Some Remarks on Perturbative Quantum Gravity

Luca Buoninfante

4th International FLAG Workshop:
The Quantum and Gravity
September 10, 2024

Radboud Universiteit

MARIE CURIE ACTIONS

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$$D = 1 + 3$$

$$(- \; + \; + +)$$

$$c = 1 = \hbar$$

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Main messages

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3. Don't be afraid of ghosts!

Main messages

1. Perturbative Quantum Gravity is NOT just EFT of Einstein's General Relativity [Quadratic Gravity]

2. The criterion of 'strict' renormalizability is very predictive also when applied to gravitational interaction

3. Don't be afraid of ghosts! [safely hidden in the second Riemann sheet]

[L.B., arXiv:2410.XXXXX]

Introduction

Einstein's General Relativity (non-renormalizable):

$$S_{EH} = \frac{M_p^2}{2} \int d^4x \sqrt{-g} R$$

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Anologue case: Fermi theory of weak interaction

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Quadratic Gravity (renormalizable in $D = 4$):

[Stelle PRD (1977)]

$$S = \frac{1}{2} \int d^4x \sqrt{-g} \left(M_p^2 R + \frac{\alpha}{6} R^2 - \frac{\beta}{2} C_{\mu\nu\rho\sigma} C^{\mu\nu\rho\sigma} \right)$$

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[Thanks to S. Silveravalle and A. Held for introducing it!]

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Massive spin-0: $m_0^2 = \frac{M_p^2}{\alpha}$

Massive spin-2 ghost: $m_2^2 = \frac{M_p^2}{\beta}$

$$G_{\mu\nu\rho\sigma}(p^2) = \frac{m_2^2 \mathrm{P}^{(2)}_{\mu\nu\rho\sigma}}{p^2(p^2+m_2^2)} - \frac{m_0^2 \mathrm{P}^{(0)}_{\mu\nu\rho\sigma}}{2p^2(p^2+m_0^2)} = G^{(EH)}_{\mu\nu\rho\sigma} + \frac{1}{2}\frac{\mathrm{P}^{(0)}_{\mu\nu\rho\sigma}}{p^2+m_0^2} - \frac{\mathrm{P}^{(2)}_{\mu\nu\rho\sigma}}{p^2+m_2^2}$$

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Massive spin-2 ghost: $m_2^2 = \frac{M_p^2}{\beta}$

That's what I mean with Perturbative Quantum Gravity !

Remark: 'EFT of GR' vs 'Quadratic Gravity'

EFT of Einstein's general relativity (non-renormalizable):

$$S = \int d^4x \sqrt{-g}\left(\frac{M_p^2}{2}R + aR^2 + bR_{\mu\nu}R^{\mu\nu} + \frac{c}{M_p^2}\mathcal{R}^3 + \cdots + \frac{d}{M_p^{2(n-2)}}\mathcal{R}^n + \cdots\right)$$

Propagator: $\sim 1/p^2$

Loop expansion: $\sim \left(\frac{1}{M_p^2}\right)^{L-1}$ and positive powers of $a, b, c, \ldots d, \ldots$

Quadratic Gravity (renormalizable):

$$S = \frac{1}{2}\int d^4x \sqrt{-g}\left(M_p^2 R + \frac{\alpha}{6}R^2 - \frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

Propagator: $\sim 1/p^4$

Loop expansion: $\sim \left(\frac{1}{\alpha}\right)^{L-1}, \quad \left(\frac{1}{\beta}\right)^{L-1}$

Motivations for Quadratic Gravity

$$S = \frac{1}{2}\int d^4x\,\sqrt{-g}\left(M_p^2 R + \frac{\alpha}{6}R^2 - \frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

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- A 'unique' strictly renormalizable QFT of gravity in $D = 4$
- New physics in the sub-Planckian regime [unlike EFT of GR]:

$$m_0 = \frac{M_p}{\sqrt{\alpha}}, \qquad m_2 = \frac{M_p}{\sqrt{\beta}}, \qquad \alpha, \beta \gg 1 \quad \Longrightarrow \quad m_0, m_2 \ll M_p$$

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- Inflation for free: $\alpha \sim 10^{10}$ ($m_0 \sim 10^{13}$GeV) **[Starobinsky 1980+]**
- Future measurement of tensor-to-scalar ratio can constrain β
  [Ivanov, Tokareva 2016; Salvio 2017+; Anselmi, Bianchi, Piva 2019+]

Quadratic Gravity

$$S = \frac{1}{2}\int d^4x\,\sqrt{-g}\left(M_p^2 R + \frac{\alpha}{6}R^2 - \frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

What about the spin-2 ghost?

$$G^{(2)}_{\mu\nu\rho\sigma}(p^2) = \mathrm{P}^{(2)}_{\mu\nu\rho\sigma}\frac{m_2^2}{p^2(p^2+m_2^2)} = \mathrm{P}^{(2)}_{\mu\nu\rho\sigma}\left[\frac{1}{p^2} - \frac{1}{p^2+m_2^2}\right]$$

Remarks on the ghost

Scalar toy model:

ghost

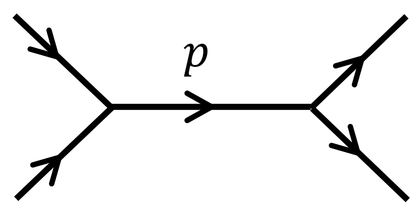
$$\mathcal{L} = \frac{1}{2}\phi\Box\left(1 - \frac{\Box}{m^2}\right)\phi - V(\phi) \quad \Longrightarrow \quad G(p^2) = \frac{1}{p^2} - \frac{1}{p^2 + m^2}$$

Optical theorem:

$$S^{+}S = 1, \qquad S = 1 + iT,$$

$$1 = \sum_{\{n\}} c_n \, |n\rangle\langle n| \qquad \Longrightarrow \qquad 2Im\{\langle\psi|T|\psi\rangle\} = \sum_{\{n\}} c_n |\langle n|T|\psi\rangle|^2$$

Tree-level example ($V \sim \phi^3$):



Remarks on the ghost

$$G(p^2) = \frac{1}{p^2 - i\epsilon} - \frac{1}{p^2 + m^2 - ib\varepsilon}$$

$$(\epsilon > 0,\ \varepsilon > 0) \qquad b = \pm 1$$

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Tree level example:

$$Im\{\langle\psi|T|\psi\rangle\} \sim \pi \, [\delta(p^2) - sign(b)\delta(p^2 + m^2)]$$

Remarks on the ghost

Causal propagation & positive norms

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$$S^{+}S = 1, \qquad S = 1 + iT,$$

$$1 = \sum_{\{n\}} c_n\, |n\rangle\langle n|, \qquad c_n^{normal} > 0$$

$$c_n^{ghost} > 0$$

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$$2Im\{\langle\psi|T|\psi\rangle\} = \sum_{\{n\}} c_n |\langle n|T|\psi\rangle|^2 \geq 0$$

Tree level example:

$$Im\{\langle\psi|T|\psi\rangle\} \sim \pi\,[\delta(p^2) - \delta(p^2 + m^2)]$$

Unitarity is violated!

Remarks on the ghost

1) Causal propagation & negative norms

[Lee, Wick, Salvio, Strumia, Holdom…]

$$G(p^2) = \frac{1}{p^2 - i\epsilon} - \frac{1}{p^2 + m^2 - i\varepsilon}$$

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$$c_n^{ghost} < 0$$

Optical theorem:

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Tree level example:

$$Im\{\langle\psi|T|\psi\rangle\} \sim \pi\,[\delta(p^2) - \delta(p^2 + m^2)]$$

Unitarity is preserved!

Remarks on the ghost

2) Acausal propagation & positive norms

[Donoghue, Menezes,...]

$$G(p^2) = \frac{1}{p^2 - i\epsilon} - \frac{1}{p^2 + m^2 + i\varepsilon}$$

$$(\epsilon > 0,\ \varepsilon > 0) \qquad b = -1$$

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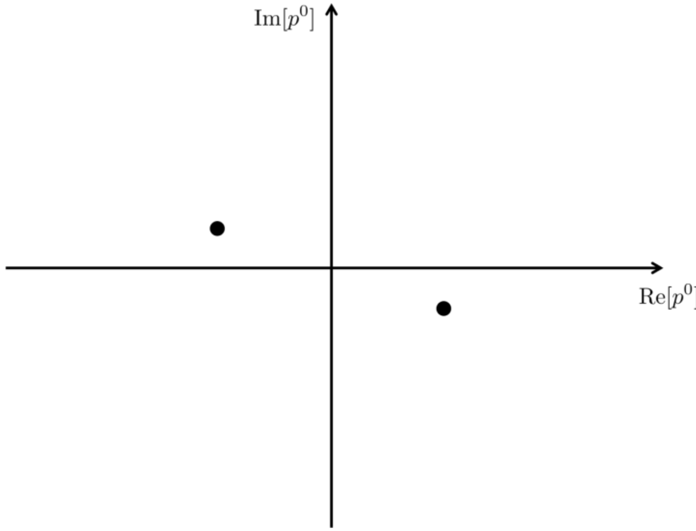
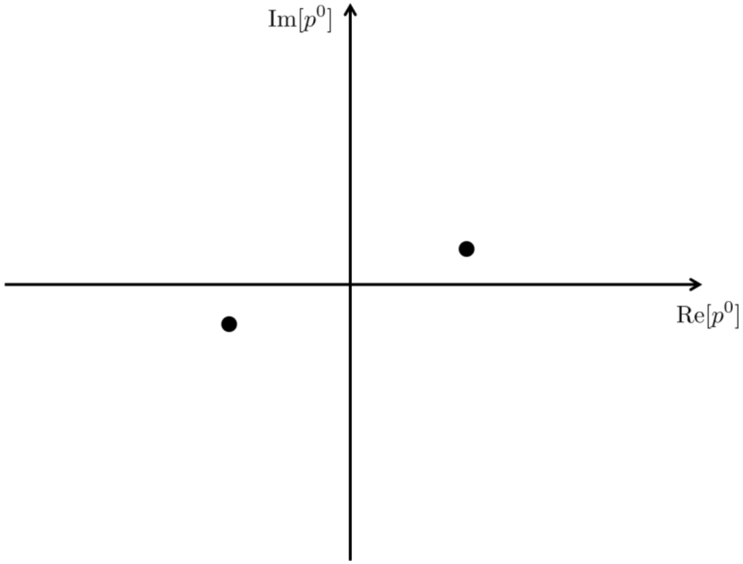
Remarks on the ghost

1) Causal propagation & negative norms
   Feynman ghost



2) Acausal propagation & positive norms
   Anti-Feynman ghost



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Is the spin-2 massive ghost part of the asymptotic states?

If YES: stabilities might still be there…

If NO: stabilities could be avoided!

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Spin-2 ghost gets a width through quantum effects:

[Tomboulis 1977-1986; Salvio, Strumia et al. 2017+, Anselmi, Piva 2018+; Donoghue, Menezes 2018+]

$$G^{(2,gh)}_{\mu\nu\rho\sigma}(p^2) = \mathrm{P}^{(2)}_{\mu\nu\rho\sigma}\frac{-1}{p^2+m_2^2+\Sigma(p^2)}, \qquad \Sigma(p^2) = Re[\Sigma(p^2)] + i\, Im[\Sigma(p^2)]$$

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Narrow width approximation ($\Gamma \ll m_2$):

$$\bar{G}^{(2,gh)}_{\mu\nu\rho\sigma}(s) \sim \mathrm{P}^{(2)}_{\mu\nu\rho\sigma}\frac{-1}{-s+m_2^2+i\, m_2\Gamma} \qquad s \equiv -p^2, \quad \Gamma \equiv \frac{Im[\Sigma(m_2^2)]}{m_2} > 0$$

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Ghost as an unstable resonance? If YES, Veltman projection works!

[Veltman 1963]

Ghost as an unstable resonance?

A recent critique [Kubo & Kugo, arXiv:2402.15956]

Toy model:

$$\mathcal{L} = \frac{1}{2}\phi(\Box - \mu^2)\phi + a\frac{1}{2}\chi(\Box - m_0^2)\chi - g\chi\phi^2 + \cdots, \qquad m \geq 2\mu$$

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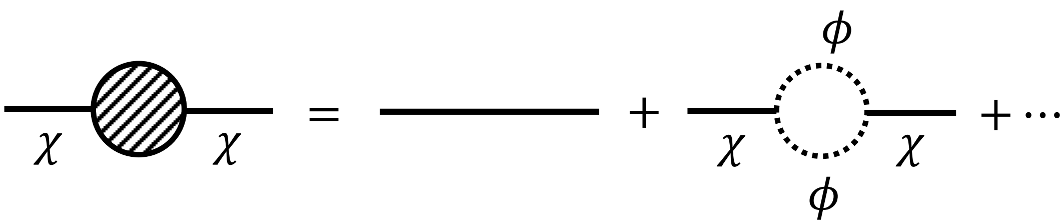
Resummed χ-propagator:



$$\bar{G}_\chi(s) = \frac{a}{-s + m_0^2 - a\Sigma(s)} \quad \longrightarrow \quad \frac{a}{-s + m^2 - i\, a\, m\Gamma}$$

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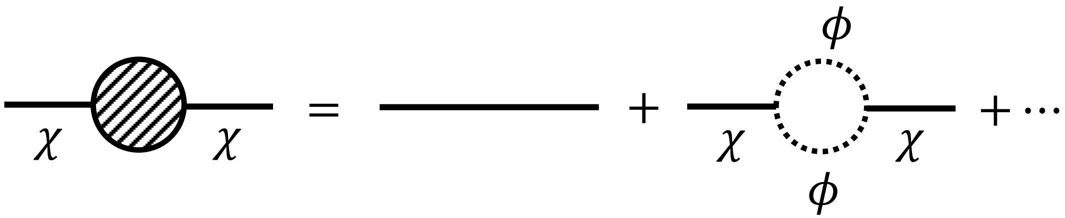

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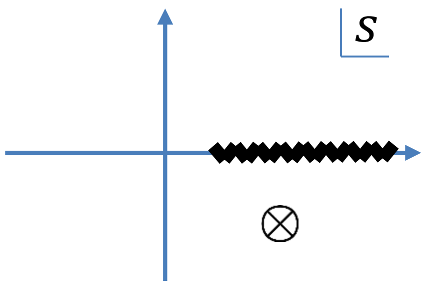
- Standard case (no ghost): $a = 1$, $s = m^2 - i\, m\Gamma$

  The pole appears in the second Riemann sheet



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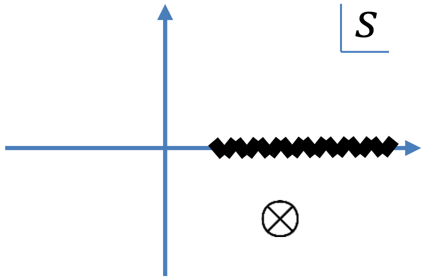

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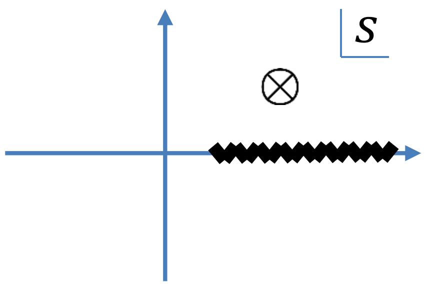
- Case with ghost: $a = -1,\quad s = m^2 + i\,m\Gamma$



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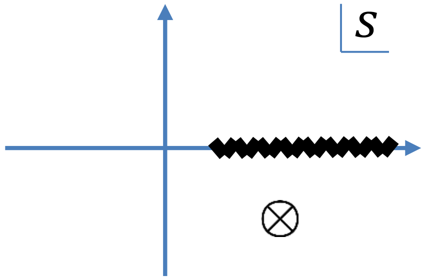

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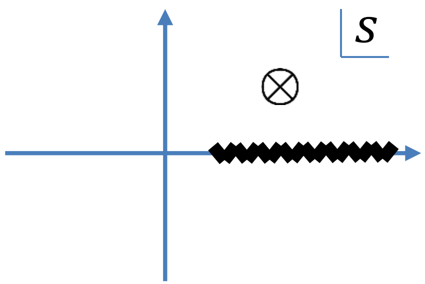

Question: does the ghost pole appear in the first or second Riemann sheet?

Ghost as an unstable resonance?

[L.B., arXiv:2410.XXXXX]

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Pole: $s = m^2 - i\, a\, m\Gamma$

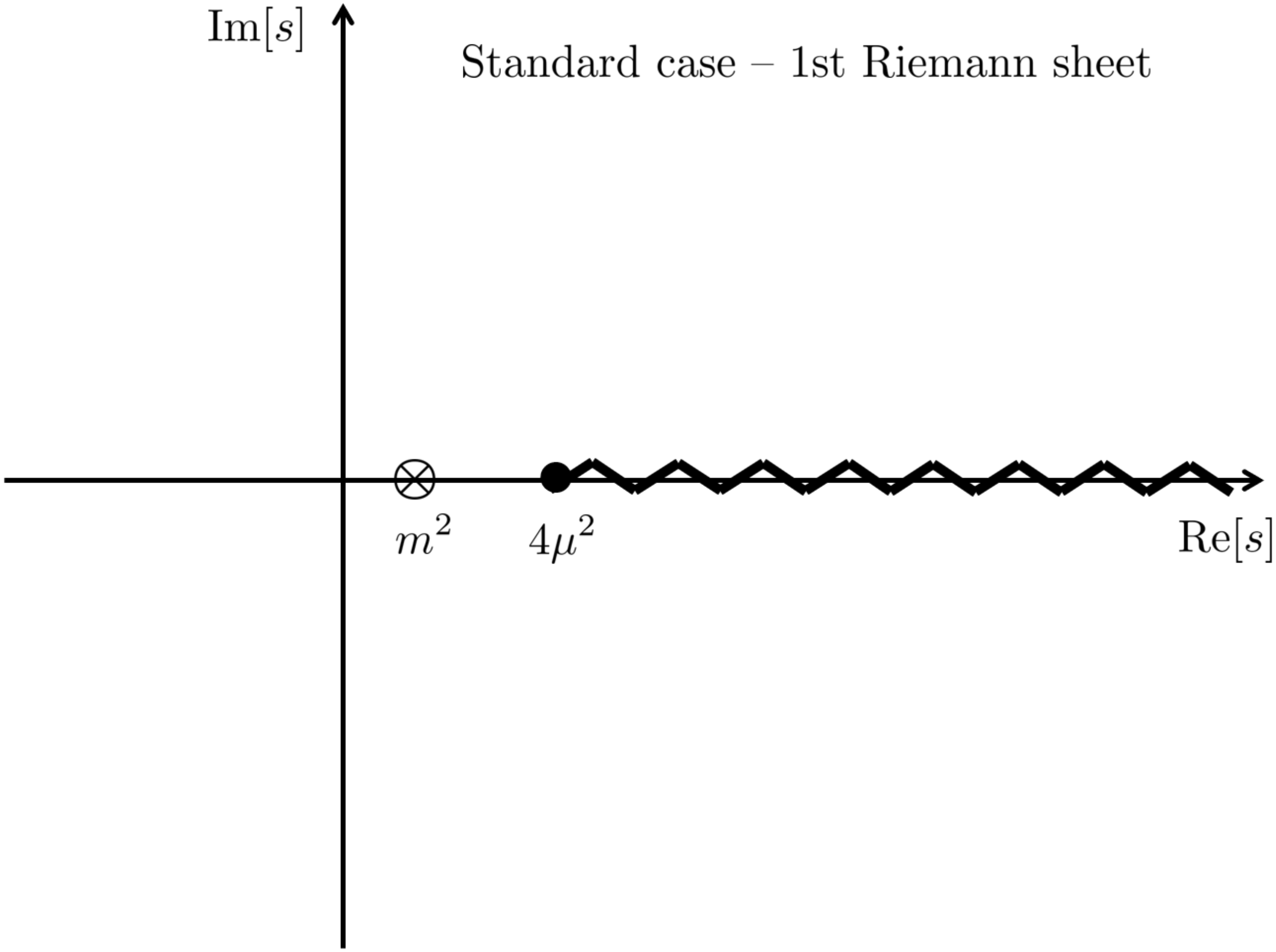
Question: does the ghost pole appear in the first or second Riemann sheet?

$$Im\,[\Sigma(s)] = \frac{g^2}{16\pi}\sqrt{1 - \frac{4\mu^2}{s}},$$

$$-s + m^2 = a\, i\, \frac{g^2}{16\pi}\sqrt{1 - \frac{4\mu^2}{s}}, \qquad s = m^2 e^{i\theta}$$

Ghost: stable or unstable?

Standard case (no ghost): $\bar{G}_\chi(s) = \lim_{\varepsilon \to 0^+} \bar{G}_\chi(s + i\varepsilon), \quad s \in \mathbb{R}$



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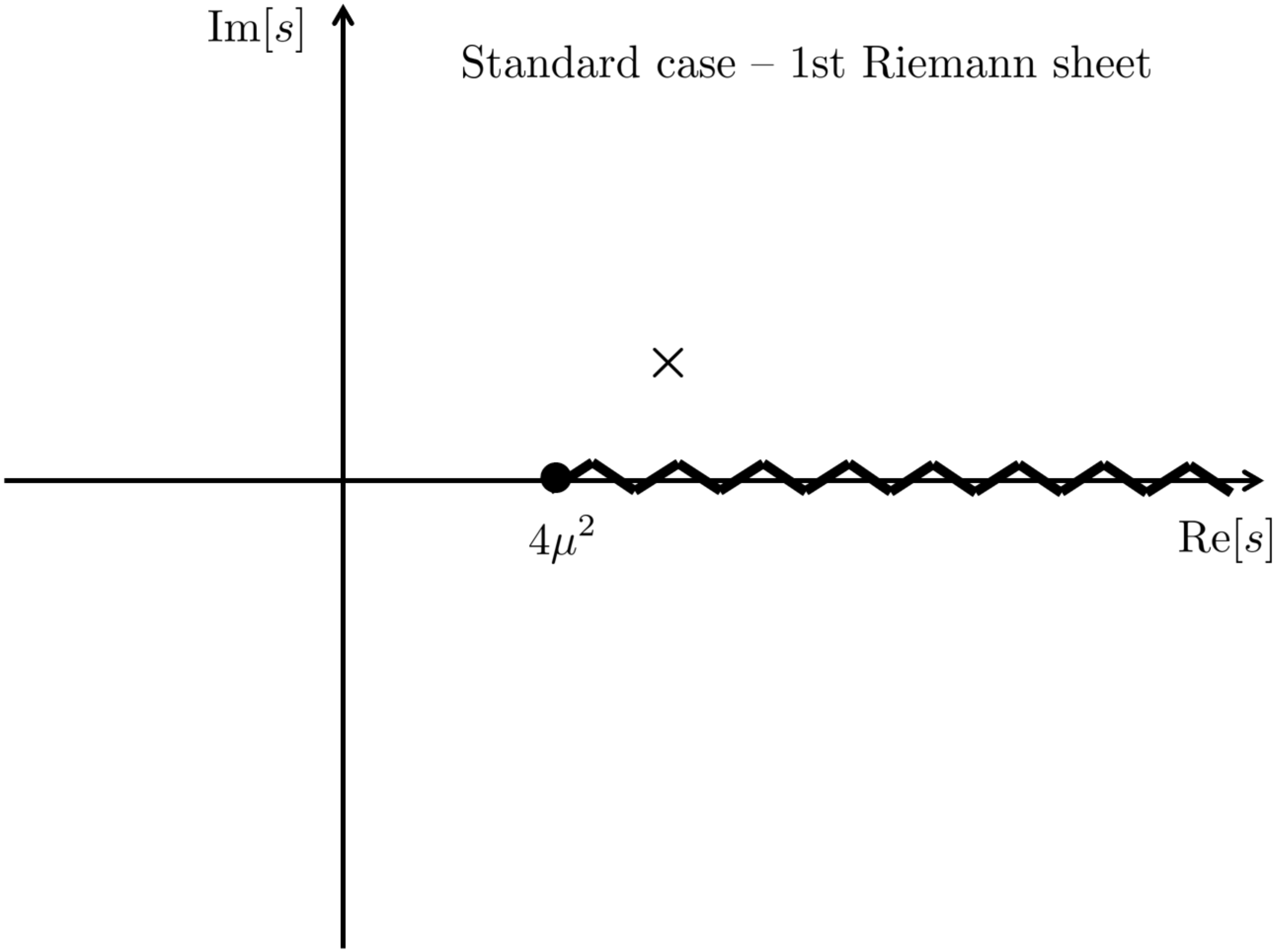

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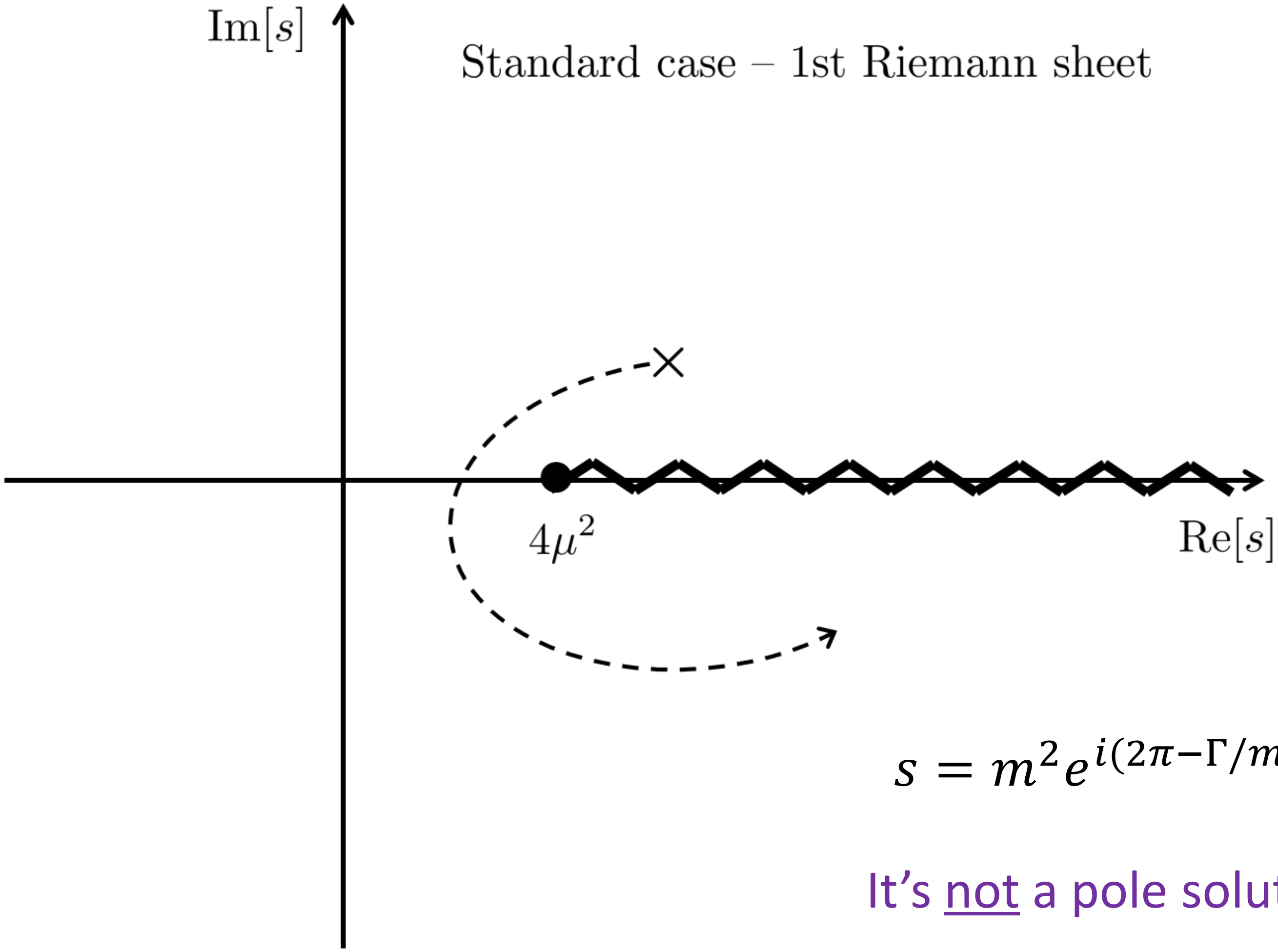

$$s = m^2 e^{i(2\pi - \Gamma/m)}$$

It's not a pole solution!

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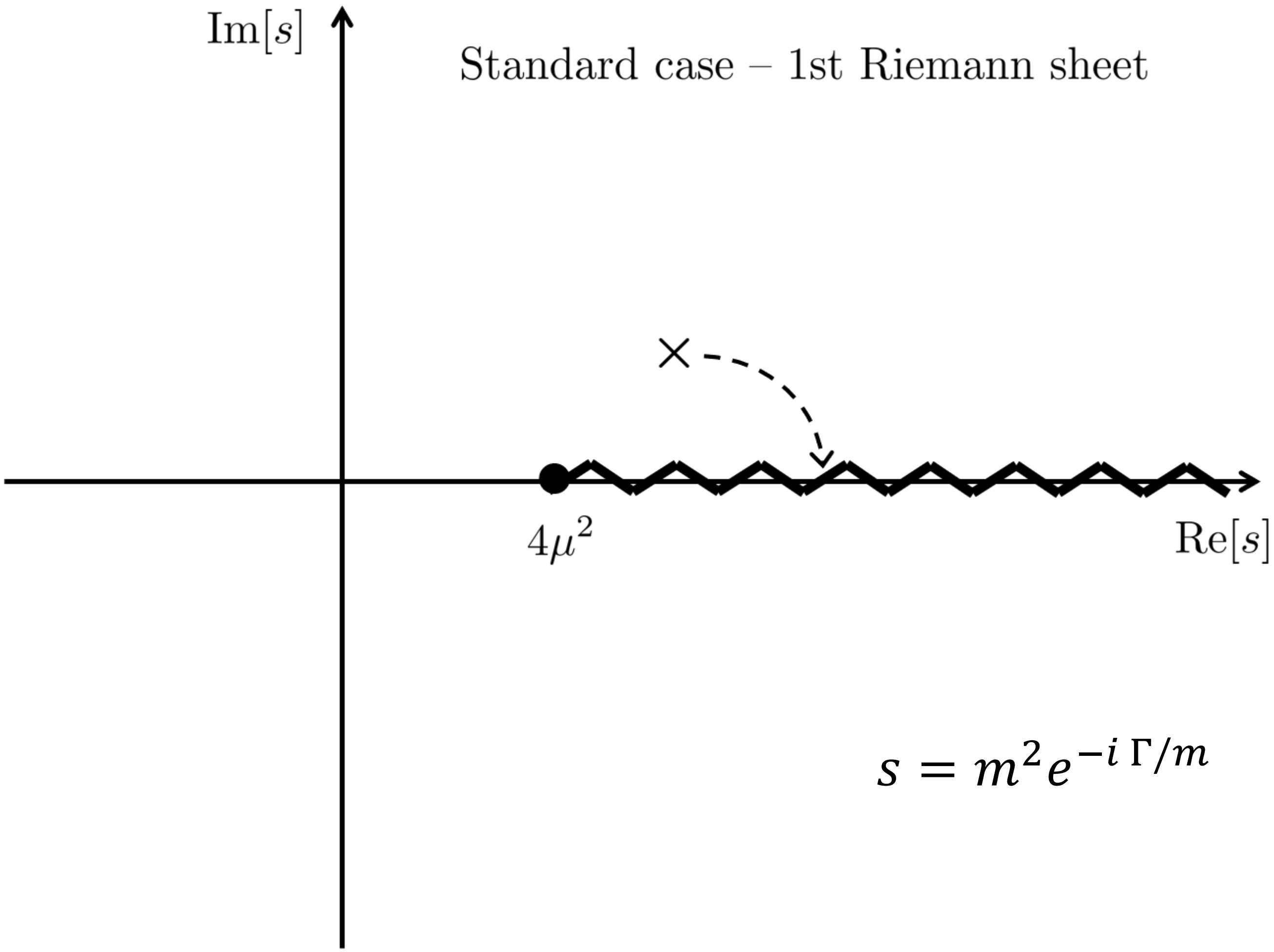

$$s = m^2 e^{-i\,\Gamma/m}$$

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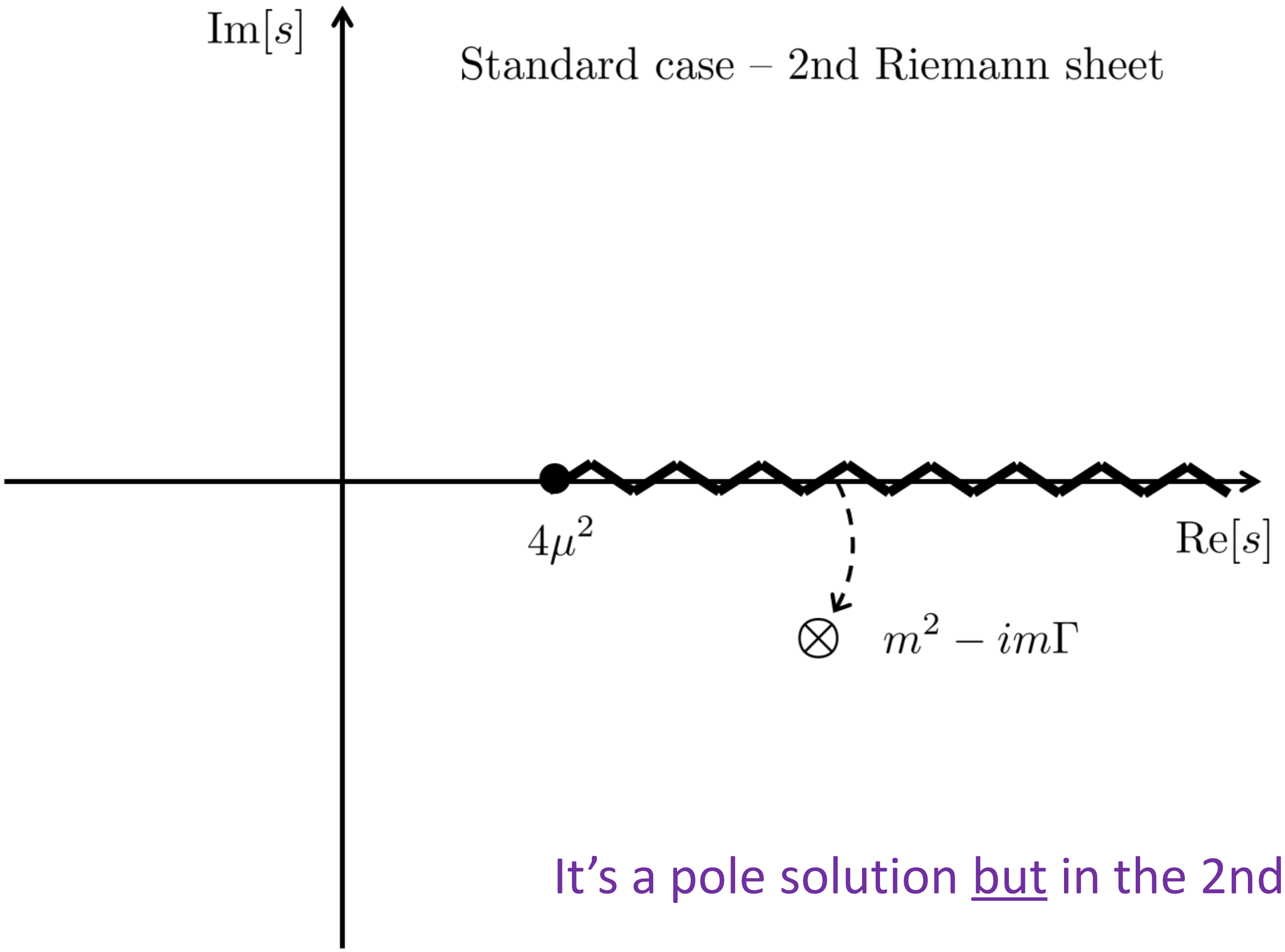
Standard case (no ghost): $\qquad \bar{G}_\chi(s) = \lim_{\varepsilon \to 0^+} \bar{G}_\chi(s + i\varepsilon), \quad s \in \mathbb{R}$



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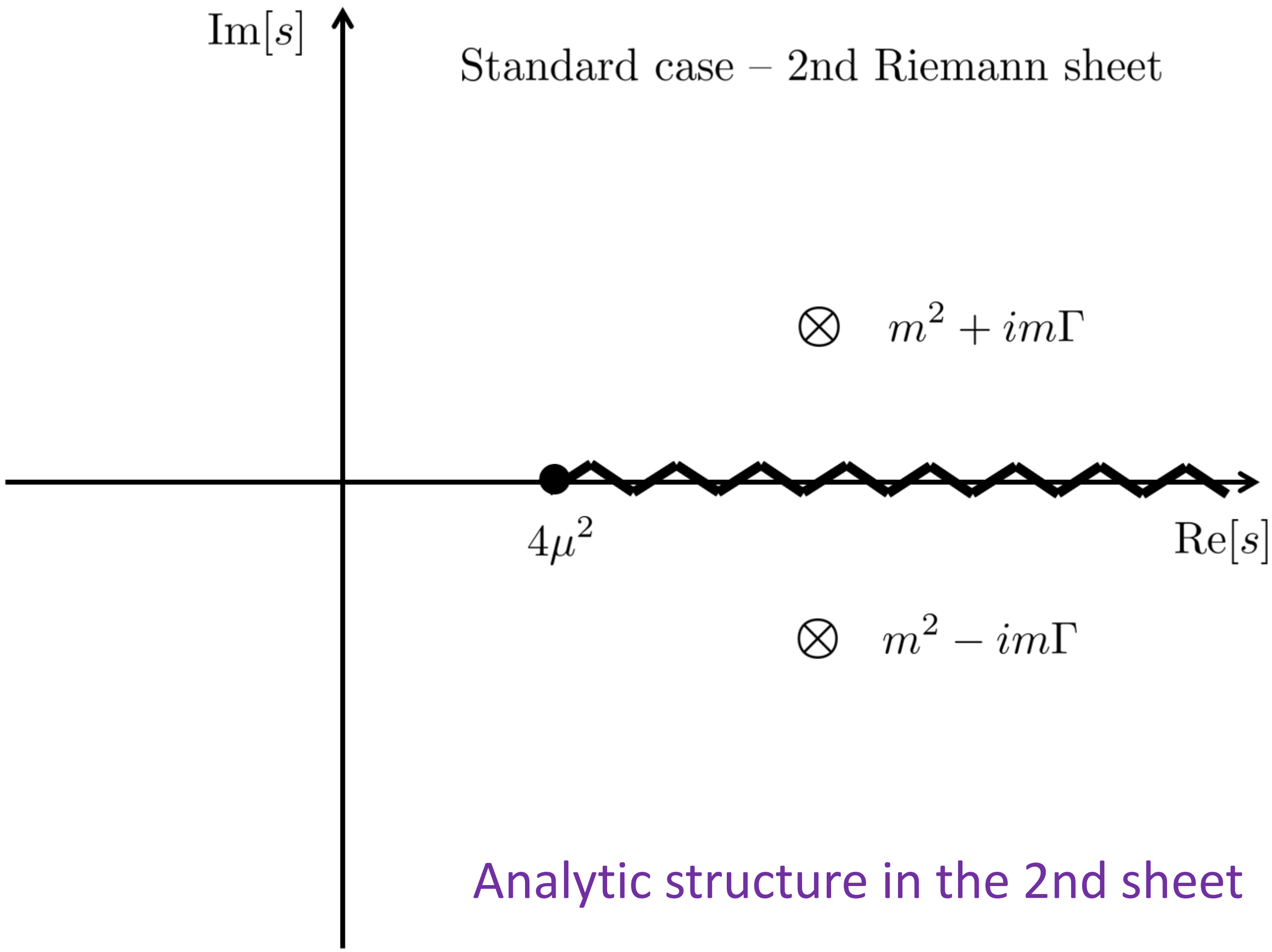

Analytic structure in the 2nd sheet

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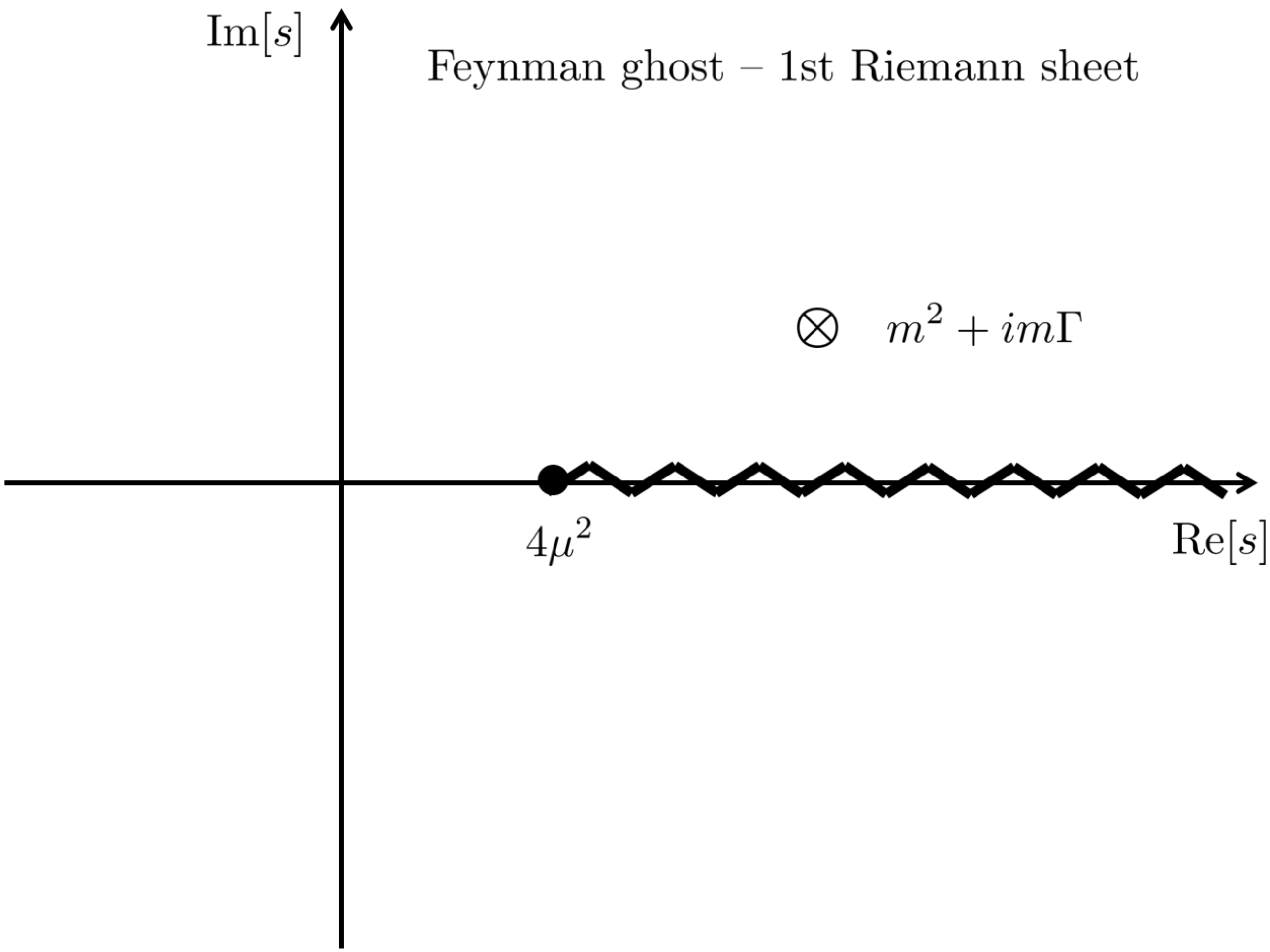

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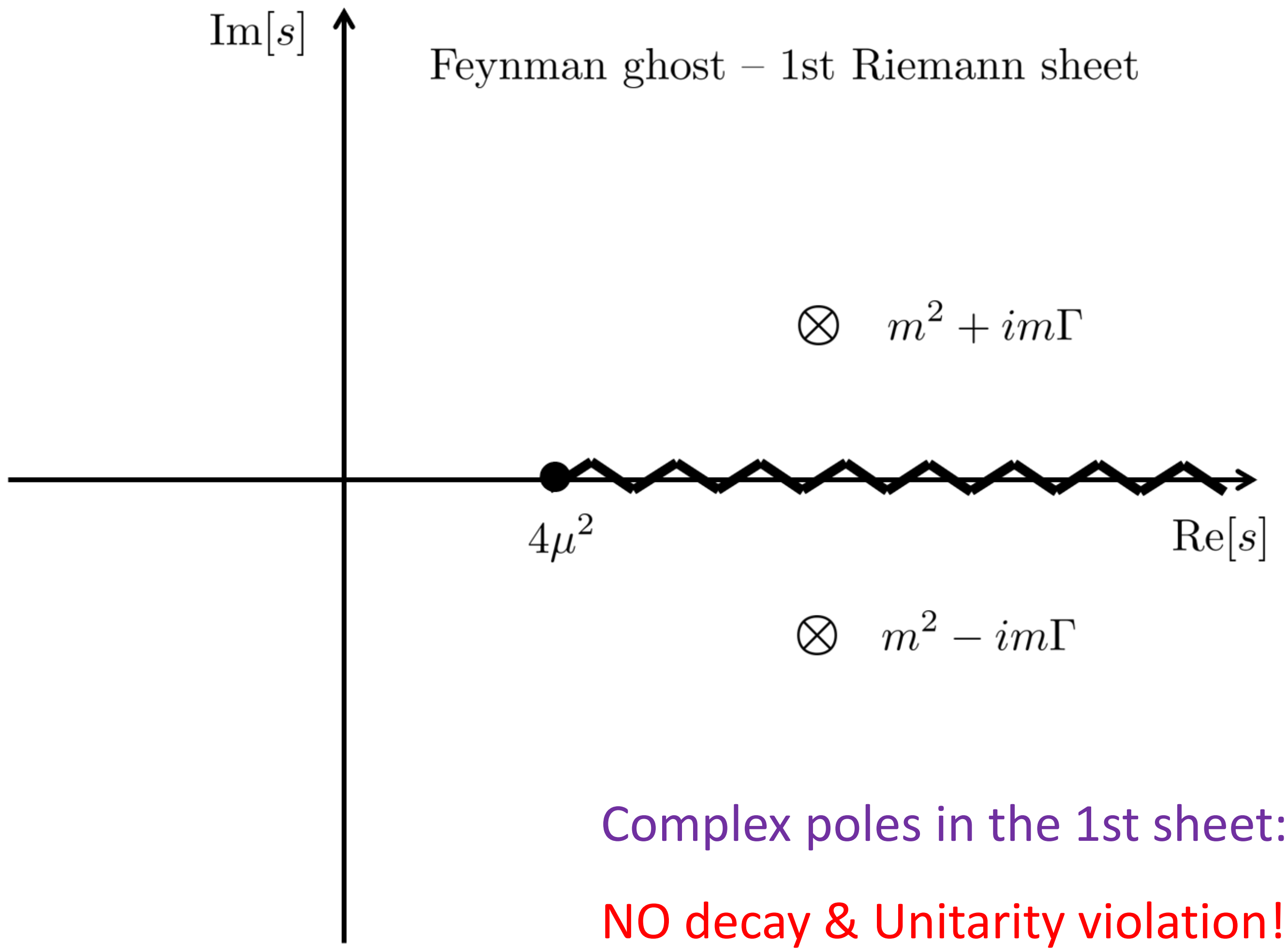


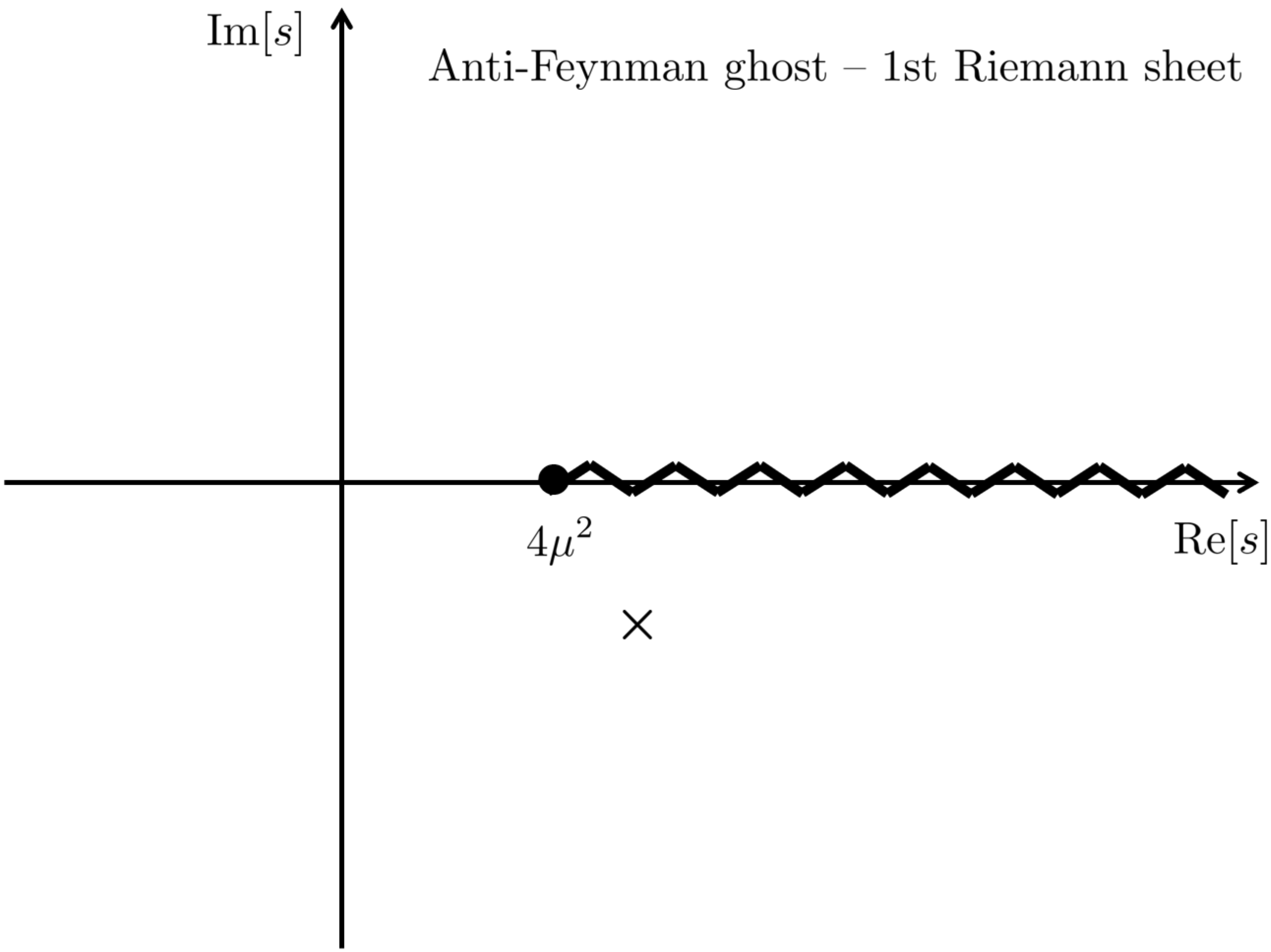
Complex poles in the 1st sheet:

NO decay & Unitarity violation!

Ghost: stable or unstable?

Anti-Feynman ghost:

$$\bar{G}_\chi(s) = \lim_{\varepsilon \to 0^+} \bar{G}_\chi(s - i\varepsilon), \quad s \in \mathbb{R}$$



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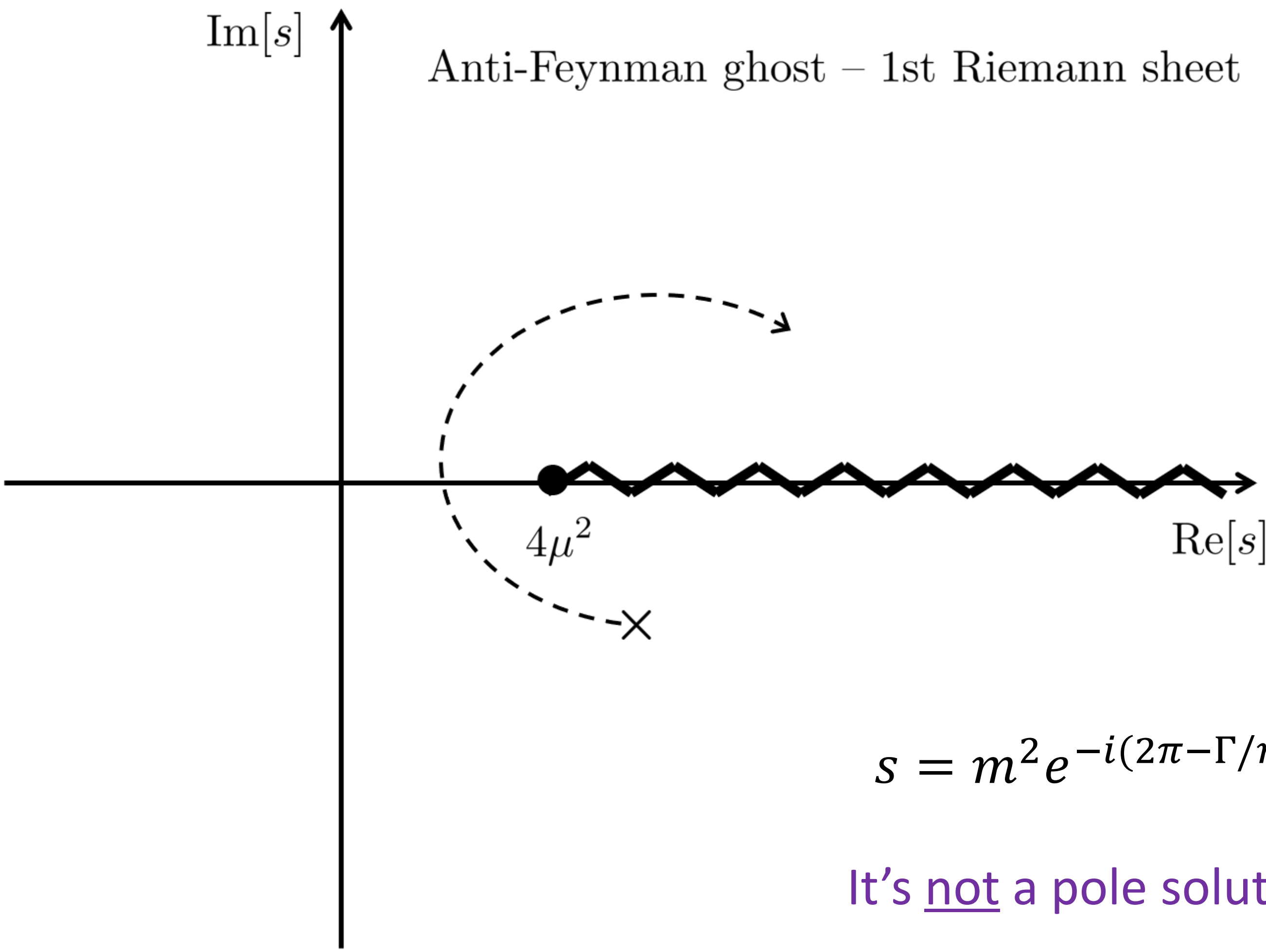

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It's <u>not</u> a pole solution!

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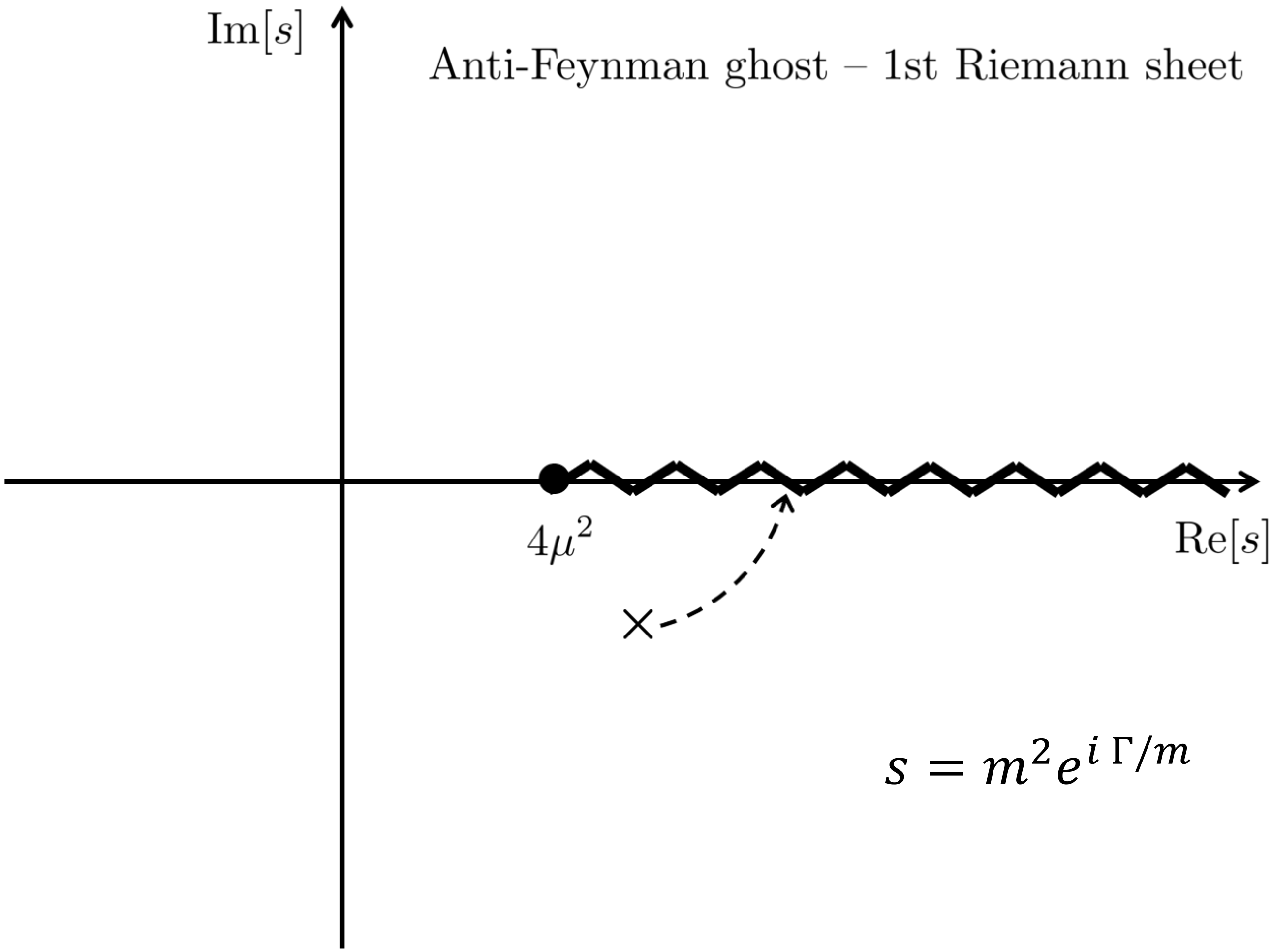

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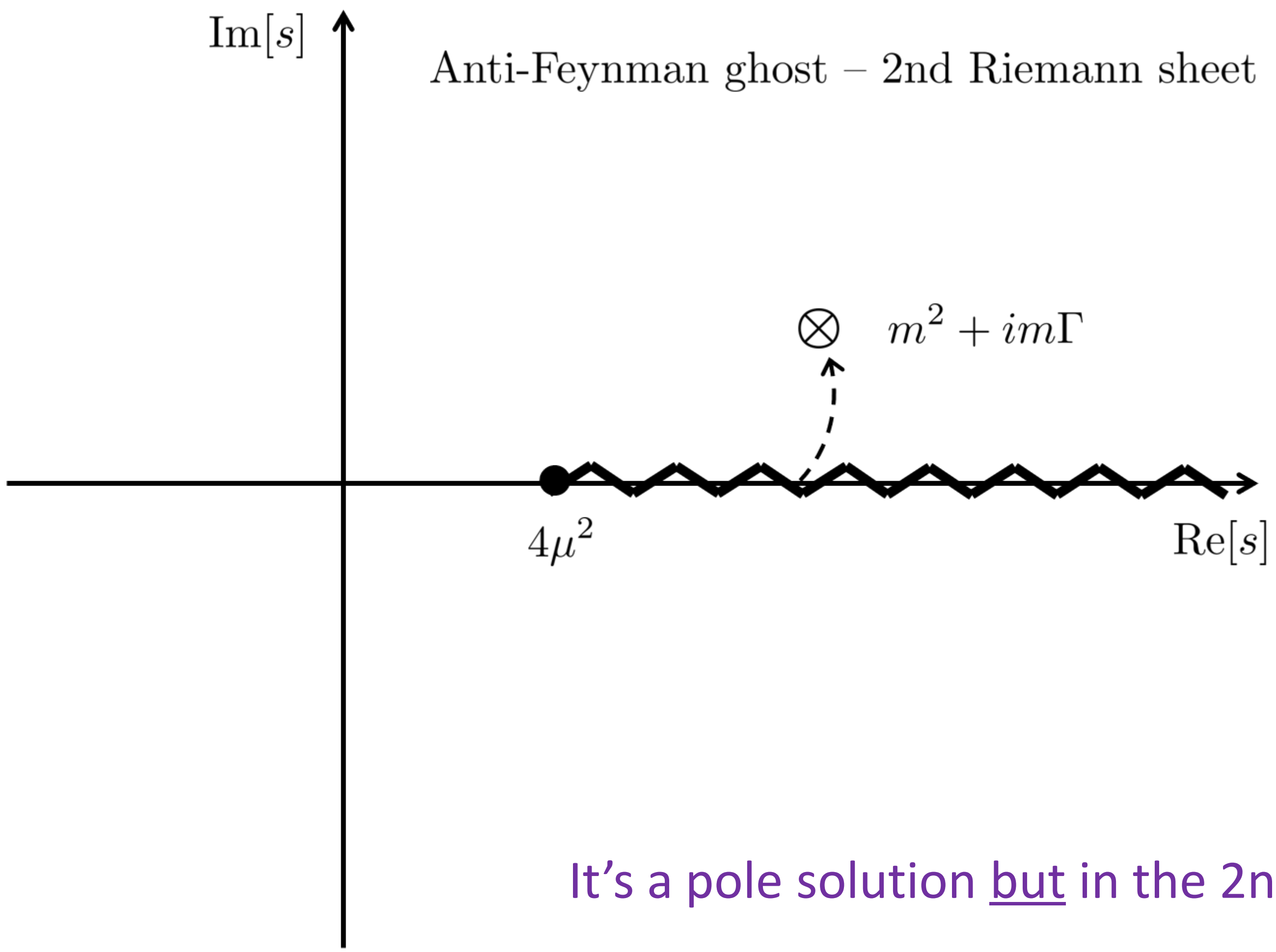

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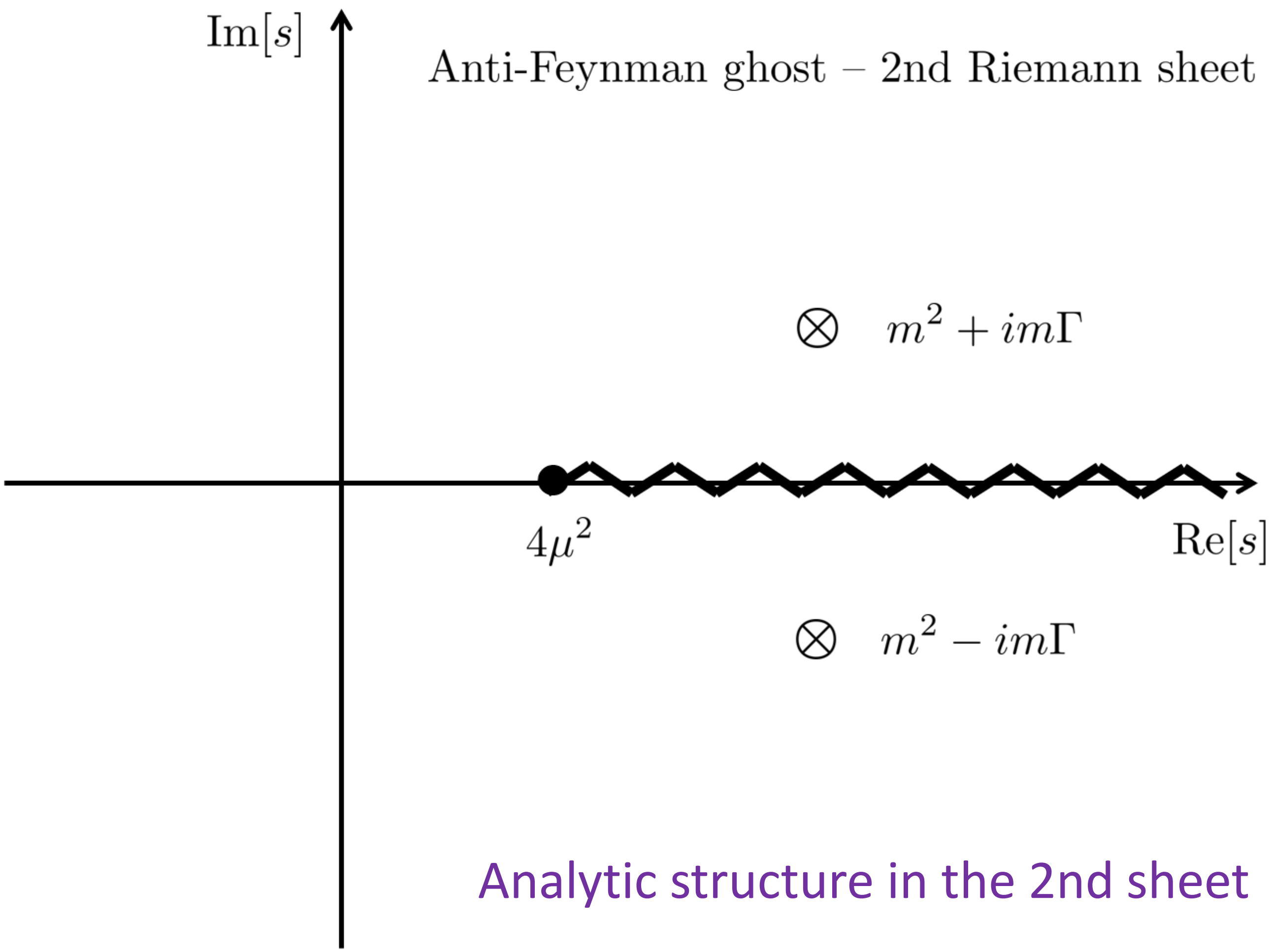

Analytic structure in the 2nd sheet

Ghost: stable or unstable?

<u>Feynman ghost</u>:

- no decay
- always an asymptotic state
- ‘physical’ complex masses (1st sheet)
- no unitarity

<u>Anti-Feynman ghost</u>:

- it can decay !
- no longer an asymptotic state !
- it can be projected out of the set of intermediate states !
- unitarity is satisfied !

Possible implications

- (anti-Feynman) ghost in Lee-Wick theories can decay

  [Lee & Wick 1969, Coleman 1969, Nakanishi 1970, Boulware & Gross 1984, Donoghue & Menezes 2018+]

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- (anti-Feynman) ghost in Lee-Wick theories can decay

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- Spin-2 (anti-Feynman) ghost in Quadratic Gravity can decay

  [Tomboulis 1977+; Tomboulis & Antoniadis 1986; Donoghue & Menezes 2018+]

$$\Gamma \sim \frac{m_2^3}{M_p^2}, \qquad \tau_{life} \sim \frac{1}{\Gamma} \sim \frac{M_p^2}{m_2^3}$$

If $m_2 \geq 2m_0$ & $m_0 \sim 10^{13} GeV \implies \tau_{life} \lesssim 10^{-3} GeV^{-1} \sim 10^{-28}\mathrm{sec}$

Conclusions

$$S = \frac{1}{2}\int d^4x \sqrt{-g}\left(M_p^2 R + \frac{\alpha}{6}R^2 - \frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

1. Quadratic Gravity is a unique (most conservative) QFT of gravity

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2. Future cosmological observation can constrain/falsify the theory

   Tensor-to-scalar ratio: [Sasaki et al. 2012; Salvio 2017+; Anselmi, Bianchi, Piva 2020]

$$r = \frac{24}{N_e^2}\frac{m_2^2}{m_0^2 + 2m_2^2},$$

If $N_e \sim 60\,,\, m_2 \geq 2m_0 \quad \Longrightarrow \quad 3\times 10^{-3} \lesssim r \lesssim 3.3\times 10^{-3}$

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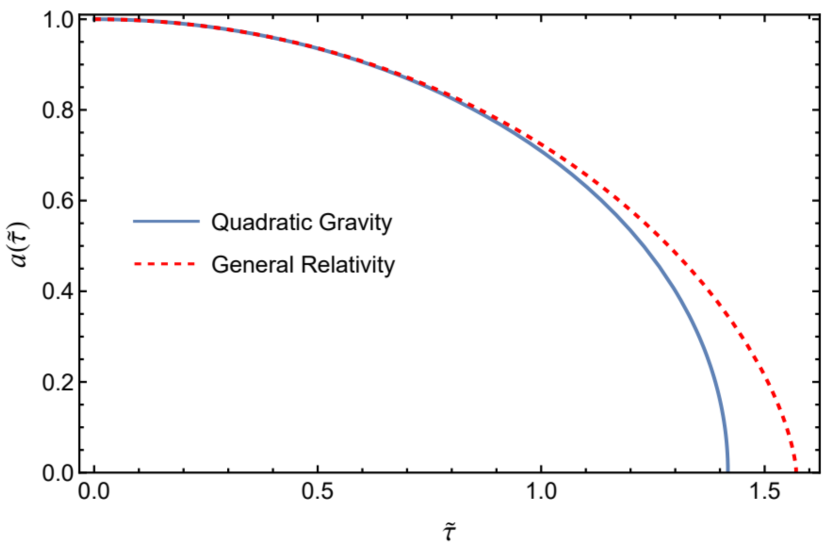
5. There is still a lot to learn! [E.g. see Roberto Percacci's talk tomorrow]

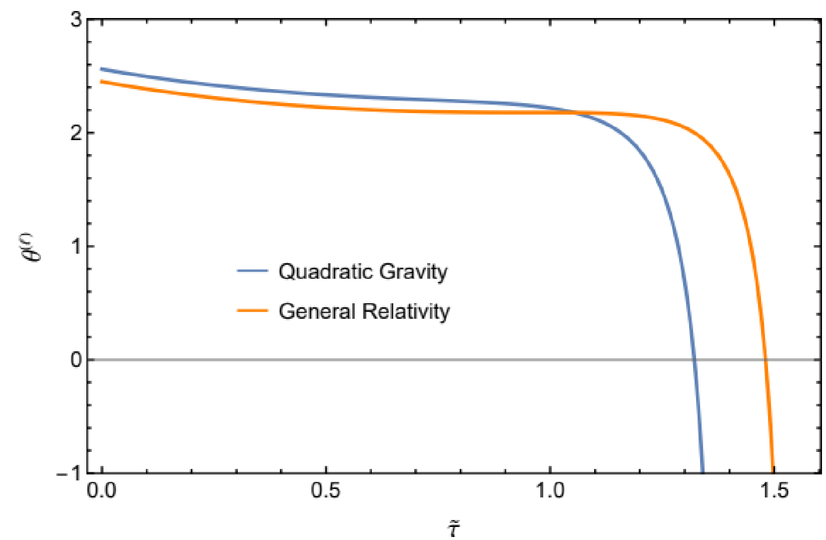
…Extra Slides…

Dust collapse in Quadratic Gravity

[L.B., Di Filippo, Kolar, Saueressig, arXiv:2410.XXXXX]

- Apparent horizon forms
- Singularity is reached in a finite comoving time (faster than in GR)
- The exterior solution cannot be horizonless





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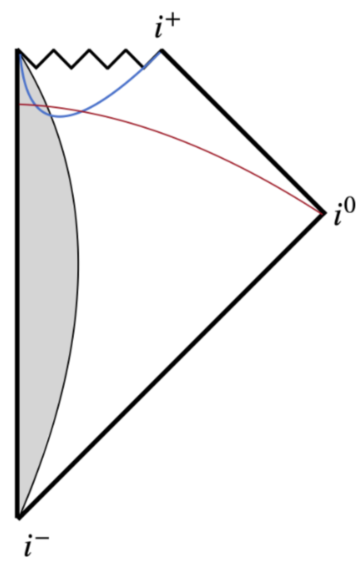

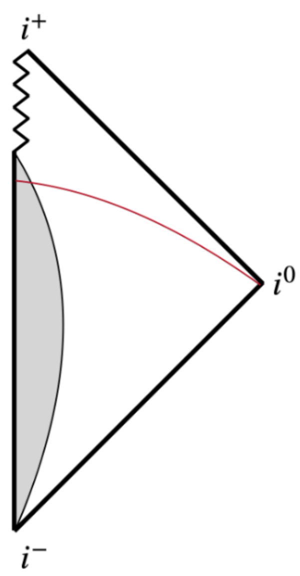

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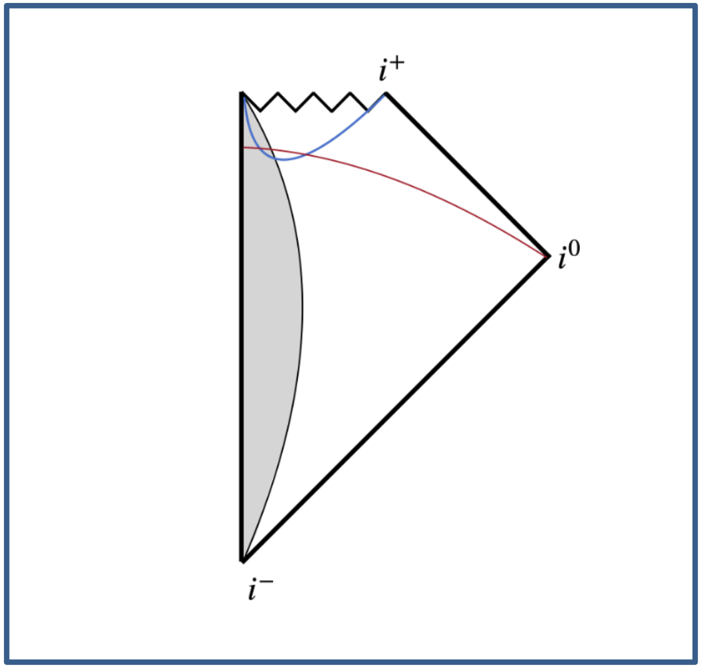

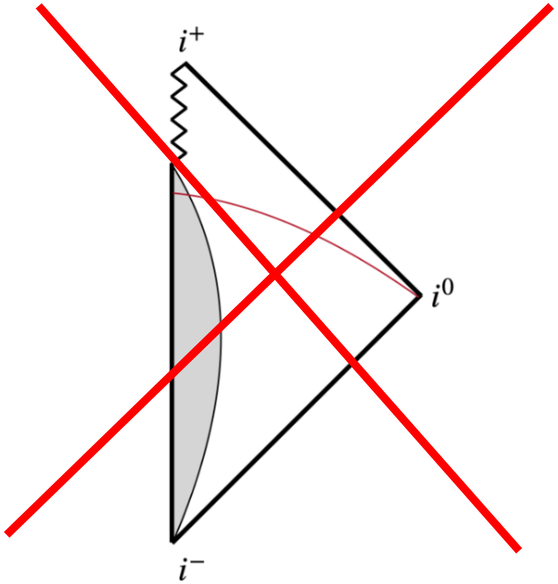

Remarks on the quantization(s)

- May be OK in the case of 2-derivative ghosts
- Still don't understand the case of 4-derivative ghosts:

$$G(x) = \int \frac{d^4p}{(2\pi)^4} e^{ip\cdot x} \frac{M^2}{(p^2 - i\epsilon)(p^2 + M^2 + i\varepsilon)}$$

$$= \int \frac{d^4p}{(2\pi)^4} e^{ip\cdot x} \left[\frac{1}{p^2 - i\epsilon} - \frac{1}{p^2 + M^2 + i\varepsilon}\right]$$

- 2 prescriptions together **but** only 1 type of Wick rotation
- Good behavior of position-space propagator is necessary for a good UV behavior of loop integrals

$$\int \frac{d^4k}{(2\pi)^4} \frac{1}{k^2(k^2 + M^2)} \frac{1}{k^2(k^2 + M^2)} \sim \int d^4x\, G(x) G(-x)$$

Motivations

Einstein's General Relativity:

$$S_{EH} = \frac{M_p^2}{2} \int d^4x \sqrt{-g}(R - 2\Lambda)$$

Issues:

1. Theoretical: perturbatively non-renormalizable
2. Observational: cannot explain CMB anisotropies (early times physics)

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Simplest model to explain 2. :

$$S = S_{EH} + S_\phi + \cdots, \qquad S_\phi \equiv \text{inflaton action}$$

Motivations

How to select ‘‘fundamental’’ Lagrangians?

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How to select ''fundamental'' Lagrangians?

- Guidance from Nature:
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‘Unique’ (strictly) renormalizable QFT of gravity in $D = 4$:

[Stelle, PRD (1977)]

$$S = \frac{1}{2}\int d^4x\,\sqrt{-g}\left(M_p^2(R - 2\Lambda) + \frac{\alpha}{6}R^2 - \frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

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massive spin-0, $\alpha \sim 10^{10}$

Natural explanation for inflation!

[Starobinsky, 1980+]

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Spin-2 massive ghost
[Salvio & Strumia 2015+; Anselmi & Piva 2017+; Donoghue & Menezes 2018+; Holdom 2015+, etc…]

Quadratic Gravity as Quantum Gravity

$$S = \frac{1}{2}\int d^4x\,\sqrt{-g}\left(M_p^2(R-2\Lambda)+\frac{\alpha}{6}R^2-\frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

Cosmological constant: $\Lambda \sim 10^{-122} M_p^2$

Natural candidate for inflaton: $\alpha \sim 10^{10}$

[Starobinsky, 1980+]

In my opinion, if we accept these facts very <u>important implications</u> follow:

1. The framework of perturbative QFT and the criterion of renormalizability (as a tool to select theories) are quite successful also when applied to gravity!

2. CMB observations have provided for the first time a <u>test</u> of higher-curvature gravity and an ‘indirect’ proof of quantized gravity (the scalar field is a gravitational dof)!!

3. Contrary to some beliefs, Starobinsky inflation is <u>not</u> just a model!

Motivations

<u>Obvious question:</u> What about the spin-2 massive ghost?

1. Throw the entire theory away just because maybe we don't know how to deal with the spin-2 ghost?

2. Or, instead, after appreciating the achievements described before, should we feel very motivated to understand the role of the ghost at a deeper level?

I opt for the 2nd option!

Question

$$S = \frac{1}{2}\int d^4x\,\sqrt{-g}\left(M_p^2(R-2\Lambda)+\frac{\alpha}{6}R^2-\frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

$$\lim_{\beta,\alpha\to\infty} S = \ ?$$

Simpler example

$$S = -\frac{1}{4g^2}\int d^4x\, \hat{F}^a_{\mu\nu}\hat{F}^{a\mu\nu}, \qquad \hat{F}^a_{\mu\nu} = \partial_\mu \hat{A}^a_\nu - \partial_\nu \hat{A}^a_\mu + f^{abc}\hat{A}^b_\mu\,\hat{A}^c_\nu$$

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Canonically normalized field:

$$A^a_\mu = \frac{1}{g}\hat{A}^a_\mu \qquad \Longrightarrow \qquad S = -\frac{1}{4}\int d^4x\, F^a_{\mu\nu}F^{a\mu\nu},$$

$$F^a_{\mu\nu} = \partial_\mu A^a_\nu - \partial_\nu A^a_\mu + g f^{abc} A^b_\mu\, A^c_\nu$$

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Limit:

$$\lim_{g\to 0} S = \frac{1}{2}\int d^4x\, A^a_\mu(\eta^{\mu\nu}\Box - \partial^\mu\partial^\nu)A^a_\nu \qquad (A^a_\mu = \text{fixed})$$

Question

$$S = \frac{1}{2}\int d^4x\,\sqrt{-g}\left(M_p^2(R-2\Lambda) + \frac{\alpha}{6}R^2 - \frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

$$\lim_{\beta,\alpha\to\infty} S = \ ?$$

1. Identify the canonically normalized fields
2. Identify the interaction couplings
3. Determine the structure of the particle spectrum

Additional spin-0 field

$$S[g,\phi] = \frac{\bar{M}_p^2}{2}\int d^4x\,\sqrt{-g}(R-2\Lambda) - \frac{\beta}{4}\int d^4x\,\sqrt{-g}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma} + S_0[g,\phi],$$

$$S_0[g,\phi] = \int d^4x\,\sqrt{-g}\left[-\frac{1}{2}\nabla^\mu\phi\nabla_\mu\phi - \frac{m_0^2}{2}\frac{3\bar{M}_p^2}{2}\left(1-e^{\sqrt{2/3}\phi/\bar{M}_p}\right)^2\right]$$

$$\bar{M}_p^2 \equiv M_p^2 + \frac{4}{3}\alpha\Lambda$$

Shifted Planck Mass when $\Lambda \neq 0$

$$m_0^2 \equiv \frac{M_p^2}{\alpha}$$

Mass of the spin-0 field

Additional spin-2 field

Spin-2 field $f_{\mu\nu}$:   [Kaku et al. (1977); Hindawi et al. (1996); Tekin (2016); Anselmi & Piva (2018)]

$$S[g,\phi,f] = \frac{\tilde{M}_p^2}{2}\int d^4x\,\sqrt{-g}(R-2\Lambda) + S_0[g,\phi]$$
$$-\int d^4x\,\sqrt{-g}\left[\tilde{M}_p\left(G_{\mu\nu}+\Lambda g_{\mu\nu}\right)f^{\mu\nu} - \frac{m_2^2}{2}\left(f_{\mu\nu}f^{\mu\nu}-f^2\right)\right]$$

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Diagonalizaton: $g_{\mu\nu} \to g_{\mu\nu} - \frac{2}{\widetilde{M}_p} f_{\mu\nu}$

$$S[g,\phi,f] = \frac{\widetilde{M}_p^2}{2}\int d^4x\,\sqrt{-g}(R-2\Lambda) + S_0\left[g - 2f/\widetilde{M}_p,\phi\right] + S_2[g,f],$$

$$\widetilde{M}_p^2 \equiv M_p^2 + \frac{2}{3}(2\alpha+\beta)\Lambda$$

Shifted Planck Mass when $\Lambda \neq 0$

Additional spin-2 field

$$S_2[g,f] = -S_{FP}[g,f] - \int d^4x\sqrt{-g}\left[\left(2f_\mu^\rho f_{\rho\nu} - f f_{\mu\nu}\right)R^{\mu\nu} + \left(\Lambda - \frac{R}{2}\right)\left(f_{\mu\nu}f^{\mu\nu} - \frac{1}{2}f^2\right)\right]$$

$$-\frac{1}{2}\frac{m_2^2}{\widetilde{M}_p}\int d^4x\sqrt{-g}\left[5f_{\mu\nu}f^{\mu\nu}f - 4f^{\mu\nu}f_\mu^\rho f_{\rho\nu} - f^3\right]$$

$$+\frac{8}{3}\frac{1}{M_p^2}\frac{1}{\widetilde{M}_p}\int d^4x d^4y d^4z \frac{\delta^{(3)}S_{EH}}{\delta g_{\mu\nu}(x)\delta g_{\rho\sigma}(y)\delta g_{\alpha\beta}(z)} f_{\mu\nu}(x) f_{\rho\sigma}(y) f_{\alpha\beta}(z)$$

$$+O(f^4)$$

$S_{PF}[g,f]$ is a covariantized Fierz-Pauli action for $f_{\mu\nu}$ with mass m_2^2

$$m_2^2 = \frac{\widetilde{M}_p^2}{\beta} = \frac{M_p^2}{\beta} + \frac{2}{3}\left(2\frac{\alpha}{\beta} + 1\right)\Lambda$$

spin-2 ghost mass depends on Λ !

$$m_2^2 \geq \frac{2}{3}\Lambda$$

$$(\Lambda \geq 0, \beta > 0)$$

Couplings

n-point interaction couplings:

$$\sim \left(\frac{1}{\widetilde{M}_p}\right)^{n-2} = \left(\frac{1}{M_p}\right)^{n-2} \left(\frac{1}{1+2\Lambda(2\alpha+\beta)/3M_p^2}\right)^{\frac{n-2}{2}}$$

$$\sim \left(\frac{1}{\bar{M}_p}\right)^{n-2} = \left(\frac{1}{M_p}\right)^{n-2} \left(\frac{1}{1+4\Lambda\alpha/3M_p^2}\right)^{\frac{n-2}{2}}$$

Couplings dependence on Λ $\Longrightarrow$ additional dependence on β and α !

Degrees of freedom

Linear analysis:

$$\frac{\delta S}{\delta f^{\mu\nu}} = 0 \quad \Leftrightarrow \quad \widetilde{M}_p\left(G_{\mu\nu} + \Lambda g_{\mu\nu}\right) = m_2^2\left(f_{\mu\nu} - g_{\mu\nu} f\right)$$

4 Constraints:

$$\boxed{\nabla_\mu f^\mu_\nu = \nabla_\nu f}$$

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4 Constraints:

$$\boxed{\nabla_\mu f^\mu_\nu = \nabla_\nu f}$$

$$g^{\mu\nu}\frac{\delta S}{\delta g^{\mu\nu}} = 0 \quad \Leftrightarrow \quad \widetilde{M}_p\left(m_2^2 - \frac{2}{3}\Lambda\right) f = 0$$

1 trace constraint:

$$\boxed{f = 0} \qquad (m_2^2 - 2\Lambda/3 \neq 0)$$

Degrees of freedom

Linear analysis:

$$\frac{\delta S}{\delta f^{\mu\nu}} = 0 \quad \Leftrightarrow \quad \widetilde{M}_p\left(G_{\mu\nu} + \Lambda g_{\mu\nu}\right) = m_2^2\left(f_{\mu\nu} - g_{\mu\nu} f\right)$$

4 Constraints: $\boxed{\nabla_\mu f^\mu_\nu = \nabla_\nu f}$

$$g^{\mu\nu}\frac{\delta S}{\delta g^{\mu\nu}} = 0 \quad \Leftrightarrow \quad \widetilde{M}_p\left(m_2^2 - \frac{2}{3}\Lambda\right) f = 0$$

1 trace constraint: $\boxed{f = 0}$ $\quad (m_2^2 - 2\Lambda/3 \neq 0)$

dof $f_{\mu\nu}$: $\quad 10 - 4 - 1 = 5 \quad \Longrightarrow$ massive spin-2 with 5 helicities

Degrees of freedom: remark

$$\widetilde{M}_p\left(m_2^2 - \frac{2}{3}\Lambda\right) f = 0, \qquad\qquad m_2^2 - \frac{2}{3}\Lambda = 0 \quad ?$$

In this case the linear trace equation vanishes <u>identically</u>

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In this case the linear trace equation vanishes identically

Two possibilities:

1. $\Lambda = 0$: $\quad m_2^2 = 0$ (massless)
2. $\Lambda \neq 0$: $\quad m_2^2 = \frac{2}{3}\Lambda$ (partially massless)

These cases need a separate discussion!

Degrees of freedom: remark

$$\widetilde{M}_p \left(m_2^2 - \frac{2}{3}\Lambda \right) f = 0, \qquad m_2^2 - \frac{2}{3}\Lambda = 0 \quad ?$$

In this case the linear trace equation vanishes identically

Two possibilities:

1. $\Lambda = 0$: $m_2^2 = 0$ (massless)
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These cases need a separate discussion!

NB:

$$\beta \longrightarrow \infty \quad \Longrightarrow \quad m_2^2 - \frac{2\Lambda}{3} = \frac{\overline{M}_p^2}{\beta} \longrightarrow 0$$

Case $\Lambda = 0$

Stückelberg formalism:

$$f_{\mu\nu} = \varphi_{\mu\nu} + \frac{1}{m_2}\left(\nabla_\mu A_\nu + \nabla_\nu A_\mu\right) + \frac{2}{m_2^2}\nabla_\mu \nabla_\nu \chi$$

The limit is regular in $D = 4$:

$$\lim_{\beta\to\infty} S = \frac{M_p^2}{2}\int d^4x\sqrt{-g}\,R - M_p\int d^4x\sqrt{-g}G_{\mu\nu}\varphi^{\mu\nu} + \frac{1}{4}\int d^4x\sqrt{-g}F^{\mu\nu}F_{\mu\nu} + S_{\phi\chi}[g,\phi,\chi]$$

$$S_{\phi\chi}[g,\phi,\chi] = \int d^4x\sqrt{-g}\left[\frac{1}{2}e^{-\sqrt{2/3}\chi/M_p}\left(\nabla_\mu\chi\nabla^\mu\chi - \nabla_\mu\phi\nabla^\mu\phi\right) - \frac{m_0^2}{2}\frac{3M_p^2}{2}e^{-2\sqrt{2/3}\chi/M_p}\left(1 - e^{\sqrt{2/3}\phi/M_p}\right)^2\right]$$

$f_{\mu\nu}$ splits into 5 interacting massless ghost-like dofs ($\pm 2,\ \pm 1,\ 0$)

[Saravani & Hinterbichler (2015)]

Case $\Lambda > 0$

It is NOT a massless limit:

$$\beta \to \infty \quad \Longrightarrow \quad m_2^2 = \frac{\bar{M}_p^2}{\beta} + \frac{2}{3}\Lambda \longrightarrow \frac{2}{3}\Lambda$$

Case $\Lambda > 0$

It is NOT a massless limit:

$$\beta \to \infty \qquad \Longrightarrow \qquad m_2^2 = \frac{\bar{M}_p^2}{\beta} + \frac{2}{3}\Lambda \longrightarrow \frac{2}{3}\Lambda$$

NB: In Massive Gravity theories this limit is known as *partially massless limit* and in general may lead to strong coupling! [de Rham et al. (2018)]

Case $\Lambda > 0$

$$\lim_{\beta\to\infty} S = S_{EH}^{(2)}[\bar{g},h] - S_{FP}^{(m^2=2\Lambda/3)}[\bar{g},\varphi] + S_{\phi\chi}$$

$$S_{FP}^{(m^2=2\Lambda/3)} = \int d^4x\sqrt{-g}\Big[-\frac{1}{2}\nabla_\rho\varphi_{\mu\nu}\nabla^\rho\varphi^{\mu\nu} + \nabla_\rho\varphi_{\mu\nu}\nabla^\mu\varphi^{\rho\nu} - \nabla_\mu\varphi\nabla_\nu\varphi^{\mu\nu} + \frac{1}{2}\nabla_\rho\varphi\nabla^\rho\varphi$$
$$+\Lambda\left(\varphi_{\mu\nu}\varphi^{\mu\nu} - \frac{1}{2}\varphi^2\right) - \frac{\Lambda}{3}\left(\varphi_{\mu\nu}\varphi^{\mu\nu} - \varphi^2\right)\Big]$$

$$S_{\phi\chi}[g,\phi,\chi] = \int d^4x\sqrt{-g}\Big[\frac{1}{2}e^{-\sqrt{2/3}\chi/\bar{M}_p}\left(\nabla_\mu\chi\nabla^\mu\chi - \nabla_\mu\phi\nabla^\mu\phi\right)$$
$$-\Lambda\,\bar{M}_p^2\left(1 - e^{\sqrt{2/3}\phi/\bar{M}_p}\right)^2 - \frac{m_0^2}{2}\frac{3\bar{M}_p^2}{2}e^{-2\sqrt{2/3}\chi/\bar{M}_p}\left(1 - e^{\sqrt{2/3}\phi/\bar{M}_p}\right)^2\Big]$$

- $f_{\mu\nu}$ splits into 1 partially massless graviton (4 dof) + 1 scalar dof
- Enhanced symmetry: $\delta\varphi_{\mu\nu} = \nabla_\mu\nabla_\nu\xi(x) + \frac{\Lambda}{3}\xi(x)$
- The compatible metric background is $\bar{R}_{\mu\nu} = \Lambda\bar{g}_{\mu\nu}$

Case $\Lambda > 0$

We can also take the limit $\alpha \longrightarrow \infty$ and kill of the interactions in the spin-0 sector

In summary, if the cosmological constant is non-zero (and positive), in the limits $\beta, \alpha \longrightarrow \infty$ we get a free theory whose degrees of freedom are

massless graviton (2 dofs) + partially massless graviton (4 dofs) + 2 scalars

Summary

$$S = \frac{1}{2}\int d^4x\,\sqrt{-g}\left(M_p^2(R-2\Lambda)+\frac{\alpha}{6}R^2-\frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

- There are strong theoretical and experimental motivations to work on Quadratic Gravity
- Some remarks on its quantization(s) [things I still don’t fully understand]
- I asked the question: $\lim_{\beta,\alpha\to\infty} S = \,?$

Summary

$$S = \frac{1}{2}\int d^4x\,\sqrt{-g}\left(M_p^2(R-2\Lambda) + \frac{\alpha}{6}R^2 - \frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

- Renormalizability + ghost-like nature of spin-2 $\Longrightarrow$ limit $\beta \to \infty$ is regular
- The limits $\beta, \alpha \to \infty$ depend non-trivially on Λ
- When $\Lambda \neq 0$: structure of degrees of freedom is different (but same number); the limits kill all the interactions

Physical implications?

$$S = \frac{1}{2}\int d^4x\,\sqrt{-g}\left(M_p^2(R-2\Lambda) + \frac{\alpha}{6}R^2 - \frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

- Can the result of the limits $\beta, \alpha \to \infty$ help understand the high-energy behavior of the spin-2 ghost?
- Does a $\Lambda \neq 0$ affect current quantization approaches to Quadratic Gravity?
- Role of the cosmological constant in quantum gravity?

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- Can the result of the limits $\beta, \alpha \to \infty$ help understand the high-energy behavior of the spin-2 ghost?
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- Role of the cosmological constant in quantum gravity?

A nice formula:

$$\Lambda = \frac{3}{2}m_2^2\,\frac{\beta - M_p^2/m_2^2}{\beta + 2M_p^2/m_0^2}$$

Case $\Lambda = 0$

It is a massless limit:

$$\beta \to \infty \qquad \Longrightarrow \qquad m_2^2 = \frac{M_p^2}{\beta} \to 0$$

NB: typically, the massless limit in theories of Massive Gravity can lead to strong coupling even below M_p. [Reviews by Hinterbichler (2011) and de Rham (2014)]

Digression on Massive Gravity with $\Lambda = 0$

$$S_{MG} = \int d^4x\sqrt{-g}\Big[-\frac{1}{2}\nabla_\rho f_{\mu\nu}\nabla^\rho f^{\mu\nu} + \nabla_\rho f_{\mu\nu}\nabla^\mu f^{\rho\nu} - \nabla_\mu f\nabla_\nu f^{\mu\nu} + \frac{1}{2}\nabla_\rho f\nabla^\rho f$$
$$-\frac{m_2^2}{2}\left(f_{\mu\nu}f^{\mu\nu} - f^2\right) + O(f^3)\Big]$$

Naively, the limit $m_2^2 \to 0$ seems to give a *massless* spin-2 with 2 dofs

$$S_{MG}^{(2)} = \int d^4x\sqrt{-g}\Big[-\frac{1}{2}\nabla_\rho f_{\mu\nu}\nabla^\rho f^{\mu\nu} + \nabla_\rho f_{\mu\nu}\nabla^\mu f^{\rho\nu} - \nabla_\mu f\nabla_\nu f^{\mu\nu} + \frac{1}{2}\nabla_\rho f\nabla^\rho f\Big]$$

Gauge symmetry gives 2 dofs

$$\delta f_{\mu\nu}(x) = \nabla_\mu\xi_\nu(x) + \nabla_\nu\xi_\mu(x),$$

Digression on Massive Gravity with $\Lambda = 0$

$$S_{MG} = \int d^4x\sqrt{-g}\Big[-\frac{1}{2}\nabla_\rho f_{\mu\nu}\nabla^\rho f^{\mu\nu} + \nabla_\rho f_{\mu\nu}\nabla^\mu f^{\rho\nu} - \nabla_\mu f\nabla_\nu f^{\mu\nu} + \frac{1}{2}\nabla_\rho f\nabla^\rho f - \frac{m_2^2}{2}\left(f_{\mu\nu}f^{\mu\nu} - f^2\right) + O(f^3)\Big]$$

Stückelberg formalism:

$$f_{\mu\nu} = \varphi_{\mu\nu} + \frac{1}{m_2}\left(\nabla_\mu A_\nu + \nabla_\nu A_\mu\right) + \frac{2}{m_2^2}\nabla_\mu\nabla_\nu\chi,$$

Gauge symmetries:

$$\delta\varphi_{\mu\nu} = \nabla_\mu\xi_\nu + \nabla_\nu\xi_\mu,$$
$$\delta A_\mu = -m_2\xi_\mu + \nabla_\mu\xi,$$
$$\delta\chi = -m_2\,\xi,$$

Massless limit (2+2+1 = 5 dofs):

$$S_{MG}[\varphi, A, \chi] = S_{FP}^{(m_2=0)}[\varphi] + \int d^4x\sqrt{-g}\left(-\frac{1}{2}F^{\mu\nu}F_{\mu\nu} - 3\nabla_\rho\chi\nabla^\rho\chi\right) + O(f^3)$$

Possible strong coupling from helicity-0 interactions: $O(f^3)\sim\frac{1}{m_2}O(\chi^3)\ \rightarrow\ \infty$

[Reviews by Hinterbichler (2011) and de Rham (2014)]

Limit $\beta \to \infty$ with $\Lambda = 0$: massless limit

Does a strong coupling (below M_p) arise in quadratic gravity?

[first asked by Hinterbichler & Saravani (2016)]

A strong coupling in the limit $m_2^2 \to 0$ (i.e., $\beta \to \infty$) can be avoided *only* in $D = 4$!

Stückelberg decomposition for $\Lambda = 0$ and in D dimensions:

$$f_{\mu\nu} = \tilde{f}_{\mu\nu} + \frac{1}{m_2}(\nabla_\mu \tilde{A}_\nu + \nabla_\nu \tilde{A}_\mu), \qquad \tilde{A}_\mu = A_\mu + \frac{1}{m_2}\nabla_\mu \chi$$

$$\Longrightarrow S_2'[g,f] = \frac{M_p^{D-2}}{2}\int d^D x\sqrt{-g}R + \int d^D x\sqrt{-g}\left[-M_p^{\frac{D-2}{2}} G_{\mu\nu} f^{\mu\nu} + \frac{m_2^2}{2}(f_{\mu\nu}f^{\mu\nu} - f^2)\right]$$

$$= \frac{M_p^{D-2}}{2}\int d^D x\sqrt{-g}R + \int d^D x\sqrt{-g}\Big[-M_p^{\frac{D-2}{2}} G_{\mu\nu} \tilde{f}^{\mu\nu} + \frac{m_2^2}{2}(\tilde{f}_{\mu\nu}\tilde{f}^{\mu\nu} - \tilde{f}^2)$$

$$+\frac{1}{2}F^{\mu\nu}F_{\mu\nu} + 2m_2 \tilde{f}^{\mu\nu}(\nabla_\mu \tilde{A}_\nu - g_{\mu\nu}\nabla^\rho \tilde{A}_\rho) - 2R^{\mu\nu}\tilde{A}_\mu \tilde{A}_\nu\Big],$$

Possible strong coupling from $R^{\mu\nu}\tilde{A}_\mu \tilde{A}_\nu \sim \frac{1}{m_2^2} R^{\mu\nu}\nabla_\mu\chi\nabla_\nu\chi$???

Limit $\beta \to \infty$ with $\Lambda = 0$: massless limit

Make a field redefinition:

$$\tilde{f}_{\mu\nu} \to \tilde{f}_{\mu\nu} + a\,\tilde{A}_\mu \tilde{A}_\nu + b\, g_{\mu\nu}\, \tilde{A}_\rho \tilde{A}^\rho$$

In the massless limit $m_2^2 \to 0$ $(\beta \to \infty,\ \Lambda = 0)$ we get

$$\Longrightarrow \quad S_2'[g,\tilde{f},\tilde{A}] = \frac{M_p^{D-2}}{2}\int d^D x\,\sqrt{-g}R + \int d^D x\,\sqrt{-g}\Big[-M_p^{\frac{D-2}{2}} G_{\mu\nu}\tilde{f}^{\mu\nu} + \frac{m_2^2}{2}(\tilde{f}_{\mu\nu}\tilde{f}^{\mu\nu} - \tilde{f}^2) + \frac{1}{2}F^{\mu\nu}F_{\mu\nu}$$

$$+2m_2\tilde{f}^{\mu\nu}\left(\nabla_\mu \tilde{A}_\nu - g_{\mu\nu}\,\nabla^\rho \tilde{A}_\rho\right) + m_2^2 a\,\tilde{f}^{\mu\nu}\tilde{A}_\mu\tilde{A}_\nu + m_2^2[b(1-D) - a]\tilde{f}\tilde{A}_\rho\tilde{A}^\rho$$

$$-\left(aM_p^{\frac{D-2}{2}} + 2\right)R^{\mu\nu}\tilde{A}_\mu\tilde{A}_\nu + M_p^{\frac{D-2}{2}}\left(\left(1-\tfrac{D}{2}\right)b - \tfrac{a}{2}\right)R\tilde{A}_\rho\tilde{A}^\rho$$

$$-m_2(2b(1-D) - 3a)\tilde{A}_\mu\tilde{A}_\nu\nabla^\mu\tilde{A}^\nu - \frac{m_2^2}{2}(b^2 D(1-D) + 2ab(1-D))\left(\tilde{A}_\rho\tilde{A}^\rho\right)^2\Big]$$

<u>4 conditions</u> to avoid strong coupling in the massless limit:

$$aM_p^{\frac{D-2}{2}} + 2 = 0, \qquad 2b(1-D) - 3a = 0,$$

$$\left(1-\tfrac{D}{2}\right)b - \tfrac{a}{2} = 0, \qquad b^2 D(1-D) + 2ab(1-D) = 0$$

can be simultaneously satisfied *only* in $D = 4$!!!

$$a = -\frac{2}{M_p} = -2b,$$

Digression on Massive Gravity with $\Lambda > 0$

$$S_{MG} = \int d^4x \sqrt{-g}\Big[-\frac{1}{2}\nabla_\rho f_{\mu\nu}\nabla^\rho f^{\mu\nu} + \nabla_\rho f_{\mu\nu}\nabla^\mu f^{\rho\nu} - \nabla_\mu f \nabla_\nu f^{\mu\nu} + \frac{1}{2}\nabla_\rho f \nabla^\rho f$$
$$+\Lambda\left(f_{\mu\nu}f^{\mu\nu} - \tfrac{1}{2}f^2\right) - \tfrac{m_2^2}{2}\left(f_{\mu\nu}f^{\mu\nu} - f^2\right) + O(f^3)\Big]$$

Naively, the limit $m_2^2 \to \frac{2}{3}\Lambda$ seems to give a *partially massless* spin-2 with <u>4 dofs</u>

$$S^{(2)}_{MG} = \int d^4x \sqrt{-g}\Big[-\frac{1}{2}\nabla_\rho f_{\mu\nu}\nabla^\rho f^{\mu\nu} + \nabla_\rho f_{\mu\nu}\nabla^\mu f^{\rho\nu} - \nabla_\mu f \nabla_\nu f^{\mu\nu} + \frac{1}{2}\nabla_\rho f \nabla^\rho f$$
$$+\Lambda\left(f_{\mu\nu}f^{\mu\nu} - \tfrac{1}{2}f^2\right) - \tfrac{\Lambda}{3}\left(f_{\mu\nu}f^{\mu\nu} - f^2\right)\Big]$$

Scalar gauge symmetry (10-4-2=4 dofs)

$$\delta f_{\mu\nu}(x) = \nabla_\mu\nabla_\nu\zeta(x) + \frac{\Lambda}{3}g_{\mu\nu}\zeta(x),$$

[Deser, Nepomechie (1984); Deser, Waldron (2001+); etc]

Digression on Massive Gravity with $\Lambda > 0$

$$S_{MG} = \int d^4x\sqrt{-g}\Big[-\frac{1}{2}\nabla_\rho f_{\mu\nu}\nabla^\rho f^{\mu\nu} + \nabla_\rho f_{\mu\nu}\nabla^\mu f^{\rho\nu} - \nabla_\mu f \nabla_\nu f^{\mu\nu} + \frac{1}{2}\nabla_\rho f \nabla^\rho f$$
$$+\Lambda\left(f_{\mu\nu}f^{\mu\nu} - \tfrac{1}{2}f^2\right) - \tfrac{m_2^2}{2}\left(f_{\mu\nu}f^{\mu\nu} - f^2\right) + O(f^3)\Big]$$

Stückelberg trick

$$f_{\mu\nu} = \varphi_{\mu\nu} + \sqrt{\frac{3}{\Lambda}}\,\frac{1}{\Delta}\left(\nabla_\mu\nabla_\nu\chi + g_{\mu\nu}\frac{\Lambda}{3}\chi\right)$$

$$\Delta \equiv m_2^2 - \frac{2}{3}\Lambda$$

Gauge symmetries:

$$\delta\varphi_{\mu\nu} = \nabla_\mu\nabla_\nu\zeta + \frac{\Lambda}{3}g_{\mu\nu}\zeta,$$

$$\delta\chi = -\sqrt{\frac{\Lambda}{3}}\,\Delta\,\zeta,$$

Partially massless limit $\Delta \to 0$ (4+1=5 dofs):

$$S_{MG}[\varphi,\chi] = S_{FP}^{(\Delta=0)}[\varphi] + 3\int d^4x\sqrt{-g}\left(-\frac{1}{2}\nabla_\rho\chi\nabla^\rho\chi - \frac{m_\chi^2}{2}\chi^2\right) + O(f^3), \qquad m_\chi^2 \equiv -\frac{4}{3}\Lambda$$

Possible strong coupling from χ interactions: $\quad O(f^3) \sim \frac{1}{\Delta}O(\chi^3) \;\to\; \infty$

[de Rham, Hinterbichler, Johnson (2018)]

Case $\Lambda > 0$

Stückelberg formalism:

$$f_{\mu\nu} = \varphi_{\mu\nu} + \frac{1}{\sqrt{\Lambda}\sqrt{m_2^2 - 2\Lambda/3}}\left(\nabla_\mu\nabla_\nu\chi + \frac{\Lambda}{3}g_{\mu\nu}\chi\right)$$

$$\lim_{\beta\to\infty} S = \frac{\widetilde{M}_p^2}{2}\int d^4x\,\sqrt{-g}(R - 2\Lambda)$$
$$+\int d^4x\,\sqrt{-g}\left[-\widetilde{M}_p\left(G_{\mu\nu} + \Lambda g_{\mu\nu}\right)\varphi^{\mu\nu} + \frac{\Lambda}{3}\left(\varphi_{\mu\nu}\varphi^{\mu\nu} - \varphi^2\right)\right] + S_{\phi\chi}$$

$$S_{\phi\chi}[g,\phi,\chi] = \int d^4x\,\sqrt{-g}\left[\frac{1}{2}e^{-\sqrt{2/3}\chi/\bar{M}_p}\left(\nabla_\mu\chi\nabla^\mu\chi - \nabla_\mu\phi\nabla^\mu\phi\right)\right.$$
$$\left. -\Lambda\,\bar{M}_p^2\left(1 - e^{\sqrt{2/3}\phi/\bar{M}_p}\right)^2 - \frac{m_0^2}{2}\frac{3\bar{M}_p^2}{2}e^{-2\sqrt{2/3}\chi/\bar{M}_p}\left(1 - e^{\sqrt{2/3}\phi/\bar{M}_p}\right)^2\right]$$

$f_{\mu\nu}$ splits into 1 partially massless graviton (4 dof) + 1 scalar dof

Case $\Lambda > 0$

Interaction couplings: $\frac{1}{\tilde{M}_p} \sim \frac{1}{\sqrt{\Lambda\beta}}$ for spin-2 sector & $\frac{1}{\bar{M}_p}$ for spin-0 sector

- Expand in $\varphi_{\mu\nu}$ and $g_{\mu\nu} = \bar{g}_{\mu\nu} + \frac{2}{\tilde{M}_p} h_{\mu\nu}$
- Diagonalize kinetic term for $h_{\mu\nu}$ and $\varphi_{\mu\nu}$

$$\lim_{\beta\to\infty} S = S_{EH}^{(2)}[\bar{g}, h] - S_2^{(2)}[\bar{g}, \varphi] \,|_{m_2^2=\frac{2}{3}\Lambda} + S_{\phi\chi}[\bar{g}, \phi, \chi]$$

Spin-2 sector completely decouples!

The only compatible metric background is

$$\bar{R}_{\mu\nu} = \Lambda \bar{g}_{\mu\nu}$$

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Spin-2 sector completely decouples!

The only compatible metric background is

$$\bar{R}_{\mu\nu} = \Lambda \bar{g}_{\mu\nu}$$

Case $\Lambda > 0$

Resulting interacting theory:

$$S_{\phi\chi}[g,\phi,\chi] = \int d^4x\,\sqrt{-g}\left[\frac{1}{2}e^{-\sqrt{2/3}\chi/\bar{M}_p}\left(\nabla_\mu\chi\nabla^\mu\chi - \nabla_\mu\phi\nabla^\mu\phi\right)\right.$$

$$\left. -\Lambda\,\bar{M}_p^2\left(1 - e^{\sqrt{2/3}\phi/\bar{M}_p}\right)^2 - \frac{m_0^2}{2}\frac{3\bar{M}_p^2}{2}e^{-2\sqrt{2/3}\chi/\bar{M}_p}\left(1 - e^{\sqrt{2/3}\phi/\bar{M}_p}\right)^2\right]$$

$$\bar{M}_p^2 = M_p^2 + \frac{4}{3}\alpha\Lambda$$

Case $\Lambda > 0$

Resulting interacting theory:

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$$\bar{M}_p^2 = M_p^2 + \frac{4}{3}\alpha\Lambda$$

Field redefinitions: $\phi = -\sqrt{\frac{3}{2}}\bar{M}_p\log\left(\frac{\tilde{\chi}+\tilde{\phi}}{\tilde{\chi}-\tilde{\phi}}\right)$, $\chi = -\sqrt{\frac{3}{2}}\bar{M}_p\log\left(\frac{\tilde{\chi}^2-\tilde{\phi}^2}{6\bar{M}_p^2}\right)$

$$S_{\phi\chi}[g,\tilde{\phi},\tilde{\chi}] = \int d^4x\sqrt{-g}\left[\frac{1}{2}\left(\partial_\mu\tilde{\chi}\partial^\mu\tilde{\chi} - \partial_\mu\tilde{\phi}\partial^\mu\tilde{\phi}\right) - V(\tilde{\phi},\tilde{\chi})\right]$$

$$V(\tilde{\phi},\tilde{\chi}) = \frac{\Lambda}{36\bar{M}_p^2}\left(\tilde{\chi}^2-\tilde{\phi}^2-6\bar{M}_p^2\right)^2 + \frac{m_0^2}{36\bar{M}_p^2}\tilde{\phi}^2\left(\tilde{\chi}+\tilde{\phi}\right)^2$$

Spectrum

$$S = \frac{1}{2}\int d^4x\,\sqrt{-g}\left(M_p^2(R - 2\Lambda) + \frac{\alpha}{6}R^2 - \frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

Massless graviton:

$g_{\mu\nu}$ (2 dofs)

Massive spin-0:

ϕ (1 dof)

$$m_0^2 = \frac{M_p^2}{\alpha}$$

Massive spin-2 ghost:

$f_{\mu\nu}$ (5 dofs)

$$m_2^2 = \frac{\widetilde{M}_p^2}{\beta}$$

$$\widetilde{M}_p^2 = M_p^2 + \frac{2}{3}\Lambda\left(2\frac{\alpha}{\beta} + 1\right)$$

Spectrum

$$S = \frac{1}{2}\int d^4x\sqrt{-g}\left(M_p^2(R-2\Lambda) + \frac{\alpha}{6}R^2 - \frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

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$f_{\mu\nu}$ (5 dofs)

$$m_2^2 = \frac{\widetilde{M}_p^2}{\beta}$$

$$\widetilde{M}_p^2 = M_p^2 + \frac{2}{3}\Lambda\left(2\frac{\alpha}{\beta} + 1\right)$$

$\beta \longrightarrow \infty \quad \Longrightarrow$

$\Lambda = 0$: $\quad m_2^2 \longrightarrow 0$ (massless limit)

$\Lambda \neq 0$: $\quad m_2^2 \longrightarrow \frac{2}{3}\Lambda$ (partially massless limit)

Action in canonical form

$$S[g,\phi,f] = \frac{\widetilde{M}_p^2}{2}\int d^4x\sqrt{-g}(R-2\Lambda) + S_0[g,\phi] + S_2[g,f],$$

$$S_0[g,\phi] = \int d^4x\sqrt{-g}\left[-\frac{1}{2}\nabla^\mu\phi\nabla_\mu\phi - \frac{m_0^2}{2}\frac{3\overline{M}_p^2}{2}\left(1-e^{\sqrt{2/3}\phi/\overline{M}_p}\right)^2\right]\Bigg|_{g-2f/\widetilde{M}_p}$$

$$S_2[g,f] = -S_{PF}[g,f] - \int d^4x\sqrt{-g}\left[\left(2f_\mu^\rho f_{\rho\nu} - f f_{\mu\nu}\right)R^{\mu\nu} + \left(\Lambda - \frac{R}{2}\right)\left(f_{\mu\nu}f^{\mu\nu} - \frac{1}{2}f^2\right)\right]$$

$$-\frac{1}{2}\frac{m_2^2}{\widetilde{M}_p}\int d^4x\sqrt{-g}\left[5f_{\mu\nu}f^{\mu\nu}f - 4f^{\mu\nu}f_\mu^\rho f_{\rho\nu} - f^3\right]$$

$$+\frac{8}{3}\frac{1}{M_p^2}\frac{1}{\widetilde{M}_p}\int d^4x d^4y d^4z\frac{\delta^{(3)}S_{EH}}{\delta g_{\mu\nu}(x)\delta g_{\rho\sigma}(y)\delta g_{\alpha\beta}(z)}f_{\mu\nu}(x)f_{\rho\sigma}(y)f_{\alpha\beta}(z) + O(f^4)$$

$$\overline{M}_p^2 \equiv M_p^2 + \frac{4}{3}\alpha\Lambda$$

$$\widetilde{M}_p^2 \equiv M_p^2 + \frac{2}{3}(2\alpha+\beta)\Lambda$$

Couplings

spin-2 sector coupling:

$$\frac{1}{\widetilde{M}_p} = \frac{1}{M_p}\left(\frac{1}{1 + 2\Lambda(2\alpha + \beta)/3M_p^2}\right)^{\frac{1}{2}}$$

spin-0 sector coupling:

$$\frac{1}{\overline{M}_p} = \frac{1}{M_p}\left(\frac{1}{1 + 4\Lambda\alpha/3M_p^2}\right)^{\frac{1}{2}}$$

Couplings dependence on Λ $\Longrightarrow$ additional dependence on α, β !

Features of the limit $\beta \longrightarrow \infty$

$$S = \frac{1}{2}\int d^4x\sqrt{-g}\left(M_p^2(R-2\Lambda)+\frac{\alpha}{6}R^2-\frac{\beta}{2}C_{\mu\nu\rho\sigma}C^{\mu\nu\rho\sigma}\right)$$

$$\lim_{\beta\to\infty} S = \ ?$$

- The limit distinguishes $\Lambda = 0$ & $\Lambda \neq 0$
  (couplings, particle spectrum, enhanced gauge symmetry)
- Limit is regular only in $D = 4$
- When $\Lambda \neq 0$ the resulting theory is much simpler

Result of the limit $\beta \longrightarrow \infty$ $(\Lambda > 0)$

- Massless spin-2 and $\pm 2, \pm 1$ helicities of massive spin-2 ghost decouple
- Massive spin-0 $(\tilde{\phi})$ & helicity-0 $(\tilde{\chi})$ of spin-2 ghost survive

Result of the limit $\beta \longrightarrow \infty$ $(\Lambda > 0)$

- Massless spin-2 and $\pm 2, \pm 1$ helicities of massive spin-2 ghost decouple
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$$S_{\phi\chi}[g, \tilde{\phi}, \tilde{\chi}] = \int d^4x \sqrt{-g} \left[\frac{1}{2} \left(\partial_\mu \tilde{\chi} \partial^\mu \tilde{\chi} - \partial_\mu \tilde{\phi} \partial^\mu \tilde{\phi} \right) - V(\tilde{\phi}, \tilde{\chi}) \right]$$

$$V(\tilde{\phi}, \tilde{\chi}) = \frac{\Lambda}{36 \bar{M}_p^2} \left(\tilde{\chi}^2 - \tilde{\phi}^2 - 6 \bar{M}_p^2 \right)^2 + \frac{m_0^2}{12 \bar{M}_p^2} \tilde{\phi}^2 \left(\tilde{\chi} + \tilde{\phi} \right)^2$$

$$\bar{M}_p^2 = M_p^2 + \frac{4}{3} \alpha \Lambda$$