Searching for ALP Dark Matter with a 1000 km baseline interferometer

19th Patras Workshop on axions, WIMPs and WISPs

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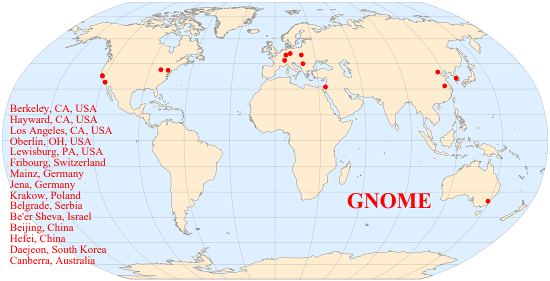
Contents

What is a GNOME?[1]

- **G**lobal **N**etwork of **O**ptical **M**agnetometers for **E**xotic physics searches
- Looking for transient and background dark matter signals
- Sensitive to Axion-nucleon coupling:

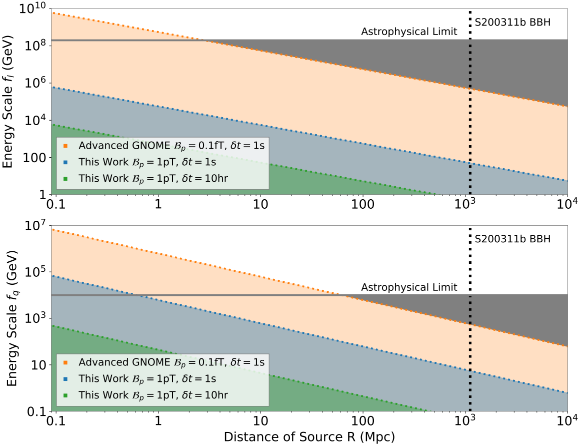
$$\mathcal{H}_N = g_{aNN} \nabla a \cdot \sigma_N \,,$$
$$\mathcal{H}_P = g_{aPP} \nabla a \cdot \sigma_P \,,$$

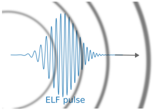


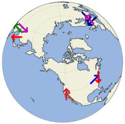
[1] Phys.Dark Univ. 22 (2018), 162-180

What can a GNOME do?[2] Look for ELFs[3]

- Exotic Low-mass Field (ELF) search with multi-messenger astronomy







- High energy astrophysical events detected by GW detectors

[2]Afach et al. ANNALEN DER PHYSIK 2023, 2300083

[3]Khamis et al. arXiv: 2407.13919

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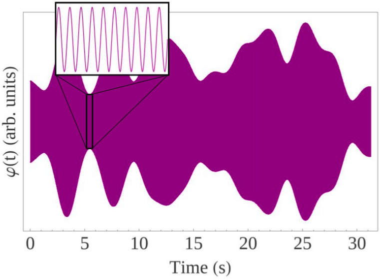
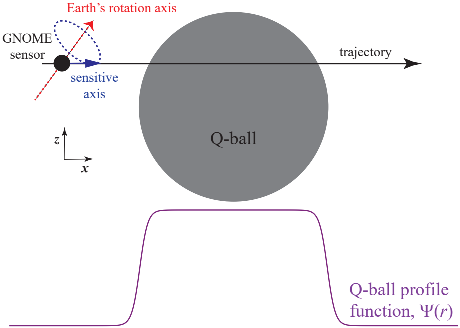
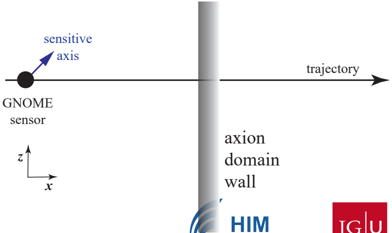
What can a GNOME do?

- Stochastic ALP DM field fluctuations
- Axion Domain Walls[4]
- Q-balls
- and much more!







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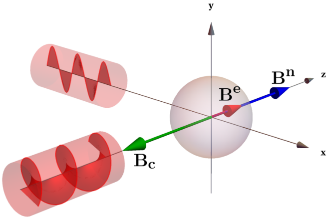
[4]Afach et al. Nat. Phys. 17, 1396–1401 (2021).

Advanced GNOME: K-Rb-^{3}He Comagnetometer

- Hot vapour cell with K, Rb and He magnetically shielded
- Polarize Rb electron $\rightarrow$ K electron and He nucleus polarization



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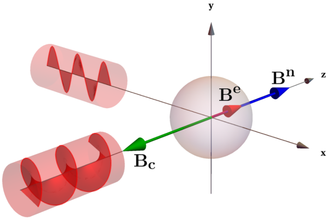
- Apply a compensation field
- More sensitive to spin couplings, including rotations and exotic interactions



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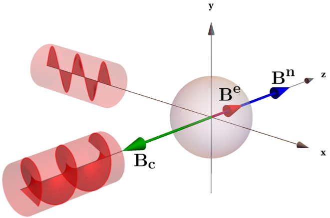
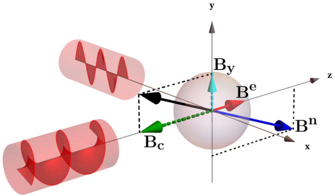


Perturbation



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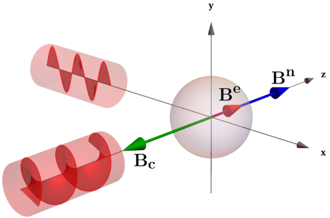

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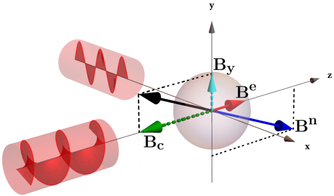
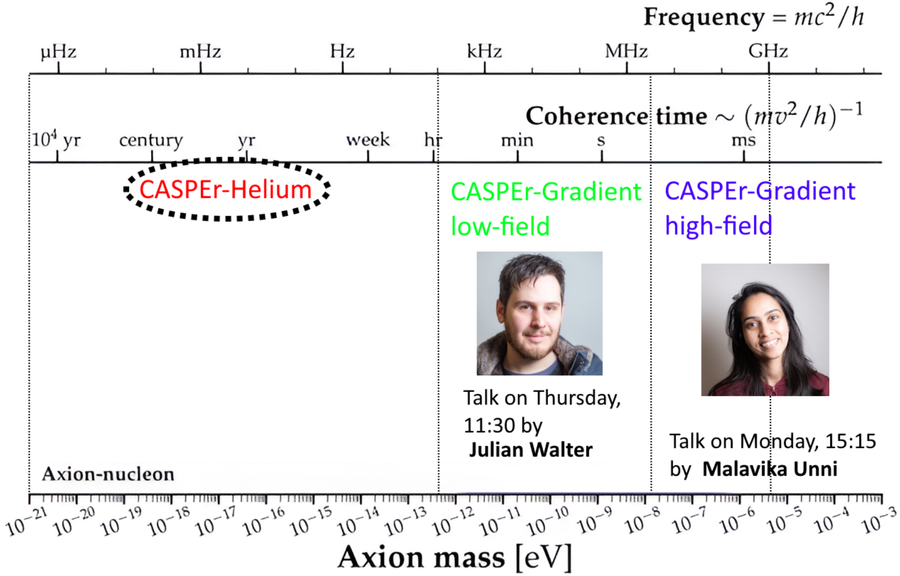


How sensitive are they?

Most stringent constraints on ALP DM at $\mathcal{O}(1)$ Hz. What about lower frequencies?

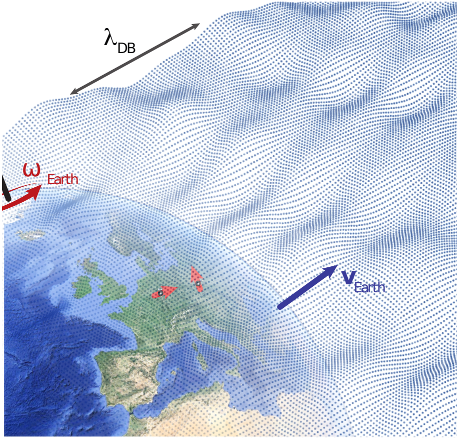
Contents

The Cosmic Axion Spin Precession Experiment (CASPEr)



Two comagnetometers as an interferometer

- Situated in Mainz and Krakow, $\sim 1000\,\mathrm{km}$ apart
- Time synchronized measurement
- Lower frequency regime $\rightarrow$ coherent signal
- We calibrate the frequency response of the comagnetometers[5] every 25 h



[5] Padniuk et al. Phys. Rev. Research 6, 013339

3D Gradient of the ALP field

- Spread of frequencies

$$\Delta\omega \approx \omega_a \frac{v_0^2}{c^2} \approx \omega_a \times 10^{-6}$$

- Coherence time
  $\tau \sim 1/\Delta\omega$

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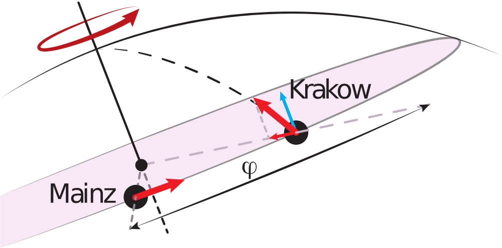

$$\nabla a(t) \sim \sum_n^N \boldsymbol{v}_n \cos(\omega_a t + \phi_n)$$

$$= \hat{\boldsymbol{x}}\alpha_x \cos(\omega_a t + \phi_x) + \hat{\boldsymbol{y}}\alpha_y \cos(\omega_a t + \phi_y) + \hat{\boldsymbol{z}}\alpha_z \cos(\omega_a t + \phi_z)$$

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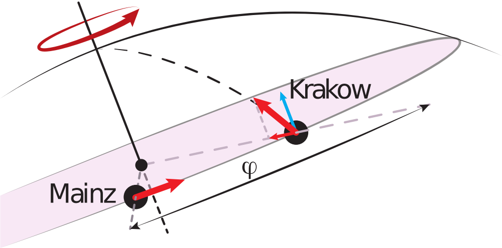

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- $\nabla a(t)$ depends on six random parameters:
  - α: Rayleigh distributed random number
  - ϕ: phase of the field in each orthogonal direction

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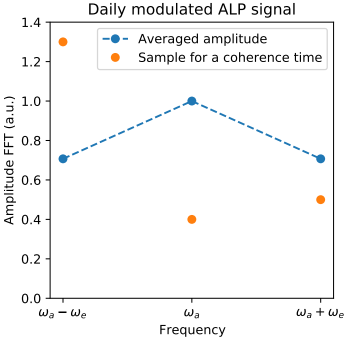
ALP signature in the frequency domain

- A carrier at ω_a and two sidebands at $\omega_a \pm \omega_e$



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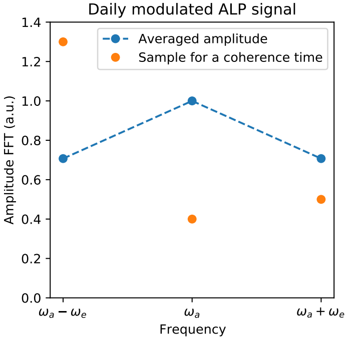
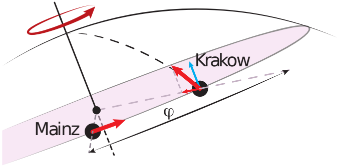
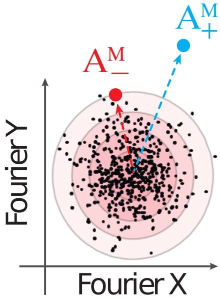


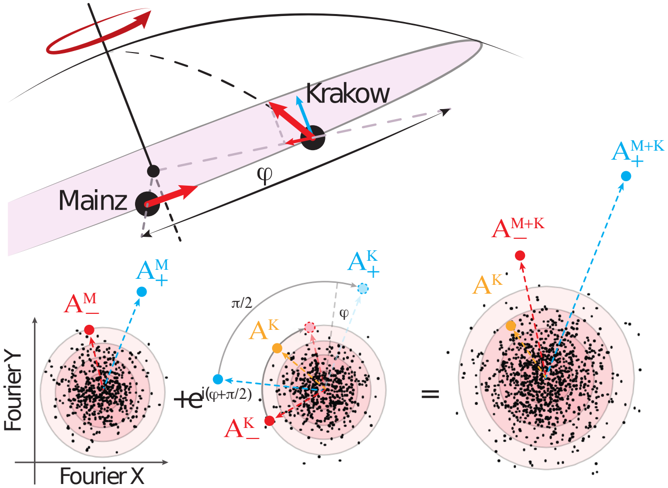




- Sidebands are in general asymmetric!

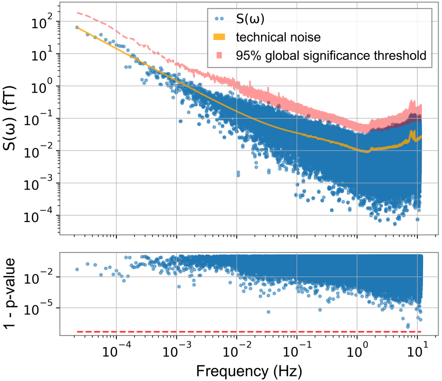
Search strategy

- We combine the ALP signatures properly shifted



Search results

- No ALP candidate is found



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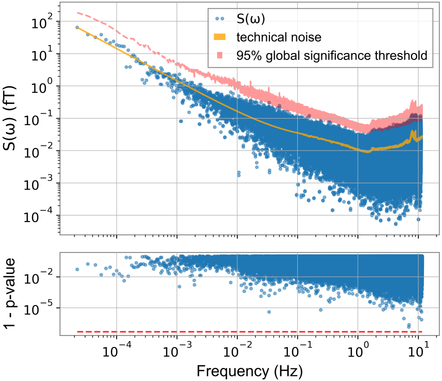
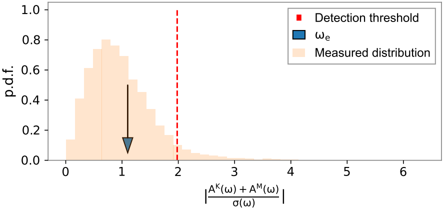
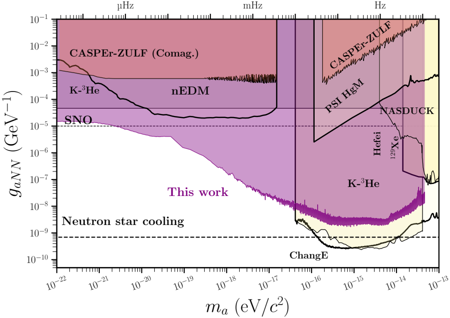


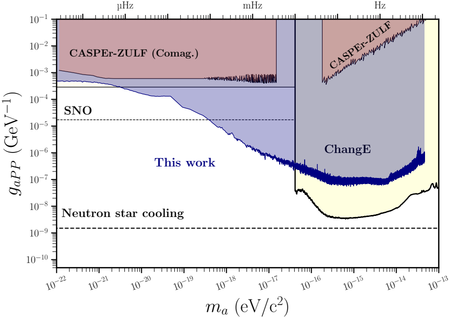
- Independent analysis of amplitudes at ω_e



Exclusions plots for proton and neutron coupling

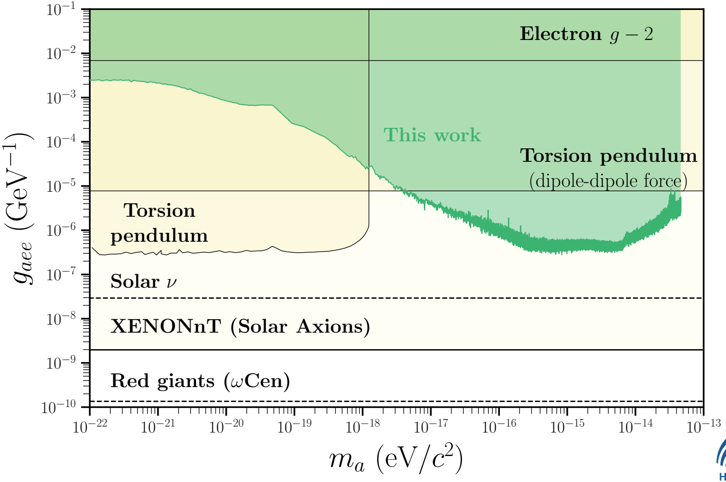
- Constraints rescaled by nuclear spin content of ^{3}He
- Reduction of sensitivity due to incoherence of the field for frequencies $> 10^{-2}$ Hz





Exclusion plot for electron coupling

- Assuming that electrons in the shield generates a magnetic field in opposite direction



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Outlook

- We present a search in the ultra-low ALP mass range
- It extends for nine orders of magnitude in laboratory unconstrained space in neutron, proton and electron coupling.
- The experimental set up is based on two comagnetometers in separate locations.
- This work is part of the CASPEr family of experiments looking for ALP halo through spin interactions
- The comagnetometers are part of Advanced GNOME and will run together as a network to look for transient DM events (ELFs, axion domain walls, Q-balls, ...)

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Acknowledgements

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Dmitry Budker

Alexander Sushkov

Szymon Pustelny

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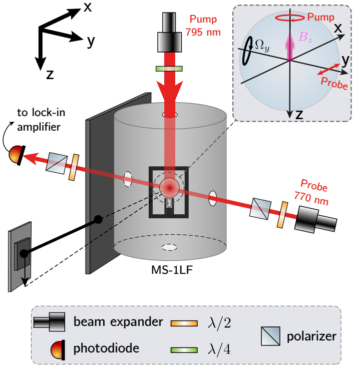
Polarization dynamics in a comagnetometer

- Frequency response for arbitrary perturbation:

$$\mathcal{F}^r_\pm = -a\frac{\omega_n(\alpha_e - \alpha_n) + (\pm\omega + \gamma_n\Delta_{B_z} - i|R_n|)\alpha_e}{(\pm\omega + \omega_e + \Delta_{B_z}\gamma_e/q - i|R_e|)(\pm\omega + \omega_n + \gamma_n\Delta_{B_z} - i|R_n|) - \omega_e\omega_n}$$



- α = interaction coupling
- a = amplitude
- Δ_{B_z} = detuning from compensation point
- R = Relaxation rate
- ω = Larmor frequency

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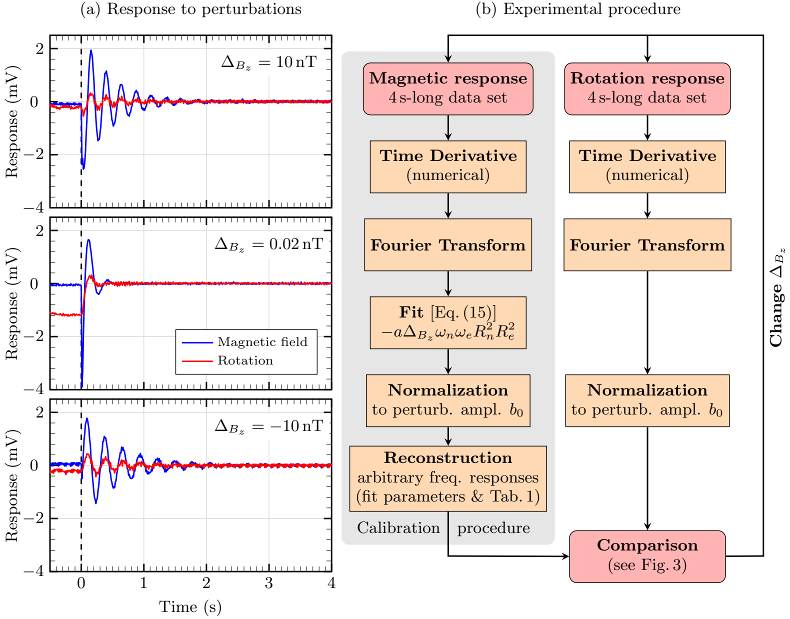
Comagnetometer response calibration routine



Comagnetometer response demonstration