

MOND: An alternative to particle dark matter

Federico Lelli

INAF - Arcetri Astrophysical Observatory



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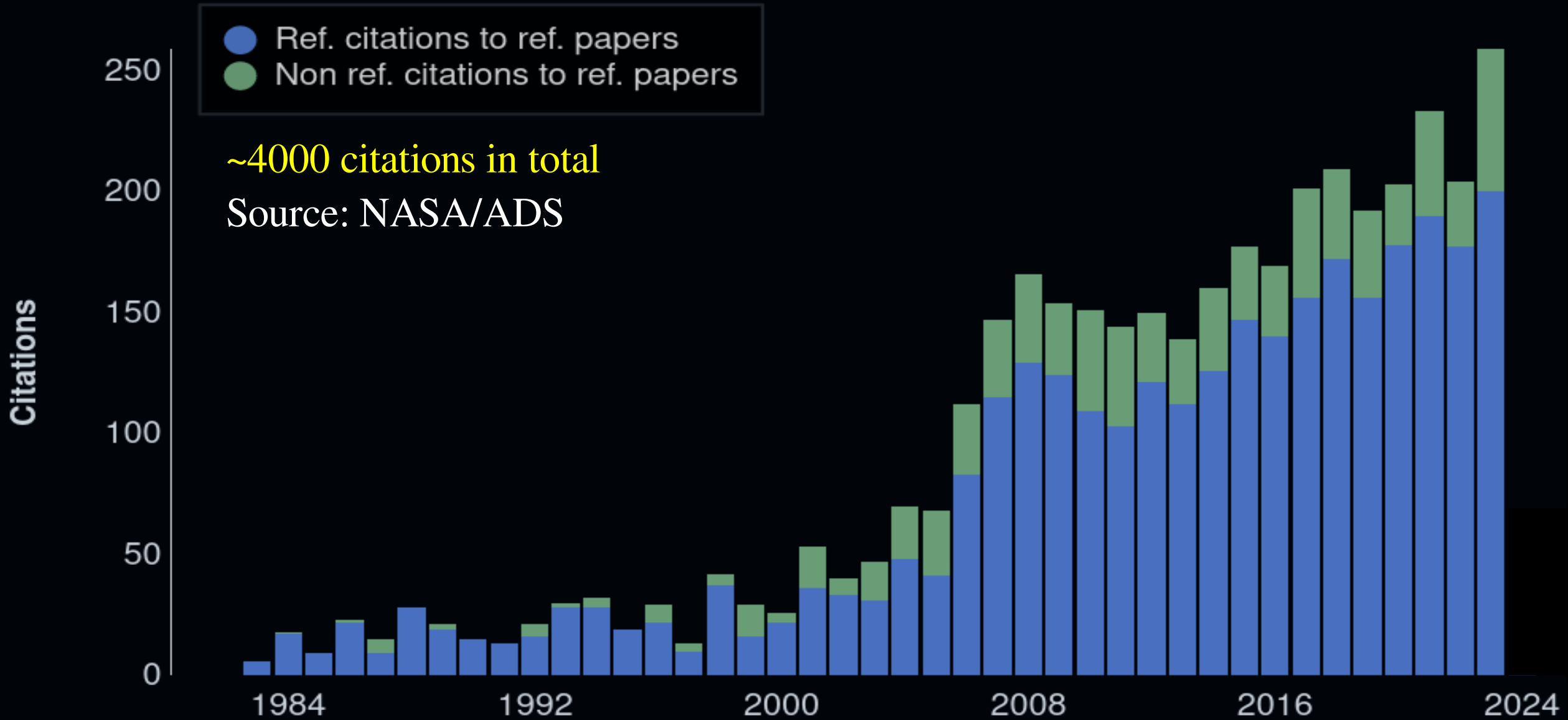


Proposed by Moderhai Milgrom (1983a, b, c)

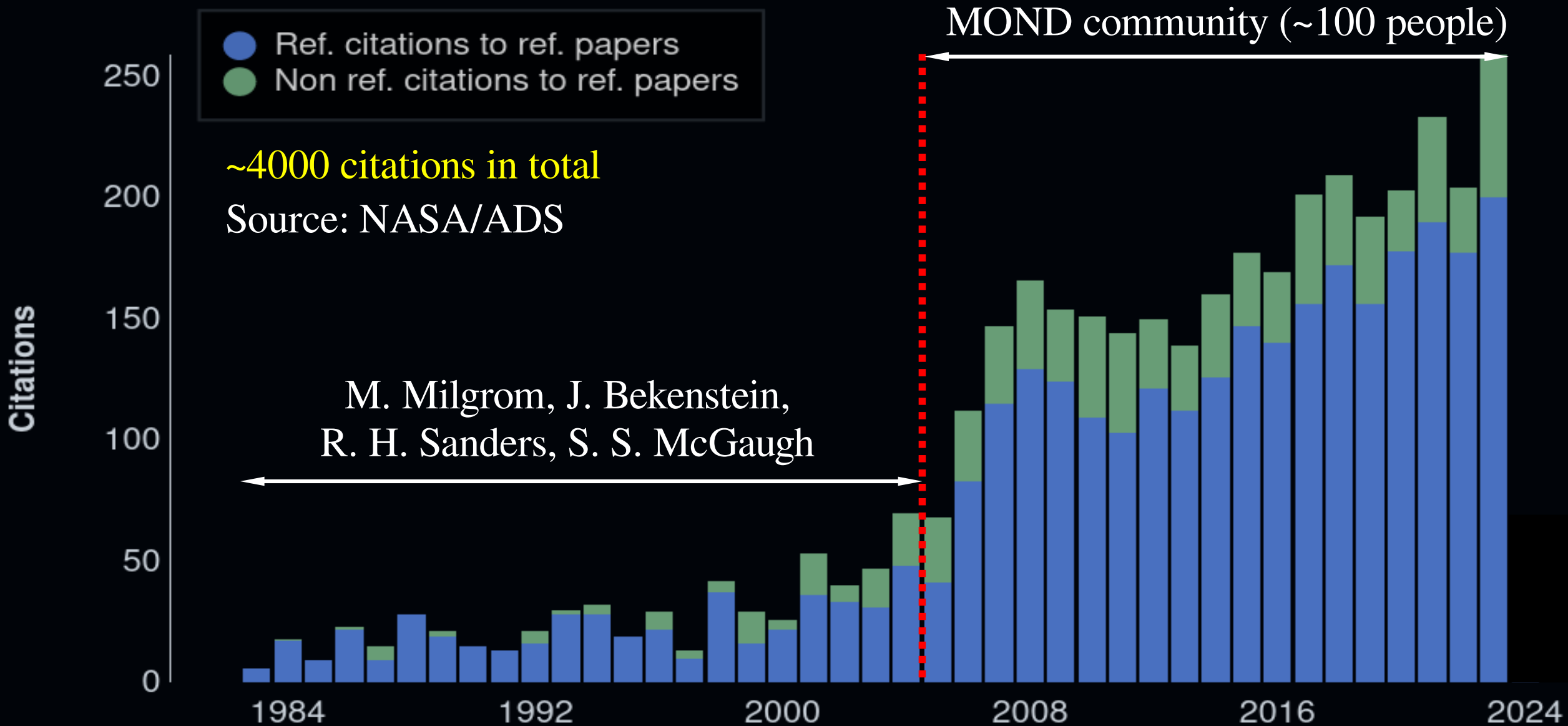
MOND still stands after more than 40 years.

Multifaceted paradigm including different theories
at both the non-relativistic and relativistic levels.

Citations to the original MOND trilogy (Milgrom 1983a, b, c)



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Talk Outline

1. General MOND paradigm
2. MOND in nearby galaxies
3. MOND in galaxy clusters
4. MOND cosmology
5. Specific MOND theories

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MOND postulates at the non-relativistic level

1) **New constant of Physics:** $a_0 \simeq 10^{-10} \text{ m/s}^2$

similar role as c in Relativity and \hbar in Quantum Mechanics

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$\vec{a} = \frac{d^2 \vec{x}}{dt^2}$ kinetic (observed) acceleration of a particle

$\vec{g}_N = -\vec{\nabla} \phi_N$ Newtonian gravitational field (from the Poisson's equation)

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Or impose scale invariance (Milgrom 2009, ApJ): $(\vec{x}, t) \rightarrow (\lambda \vec{x}, \lambda t)$ V is invariant!

A MODIFICATION OF THE NEWTONIAN DYNAMICS: IMPLICATIONS FOR GALAXIES¹

M. MILGROM

Department of Physics, Weizmann Institute, Rehovot, Israel; and The Institute for Advanced Study

Received 1982 February 4; accepted 1982 December 28

ABSTRACT

I use a modified form of the Newtonian dynamics (inertia and/or gravity) to describe the motion of bodies in the gravitational fields of galaxies, *assuming that galaxies contain no hidden mass*, with the following main results.

- ① The Keplerian, circular velocity around a finite galaxy becomes independent of r at large radii, thus resulting in asymptotically flat velocity curves.
- ② The asymptotic circular velocity (V_∞) is determined only by the total mass of the galaxy (M): $V_\infty^4 = a_0 GM$, where a_0 is an acceleration constant appearing in the modified dynamics. This relation is consistent with the observed Tully-Fisher relation if one uses a luminosity parameter which is proportional to the observable mass.
- ③ The discrepancy between the dynamically determined Oort density in the solar neighborhood and the density of observed matter disappears.
- ④ The rotation curve of a galaxy can remain flat down to very small radii, as observed, only if the galaxy's average surface density Σ falls in some narrow range of values which agrees with the Fish and Freeman laws. For smaller values of Σ , the velocity rises more slowly to the asymptotic value.
- ⑤ The value of the acceleration constant, a_0 , determined in a few independent ways is approximately $2 \times 10^{-8} (H_0/50 \text{ km s}^{-1} \text{ Mpc}^{-1})^2 \text{ cm s}^{-2}$, which is of the order of $CH_0 = 5 \times 10^{-8} (H_0/50 \text{ km s}^{-1} \text{ Mpc}^{-1}) \text{ cm s}^{-2}$.

The main predictions are:

- ① Rotation curves calculated on the basis of the *observed* mass distribution and the modified dynamics should agree with the observed velocity curves.
- ② The $V_\infty^4 = a_0 GM$ relation should hold exactly.
- ③ An analog of the Oort discrepancy should exist in all galaxies and become more severe with increasing r in a predictable way.

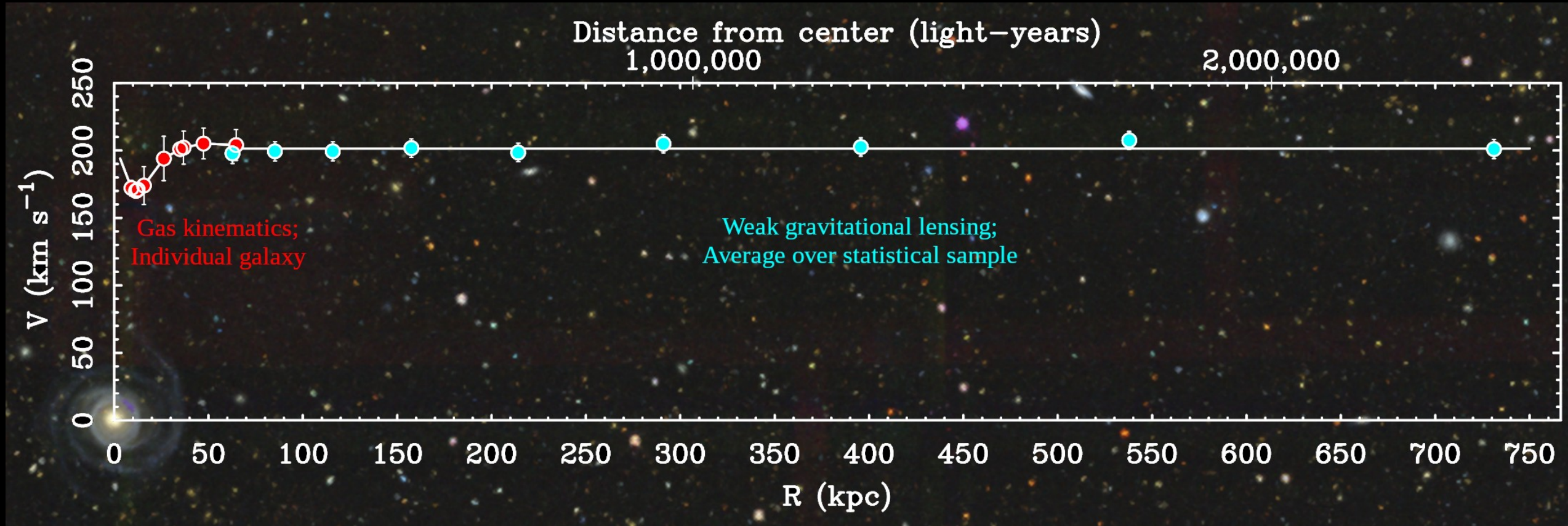


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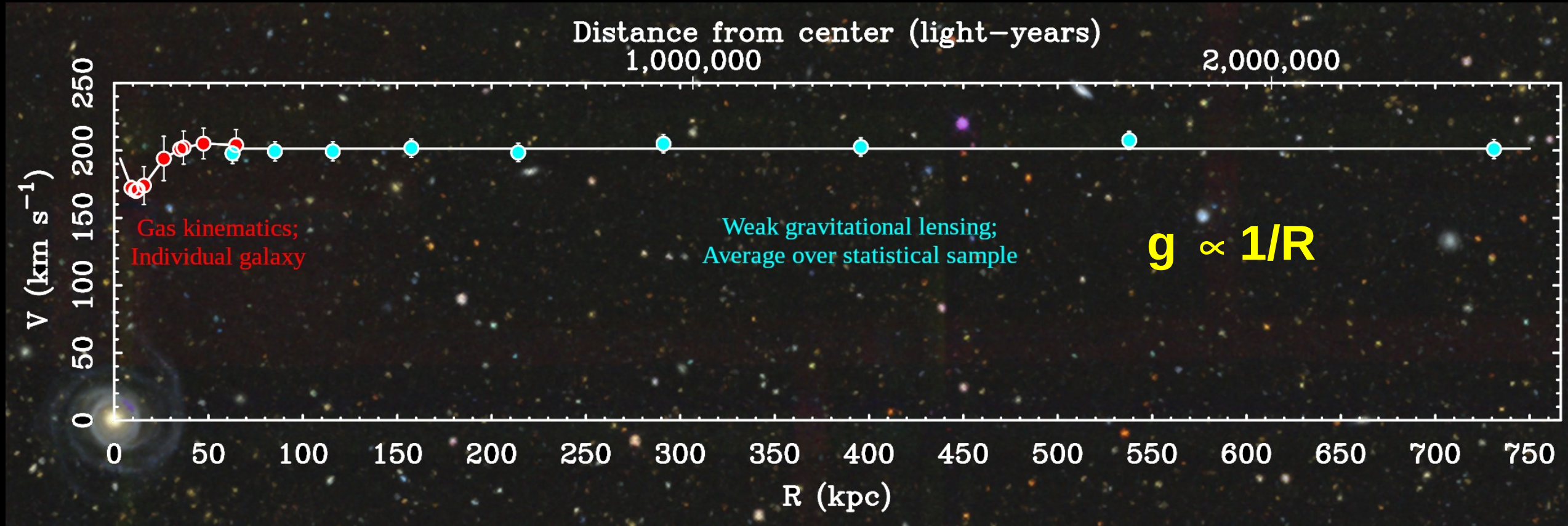
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See Mistele+2024, ApJ

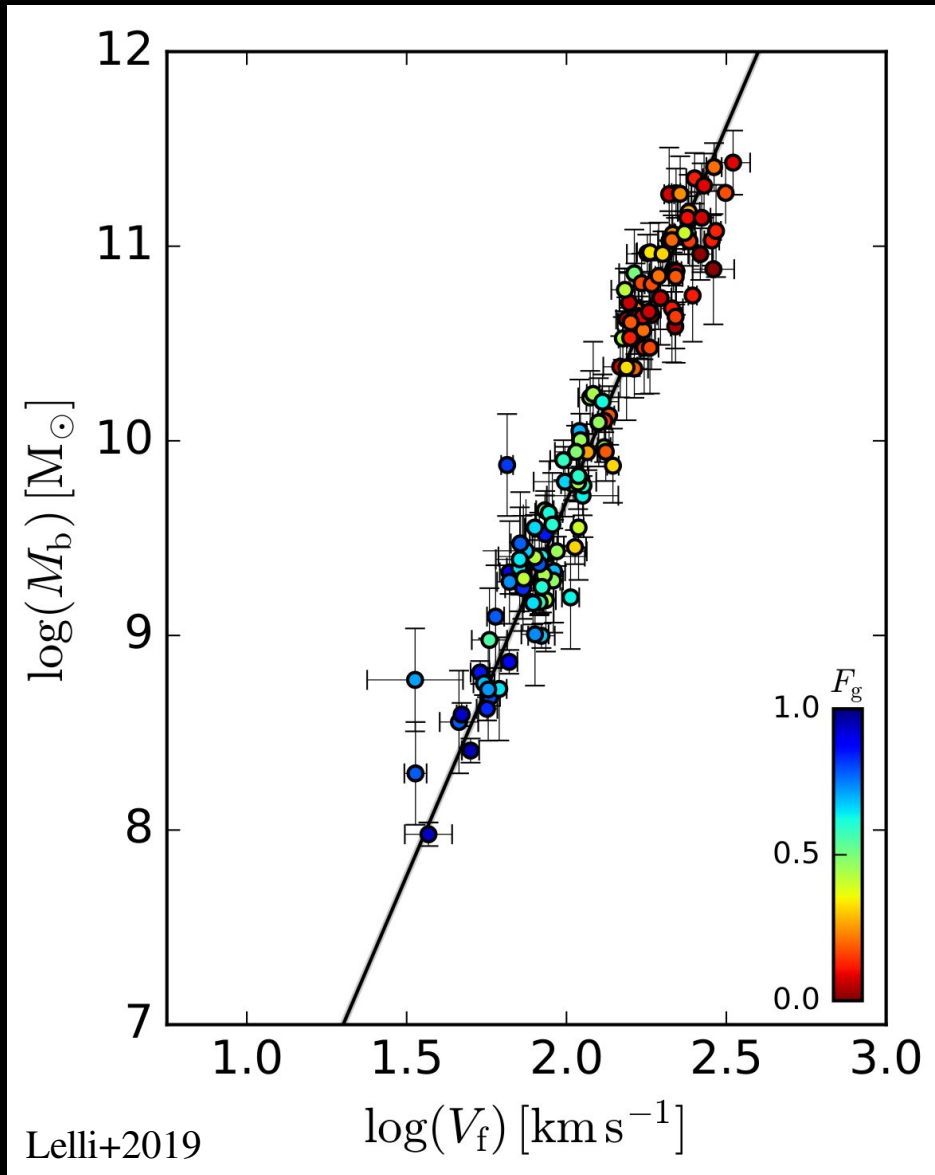
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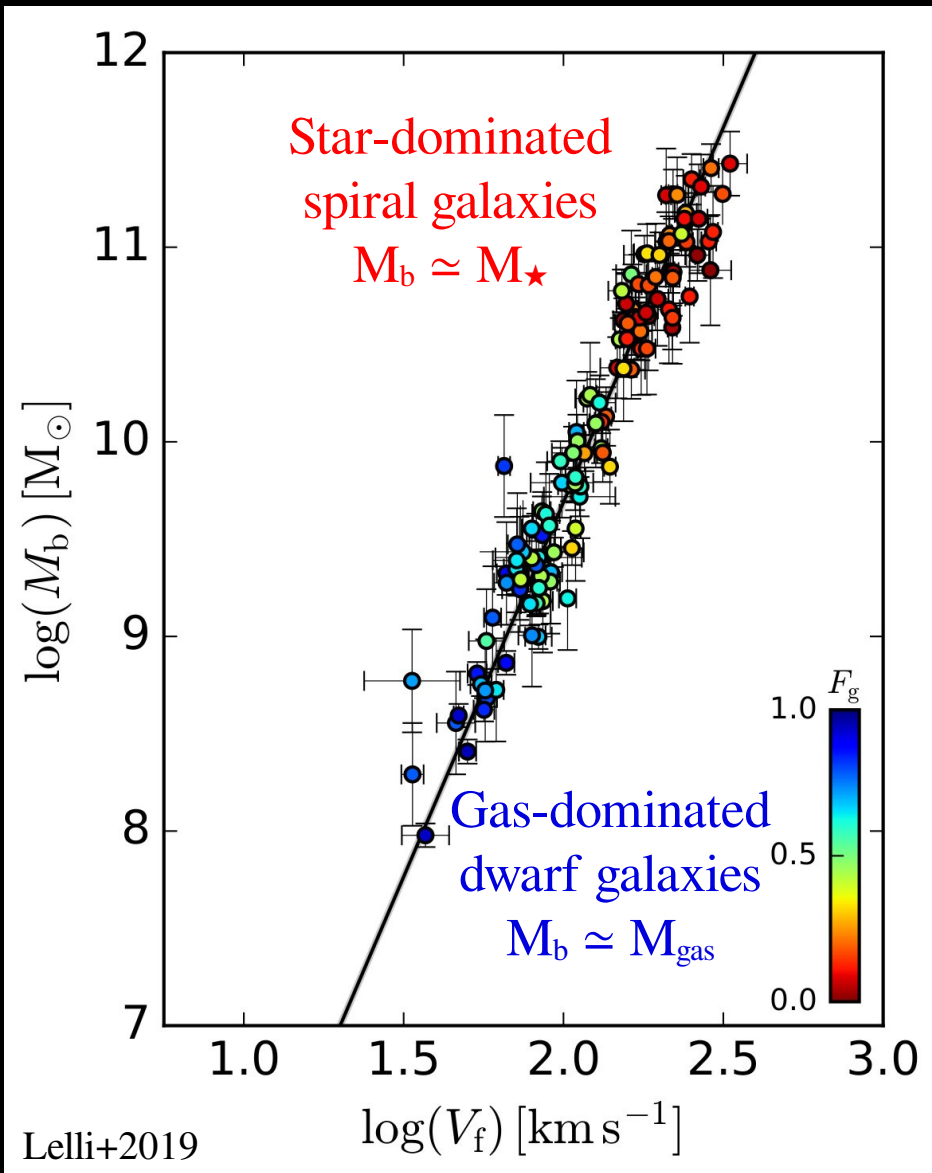


Tully-Fisher relation (1977, A&A):

HI linewidth vs Luminosity for bright spirals

McGaugh+(2000, 2005, 2010), Verheijen+2001, Lelli+(2016, 2019, 2022), Ponomareva+2018, Schombert+(2020), Di Teodoro (2021, 2022), and many more.

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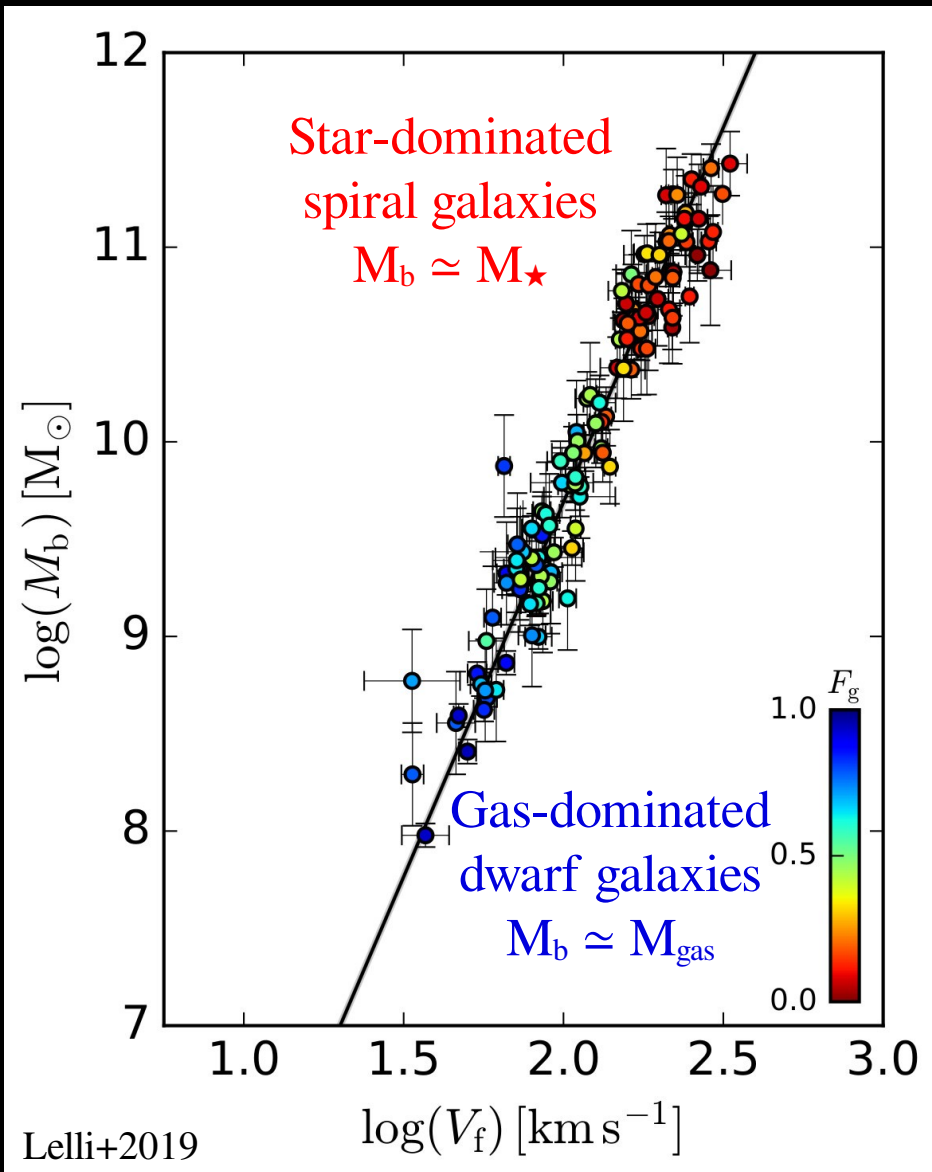
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Four independent predictions in one equation:

(i) Key quantities are V_f and M_b (stars+gas) \rightarrow OK

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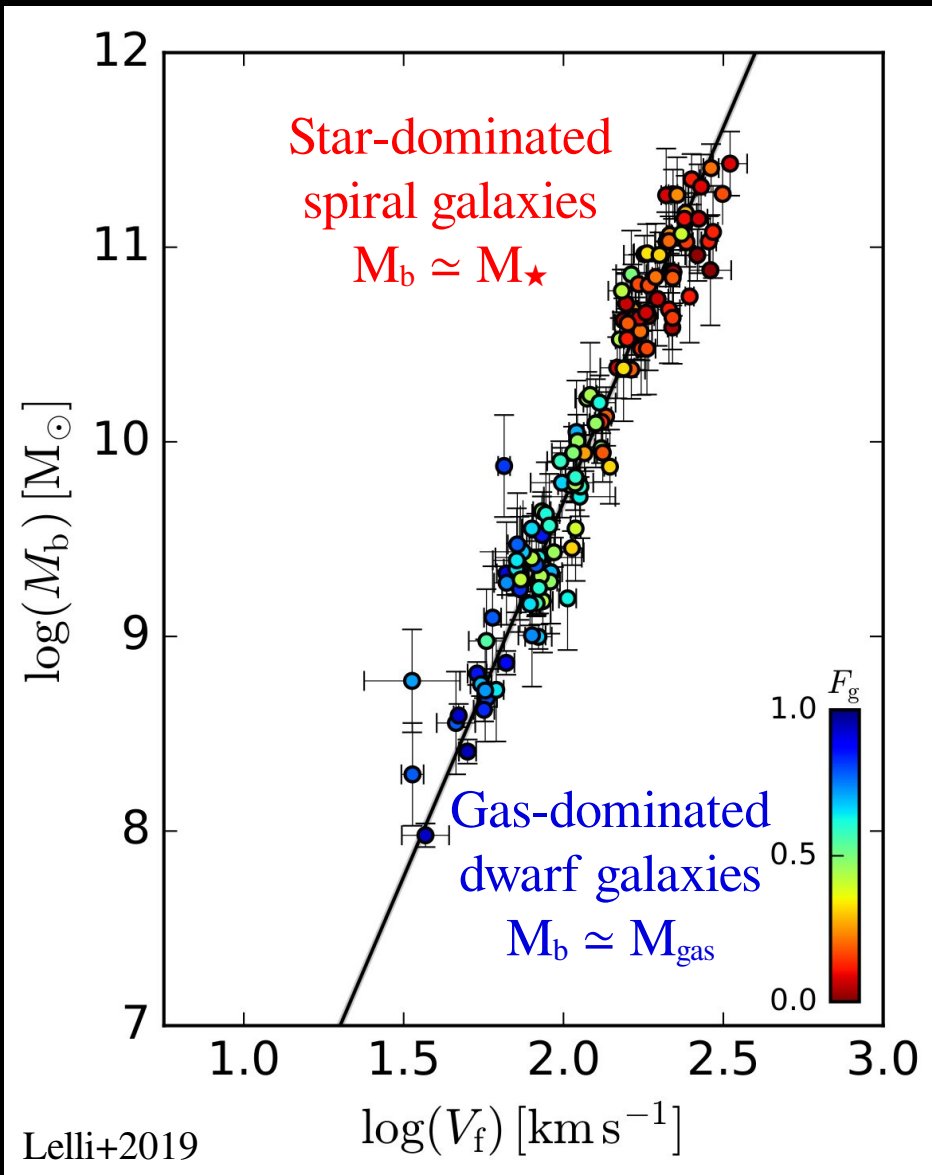
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(ii) Slope should be exactly 4 \rightarrow **OK**

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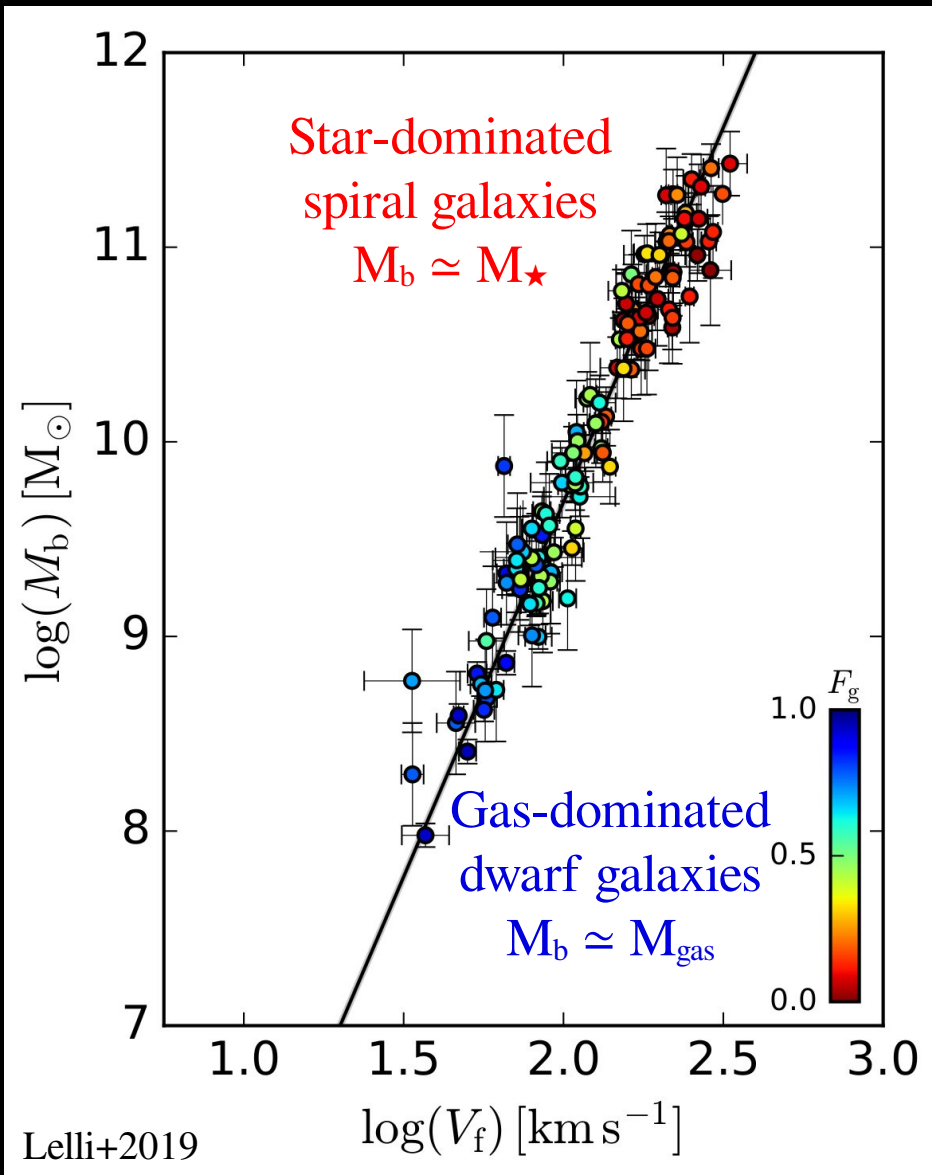
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(see Lelli 2022, Nature Astronomy)

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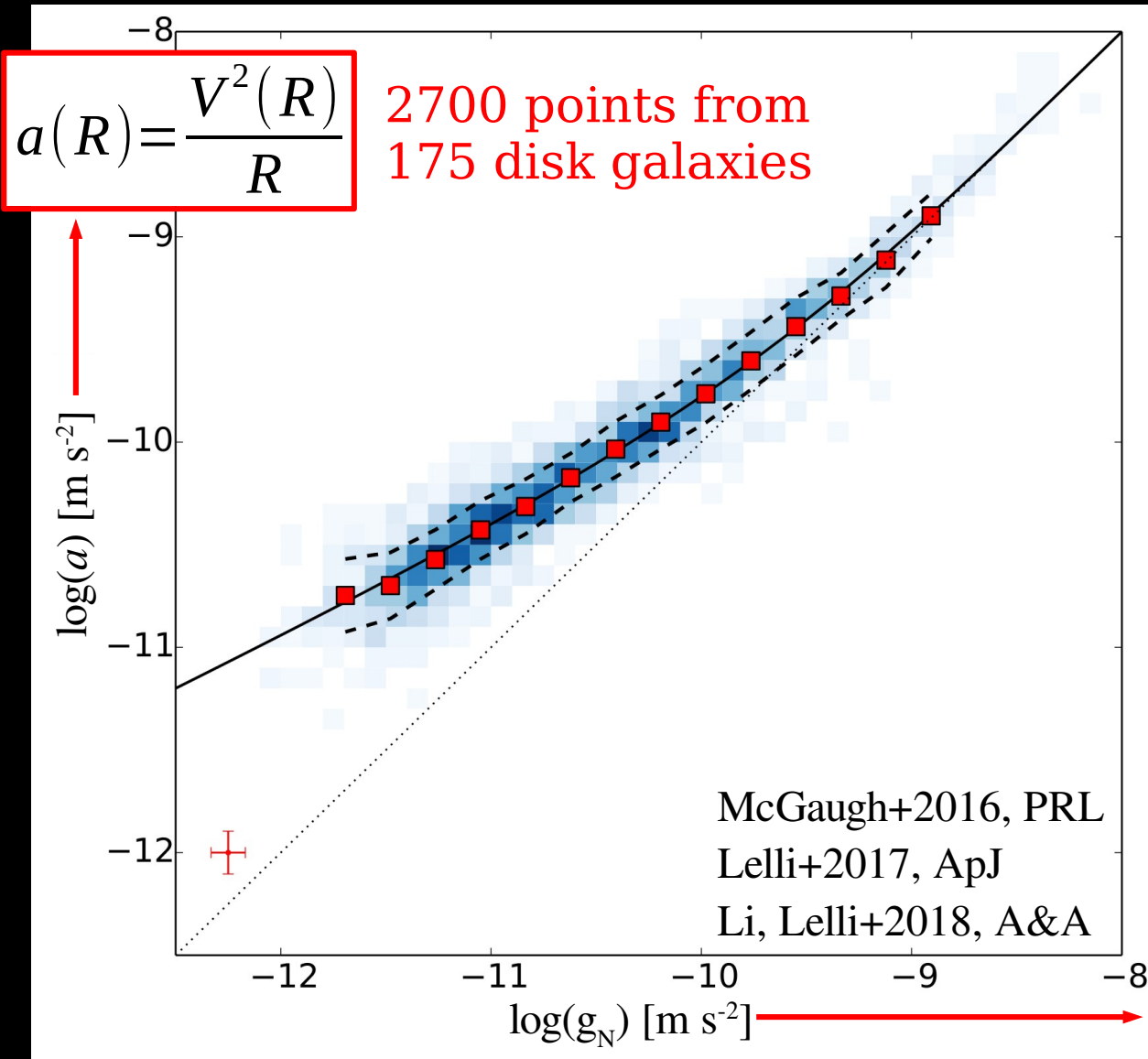
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- (iii) Normalization is $a_0 G$ \rightarrow **OK** with other estimates
(see Lelli 2022, Nature Astronomy)
- (iv) **NO** dependence on other properties (e.g., R_e , Σ_e),
so **NO** intrinsic scatter along the relation \rightarrow **OK**

McGaugh+(2000, 2005, 2010), Verheijen+2001, Lelli+(2016, 2019, 2022), Ponomareva+2018, Schombert+(2020), Di Teodoro (2021, 2022), and many more.

(3) Rotation curves can be predicted from the baryon distribution

Radial Acceleration Relation (RAR)

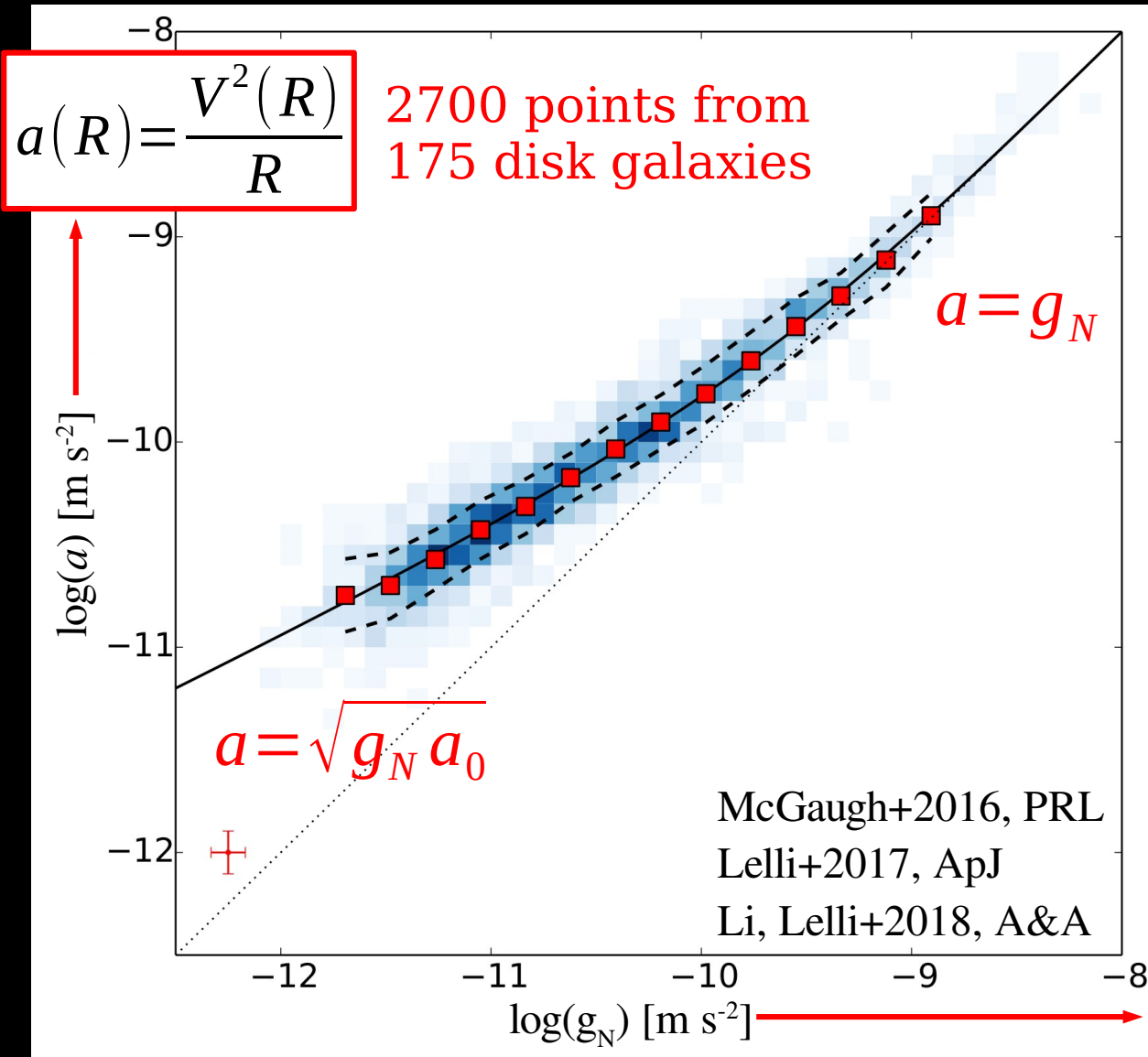
(i) Fully empirical, independent of MOND



$$\nabla^2 \Phi_N(R, z) = 4\pi G \rho_b(R, z)$$

$$g_N(R, z=0) = -\nabla \Phi_N(R, z=0)$$

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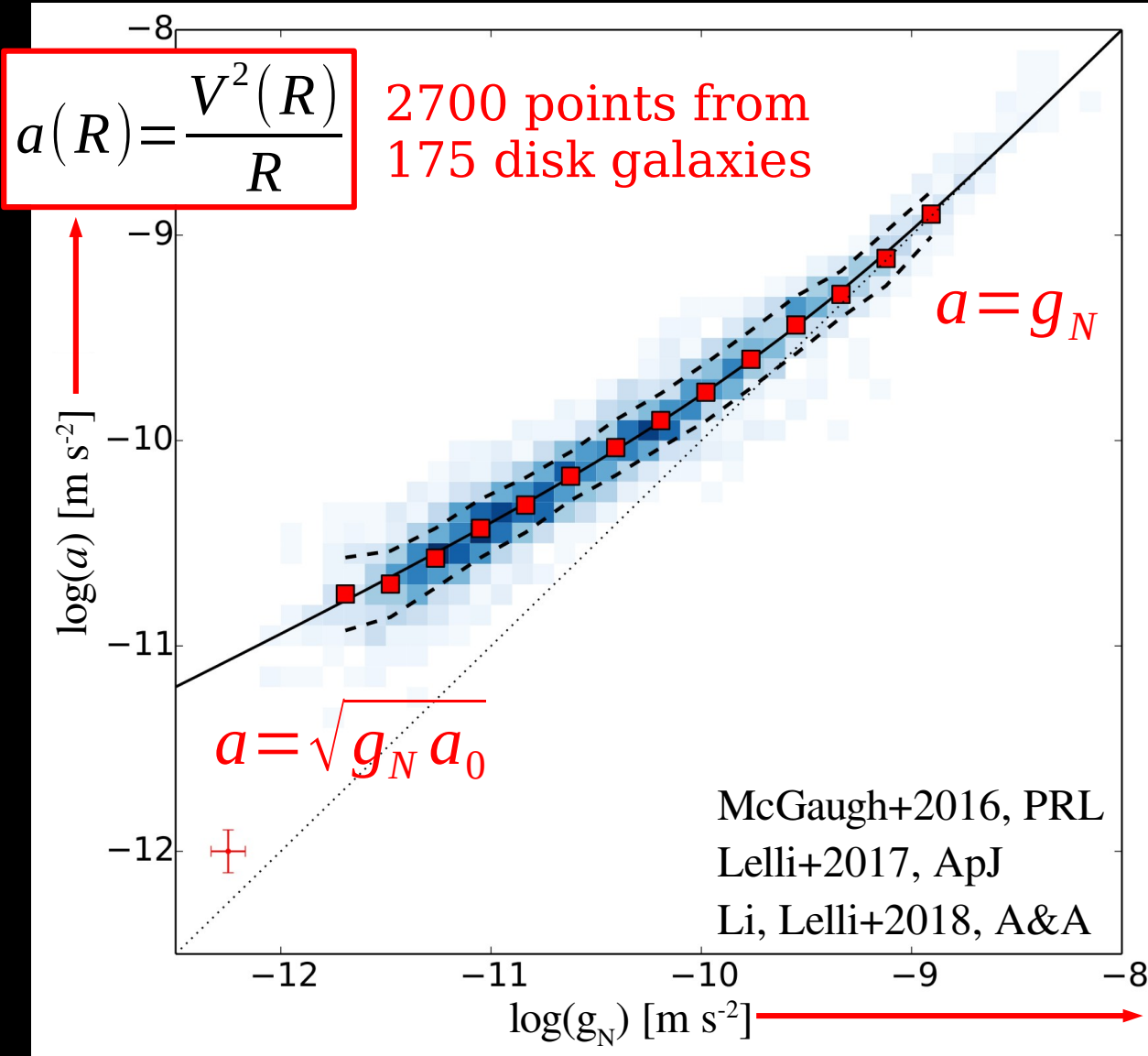
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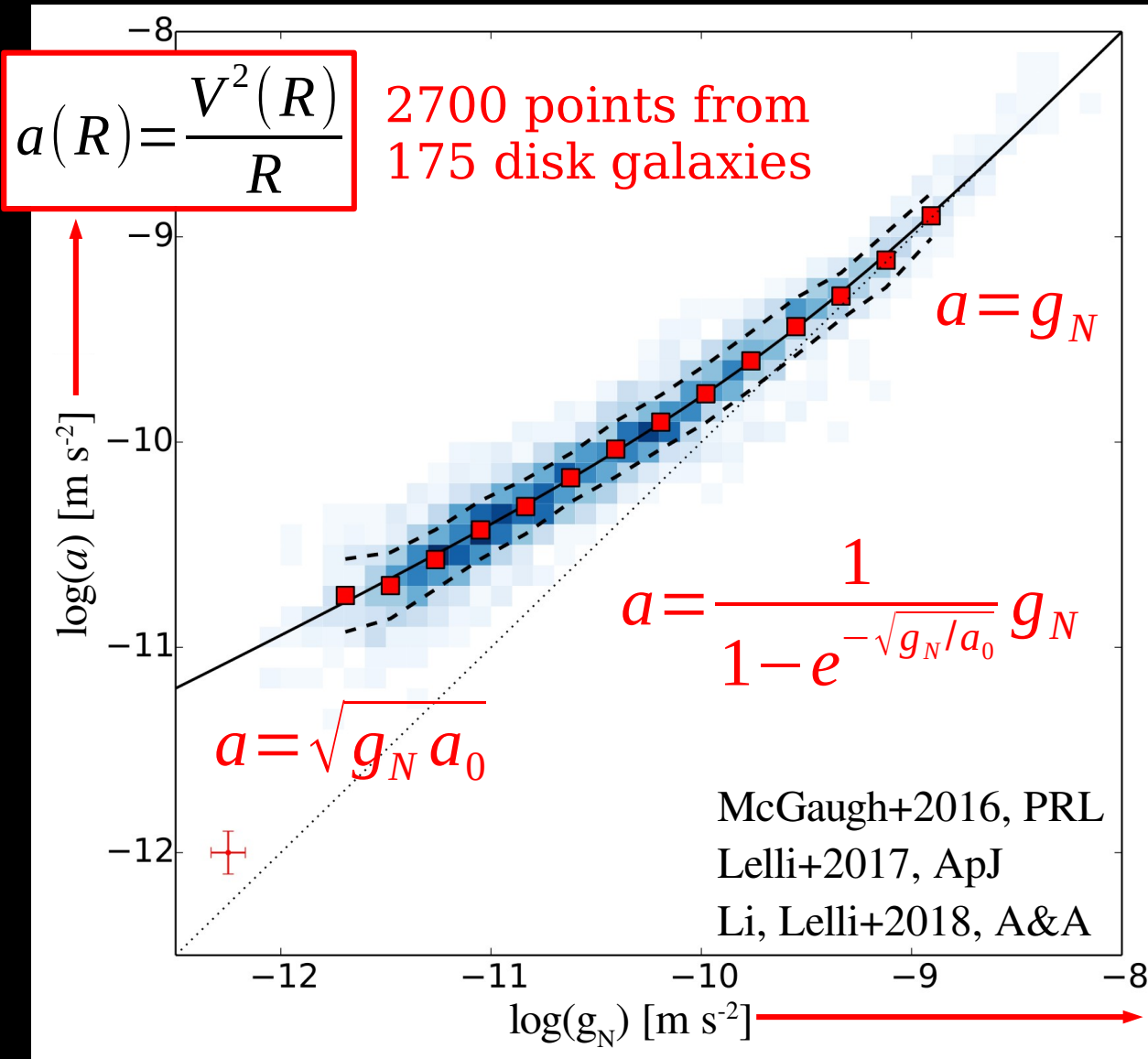
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- (iv) MOND interpolation function μ or ν

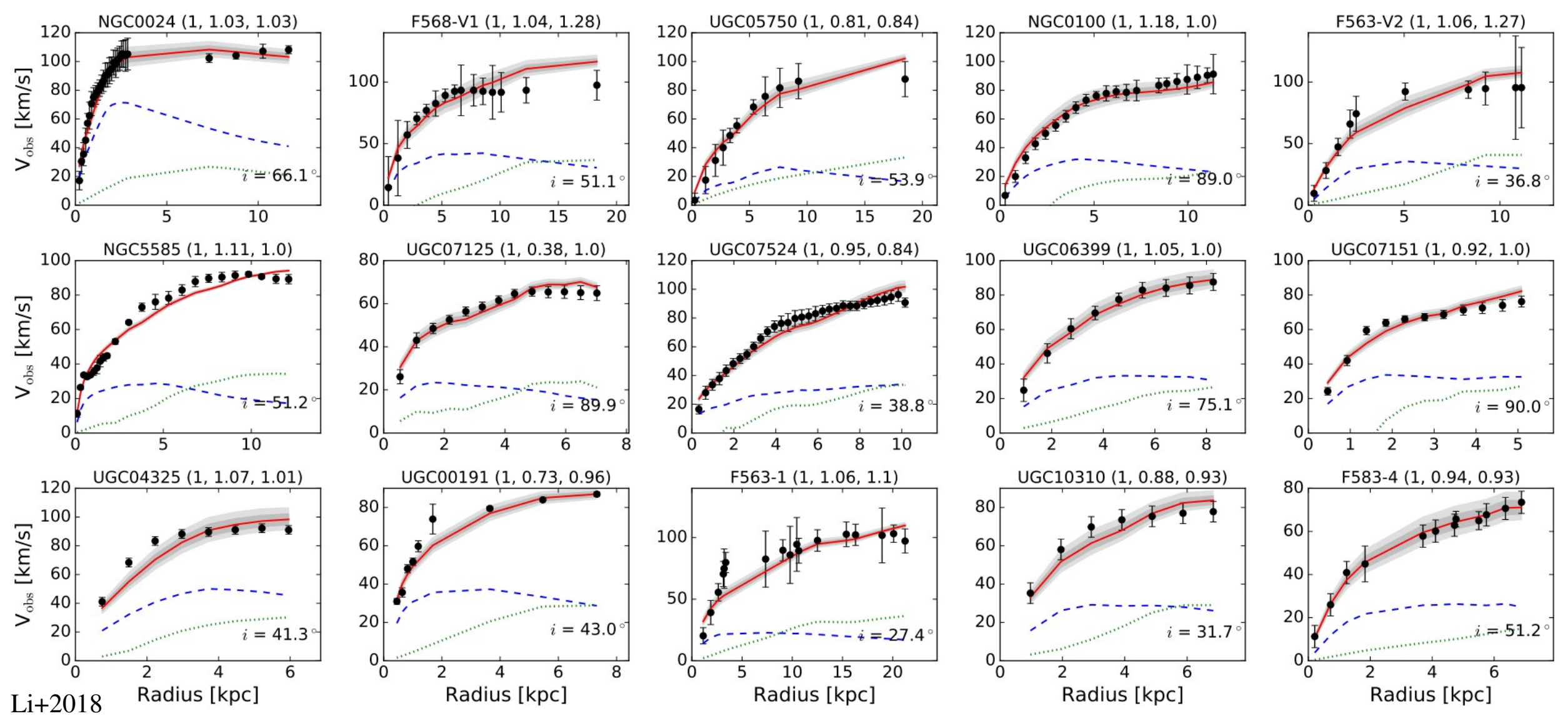
$$a \mu\left(\frac{a}{a_0}\right) = g_N \iff a = \nu\left(\frac{g_N}{a_0}\right) g_N$$

We can now assume $\nu(g_N/a_0)$ and predict rotation curves given ρ_b (within the errors)

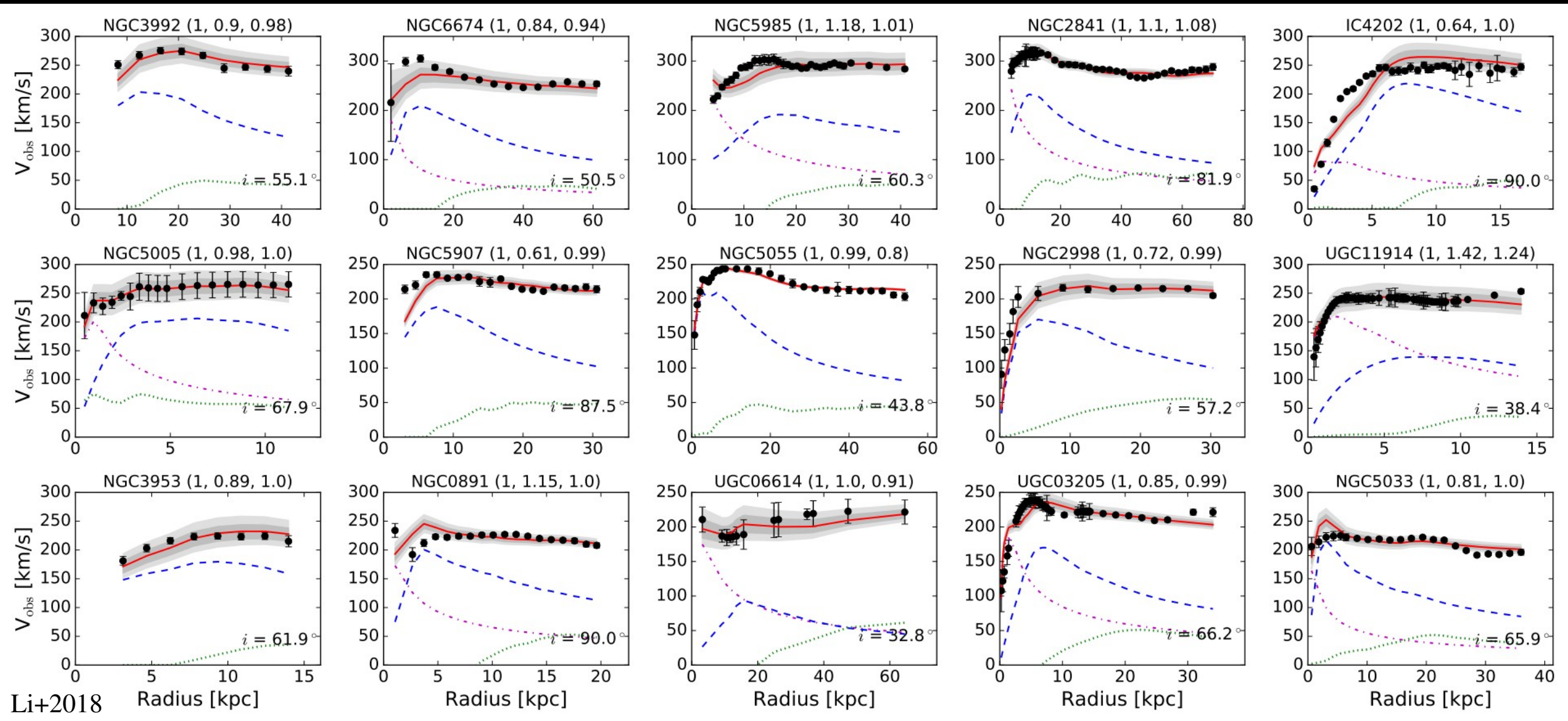
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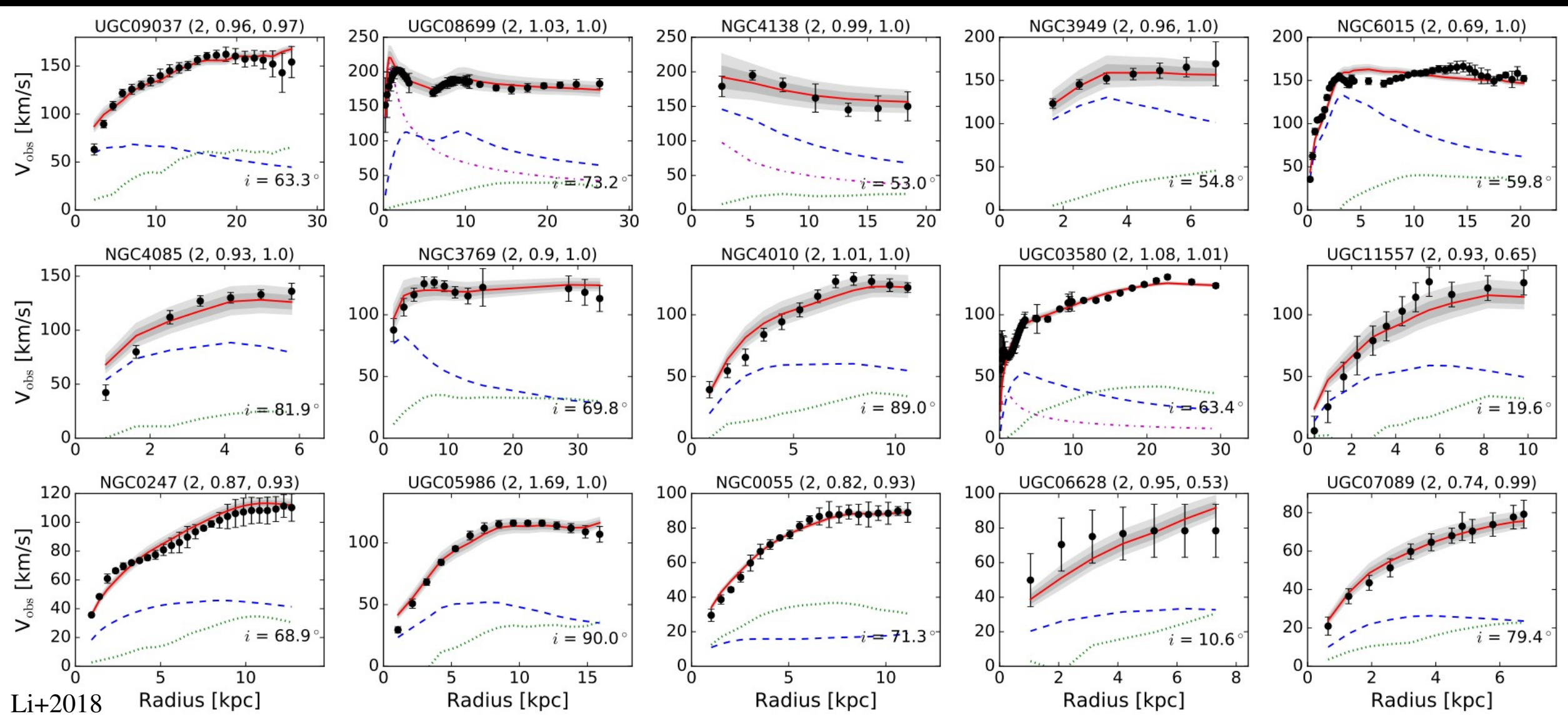


(3) Rotation curves can be predicted from the baryon distribution

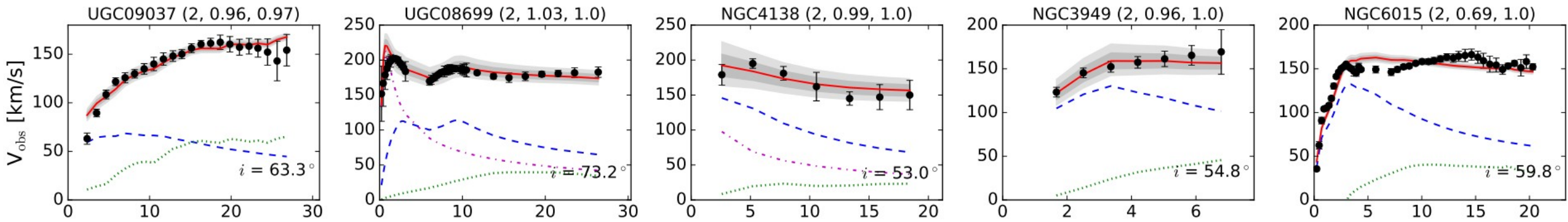


Li+2018 Radius [kpc]

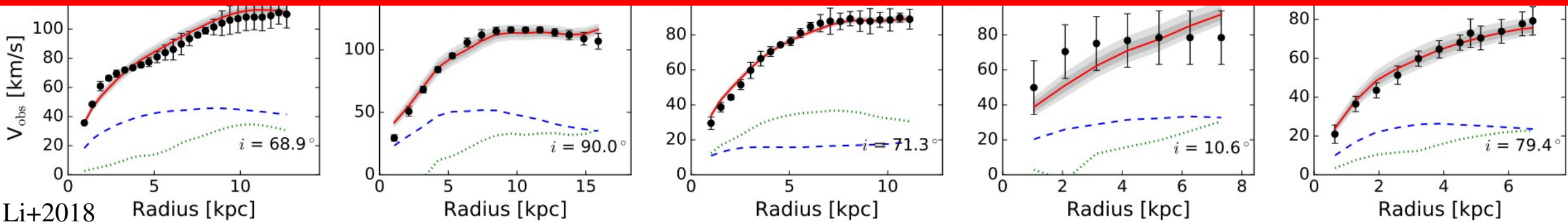
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For the full sample of 175 galaxies, see Li, Lelli, McGaugh et al. (2018) or the SPARC database (<http://astroweb.cwru.edu/SPARC/>)



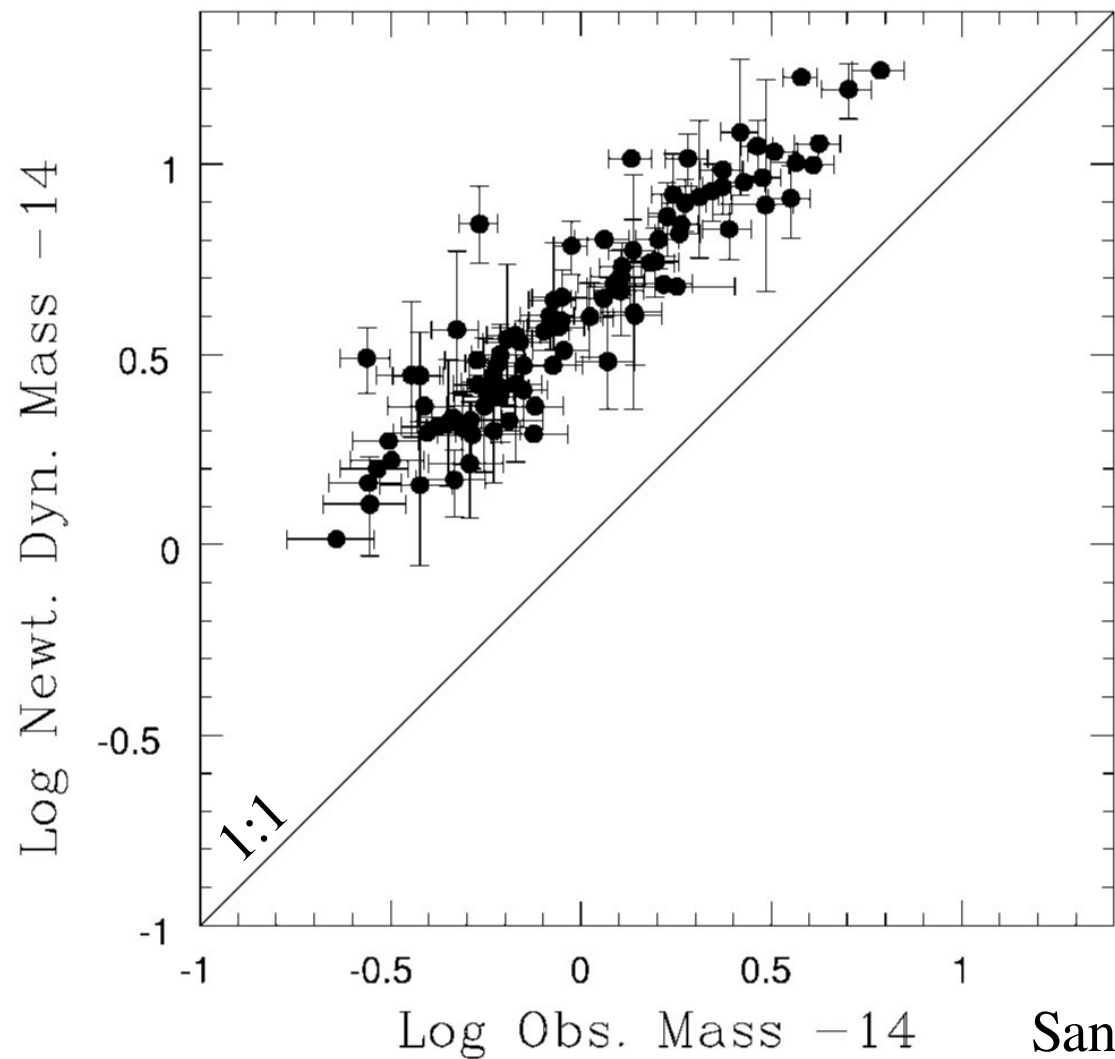
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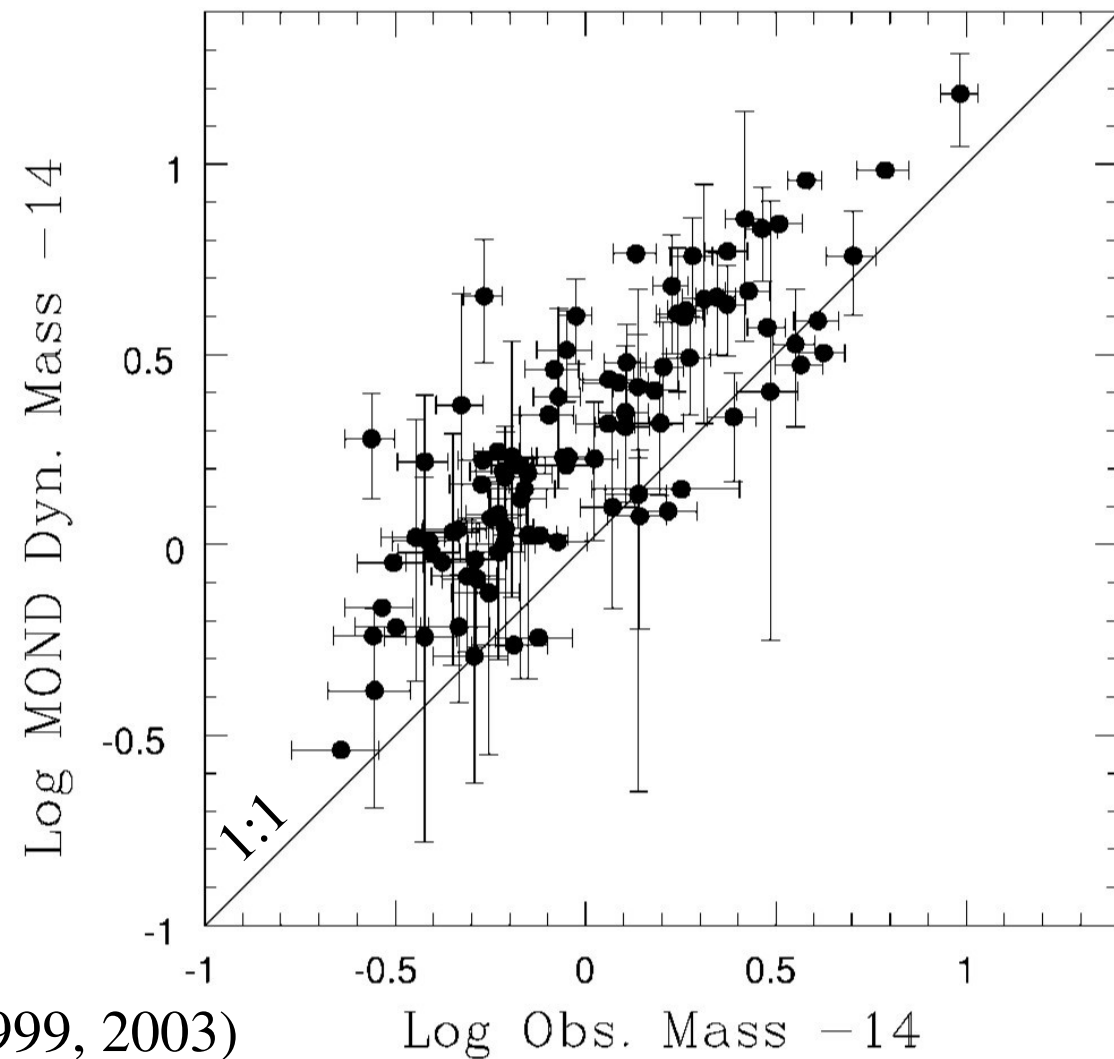
Galaxy Clusters: Long-standing problem for MOND

Newtonian analysis: $M_{\text{dyn}}/M_{\text{bar}} \simeq 5$

MOND analysis: $M_{\text{dyn}}/M_{\text{bar}} \simeq 2$



Sanders (1999, 2003)



Log Obs. Mass -14

Nature of the missing mass?

- **Dark baryons** such as compact clouds of cold gas (Milgrom 2008)

For HI clouds ($T \sim 10^4$ K) $\rightarrow M_{c1} < 10^5 M_{\odot}$ and $R_{c1} < 50$ pc

(Kelleher & Lelli 2024)

Below current HI detection limits; SKA may detect them.

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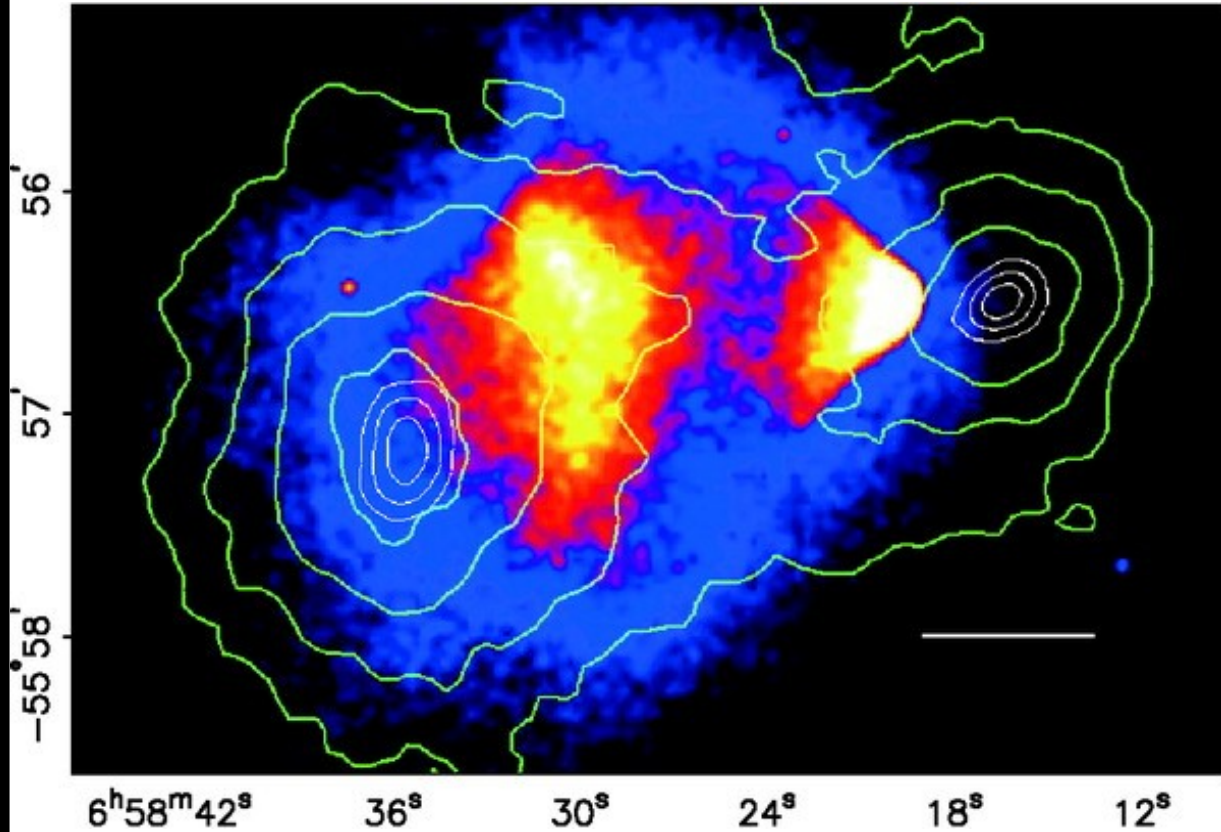
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- **Sterile neutrinos** of ~ 10 eV (Angus 2008, 2009) but CMB power spectrum may be problematic (Thomas+2016, Kopp+2018, Ilic+2021)

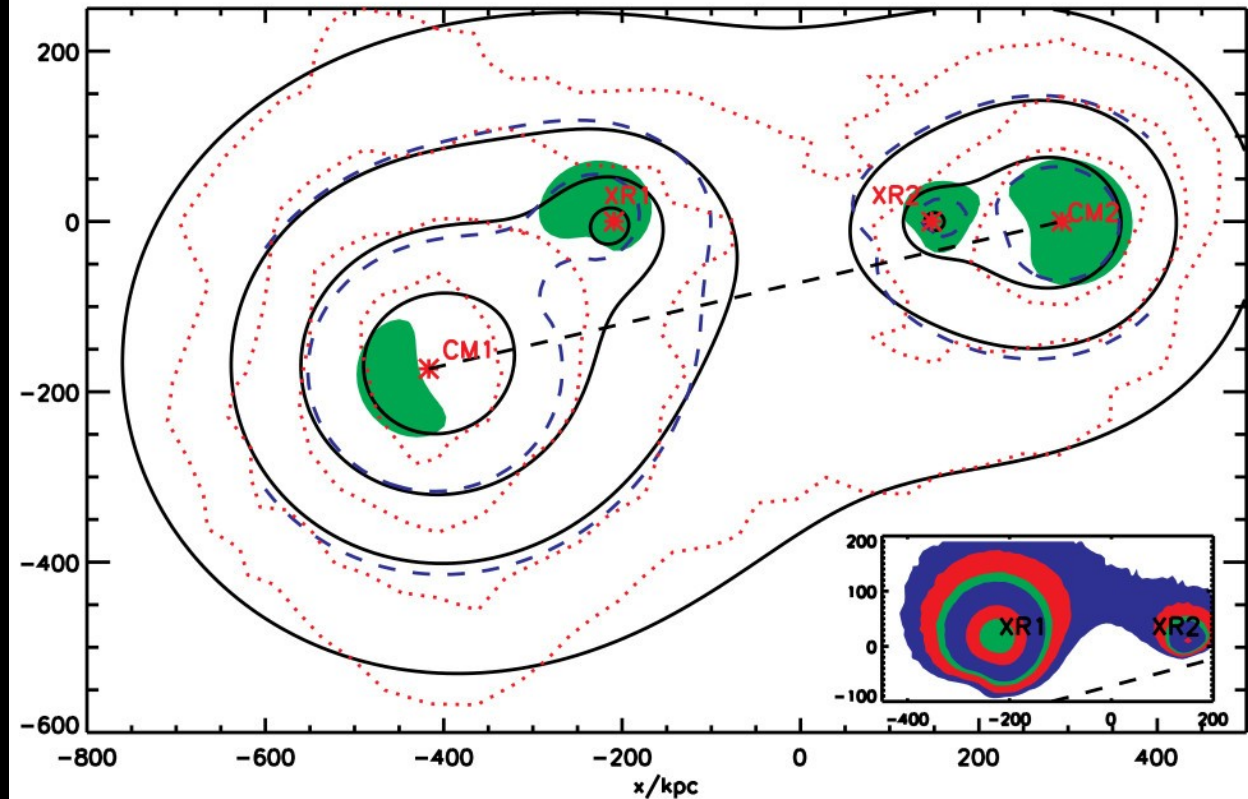
Bullet cluster in MOND: missing mass must be collisionless

OBSERVATIONS (Clowe+2004, ApJ)



Green: Observed lensing map (total mass)
Blue/Red/Yellow: X-ray emission (hot gas)

MOND (Angus+2006, MNRAS; Angus+2007, ApJ)



Red: Observed lensing convergence map
Black: TeVeS model with 2eV neutrinos
Blue: total surface densities (baryons+ ν)

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Cosmology with modified Newtonian dynamics (MOND)

R. H. Sanders

Kapteyn Astronomical Institute, Groningen, The Netherlands

Accepted 1998 January 13. Received 1997 December 22; in original form 1997 October 17

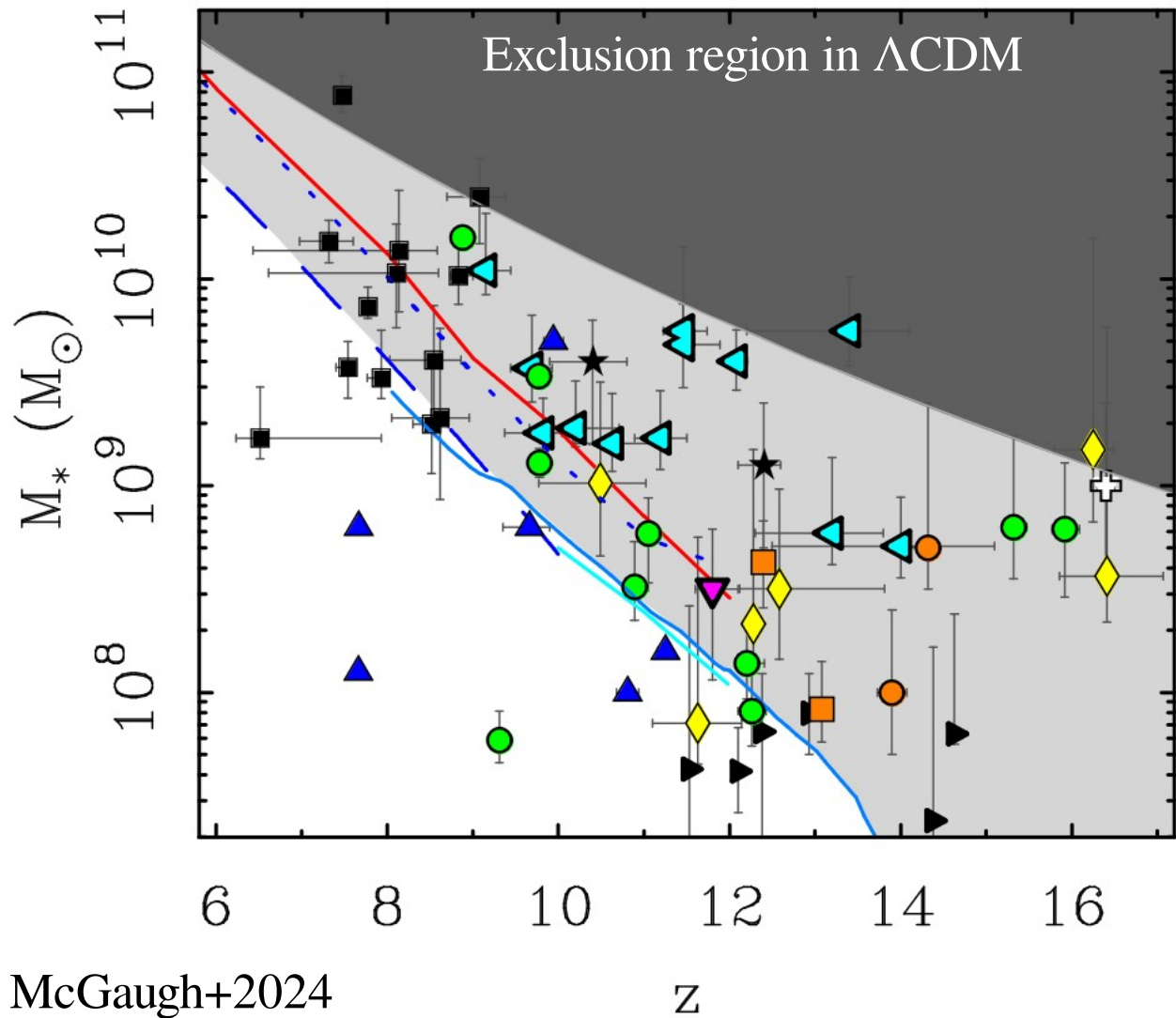
ABSTRACT

It is well known that the application of Newtonian dynamics to an expanding spherical region leads to the correct relativistic expression (the Friedmann equation) for the evolution of the cosmic scalefactor. Here, the cosmological implications of Milgrom's modified Newtonian dynamics (MOND) are considered by means of a similar procedure. Earlier work by Felten demonstrated that in a region dominated by modified dynamics the expansion cannot be uniform (separations cannot be expressed in terms of a scalefactor) and that any such region will eventually recollapse regardless of the initial expansion velocity and mean density. Here I show that, because of the acceleration threshold for the MOND phenomenology, a region dominated by MOND will have a finite size which, in the earlier Universe ($z > 3$), is smaller than the horizon scale. Therefore, uniform expansion and homogeneity on the horizon scale are consistent with MOND-dominated non-uniform expansion and the development of inhomogeneities on smaller scales. In the radiation-dominated era, the amplitude of MOND-induced inhomogeneities is much smaller than that implied by observations of the cosmic background radiation, and the thermal and dynamical history of the Universe is identical to that of the standard big bang model. In particular, the standard results for primordial nucleosynthesis are retained. When matter first dominates the energy density of the Universe, the cosmology diverges from that of the standard model. Objects of galaxy mass are the first virialized objects to form (by $z = 10$), and larger structure develops rapidly. At present, the Universe would be inhomogeneous out to a substantial fraction of the Hubble radius.

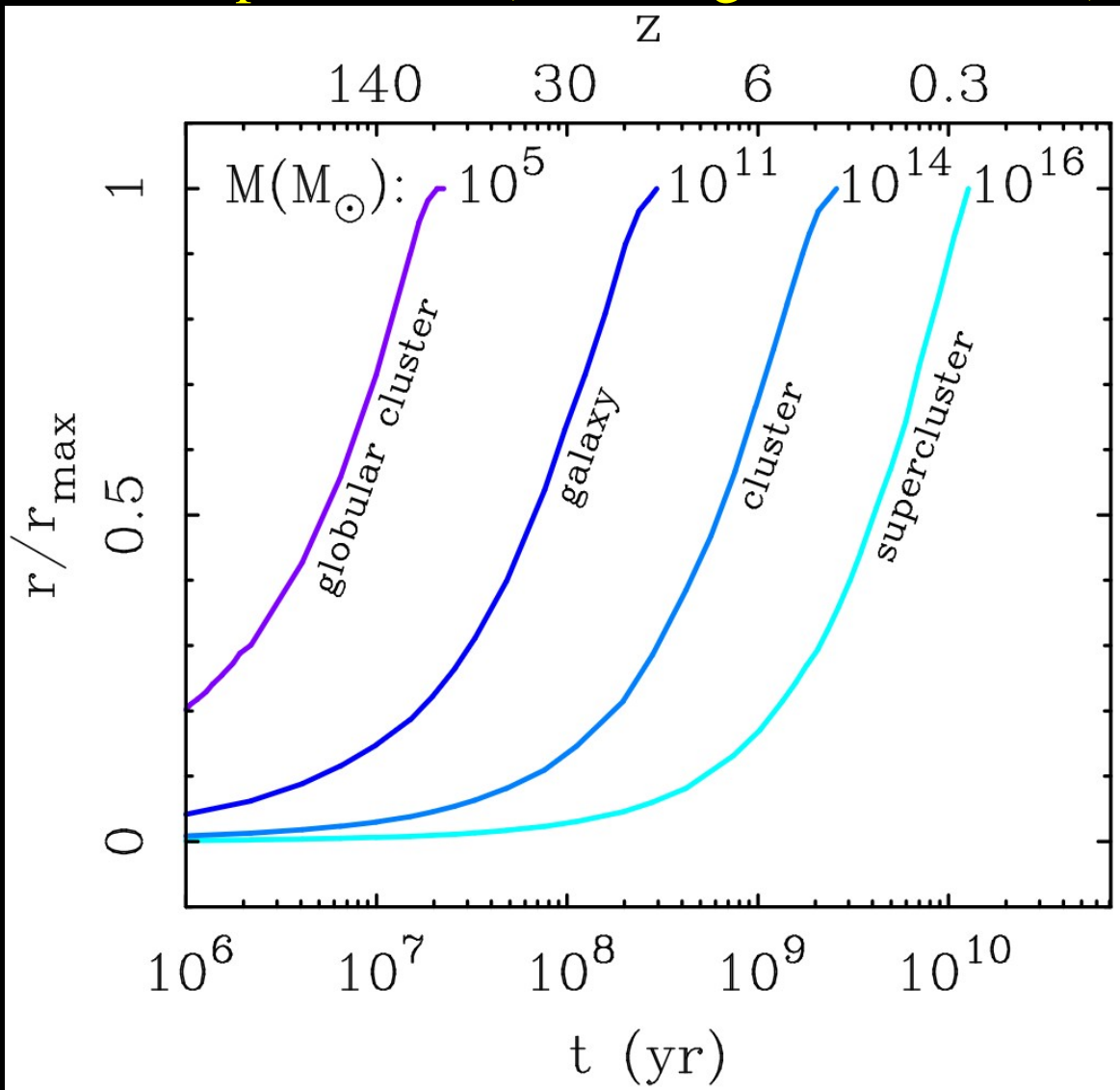


Massive galaxies at $z > 10$ is the new normal!

Compilation of JWST observations



MOND prediction (following Sanders 1998)



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```
graph TD; A[MOND paradigm] --> B[Modified Gravity (→ ∇²Φ = 4πGρ)]; A --> C[Modified Inertia (→ F = ma)];
```

Modified Gravity ($\rightarrow \nabla^2\Phi = 4\pi G\rho$)

Modified Inertia ($\rightarrow F = ma$)

MOND paradigm

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Stratified Scalar-Tensor theory (Sanders 1997, 2011)

TeVes: Tensor-Vector-Scalar (Bekenstein 2004)

BIMOND: bimetric theory (Milgrom 2009, 2022)

Non-local metric theories (Deffayet+2011, +2014)

AeST: Aether-Scalar-Tensor (Skordis & Zlosnik 2021)

Khronon: Scalar-Tensor (Blanchet & Skordis 2024)

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Time non-local theories

(Milgrom 1994, 2022, 2023)

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Heuristic ideas:

Mach's principle for inertia ?

$a_0 \simeq c \Lambda^{1/2} \rightarrow$ quantum vacuum? (Milgrom 1999)

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MOND-like DM models:

Bipolar DM (Blanchet+2008, 2009, 2015, 2017)

Superfluid DM (Berezhiani & Khoury 2015)

Baryon-interacting DM (Famaey+2018, +2020)

MOND paradigm

Modified Gravity ($\rightarrow \nabla^2 \Phi = 4\pi G \rho$)

Non-relativistic Lagrangian theories:

AQUAL (Bekenstein & Milgrom 1984)

QUMOND (Milgrom 2010, 2023)

TRIMOND (Milgrom 2022)

Relativistic Lagrangian theories:

Stratified Scalar-Tensor theory (Sanders 1997, 2011)

~~TeV~~S: Tensor-Vector-Scalar (Bekenstein 2004)

BIMOND: bimetric theory (Milgrom 2009, 2022)

Non-local metric theories (Deffayet+2011, +2014)

AeST: Aether-Scalar-Tensor (Skordis & Zlosnik 2021)

Khronon: Scalar-Tensor (Blanchet & Skordis 2024)

Modified Inertia ($\rightarrow F = ma$)

Time non-local theories

(Milgrom 1994, 2022, 2023)

Heuristic ideas:

Mach's principle for inertia ?

$a_0 \simeq c \Lambda^{1/2} \rightarrow$ quantum vacuum? (Milgrom 1999)

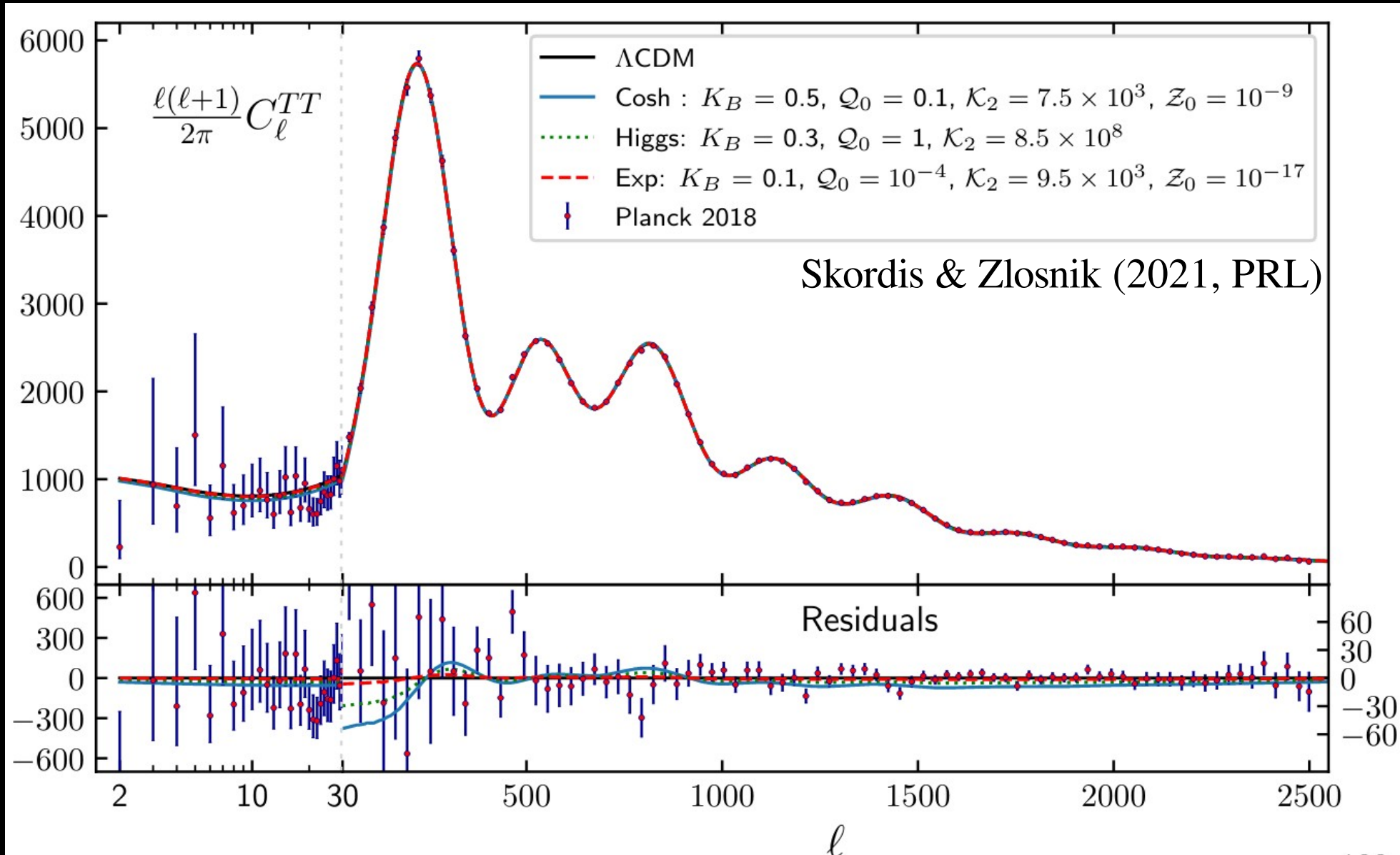
MOND-like DM models:

Bipolar DM (Blanchet+2008, 2009, 2015, 2017)

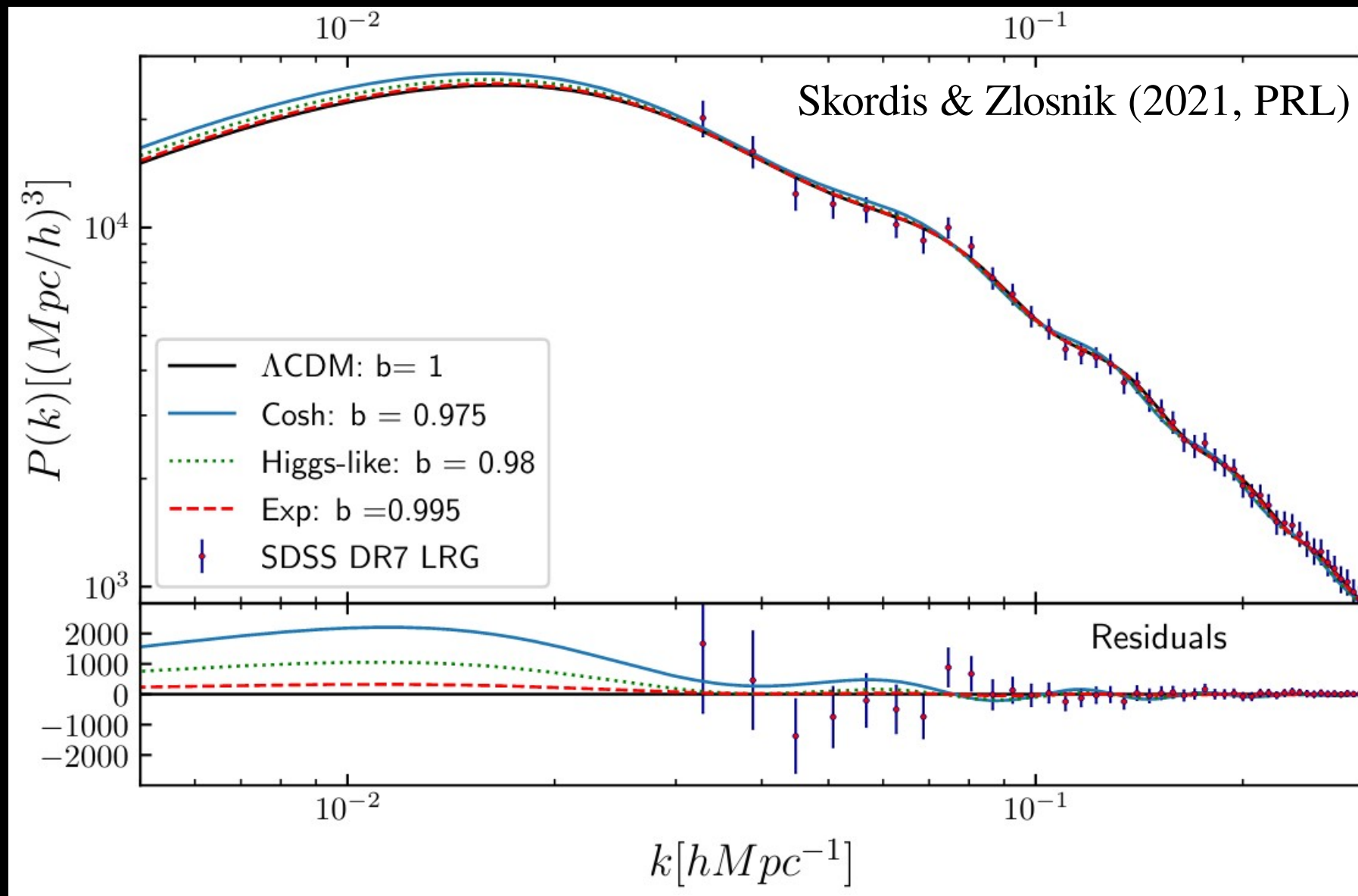
Superfluid DM (Berezhiani & Khoury 2015)

Baryon-interacting DM (Famaey+2018, +2020)

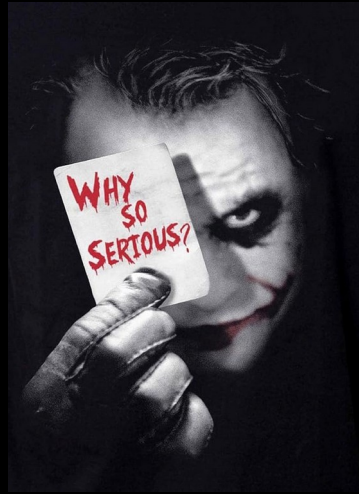
AeST theory: CMB Power Spectrum



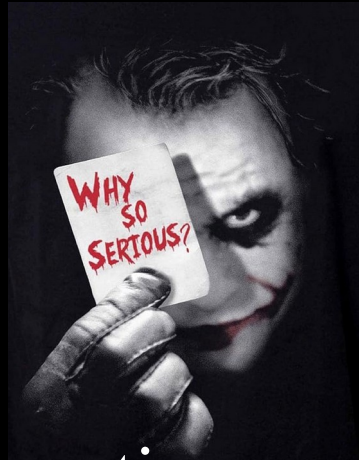
AeST theory: Linear Matter Power Spectrum



Why taking MOND seriously?

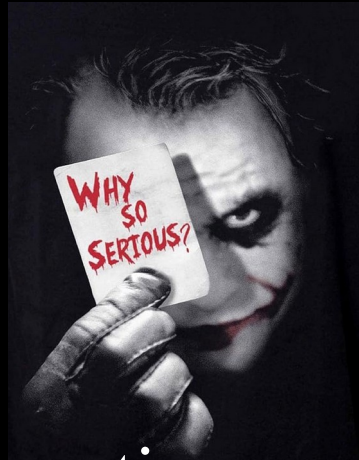


Why taking MOND seriously?



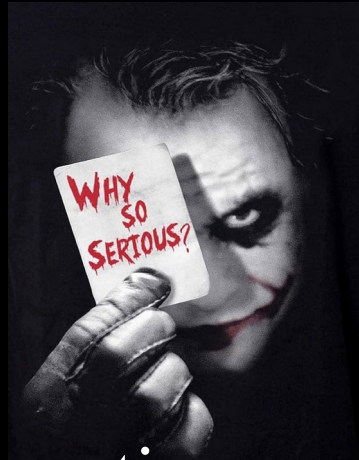
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- Even if MOND is wrong (or DM is detected), there must be a **baryon – DM coupling in galaxies acting in a MOND way**. It tells us on the nature of DM.
- Over the past ~10 years, **substantial progress in MOND research** in both observations and theory (new tests, new ideas, new relativistic theories).

Summary: Status of MOND at Various Scales

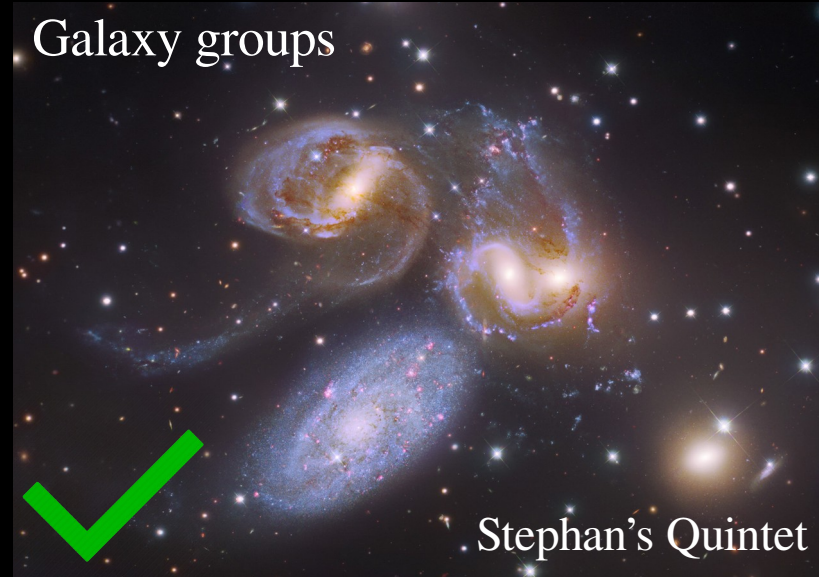
Galaxy Scales (~1-100 kpc)

Rotation-supported galaxies



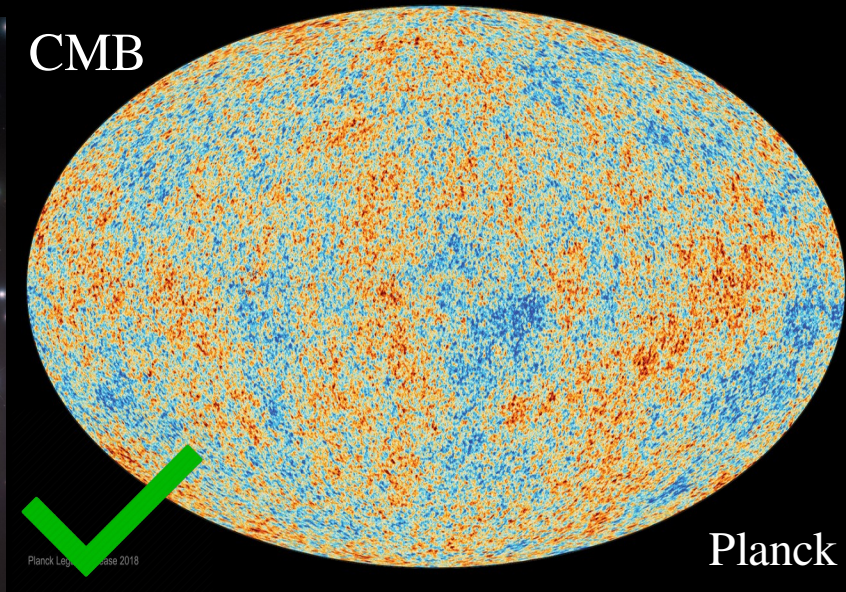
Groups/Clusters Scales (~1-5 Mpc)

Galaxy groups

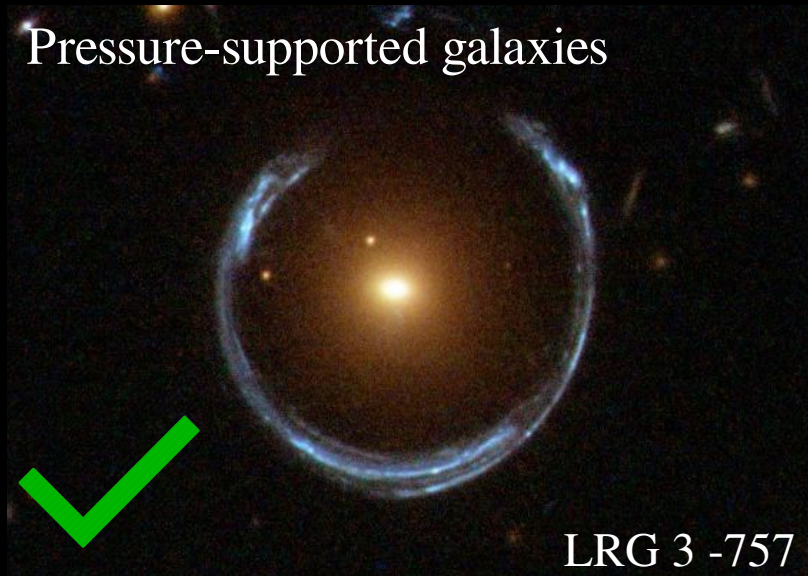


Cosmological Scales (>100 Mpc)

CMB



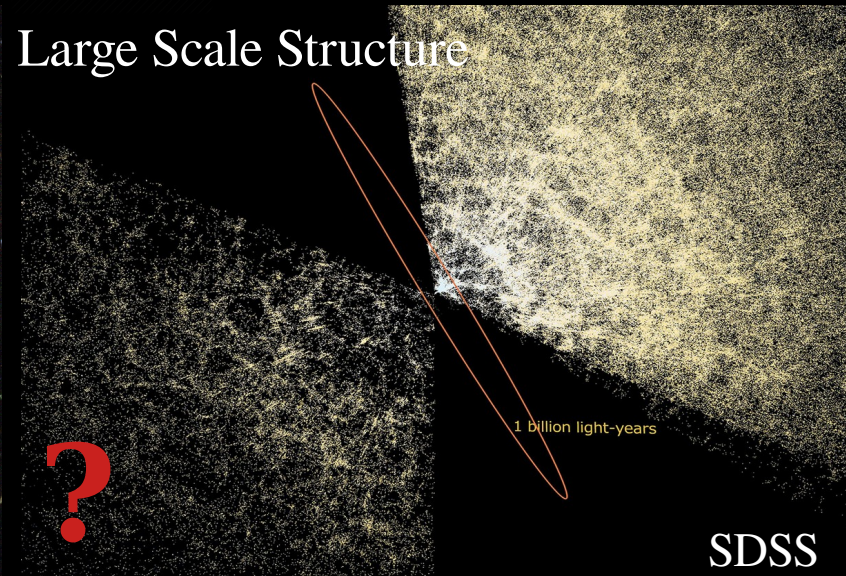
Pressure-supported galaxies



Galaxy clusters

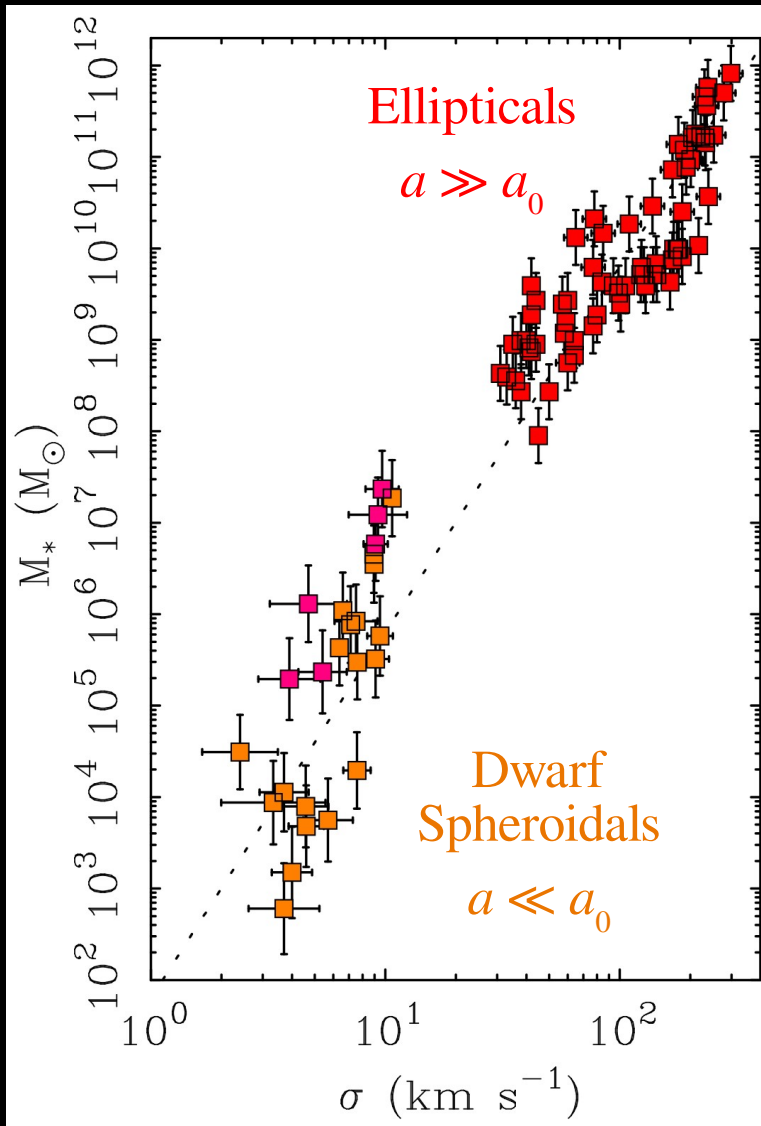


Large Scale Structure



More Slides

(4) $\sigma_V^4 \simeq a_0 G M_b$ for quasi-isothermal systems \rightarrow pressure-supported gals



Faber-Jackson relation (1976, ApJ) for ellipticals

Three a-priori independent predictions in one equation:

(i) Slope should be exactly 4 \rightarrow **OK**

(ii) Normalization is $a_0 G \rightarrow$ **OK** with BTFR estimate!

(iii) No dependence on other quantities **IF** $a \ll a_0 \rightarrow$ **OK**

σ_V is measured at $R < R_e$ (containing half luminosity):

For dwarf spheroidals: $a \ll a_0$ at $R < R_e \rightarrow$ MOND regime

For giant ellipticals: $a \gg a_0$ at $R < R_e \rightarrow$ Newtonian regime

$$\frac{\sigma_V^2}{R} \simeq \frac{G M}{R^2} \quad \rightarrow \quad M \simeq \sigma_V^2 R_e \quad \text{Fundamental plane of ellipticals}$$

(Djorgovski & Davis 1987; Dressler 1987)

Interacting & Merging Galaxies: The Antennae

Observations

Blue = Gas

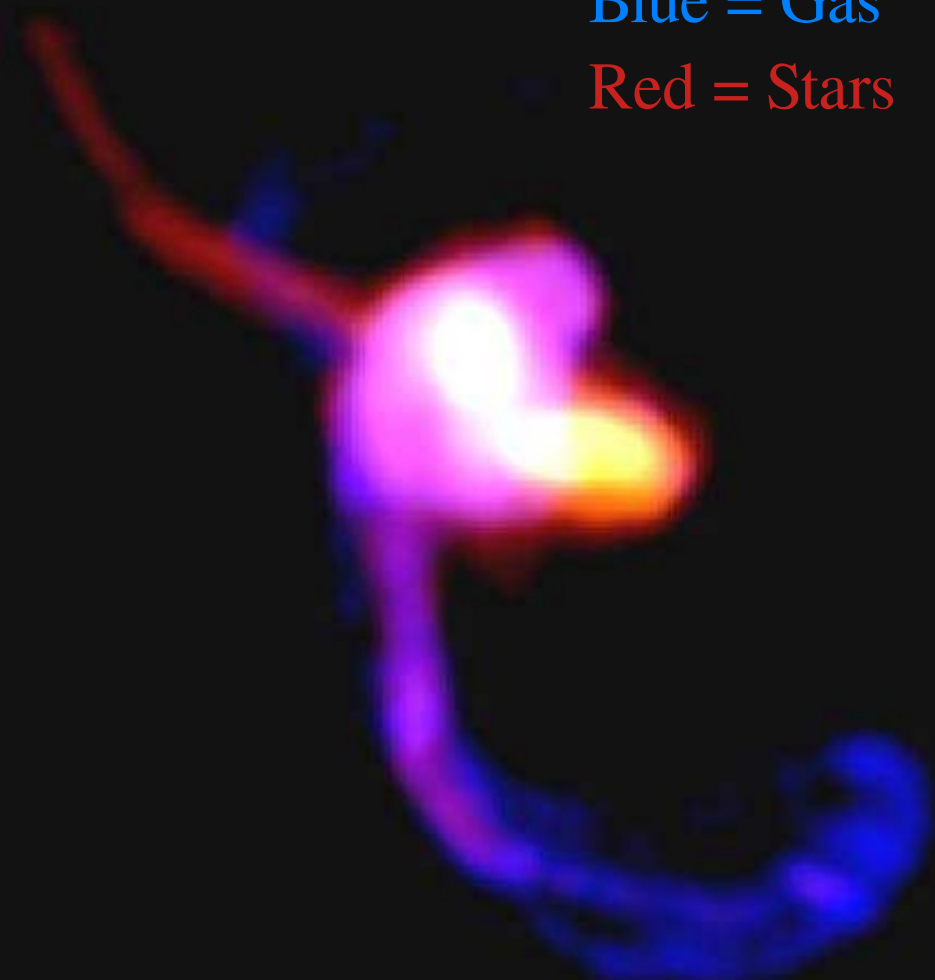
Green = Stars



Simulations

Blue = Gas

Red = Stars



Tiret & Combes (2008, ASPC)

Galaxy Clusters: bound systems with ~100-1000 galaxies

Observed baryon budget:

~10% galaxies (optical & NIR)

~90% hot ionized gas (X rays)

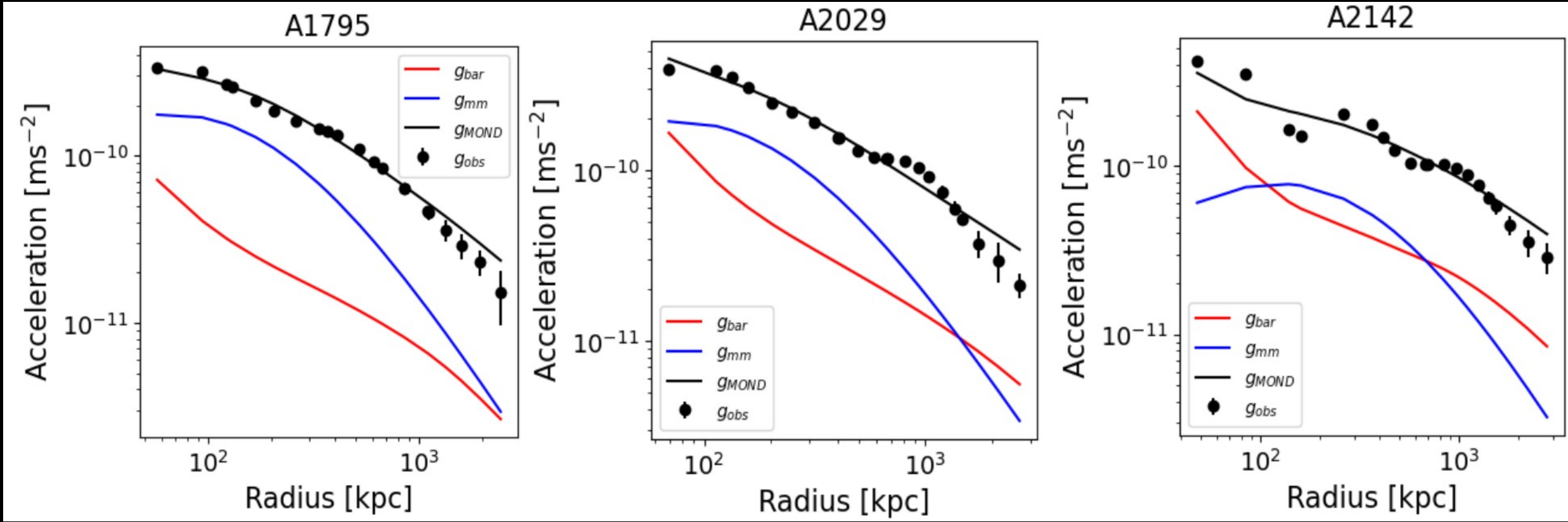
Sphere in Hydrostatic Equilibrium

$$g_{obs} = \frac{-1}{\rho_{gas}} \frac{\partial P_{gas}}{\partial r} \quad P_{gas} = \frac{k_B \rho_{gas} T_{gas}}{w m_p}$$

$$\Rightarrow g_{obs} = \frac{-k_B}{w m_p \rho_{gas}} \frac{\partial}{\partial r} (\rho_{gas} T_{gas})$$

XLSSC 006

MOND fits with a missing mass component (Kelleher & Lelli 2024)



Density profile of missing mass:

$$\rho_{mm}(r) = \frac{\rho_0}{(1+r/r_s)^4}$$

Converged (physical) total mass:

$$\Rightarrow M_{mm,tot} = \int_0^{\infty} \rho(r) 4\pi r^2 dr = \frac{4\pi}{3} \rho_0 r_s^3$$

MOND – Cosmology Connection?

Two numerical coincidences (Milgrom 1983a, ApJ; Milgrom 1999, PhLA):

$$a_0 \simeq \frac{H_0 \cdot c}{2\pi} \quad H_0 = \text{Hubble constant} \rightarrow \text{maybe } a_0(t) \sim H(t) ?$$

$$a_0 \simeq \frac{c^2 \sqrt{\Lambda/3}}{2\pi} \quad \Lambda = \text{Cosmological constant} \rightarrow \text{relation to Dark Energy?}$$

IF this numerology has some deeper, fundamental meaning:

Either the state of the Universe at large enters in local dynamics, or the same parameters enters both Cosmology (Λ) and local dynamics (a_0).

AQUAL = Aquadratic Lagrangian (Bekenstein & Milgrom 1984, ApJ)

$$S_N = \int dt L_N = \int dt d^3 x \left(\rho \frac{V^2}{2} - \frac{|\vec{\nabla} \Phi|^2}{8\pi G} - \rho \Phi \right) \quad \text{Lagrangian is quadratic in } \nabla \Phi$$

$$\downarrow$$

$$-\frac{a_0^2}{8\pi G} F\left(\frac{|\vec{\nabla} \Phi|^2}{a_0^2}\right) \quad \begin{aligned} F(z) &\rightarrow z \text{ for } z = |\nabla \Phi|^2/a_0^2 \gg 1 \\ F(z) &\rightarrow z^{3/2} \text{ for } z = |\nabla \Phi|^2/a_0^2 \ll 1 \end{aligned}$$

$$\frac{\delta S}{\delta \Phi} = 0 \rightarrow \nabla \cdot \left[\mu \left(\frac{|\vec{\nabla} \Phi|^2}{a_0} \right) \vec{\nabla} \Phi \right] = 4\pi G \rho \quad \text{Modified Poisson's Equation}$$

$$\mu(x) = \frac{dF(z)}{dz} \quad z = x^2 \quad F(z) \text{ provides the interpolation function } \mu = v^{-1}$$

QUMOND = Quasi-Linear MOND (Milgrom 2010, MNRAS)

$$S_N = \int dt L_N = \int dt d^3 x \left(\rho \frac{V^2}{2} - \frac{|\vec{\nabla} \Phi|^2}{8\pi G} - \rho \Phi \right) \quad \text{Single gravitational potential } \Phi$$

$$\frac{-1}{8\pi G} \left[2 \vec{\nabla} \Phi \cdot \vec{\nabla} \Phi_N - a_0^2 Q \left(\frac{|\vec{\nabla} \Phi_N|^2}{a_0^2} \right) \right] \quad \text{Two potentials: } \Phi \text{ and } \Phi_N!$$

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Principle of least Action varying Φ , Φ_N and $\vec{x} \rightarrow$ set of 3 equations

$$\nabla^2 \Phi_N = 4\pi G \rho \longrightarrow \text{Standard, linear Poisson's equation for } \Phi_N$$

$$\nabla^2 \Phi = \vec{\nabla} \cdot \left[v \left(\frac{|\vec{\nabla} \Phi_N|}{a_0} \right) \vec{\nabla} \Phi_N \right] \longrightarrow \text{Non-linear step: get } \Phi \text{ from } \Phi_N \quad v(\sqrt{x}) = \frac{dQ(x)}{dx}$$

$$\vec{a} = -\vec{\nabla} \Phi \longrightarrow \text{Acceleration/force set by second potential } \Phi$$

Relativistic MOND \rightarrow add degrees of freedom (new fields) to GR

- Tensor $g_{\mu\nu}$ \rightarrow Einstein's metric
- Scalar ϕ \rightarrow for the DM effect in the NR limit (Bekenstein & Milgrom 1984, ApJ)
- Vector A^μ \rightarrow for gravitational lensing (Sanders 1997, ApJ; Bekenstein 2004, PRD)
- Free function(s) \rightarrow interpolation function(s) (non-fundamental effective theories?)

AeST: Aether Scalar-Tensor theory (Skordis & Zlosnik 2019, PRD; 2021, PRL; 2022, PRD)

(1) Grav. Waves: $B^\mu = e^{-2\phi} A^\mu$ (timelike) \rightarrow theories $\{B_\mu, g_{\mu\nu}\} \rightarrow c_{\text{GW}} = c_{\text{EM}}$

(2) Cosmology: k -essence term for $\phi \rightarrow \rho_\phi \propto a(t)^{-3}$ (like dust or CDM)

MOND as Modified Inertia (Milgrom 1994, 1999, 2022)

$$\vec{A}[\vec{x}(t); a_0] = -\vec{\nabla} \Phi_N \quad \bar{A} \text{ is a functional of the full trajectory } \bar{x}(t)$$

For $a \gg a_0$, $A \rightarrow a = d^2x/dt^2$ (Newton's 2nd law)

No full theory yet, but two general results:

(1) Imposing Newtonian and MOND limits + Galilei Invariance $\vec{x}(t) \rightarrow \vec{x}(t) + \vec{v}_0 t$

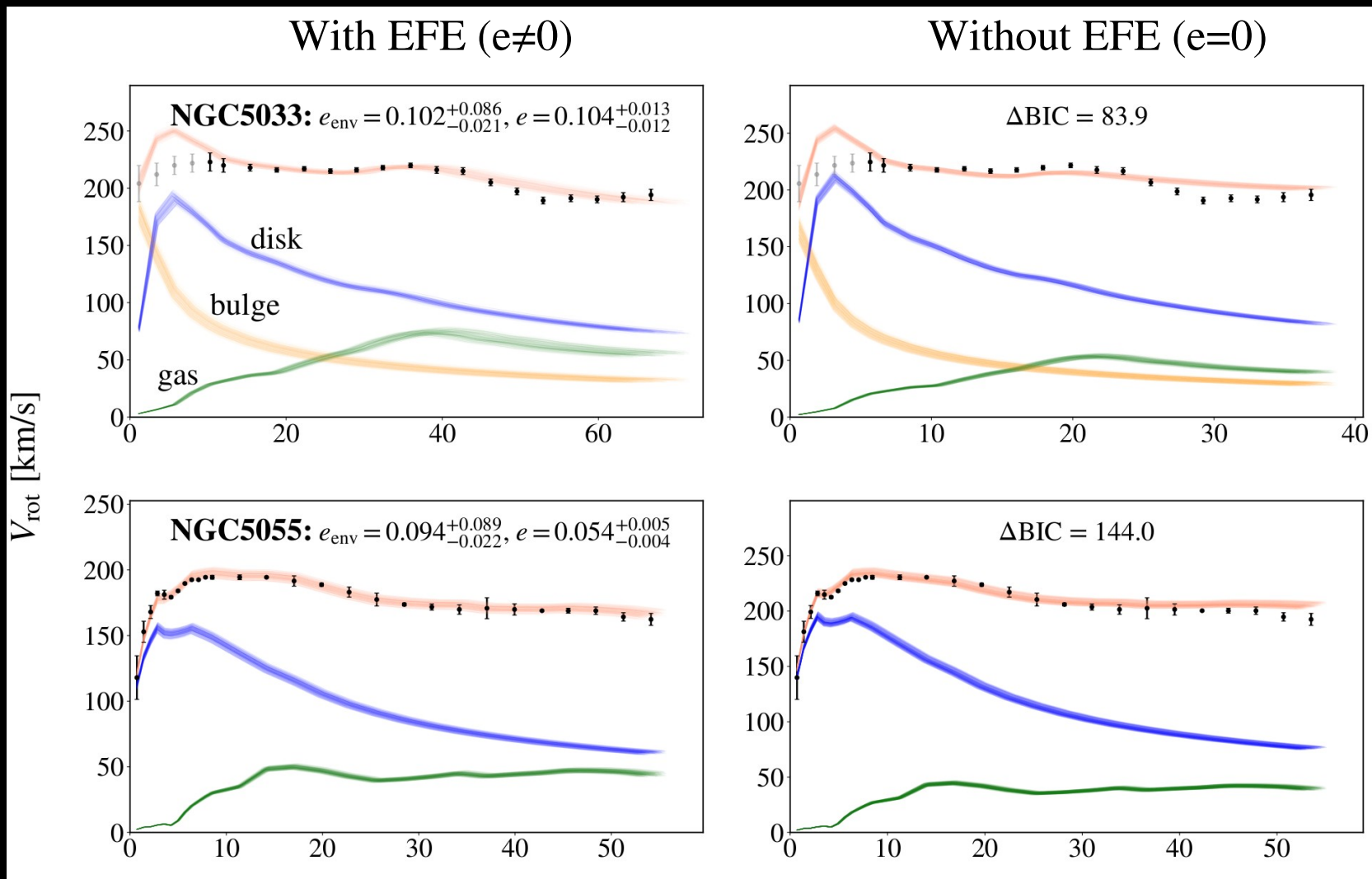
Theory is **time non-local**: $\vec{A}[\vec{x}(t), a_0] \neq F\left(\frac{d^i \vec{x}}{dt^i}; i=1, 2, \dots, N\right)$

Accelerations at (\bar{x}, t) depend on the **full orbital history**!

(2) For purely circular orbits: $\vec{a} \mu\left(\frac{a}{a_0}\right) = \vec{g}_N$ holds exactly (RAR for disk galaxies)

The **interpolation function** is a **derived concept** valid for circular orbits!

EFE is weak: individual detections only in extreme cases



NGC 5033



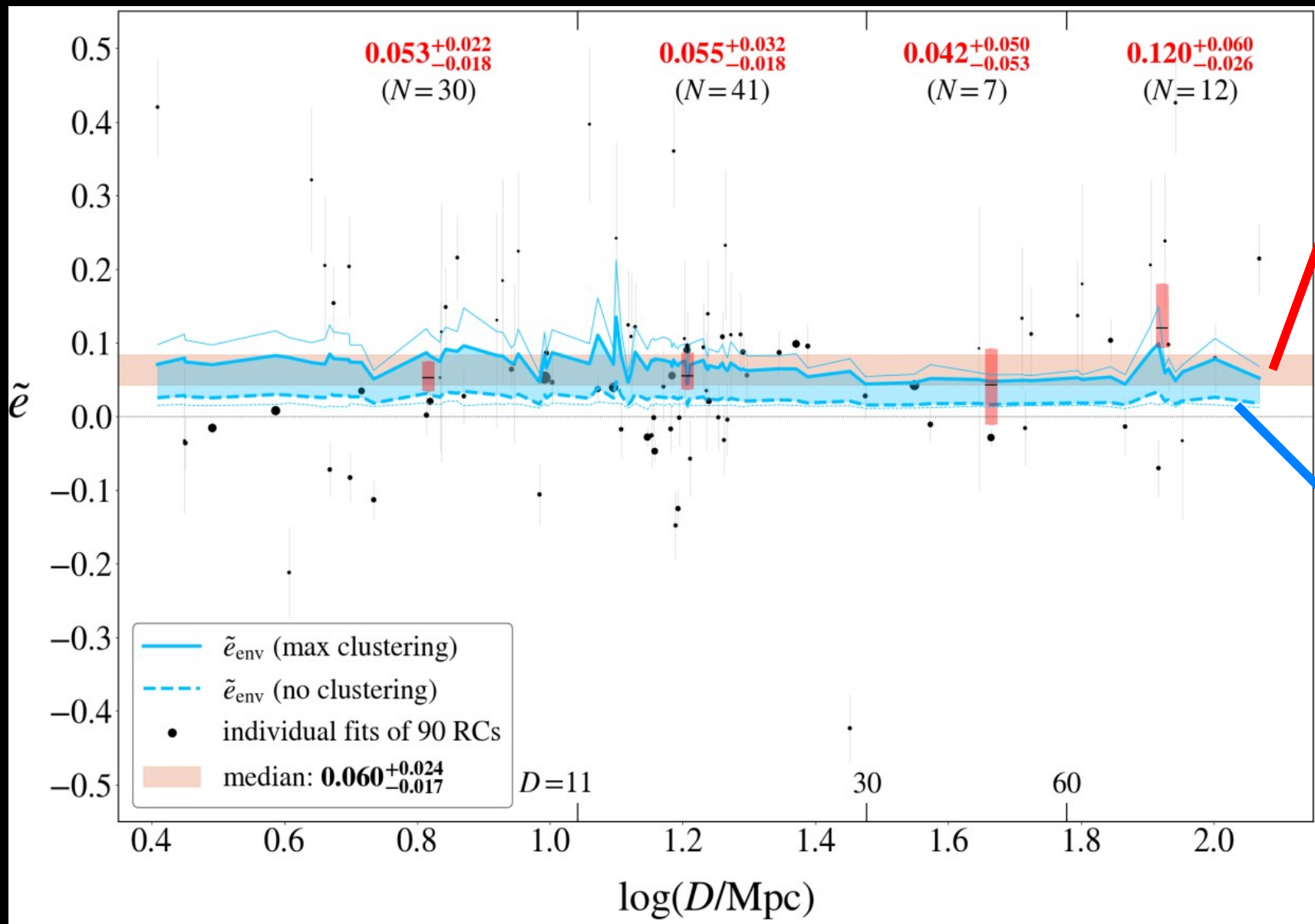
NGC 5055



Chae+2020, 2021

Statistical approach: $EFE > 0$ at $>4\sigma$ and agrees with LSS

$$\frac{g_{ext}}{a_0}$$



From Rotation
Curve Fits

From Baryon
Large Scale
Structure

Chae+2020, 2021

Galaxy Clusters on the Radial Acceleration Relation

