NEWS ON LEPTOGENESIS

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INVISIBLES24

LEPTOGENESIS

GREATEST HITS

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DISCLAIMER: this is NOT a comprehensive review of leptogenesis

See for instance:

Buchmüller, Peccei, Yanagida, 2005; Davidson, Nardi, Nir, 2008; Fong, Nardi, Riotto, 2012; Garbrecht, Molinaro et al., 2018

And also Talks by Miguel Vanvlasselaer and Ninetta Saviano

Outline

- Introduction
- Leptogenesis via HNL decay
- Low scale leptogenesis beyond the seesaw
- ν electroweak baryogenesis
- Leptogenesis via heavy neutrino oscillations
- Summary and outlook

1. Introduction

- Baryon number density: determined from
  - Big Bang Nucleosynthesis: primordial abundances of light elements (D, ^3He, ^4He, ^7Li) mainly depend on one parameter, n_B/n_γ
  - CMB anisotropies

$$Y_B = \frac{n_B - n_{\bar{B}}}{s} \simeq \frac{n_B}{s} = (8.66 \pm 0.01) \times 10^{-11}$$

Impressive consistency between both determinations, completely independent !

- Baryon asymmetry
  - nucleons and antinucleons were in thermal equilibrium up to $T_{fo} \approx 22$ MeV, when $\Gamma_{ann} < H$
  - If the Universe were locally baryon-symmetric:

    $Y_{Bfo} < 10^{-20}$ → there was a baryon asymmetry
- Sakharov's conditions to dinamically generate the baryon asymmetry (BAU)
  - Baryon number violation
  - C and CP violation
  - Departure from thermal equilibrium

2. Leptogenesis via HNL decay

- BAU generated in the decay of heavy Majorana neutrinos: Fukugita, Yanagida, 1986

  - Out of equilibrium decay
  - L and CP violating interactions ➔ lepton asymmetry, ΔL
  - (B+L)-violating, but (B-L) conserving, non-perturbative sphaleron interactions ΔL ➔ ΔB

- Non-equilibrium process ➔ Boltzmann eqs.

$$\frac{dY_{N_1}}{dz} = -\left(\frac{Y_{N_1}}{Y_{N_1}^{eq}} - 1\right)(D+S)$$

$$\frac{dY_{B-L}}{dz} = -\epsilon\left(\frac{Y_{N_1}}{Y_{N_1}^{eq}} - 1\right) D - Y_{B-L} W$$

$$z \equiv M_1/T$$

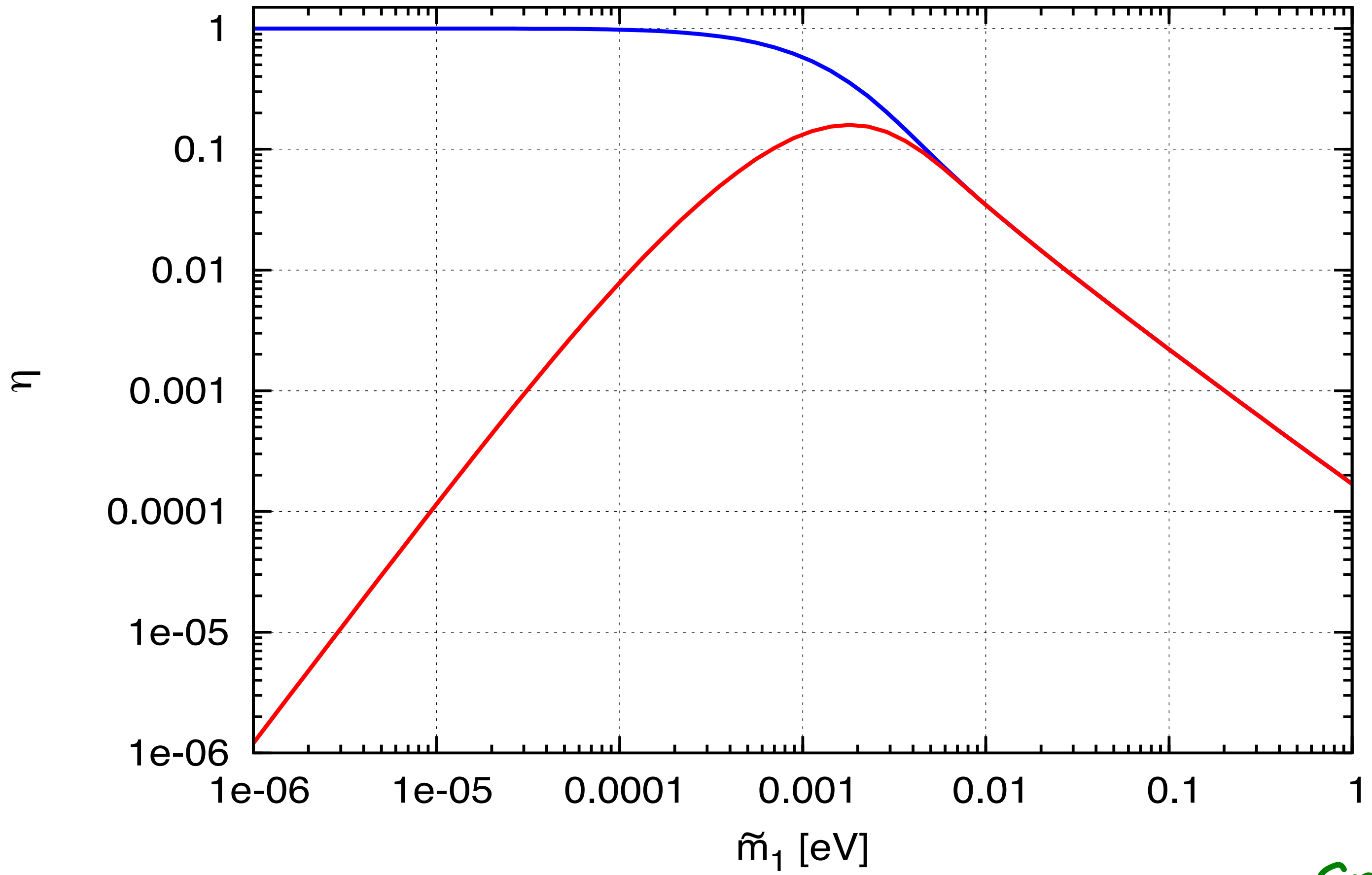
$$\epsilon = \sum_\alpha \epsilon_{\alpha 1} = \sum_\alpha \frac{\Gamma(N_1 \to \ell_\alpha h) - \Gamma(N_1 \to \bar{\ell}_\alpha \bar{h})}{\sum_\beta \Gamma(N_1 \to \ell_\beta h) + \Gamma(N_1 \to \bar{\ell}_\beta \bar{h})}$$

- Final baryon asymmetry:

$$Y_B = -\kappa\,\epsilon\,\eta$$

$$\kappa = \frac{28}{79} Y^{eq}_{N_1}(T \gg M_1) \sim 10^{-3}$$

η = efficiency : $0 \leq \eta \leq 1$



— $Y_N^i = 0$ — $Y_N^i = Y_N^{eq}$

Credit: J. Racker

$$\tilde{m}_1 \equiv \frac{(\lambda^\dagger \lambda)_{11} v^2}{M_1}$$

Related to the contribution of N_1 to light neutrino masses

- η maximum for

$$\tilde{m}_1 = m_* = \frac{16}{3\sqrt{5}}\pi^{5/2}\sqrt{g_*}\frac{v^2}{M_P} \sim 10^{-3}\,\mathrm{eV}$$

m_*, defined by: $$\frac{\Gamma_N}{H(T=M_1)} = \frac{\tilde{m}_1}{m_*}$$

determines the amount of departure from thermal equilibrium and the strength of the washouts:

$\tilde{m}_1 \gg m_*$ ➔ strong washout:

- independence of initial conditions, $\eta \propto 1/\tilde{m}_1$

$\tilde{m}_1 \ll m_*$ ➔ weak washout:

- depends on initial conditions, if $Y_N^i = 0 \;\rightarrow \eta \propto \tilde{m}_1^2$

- Hierarchical heavy neutrinos: $\epsilon \sim \frac{3}{16\pi} \frac{\lambda_{\alpha 2}^2}{M_2} M_1$

- Connection to light neutrino masses (type I seesaw):

$$|\epsilon| \leq \epsilon_{DI} = \frac{3}{16\pi} \frac{M_1}{v^2} (m_3 - m_1)$$

Davidson, Ibarra, 2002

➔ $M_1 \gtrsim 10^9 \, \mathrm{GeV}$

Detailed numerical analysis solving BEs:

➔ $M_1 \gtrsim (4\times) 10^8 \mathrm{GeV}$ for fine-tuned regions

Hambye et al. 2004

➔ bound on light neutrino masses, $m_\nu < 0.15$ eV

Buchmüller, Di Bari, Plümacher, 2004

Flavour effects

- At $T \lesssim 10^{12}$ GeV, the τ Yukawa interaction is fast, and there are (in general) 2 lepton flavour asymmetries evolving almost independently

- At $T \lesssim 10^{9}$ GeV, both τ and μ Yukawa interactions are in equilibrium ➔ 3 independent lepton flavour asymmetries, $Y_{\Delta(B/3-L\alpha)}$

Barbieri et al. 2000; Endoh et al. 2004; Abada et al. 2006; Nardi et al. 2006

Some consequences:

★Flavoured asymmetries ε_α depend on U_{PMNS} phases although in general leptogenesis is "insensitive" to them, even in SUSY Davidson, Garayoa, Palorini, NR, 2007

★ Bound on light neutrino masses $m_\nu < 0.15$ eV evaded

★ N_2 leptogenesis can survive N_1 washouts more easily:

➔ relevant for SO(10) models which predict $M_1 \ll 10^9$ GeV

★ Leptogenesis possible with $\epsilon = \sum_\alpha \epsilon_\alpha = 0$

➔ relevant for models with small B-L violation (inverse seesaw)

$$\epsilon_{\alpha i} \equiv \frac{\Gamma(N_i \to \ell_\alpha h) - \Gamma(N_i \to \bar{\ell}_\alpha \bar{h})}{\sum_\beta \Gamma(N_i \to \ell_\beta h) + \Gamma(N_i \to \bar{\ell}_\beta \bar{h})} = \epsilon^{\not L}_{\alpha i} + \epsilon^{L}_{\alpha i}$$

where:

Covi, Roulet, Vissani, 1996

$$\epsilon^{\not L}_{\alpha i} = \sum_{j \neq i} f(a_j) \mathrm{Im}[\lambda^*_{\alpha j} \lambda_{\alpha i} (\lambda^\dagger \lambda)_{ji}]$$

$$a_j \equiv M_j^2 / M_i^2$$

$$\epsilon^{L}_{\alpha i} = \sum_{j \neq i} g(a_j) \mathrm{Im}[\lambda^*_{\alpha j} \lambda_{\alpha i} (\lambda^\dagger \lambda)_{ij}]$$

related to L conserving d=6 operators ➔ escape the DI bound because they are not linked to neutrino masses (LNV d=5 Weinberg operator)

Resonant leptogenesis

- Enhancement of the CP asymmetry for degenerate neutrinos, $M_2 - M_1 \approx \Gamma_2$

$$|\epsilon| \sim \frac{1}{2} \frac{\mathrm{Im}[(\lambda^\dagger \lambda)^2_{21}]}{(\lambda^\dagger \lambda)_{11} (\lambda^\dagger \lambda)_{22}} \leq \frac{1}{2}$$

Covi, Roulet, 1997; Pilaftsis, 1997; Pilaftsis, Underwood 2004

➔ EW scale e-,μ-,τ- leptogenesis with observable LFV

Pilaftsis 2005; Deppisch, Pilaftsis 2011

- Approximate flavour symmetries + universal RHN masses at the GUT scale ➔ heavy neutrino mass splittings radiatively generated

Non-equilibrium QFT:

Kadanoff-Baym equations for spectral functions and statistical propagators of leptons and Majorana neutrinos Anisimov et al., 2011; Drewes et al. 2013

★ Free from double counting problem

★ Relevant in the weak washout regime

★ Very important for resonant leptogenesis: supression wrt classical Boltzmann result.

Garny, et al., 2013

- If $M_2 - M_1 \approx \Gamma_2$ ➔ sterile neutrino oscillations

- Taken into account using "Flavour covariant transport equations" ➔ density matrix formalism

$$\dot{\rho} = -i[H, \rho]$$

Dev et al., 2014 (109 p.)

- Identify mixing contribution from diagonal ρ_N and heavy neutrino oscillation contribution from off-diagonal $(\rho_N)_{12}$

- Also in the Kadanoff-Baym approach Dev et al., 2015

3. Low scale leptogenesis beyond the seesaw

Baryogenesis from particle decays or annihilations at low T (TeV scale)

- Potentially testable mechanisms for baryogenesis
- Avoid hierarchy and gravitino (in SUSY) problems
- High scale baryogenesis disfavored if L violating processes were observed at LHC
- Common origin of dark and baryon matter?

$$\frac{\Omega_{DM}}{\Omega_B} \approx 5$$

(see A. Hook's talk)

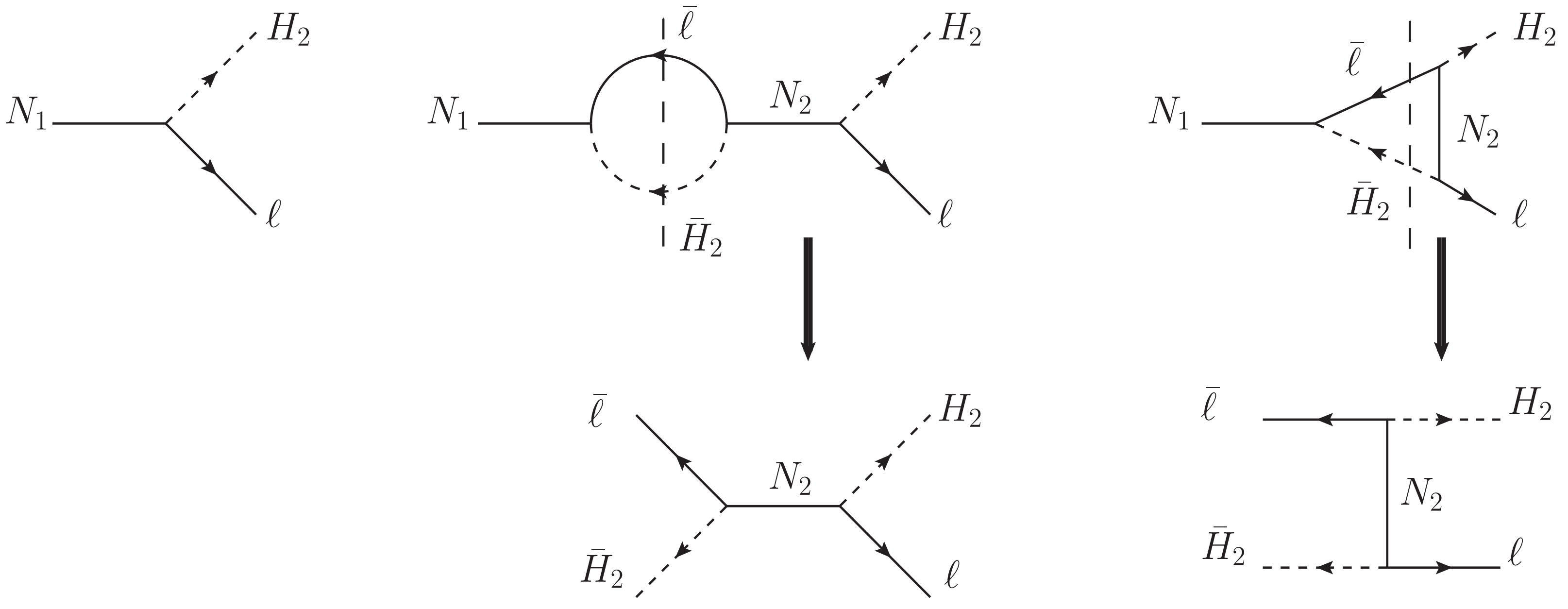
- Radiative neutrino mass models (like scotogenic model) ➔ no DI bound on ε

- Generic problem of baryogenesis at low scales, independent of the connection to neutrino masses: fast washout of the asymmetries

$$\epsilon \propto \frac{\lambda_{\alpha 2}^2}{M_2} M_1$$

if $M_2 \gg M_1$

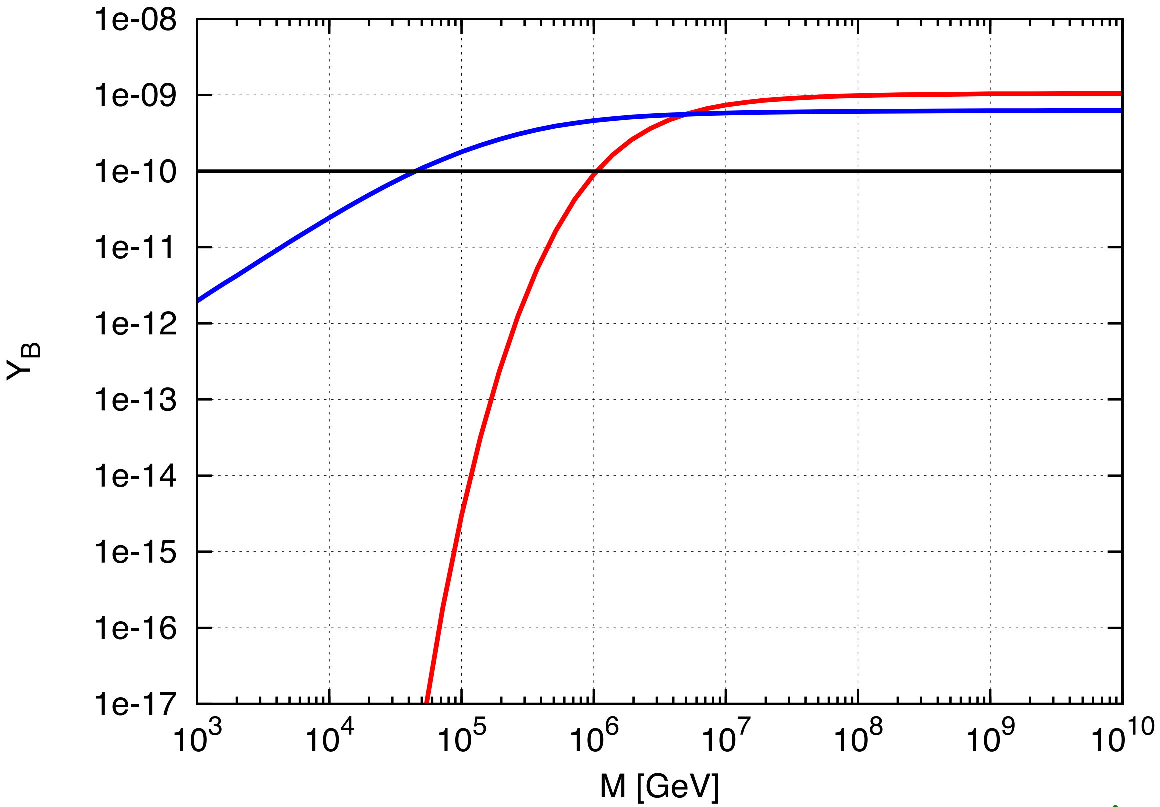
washouts $\propto \left(\frac{\lambda_{\alpha 2}^2}{M_2} \right)^2$



Credit: J. Racker, 2014

_____ without washouts _____ with washouts

$$H_2\ell \leftrightarrow \overline{H}_2\bar{\ell}, \qquad \ell\ell \leftrightarrow \overline{H}_2\overline{H}_2$$



$$\frac{\Gamma_N}{H(T=M_N)} = const.$$

$$\epsilon_{CP} = const.$$

Courtesy of J. Racker, 2014

Ways out:

- Late decay: new interactions needed to produce N_1
  [Plumacher 1997; Racker, Roulet 2009, Cui,Sundrum 2012]
- Resonant leptogenesis
- Massive decay products:
  [Cui,Randall,Shuve, 2012; J. Racker, 2014]
- ALP Leptogenesis [talk by Martina Cataldi]

4. ν electroweak baryogenesis

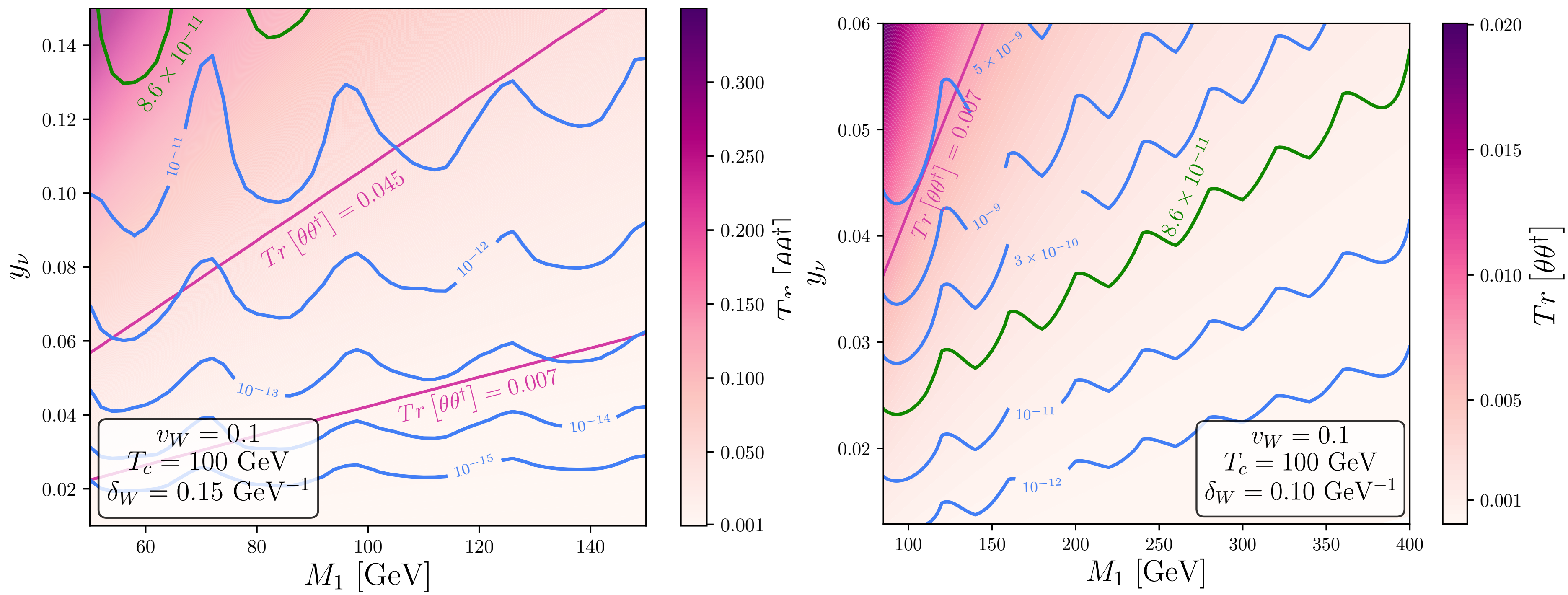
- Extra SM singlet scalar provides both: strongly first order EW phase transition (SFOPT) and sterile neutrino (Dirac) masses
- Inverse or linear seesaw ➔ large neutrino Yukawa couplings

$$\mathcal{L} = -\overline{L}_L \tilde{H} Y_\nu N_R - \overline{N}_L \phi Y_N N_R + h.c. - V(\phi^* \phi, H^\dagger H)$$

- Profiles of the vevs $v_H(z)$ and $v_\phi(z)$ along the bubble wall must be different

P. Hernández, NR, 1997

Importance of flavour effects: $\theta \equiv m_D M_N^{-1} = \frac{1}{\sqrt{2}} Y_\nu v_H Y_N^{-1} v_\phi$



Unflavoured　　　With flavour

Regions of SFOPT consistent with current experimental bounds also identified

E. Fernández-Martínez et al. 2020, 2023

5. Leptogenesis via sterile neutrino oscillations

Akhmedov, Rubakov, Smirnov, 1998; Asaka,Shaposhnikov 2005; Antusch et al. 2018; Hernández et al. 2016; Drewes et al. 2018; Abada et al. 2019; Domcke et al. 2020; etc

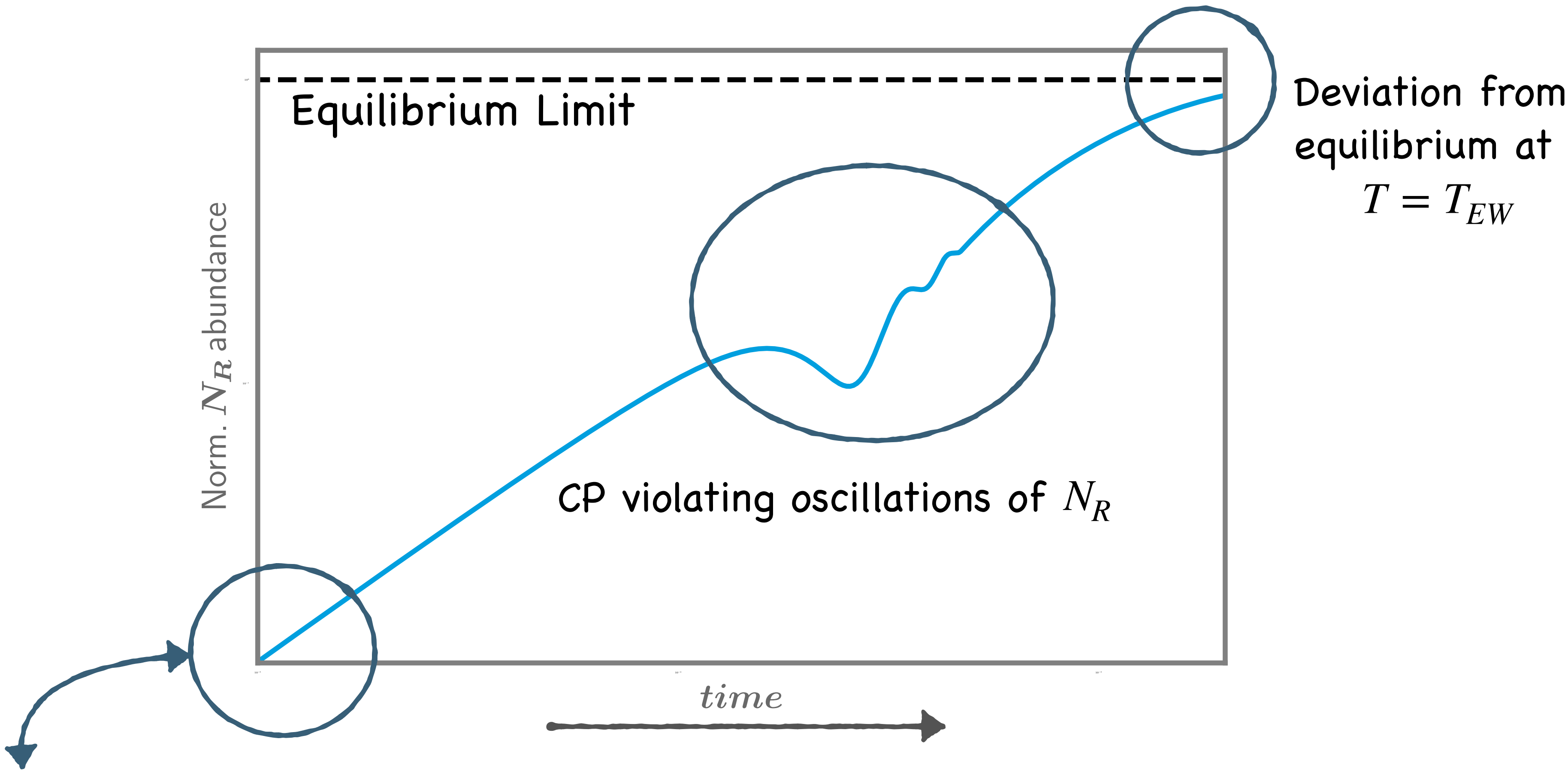
Sakharov conditions for baryogenesis:

- CP violating phases in Y, M
- B violated by sphaleron processes at $T > T_{EW}$
- At least one sterile neutrino does **NOT** equilibrate by T_{EW} , i.e. for some rate

$$\Gamma_i (T_{EW}) \leq H_u(T_{EW}) = T_{EW}^2 / M_P^*$$

Fulfilled for M = O(GeV), Y ~ 10^{-6} – 10^{-7}, in the correct range to explain neutrino masses ! Freeze-in baryogenesis

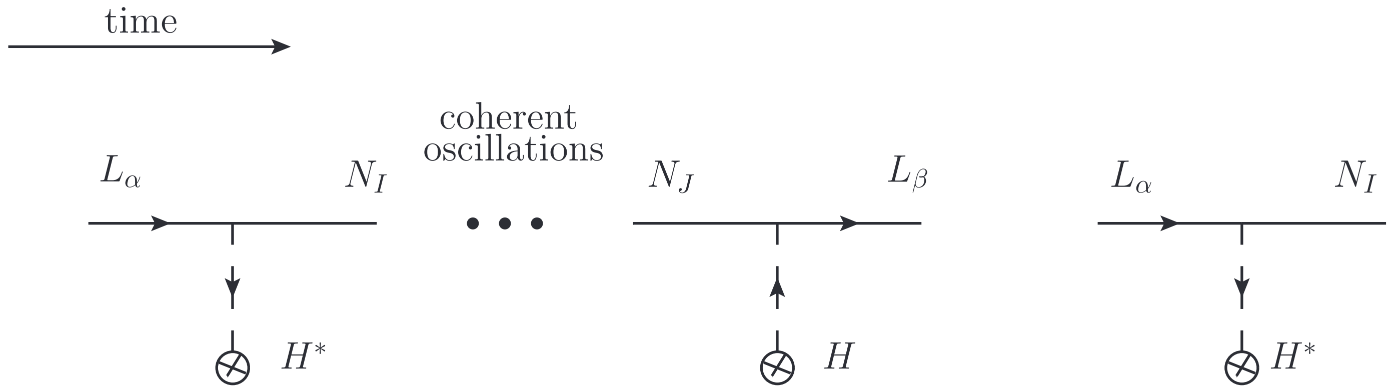
Schematic evolution of N_R abundance



Courtesy of S. Sandner

- Basic stages:

Shuve, Yavin 2014



$Y_{\Delta L_1} = 0$	$Y_{\Delta L_1} > 0$	$Y_{\Delta L_1} > 0$
$Y_{\Delta L_2}, Y_{\Delta L_3} = 0$	$Y_{\Delta L_2}, Y_{\Delta L_3} < 0$	$Y_{\Delta L_2}, Y_{\Delta L_3} < 0$
$\sum_\alpha Y_{\Delta L_\alpha} = 0$	$\sum_\alpha Y_{\Delta L_\alpha} = 0$	$\sum_\alpha Y_{\Delta L_\alpha} \neq 0$
Out of equilibrium HNL production	Asymmetries in lepton flavours	Different washout of flavoured asymmetries

- Inclusion of LNV (helicity conserving, HC) rates, suppressed by $(M/T)^2$

Hambye, Teresi 2016,2017

Density matrix formalism(*)

Raffelt & Sigl, 1993

$$\dot{\rho} = -i[H,\rho] - \frac{1}{2}\{\Gamma^a,\rho\} + \frac{1}{2}\{\Gamma^p,\rho_{eq}-\rho\}$$

- Hamiltonian term: $H = \frac{M^2}{2k^0} + \frac{T^2}{8k^0} Y^\dagger Y$
- Annihilation and production rates of the N's: Γ^a, Γ^p
- For antineutrinos: $\bar{\rho}$, H ⟹ H*
- Diagonal density matrix for SM leptons, which are in thermal equilibrium, with chemical potential

$$f_\alpha(k^0) = \frac{1}{e^{(k^0-\mu_\alpha)/T}+1}$$

- For antileptons $\mu_\alpha \Longrightarrow -\mu_\alpha$

(*)Similar results in Closed-time-path formalism Drewes et al.

Time scales and slow modes

$$\dot{\rho} = -i[H,\rho] - \frac{1}{2}\{\Gamma^a,\rho\} + \frac{1}{2}\{\Gamma^p,\rho_{eq}-\rho\}$$

- Annihilation and production rates of the N's: at T >> M,

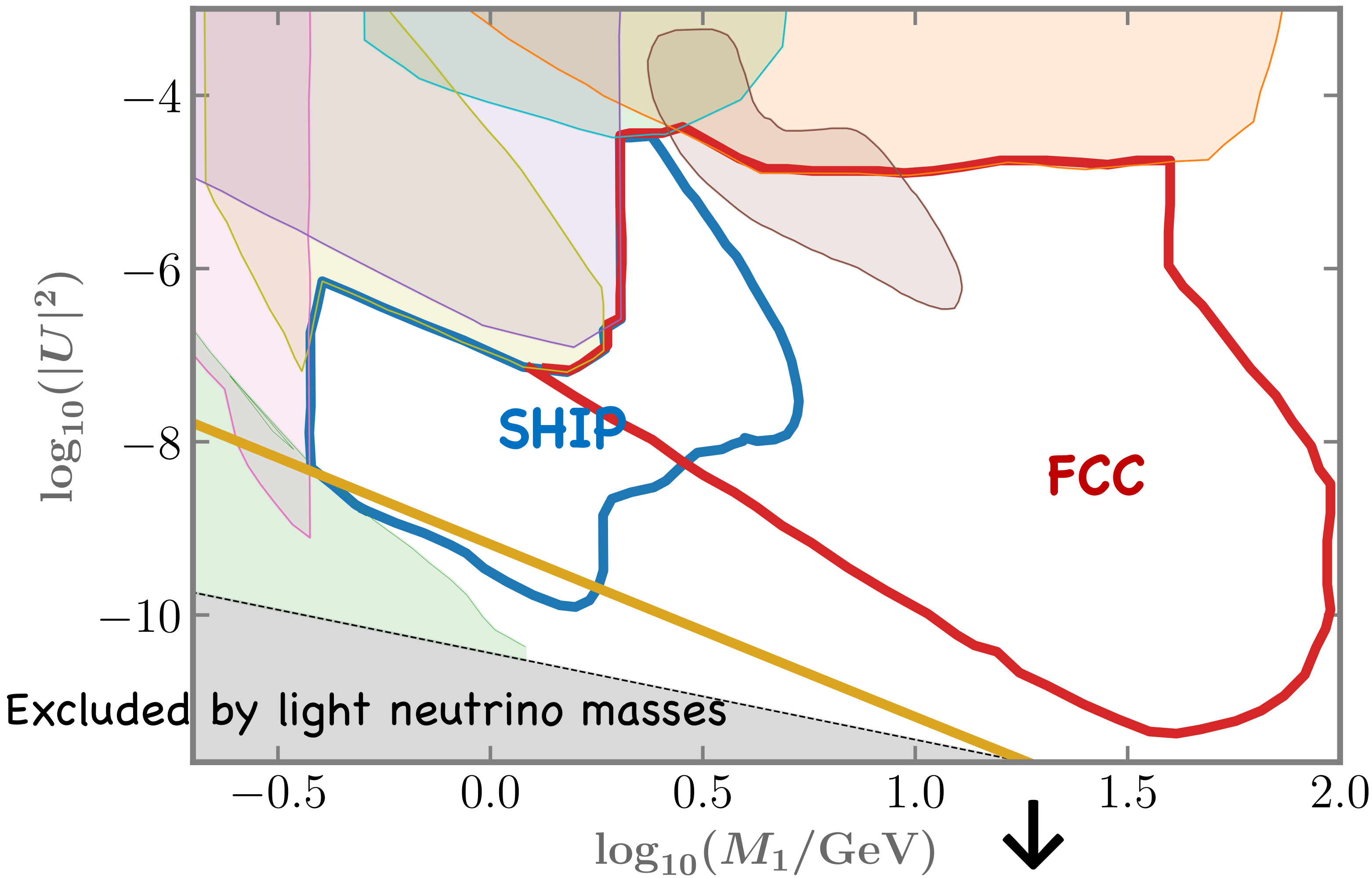
$$\Gamma(T) \propto \mathrm{Tr}[YY^\dagger]\,T$$

Ghiglieri, Laine, 2017

- Flavoured rates: $\Gamma_\alpha(T) \propto \epsilon_\alpha \Gamma(T)$ $\epsilon_\alpha \equiv \frac{(YY^\dagger)_{\alpha\alpha}}{\mathrm{Tr}[YY^\dagger]}$
- Oscillation rate: $\Gamma_{osc}(T) \propto \frac{\Delta M^2}{T}$
- LNV (HC) rate: $\Gamma_M \propto (M/T)^2\Gamma(T)$

N=2 HNL

Hernández, López-Pavón, NR, Sandner, 2022



$\Gamma\,(T_{EW}) = H_u(T_{EW})$

Approximately conserved lepton number limit

- Inverse seesaw Wyler, Wolfenstein 1983;Mohapatra, Valle 1986

$$M = \begin{pmatrix} \mu_1 & \Lambda \\ \Lambda & \mu_2 \end{pmatrix}. \qquad Y = \begin{pmatrix} y_e & y'_e e^{i\beta'_e} \\ y_\mu & y'_\mu e^{i\beta'_\mu} \\ y_\tau & y'_\tau e^{i\beta'_\tau} \end{pmatrix}$$

$y'_\alpha << y_\alpha$, $\mu_i << \Lambda$

- Once neutrino masses and mixings are fixed, there are 6 free parameters:

- $y^2 \equiv \sum_\alpha y_\alpha^2$, or, equivalently $U^2 \simeq \frac{y^2 v^2}{2\Lambda^2}$

- Three independent phases ($\mu_1 = \mu_2 \equiv \mu$ can be chosen real)

- In terms of physical HNL masses:

$\Lambda = (M_1 + M_2)/2 = M$, $\mu = (M_2 - M_1)/2 = \Delta M/2$

Analytic approach

- Identify the non-thermal modes and their characteristic time scales
- Solve the equations perturbatively, exploiting the weakly coupled modes
- Identify the CP-invariants that control Y_B

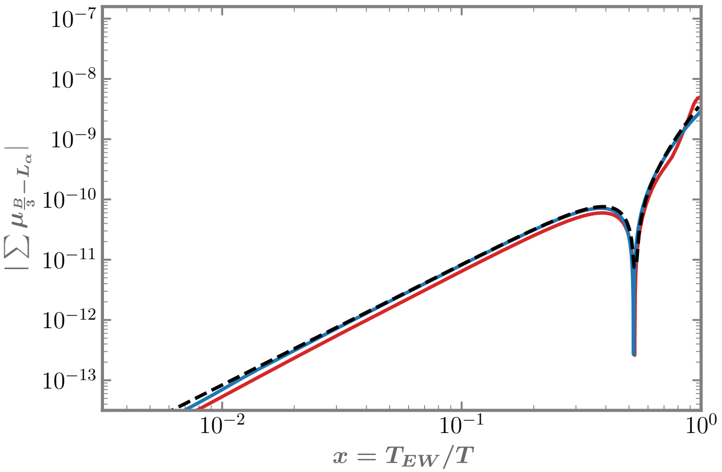
$$I_0 = \text{Im}\left(\text{Tr}\left[Y^\dagger Y M^\dagger M Y^\dagger Y_\ell Y_\ell^\dagger Y\right]\right)$$

$$\rightarrow \sum_\alpha y_{\ell_\alpha}^2 \sum_{i<j} (M_j^2 - M_i^2)\, \text{Im}\left[Y_{\alpha j}^* Y_{\alpha i} \left(Y^\dagger Y\right)_{ij}\right] \equiv \sum_\alpha y_{\ell_\alpha}^2 \Delta_\alpha$$

$$I_1 = \text{Im}\left(\text{Tr}\left[Y^\dagger Y M^\dagger M M^* Y^T Y^* M\right]\right)$$

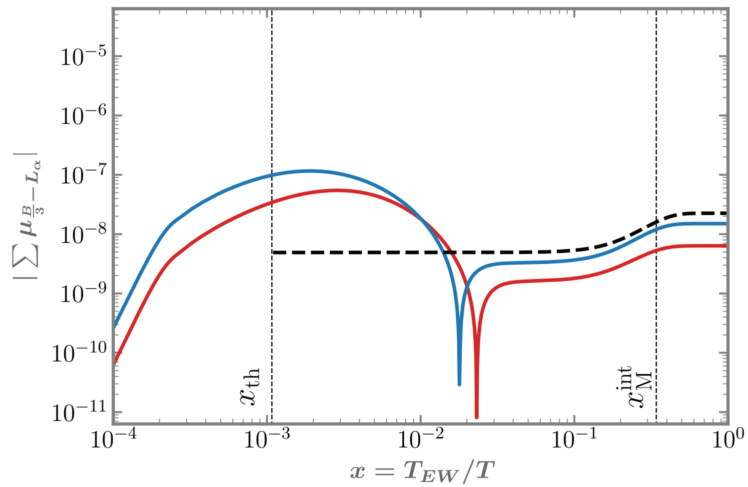
$$\rightarrow \sum_\alpha \sum_{i<j} (M_j^2 - M_i^2)\, M_i M_j \text{Im}\left[Y_{\alpha j} Y_{\alpha i}^* \left(Y^\dagger Y\right)_{ij}\right] \equiv \sum_\alpha \Delta_\alpha^M$$

Overdamped regime:



$$\epsilon \propto \frac{\Delta M^2/T}{\Gamma(T)} \ll 1$$

Slow flavour α :



$\Gamma_\alpha(T_{EW}) < H_u(T_{EW}) < \Gamma(T_{EW})$

---- analytical solution:

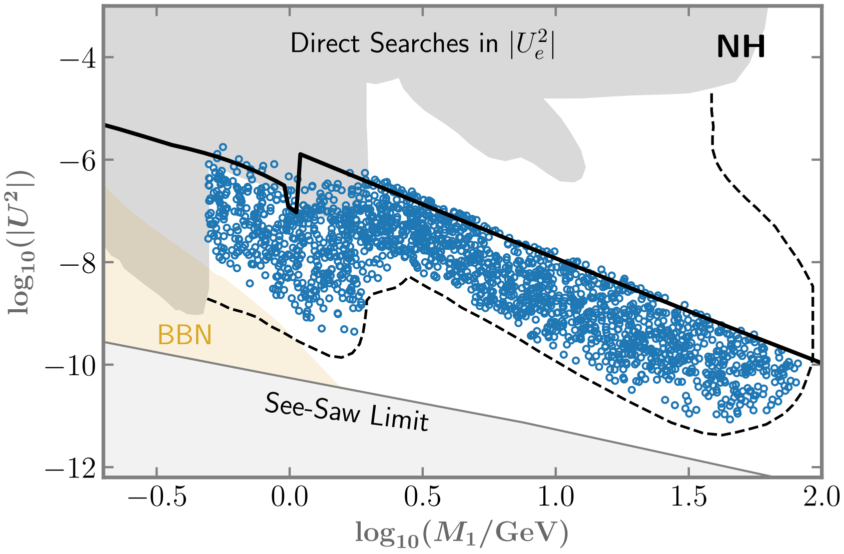
Perturbing in y' and in $(M/T)^2$

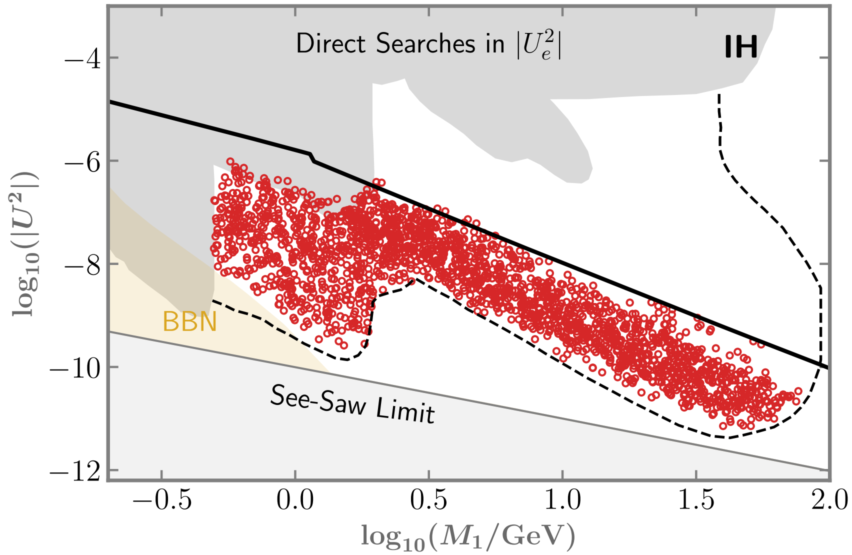
Linearized equations

—— numerical solution with same approximations

—— full numerical solution

Numerical scan





Absolute upper bound on U^2 from the overdamped regime:

- Weak LNV
  $M \lesssim O(1\ \mathrm{GeV})$

$$(U^2)_{\mathrm{ov}}^{\mathrm{wLNV}} \lesssim 4(17) \times 10^{-7} \left(\frac{1\,\mathrm{GeV}}{M}\right)^{4/3} \quad \mathrm{NH(IH)}$$

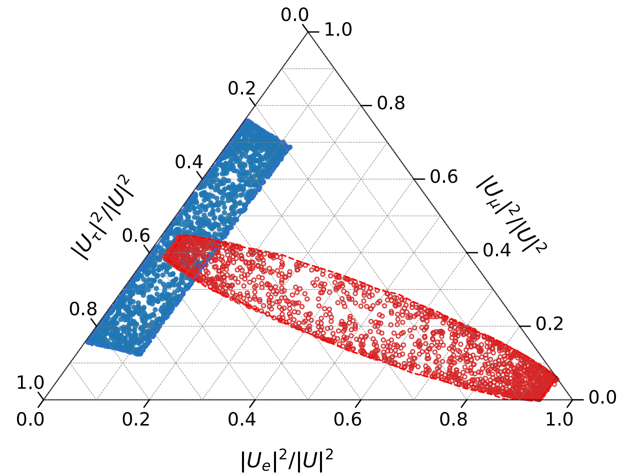
- Strong LNV
  $M \gtrsim O(1\ \mathrm{GeV})$

$$(U^2)_{\mathrm{ov}}^{\mathrm{sLNV}} \lesssim 16(2.3) \times 10^{-7} \left(\frac{1\,\mathrm{GeV}}{M}\right)^{28/13} \quad \mathrm{NH(IH)}$$

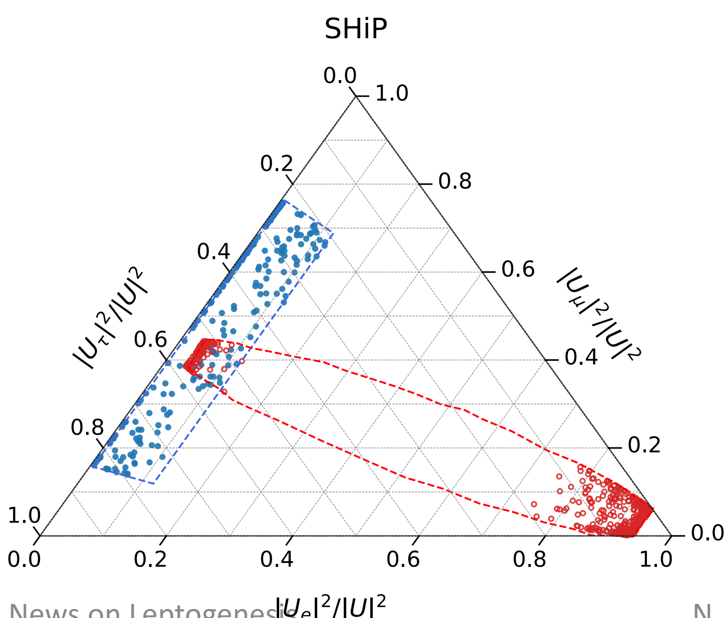
Relation to other observables

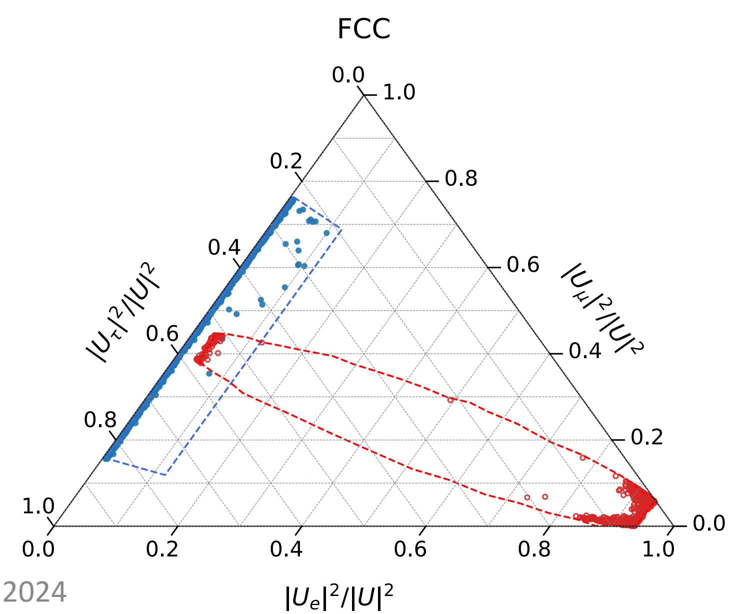
1. HNL flavour mixing

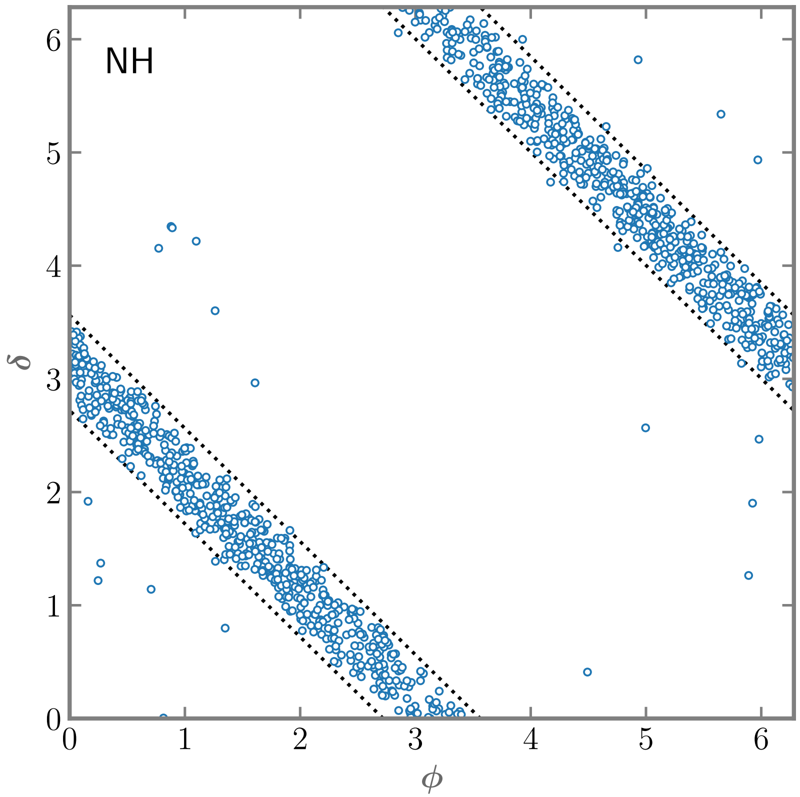
- Full scan: NH and IH

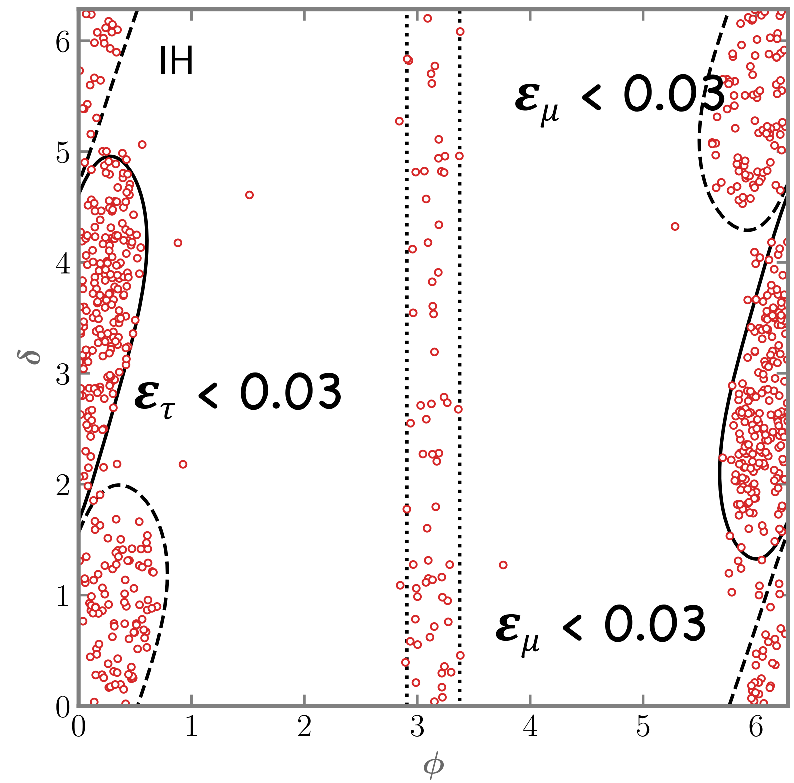


- $\Delta M/M = 10^{-2}$









$\varepsilon_e < 0.01$ ($\delta+\phi \approx \pi$, 3π)

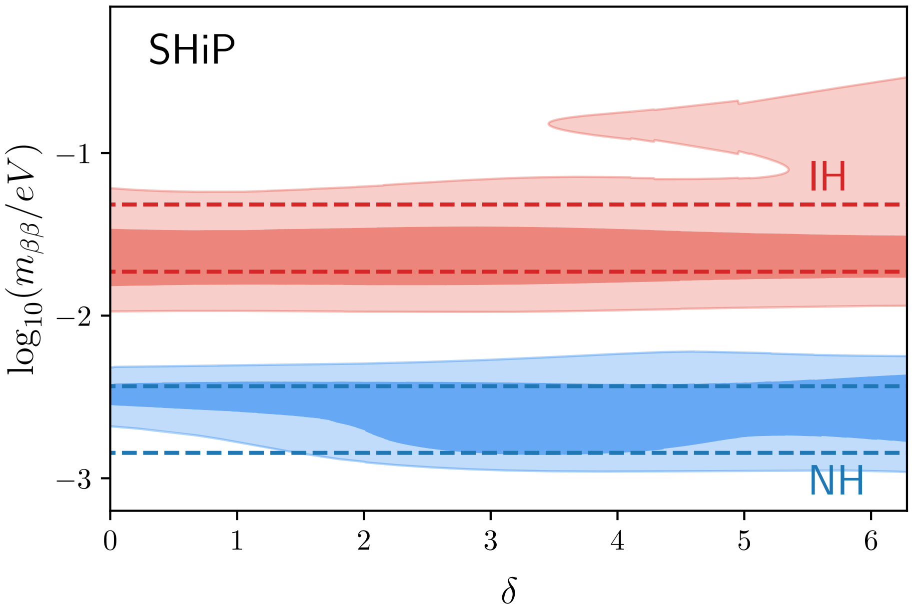
$\varepsilon_e < 0.05$

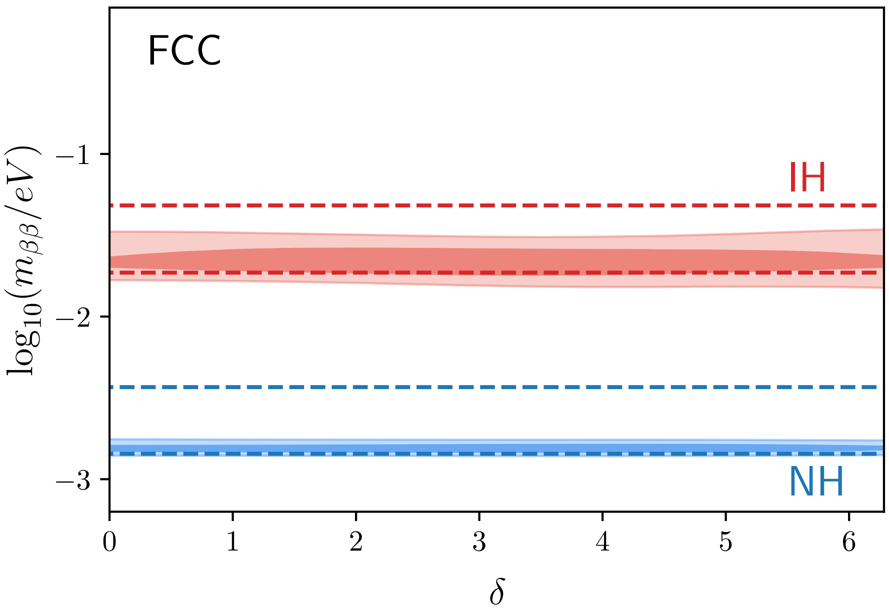
($\phi \approx \pi$)

$$\epsilon_\alpha \equiv \frac{(YY^\dagger)_{\alpha\alpha}}{\mathrm{Tr}[YY^\dagger]}$$

2. Neutrinoless double beta decay: $\Delta M/M = 10^{-2}$

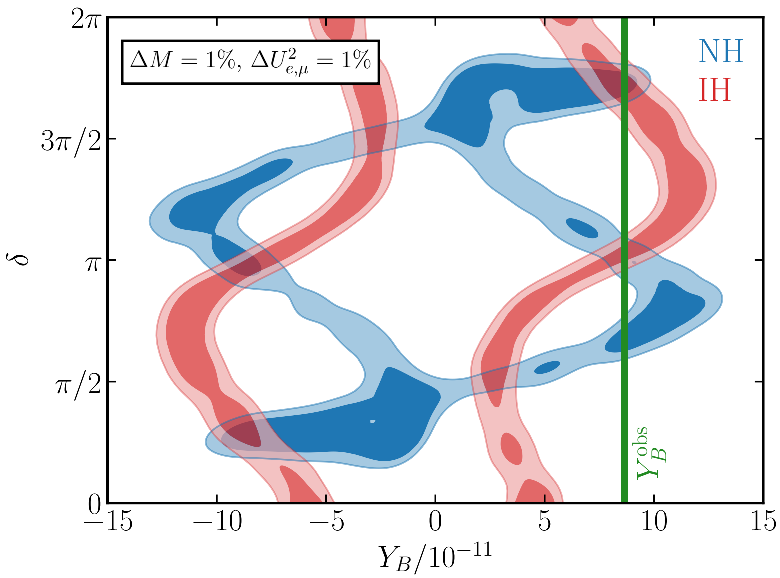
Effect of HNL only in SHIP range

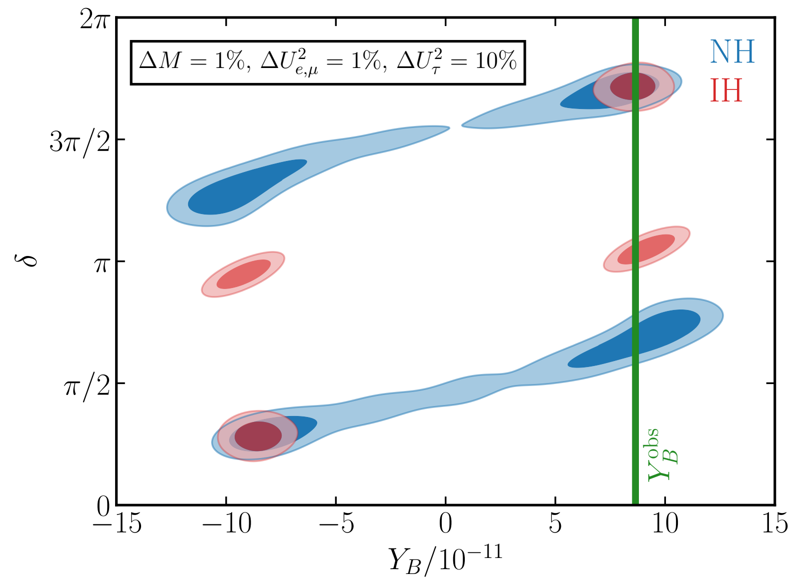


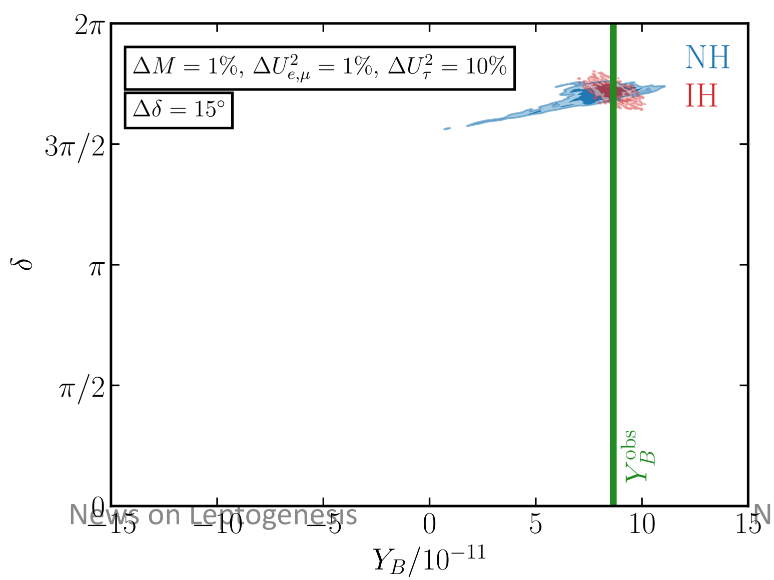


$\mu \simeq 0$ case:

Once active neutrino masses and mixings are fixed, only 4 free parameters: M, U^2 (or y^2), and PMNS phases, (δ, ϕ)







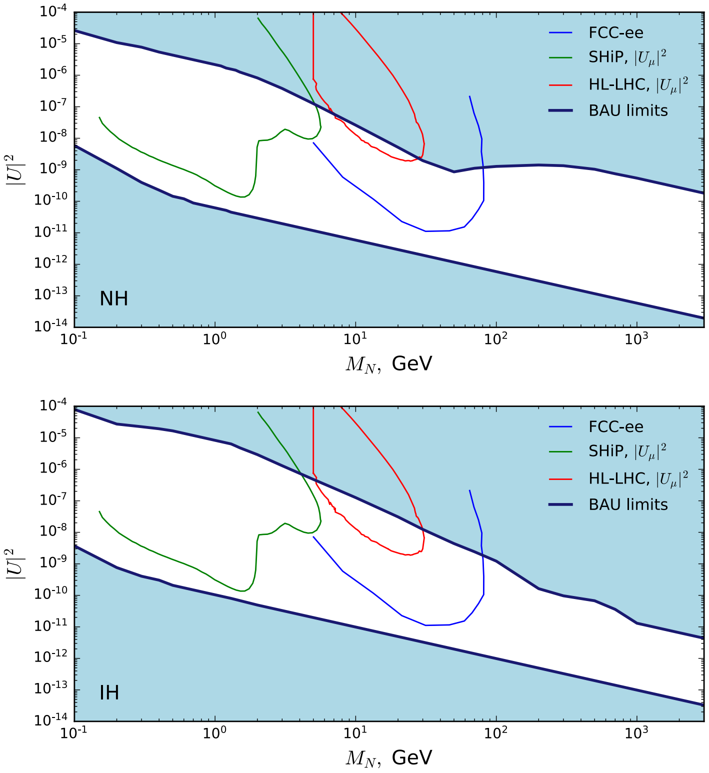
ϕ determined by Y_B

$$\Delta M \gg \Gamma$$

HNL oscillations can not be observed at FCC but LNV processes can

Extension to larger HNL masses:

unifying resonant leptogenesis and baryogenesis via neutrino oscillations



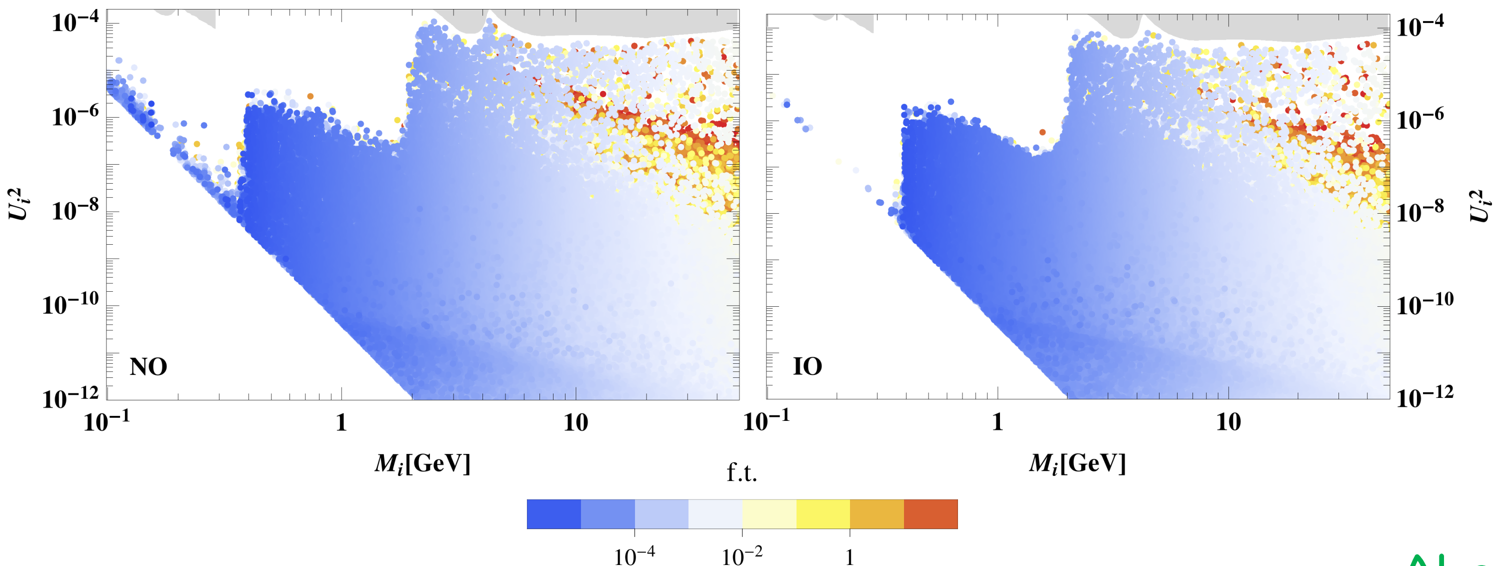
Klaric et al. 2020, 21

N=3 HNL

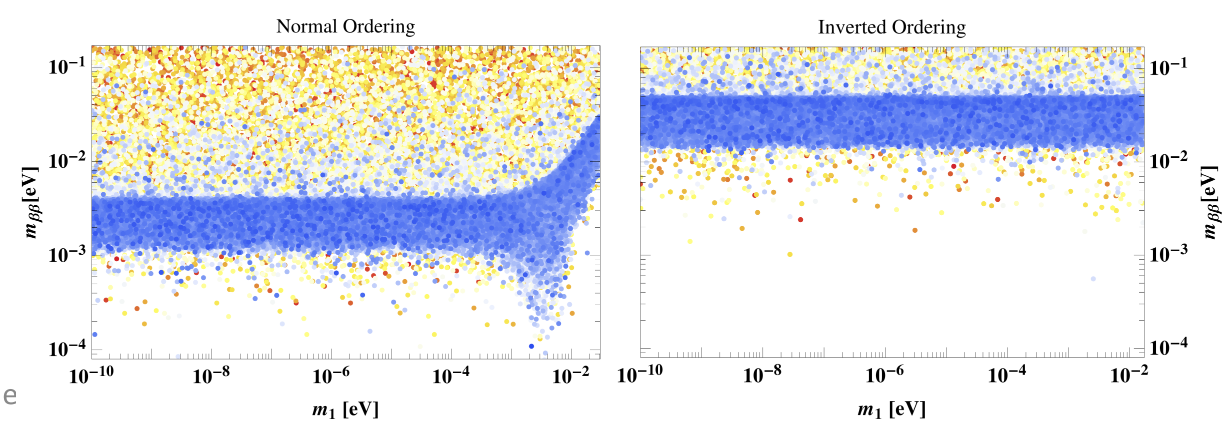
Larger values of HNL mass splittings allowed

NH

IH



Abada et al. 2019



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6. Summary and outlook

- Thermal leptogenesis in generic type I seesaw: simple, appealing … but difficult to test.

- Explore low energy alternatives:

- ν electroweak baryogenesis
- Progress understanding leptogenesis via sterile neutrino oscillations, interesting implications for low energy observables
- Unified description of resonant leptogenesis and baryogenesis via neutrino osillations under way

Backup slides

Parameter scan

Antusch et al. 2018; Abada et al. 2019; Klaric´ et al. 2020, 2021; Drewes et al. 2022

- Nested sampling algorithm UltraNest

$$\log(\mathcal{L}) = -\frac{1}{2}\left(\frac{Y_B(T_{\rm EW}) - Y_B^{\rm exp}}{\sigma_{Y_B^{\rm exp}}}\right) \qquad Y_B^{\rm exp} = (8.66 \pm 0.05) \times 10^{-11}$$

- Priors:

$\log_{10}(M_1)$	$\log_{10}(\Delta M/M_1)$	$\log_{10}(y)$	θ	δ	α
$[-1, 2]$	$[-14, -1]$	$[-8, -4]$	$[0, 2\pi]$	$[0, 2\pi]$	$[0, 2\pi]$

- y'/y < 0.1 , to ensure approximate LNC limit
- Restricted to region testable at SHIP, FCC-ee.
- Additional scan for NH with log priors in all angles, in range [-5,-2]

4. $\mu \simeq 0$ case

$$M = \begin{pmatrix} 0 & \Lambda \\ \Lambda & 0 \end{pmatrix} \qquad Y = \begin{pmatrix} y_e & y'_e e^{i\beta'_e} \\ y_\mu & y'_\mu e^{i\beta'_\mu} \\ y_\tau & y'_\tau e^{i\beta'_\tau} \end{pmatrix}$$

- Sterile neutrinos degenerate at T > T_{EW}, except for small loop correction

$$\Delta\mu \propto yy'\rho M/(4\pi)^2 \ll yy'\rho T^2/(8M)$$

- At T=0:

$$\Delta M_{NH} = |m_3| - |m_2| = \sqrt{\Delta m^2_{atm}} - \sqrt{\Delta m^2_{sol}}$$

$$\Delta M_{IH} = |m_2| - |m_1| = \sqrt{\Delta m^2_{atm}} - \sqrt{\Delta m^2_{atm} - \Delta m^2_{sol}}$$

- Once active neutrino masses and mixings are fixed, only 4 free parameters: M, U^2 (or y^2), and PMNS phases, (δ, ϕ)

CP violating flavour basis invariants

- All previous CP invariants vanish in the $\mu \simeq 0$ limit
- Higher order in the Yukawa couplings:

$$\tilde{I}_0 \equiv \mathrm{Im}\left(\mathrm{Tr}\left[Y^\dagger Y M_R^* Y^T Y^* M_R Y^\dagger Y_\ell Y_\ell^\dagger Y\right]\right) \equiv \sum_\alpha y_{\ell_\alpha}^2 \Delta_\alpha$$

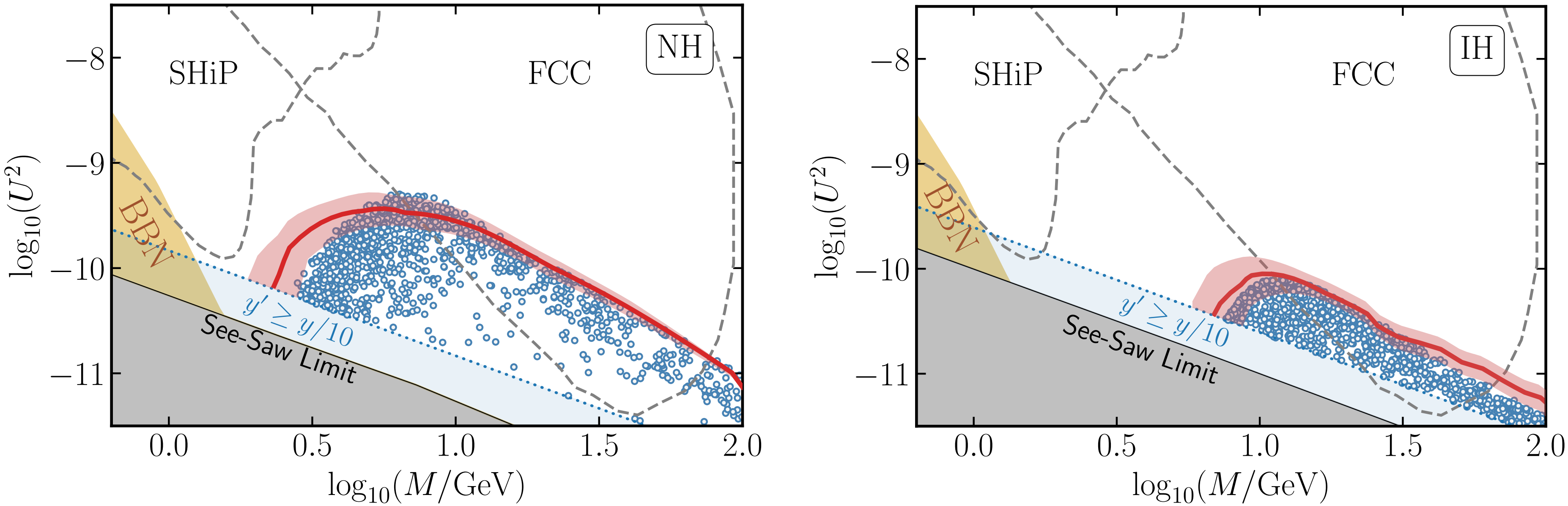
with

$$\Delta_\alpha = \mathrm{Im}\left[\left(Y Y^\dagger Y M_R^* Y^T Y^* M_R Y^\dagger\right)_{\alpha\alpha}\right] \qquad \sum_\alpha \Delta_\alpha = 0$$

We find contribution from weak flavour:

$$\Delta_\alpha^{\mathrm{fw}} = \frac{1}{\mathrm{Tr}\left(\mathrm{Y}^\dagger \mathrm{Y}\right)^2} \Delta_\alpha \,,$$

$\mu \simeq 0$ case

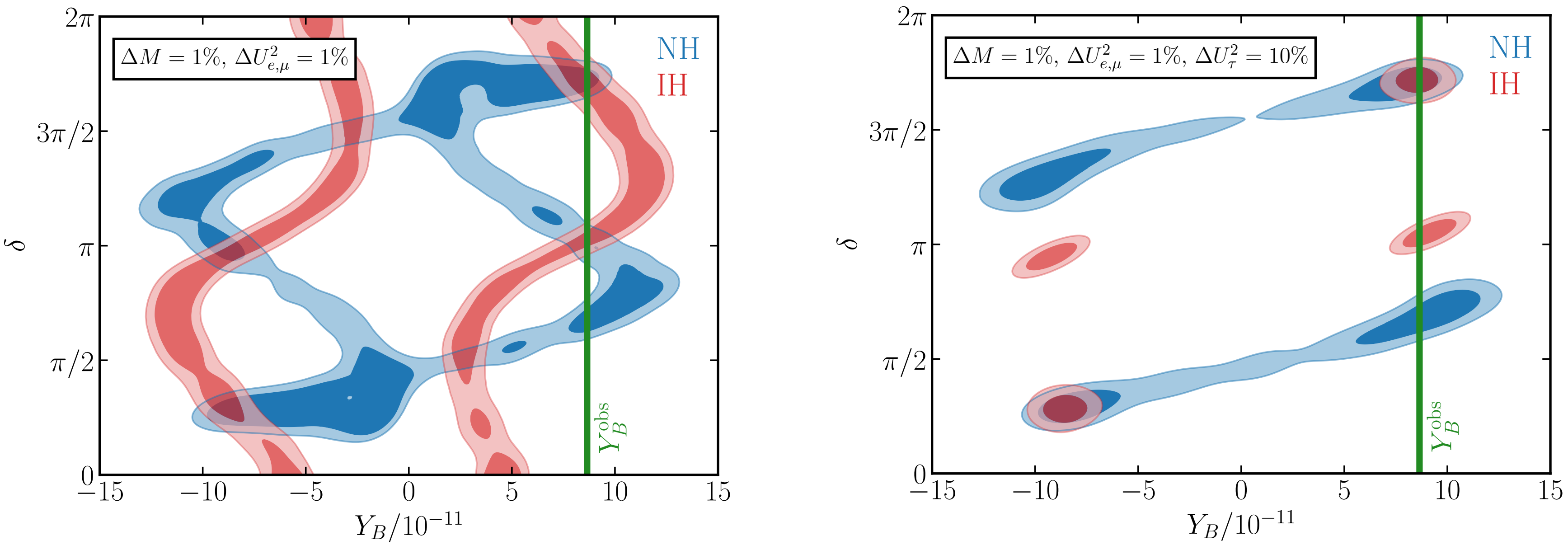


Observable at FCC only in sHC regime

$$Y_B = -1.5 \times 10^{-25} \left(\frac{\mathrm{GeV}}{M}\right)^2 \left(\frac{1}{U^2}\right)^2 f_\alpha^{\mathrm{H}} \quad U^2 \geq 1 \times 10^{-6} \left(\frac{1\,\mathrm{GeV}}{M}\right)^4$$

$$f_e^{\mathrm{NH}} = -\frac{\sqrt{r}}{2}\theta_{13} s_{12} \sin(\delta + \phi)\,, \qquad f_\mu^{\mathrm{IH}} = f_\tau^{\mathrm{IH}} = -\frac{r^2}{8} c_{12} s_{12} \sin\phi$$

Numerical likelihood inference in the case of measuring HNL-active neutrino mixings



	$M^{\rm true}/{\rm GeV}$	$(U_e^2)_{\rm true}$	$(U_\mu^2)_{\rm true}$	$(U_\tau^2)_{\rm true}$	$\delta^{\rm true}/{\rm rad}$
NH	31.60	2.843×10^{-12}	1.087×10^{-11}	1.234×10^{-11}	5.396
IH	20.731	3.291×10^{-11}	4.823×10^{-12}	3.465×10^{-12}	5.402