Constraining the Cosmic Neutrino Background with NGC 1068

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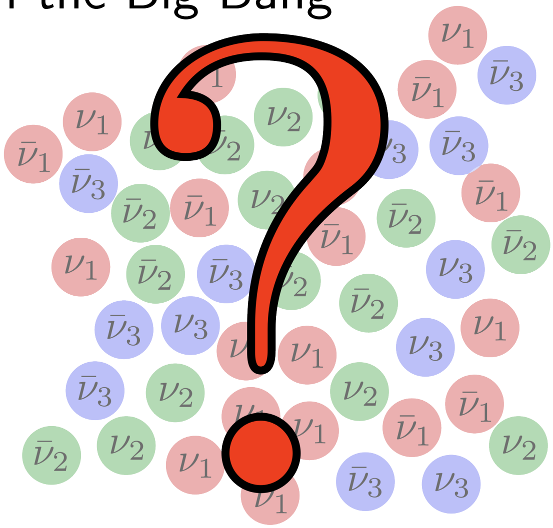
The Cosmic Neutrino Background

- The universe is filled with a sea of neutrinos
- Neutrinos decouple in the early universe
- ΛCDM: ~300 neutrinos per cm^3 left over from the Big Bang
- What we could learn about:
  - Early Universe Physics
  - BSM Neutrino Physics



Relic Neutrino Overabundance

What are the experimental bounds on the CνB?

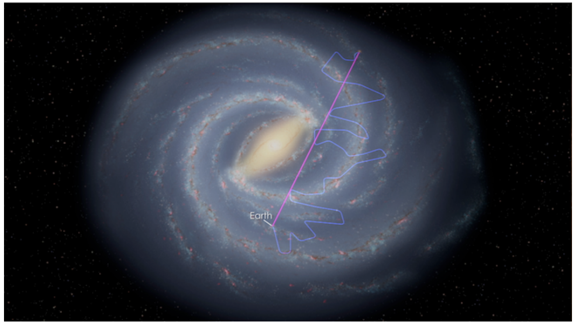
$$\eta = \frac{n}{(56\,\mathrm{cm}^{-3})}$$

- KATRIN Experiment: $\eta < 1.94 \times 10^{11}$

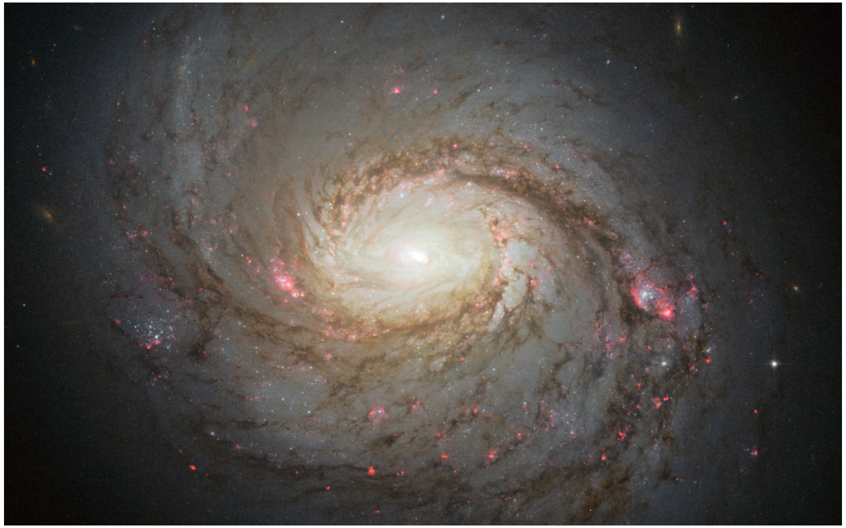
  KATRIN Collaboration, 10.1103/PhysRevLett.129.011806

- Cosmic Rays: $\eta \lesssim 10^{11}$

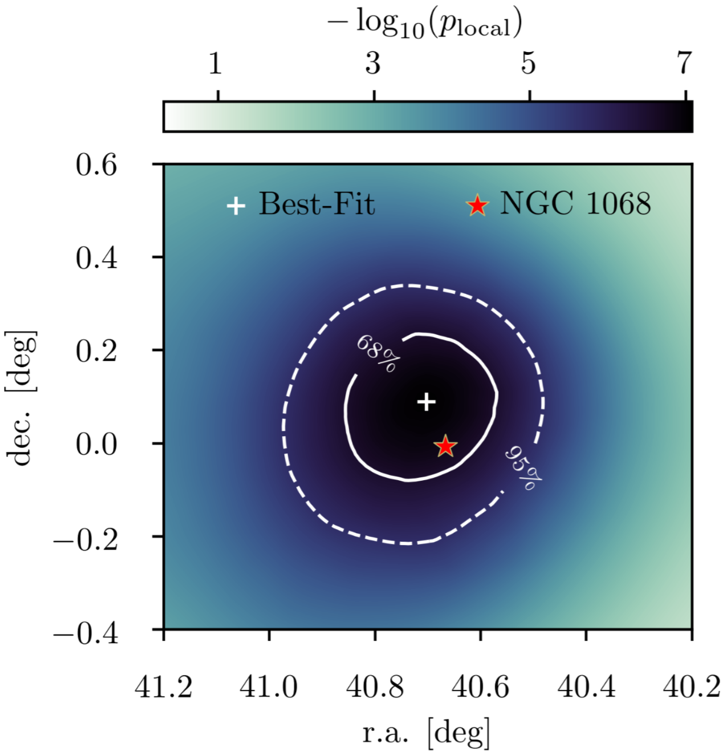
  Mar Císcar-Monsalvatje et. al., 2402.00985



NGC 1068



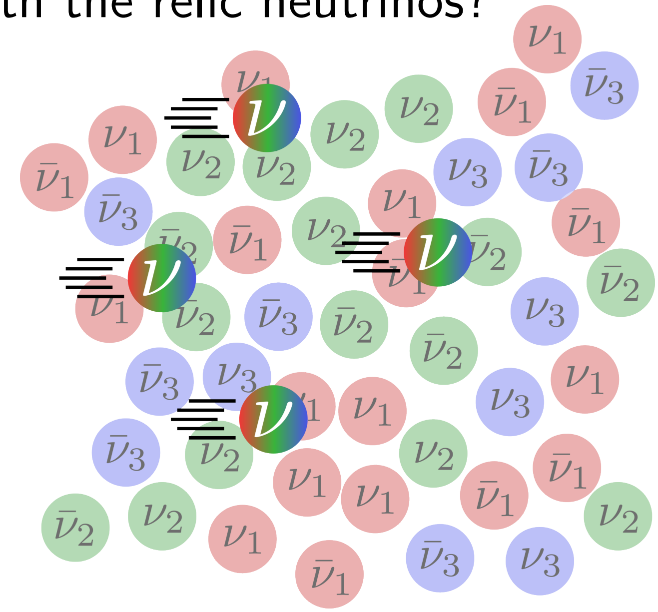
- Galaxy with an active galactic nuclei (AGN)
- Around 14 Mpc from the Milky Way
- Most significant point-source at IceCube



IceCube Collaboration 10.1126/science.abg3395

The Cosmic Neutrino Background

- Neutrinos from NGC 1068 are travelling through the CνB
- What if they interact with the relic neutrinos?



Transport Equation

Need to solve the transport equation for the flux:

$$\frac{\partial \Phi_i(r,E)}{\partial r} = -\Phi_i(r,E)\sum_j n_j \sigma_{ij}(E) + \sum_{j,k,l} n_k \int_E^\infty dE' \Phi_j(r,E') \frac{d\sigma_{jk\to il}}{dE}(E',E)$$

Transport Equation

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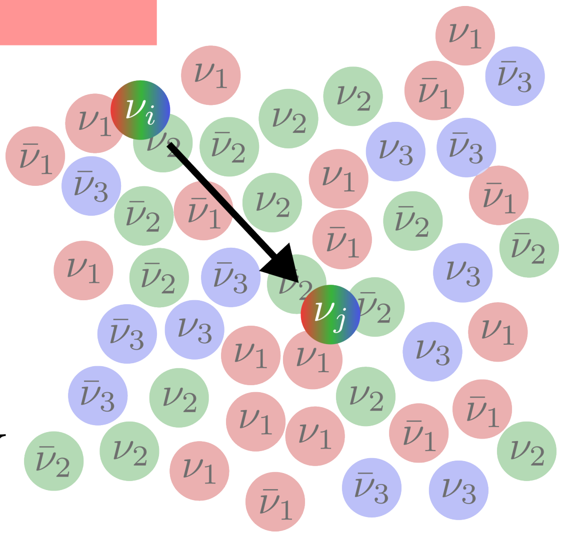
$$\frac{\partial \Phi_i(r,E)}{\partial r} = -\Phi_i(r,E)\sum_j n_j \sigma_{ij}(E)$$

Φ : Flux
n : Num. Density
σ : SM Cross-Section

Depletion Term

- $\nu\nu \to \nu\nu$
- $\bar{\nu}\nu \to \bar{\nu}\nu$
- $\bar{\nu}\nu \to e^+e^-$

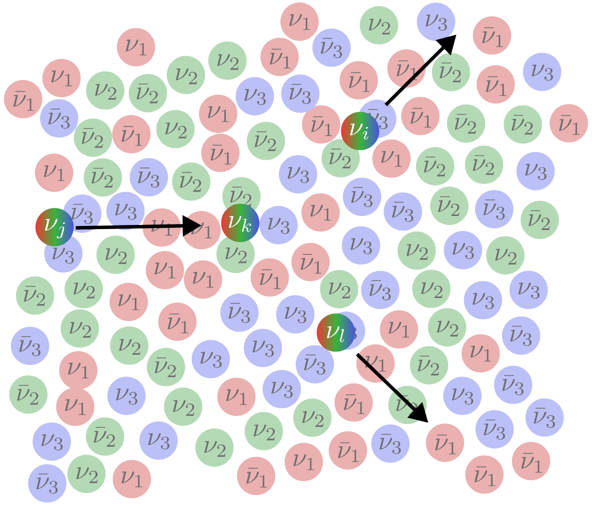
$$\sqrt{s} = \sqrt{2m_j E} \approx \text{keV} - \text{MeV}$$



Transport Equation

Need to solve the transport equation for the flux:

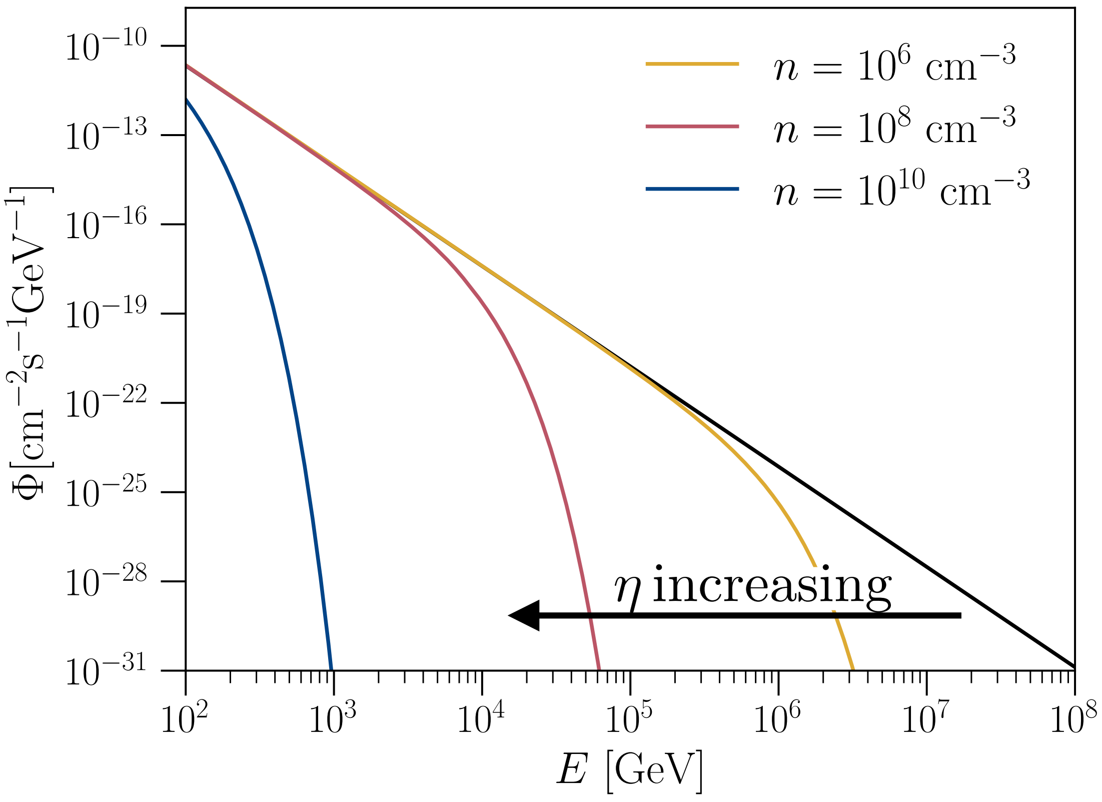
$$\frac{\partial \Phi_i(r,E)}{\partial r} = + \sum_{j,k,l} n_k \int_E^{\infty} dE' \Phi_j(r,E') \frac{d\sigma_{jk\to il}}{dE}(E',E)$$



Regeneration Term = Upscattering + Downscattering

$jk \to il$

Fluxes at Earth



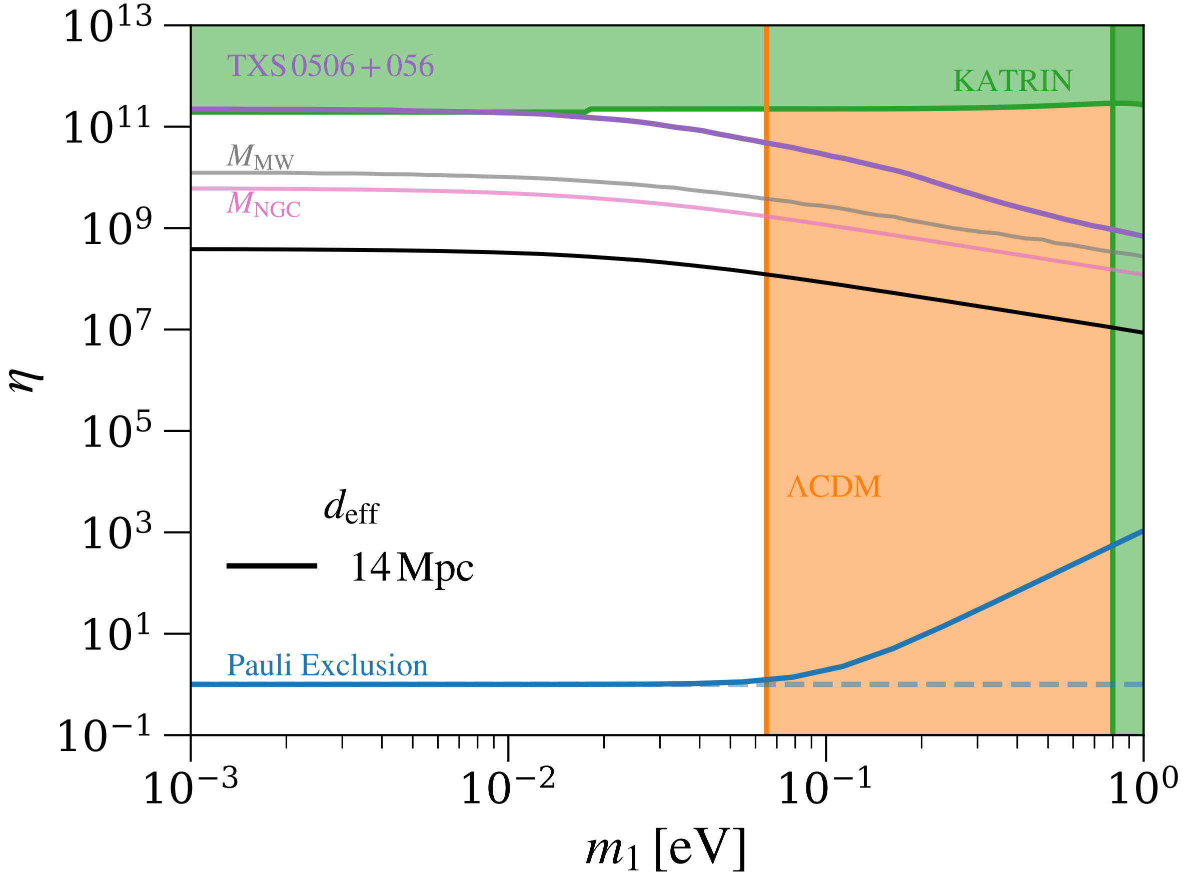
Analysis

Log-likelihood analysis using IceCube public datasets:

$$TS = -2\Delta \log \mathcal{L} = -2 \log \left(\frac{\mathcal{L}(\gamma, \eta, n_s | \mathbf{x}_i, N)}{\mathcal{L}_0} \right)$$

Take power law flux as null hypothesis

Results



95% C.L

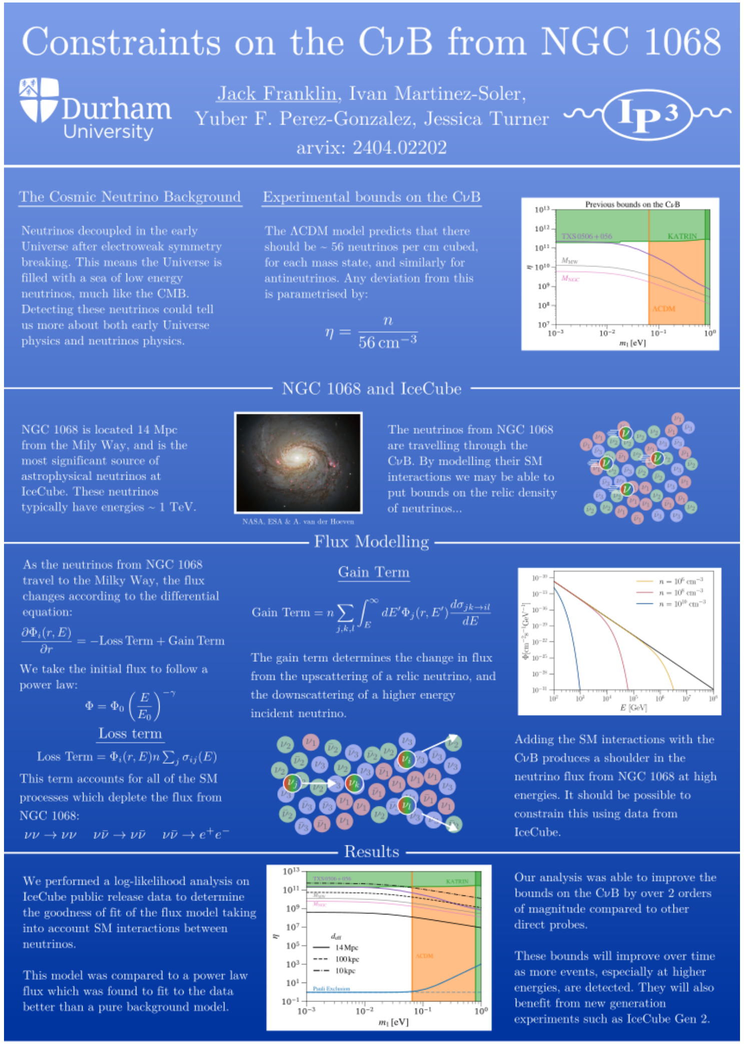
- CνB Overdensity:

$$\eta < 3.9 \times 10^8$$

Conclusion

- Direct observation constraints improved by over 2 orders of magnitude!
- A lot of constraining power still available right now with IceCube's improved analysis techniques
- Future improvements from:
  - More events
  - Higher energy neutrinos
- Extension to this work could also constrain neutrino NSIs

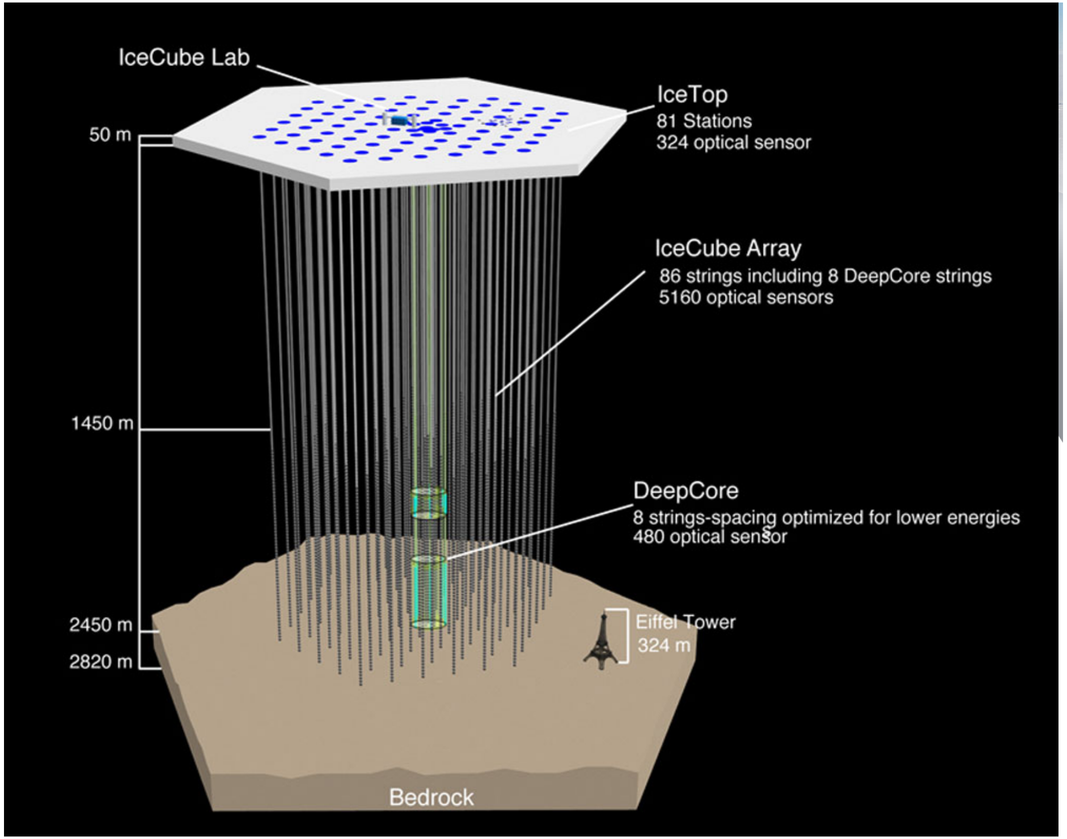
If you have any more questions,
please come see me at my poster!



Backup Slides...

The IceCube Experiment

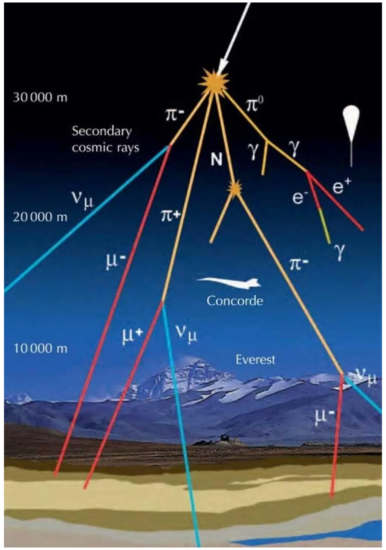
- Neutrino Observatory in Antarctica
- Uses ice as a medium for detecting neutrinos
- Consists of 86 “strings” of light-detecting modules



Neutrino Sources at IceCube

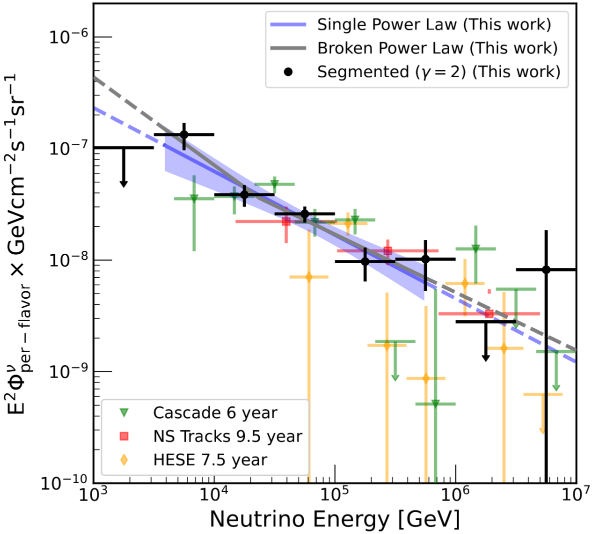
Where do the neutrinos that IceCube observes come from?

Atmospheric Neutrinos

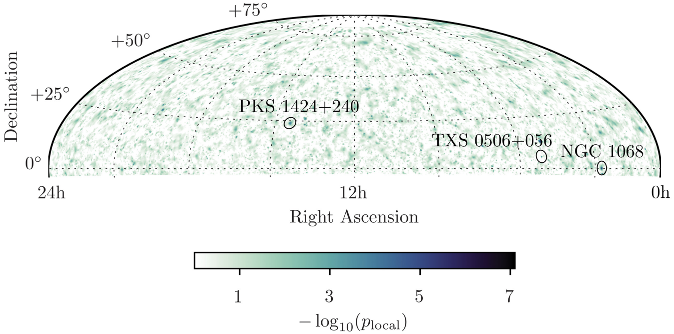


10 MeV ~ PeV

Diffuse Astrophysical Neutrinos



Point-source Neutrinos



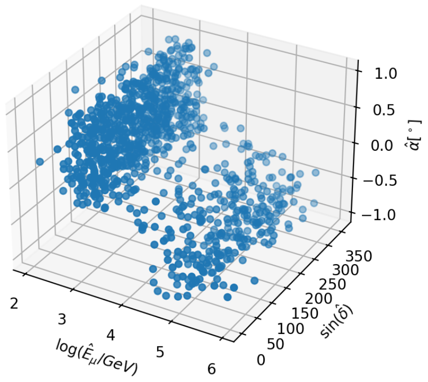
100 GeV ~ PeV

Finding Point Sources

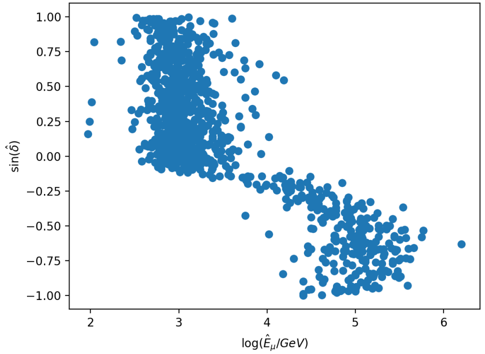
Probability Density Functions

- Background events have no dependence on right ascension
- There are ~100,000 events, of which <100 are signal
- The background pdf ~ pdf of **all** events

$$f_B(\hat{E}_{\mu,i}, \hat{\mathbf{d}}_\mathbf{i}, \hat{\sigma}_i) = \frac{1}{2\pi} f_B(\hat{E}_{\mu,i}, \sin\hat{\delta}_i)$$



"Flatten"



The Cosmic Neutrino Background

Could they interact?

Mean free path: $\lambda = \frac{1}{n\sigma}$, $\sigma \approx G_F^2 s = 2G_F^2 E_\nu m_\nu$

$$\frac{L}{\lambda} \approx 1.5 \times 10^{-8} \left(\frac{L}{14.4\,\mathrm{Mpc}} \right) \left(\frac{n}{56\,\mathrm{cm}^{-3}} \right) \left(\frac{E_\nu}{1\,\mathrm{TeV}} \right) \left(\frac{m_\nu}{1\,\mathrm{meV}} \right)$$

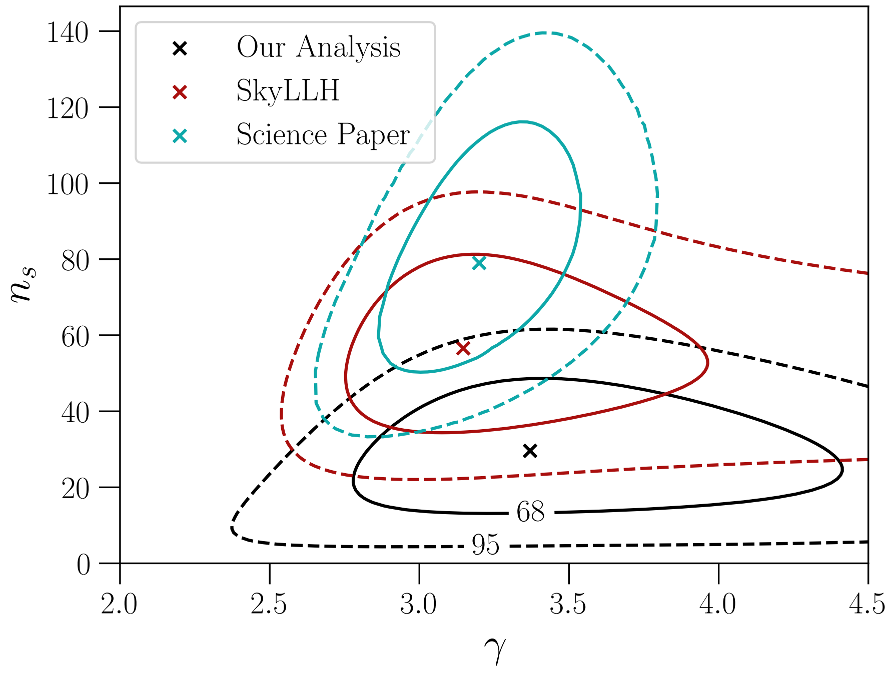
Point Source Analysis Results

Science Paper:

- New data
- Better energy reconstruction
- More accurate pdfs

SkyLLH:

- Includes data pre IC86II

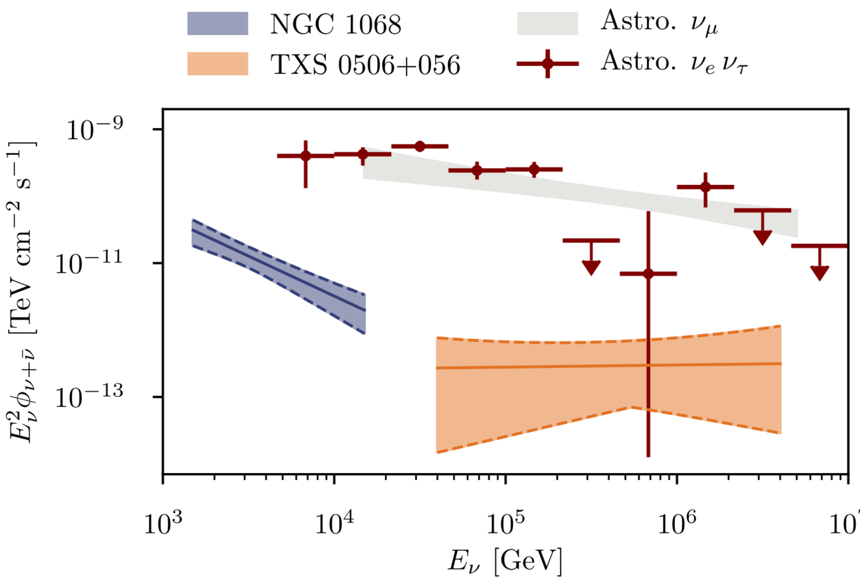


Initial Flux

- Parametrise Initial Flux with a Power Law (PL):

$$\Phi = \Phi_0 \left(\frac{E}{E_0} \right)^{-\gamma}$$

Φ_0 : Normalisation at E_0
E_0 : Reference energy (1 TeV)
γ : Spectral index



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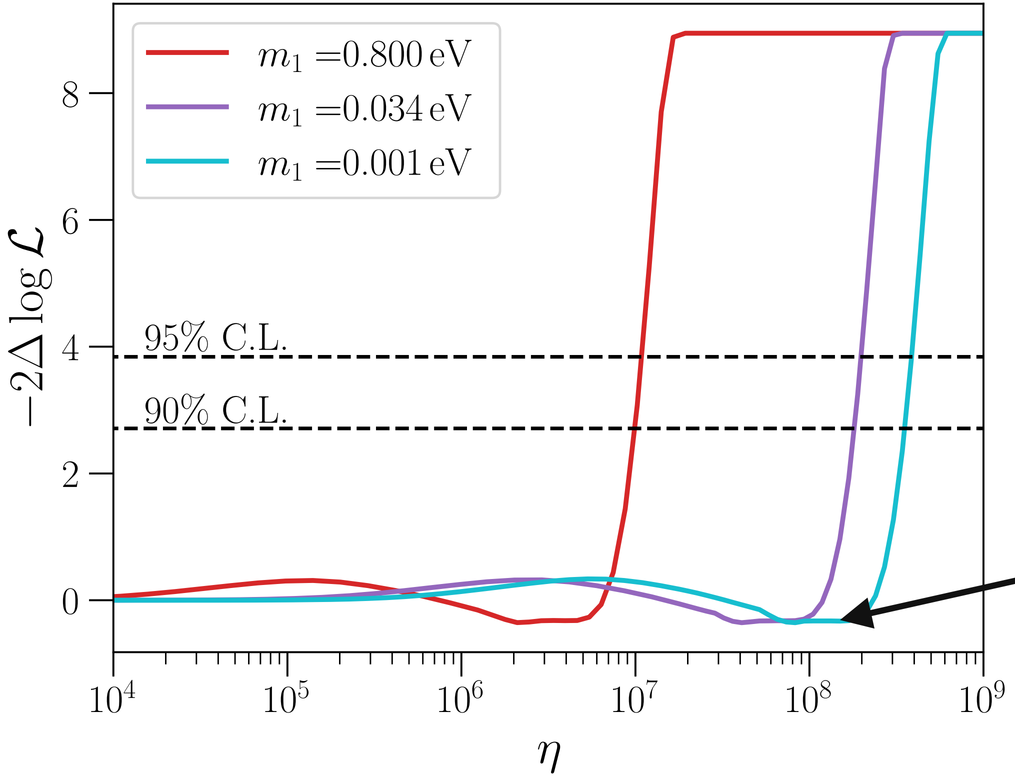
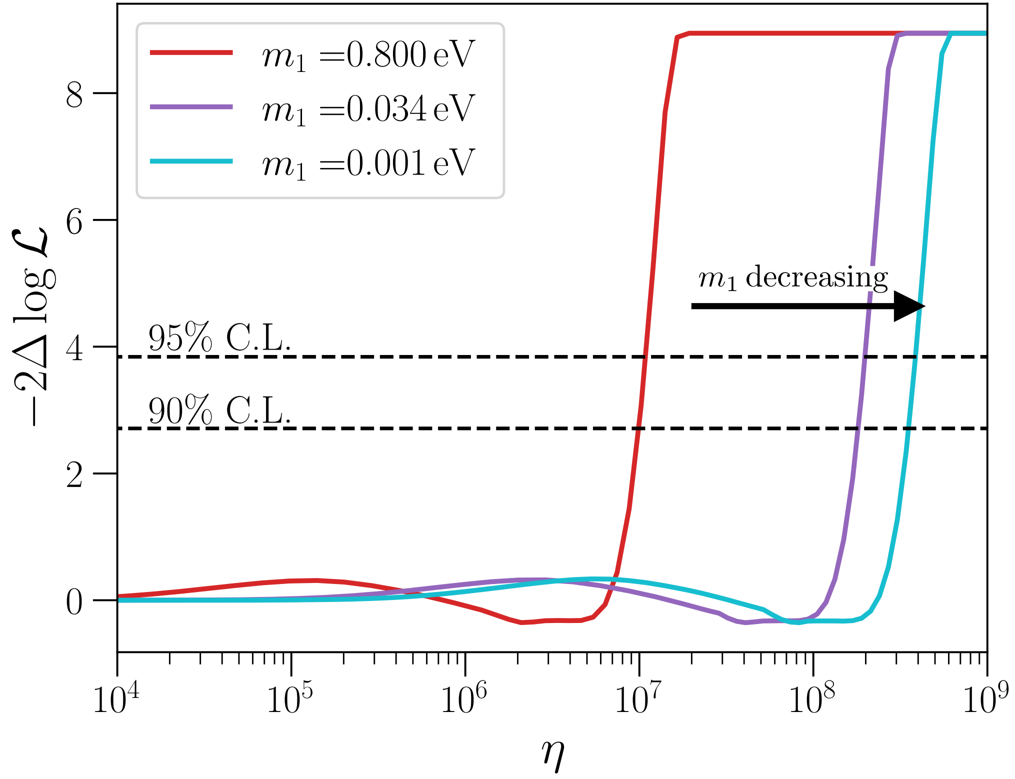
Results



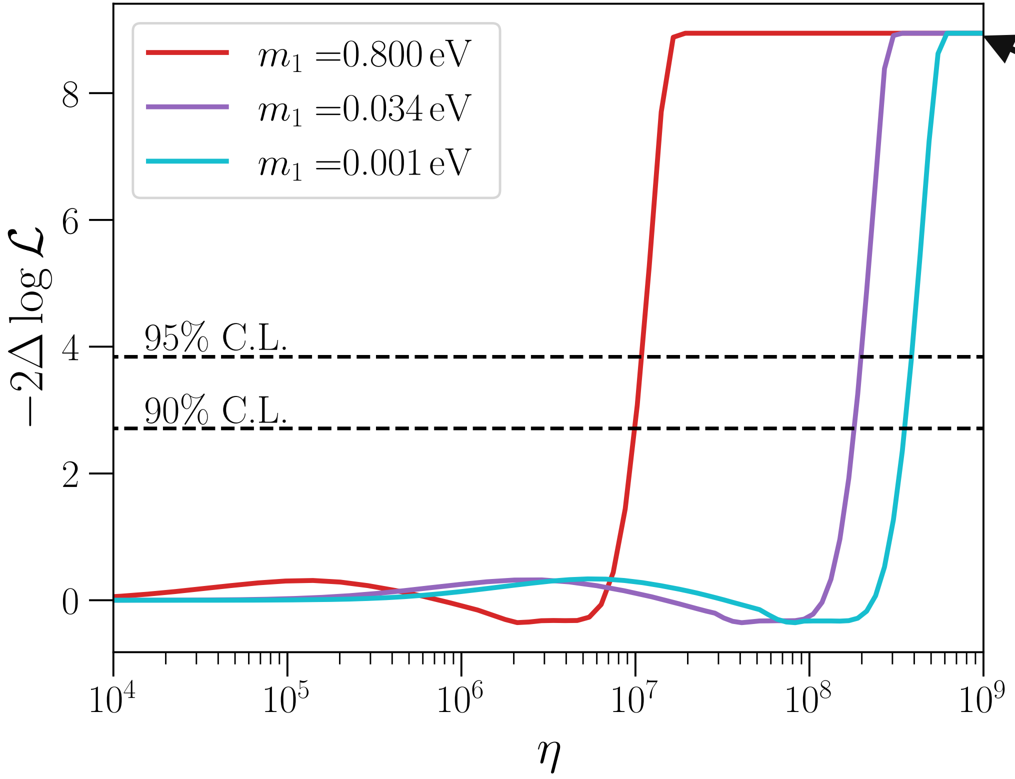
Slight preference due to lack of higher energy events. Well below 1 sigma.

Results



m_1 =0.800 eV
m_1 =0.034 eV
m_1 =0.001 eV
m_1 decreasing
95% C.L.
90% C.L.
$-2\Delta \log \mathcal{L}$
8
6
4
2
0
10^4
10^5
10^6
10^7
10^8
10^9
η

Results



Limited by statistics and methods in public data