(\frac{R_{\rm SN}}{{\rm yr}^{-1}}\right)^{\beta}\left(1 + A\left(\frac{R_{\rm SN}}{{\rm yr}^{-1}}\right)^{\beta}\right)^{-1}$$

$$A = 0.7^{+0.3}_{-0.2} \quad \beta = 0.39 \pm 0.07$$

Calorimetric Fraction and Star Formation Rate

Average $F_{\rm cal}$ between $10 - 10^4\,\mathrm{GeV}$ for CRs



$$F_{\rm cal} = A \left(\frac{R_{\rm SN}}{\rm yr^{-1}}\right)^{\beta} \left(1 + A\left(\frac{R_{\rm SN}}{\rm yr^{-1}}\right)^{\beta}\right)^{-1}$$

$$A = 0.7^{+0.3}_{-0.2} \quad \beta = 0.39 \pm 0.07$$

Powerful Starburst: ARP 299, ARP 220

$\mathrm{SFR} \sim 80 - 200\,\mathrm{M_\odot\,yr^{-1}}$

At low SFR, the local galaxies dominated (SMC,LMC, M31, M33)

$\mathrm{SFR} \sim 10^{-2} - 10^{-1}\,\mathrm{M_\odot\,yr^{-1}}$

Mild Starburst: M82, NGC 253, NGC 1068

$\mathrm{SFR} \sim 5 - 20\,\mathrm{M_\odot\,yr^{-1}}$

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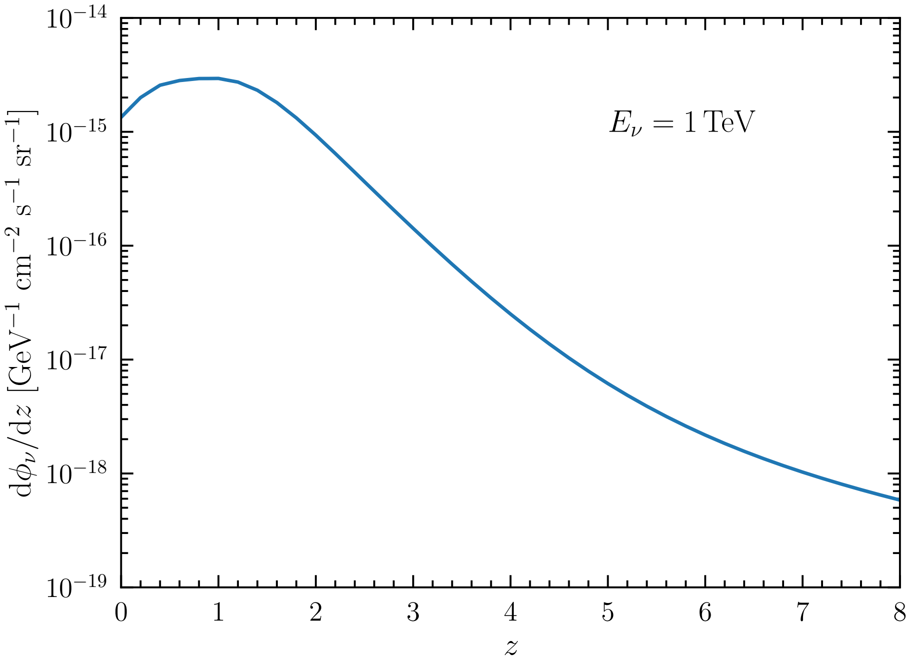
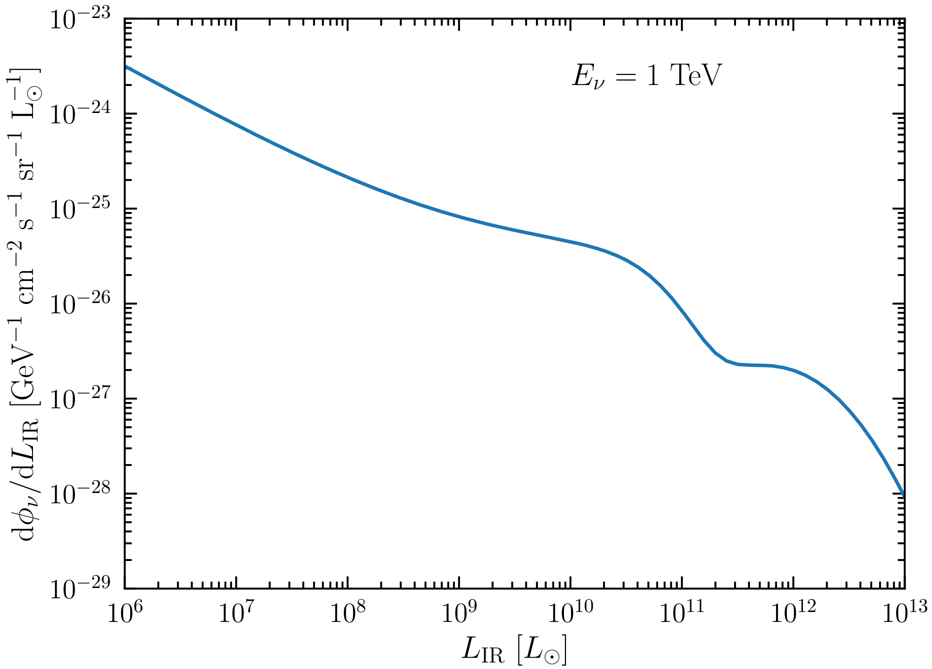
Properties of the Neutrino Flux

Redshift Distribution of the Neutrino Flux



Distant sources dominate the emission (peaking at $z \simeq 1$)

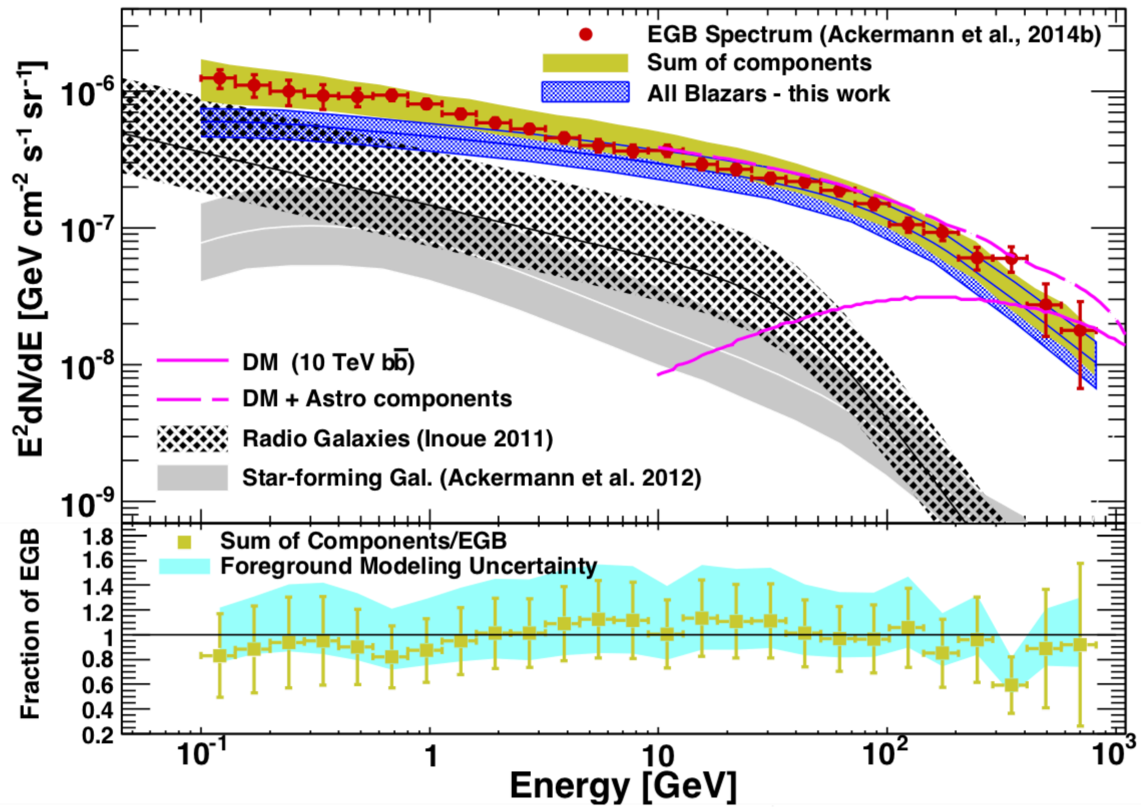
IR Luminosity Distribution of the Neutrir



Powerful SBGs $(\mathrm{SFR} > 100\,\mathrm{M}_\odot\,\mathrm{yr}^{-1})$ dominate the emission (> 50% of the emission)

extragalactic gamma-ray background

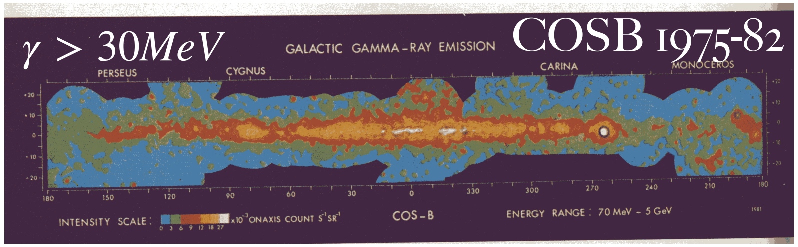
Ajello et al., APJL 800 (2015)

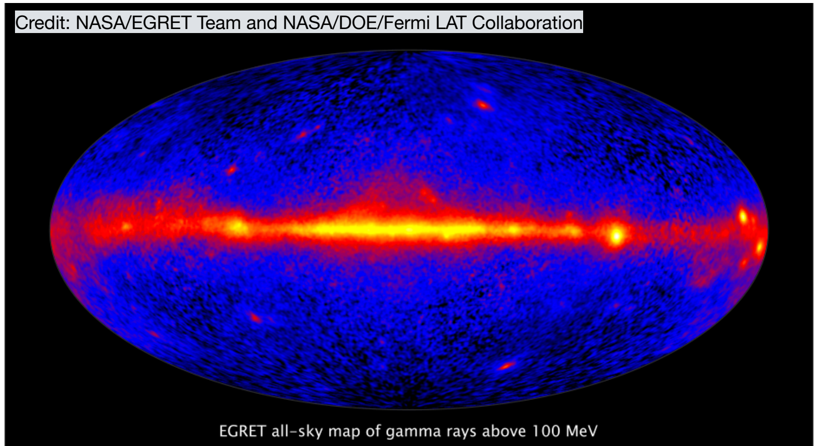


- Fermi-LAT resolved many individual sources belonging to different classes, Blazars dominates the EG samples.
- Limit on PS above 50 GeV varies from 68% (Lisanti et al. 2016) to 86% (Ackermann et al. 2016) of the EGB

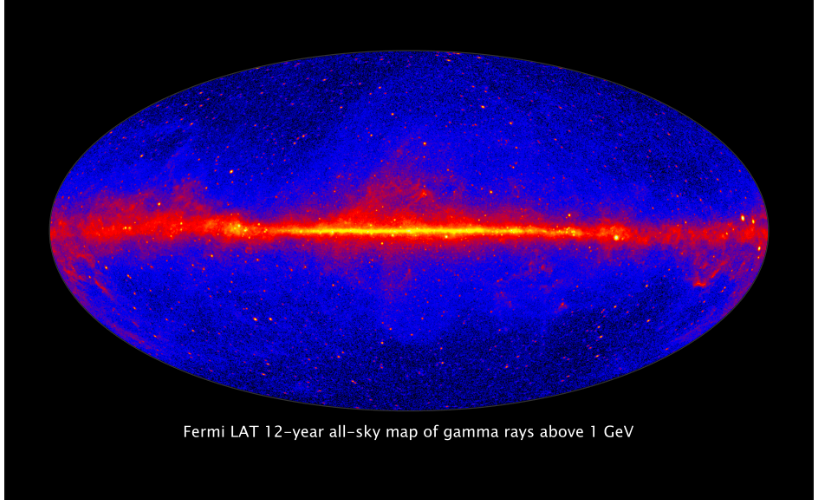
Starforming and Starburst galaxies gamma-ray component needs a better definition due to the small number of resolved ones at HE

Gamma-ray emission from the Milky Way



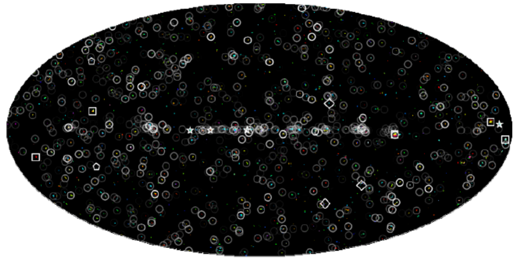


EGRET all-sky map of gamma rays above 100 MeV



Fermi LAT 12-year all-sky map of gamma rays above 1 GeV

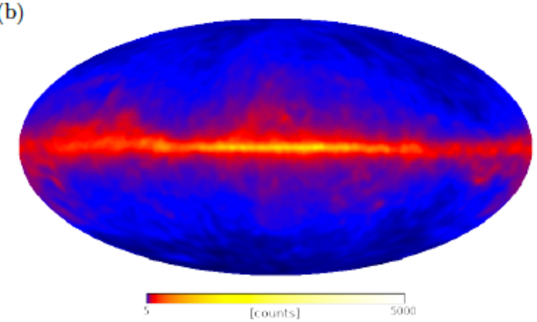
Selig et al. A&A 2014



Resolved sources

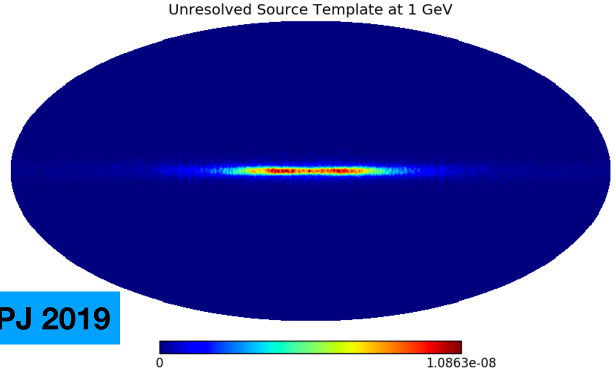
Selig et al. A&A 2014

(b)



Diffuse emission

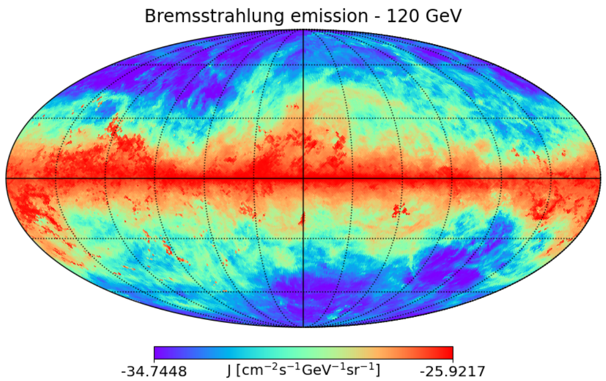
Unresolved Source Template at 1 GeV



Unresolved sources

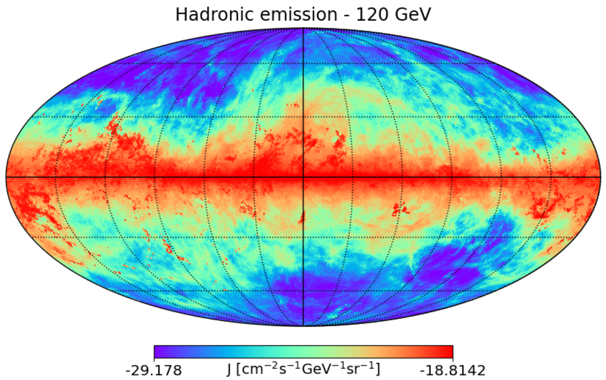
Fermi-LAT coll. APJ 2019

Galactic diffuse gamma-ray emission



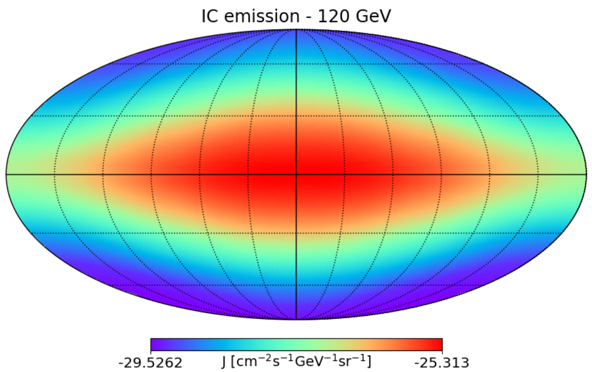
$$e + N \rightarrow e^{'} + \gamma + N^{'}$$

Bremsstrahlung emission follows the ISM gas distribution



$$p + p \rightarrow \pi_0 \pi_+ \pi_-$$
$$\pi^+ \rightarrow \mu^+ + \nu_\mu$$
$$\mu^+ \rightarrow \bar{\nu}_\mu + \nu_e + e^+$$

Diffuse emission totally correlated with the propagation of cosmic rays dominated by protons and He. Hadronic emission follows ISM gas distribution as well.



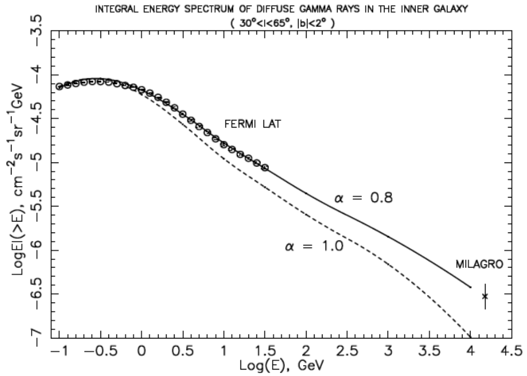
$$e + \gamma \rightarrow e^{'} + \gamma^{'}$$

IC emission depends on the energy density of the **ISRFs**

Inhomogenous diffusion model

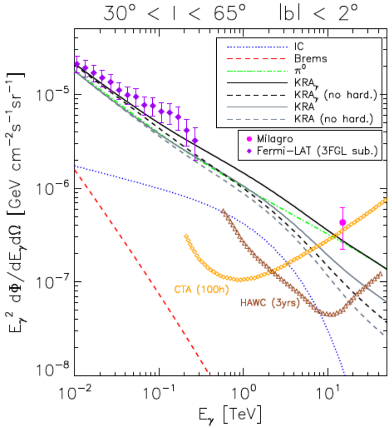
$D \propto E^{\delta(R)}$ with $\delta(R) = \delta_0 + \delta_a(R)$ Diffusion coefficient Change toward the Galactic center

- Erlykin & Wolfendale (APh 2012)

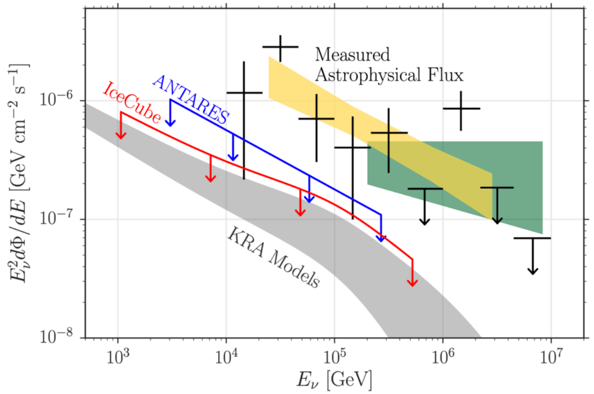


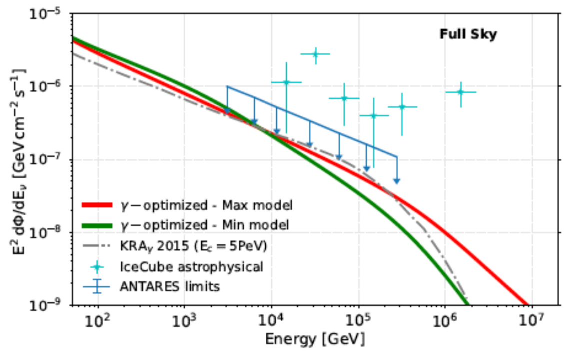
- Gaggero et al. (PRD 2015) ,
  Gaggero el al. (APJ letter 2015)

KRA-γ with CR cutoff at 5 and 50 PeV



maps produced for γ and ν

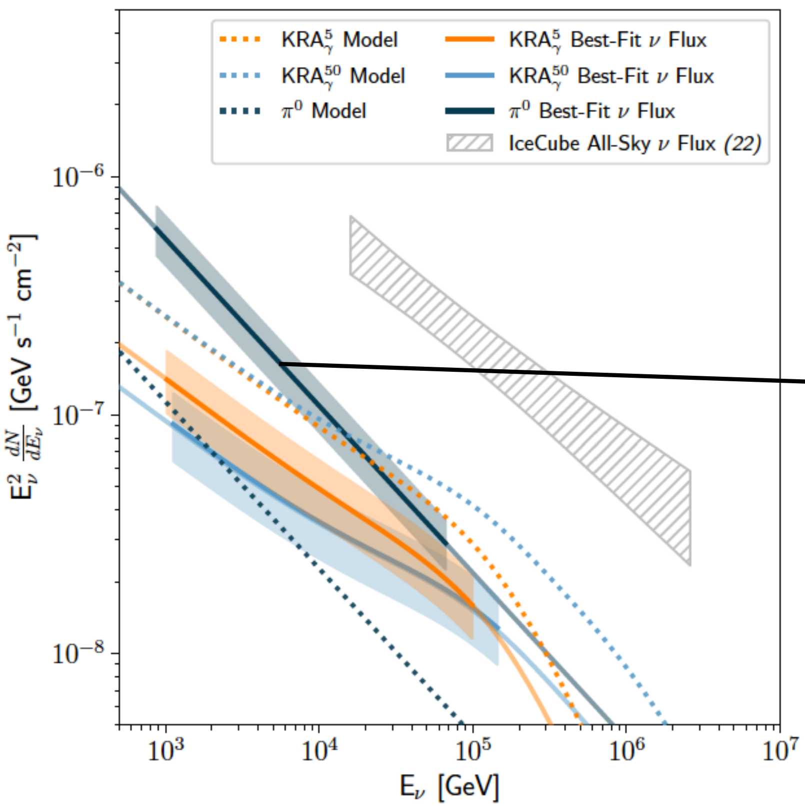




- De la Torre et al. (A&A 2021) , De la Torre et al. (Frontiers 2022)

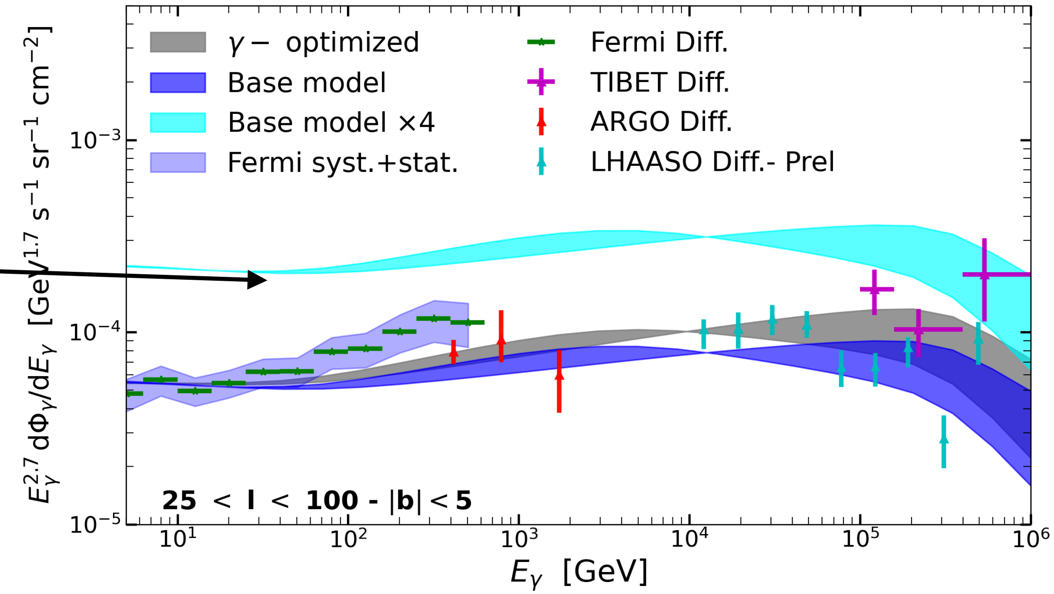
A better look to the IceCube results

IceCube coll.
Science 2023



Base Model mostly equal to what is
is called π_0 model in the IceCube Science 2023

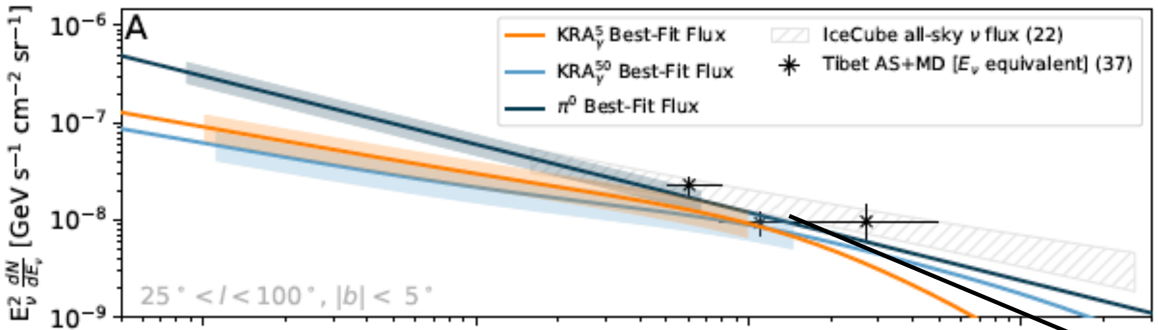
De la Torre et al.
A&A 2023



The best fit normalization of the π_0 model (4 times the expected value) present a strong tension with Galactic diffuse Fermi-LAT observations.

A better look to the IceCube results

IceCube coll.
Science 2023

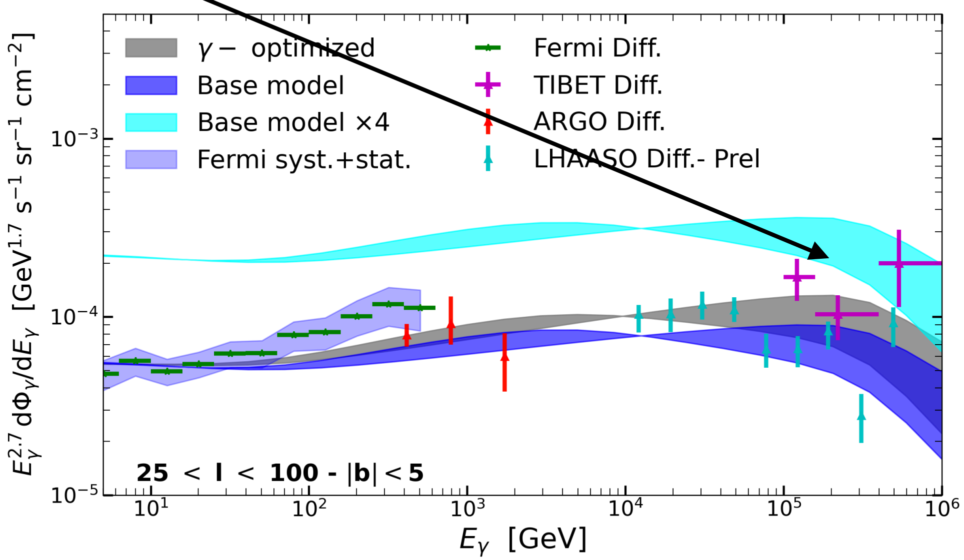


Caveat: the best-fit of observed ν is not obtained for the Tibet region, IceCube coll. rescaled of the measured flux

De la Torre et al.
A&A 2023

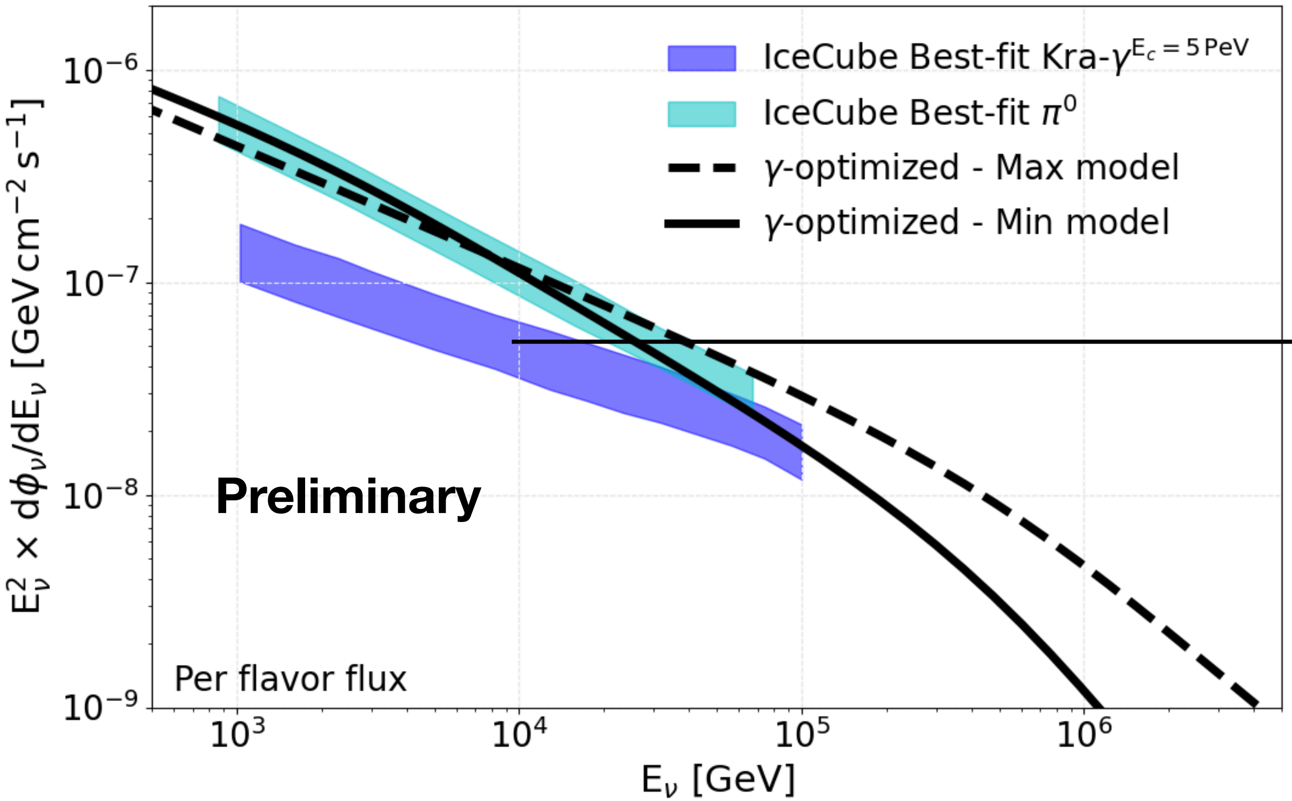
Even though this extrapolation can partially fit the Tibet results, will completely miss the LHAASO observations in the same energy range.

The best fit of the π_0 model is incompatible with gamma-ray observations at low and high energy



ν expectations from the updated KRA-gamma

The expected new full sky ν SED in comparison with IceCube



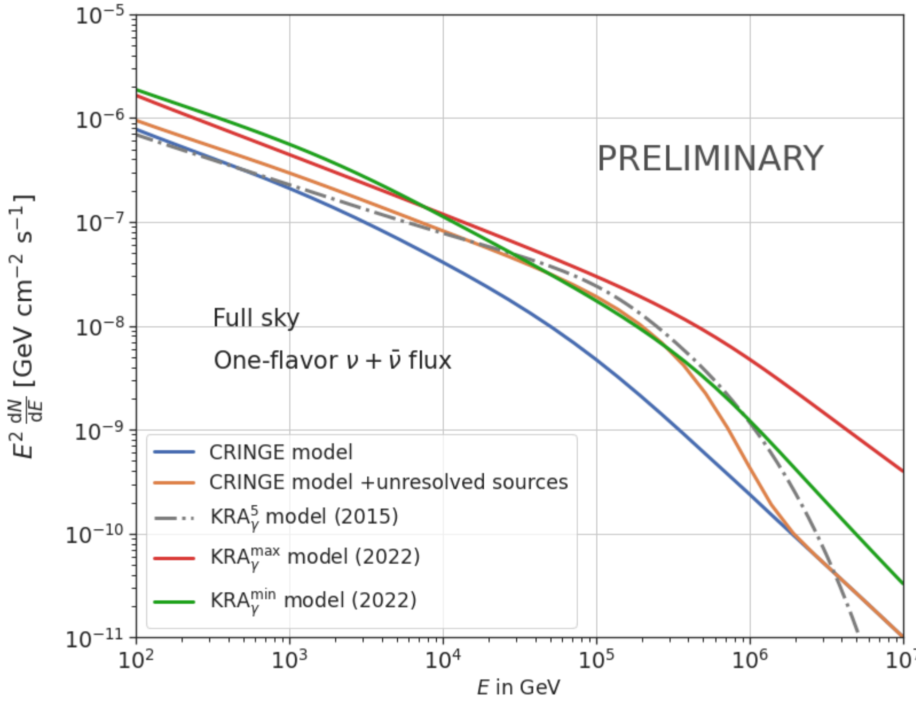
Remember that the KRA-γ 5 best-fit is 0.55 the model expectations

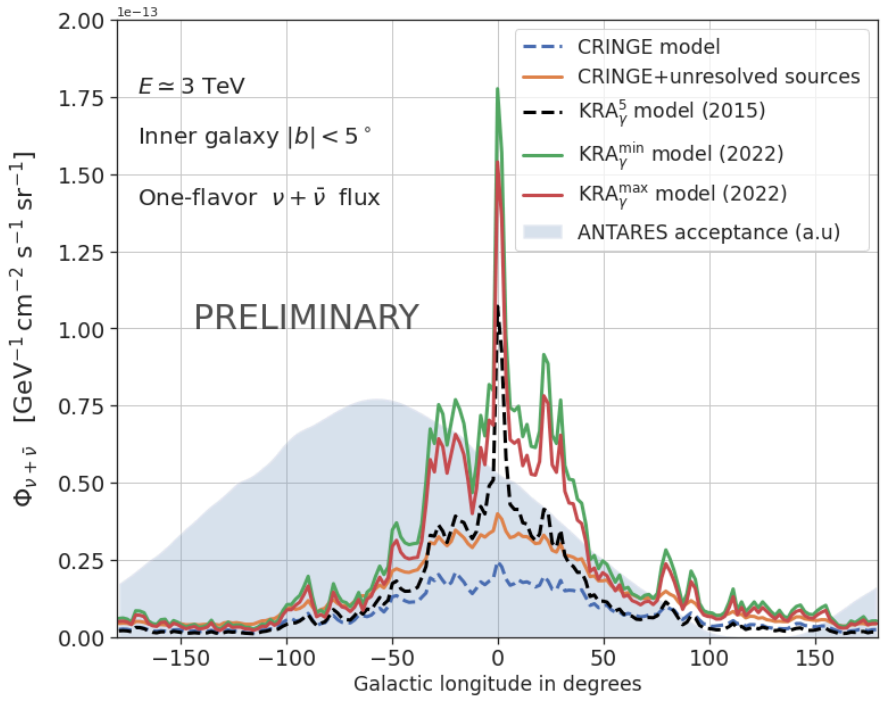
π_0 neutrino best fit and the new expectations from MIN and MAX models suggest that the Fermi-LAT spatial template can agree with diffuse γ-ray and ν observations only if an hardening of the CR toward the Galactic center is assumed ($D \propto E^{\delta(R)}$).

Template fitting of the new KRA-γ with ANTARES

The new template fitting of ANTARES using track-like and shower-like events from 2007 to 2022, 4541 days of live time

Cartraud T. For the Antares coll. ICRC 2023

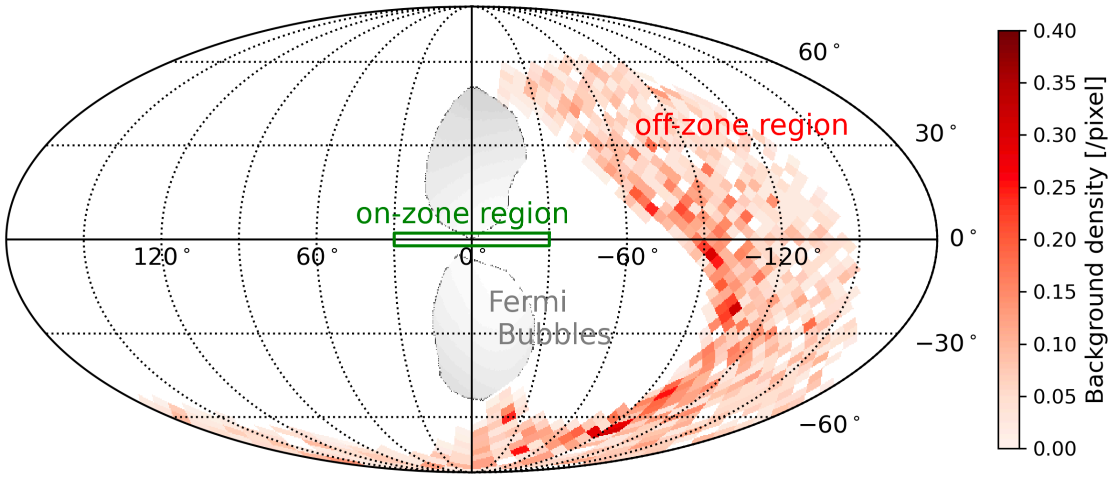


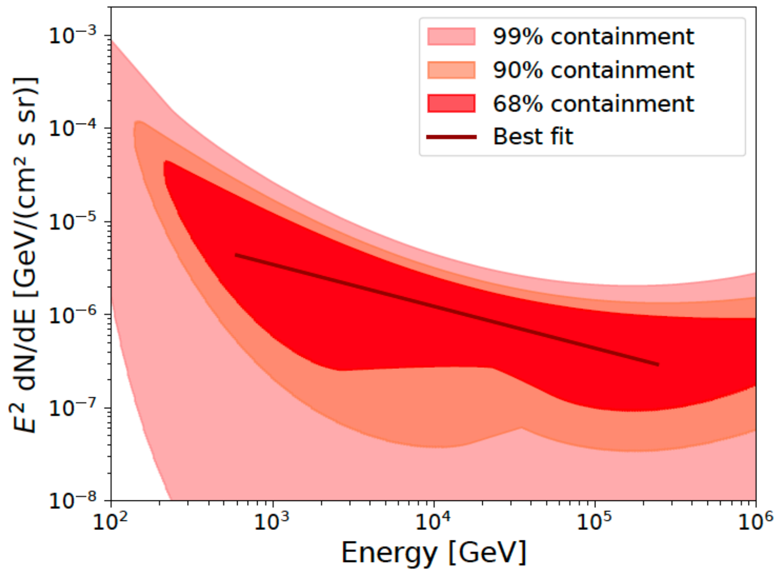


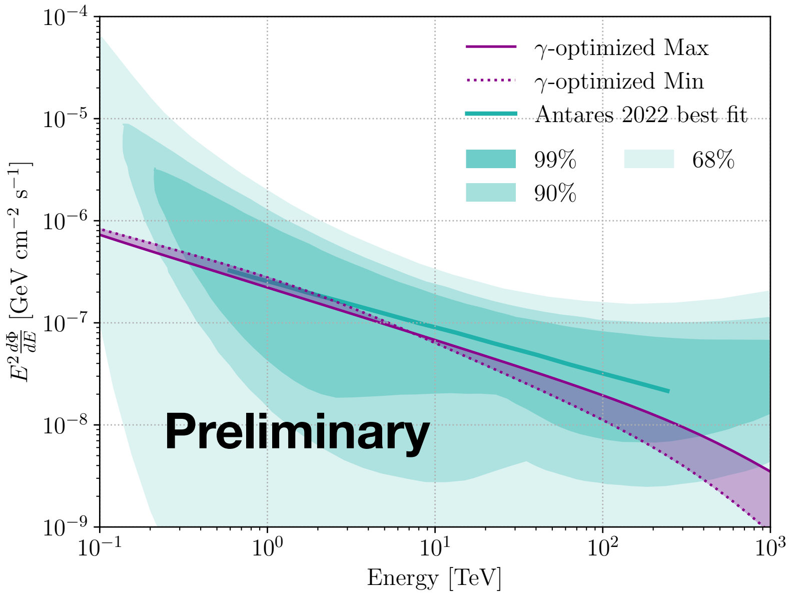
The good acceptance of ANTARES experiment for the central part of our Galaxy, makes is answer a crucial probe of the neutrino flux arriving for this region of the sky.

Comparison with ANTARES ridge observation

Phys Letter B 2023
Antares Coll.





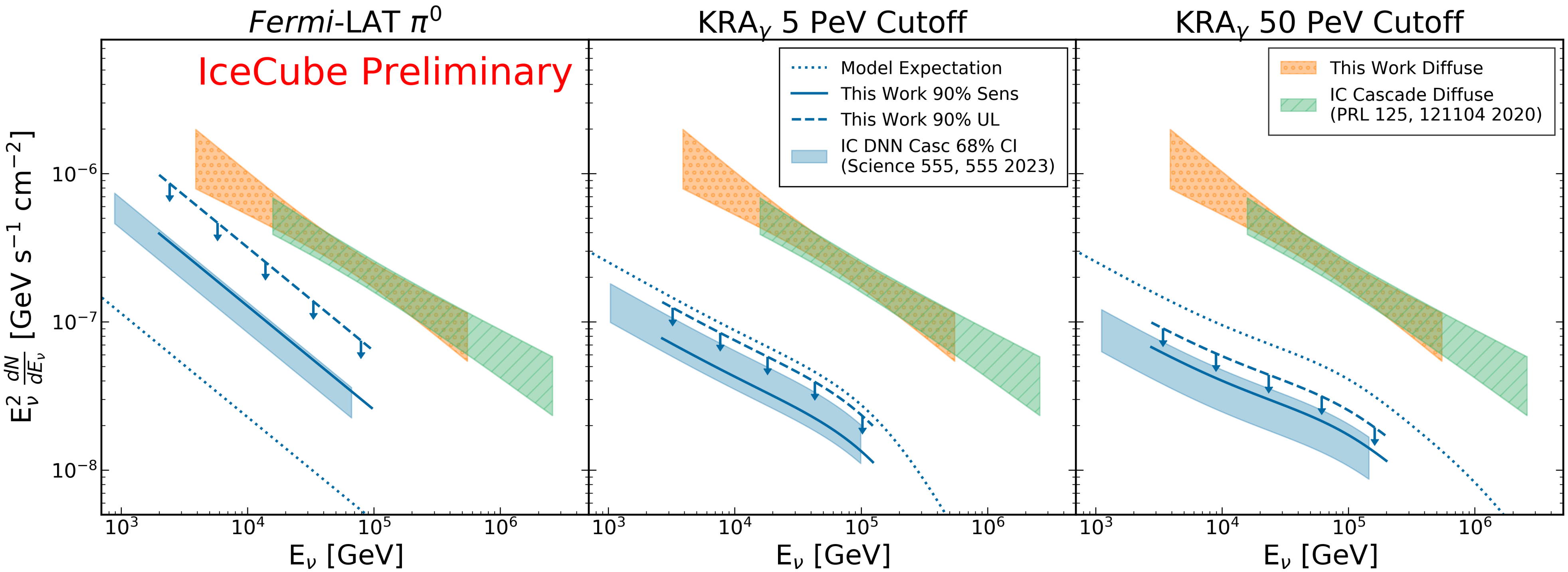


The Updated KRA_γ model expectations for the Galactic ridge (l< 30°, b<2°) results to be in agreement with the Antares confidence regions

IceCube analysis with starting tracks 2008 -2018

Silva, Mancina IceCube coll. POS(ICRC2023) 1008

Mancina talk TeVpa23

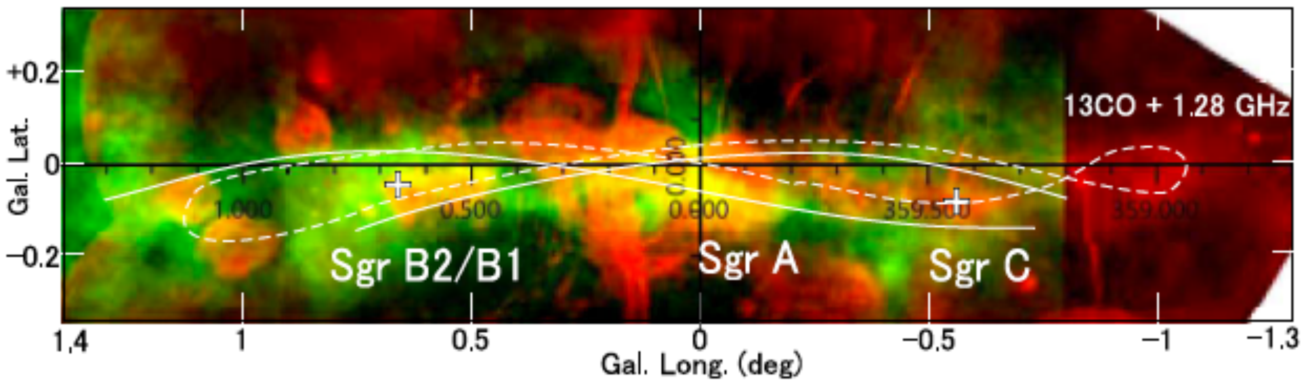


Starting track events IceCube analysis compatible with Cascade analysis, however any significant excess visible, KRA-γ with 50 PeV cutoff quite constrained.

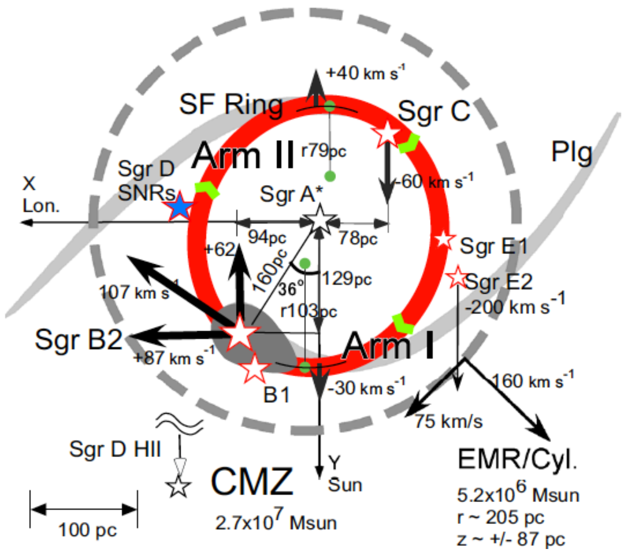
The central molecular zone

Contain 5% of the total mass of Milky Way

MNRAS 2022
Yoshiaki S.



CMZ	CHZ
207 pc radius	320 pc radius
28 pc high	70 pc high



The Central Molecular Zone will be the Gold region to test the cosmic-ray sea physics through neutrinos, already done in the past with HE (Fermi_LAT) and VHE gamma rays (HESS) (Gaggero et al. PRL 2017)

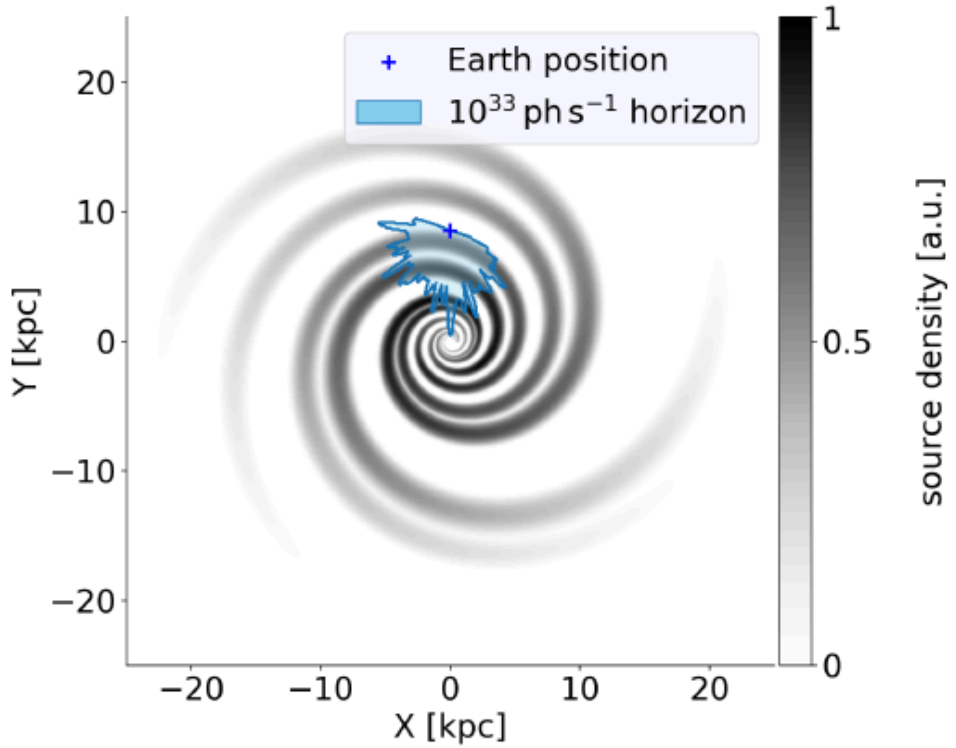
The region where more gas is concentrated and where the spectral assumptions of different models have the large discrepancy.

With the incoming KM3NeT telescope, can this extended region solve the puzzle diffuse + unresolved sources ?

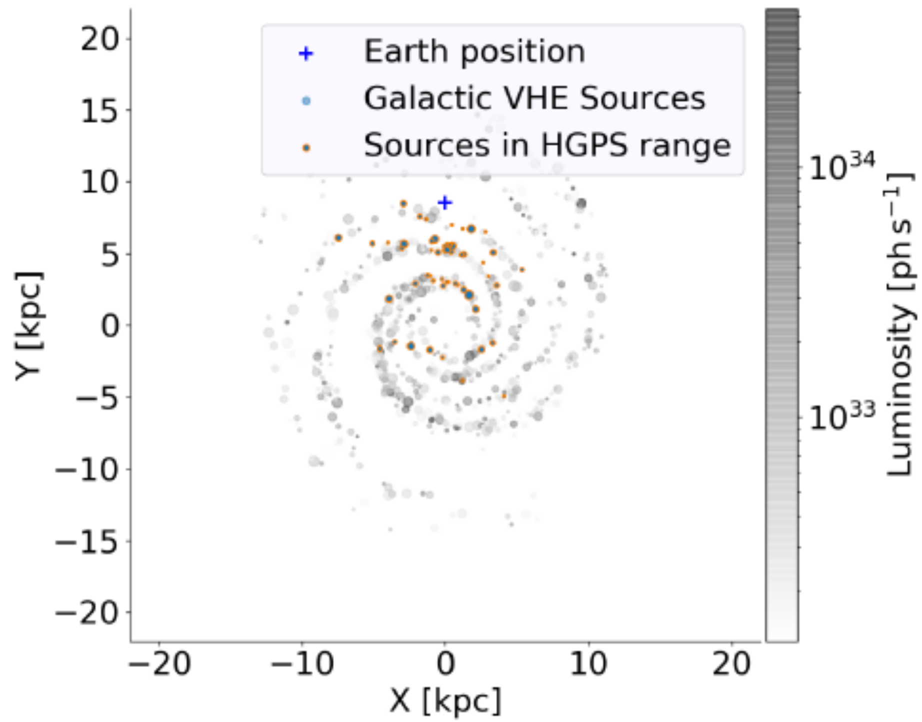
Considering unresolved source contribution

Steppa,Egbert
A&A 2023

Resolved source horizon for HESS



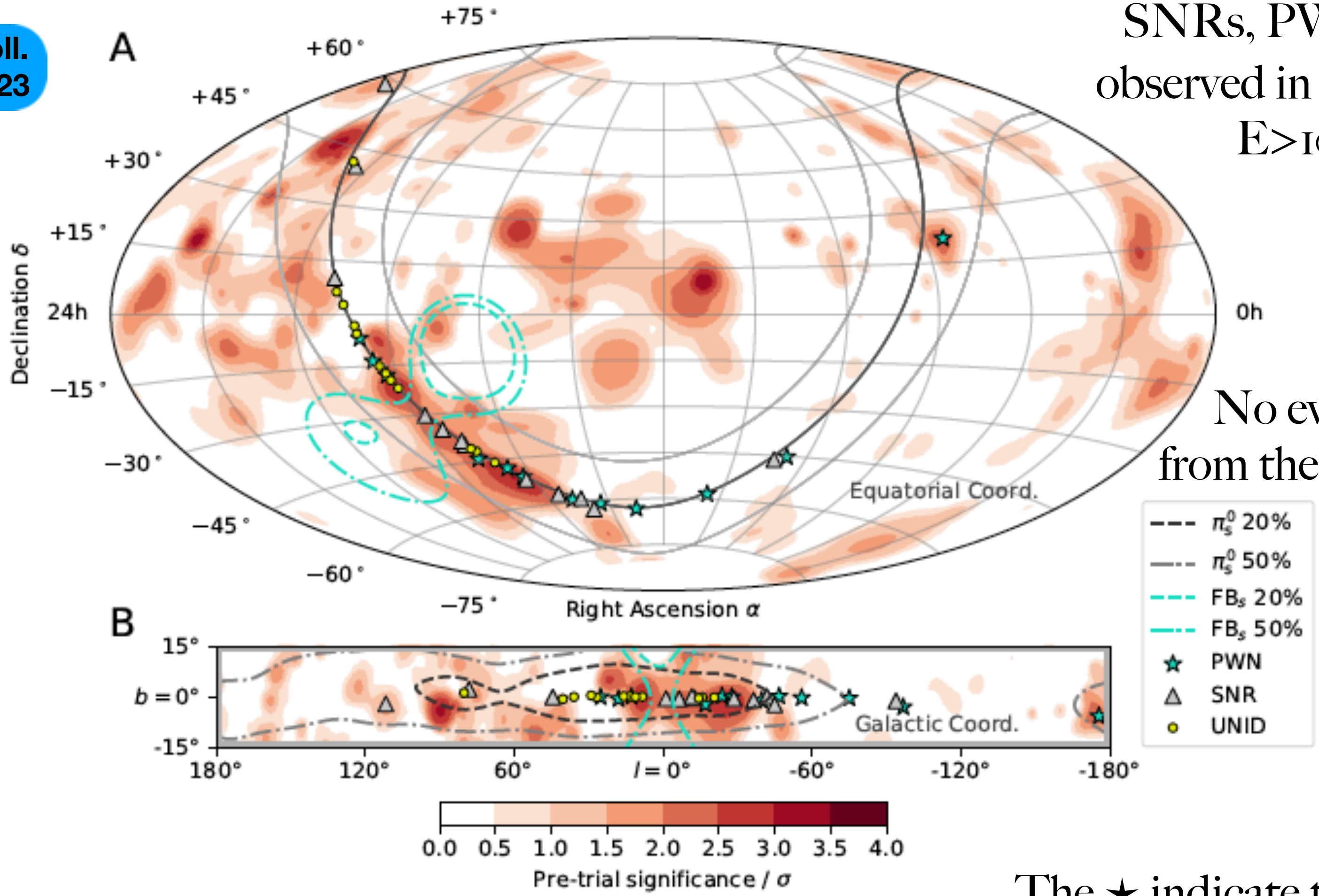
Realization of a synthetic population of VHE gamma-ray sources



Depending on the model, the HGPS sample accounts for (68–87)% of the emission of the population in the scanned region. This suggests that unresolved sources represent a critical component of the diffuse emission measurable in the HGPS. This extra component is taken into account to tune the Min and Max diffuse models. <u>Unresolved source component strongly dependent from the energy considered and from the experiment used.</u>

Search for possible known Galactic ν sources

IceCube coll. Science 2023



SNRs, PWNs, UNIDs observed in γ with energies E>100 GeV

No evidence of ν excess from the selected populations

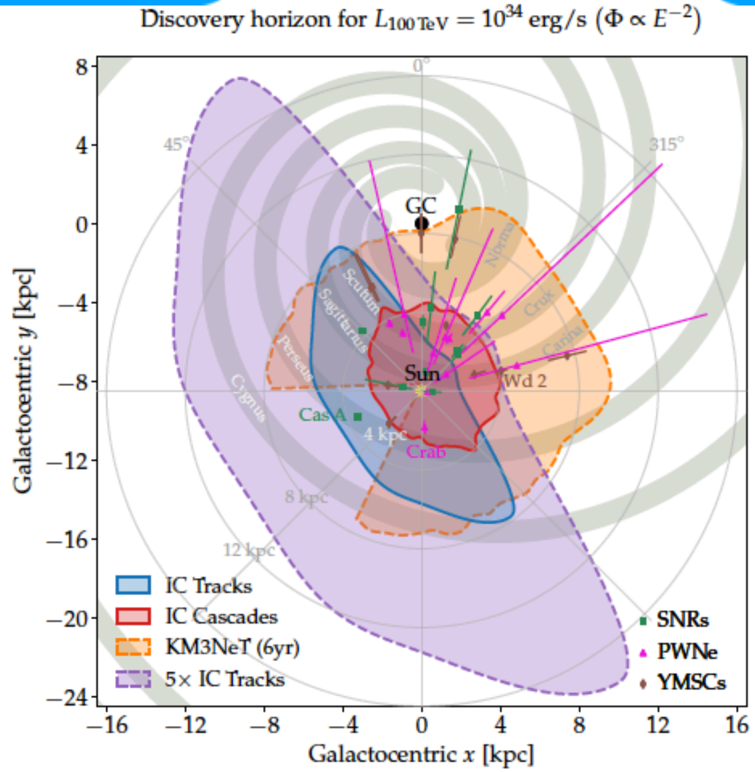
Catalog stacking analyses	p-value
SNR	5.90×10^{-4} (3.24σ)*
PWN	5.93×10^{-4} (3.24σ)*
UNID	3.39×10^{-4} (3.40σ)*

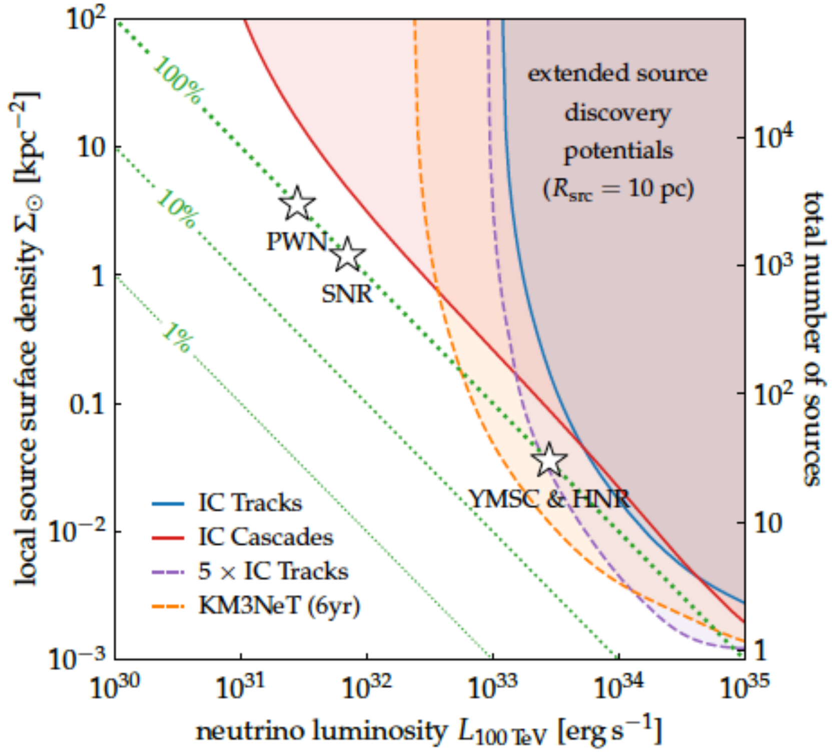
The ★ indicate that the significance is consistent with the diffuse emission template searches

Possible unresolved Galactic ν sources

Ambrosone et al. arXiv 2306.17285

TEVpa talk of Kathrine Groth





As showed in this work the actual ν telescopes and the incoming ones have a limited capabilities to resolve the known neutrino point-like populations, pointing to a possible additional quasi diffuse ν flux. However we don't know the amount hadronic production still associated to the position of these sources.