A degeneracy between the effect of dark matter and strongly interacting matter at high densities

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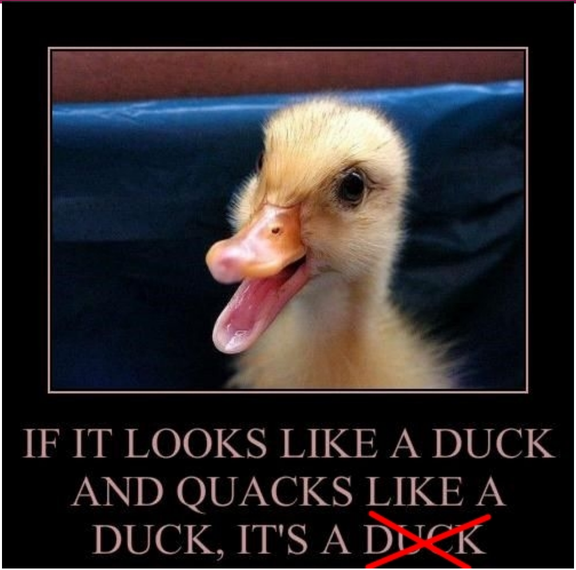
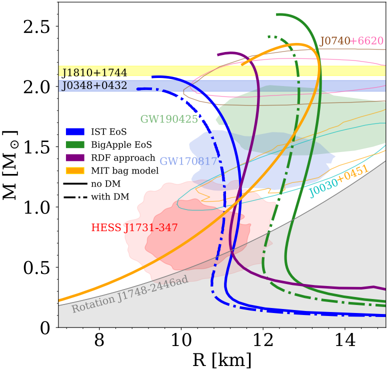
Sapienza University of Rome
12 – 16 February, 2024


IF IT LOOKS LIKE A DUCK
AND QUACKS LIKE A
DUCK, IT'S A DUCK

Recap of existing constraints on NS matter EoS

- In NSs and NS mergers we can probe the EoS at densities up to $\sim 5n_0$
- Gravitational-wave inference of GW170817 and HIC suggest soft EoS at $2-3n_0$
- HESS J1731-347 favours a very soft EoS at $2-2.5n_0$
- To reach $2M_\odot$ the EoS should be stiff enough at $> 3n_0$

Oliinychenko et al. (2023)
Danielewicz et al. (2002)
Demorest et al. (2010)
Antoniadis et al. (2013)
Doroshenko et al. (2022)



Sagun et al. (2023)

DM accumulation regimes

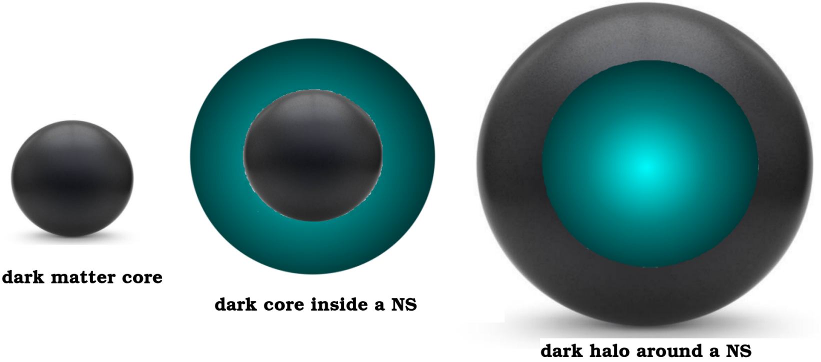
- Progenitor
  During the star formation stage the initial mixture of DM and BM contracting to form the progenitor star. Trapped DM undergoes scattering processes with baryons leading to its kinetic energy loss and thermalisation.
- Main sequence (MS) star
  From this stage of star evolution accretion rate increases due to big gravitational potential of the star. In the most central Galaxy region $M_{acc} \approx 10^{-5} M_{\odot} - 10^{-9} M_{\odot}$.
- Supernova explosion & formation of a proto-NS
  The newly-born NS should be surrounded by the dense cloud of DM particles with the temperature and radius that corresponds to the last stage of MS star evolution, i.e. a star with a silicone core.
  Kouvaris & Tinyakov 2010
  In addition, a significant amount of DM can be produced during the supernova explosion and mostly remain trapped inside the star.
- Equilibrated NS

$$M_{acc} \approx 10^{-14} \left(\frac{\rho_\chi}{0.3 \frac{GeV}{cm^3}} \right) \left(\frac{\sigma_{\chi n}}{10^{-45} cm^2} \right) \left(\frac{t}{Gyr} \right) M_{\odot}, \quad (1)$$

In the most central Galaxy region $M_{acc} \approx 10^{-5} M_{\odot} - 10^{-8} M_{\odot}$.

- Rapid DM accumulation
  A rapid DM accumulation could occur while passing through an extremely dense regions with primordial DM clumps
  Bramante et al. (2022)

DM and NS structure

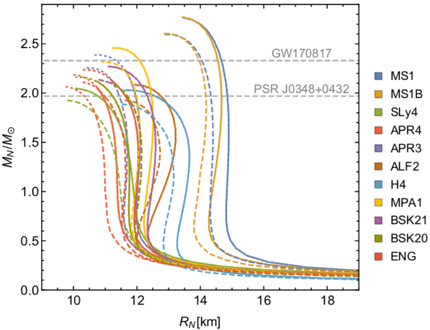


Dark matter and baryon components do not expel each other but overlap due to absence of non-gravitational interaction

Effect of DM on Mass and Radius

- DM core ⇒ decrease of the maximum mass and observed stellar radius
- DM halo ⇒ increase of the maximum mass and the outermost radius

Ciarcelluti & Sandin 2011; Nelson+ 2019; Deliyergiyev+ 2019; Ivanytskyi+2020; Das+ 2020; Del Popolo+ 2020; Karkevandi+ 2022



DM core contributing to 5% of the total NS mass

$\sqrt{\sigma_D}/m_D^3 = 0.05\,\mathrm{GeV}^{-2}$

Ellis+ 2018

TOV equations - two fluid system

2 TOV equations:

$$\frac{dp_B}{dr} = -\frac{(\epsilon_B + p_B)(M + 4\pi r^3 p)}{r^2\,(1 - 2M/r)}$$

$$\frac{dp_D}{dr} = -\frac{(\epsilon_D + p_D)(M + 4\pi r^3 p)}{r^2\,(1 - 2M/r)}$$

BM and DM are coupled only through gravity, and their energy-momentum tensors are conserved separately

total pressure $p(r) = p_B(r) + p_D(r)$

gravitational mass $M(r) = M_B(r) + M_D(r)$, where $M_j(r) = 4\pi \int_0^r \epsilon_j(r')r'^2 dr'$ (j=B,D)

$M_T = M_B(R_B) + M_D(R_D)$ - total gravitational mass

Fraction of DM inside the star:

$$f_\chi = \frac{M_D(R_D)}{M_T}$$

Asymmetric Bosonic Dark Matter

The minimal Lagrangian includes the complex scalar χ and real vector ω^μ fields, which are coupled through the covariant derivative $D^\mu = \partial^\mu - ig\omega^\mu$ with g being the corresponding coupling constant

$$\mathcal{L} = (D_\mu \chi)^* D^\mu \chi - m_\chi^2 \chi^* \chi - \frac{\Omega_{\mu\nu}\Omega^{\mu\nu}}{4} + \frac{m_\omega^2 \omega_\mu \omega^\mu}{2} \quad (2)$$

where $\Omega^{\mu\nu} = \partial^\mu \omega^\nu - \partial^\nu \omega^\mu$ and m_ω is the vector field mass.

Using a mean field approximation for ω, we get

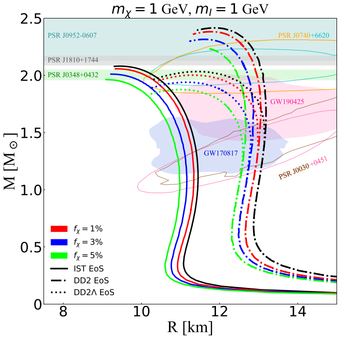
$$p_\chi = \frac{m_I^2}{4}\left(m_\chi^2 - \mu_\chi \sqrt{2m_\chi^2 - \mu_\chi^2}\right) \quad (3)$$

$$\varepsilon_\chi = \frac{m_I^2}{4}\left(\frac{\mu_\chi^3}{\sqrt{2m_\chi^2 - \mu_\chi^2}} - m_\chi^2\right) \quad (4)$$

Chemical potential is limited

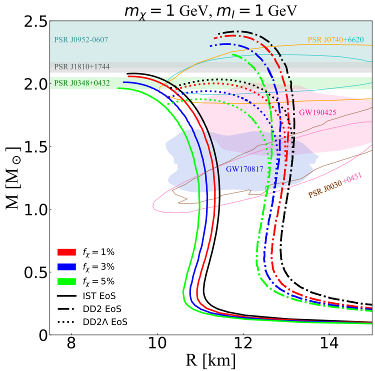
$\mu_\chi \in [m_\chi, \sqrt{2}m_\chi]$, m_χ - boson mass

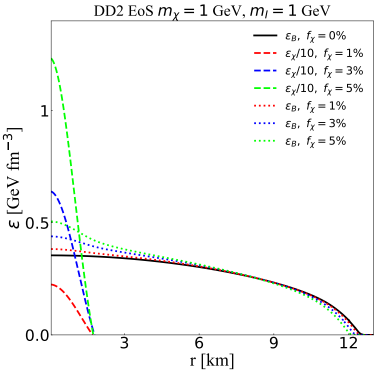
$m_I = \frac{m_\omega}{g}$ - interaction scale



Giangrandi+ 2022

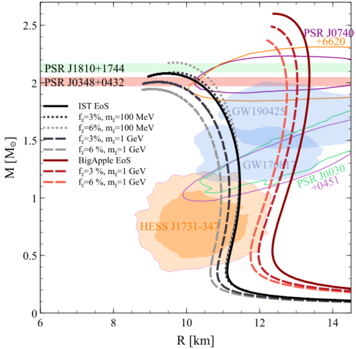
DM admixed NSs

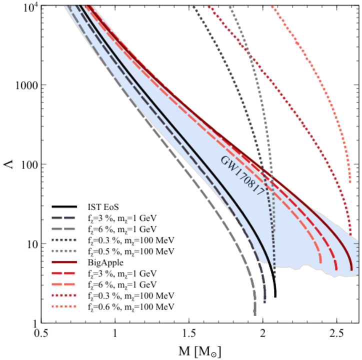




Giangrandi+ 2022

DM admixed NSs





Tidal deformability parameter

$$\Lambda = \frac{2}{3} k_2 \left(\frac{R_{\text{outermost}}}{M_{\text{tot}}} \right)^5$$

k_2 – Love's number

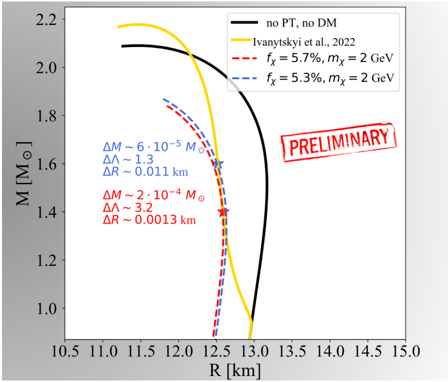
- $R_{outermost} = R_B \geq R_D$ - DM core
- $R_{outermost} = R_D > R_B$ - DM halo

Speed of sound should be calculated for two-fluid system **Giangrandi+ 2022**

Degeneracy between the DM and QGP cores

- DM and QGP cores may present undistinguishable mass, radius and tidal deformability;

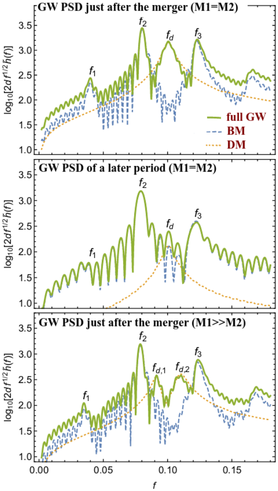
How to split this degeneracy?



Sagun et al. 2024 In prep.

An accumulated DM inside compact stars could mimic an apparent stiffening of strongly interacting matter equation of state and constraints we impose on it at high densities.

Effect of DM on GW waveform



Giudice+ 2016; Ellis+ 2018; Bezares+ 2019

The DM cores may produce a supplementary peak in the characteristic GW spectrum of NS mergers, which can be clearly distinguished from the features induced by the baryon component

Initial setups

- DM is treated as a Relativistic Fermi gas of particles with mass m_{DM} and spin one-half

  Ivanytskyi+ 2020
- BM is described by Sly4 EoS
- Both DM core and halo configurations along with pure BM NSs
- Different DM mass fractions
- Different resolutions: 128, 144 and 192 points
- Two different total mass to better study the DM effects on the post-merger phase
- Quasi-equilibrium configuration obtained through the sgrid code.

ID	m_{DM}	f_{DM}	M_{tot}	Configuration
0	-	0%	2.4 [M_{solar}]	Pure BM
1	-	0%	2.8 [M_{solar}]	Pure BM
2	1 GeV	3%	2.4 [M_{solar}]	DM core
3	1 GeV	3%	2.8 [M_{solar}]	DM core
4	1 GeV	15%	2.4 [M_{solar}]	DM core
5	1 GeV	15%	2.8 [M_{solar}]	DM core
6	0.17 GeV	0.5%	2.4 [M_{solar}]	DM halo
7	0.17 GeV	0.5%	2.8 [M_{solar}]	DM halo

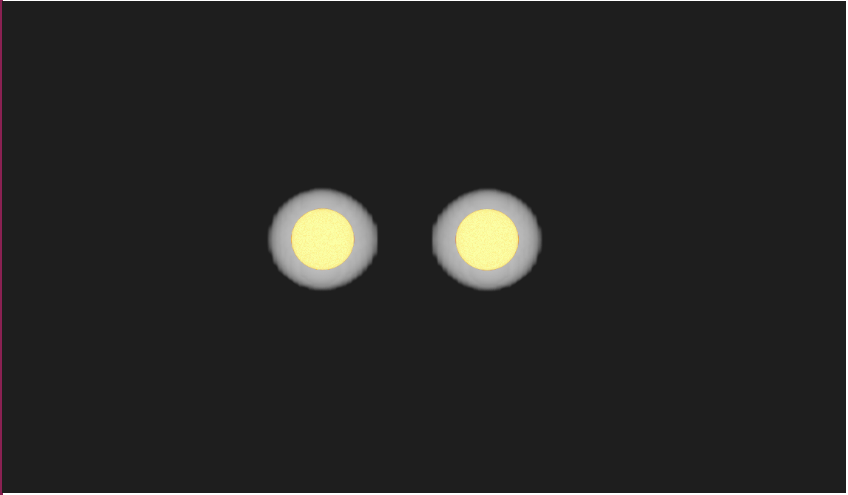
Rüter+ 2023; Giangrandi+ 2024 (In prep)

Mergers of Dark Matter Admixed Neutron Stars: core

- Dark Core configuration
- Baryonic matter: SLy4 EoS
- Dark matter: 1 GeV fermions, 5 % of mass
- $1.4\ M_\odot + 1.4\ M_\odot$
- Eccentricity ≈ 0
- Non-spinning Stars

Rüter+ 2023; Giangrandi+ 2024 (In prep)

Mergers of Dark Matter Admixed Neutron Stars: halo



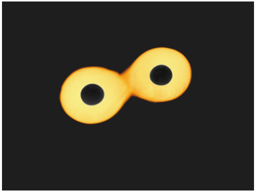
Rüter+ 2023; Giangrandi+ 2024 (In prep)

Dark matter core simulations

- Higher DM fraction leads to a longer inspiral likely due to a lower deformability of DM-admixed NSs;
- Faster formation of the BH after the merger and harder to eject material from the bulk of the stars prior to the BH formation;
- The lack of DM ejecta and debris disks is related to its concentration in the NS core;
- DM component might remain gravitationally bound after the merger of BM and orbit the center of the remnant with an orbital separation of a few km;
- The orbital separations of typically a few km is resulted in a kHz-band GW signal that could be sought in GW searches;

Bauswein et al. (2023)

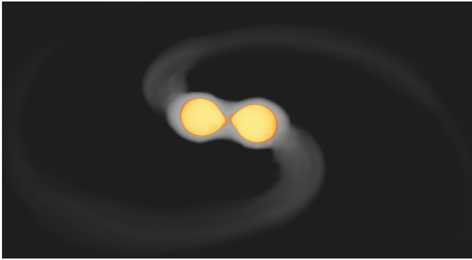
- The DM core and a host star are likely to spin at different rotational frequencies just after the merger due to the absence of non-gravitational interaction. Further on, they may synchronise via the gravitational angular momentum transfer, including tidal effects;
- DM core favours a formation of a one-arm spiral instability.



Rüter+ 2023; Giangrandi+ 2024 (In prep)

Dark matter halo simulations

- At t=0, two DM-admixed stars have still not touched each other;
- Higher DM fractions leads to more extended DM haloes, and, consequently, to higher tidal deformabilities;



Rüter+ 2023; Giangrandi+ 2024 (In prep)

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Conclusions

- DM can be accumulated in the core of a NS $\Rightarrow$ significant decrease of the maximum mass and radius of a star.
- DM halo $\Rightarrow$ increase of the maximum mass and the outermost radius.
- The secondary component of the GW190814 binary merger or HESS J1731-347 might be a DM admixed NS.

Changing the position of the NS in the Galaxy the accretion rate of DM varies, which in turn leads to different amount of DM

$\downarrow$

different modifications of M, R, surface temperature, etc

The effect of DM could mimic the properties of strongly interacting matter

Smoking gun of the presence of DM in NSs

- **by measuring mass, radius, and moment of inertia of NSs with few-%-accuracy.**

To see this effect we need high precision measurement of M and R of compact stars as well as NS searches in the central part of the Galaxy with

radio telescopes: MeerKAT, SKA, ngVLA plan to increase radio pulsar timing and discover Galactic center pulsars.

space telescopes: NICER, ATHENA, eXTP, STROBE-X are expected to measure M and R of NSs with high accuracy.

DM core $\Rightarrow$ mass and radius reduction of NSs toward the Galaxy center
DM halo $\Rightarrow$ mass increase of NSs toward the Galaxy center
or variation of mass and radius in different parts of the Galaxy

- **by performing binary numerical-relativity simulations and kilonova ejecta for DM-admixed compacts stars for different DM candidates, their particle mass, interaction strength and fractions with the further comparison to GW and electromagnetic signals.**

Large statistics on NS-NS, NS-BH mergers by LIGO/Virgo/KAGRA would be very helpful
The smoking gun of the presence of DM could be:
supplementary peak in the characteristic GW spectrum of NS mergers; exotic waveforms; modification of the kilonova ejection;
post-merger regimes: the next generation of GW detectors, i.e., the Cosmic Explorer and Einstein Telescope.

- **by detecting objects that go in contradiction with our understanding.**

As a potential candidate for a DM-admixed NS could be the secondary component of GW190814.

- **High/low surface temperature of NSs towards the Galaxy center**

Thanks for your attention!