# constraints from the satellite dwarf galaxies of the Milky Way

Stacy Kim University of Surrey SIDM Pollica Workshop 2023

# constraints from the satellite dwarf galaxies of the Milky Way

Stacy Kim University of Surrey SIDM Pollica Workshop 2023 Stars live where densities can differ significantly from CDM:



Key classical SIDM prediction: low density central cores

Stars live where densities can differ significantly from CDM:



Thus stellar kinematics can be highly sensitive to SIDM!

as well as other processes/models with differing central densities



observed dwarf galaxy



observed dwarf galaxy with luminosity L, velocity dispersion  $\sigma$ 

Milky Way





with radial distribution of satellites, estimate total # of unseen dwarfs observed dwarf galaxy with luminosity L, velocity dispersion  $\sigma$ 

completeness radius

\_\_\_\_ Milky Way

"completeness correction"

with radial distribution of satellites, estimate total # of unseen dwarfs observed dwarf galaxy with luminosity L, velocity dispersion  $\sigma$ 

completeness radius

— Milky Way

L, σ

observed dwarf galaxy with luminosity L, velocity dispersion  $\sigma$ 

completeness radius

Milky Way

repeat for all known dwarf galaxies!

observed dwarf galaxy with luminosity L, velocity dispersion  $\sigma$ 

completeness radius

L. σ

🔶 Milky Way

repeat for all known dwarf galaxies!

L<sub>2</sub>, σ<sub>2</sub>

complete*ness radius* 

L<sub>2</sub>, σ<sub>2</sub>

 $L_1, \sigma_1$ 

L<sub>3</sub>, σ<sub>3</sub>

L<sub>2</sub>, σ<sub>2</sub>

Milky Way

 $L_1, \sigma_1$ 

with completeness corrections, -

 $L_{3}, \sigma_{3}$ 

repeat for all known dwarf galaxies!

L<sub>3</sub>, σ<sub>3</sub> L<sub>2</sub>, σ<sub>2</sub>  $L_1, \sigma_1$ sum all to estimate MW's true velocity function Milky Way

L<sub>3</sub>, σ<sub>3</sub>

L<sub>2</sub>, σ<sub>2</sub>

Milky Way

L<sub>1</sub>, σ<sub>1</sub>

sum all to estimate MW's true velocity function



# corrected velocity function







## theory vs. observations



## theory vs. observations



## theory vs. observations



## implications for self-interactions



# implications for self-interactions



# implications for self-interactions



## What about M31's satellites?



sample incomplete for velocity function analysis (but stay tuned...)











Requires concentrations that are 2-3 sigma outliers(!)

Charles, SYK+, in prep.











## constraints from M31 satellites




# constraints from M31 satellites





# Core collapse in satellites of lenses





lessons from circular orbits

subhalo central density grows if cooling > heating (recall: negative heat capacity)

evaporation is significant! can be strong enough to disrupt core-collapse



# Core collapse in satellites of lenses



Parameter space for subhalo core collapse

core collapse not feasible with constant  $\sigma$ /m need very high concentrations (median c = 10-20)

smaller c needed for isolated subhalos or <u>no evaporation</u> mimics vSIDM!



if ultra-compact substructure found, strongly favors vSIDM, inelastic SIDM, etc

Zeng+ incl. SYK 2022

# SIDM constraints from satellites

Milky Way satellite kinematics imply  $\sigma/m \leq 0.5 \text{ cm}^2/\text{g}$ 

Kinematics of M31 satellites And XXI and And XXV cannot be explained with constant cross sections

Core collapse could reestablish steeper densities, but requires unphysically high concentrations Important to include evaporation, which can suppress core-collapse

> The door is still open for other SIDM models with additional degrees of freedom! (velocity-dependent SIDM, inelastic scattering, etc.)

# EXTRAS: MW VF



Lovell+ 2020 Kim et al. 2022 (arXiv:2106.09050)



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# theoretical uncertainties

an alternative method to reduce the mismatch at 10 km/s



# theoretical uncertainties



# theoretical uncertainties

addressing the too many satellites problem  $10^{3}$  $f_{lum} = 1$ tiny halos form stars?  $z_{re} = 11.3$ below atomic cooling limit  $z_{\rm re} = 9.3$ observed later reionization?  $10^{2}$ corrected  $\mathrm{N}(>\sigma^*_{\mathrm{los}})$  $10^{1}$ mleft = 1mleft = 0.1mleft = 0.01 $10^{0}$  $10^{1}$  $\sigma_{\rm los}^* \, ({\rm km/s})$ 

#### see Graus+ 2019



Geha+, private communication









Kim et al. 2021 (arXiv:2106.09050)





 $D_{\text{Kim}et}^+ a_1.22027, (Brxiv:2106.09050)$ 









# sanity checks



#### sanity checks



# corrected luminosity function



# EXTRAS: M31


# EXTRAS: CORE COLLAPSE

#### new wrinkles in the fold

The outlook for constant cross sections is not promising!

undergoing mild core collapse

#### new wrinkles in the fold

The outlook for constant cross sections is not promising!

BUT there have been a couple recent developments...



#### new wrinkles in the fold

The outlook for constant cross sections is not promising!

BUT there have been a couple recent developments...



Core collapse can reintroduce dense cores, and even reestablish cusps.

Accelerated by tidal stripping?

Elbert+ 2015, Essig+ 2019, Nishikawa+ 2020



Energy exchange via selfinteractions leads to 'heat' flow.



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Phase 1 isothermal core forms



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#### Phase 2

core slowly loses heat to outskirts, dark matter infall to more bound orbits that are hotter than before



Energy exchange via selfinteractions leads to 'heat' flow.

Phase 1 isothermal core forms

#### Phase 2

core slowly loses heat to outskirts, dark matter infall to more bound orbits that are hotter than before

> more heat flow, more infall, runaway core collapse!

Simulating subhalos under core collapse is expensive.

scattering probability 
$$\propto \frac{\sigma_{\rm SI}}{m_{\chi}} \Delta v \ \rho \ \Delta t$$

Simulating subhalos under core collapse is expensive.

scattering probability  $\propto \frac{\sigma_{\rm SI}}{m_{\chi}} \Delta v \rho \Delta t$ increases by orders of magnitude

Simulating subhalos under core collapse is expensive.

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$$\frac{\text{scattering probability}}{\sum_{i=1}^{n} \sum_{j=1}^{n} \sum_{j=1}^{n}$$

<u>high dynamic range</u>: e.g. in substructure lenses,  $10^{13}$  M<sub> $\odot$ </sub> host (main lens) + as low as  $10^{6}$  M<sub> $\odot$ </sub> sub

Simulating subhalos under core collapse is expensive.

$$\frac{\text{scattering probability}}{m_{\chi}} \propto \frac{\sigma_{\text{SI}}}{m_{\chi}} \Delta v \rho \Delta t < 1$$

$$\frac{1}{m_{\chi}} \Delta v \rho \Delta t$$

<u>high dynamic range</u>: e.g. in substructure lenses,  $10^{13}$  M<sub> $\odot$ </sub> host (main lens) + as low as  $10^{6}$  M<sub> $\odot$ </sub> sub

We adopt a hybrid approach

analytic host + 'live' (N-body) subhalo + evaporation

host-sub interactions

that reduces computational time by orders of magnitude!









stop simulation when central density grows by 100







Lessons from circular orbits





## Core collapse in satellites of lenses

![](_page_93_Picture_1.jpeg)

![](_page_93_Figure_2.jpeg)

Lessons from circular orbits

subhalo central density grows if cooling > heating (recall: negative heat capacity)

evaporation is significant! can be strong enough to disrupt core-collapse

![](_page_93_Figure_6.jpeg)

Hybrid ccSIDM validation: mass loss

- Discrepancy < 10% for subs 1/1000 of the host
- Mostly due to missing dynamical friction
- But for smaller subhalos less significant

Can study *arbitrarily* small subhalos

![](_page_94_Figure_5.jpeg)

Adapted from slides by Carton Zeng

Hybrid ccSIDM validation: density profiles

- Good agreement w/live host simulation for both cored and core-collapsing
- Robust for the particle resolution
- Evaporation is significant

![](_page_95_Figure_4.jpeg)