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Boosting the production of sterile neutrino dark matter with self-interactions

María Dias

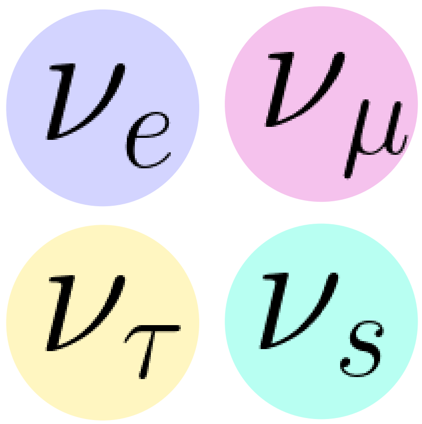
Based on arXiv: 2307.15565

In collaboration with:
Stefan Vogl

TeVPA 2023, Napoli

Sterile neutrinos as dark matter

What are sterile neutrinos?



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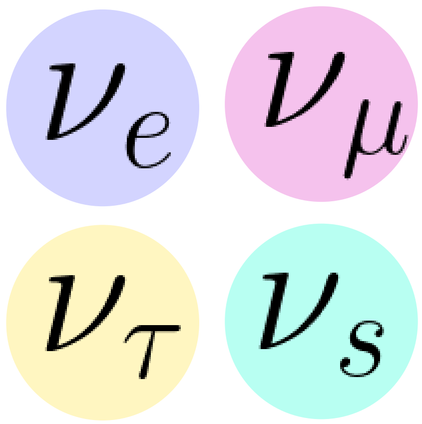

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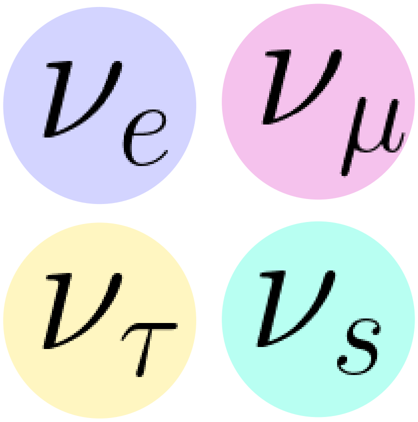

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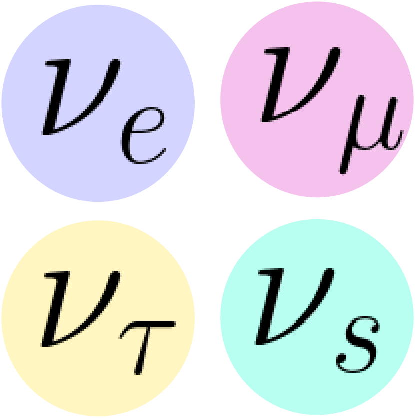

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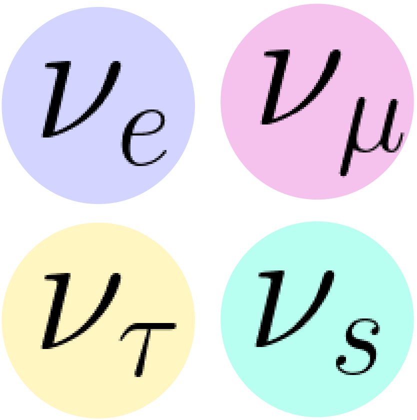

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How to extend DW?

Self-interactions

Johns and Fuller 1903.08296
Bringmann et al. 2206.10630

$$\mathcal{L}_{\text{int}} = y\ \bar{\nu}_s \nu_s \phi.$$

Production of sterile neutrinos

The evolution of the phase-space density of sterile neutrinos is given by the Boltzmann equation

$$\frac{\partial f_s}{\partial t} - Hp\frac{\partial f_s}{\partial p} = \frac{\Gamma_t}{4}\left(\frac{\omega^2 \sin^2(2\theta)}{\omega^2 \sin^2(2\theta) + \frac{\Gamma_t^2}{4} + [\omega \cos(2\theta) - V_{\text{eff}}]^2}\right)[f_a - f_s] + \mathcal{C}_s$$

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Active's distribution function

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Vacuum oscillation frequency $\sim \frac{{m_s}^2}{2p}$

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Hubble's rate

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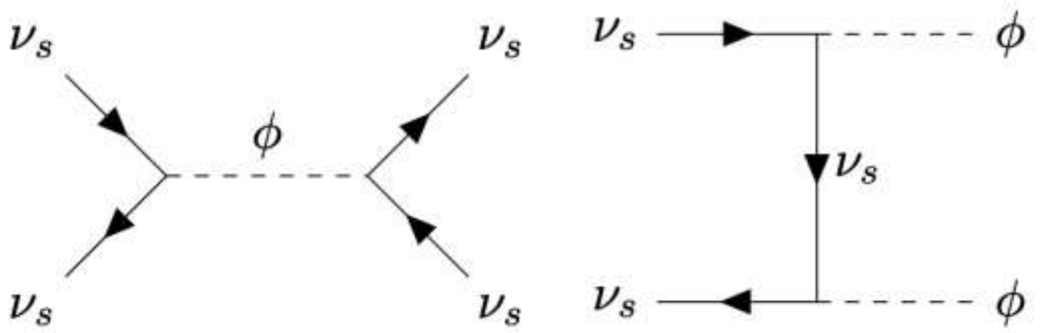
Vacuum mixing angle

Total rate

Active's distribution function

Sterile-sterile scattering processes

$$\Gamma_t = \Gamma_a + \Gamma_s$$



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Sterile's distribution function

Hubble's rate

Vacuum oscillation frequency $\sim \frac{{m_s}^2}{2p}$

Vacuum mixing angle

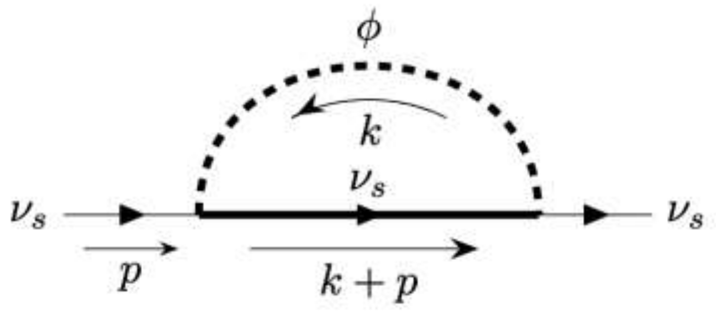
Total rate

Effective potential

Active's distribution function

Sterile-sterile scattering processes

$$V_{\text{eff}} = V_a - V_s$$



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Transition Probability

$$\langle P_m(\nu_a \leftrightarrow \nu_s)\rangle$$

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In-medium effective mixing angle

$$\sin^2(2\theta_m)$$

Production regimes

What kind of behaviour can we expect from the system?

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The relative size between the rate, the oscillaction frequency and the potential plays an essential role in the production

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Talk by Paul Frederik Depta

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$(-)$ $\; !$

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- The mass of the sterile neutrinos
- The mixing angle

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The are four parameters controlling the behavior of the system

- The mass of the sterile neutrinos
- The mixing angle
- The mass of the mediator

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The are four parameters controlling the behavior of the system

- The mass of the sterile neutrinos
- The mixing angle
- The mass of the mediator
- The size of the Yukawa coupling

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The are four parameters controlling the behavior of the system

- The mass of the sterile neutrinos
- The mixing angle
- The mass of the mediator →
- The size of the Yukawa coupling

$m_\phi \gg T$	$V_s < 0$
$m_\phi \ll T$	$V_s > 0$

Thermalization of the system

What role does the sterile collision term play?

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- C_s tries to drive the system towards equilibrium
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- The system will cool down

$$\Gamma_{\nu_s\nu_s\leftrightarrow\phi\phi}\, f_s \left(1 - \frac{f_s^2}{f_{\rm eq}^2}\right)$$

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The process stops itself

Thermalization of the system

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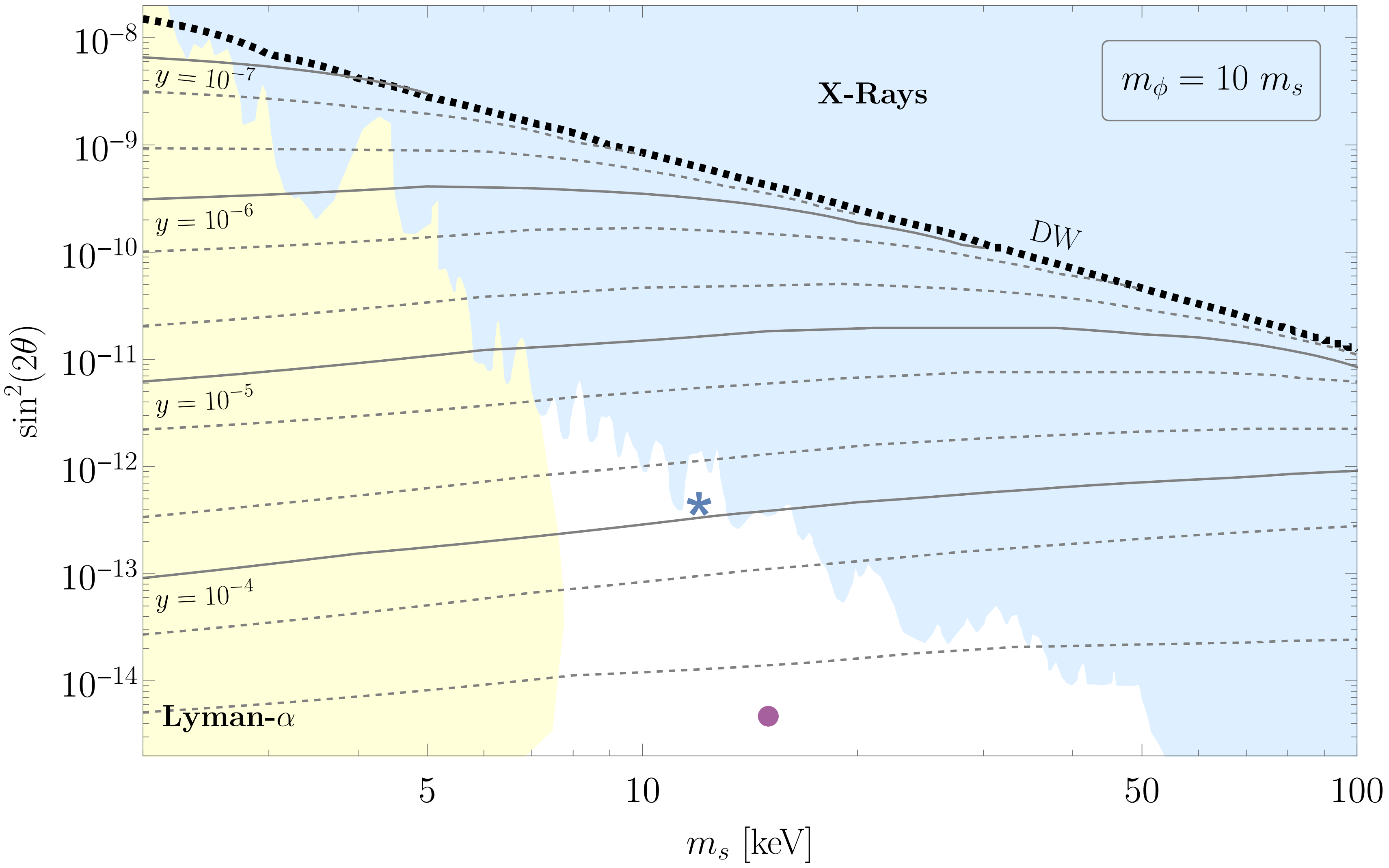
Only important for low and intermediate mediator masses

The parameter space



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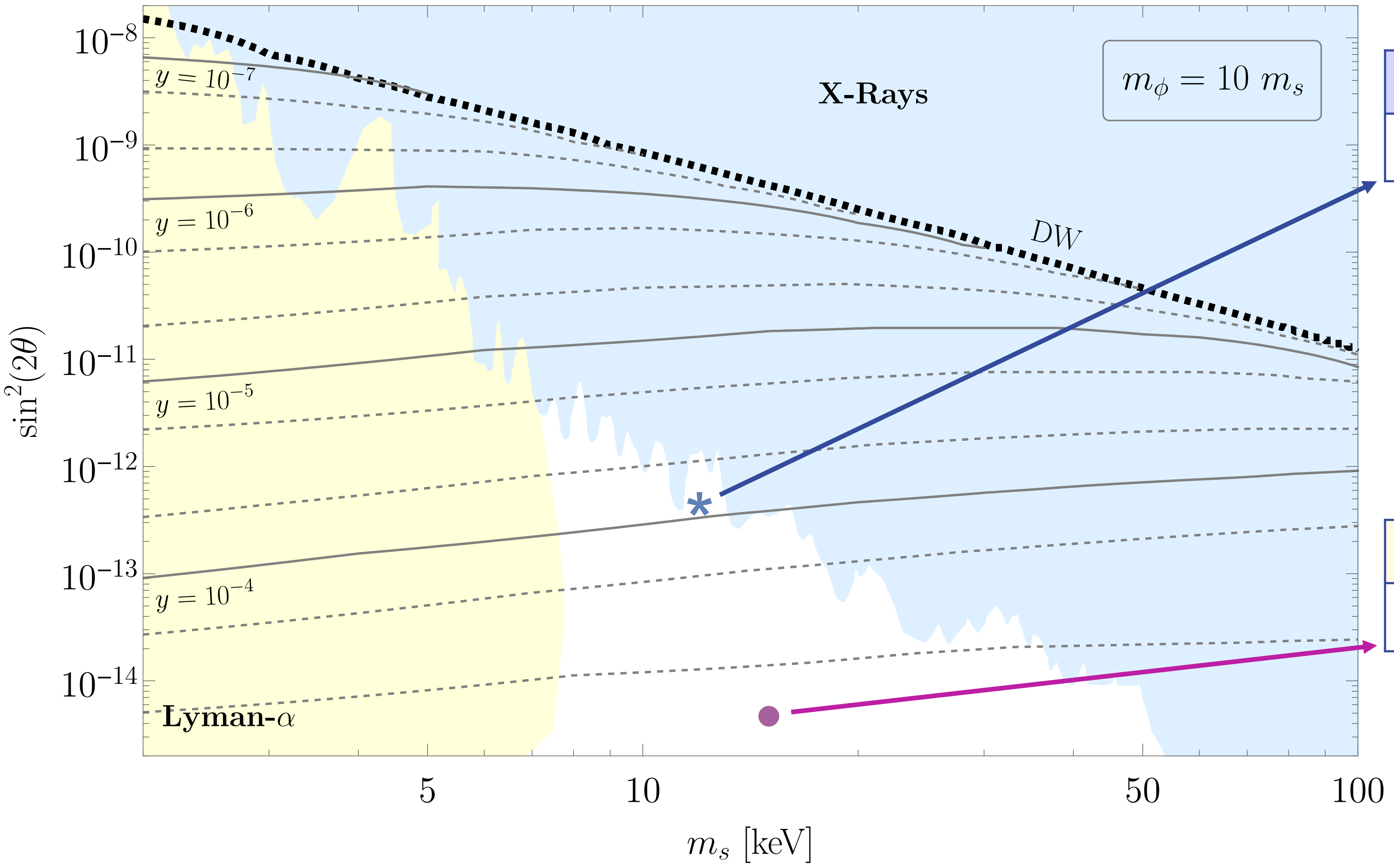

m_s	$\sin^2(2\theta)$
12 keV	$5\ x\ 10^{-13}$

m_s	$\sin^2(2\theta)$
15 keV	$4.68\ x\ 10^{-15}$

One heavy-mediator benchmark

m_s	m_ϕ	y	$\sin^2(2\theta)$
12 keV	1.5 GeV	$6.92\ x\ 10^{-2}$	$5\ x\ 10^{-13}$

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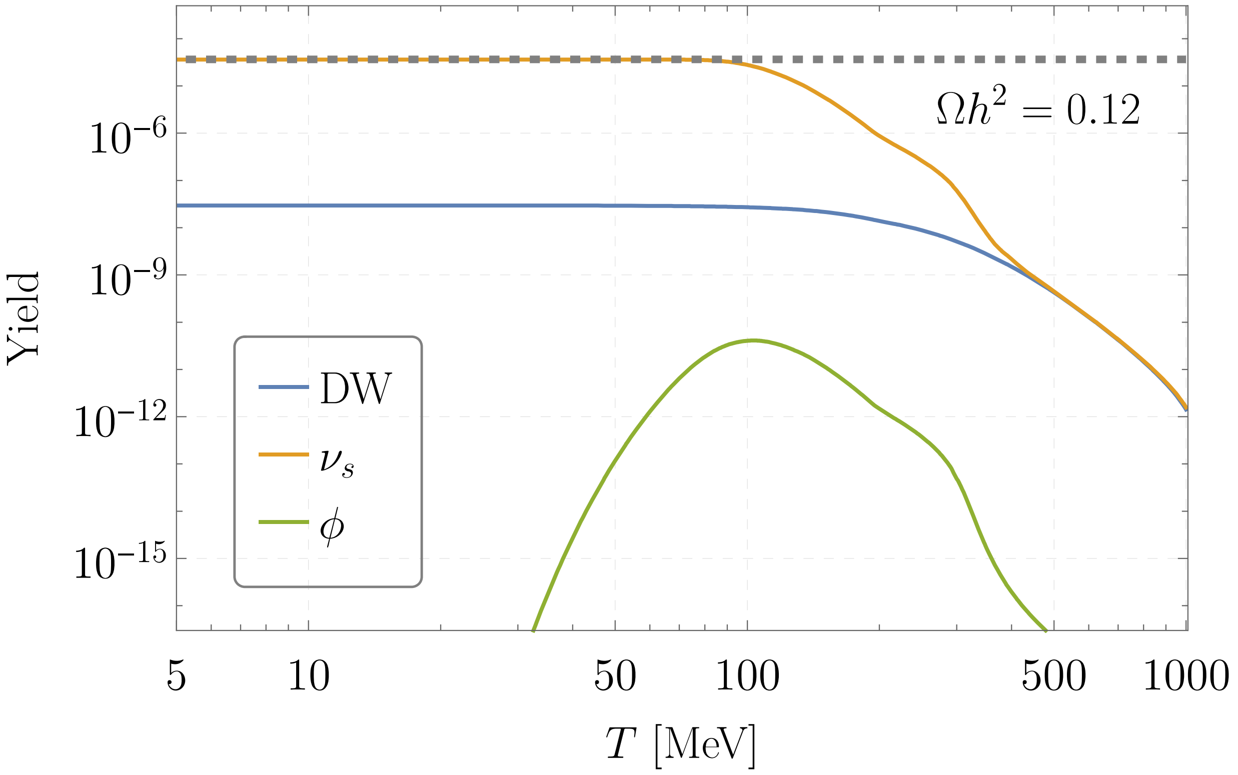

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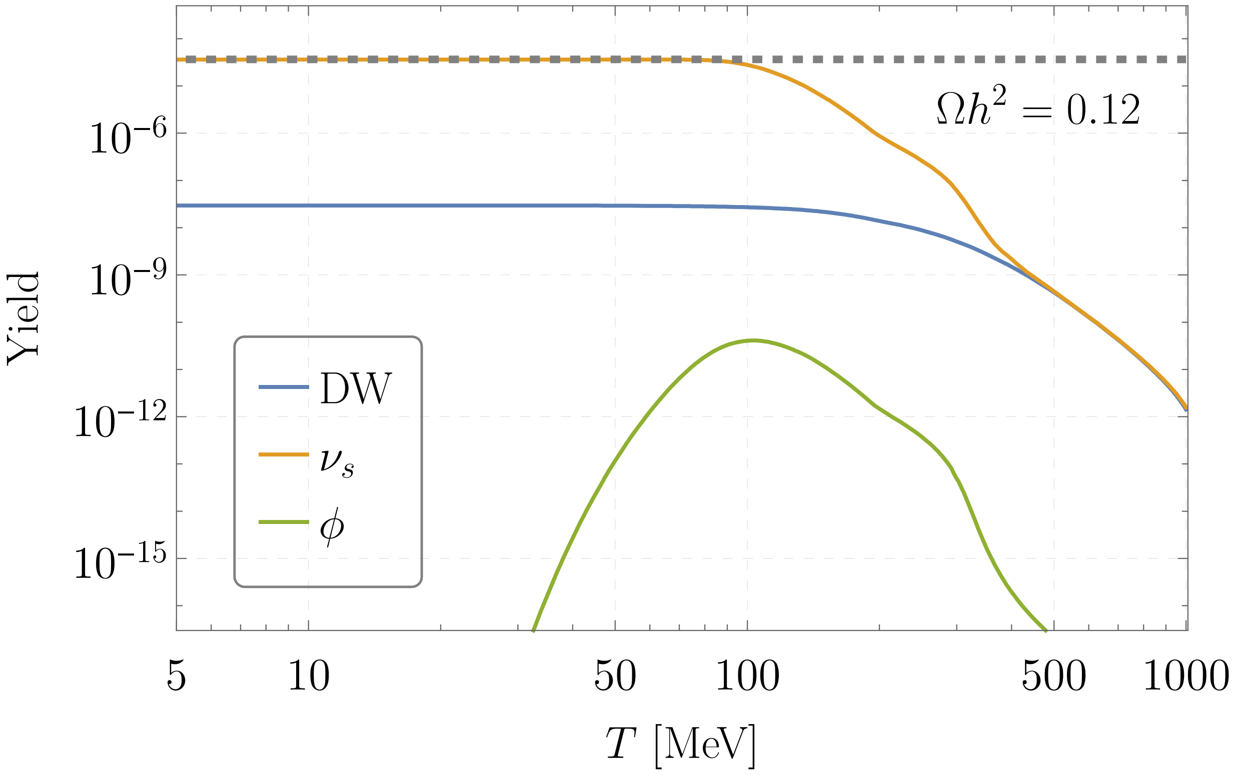

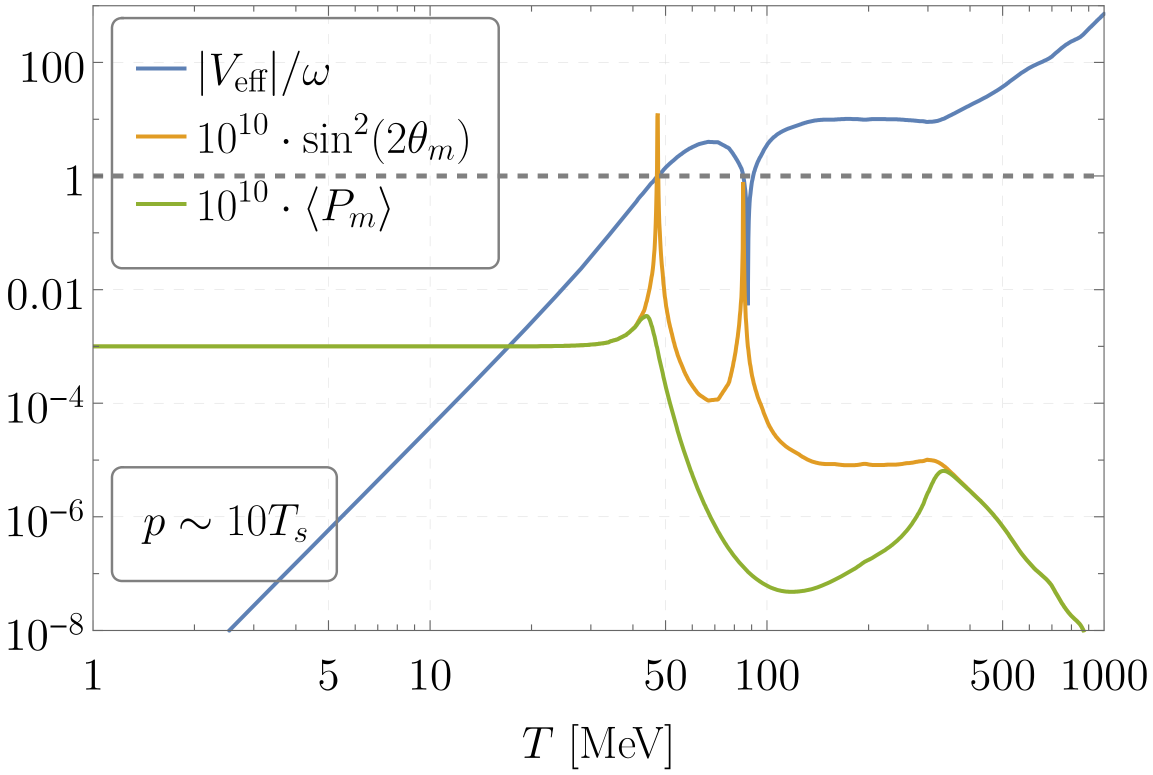

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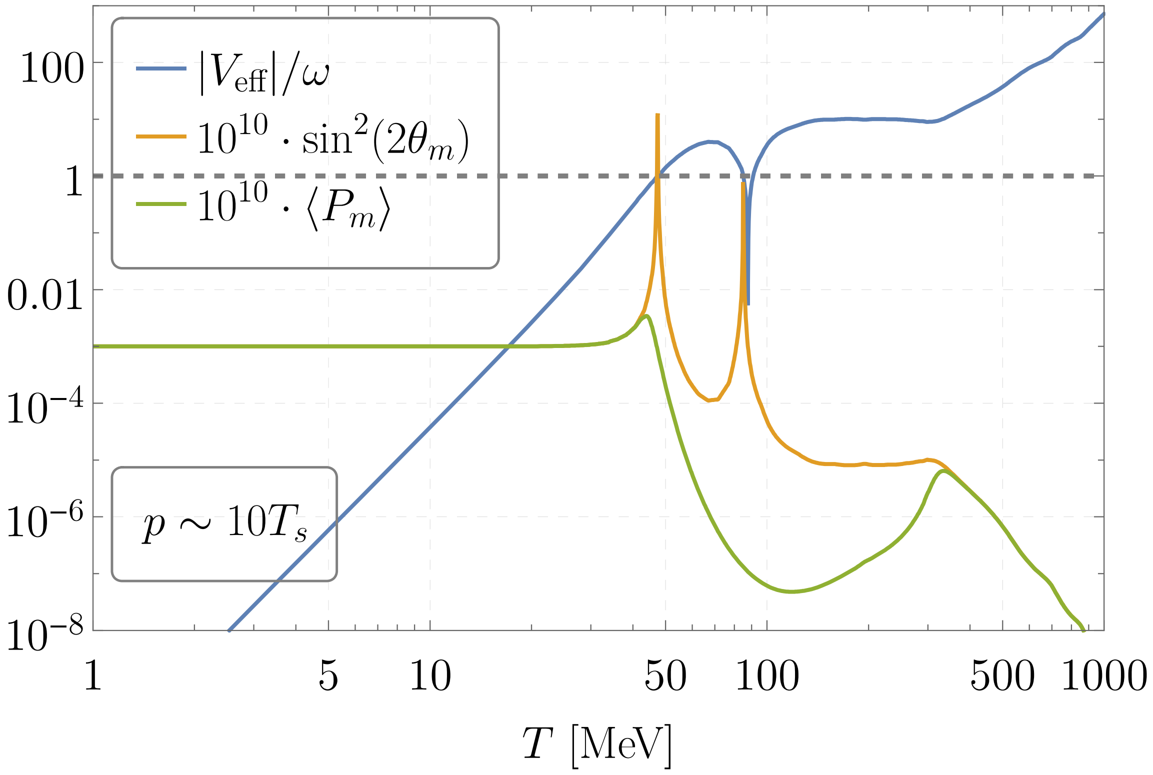

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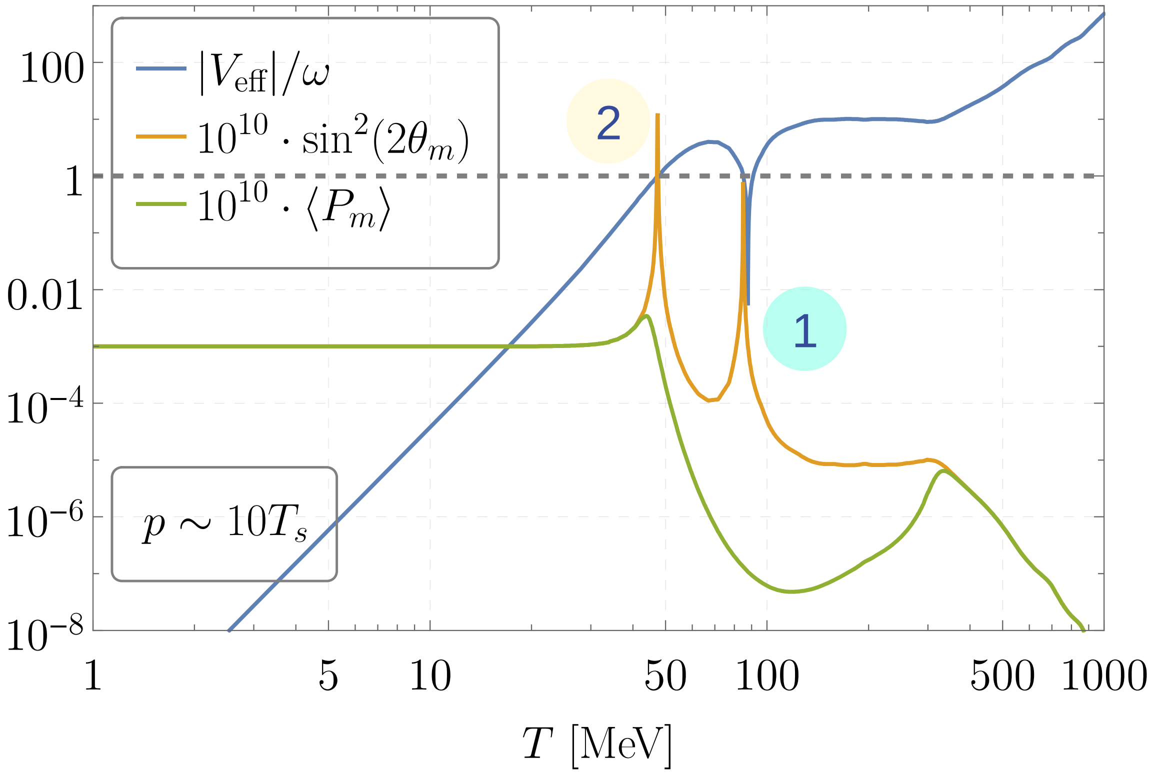

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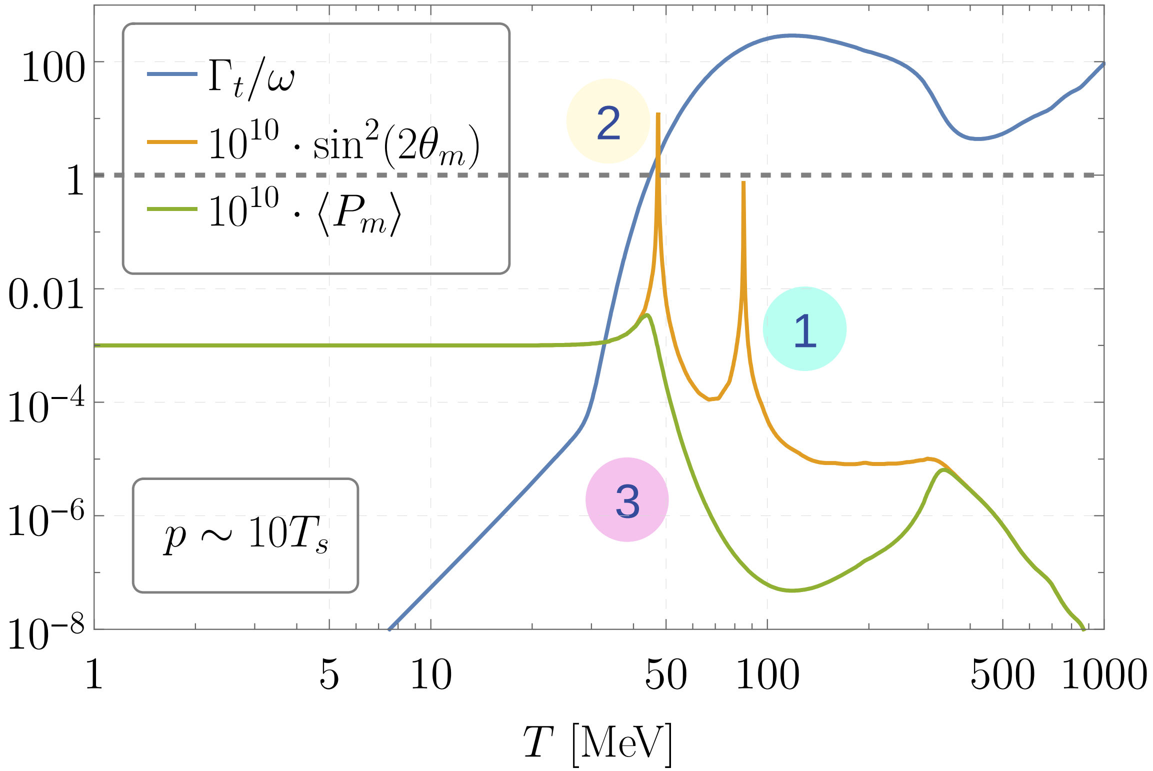
3 The resonances are regulated by quantum damping

$$\frac{\partial f_s}{\partial t} - Hp\frac{\partial f_s}{\partial p} \sim \frac{\omega^2 \sin^2(2\theta)}{\Gamma_t}[f_a - f_s] + \mathcal{C}_s$$



A benchmark with partial thermalization

m_s	m_ϕ	y	$\sin^2(2\theta)$
15 keV	150 keV	$7.5\ x\ 10^{-4}$	$4.68\ x\ 10^{-15}$

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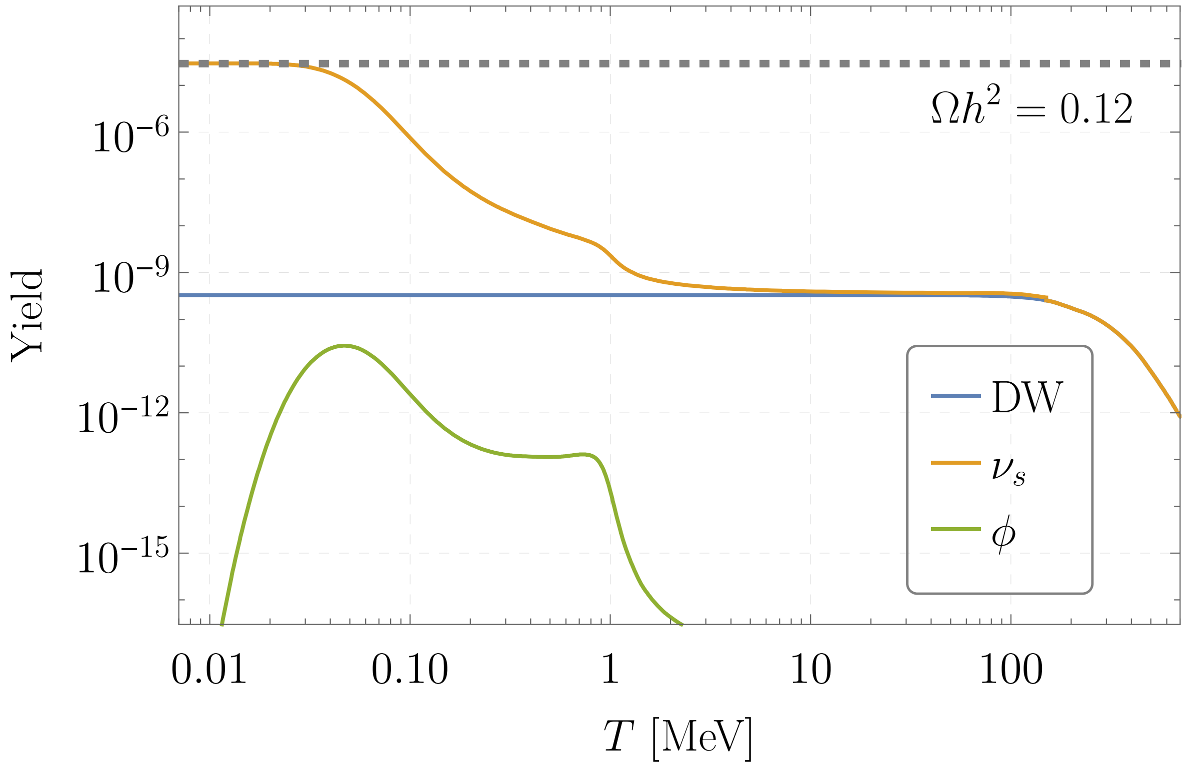

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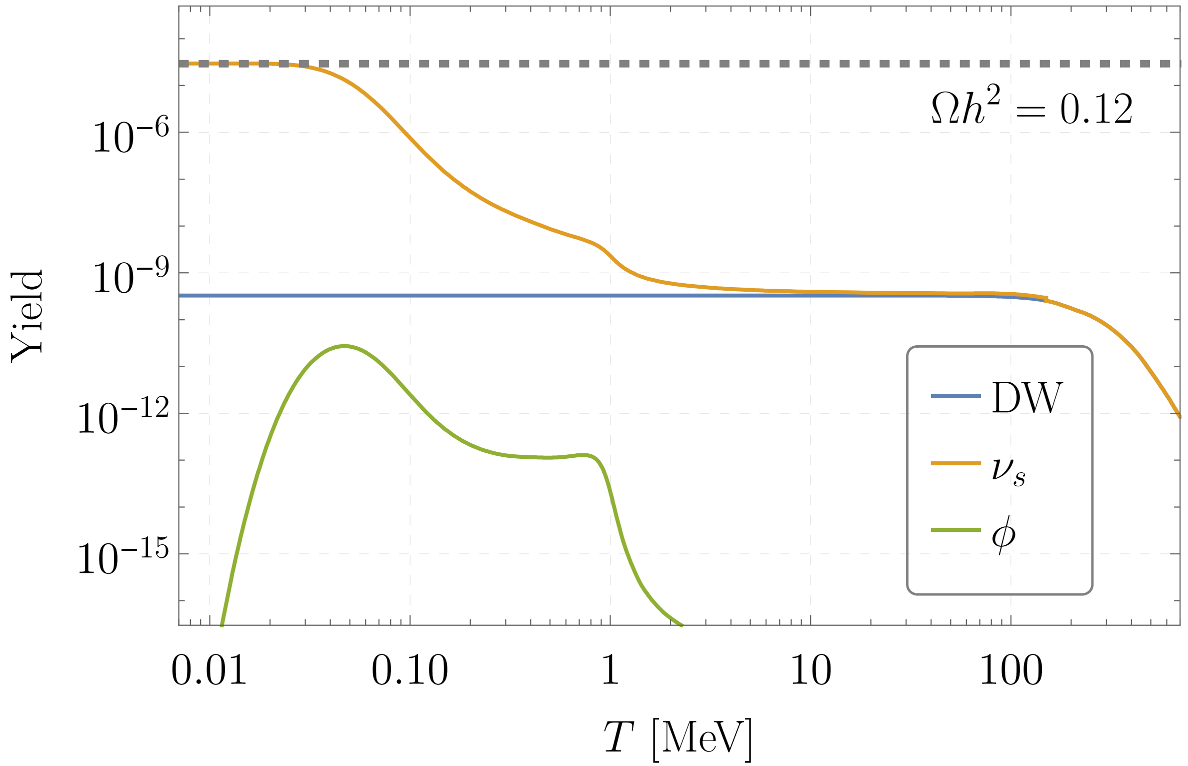

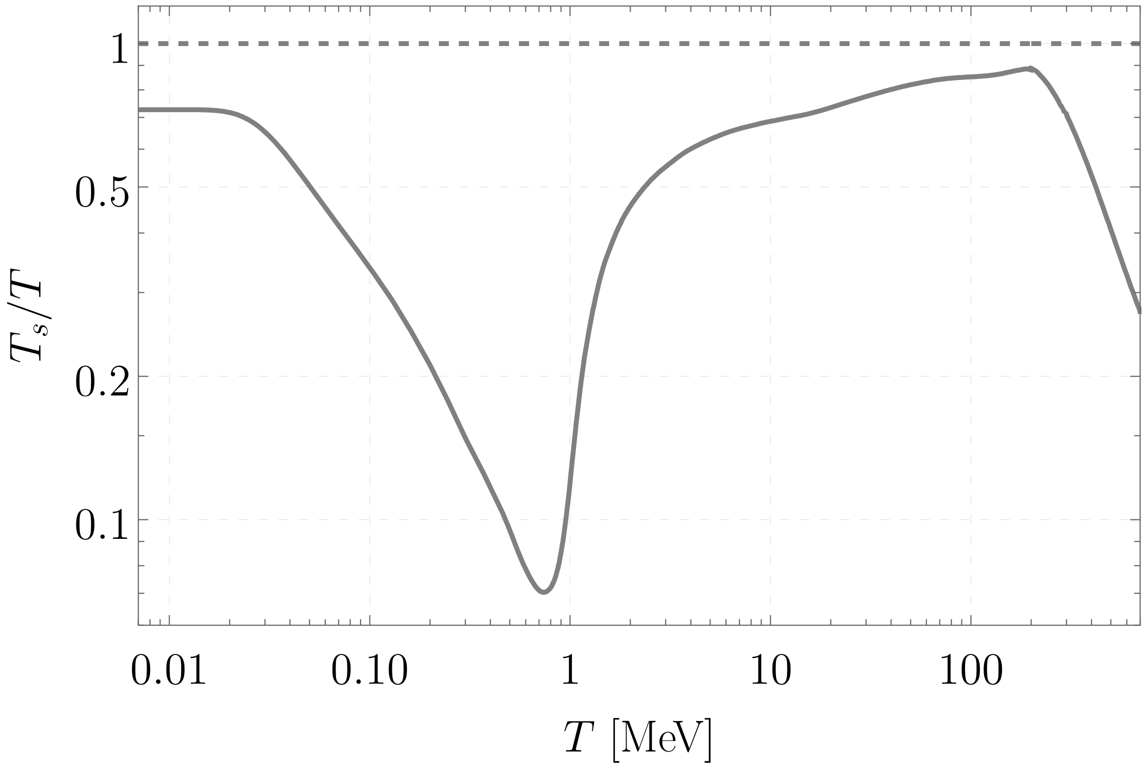

Overview of production regimes

$$m_\phi \lesssim 100\ \text{MeV}$$

- Simplified Boltzmann equation
- Initial DW production

Bringmann et al. 2206.10630

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This work 2307.15565

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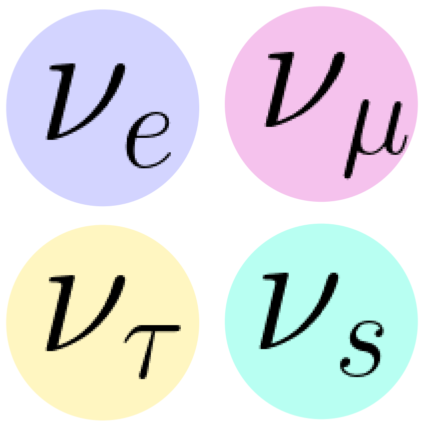
- Runaway production

Johns and Fuller 1903.08296

Conclusions

- Sterile neutrinos are attractive BSM candidates



Conclusions

- Sterile neutrinos are attractive BSM candidates
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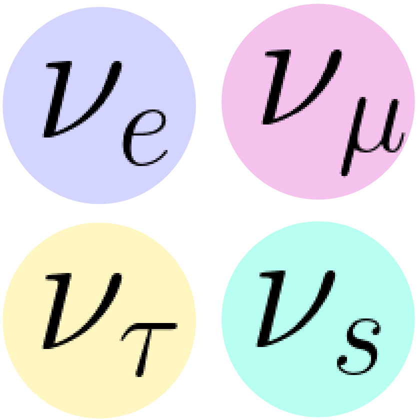
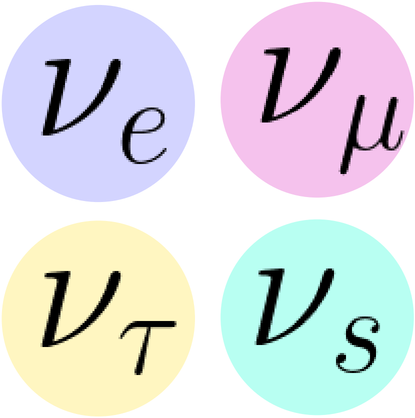


Conclusions

- Sterile neutrinos are attractive BSM candidates
- The canonical DW mechanism is mostly excluded
- Self-interactions among the sterile neutrinos can open new portions of parameter space



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Thank you for your attention

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Backup

The rates

The rates are computed in the following way

$$\Gamma_{\nu_s\nu_s\leftrightarrow kk} = \frac{1}{2E_1}\int \mathrm{d}\Pi_2\mathrm{d}\Pi_3\mathrm{d}\Pi_4(2\pi)^4\delta^4(p_i - p_f)|\mathcal{M}|^2_{\nu_s\nu_s\leftrightarrow kk}|f_s$$

The nine integrals can usually be reduced to two integrals that must be solved numerically

$$\Gamma_{\nu_s\nu_s\leftrightarrow\nu_s\nu_s} = \begin{cases} \frac{3y^4T_s^2}{2\pi^3 p}e^{\frac{\mu}{T_s}} + \frac{y^2T_s m_\phi^2}{2\pi p^2}e^{-\frac{m_\phi^2}{4pT_s}+\frac{\mu}{T_s}} & pT_s > \frac{3m_\phi^2}{2\sqrt{10}} \\ \frac{20y^4pT_s^4}{3\pi^3 m_\phi^4}e^{\frac{\mu}{T_s}} + \frac{y^2T_s m_\phi^2}{2\pi p^2}e^{-\frac{m_\phi^2}{4pT_s}+\frac{\mu}{T_s}} & pT_s \leq \frac{3m_\phi^2}{2\sqrt{10}} \end{cases}$$

The potential

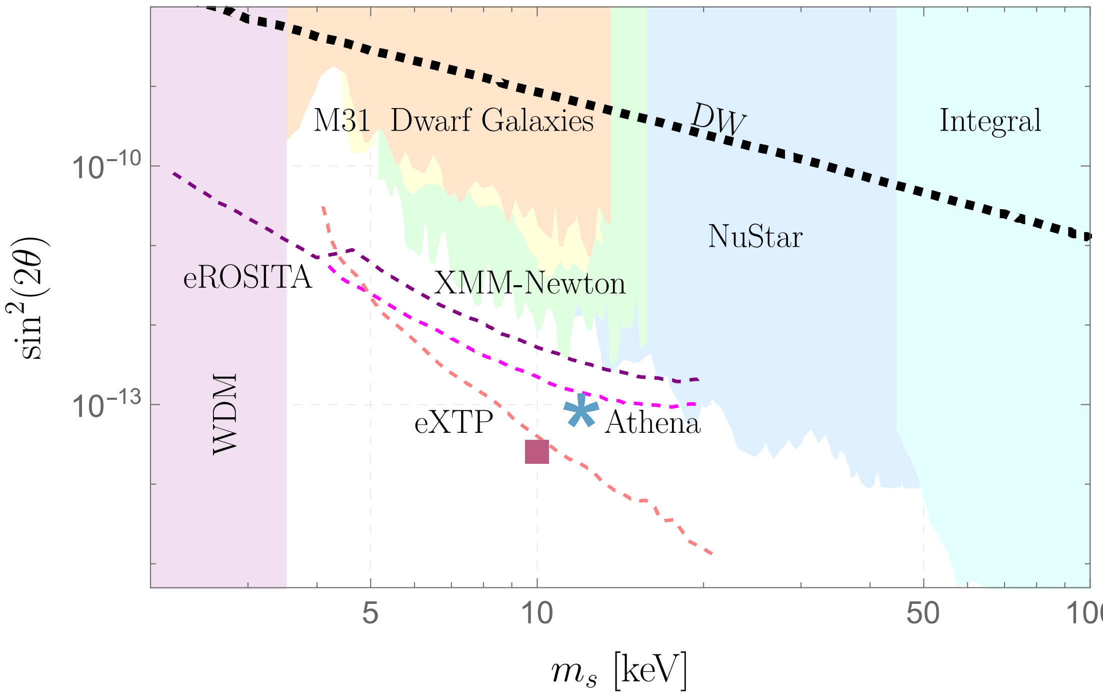
The potential modifies the dispersion relation of the SN

$$E = |\mathbf{p}| + \frac{m^2}{2|\mathbf{p}|} + V_{\text{eff}}$$

In the limiting cases the potential takes the form

$$V_s(p) = \begin{cases} \frac{y^2 T_s^2}{2\pi^2 p} e^{\frac{\mu_s}{T_s}} & T_s \gg m_\phi \\ -\frac{16 y^2 p T_s^4}{\pi^2 m_\phi^4} e^{\frac{\mu_s}{T_s}} & T_s \ll m_\phi \end{cases}$$

The parameter space



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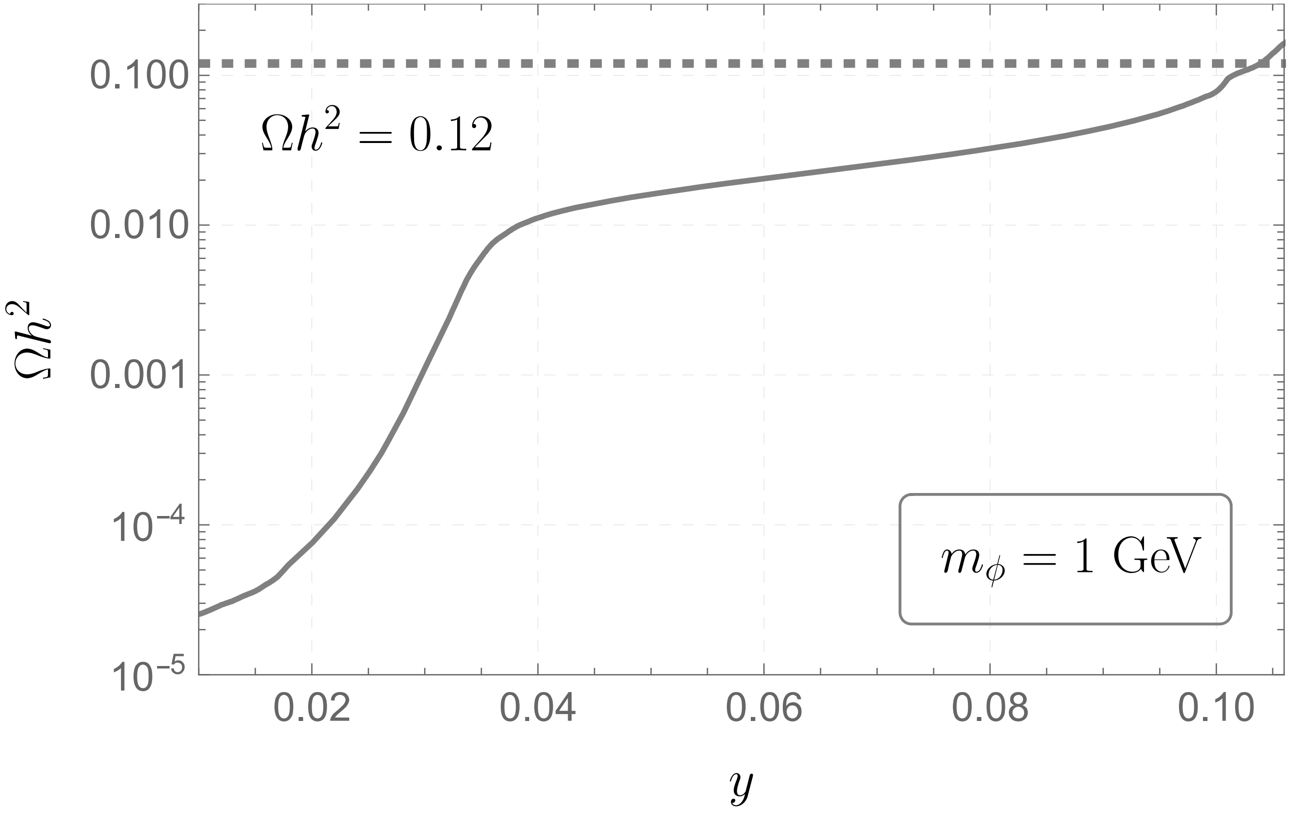

→ A heavy mediator implies a relatively large Yukawa

Observational constraints

X-ray searches

- The easiest observable decay for the SN is a one-loop process.

$$\Gamma_{\rm decay} \sim \left[\frac{\sin^2(2\theta)}{10^{-7}}\right]\left(\frac{m_s}{1\ \rm keV}\right)^5$$

For keV sterile neutrinos, the resultant photon is in the X-ray band

- There are several searches for this type of decay

Structure formation

- Sterile neutrinos are warm dark matter candidates.

- Structure will be smeared on scales below the free-streaming length of neutrinos.