08 July 2024

15th International Workshop on the Identification of Dark Matter 2024

Minimal Dark Matter Freeze-in with Low Reheating Temperatures & Implications for Direct Detection

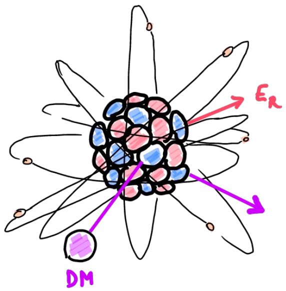
Gabriele Montefalcone

Weinberg Institute for Theoretical Physics, University of Texas at Austin

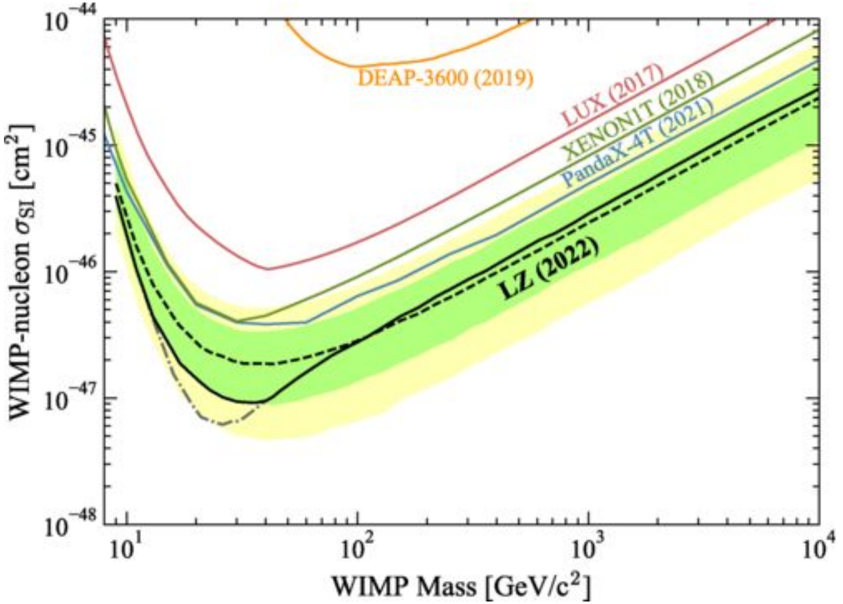
Based on work with Katherine Freese, Kimberly Boddy & Barmak Shams Es Haghi (arXiv:2405.06226)

The University of Texas at Austin

The Current State of Direct Detection
Nuclear Recoil from WIMPs

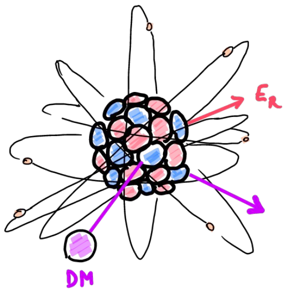


- Over the last three decades a wide-range of experimental programs have targeted the WIMP parameter space
  - increasingly constrained due to the lack of a direct detection.



LZ Collaboration 2022

The Current State of Direct Detection
Nuclear Recoil from WIMPs



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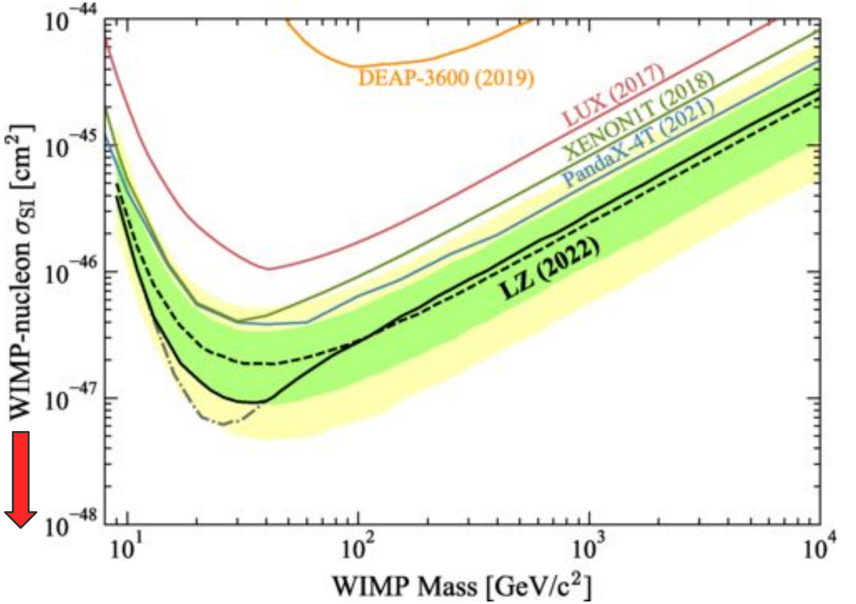
Push to **lower** cross sections

- multi-ton-scale target masses and a clear path for even larger detectors to reach the neutrino fog



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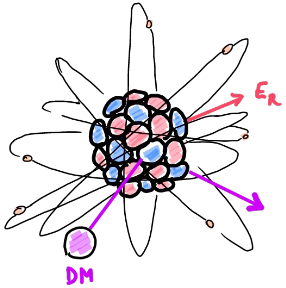

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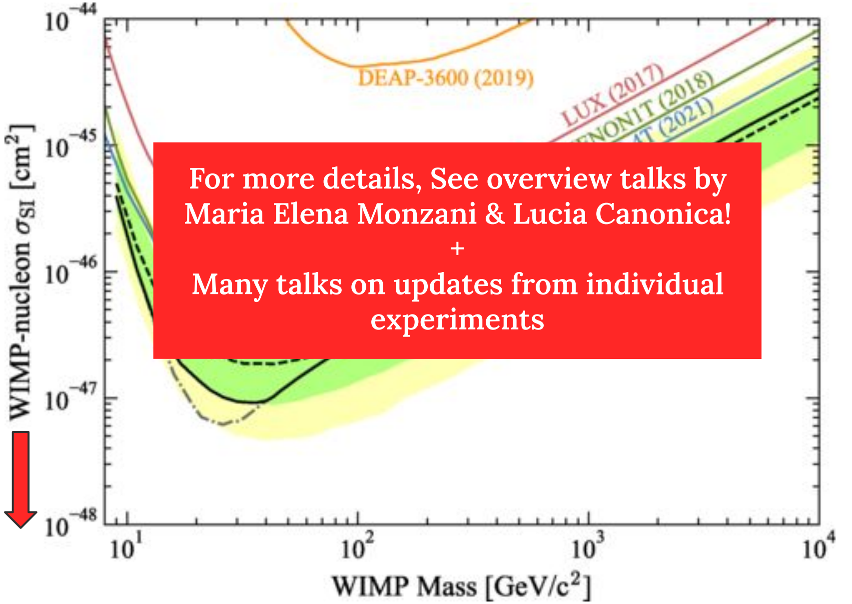
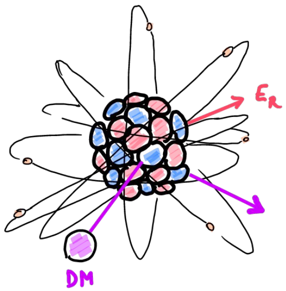


For more details, See overview talks by Maria Elena Monzani & Lucia Canonica!
+
Many talks on updates from individual experiments

LZ Collaboration 2022

Constraints on WIMPs from Direct Detection

Nuclear Recoil



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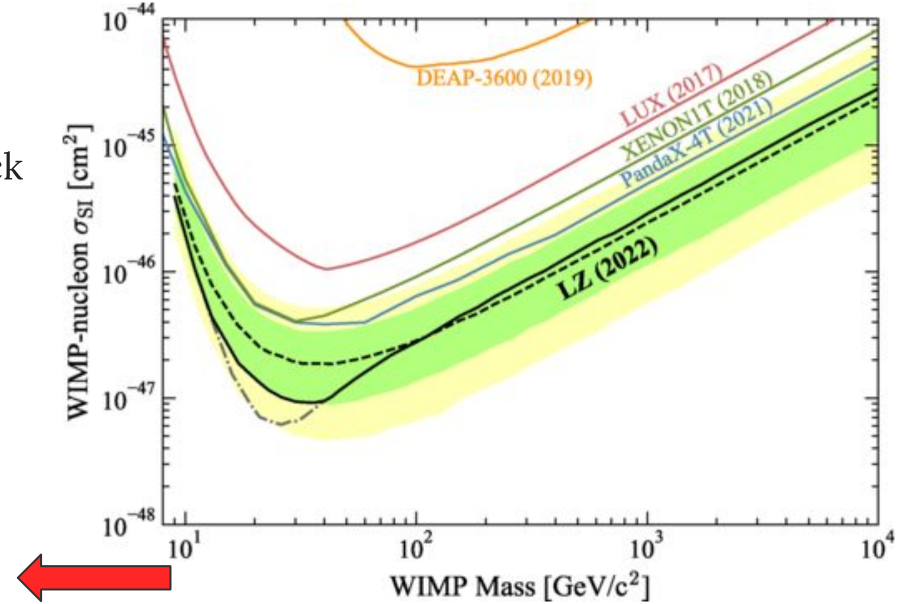
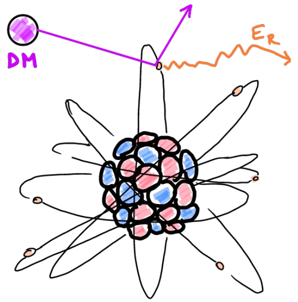
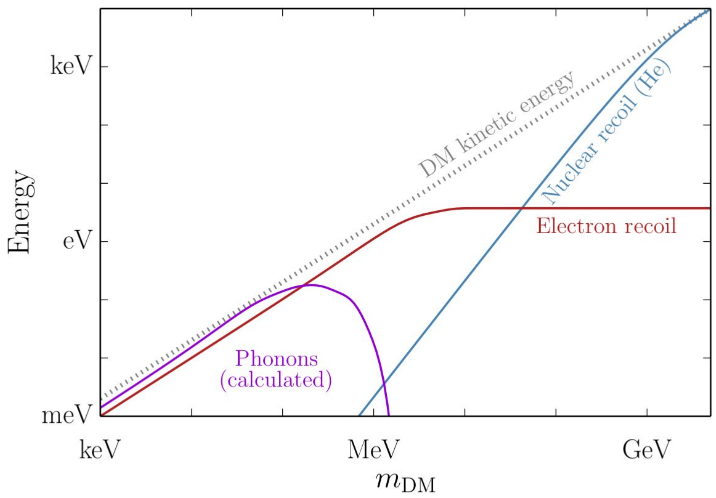
Push search to **lighter** DM candidates



Direct Detection of sub-GeV Dark Matter
Electron Recoil

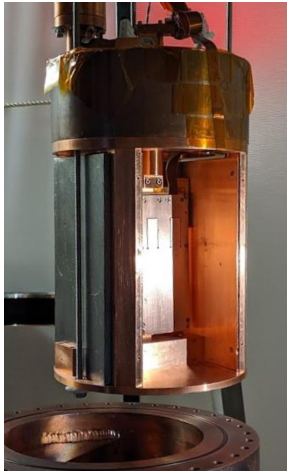


- New technology allowed recent experiments to extend their reach to sub-GeV DM masses by searching for **electronic recoils**



2019 TASI lectures by Lin

SENSEI Collaboration

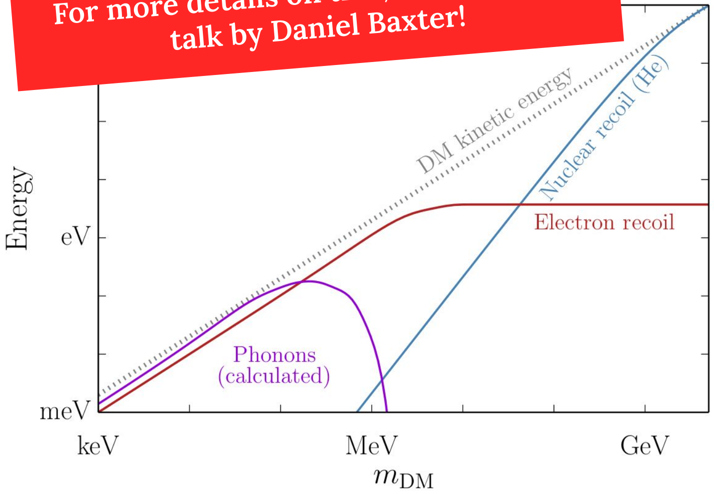


Direct Detection of sub-GeV Dark Matter

Electron Recoil

- New technology allowed recent experiments to extend their reach to sub-GeV DM masses by searching for **electronic recoils**

For more details on this, See overview talk by Daniel Baxter!



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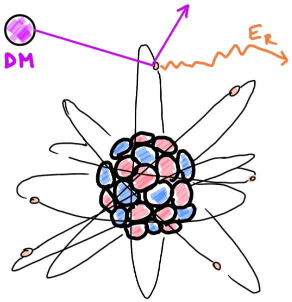

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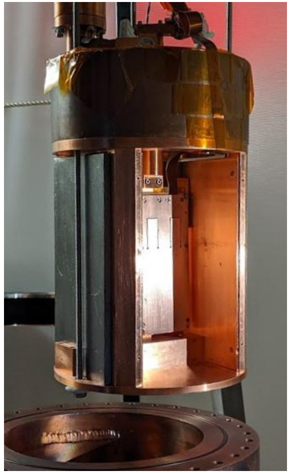


Model Building Challenges of Light Dark Matter
Necessity of a Dark Sector

- Lee-Weinberg bound: Weak scale couplings lead to an overabundance of DM for $\mathbf{m_\chi} < 1$ GeV

Lee, Weinberg 1977

→ **New BSM mediators** below the weak scale are required!

- For a sub-GeV DM candidate, if the dark sector is thermally coupled to SM, it is hard to evade CMB injection constraints.
    - Either asymmetric DM; or models with p-wave or kinematic suppression.

→ We can have a **secluded** sector (with no to negligible SM coupling)

See TASI lectures by Tongyan Lin for a review of all these constraints and the corresponding relevant papers on the subject.

Model Building Challenges of Light Dark Matter

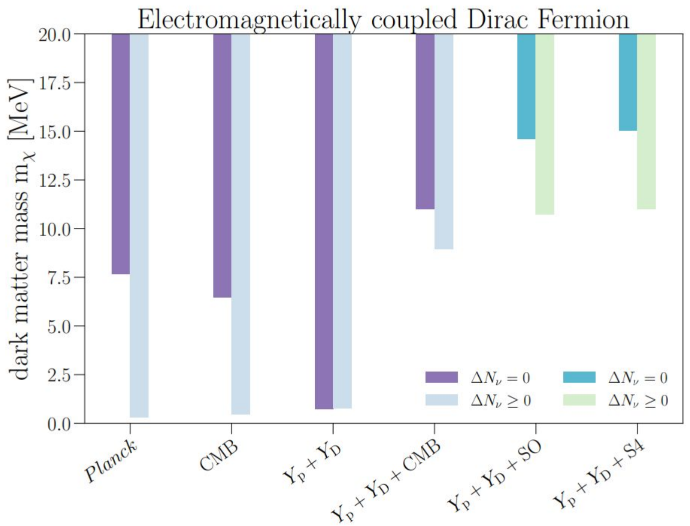
Krnjaic, McDermott, 2019;
An, Gluscevic, Calabrese, Hill, 2022

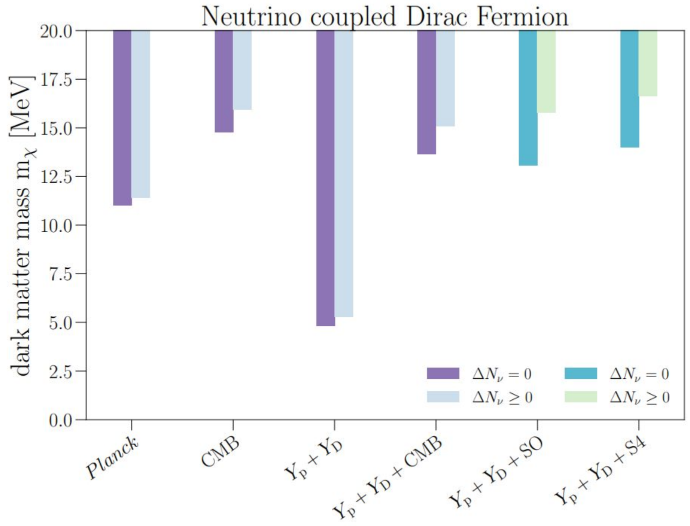
BBN constraints on thermal DM

Thermal production of **MeV DM is disallowed** by BBN

Only assumption is that DM is thermally coupled to SM

Precise constraints depend on the nature of DM particle





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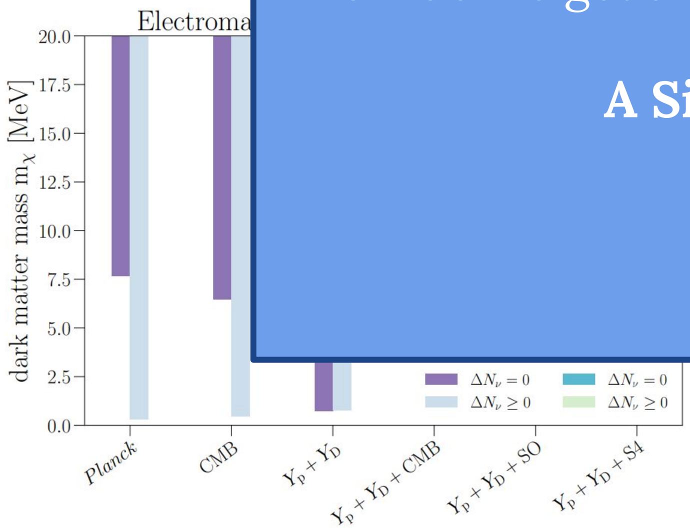
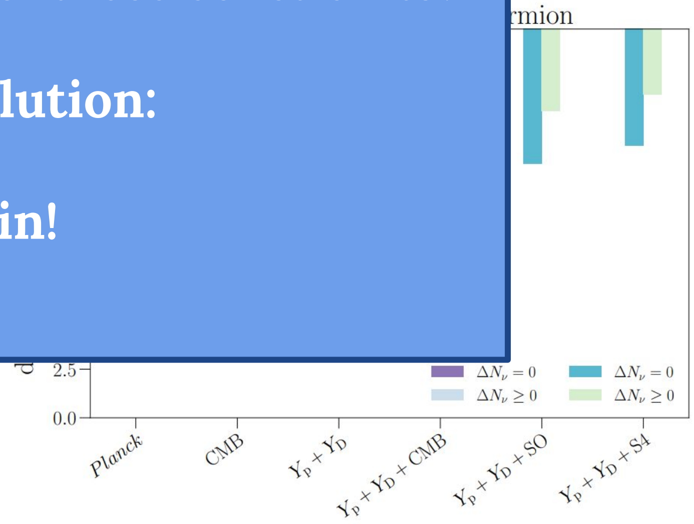


How do we get around all of these constraints?

A Simple Solution:

Freeze-in!



An alternative Scenario: Freeze-in

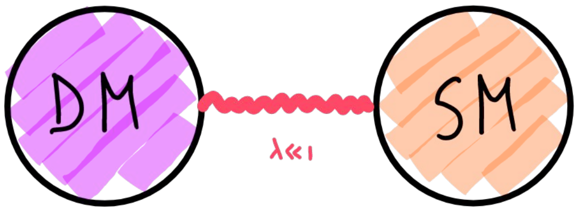
Hall, Jedamzik, March-Russell, West 2010

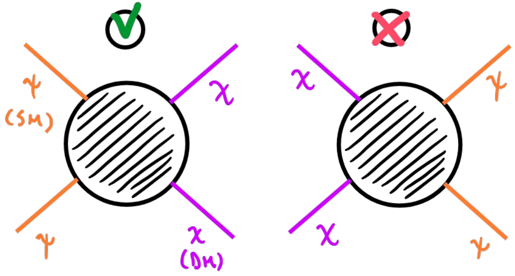
DM from a feeble interaction with SM

- **Feeble** interaction between DM and the SM so that DM is never in thermal equilibrium with the SM bath
    - Through a renormalizable operator with very small dimensionless coupling $\lambda_{\text{SM-DM}}$

- Initial DM abundance is negligible (i.e. inflaton reheats primarily the SM)

- The DM abundance is built up gradually (no inverse process!)



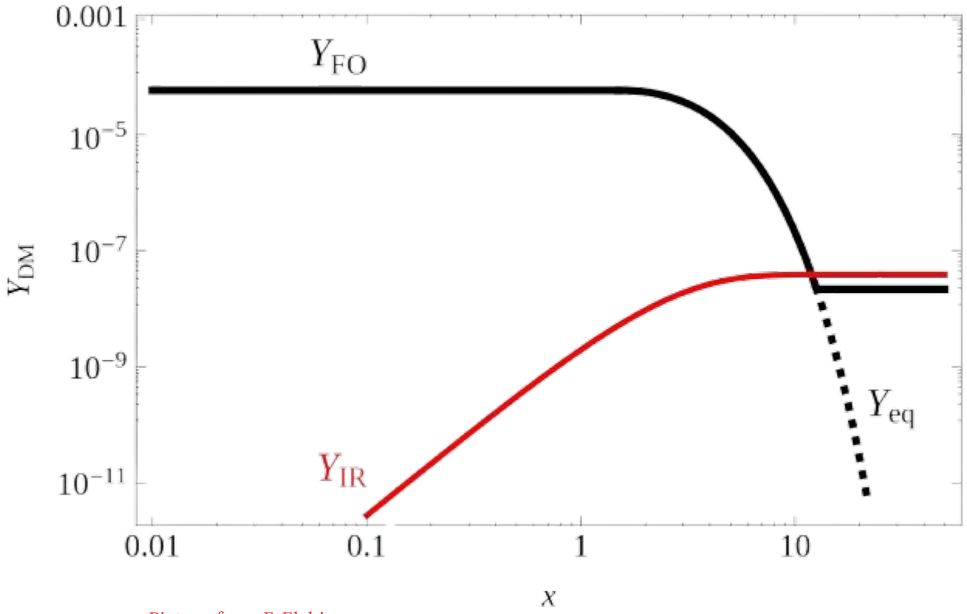


Freeze-in

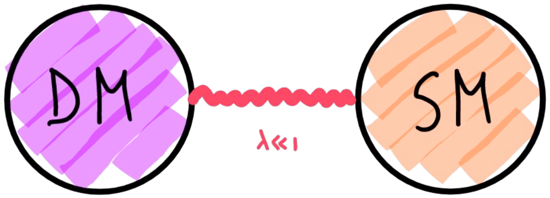
DM from a feeble interaction with SM

Hall, Jedamzik, March-Russell, West 2010

- The process is insensitive to temperatures above the DM mass
  - The DM abundance is set by lowest T, i.e. **T□m$_{\rm DM}$**



Picture from F. Elahi



$$Y_{\rm DM} \sim \lambda^2 \frac{M_{\rm pl}}{T} \sim \lambda^2 \frac{M_{\rm pl}}{m_{\rm DM}}$$

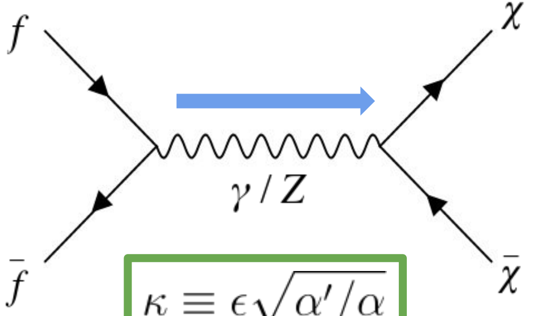
Benchmark Freeze-in Model
The Kinetic Mixing Portal

Hall, Jedamzik, March-Russell, West 2010
Chu, Hambye, Tytgat, 2012
Essig, Mardon, Volansky, 2012

- An ultralight dark photon **γ' kinetically-mixed** with the SM hypercharge

$$\mathcal{L} \supset \frac{\epsilon}{2} F'_{\mu\nu} F^{\mu\nu}$$

$$\dot{n}_\chi + 3Hn_\chi = \sum_B \langle \sigma_{B\bar{B}\to\chi\bar{\chi}} v \rangle (n_\chi^{\rm eq})^2,$$



$$\kappa \equiv \epsilon \sqrt{\alpha'/\alpha}$$

- **Target of direct detection** program!
  - Ultralight mediator leads to large enhancement of the direct detection cross section at low momentum transfers.

$$\overline{\sigma}_e = \frac{16\pi \mu_{\chi e}^2 \alpha^2 \kappa^2}{(\alpha m_e)^4},$$

Benchmark Freeze-in Model

The Kinetic Mixing Portal

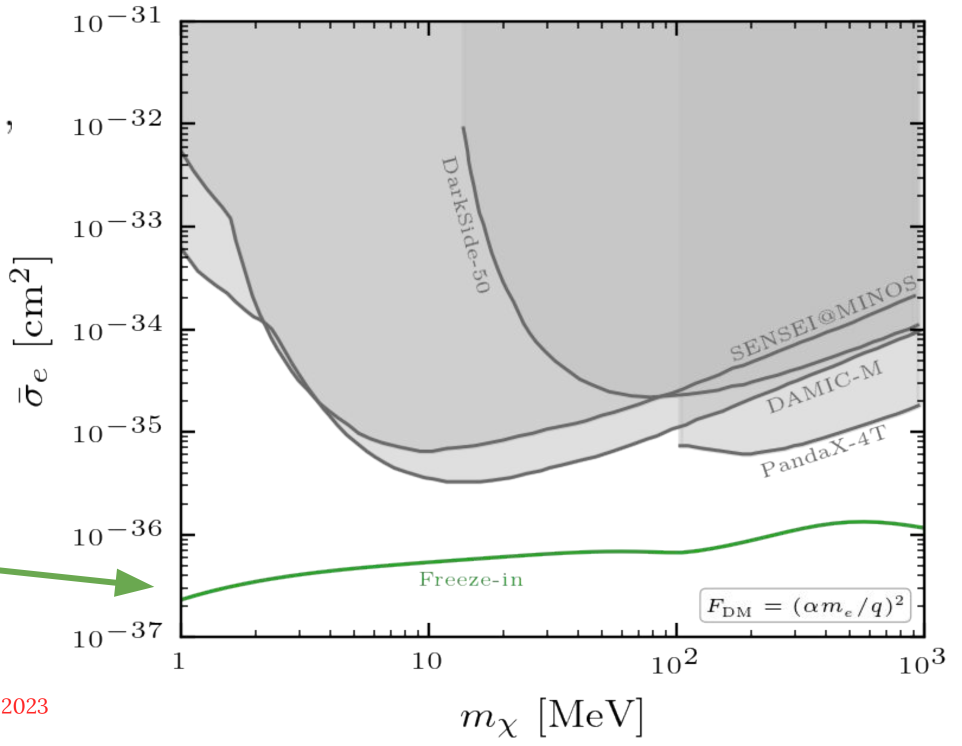
$$\bar{\sigma}_e = \frac{16\pi\mu_{\chi e}^2\alpha^2\kappa^2}{(\alpha m_e)^4},$$

$$Y_\chi(x) = \int_{x_{\rm rh}}^{x} dx' \frac{s}{Hx'} \left[\sum_B \langle\sigma_{B\bar{B}\to\chi\bar{\chi}} v\rangle (Y_\chi^{\rm eq})^2\right],$$

$$x \equiv m_\chi/T$$

- Previous work assumes $T_{rh} >> m_\chi$: $x_{rh} = 0$.
- Then, matching to the observed relic abundance today leads to

$$\kappa \equiv \epsilon\sqrt{\alpha'/\alpha} \approx \mathcal{O}(10^{-11})$$



Corrected prediction for the freeze-in benchmark by Bhattiprolu, McGehee, Pierce 2023

Benchmark Freeze-in Model

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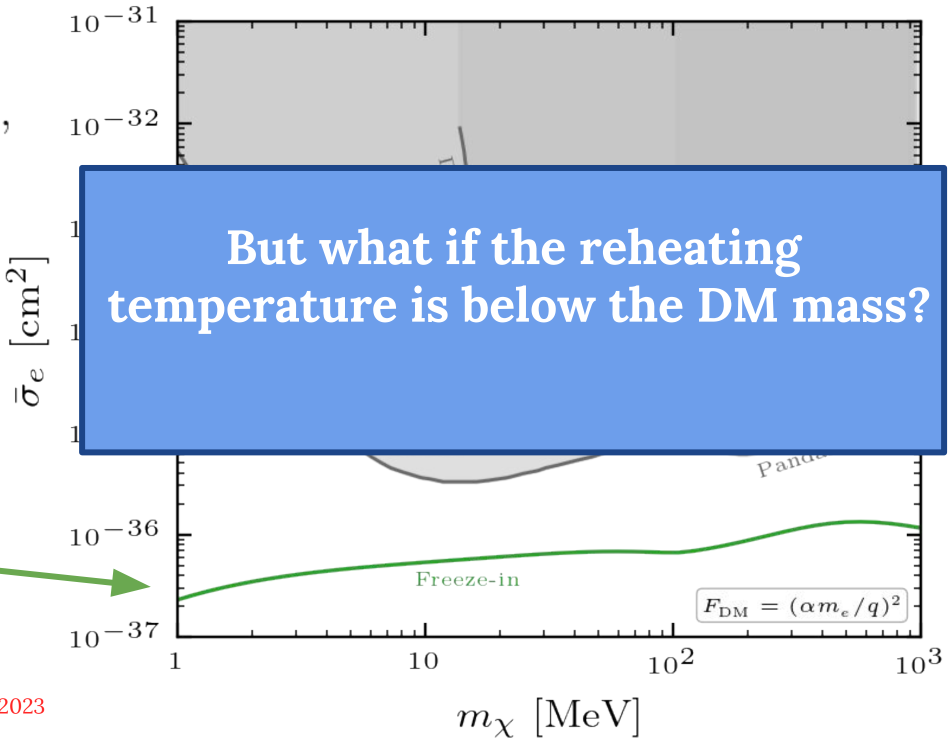


But what if the reheating temperature is below the DM mass?

Corrected prediction for the freeze-in benchmark by Bhattiprolu, McGehee, Pierce 2023

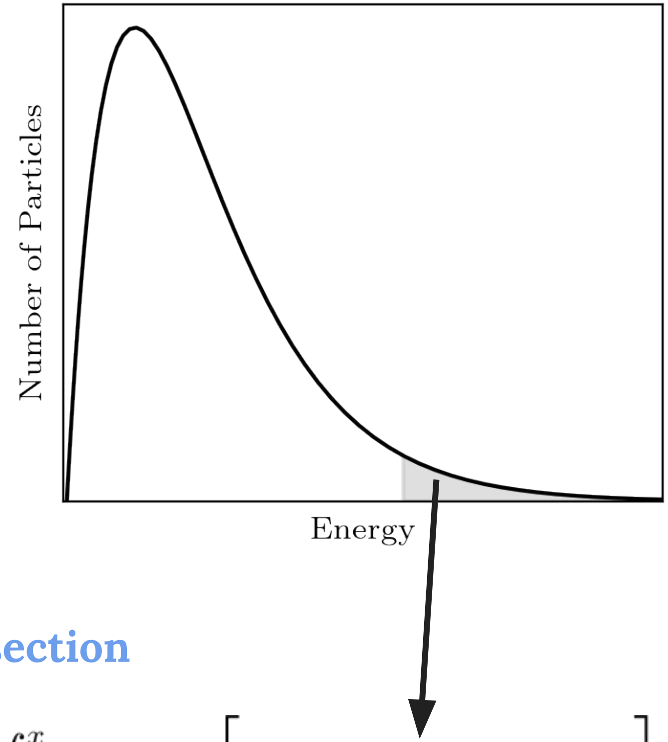
Benchmark Freeze-in Model

The Impact of the Reheating Temperature

- For $T_{rh} \ll m_\chi$: $\Gamma_{production} \sim \exp(-2m_\chi/T)$

  Kuzmin, Rubakov, 1998;
  Bringmann, Heeba, Kahlhoefer, Vangsnes 2021, Cosme, Costa, Lebedev, 2023

  - only SM particles in the tail of their velocity distributions have enough energy to annihilate into DM particles with $m_\chi \gg T$



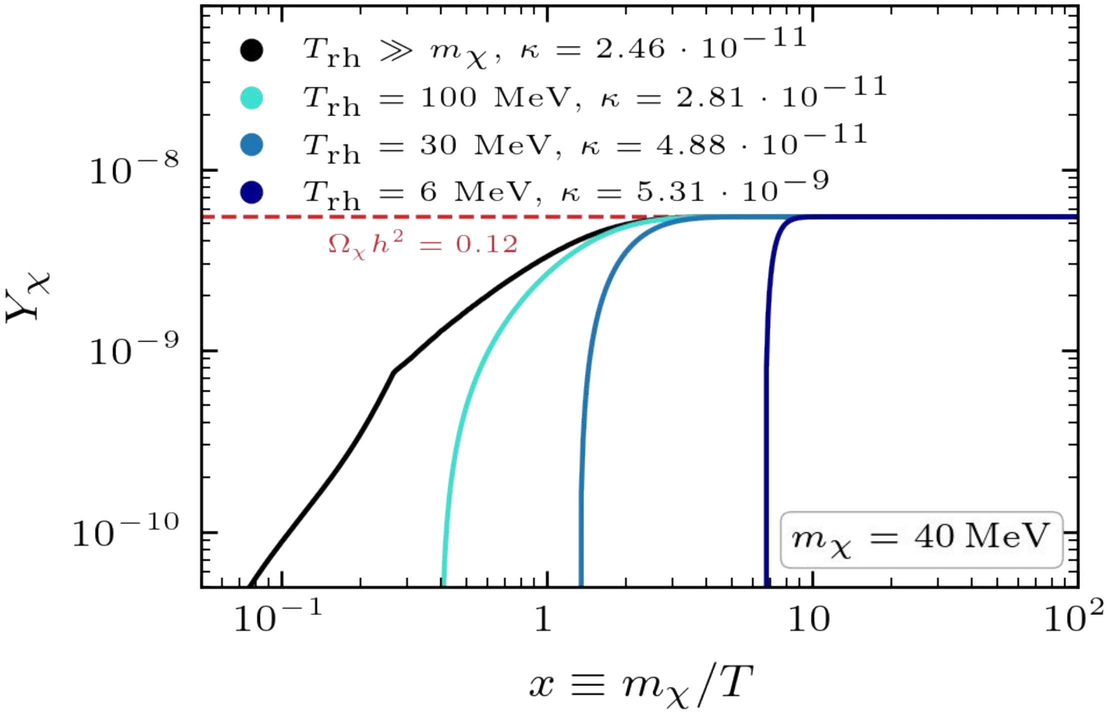
- To counteract the suppressed production and obtain the observed DM abundance today, we need:

a larger portal coupling →a larger scattering cross section

$$Y_\chi(x) = \int_{x_{\rm rh}}^{x} dx' \frac{s}{\overline{H}x'} \left[\sum_B \langle \sigma_{B\overline{B}\to\chi\overline{\chi}} v \rangle (Y_\chi^{\rm eq})^2 \right],$$

Benchmark Freeze-in Model

The Impact of the Reheating Temperature



Benchmark Freeze-in Model

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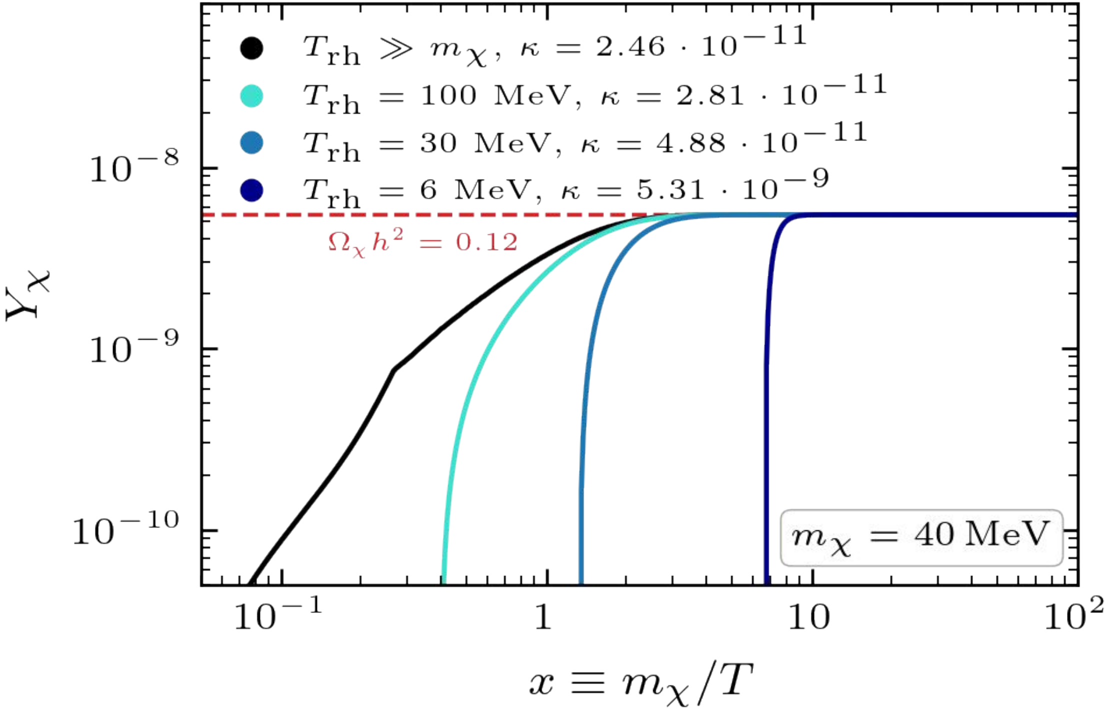


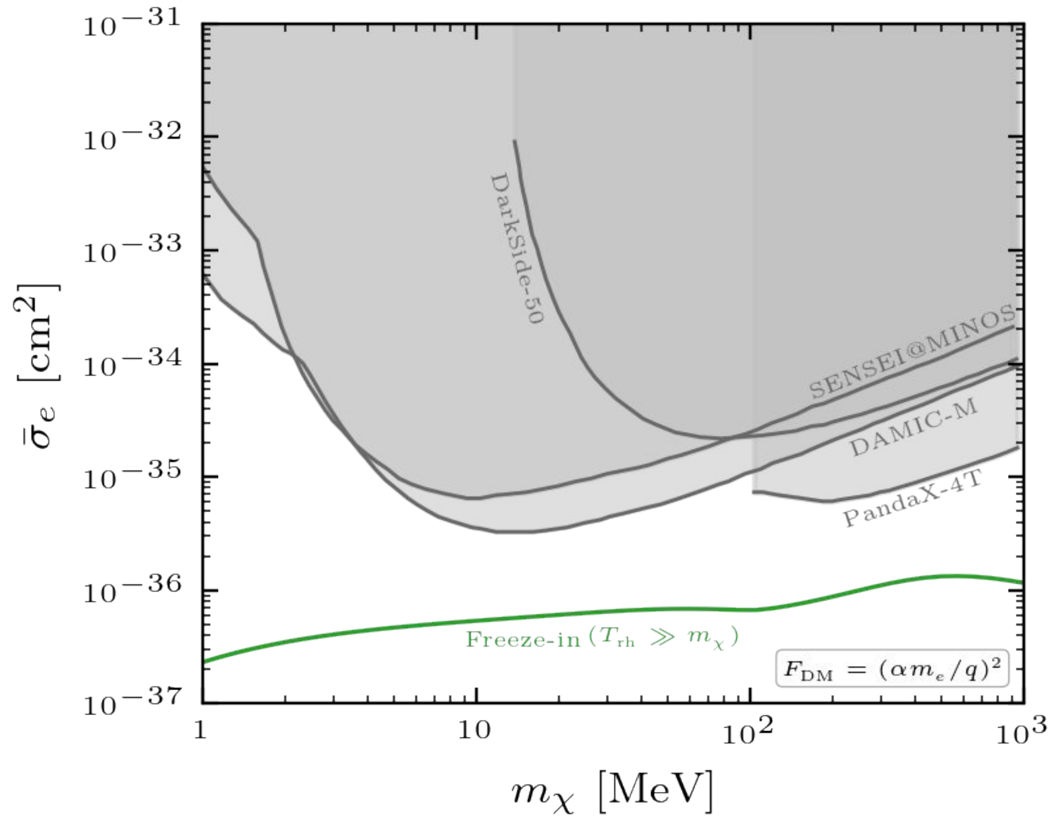
For $T_{\rm rh} << m_\chi$:

final DM abundance becomes very sensitive to $T_{\rm rh}$

$$\frac{\kappa(T_{\rm rh} \ll m_\chi)}{\kappa(T_{\rm rh} \gg m_\chi)} \sim \sqrt{x_{\rm rh}} e^{x_{\rm rh}}$$

Benchmark Freeze-in Model

The Impact of the Reheating Temperature



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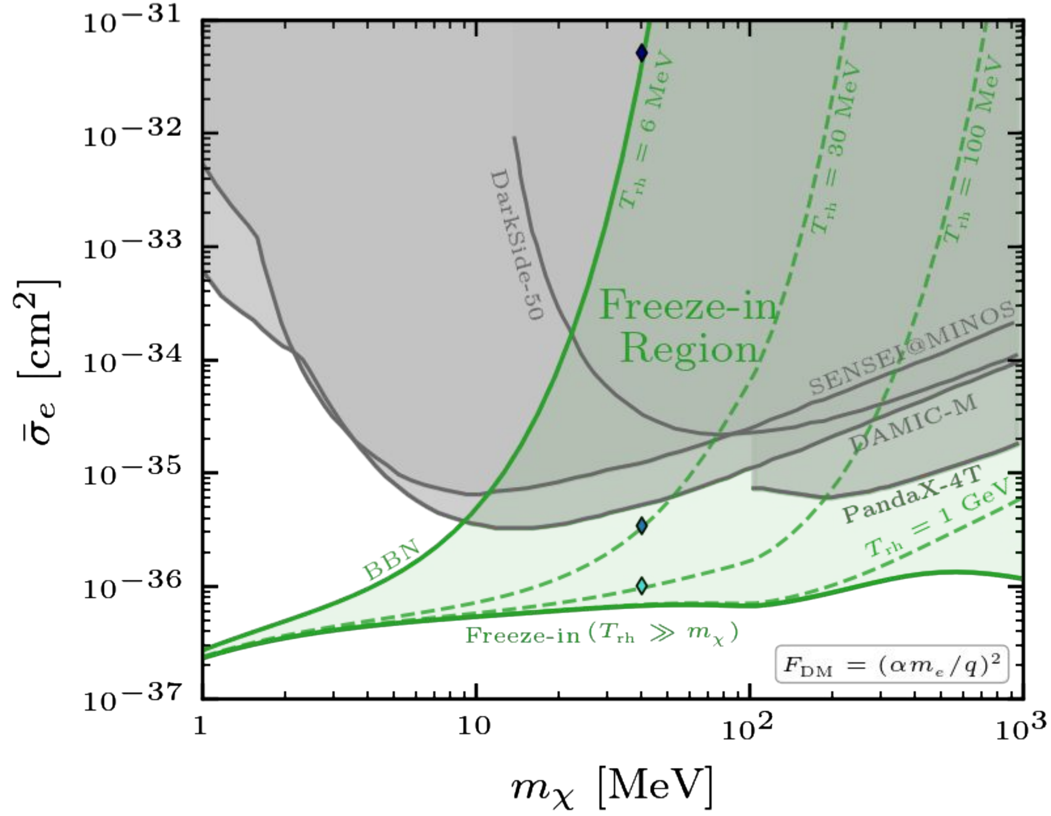
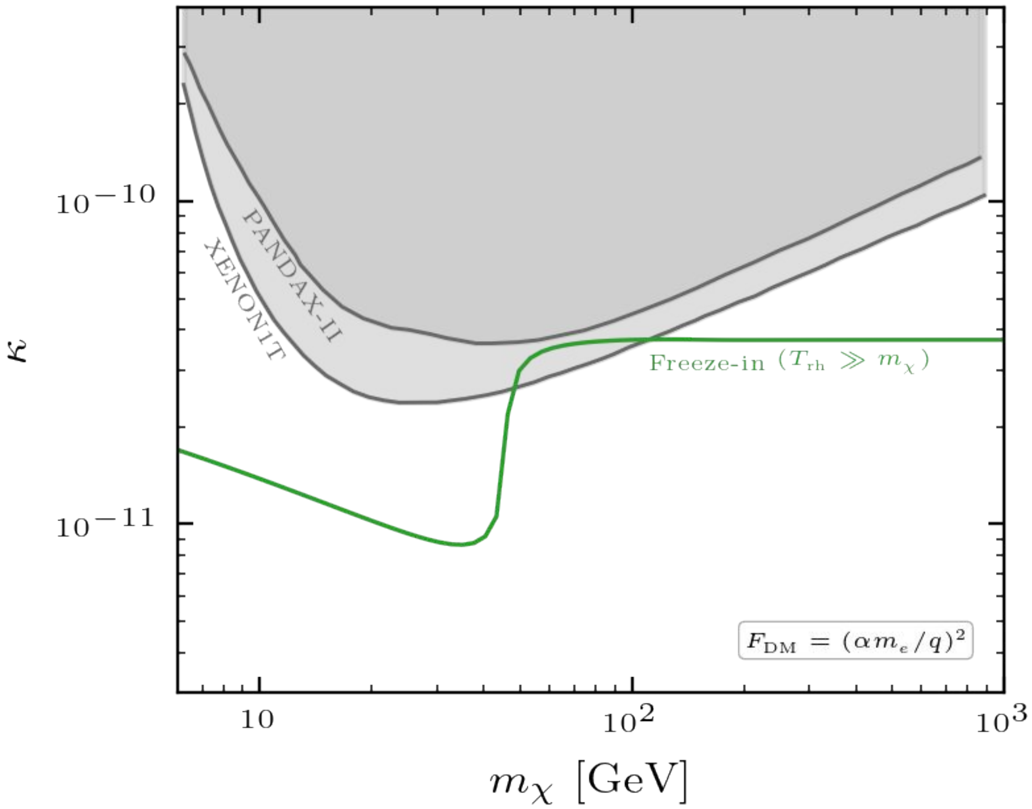


$$\frac{\kappa(T_{\rm rh} \ll m_\chi)}{\kappa(T_{\rm rh} \gg m_\chi)} \sim \sqrt{x_{\rm rh}}\, e^{x_{\rm rh}}$$

- The freeze-in benchmark should be regarded as an **extended region** defined by the reheating temperature, rather than a single curve.
- A large portion of parameter space is currently being **probed by direct detection!**

Benchmark Freeze-in Model

The Impact of the Reheating Temperature



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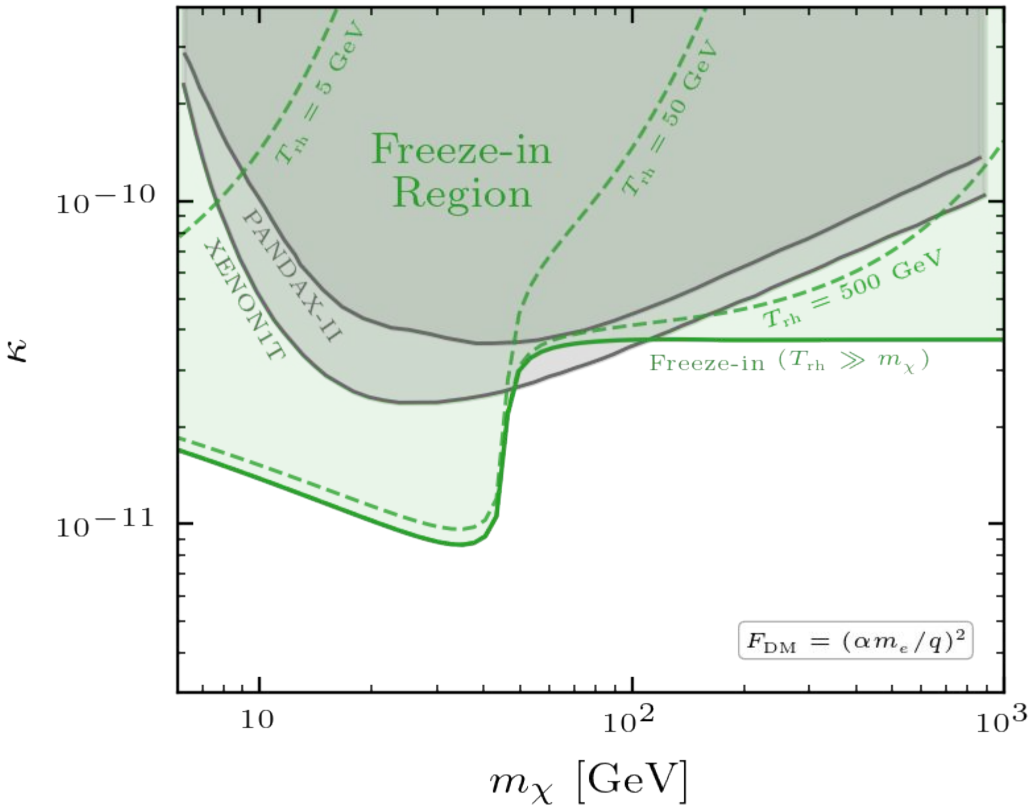


$$\frac{\kappa(T_{\rm rh} \ll m_\chi)}{\kappa(T_{\rm rh} \gg m_\chi)} \sim \sqrt{x_{\rm rh}} e^{x_{\rm rh}}$$

The same story holds for $m_\chi > 1$ GeV

A large portion of parameter space is currently being **probed by direct detection!**

Aside: Max vs Reheat Temperature

- Our work assumes that the maximum temperature of the thermal bath is equal to the reheating temperature
  - Always valid in the instantaneous reheating approximation!
  - Many examples also in the case of **finite** reheating (,)

Inflaton decays to radiation directly

Chung, Kolb, Riotto, 1998; Giudice, Kolb, Riotto, 2000; Kolb, Notari, Riotto, 2003

Inflaton decays to an unstable particle which then decays to radiation

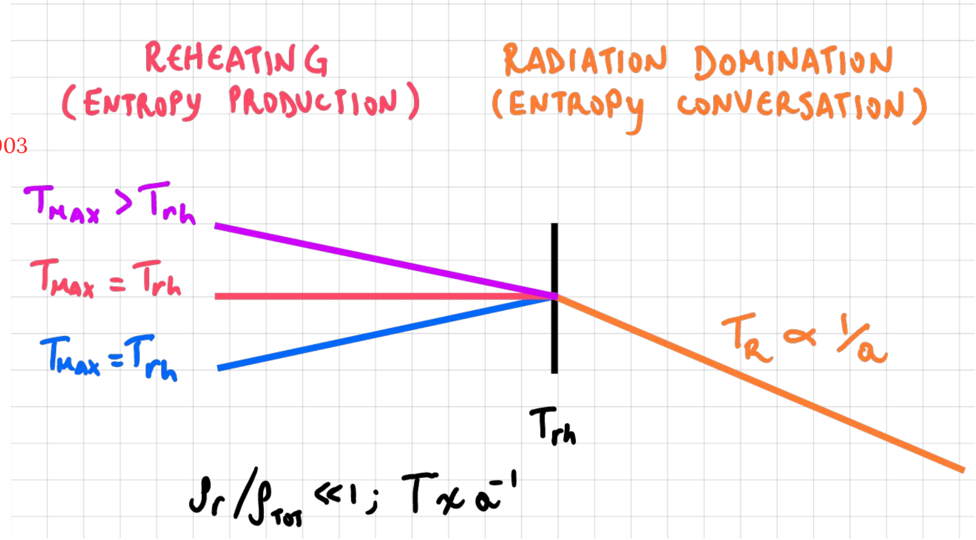
Cosme, Costa, Lebedev, 2024

Inflaton has generic dissipation rate dependent on temperature and scale factor

Co, Gonzalez, Harigaya, 2021

Resonant reheating: s-channel inflaton annihilation

Barman, Bernal, Xu, 2024



Conclusions

- We cannot neglect the impact of the reheating temperature on the benchmark freeze-in model

- For $T_{rh} << m_\chi$, DM production rate is exponentially suppressed, so that to achieve the observed relic abundance we need:

  a larger portal coupling → a larger DM-electron scattering cross section

- The freeze-in benchmark target is a **region** defined by reheating temperature rather than a single curve.

  - A large portion of parameter space is currently being **tested by direct detection!**

  - A potential future detection that lies between the current observational upper limits and the traditional freeze-in benchmark would directly probe the reheating temperature and the conditions of the universe in its earliest moments

Grazie per l'attenzione

Gabriele Montefalcone

Weinberg Institute for Theoretical Physics, University of Texas at Austin

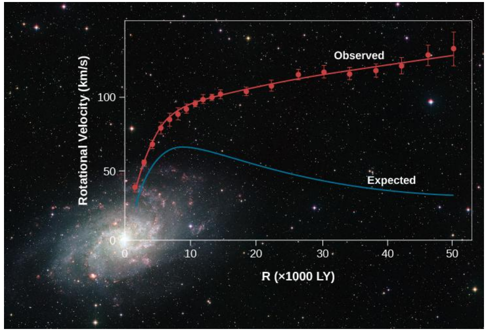
Based on work with Katherine Freese, Kimberly Boddy & Barmak Shams Es Haghi (arXiv:2405.06226)

BACK-UP SLIDES

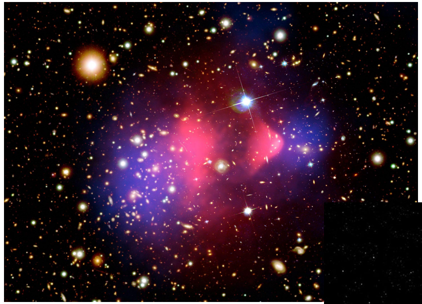
Evidence for Dark Matter (DM)

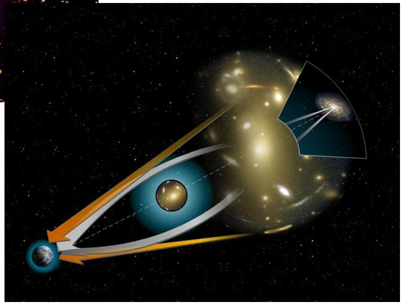
Huge amount of evidence from all scales (only from gravitational interaction)

Galactic scales

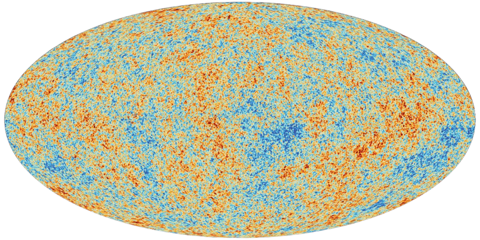


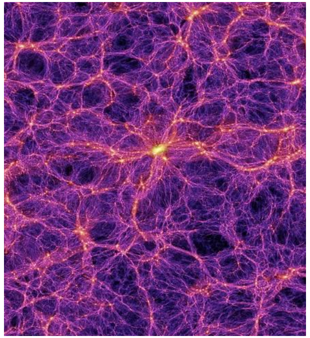
Cluster scales





Cosmological scales





What we know about DM

- Cold and Massive
- Stable/long lived
- No/weak interactions with the Standard Model (SM)
- No/weak SM charge (electric and color)
- Abundance: DM corresponds to **%25** of the energy budget in the universe today (**~5x** the amount of ordinary matter)

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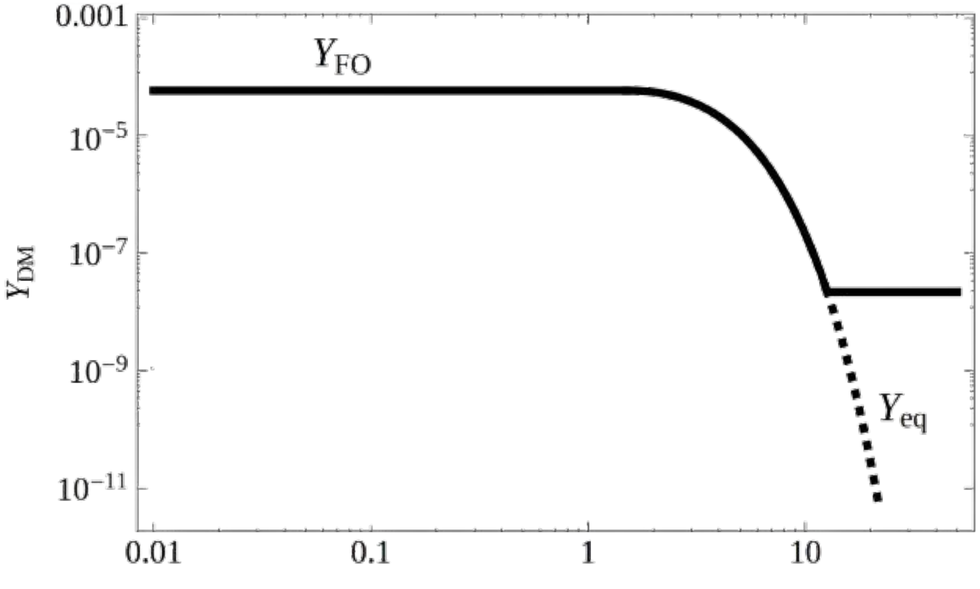
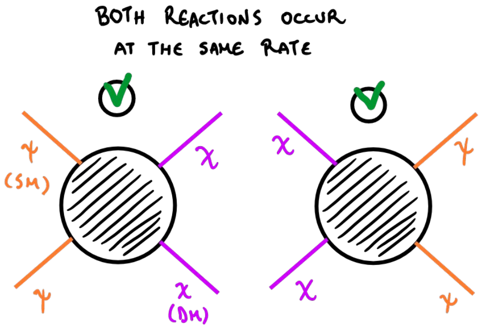
How was DM produced in the early universe?

The Canonical Freeze-out story

- DM is in thermal equilibrium with SM when $T \gg m_{DM}$
- DM freezes out at **T□m_{DM}/20**



Picture from F. Elahi



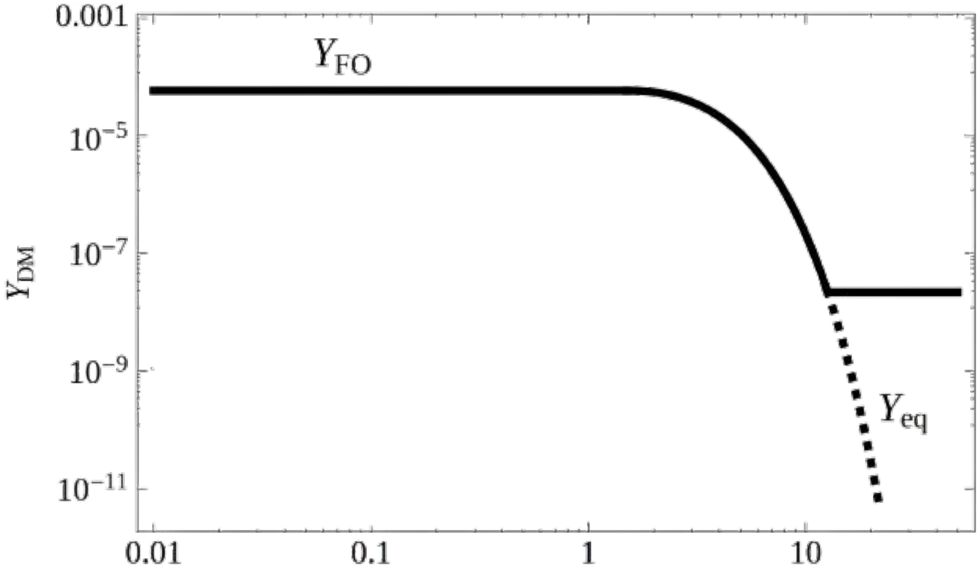
$$Y_{DM} \equiv \frac{n_{DM}}{s}$$

$$x \equiv \frac{m_{DM}}{T}$$

The Canonical Freeze-out story

The WIMP miracle!

m_{DM}□m_W and σ_{DM}□α_W^2/m_W^2 reproduces the observed DM abundance **(α_W□10^{-2}, m_W□100 GeV)**



Picture from F. Elahi

x

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Benchmark Freeze-in Model

The Kinetic Mixing Portal

$$\mathcal{L} = \mathcal{L}_{\text{SM}} - \frac{1}{4}\hat{X}_{\mu\nu}\hat{X}^{\mu\nu} + \frac{\epsilon_Y}{2}\hat{X}_{\mu\nu}\hat{B}^{\mu\nu} - e'\hat{X}_\mu\bar{\chi}\gamma^\mu\chi.$$

Dark U'(1) — SM hypercharge — dirac fermionic DM

Diagonalizing the gauge basis $\{\hat{A}_\mu, \hat{Z}_\mu, \hat{X}_\mu\}$ in terms of the mass basis $\{A_\mu, Z_\mu, A'_\mu\}$ we get

$$\hat{Z}_\mu = Z_\mu$$
$$\hat{A}_\mu = A_\mu + \epsilon A'_\mu$$
$$\hat{X}_\mu = A'_\mu - \epsilon \tan\theta_W Z_\mu$$

$$\mathcal{L} \supset -\epsilon e A'_\mu J^\mu_{\text{EM}} - e' J^\mu_{\text{DM}}\left(A'_\mu - \epsilon \tan\theta_W Z_\mu\right) + i\epsilon e\left[F'^{\mu\nu}W^+_\mu W^-_\nu - \left(\partial_\mu W^+_\nu - \partial_\nu W^+_\mu\right)A'^\mu W^{-\nu} + \left(\partial_\mu W^-_\nu - \partial_\nu W^-_\mu\right)A'^\mu W^{+\nu}\right]$$

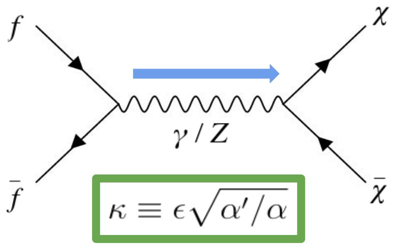
The effective kinetic mixing parameter is $\epsilon \equiv \epsilon_Y \cos\theta_W$

Below EW phase transition, we simply have: $\mathcal{L} \supset \frac{\epsilon}{2}F'_{\mu\nu}F^{\mu\nu}$

Following prescription of Bhattiprolu, McGehee, Pierce 2023

Benchmark Freeze-in Model

The Kinetic Mixing Portal



$$-\frac{\bar{H}T}{s}\frac{dY_\chi}{dT} = \sum_B \langle \sigma_{B\bar{B}\to\chi\bar{\chi}} v\rangle (Y_\chi^{\rm eq})^2,$$

$$H/\bar{H} = 1 + \frac{1}{3}\frac{d\ln g_{*,s}}{d\ln T}$$ (accounts for varying number of relativistic degrees of freedom)

Thermally averaged cross section:

$$\langle \sigma_{B\bar{B}\to\chi\bar{\chi}} v\rangle = \frac{T}{(n_\chi^{\rm eq})^2}\frac{4g_i^2}{(4\pi)^5}\int_{s_{\rm min}}^{\infty} ds\, \overline{|\mathcal{M}|}^2_{B\bar{B}\to\chi\bar{\chi}}\sqrt{1-\frac{4m_i^2}{s}}\sqrt{1-\frac{4m_\chi^2}{s}}\sqrt{s}\,K_1(\sqrt{s}/T),$$

→ We sum over all SM particles B that produce DM at tree level

$$\left\{\begin{array}{l} \{e,\ \mu,\ \tau\} \\ \{\nu_e,\ \nu_\mu,\ \nu_\tau\} \\ \{u,\ c,\ t,\ d,\ s,\ b\} \\ \{\pi^\pm, K^\pm\} \\ W^\pm \end{array}\right.$$

Following prescription of Bhattiprolu, McGehee, Pierce 2023

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$$-\frac{\bar{H}T}{s}\frac{dY_\chi}{dT} = \sum_B \langle \sigma_{B\bar{B}\to\chi\bar{\chi}} v \rangle (Y_\chi^{\rm eq})^2, \quad \langle \sigma_{B\bar{B}\to\chi\bar{\chi}} v \rangle = \frac{T}{(n_\chi^{\rm eq})^2}\frac{4g_i^2}{(4\pi)^5}\int_{s_{\rm min}}^\infty ds\, \overline{|\mathcal{M}|}^2_{B\bar{B}\to\chi\bar{\chi}} \sqrt{1-\frac{4m_i^2}{s}}\sqrt{1-\frac{4m_\chi^2}{s}}\sqrt{s}\, K_1(\sqrt{s}/T),$$

For fermions **f**: $\{e,\ \mu,\ \tau\}$ leptons $\{\nu_e,\ \nu_\mu,\ \nu_\tau\}$ neutrinos $\{u,\ c,\ t,\ d,\ s,\ b\}$ quarks

$$\overline{\mathcal{M}}^2_{f\bar{f}\to\chi\bar{\chi}} = \frac{32}{3}\pi^2\alpha^2\kappa^2 N_f\left(s+2m_\chi^2\right)\left[\frac{Q_f^2}{s^2}\left(s+2m_f^2\right) - 2Q_f V_f \tan\theta_W \frac{\left(s+2m_f^2\right)\left(s-m_Z^2\right)}{s\left[\left(s-m_Z^2\right)^2+m_Z^2\Gamma_Z^2\right]} + \tan^2\theta_W \frac{V_f^2\left(s+2m_f^2\right)+A_f^2\left(s-4m_f^2\right)}{\left(s-m_Z^2\right)^2+m_Z^2\Gamma_Z^2}\right],$$

For scalars **ϕ**: $\{\pi^\pm, K^\pm\}$

$$\overline{\mathcal{M}}^2_{\phi^+\phi^-\to\chi\bar{\chi}} = \frac{32}{3}\pi^2\alpha^2\kappa^2\left(1+\frac{2m_\chi^2}{s}\right)\left(1-\frac{4m_\phi^2}{s}\right),$$

For the **W** boson:

$$\overline{\mathcal{M}}|^2_{W^+W^-\to\chi\bar{\chi}} = \frac{8}{27}\pi^2\alpha^2\kappa^2\left(\frac{m_Z}{m_W}\right)^4 \frac{\left(s+2m_\chi^2\right)\left(s-4m_W^2\right)\left(s^2+20sm_W^2+12m_W^4\right)}{s^2\left[\left(s-m_Z^2\right)^2+m_Z^2\Gamma_Z^2\right]},$$

Following prescription of Bhattiprolu, McGehee, Pierce 2023