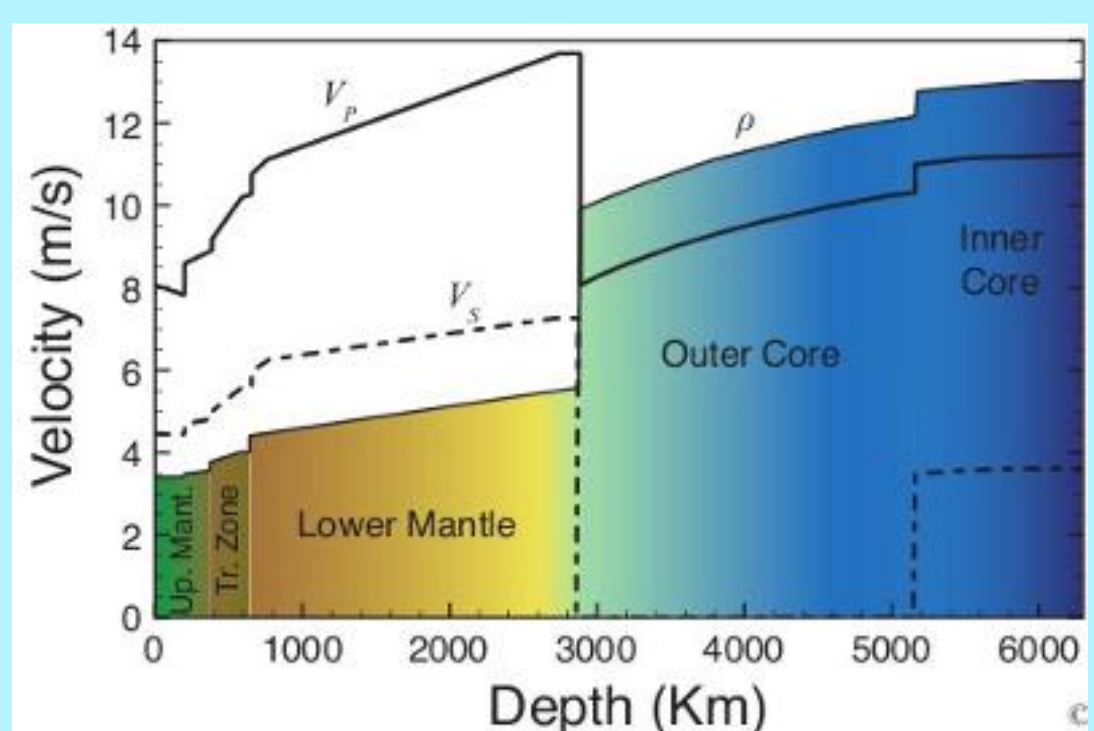


# Neutrino tomography of the Earth's lower mantle: first study with a full 3D model

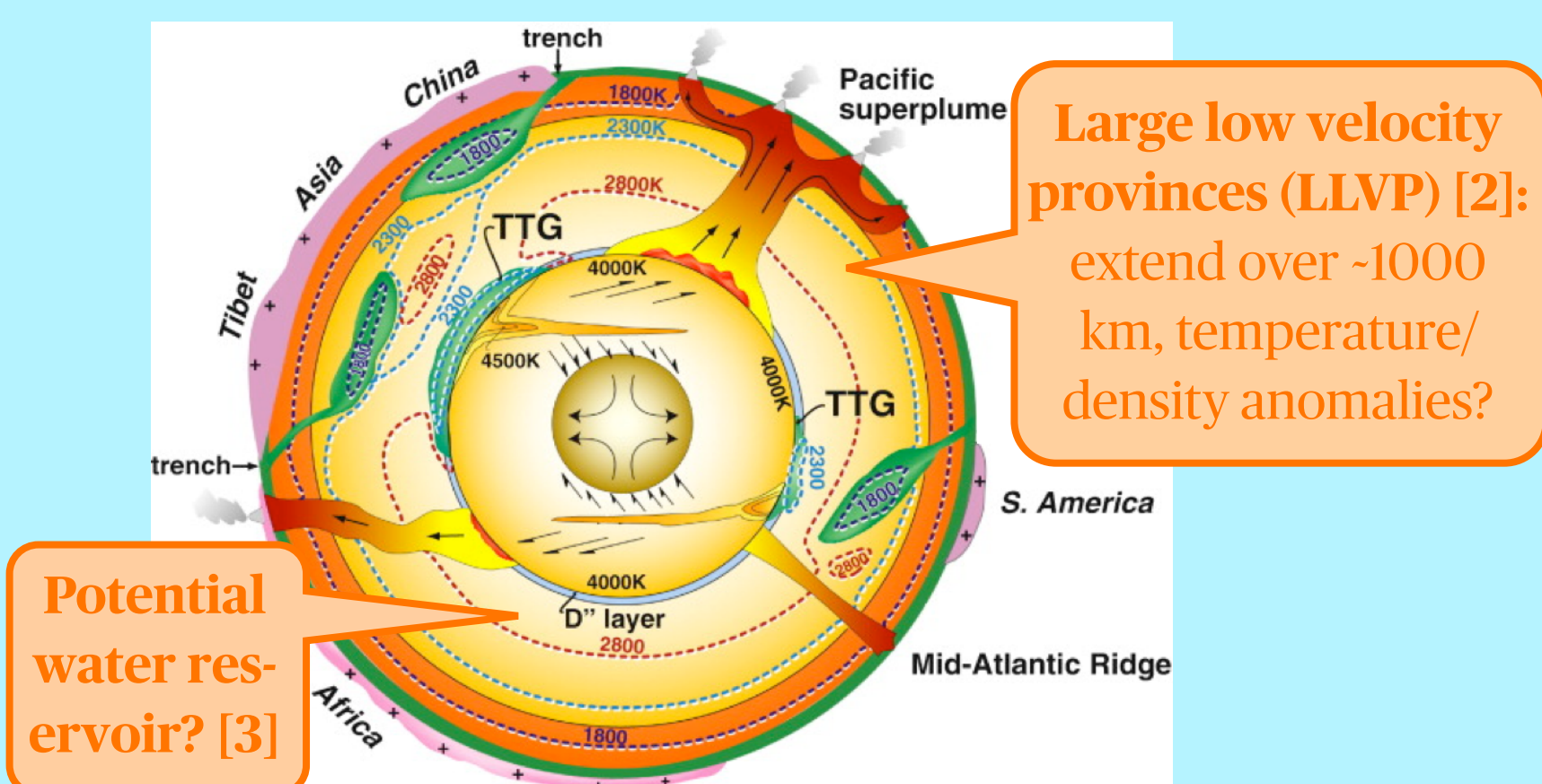
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## 1 3D structures in the Earth's mantle

Preliminary Reference Earth Model (PREM) [1]: average 1D matter density profile, known at -few % level



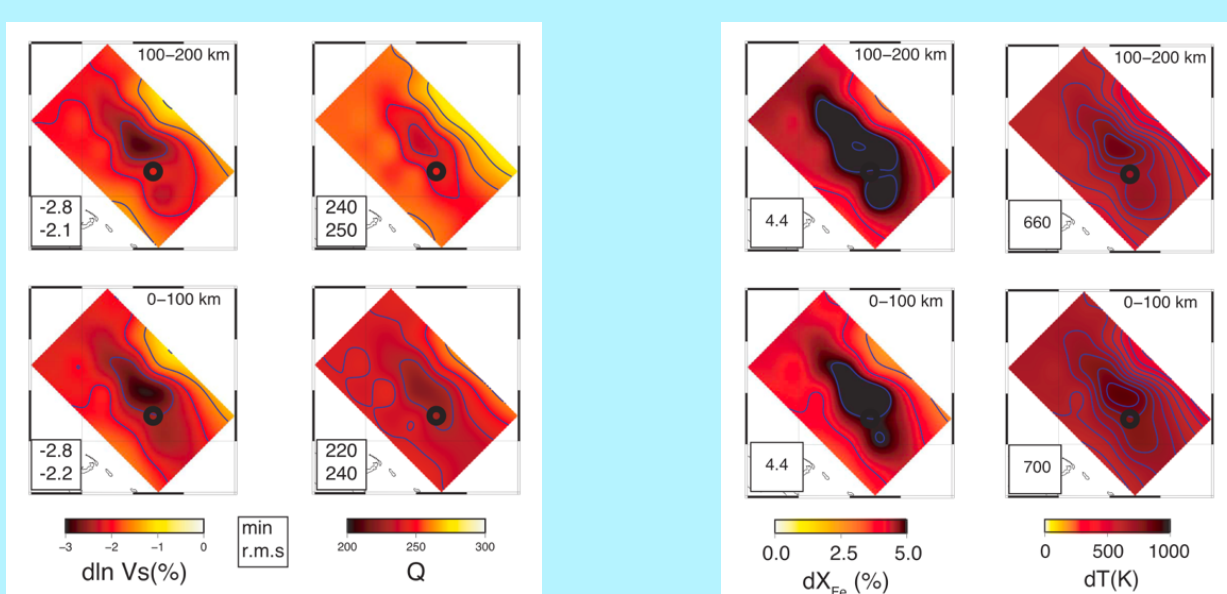
... but the inner Earth is not radially symmetric!



Potential water reservoir? [3]

Large low velocity provinces (LLVP) [2]: extend over -1000 km, temperature/density anomalies?

Possible approach to study the inner Earth: Imaging of the edge of the Hawaiian LLVP [4]



seismic data\* → interpretation in terms of Fe content and temperature

\*1D/3D imaging techniques in seismology:

- P-wave and S-wave velocities from seismic travel times
- Density from long-period normal mode analysis
- Attenuation (Q) and anisotropy study using waveform amplitudes

Disentangling thermal from compositional origin of inhomogeneities requires a multi-parameter analysis → can the electron density be a new observable?

## 2 Neutrino oscillation tomography

Atmospheric neutrinos:

- isotropic beam of known composition of  $\nu_e$  and  $\nu_\mu$  that cross the Earth [6]
- Wide range of energies: GeV-TeV
- Wide range of oscillation path lengths through the Earth: 50 km - 12.800 km

$\nu_e$  (not  $\nu_\mu, \nu_\tau$ ) interact with electrons in matter:

- Extra potential proportional to the electron density  $N_e$
- Matter effects modify flavour oscillations along the neutrino path

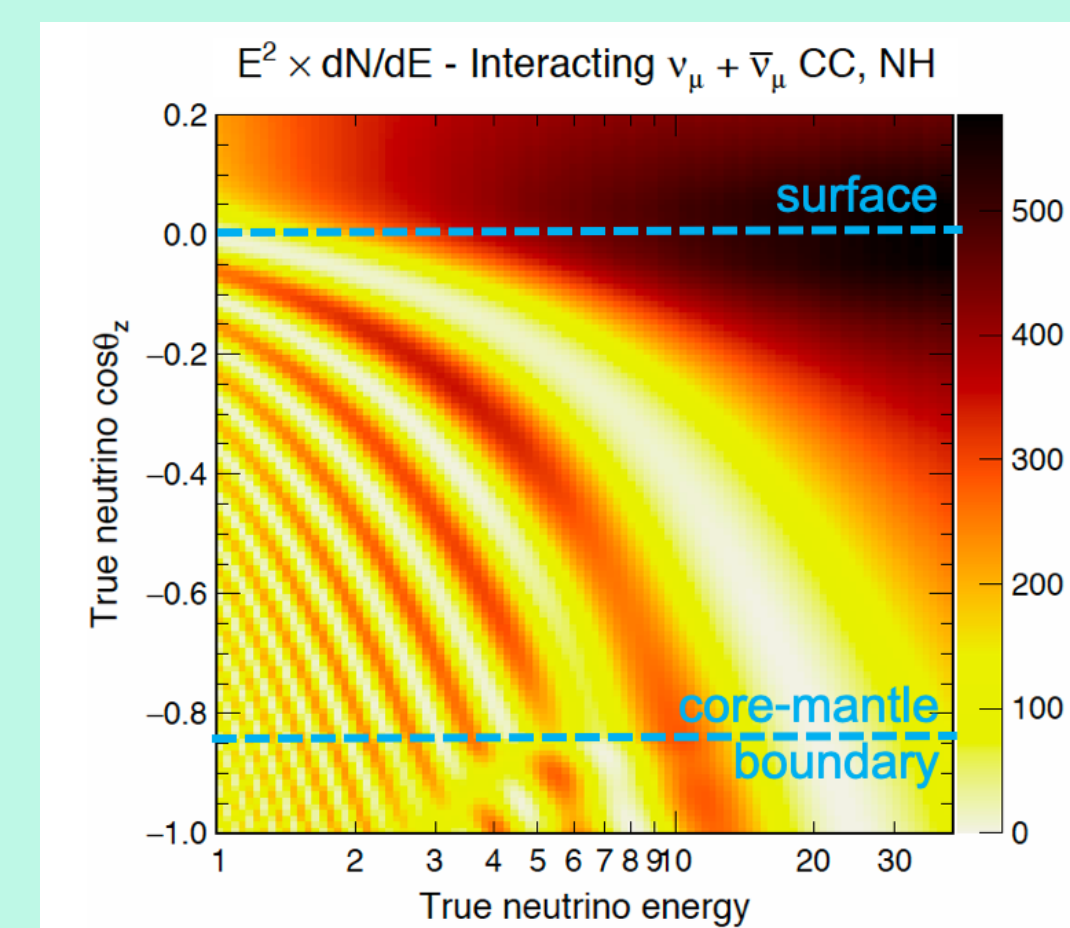


Figure taken from [5]

Maximal effect at resonance energy:

$$E_{\text{res}} = \frac{\Delta m_{31}^2 \cos 2\theta_{13}}{2\sqrt{2}G_F N_e} \approx 7 \text{ GeV} \left( \frac{4.5 \text{ g/cm}^3}{\rho_{\text{matter}}} \right) \left( \frac{\Delta m_{31}^2}{2.4 \cdot 10^{-3} \text{ eV}^2} \right) \cos 2\theta_{13} \approx 7 \text{ GeV}$$

Differential rate of interacting (anti)neutrinos of flavour  $\beta$ :

$$\frac{dN_{\beta}^{\text{int}}(E, \theta)}{dE d\theta} = \underbrace{\sigma_{\nu_{\beta}}(E)}_{\text{Cross section}} \cdot \underbrace{\sum_{\nu_{\alpha}} P_{\nu_{\alpha} \rightarrow \nu_{\beta}}(E, \theta)}_{\text{Oscillation probabilities}} \cdot \underbrace{\frac{d\Phi_{\nu_{\alpha}}(E, \theta)}{dE d\theta}}_{\text{Differential flux of atmospheric neutrinos}}$$

For mantle density

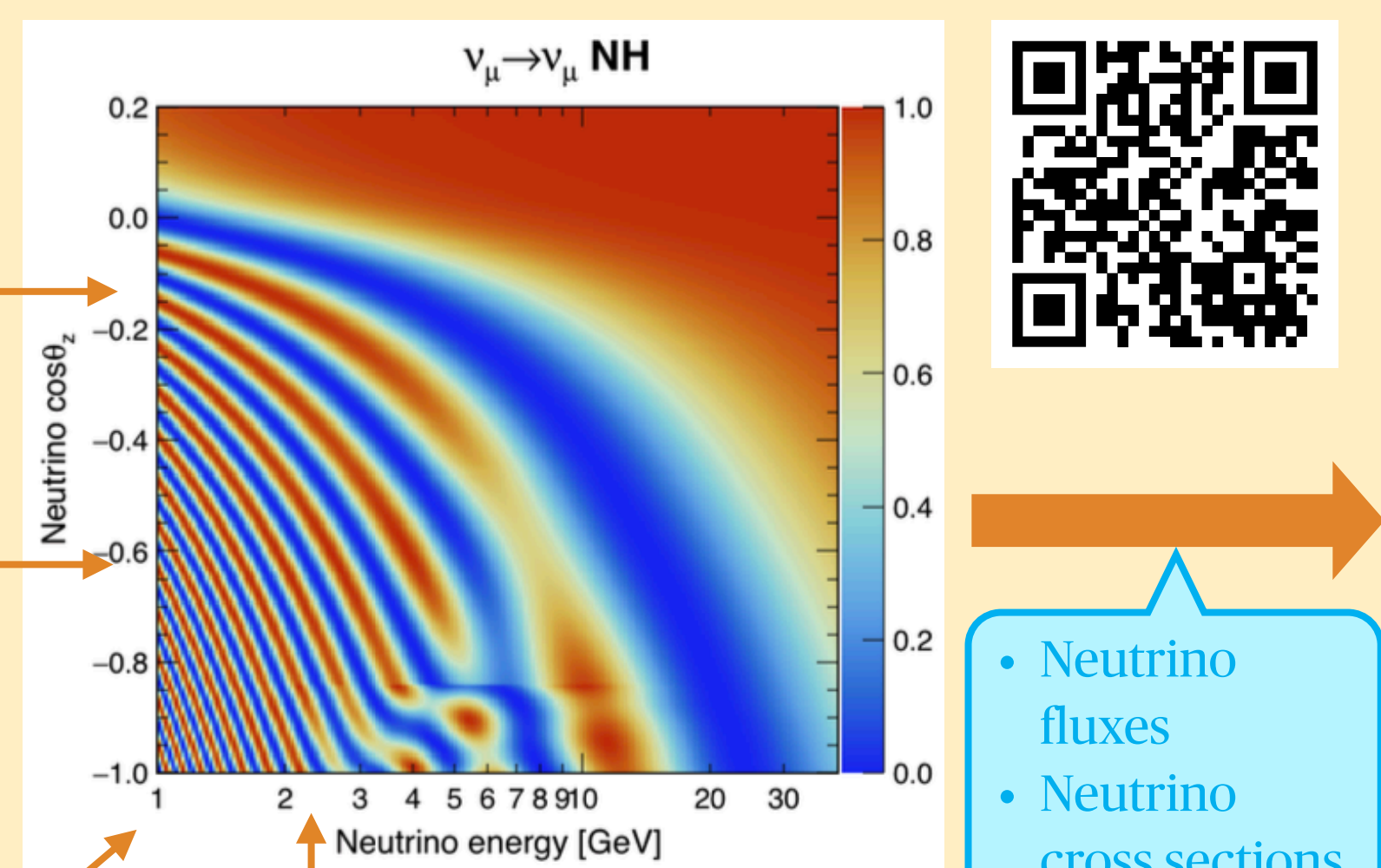
## 3 Simulations with EarthProbe

OscProb module

Neutrino energy and zenith angle sampling points

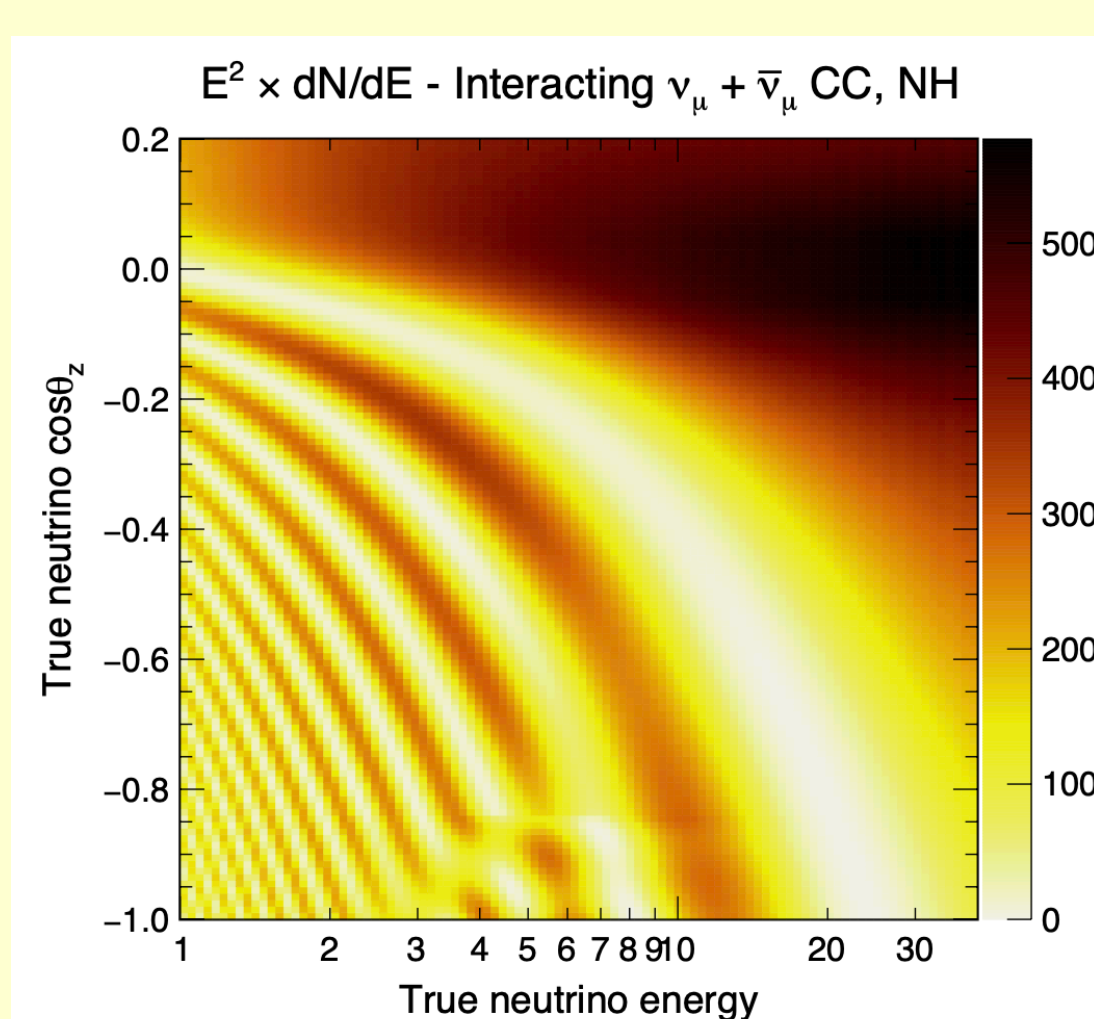
Neutrino oscillation parameters → Standard 3-flavours

Following the neutrino baselines - with layer information, followed in 3D:



Earth model ID PREM → 3D PREM with LLVPs  
 Detector position → latitude, longitude, depth

Interacting events



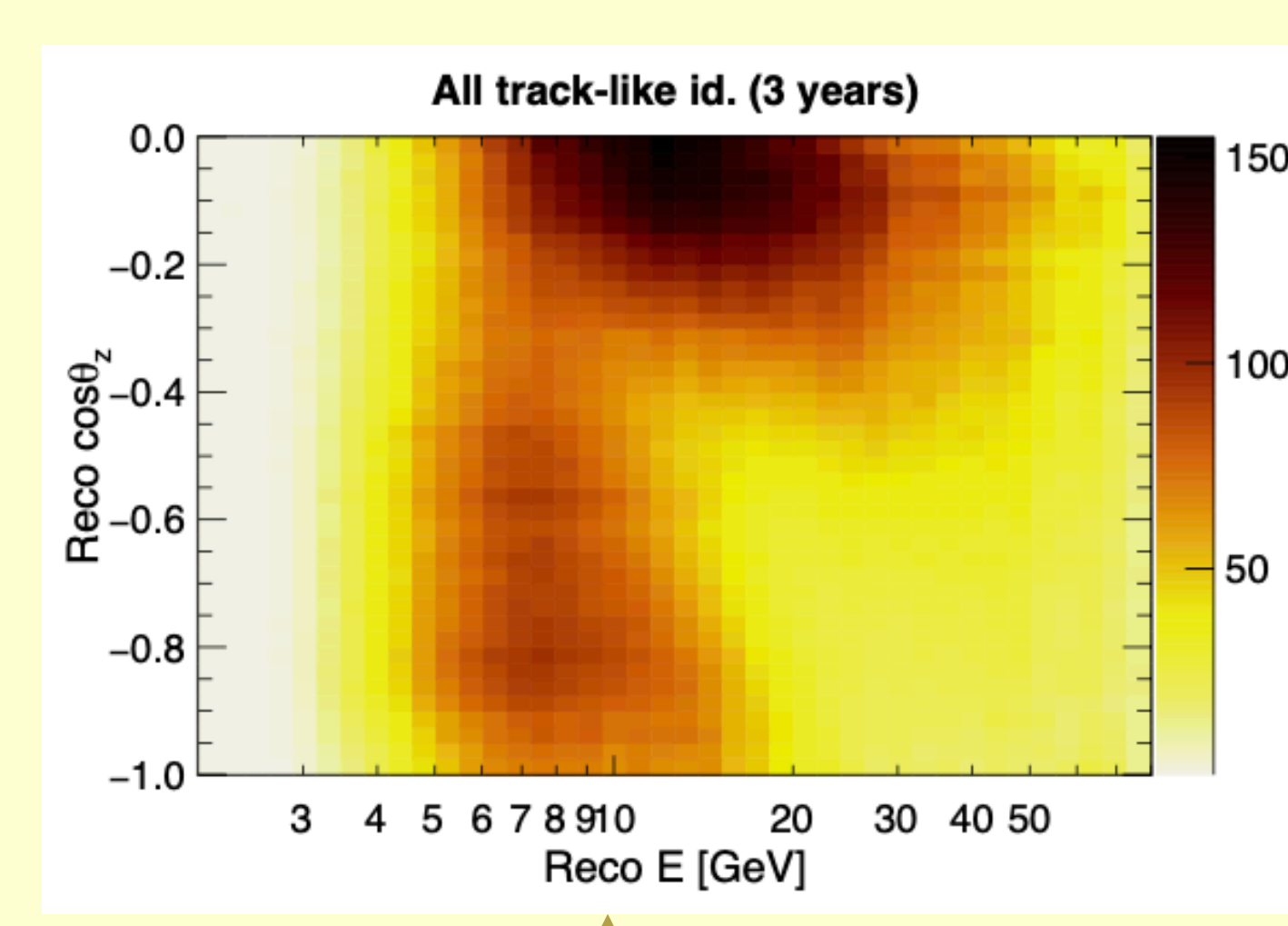
Exposure of 1 Mt · yr for a detector in sea-water

Parameterised detector response:  $E, \theta, \phi$  resolutions

Effective mass Particle identification → track vs. Shower

Binning for reconstructed neutrino energy and angle of incidence

Reconstructed events



Exposure of - 18 Mt · yr at KM3NeT/ORCA

## 4 3D implementation

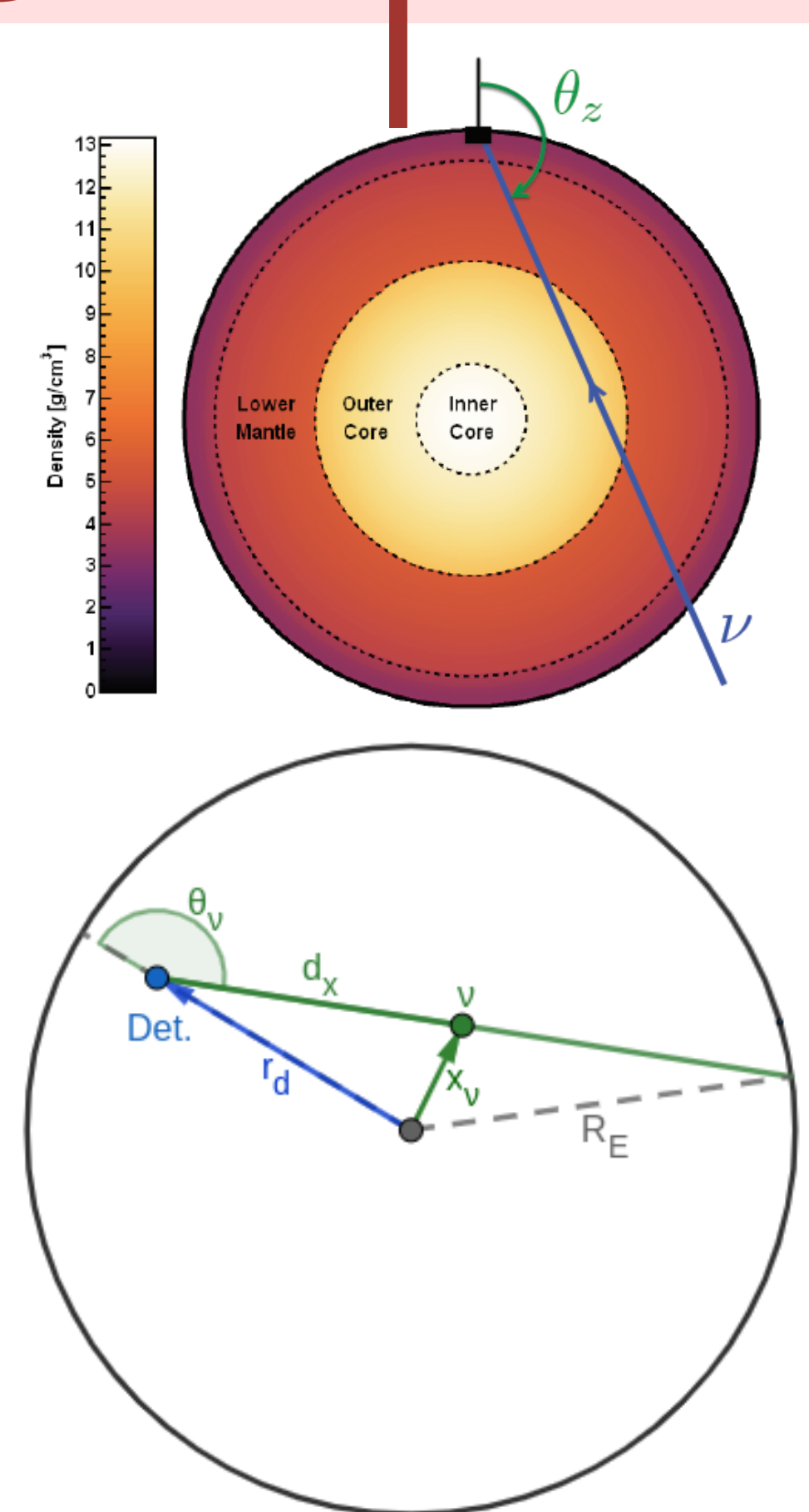
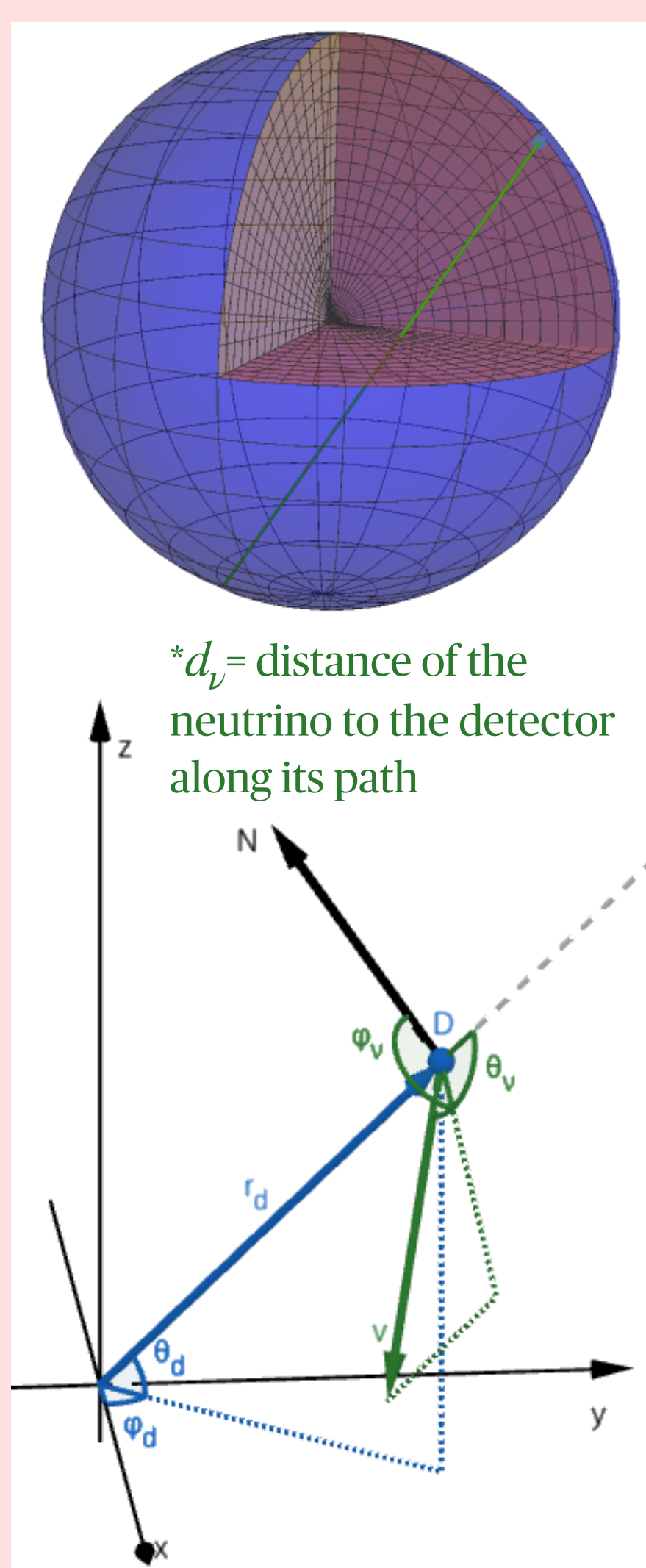


Figure taken from [7]

- Coordinates of the detector in the Earth's frame:  $d_d$  (depth),  $\theta_d$  (latitude),  $\phi_d$  (longitude)
- Coordinates of the neutrino trajectory as measured at the detector:  $d_v^*$ ,  $\theta_v$ ,  $\phi_v$

All these coordinates need to be combined to obtain: the neutrino trajectory in the Earth's coordinate frame

Trajectory is distributed among radial,  $\theta$  and  $\phi$  bins.



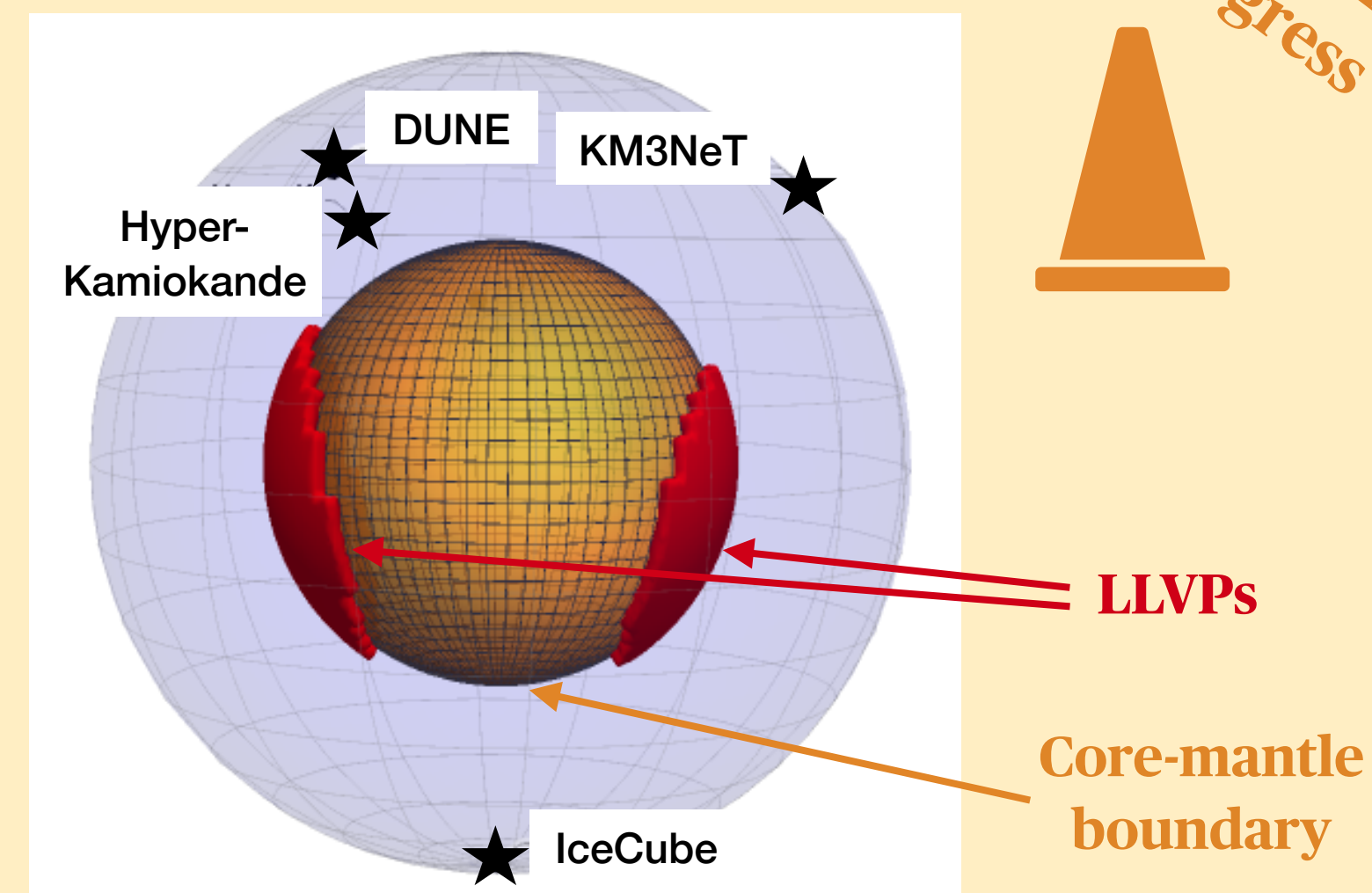
\* $d_v^*$  = distance of the neutrino to the detector along its path

## 5 Framework validation

PREM with modifiable region Example: Binned "pancake" in the location and with the size of the African and Pacific LLVPs

African LLVP: radius = 3572 km, height = 274 km, diameter = 5400 km

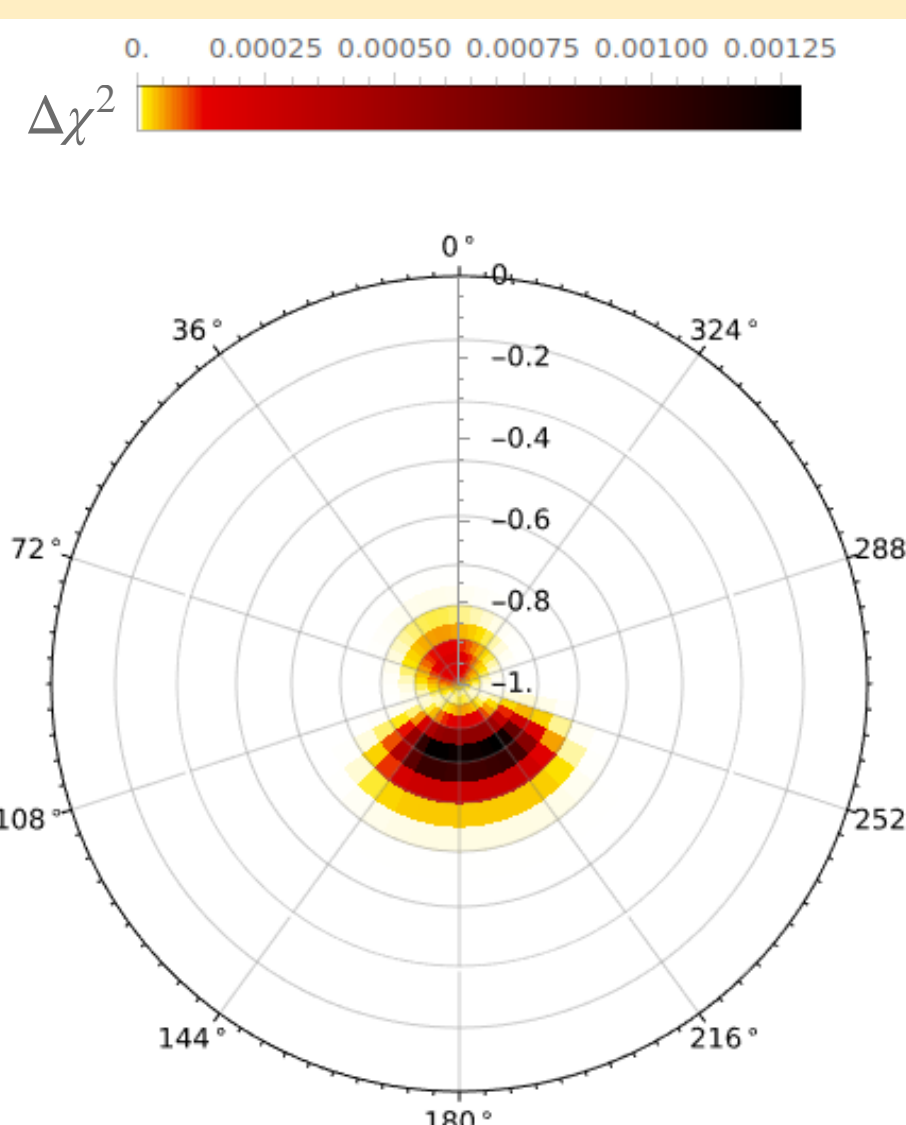
Pacific LLVP: radius = 3572 km, height = 200 km, diameter = 5000 km



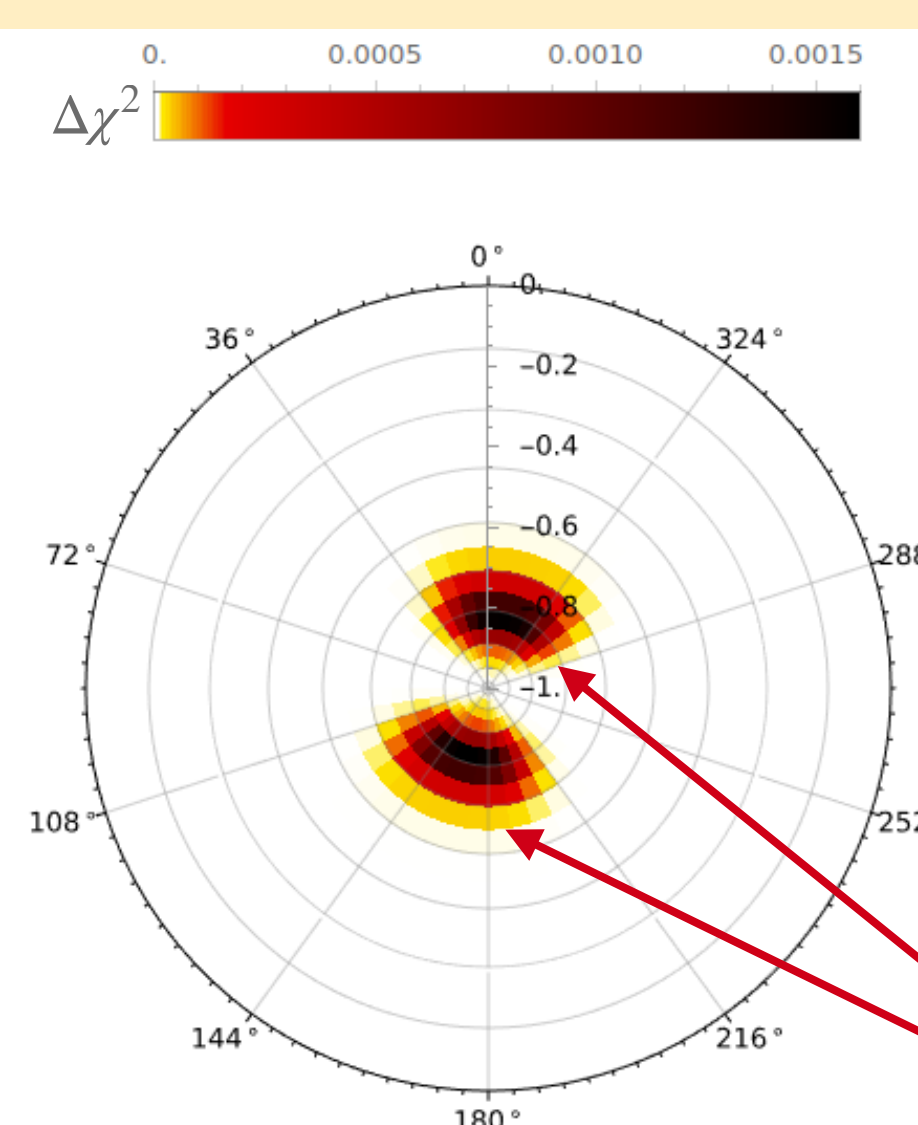
- Outlook:
- Proof of concept for the detection of large inhomogeneities in the deep Earth by neutrino detectors
  - Ongoing study exploring the requirements for next-generation neutrino detectors to achieve desired sensitivity

Sensitivity to discern between PREM and a +3% LLVP density anomaly:

\*Next-generation detector in the Mediterranean Sea: Total  $\Delta\chi^2 = 0.045$



Next-generation detector in the South Pole: Total  $\Delta\chi^2 = 0.059$



\*Next-generation detector = hypothetical detector with a size of 10 Mt, an energy resolution of <10%, an angular resolution of -7 degrees and an exposure time of 20 years.

[1] A. Dziewonski et al., Phys. Earth Planet. Int. 25, 4 (1981).  
 [2] A. McNamara et al., Tectonophysics 760 (2019).  
 [3] S. Karato, EPSL 301, 3-4 (2011).  
 [4] K. Konishi et al., J. Geophys. Res. Solid Earth 125, 2 (2020).  
 [5] S. Bourret, Ph. D. Thesis, USPC (2018).  
 [6] Honda M. et al., Phys. Rev. D 92, 2 (2015).

[7] Maderer L. et al., Front. Earth Sci. 11 (2023).

