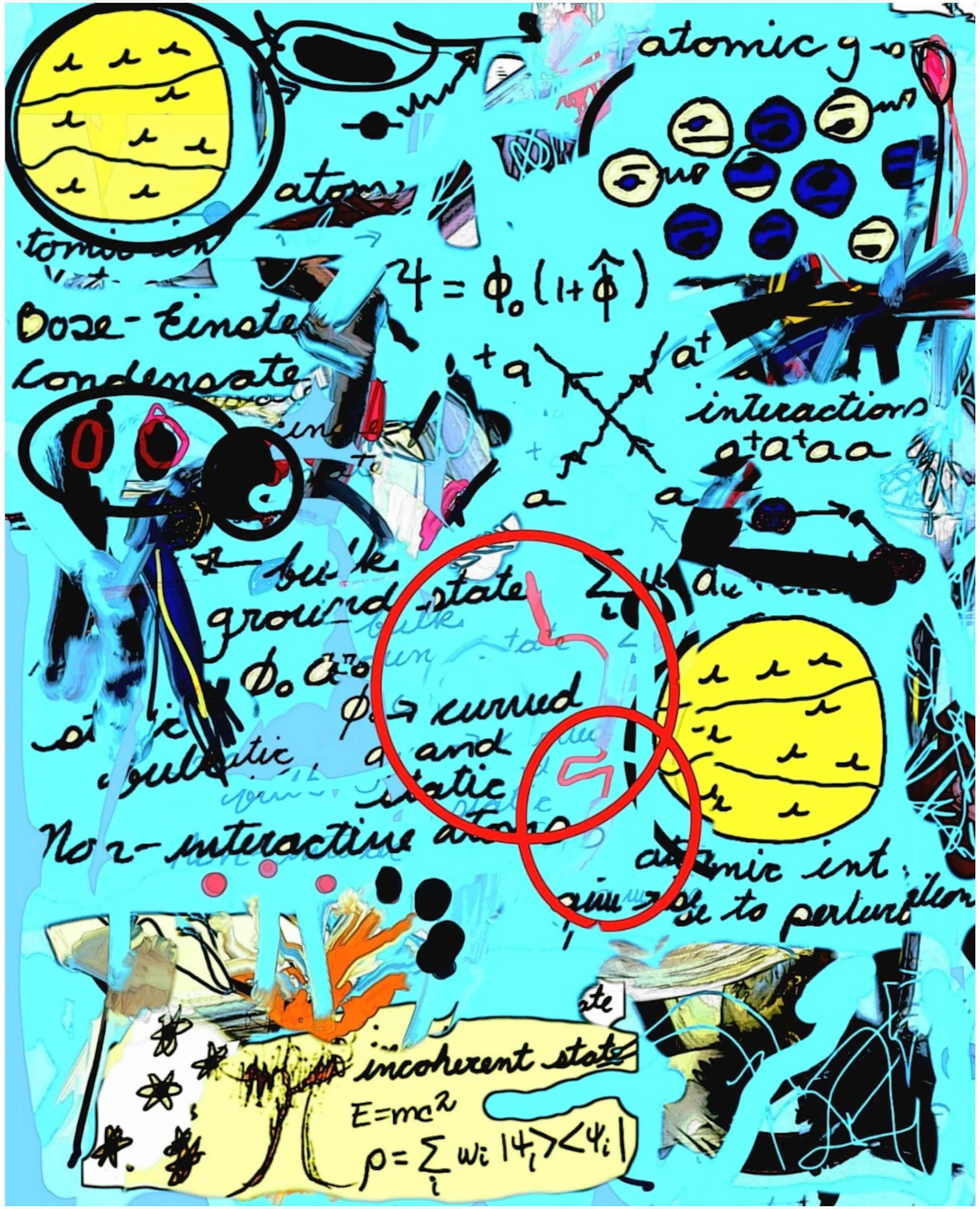
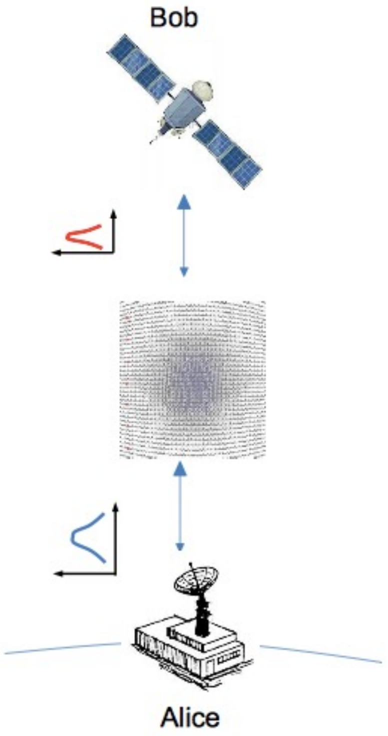


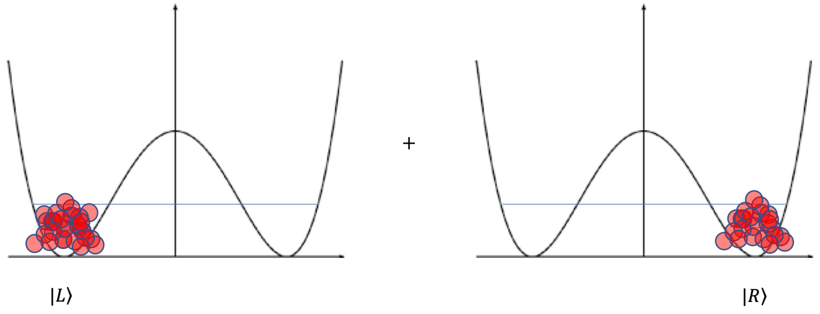
Quantum information and quantum metrology for Fundamental Physics

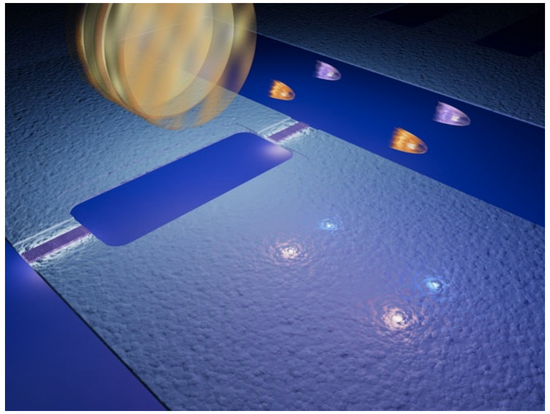
Ivette Fuentes
University of Southampton
Keble College, Oxford

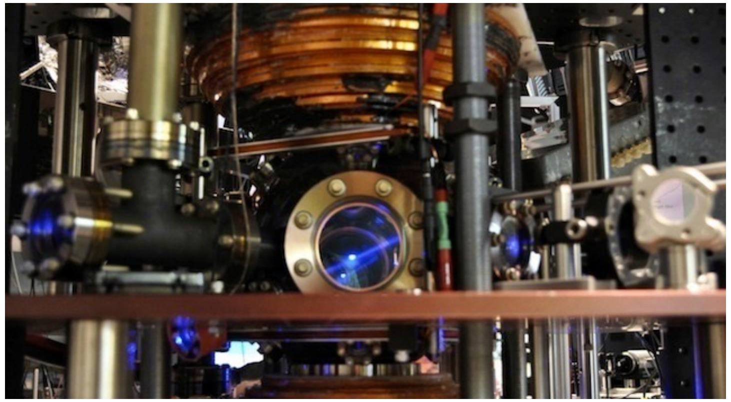
Entanglement and Relativity

How do we quantify entanglement?

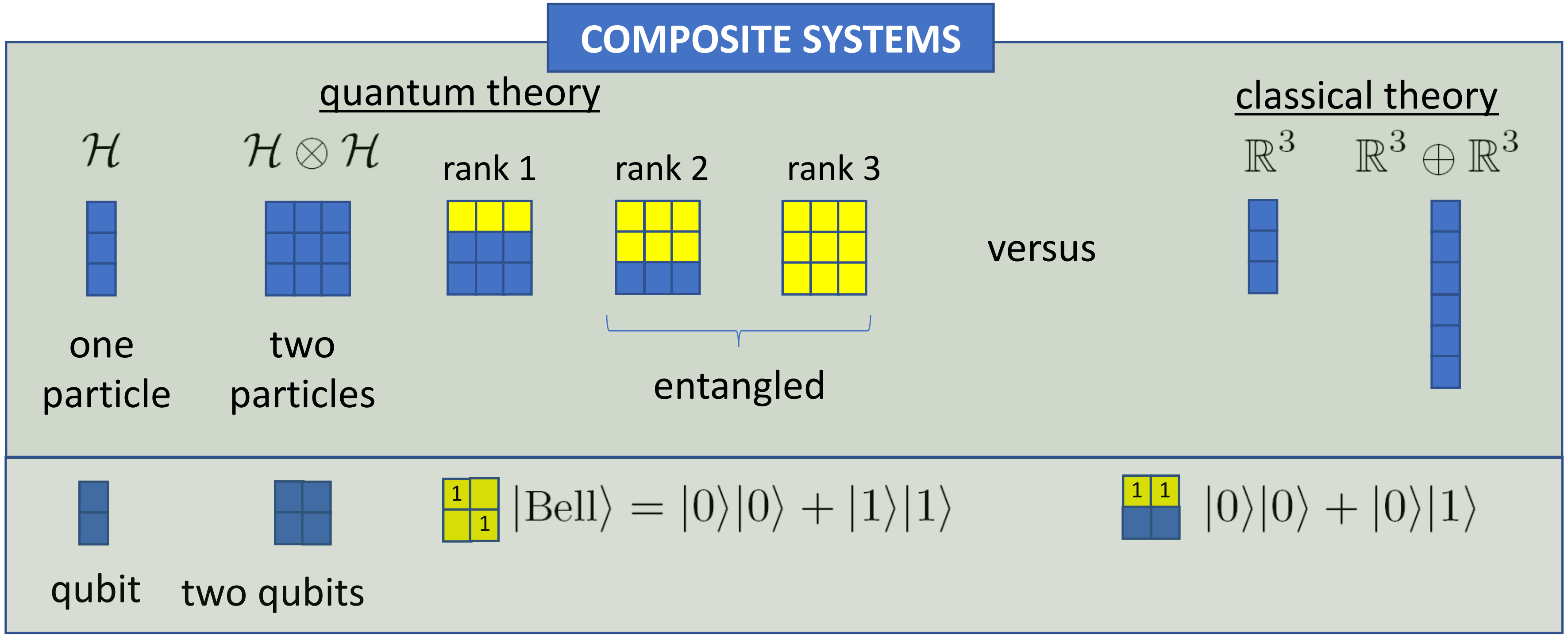








ENTANGLEMENT AND ITS USES



HOW DOES THIS WORK IN A RELATIVISTIC QUANTUM THEORY?

Quantum field theory in curved spacetime

Classical spacetime + quantum fields

Combines QT with GR at low energies scales reachable by cutting-edge experiments

Theoretical predictions

Particle creation by spacetime dynamics

Hawking radiation

Davies-Fulling-Unruh effect

We challenge the belief that only analogue systems can test QFTCS predictions

PARTICLES FROM FIELDS

Quantum field theory on curved spacetime

$\downarrow$

Quantum field fundamental • Particles derived notion (if at all)

QUANTUM FIELD THEORY

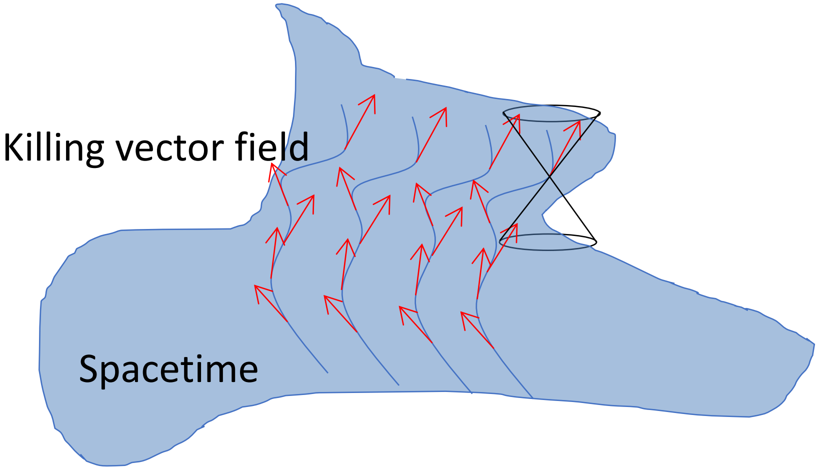
linear field equation $(\Box + m^2)\Phi = 0$

$\downarrow$

vectorspace of solutions $\mathrm{span}\{\tilde{\Phi}_k\}$

$\tilde{\Phi} \in \mathcal{H} \quad (\phi, \psi) = i \int_\Sigma d\Sigma^\mu [\psi^* \partial_\mu \phi - \phi \partial_\mu \psi^*]$

Inner product not positive definite!



PARTICLE INTERPRETATION

requires classification of modes into positive $\leftrightarrow$ negative frequency ω

$$i\partial_T \tilde{\Phi} = \omega \tilde{\Phi} \quad \left[\begin{array}{l} \tilde{\Phi}^* \text{ pos.} \\ \Updownarrow \\ \tilde{\Phi} \text{ neg.} \end{array} \right.$$

$\uparrow$

timelike Killing vector field

$\mathcal{F} = \mathbb{C} \oplus \mathcal{H} \oplus \mathcal{H} \circ \mathcal{H} \oplus \ldots \oplus \mathcal{H}^n$

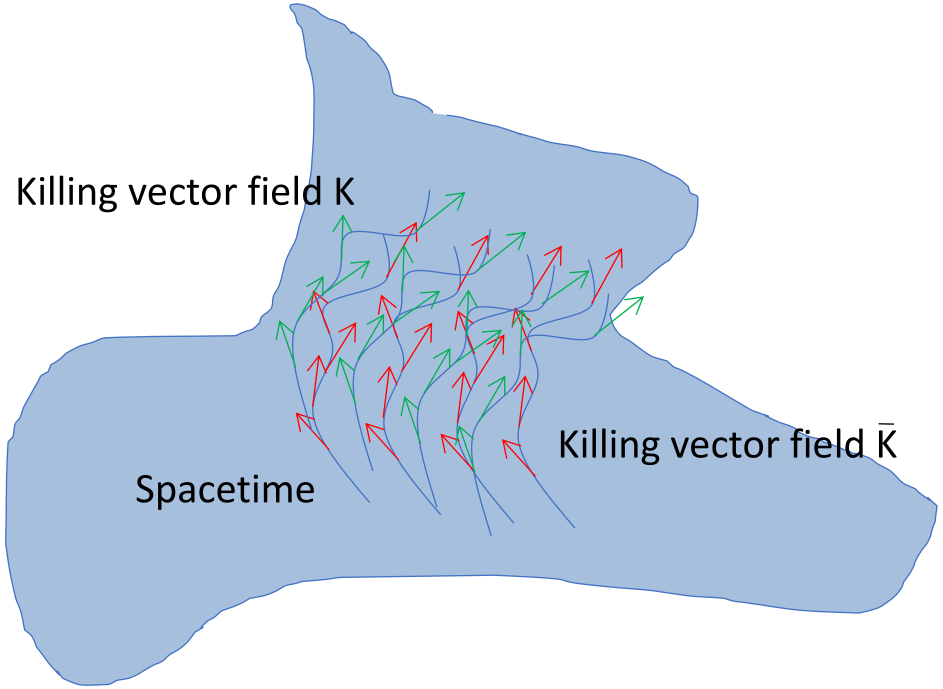
boson Fock space

$$\phi(x) = \int dk \left[\underset{\substack{\uparrow \\ \text{creation}}}{a_k^+} \tilde{\Phi}_k + \underset{\substack{\uparrow \\ \text{annihilation}}}{a_k^-} \tilde{\Phi}_k^* \right]$$

KILLING OBSERVERS

INSIGHTS

- particles present ill-defined subsystems!
- particles well-defined only for killing observers
- particle interpretation may change with change of Killing vector field



KILLING OBSERVERS

different timelike Killing vectors K and $\bar{K}$

$\Rightarrow$ different splits of basis in pos/neg

$$\{u_p, u_p^*\} \longrightarrow \{\bar{u}_p, \bar{u}_p^*\}$$

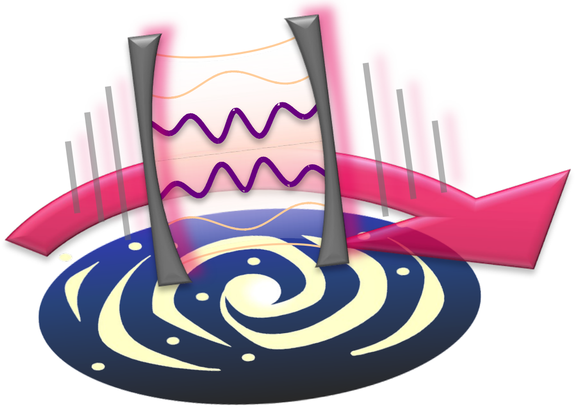
Bogoliubov transformation

$$\bar{a}_p = \int_{q\in\mathcal{P}} [\alpha_{pq}^* a_q - \beta_{pq}^* a_q^\dagger],$$

Squeezed states

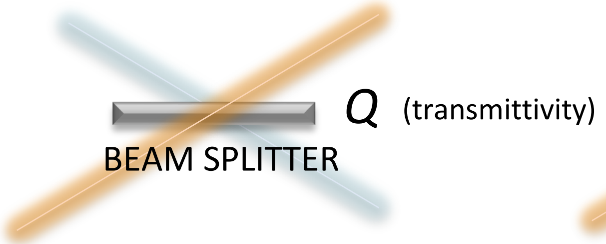
$$|0\rangle = e^{\sum_{i\neq j}\gamma_{ij}a_i^\dagger a_j^\dagger - \gamma_{ij}^* a_i a_j}|\bar{0}\rangle$$

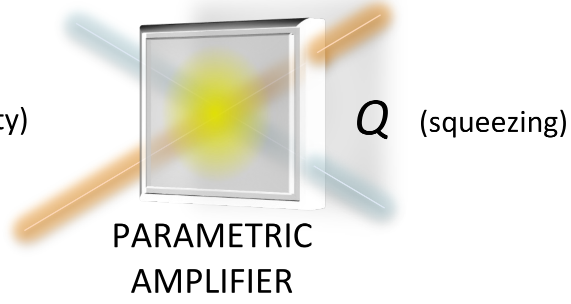
Bogoliubov transformations



Bogoliubov transformation

- Realizes a linear transformation of the modes: $\tilde{a}_m = \sum_n \left(\alpha_{mn} a_n + \beta^*_{mn} a^\dagger_n\right)$
- *Alphas*: passive terms (beam-splitter like)
- *Betas*: active terms (two-mode squeezers)

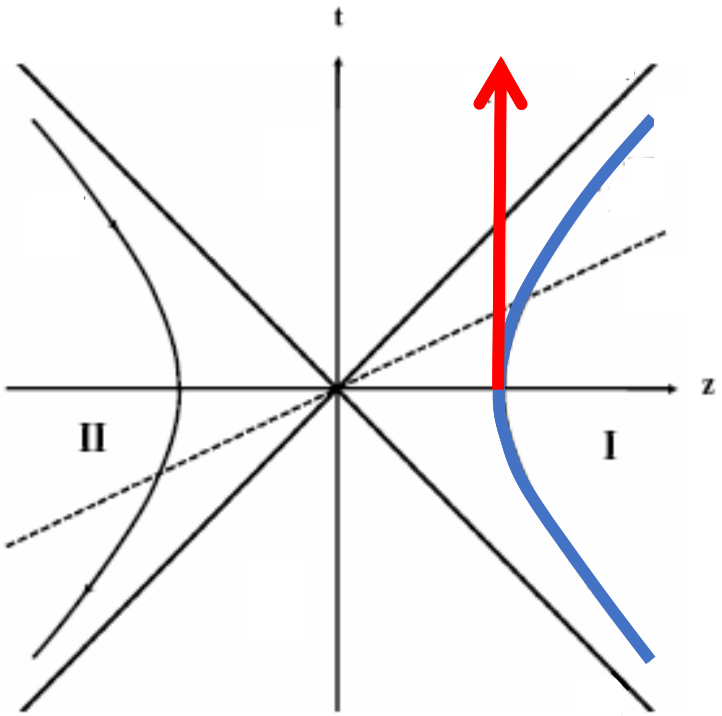




Examples: change of observer, space-time dynamics, moving cavity

EXAMPLE: UNRUH EFFECT

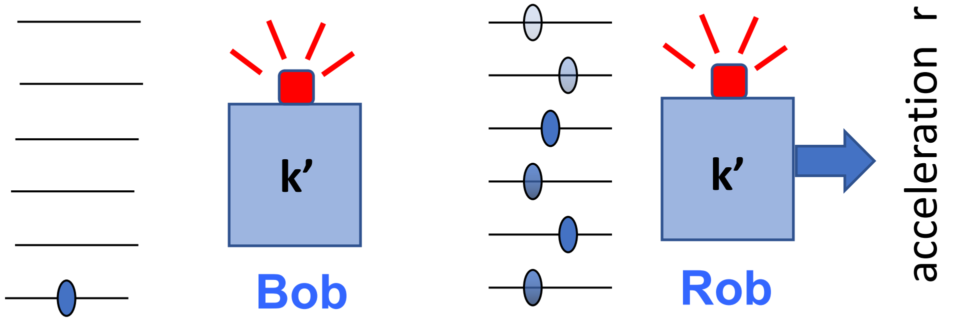
Minkowski spacetime in 1+1 dimensions (flat spacetime = no gravity!)



$$|0_k\rangle^{\mathcal{M}} \sim \frac{1}{\cosh r}\sum_{n=0}^{\infty} \tanh^n r |n_k\rangle_I |n_k\rangle_{II},$$

$$\cosh r = (1 - e^{-2\pi\Omega})^{-1/2}, \qquad \Omega = |k|c/a$$

Rob is causally disconnected from region II



trace $\Rightarrow$ thermal state

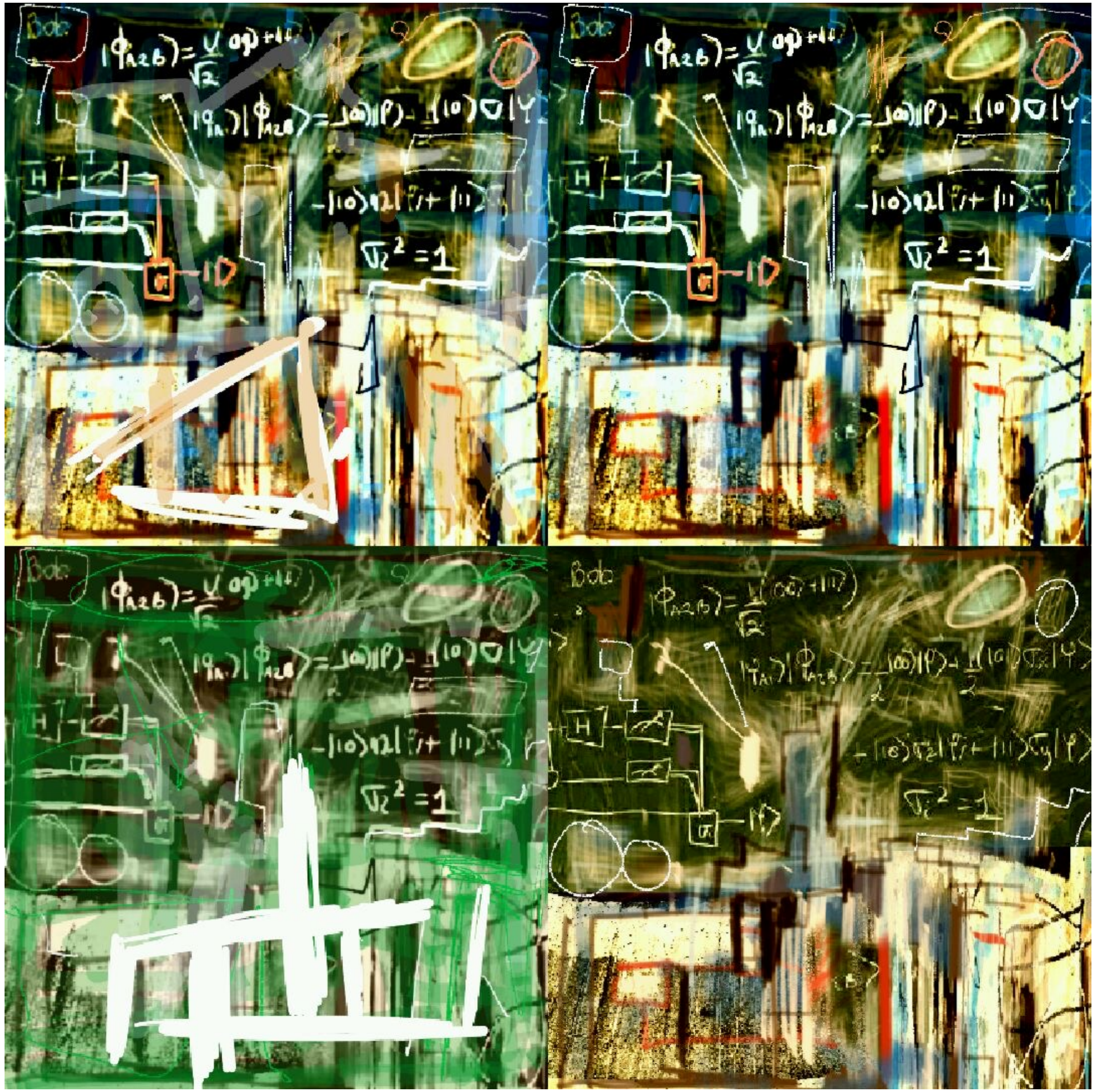
Timelike killing observers

(a) inertial observer

(b) uniformly accelerated observers

Similar effect in black holes: Hawking radiation

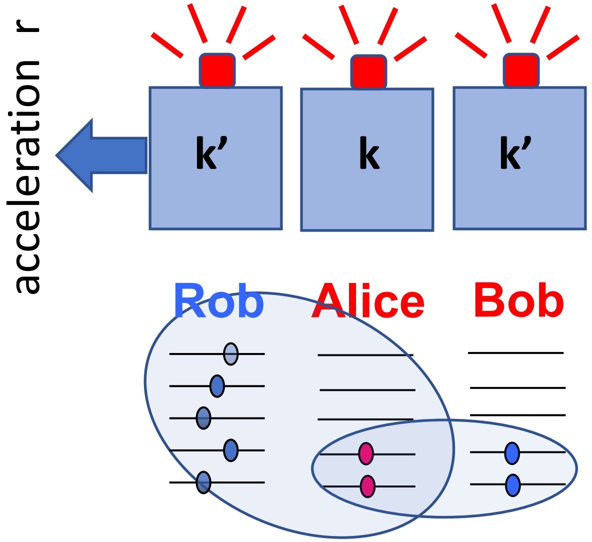
some results



FLAT SPACETIME

Alice and Rob

I. Fuentes-Schuller & R. Mann PRL (2005)



THEORETICAL PHYSICS

To Escape From Quantum Wierdness, Put the Pedal to the Metal

A. Cho, Science (2005)

Entanglement • observer-dependent
• degrades with acceleration , vanishes for ∞ acceleration

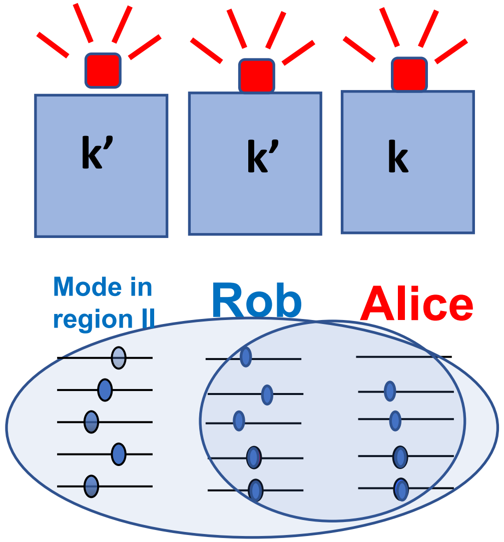
Entanglement sharing

Where did the lost entanglement between Alice and Bob go?

$$|0_k\rangle^{\mathcal{M}} \sim \frac{1}{\cosh r} \sum_{n=0}^{\infty} \tanh^n r |n_k\rangle_I |n_k\rangle_{II},$$

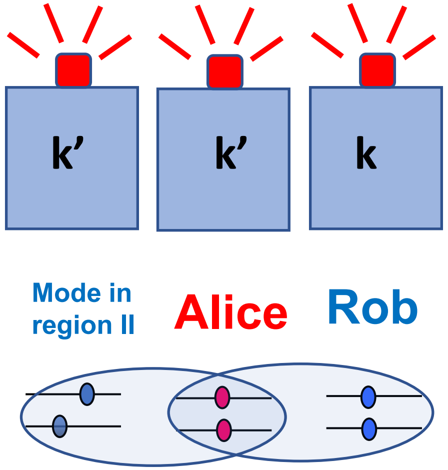
Alsing, Fuentes-S, Mann, Tessier PRA 2006
Adesso, Fuentes-S , Ericsson PRA 2007

Bosonic field



multipartite entanglement!

Fermionic field



Bipartite entanglement between Alice and mode in region II

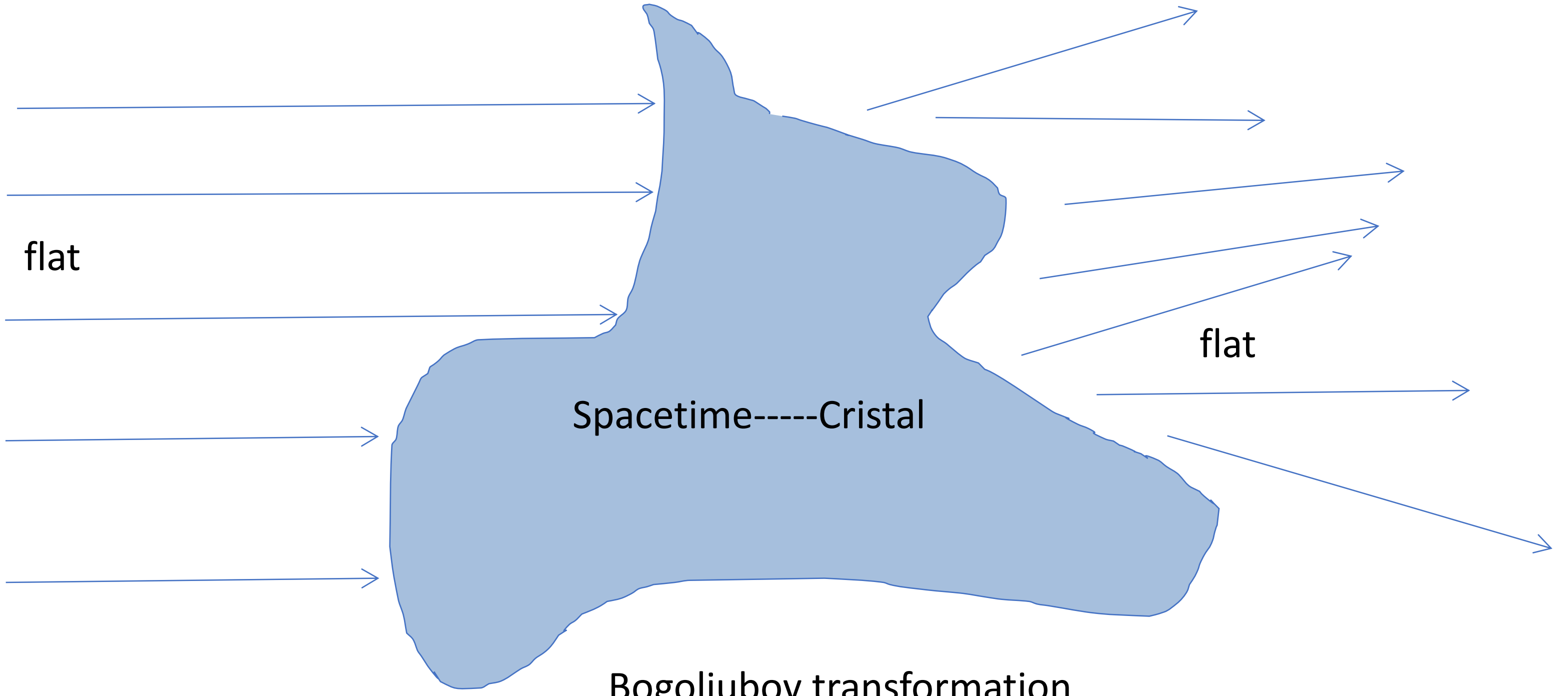
important differences between fermions and bosons.

CURVED SPACETIME

SPACETIME AS A CRISTAL

Curve spacetimes generally do not admit timelike killing vector fields…

particular spacetimes with asymptotically flat regions



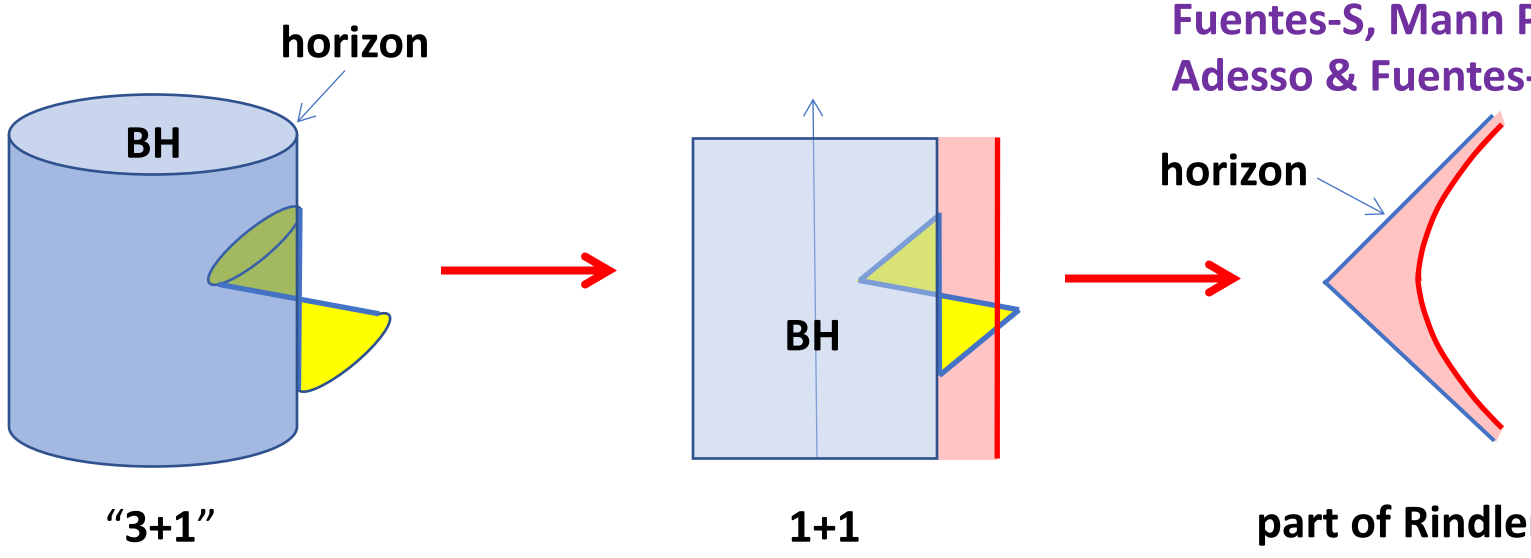
$$\{u_p, u_p^*\} \xrightarrow[\text{Squeezed states}]{\text{Bogoliubov transformation}} \{\bar{u}_p, \bar{u}_p^*\}$$

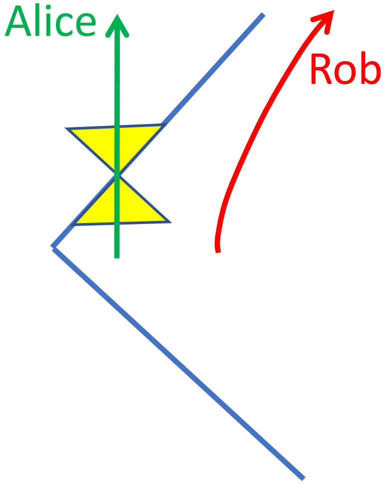
Just like in quantum optics!

BLACK HOLES
&
COSMOLOGY

Alice falls into a black hole

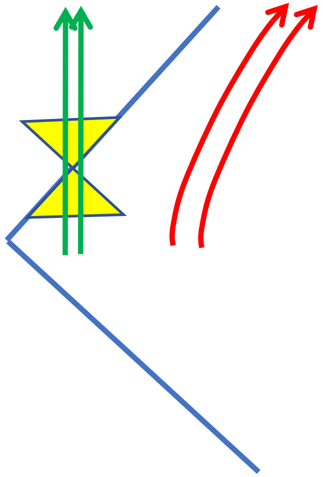
Fuentes-S, Mann PRL 2005
Adesso & Fuentes-S 2007





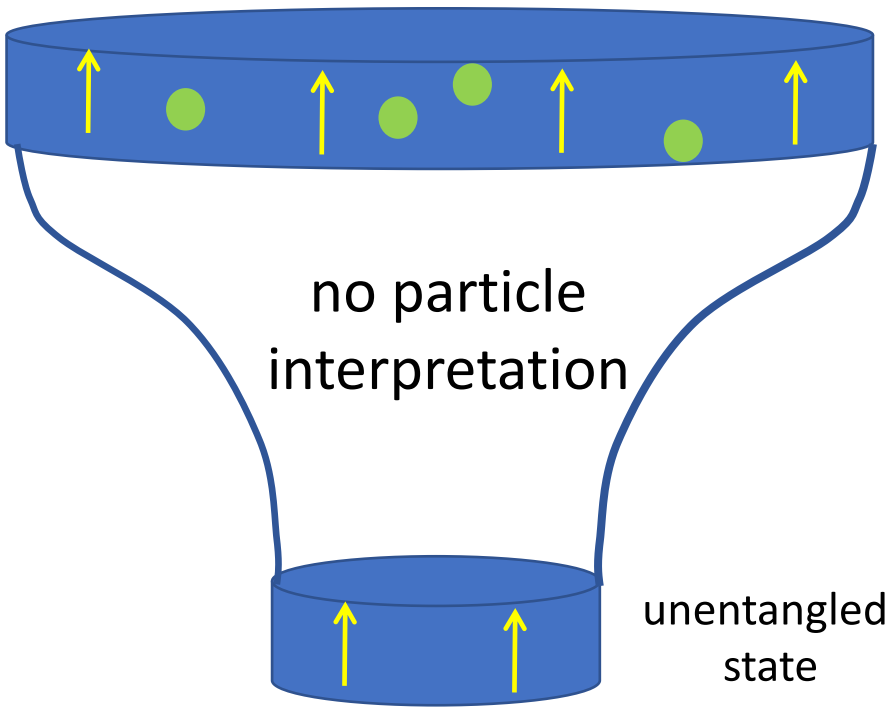
Entanglement
Classical correlations } degraded for escaping observers

Lost entanglement → multipartite entanglement between modes inside and outside the BH



Entanglement cosmology

Ball, Fuentes-S, Schuller PLA 2006



"History of the universe
encoded in entanglement"

<u>toy model</u>

expansion rate σ

expansion factor ϵ

- calculate entanglement

  asymptotic past $S = 0$

  asymptotic future $S = S(\sigma, \epsilon)$

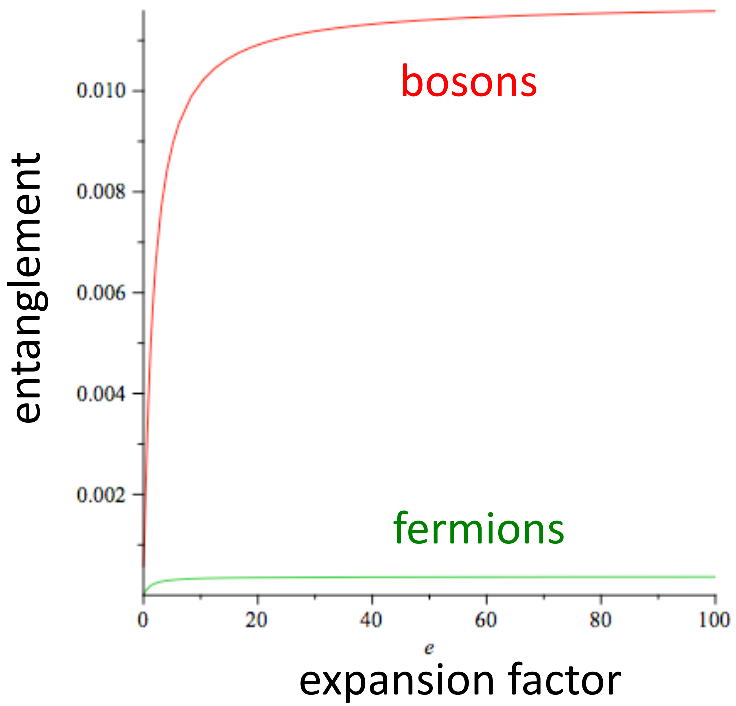
- excitingly, can solve for

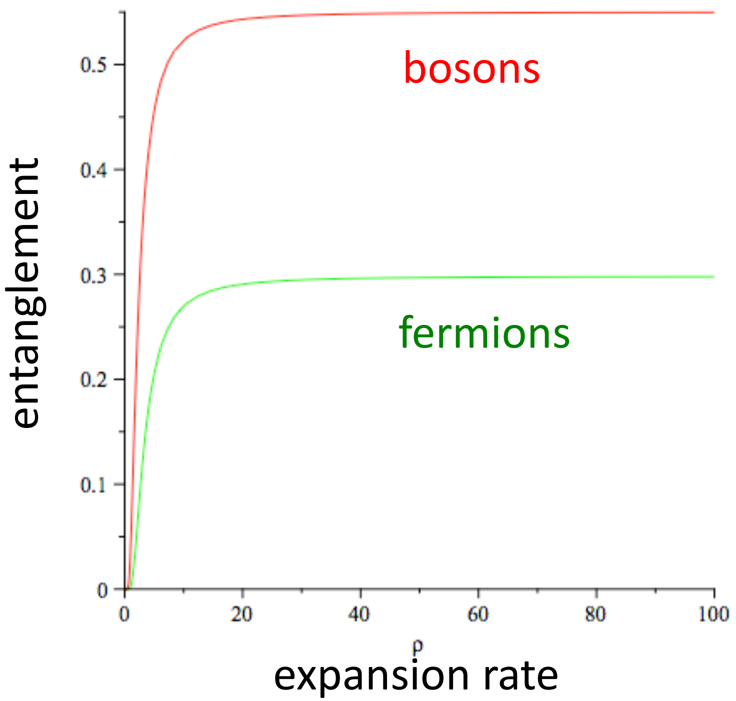
$\sigma = \sigma(S)$ $\epsilon = \epsilon(S)$

Fermionic entanglement cosmology

Fuentes, Mann, Martin-Martinez, Moradi PRD 2010

Fermionic fields in 3+1 dimensions: more realistic model





"The Universe entangles less fermionic fields"

REVIEW

Observer-dependent entanglement

Paul M Alsing[1] and Ivette Fuentes[2]

Published 18 October 2012 •

Classical and Quantum Gravity, Volume 29, Number 22

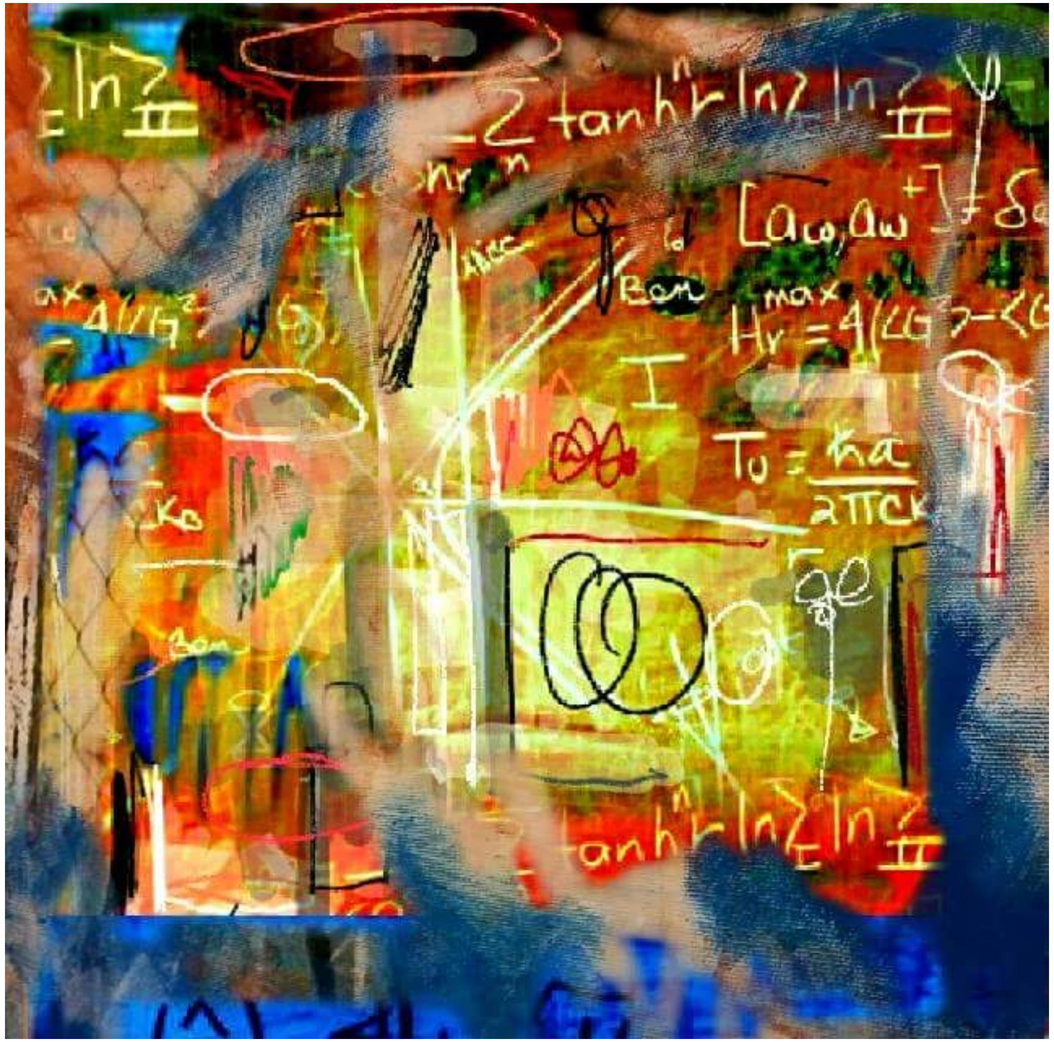
Citation Paul M Alsing and Ivette Fuentes 2012 *Class. Quantum Grav.* **29** 224001

Entanglement is observer dependent

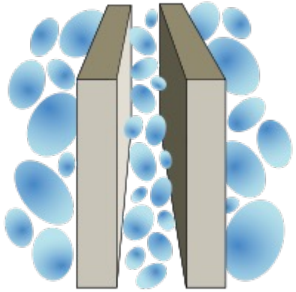
Spacetime dynamics create entanglement

Horizons (black holes) degrade entanglement

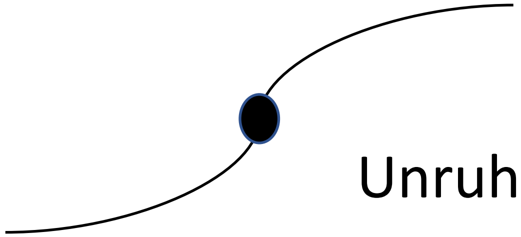
Localized systems for Relativistic Quantum Information



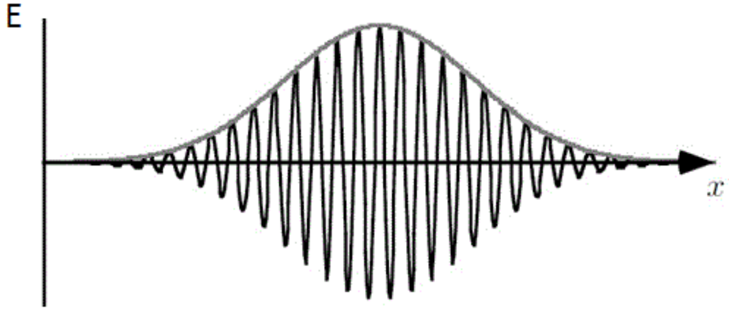
Systems for relativistic quantum information



Cavities
Trapped BECs



Unruh-DeWitt detectors



Traveling wave-packets

1. Entanglement between localized systems
    - cavities
    - detectors
    - localized wave-packets
- gravity effects on quantum properties
- Propose earth-based and space-based experiments

Covariance Matrix Formalism

Convenient for systems consisting of N bosonic modes $\longrightarrow$ N harmonic oscillators

$$\hat{x}_{(2n-1)} = \frac{1}{\sqrt{2}}(\hat{a}_n + \hat{a}_n^\dagger)$$ generalized position

$$\hat{x}_{(2n)} = \frac{-i}{\sqrt{2}}(\hat{a}_n - \hat{a}_n^\dagger)$$ generalized momentum

Collect in vector form: $\hat{X} = \begin{pmatrix} \hat{x}_1 \\ \hat{x}_2 \\ \vdots \\ \hat{x}_{2n} \end{pmatrix}$

Covariance matrix: $\sigma_{ij} = \langle \hat{x}_i \hat{x}_j + \hat{x}_j \hat{x}_i \rangle - 2 \langle \hat{x}_i \rangle \langle \hat{x}_j \rangle$

Displacement vector: $\hat{d} = \begin{pmatrix} \langle \hat{x}_1 \rangle \\ \langle \hat{x}_2 \rangle \\ \vdots \\ \langle \hat{x}_{2n} \rangle \end{pmatrix}$

$\langle \hat{x}_j \rangle \longrightarrow$ First moments

$\langle \hat{x}_i \hat{x}_j \rangle \longrightarrow$ Second moments

For Gaussian states the covariance matrix and displacement vector replace the density matrix

Gaussian states

$|0\rangle$

vacuum

$$|\alpha\rangle = e^{-|\alpha|^2/2} \sum_{n=0}^{\infty} \frac{\alpha^n}{\sqrt{n!}} |n\rangle$$

Coherent states

$$\frac{1}{\cosh^2(r)} \sum_{n=0}^{\infty} \tanh^{2n}(r) |n\rangle\langle n|,$$

Thermal states

$$\frac{1}{\sqrt{\cosh r}} \sum_{n=0}^{\infty} (-e^{i\phi} \tanh r)^n \frac{\sqrt{(2n)!}}{2^n n!} |2n\rangle$$

$$\frac{1}{\cosh(r)} \sum_n \tanh r^n \, |n_a\rangle |n_b\rangle$$

Squeezed states

Gaussian states can be described using the first and second moments.
They are defined as states for which higher moments vanish.

Example: vacuum state

$$|0\rangle = |0_1\rangle\,|0_2\rangle \ldots |0_N\rangle = |0\rangle^{\otimes N}$$

Frist moments:

$$\langle \hat{x}_i \rangle = \langle 0|\, \frac{1}{\sqrt{2}} (\hat{a}_i + \hat{a}_i^\dagger)\, |0\rangle = 0$$

Displacement vector: $\hat{d} = \begin{pmatrix} 0 \\ \vdots \\ 0 \end{pmatrix}$

Second moments:

$$\left\langle \hat{x}_i \hat{x}_j \right\rangle = \frac{1}{2} \langle 0|\, (\hat{a}_i \hat{a}_j^\dagger + \hat{a}_i^\dagger \hat{a}_j)\, |0\rangle = \delta_{ij}$$

Covariance matrix:

$$\sigma = \begin{pmatrix} 1 & \cdots & 0 \\ \vdots & \ddots & \vdots \\ 0 & \cdots & 1 \end{pmatrix} = \mathrm{I}$$

Example: single-mode squeezed coherent state

$$\mathrm{U}(r) = e^{\frac{r}{2}(\hat{a}^2 - \hat{a}^{\dagger 2})}$$

$$D(\alpha) = e^{\,(\alpha\hat{a} - \alpha^{*}\hat{a}^{\dagger})}$$

$$|\phi(r,\alpha)\rangle = U(r)D(\alpha)\,|0\rangle$$

Displacement vector: $\hat{d} = \begin{pmatrix} \alpha \\ \alpha^{*} \end{pmatrix}$

$$\mathcal{S}(r,\alpha) = U(r)D(\alpha)$$

Second moments:

$$\langle \hat{x}^2 \rangle = \langle \hat{p}^2 \rangle = \langle \phi|\,(\hat{a}\hat{a}^{\dagger} + \hat{a}^{\dagger}\hat{a})\,|\phi\rangle = \cosh(2r)$$

$$\langle \hat{x}\hat{p} \rangle = \langle \hat{p}\hat{x} \rangle = \frac{1}{2}\langle \phi|\,(\hat{a}\hat{a}^{\dagger} + \hat{a}^{\dagger}\hat{a})\,|\phi\rangle = -\sinh(2r)$$

Covariance matrix: $\sigma = \begin{pmatrix} \cosh(2r) & -\sinh(2r) \\ -\sinh(2r) & \cosh(2r) \end{pmatrix}$

Separable state: $S_a\,(r_a, \alpha_a) \otimes S_b\,(r_b, \alpha_b)|00\rangle$

Example: two-mode squeezed state

$$|\psi_{ab}\rangle = \frac{1}{\cosh(r)}\sum_n tanh^n(r)|n_a\rangle|n_b\rangle$$

Frist moments:

$$\langle\hat{x}_a\rangle = \left\langle\psi_{ab}\right| \frac{1}{\sqrt{2}}(\hat{a}_a + \hat{a}_a^\dagger)\left|\psi_{ab}\right\rangle = 0 = \langle\hat{x}_b\rangle$$

Second moments:

$$\langle\hat{x}_a\hat{x}_a\rangle = \frac{1}{2}\left\langle\psi_{ab}\right| (\hat{a}_a\hat{a}_a^\dagger + \hat{a}_a^\dagger\hat{a}_a)\left|\psi_{ab}\right\rangle = \cosh(2r)$$

Displacement vector: $\hat{d} = \begin{pmatrix} 0 \\ 0 \\ 0 \\ 0 \end{pmatrix}$

Covariance matrix:

$$\sigma = \begin{pmatrix} \cosh(2r) & 0 & \sinh(2r) & 0 \\ 0 & \cosh(2r) & 0 & -\sinh(2r) \\ \sinh(2r) & 0 & \cosh(2r) & 0 \\ 0 & -\sinh(2r) & 0 & \cosh(2r) \end{pmatrix}$$

State Evolution

$\sigma_f = S^T \sigma_i \, S$

S is the symplectic matrix
It plays an important role in QFT

$S^T \Omega S = \Omega$

symplectic form

$$\Omega = \begin{pmatrix} 0 & 1 & 0 & 0 \\ -1 & 0 & 0 & 0 \\ 0 & 0 & 0 & 1 \\ 0 & 0 & -1 & 0 \end{pmatrix}$$

$2i\Omega_{ij} = [\hat{x}_i, \hat{x}_j]$

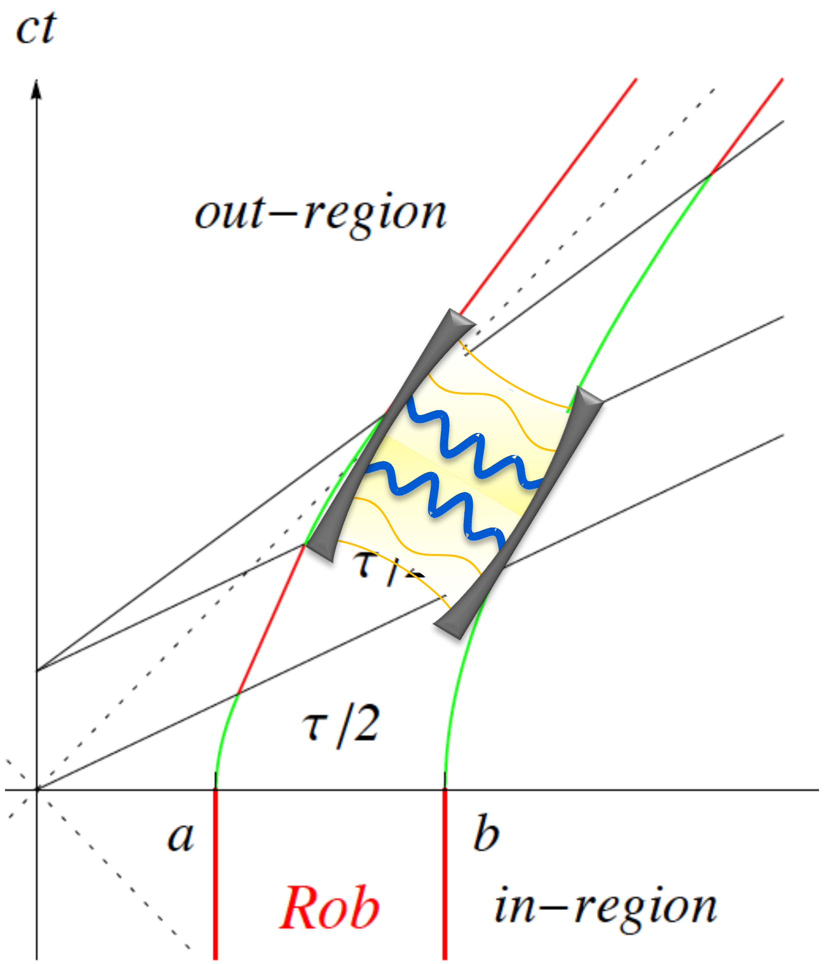
There are computable measures of entanglement if the states are Gaussian

Example:

Negativity $N(\nu_i)$ $\quad$ ν_i is the smallest eigenvalue of $\quad i\Omega\sigma^{PT}$

Partial transpose

QFT in the symplectic formalism



Articles

Entanglement generation in relativistic quantum fields

Nicolai Friis & Ivette Fuentes

Pages 22-27 | Received 06 Apr 2012, Accepted 10 Jul 2012, Published online: 20 Aug 2012

Download citation https://doi.org/10.1080/09500340.2012.712725

$$S = \begin{pmatrix} \mathcal{M}_{11} & \mathcal{M}_{12} & \mathcal{M}_{13} & \dots \\ \mathcal{M}_{21} & \mathcal{M}_{22} & \mathcal{M}_{23} & \dots \\ \mathcal{M}_{31} & \mathcal{M}_{32} & \mathcal{M}_{33} & \dots \\ \vdots & \vdots & \vdots & \ddots \end{pmatrix}$$

$$\mathcal{M}_{mn} = \begin{pmatrix} \Re(\alpha_{mn} - \beta_{mn}) & \Im(\alpha_{mn} + \beta_{mn}) \\ -\Im(\alpha_{mn} - \beta_{mn}) & \Re(\alpha_{mn} + \beta_{mn}) \end{pmatrix}$$

Example: inertial cavity

Minkowski coordinates (t, x)

$$\Box\phi(t,x) = 0$$ field equation

solutions: plane waves+ boundary

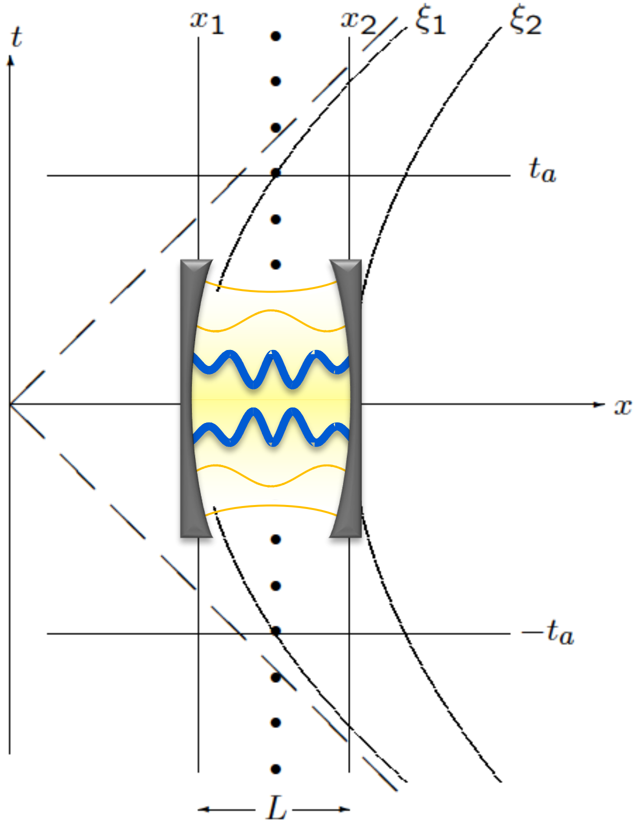
$$u_k(x,t) = \frac{1}{\sqrt{k\pi}} \sin\left[\frac{k\pi}{L}(x - x_A)\right] e^{-i\omega_k t},$$

$$\omega_k = \frac{1}{L}\sqrt{(k\pi)^2 + m^2},$$ discrete spectrum

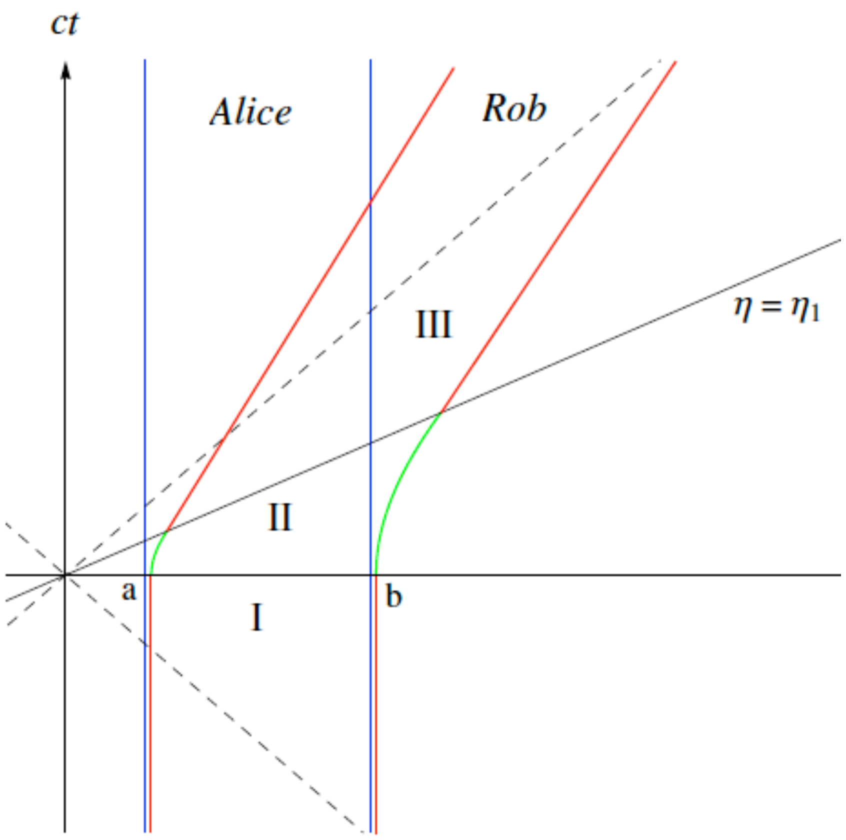
creation and annihilation operators

$$\hat{\phi}(x,t) = \sum_n (u_n(t,x)\hat{a}_n + u_n^*(t,x)\hat{a}_n^\dagger)$$

$\hat{a}_n|0\rangle = 0$ vacuum state $\qquad |m_n\rangle = \hat{a}_n^{\dagger\, m}|0\rangle$ particle states



Non-uniform motion



Voyage to Alpha Centauri: Entanglement degradation of cavity modes due to motion

David Edward Bruschi, Ivette Fuentes, and Jorma Louko
Phys. Rev. D **85**, 061701(R) – Published 15 March 2012

Bogoliubov transformations

$$\tilde{a}_m = \sum_n \left(\alpha^*_{mn}\, a_n - \beta^*_{mn}\, a^\dagger_n \right)$$

$$\alpha_{mn} = (\tilde{\phi}_m, \phi_n) \text{ and } \beta_{mn} = -(\tilde{\phi}_n, \phi^*_m)$$

Entangling Moving Cavities in Noninertial Frames

T. G. Downes, I. Fuentes, and T. C. Ralph
Phys. Rev. Lett. **106**, 210502 – Published 25 May 2011

Bogoliubov transformations

$$\begin{pmatrix} \hat{x}'_1 \\ \vdots \\ \hat{x}'_{2n} \end{pmatrix} = \begin{pmatrix} \Re(\alpha_{11} - \beta_{11}) & \cdots & \Im(\alpha_{1n} + \beta_{1n}) \\ \vdots & \ddots & \vdots \\ -\Im(\alpha_{n1} + \beta_{n1}) & \cdots & \Re(\alpha_{nn} + \beta_{nn}) \end{pmatrix} \begin{pmatrix} \hat{x}_1 \\ \vdots \\ \hat{x}_{2n} \end{pmatrix}$$

$$\alpha_{nm} = (u'_n, u_m) \qquad \beta_{nm} = (u'_n, u^*_m)$$

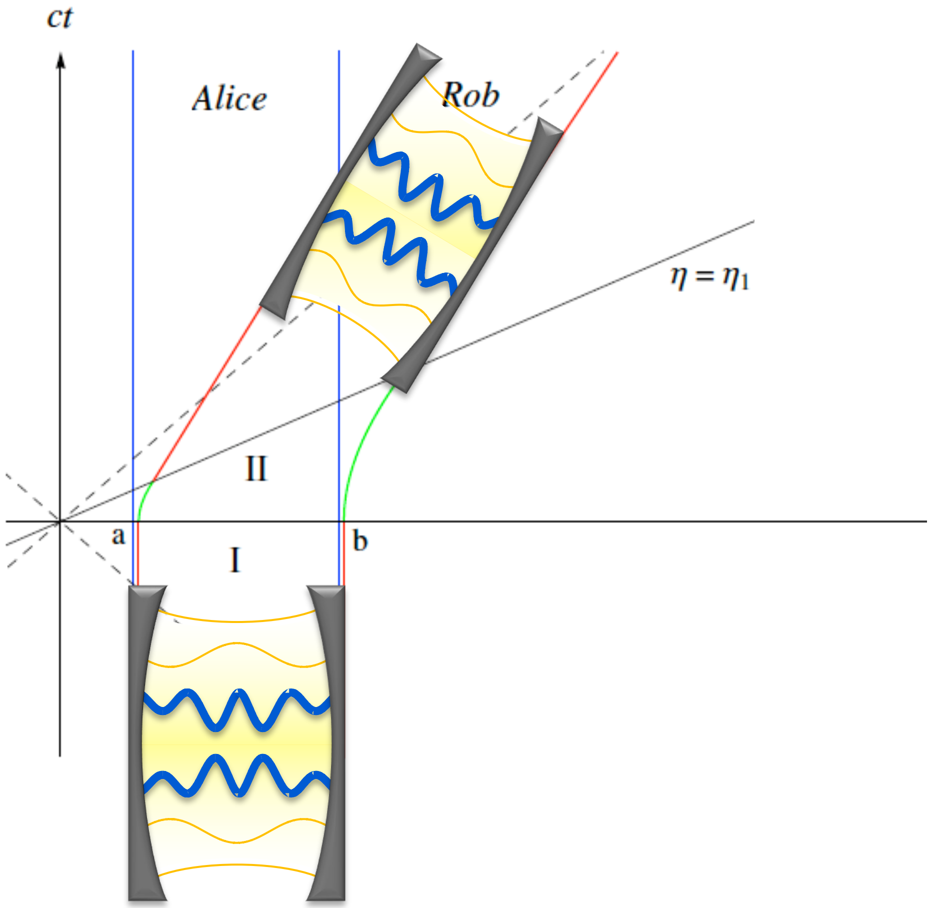
$$(\phi, \psi) = i \int_\Sigma d\Sigma^\mu [\psi^* \partial_\mu \phi - \phi \partial_\mu \psi^*]$$

Can't be computed but…

$$S = \begin{pmatrix} \mathcal{M}_{11} & \mathcal{M}_{12} & \mathcal{M}_{13} & \ldots \\ \mathcal{M}_{21} & \mathcal{M}_{22} & \mathcal{M}_{23} & \ldots \\ \mathcal{M}_{31} & \mathcal{M}_{32} & \mathcal{M}_{33} & \ldots \\ \vdots & \vdots & \vdots & \ddots \end{pmatrix}$$

$$\mathcal{M}_{mn} = \begin{pmatrix} \Re(\alpha_{mn} - \beta_{mn}) & \Im(\alpha_{mn} + \beta_{mn}) \\ -\Im(\alpha_{mn} - \beta_{mn}) & \Re(\alpha_{mn} + \beta_{mn}) \end{pmatrix}$$

Approximation…



$$h = aL$$

acceleration length

$$\alpha = \alpha^{(0)} + \alpha^{(1)} + \alpha^{(2)} + O(h^3)\,,$$

$$\beta = \beta^{(1)} + \beta^{(2)} + O(h^3)\,,$$

$h \ll 1$ computable transformations

During periods of inertial or uniformly accelerated motion

$$T_t = \left(\begin{array}{c|c} \begin{matrix} e^{i\omega_1 t} & \cdots & 0 \\ \vdots & \ddots & \vdots \\ 0 & \cdots & e^{i\omega_n t} \end{matrix} & \begin{matrix} 0 & \cdots & 0 \\ \vdots & \ddots & \vdots \\ 0 & \cdots & 0 \end{matrix} \\ \hline \begin{matrix} 0 & \cdots & 0 \\ \vdots & \ddots & \vdots \\ 0 & \cdots & 0 \end{matrix} & \begin{matrix} e^{-i\omega_1 t} & \cdots & 0 \\ \vdots & \ddots & \vdots \\ 0 & \cdots & e^{-i\omega_n t} \end{matrix} \end{array}\right)$$

$$T_\tau = \left(\begin{array}{c|c} \begin{matrix} e^{i\omega_1 \tau} & \cdots & 0 \\ \vdots & \ddots & \vdots \\ 0 & \cdots & e^{i\omega_n \tau} \end{matrix} & \begin{matrix} 0 & \cdots & 0 \\ \vdots & \ddots & \vdots \\ 0 & \cdots & 0 \end{matrix} \\ \hline \begin{matrix} 0 & \cdots & 0 \\ \vdots & \ddots & \vdots \\ 0 & \cdots & 0 \end{matrix} & \begin{matrix} e^{-i\omega_1 \tau} & \cdots & 0 \\ \vdots & \ddots & \vdots \\ 0 & \cdots & e^{-i\omega_n \tau} \end{matrix} \end{array}\right)$$

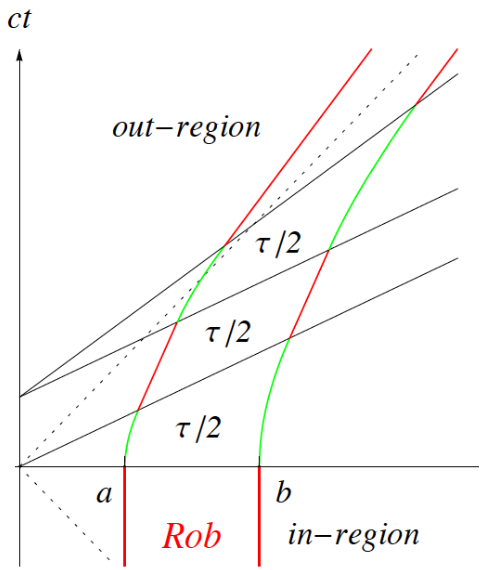
Consider the vacuum as the initial state in region 1

Region I $\quad \sigma_0 = I$

Region II $\quad \sigma_{II} = S^T I S$

Region III $\quad \sigma_{II} = S^{-1T} T_\tau^T S^T S T_\tau S^{-1}$

Basic building block



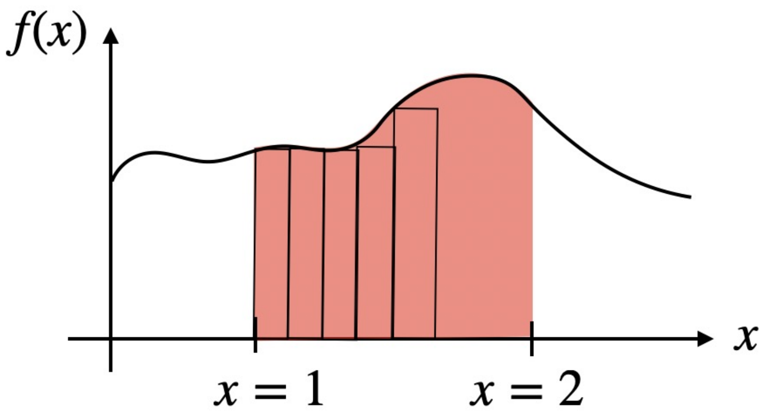
$$B = T_t S T_\tau S^{-1}$$

$$\sigma_B = B^T B$$

General trajectory: grafting $B_F = B_1\, B_2 \ldots B_n$

We can approximate any trajectory

“Integration” method for quantum fields



$$\hat{A}_{mn} = i(\omega_m - \omega_n)\hat{\alpha}_{mn} \int_{\tau_0}^{\tau_f} e^{-i(\omega_m - \omega_n)(\tau - \tau_0)}\, h(\tau)\, d\tau \ ,$$

$$\hat{B}_{mn} = i(\omega_m + \omega_n)\hat{\beta}_{mn} \int_{\tau_0}^{\tau_f} e^{-i(\omega_m + \omega_n)(\tau - \tau_0)}\, h(\tau)\, d\tau \ .$$

Regular Article - Theoretical Physics | Open Access | Published: 31 August 2020

Evolution of confined quantum scalar fields in curved spacetime. Part I

Spacetimes without boundaries or with static boundaries in a synchronous gauge

Luis C. Barbado, Ana L. Báez-Camargo & Ivette Fuentes

The European Physical Journal C **80**, Article number: 796 (2020) | Cite this article

Regular Article - Theoretical Physics | Open Access | Published: 29 October 2021

Evolution of confined quantum scalar fields in curved spacetime. Part II

Spacetimes with moving boundaries in any synchronous gauge

Luis C. Barbado, Ana L. Báez-Camargo & Ivette Fuentes

The European Physical Journal C **81**, Article number: 953 (2021) | Cite this article

Entanglement

Computable measures for Gaussian states

Separable initial states of two modes $S_a\,(s) \otimes S_b\,(s)|00\rangle$

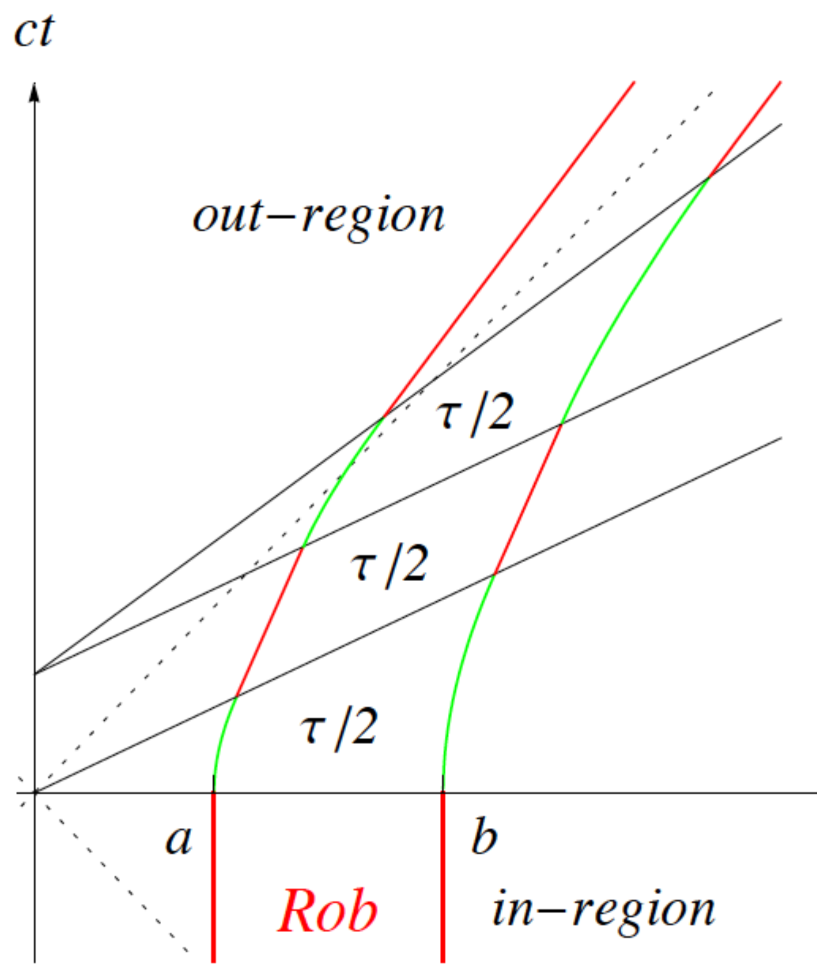
Negativity $N(\nu_i)$ ν_i is the smallest eigenvalue of $i\Omega\sigma^{PT}$

$$\mathcal{N} = (\Re\left(e^{i\omega_a t}\beta_{ab}\right)^2 + (\Im\left(e^{i\omega_a t}\beta_{ab}\right)\cosh(s) - \Im\left(e^{i\omega_a t}\alpha_{ab}\right)\sinh(s))^{1/2}$$

$$\alpha = \alpha^{(0)} + \alpha^{(1)} + \alpha^{(2)} + O(h^3)\,,$$
$$\beta = \beta^{(1)} + \beta^{(2)} + O(h^3)\,,$$

Friis and Fuentes JMO (invited) 2012

entanglement generated



Friis, Bruschi, Louko & Fuentes PRD 2012
Friis and Fuentes invited at JMO 2012
Bruschi, Louko, Faccio & Fuentes 2012

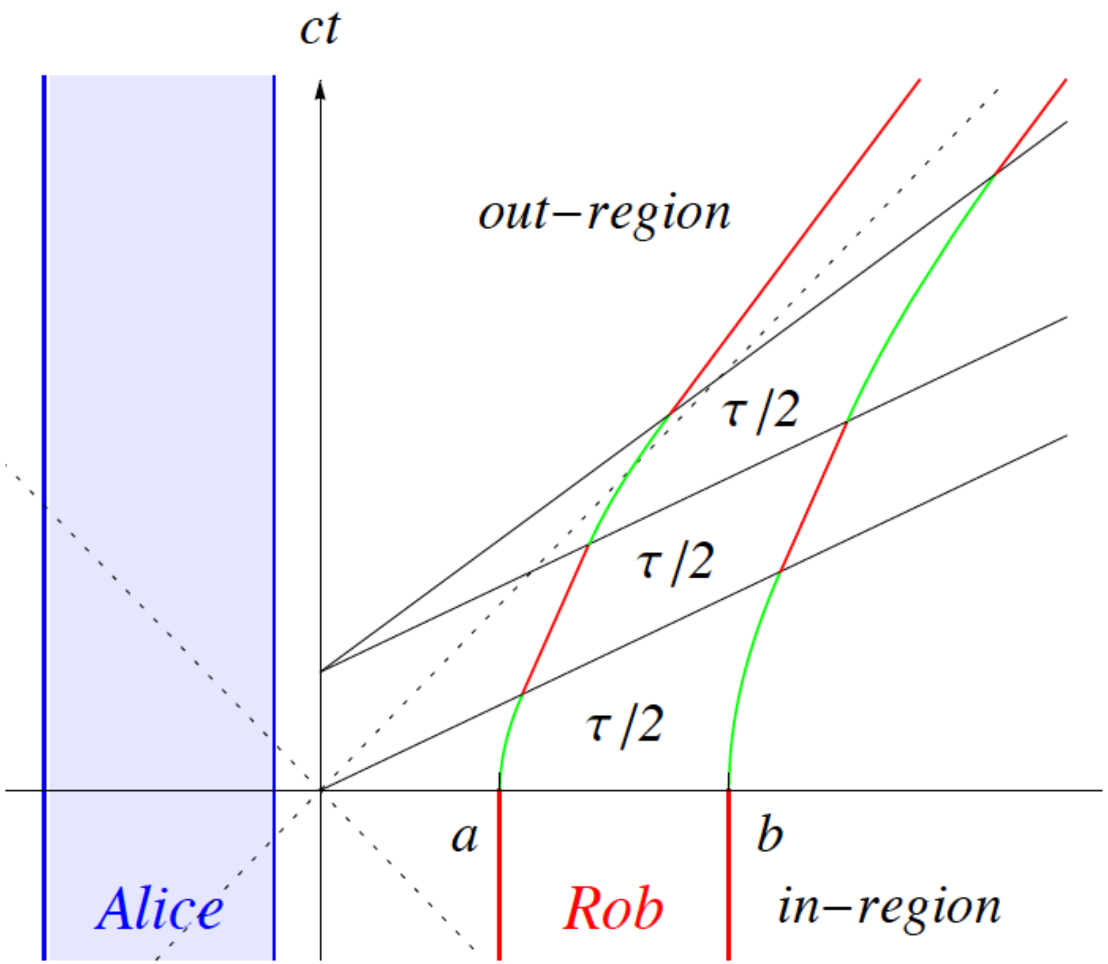
general trajectories
continuous motion
including circular acceleration

initial separable squeezed state

entanglement: negativity

$$\mathcal{N} = (\Re(e^{i\omega_a t}\beta_{ab})^2 + (\Im(e^{i\omega_a t}\beta_{ab})\cosh(s) - \Im(e^{i\omega_a t}\alpha_{ab})\sinh(s))^{1/2}$$

motion and gravity create entanglement



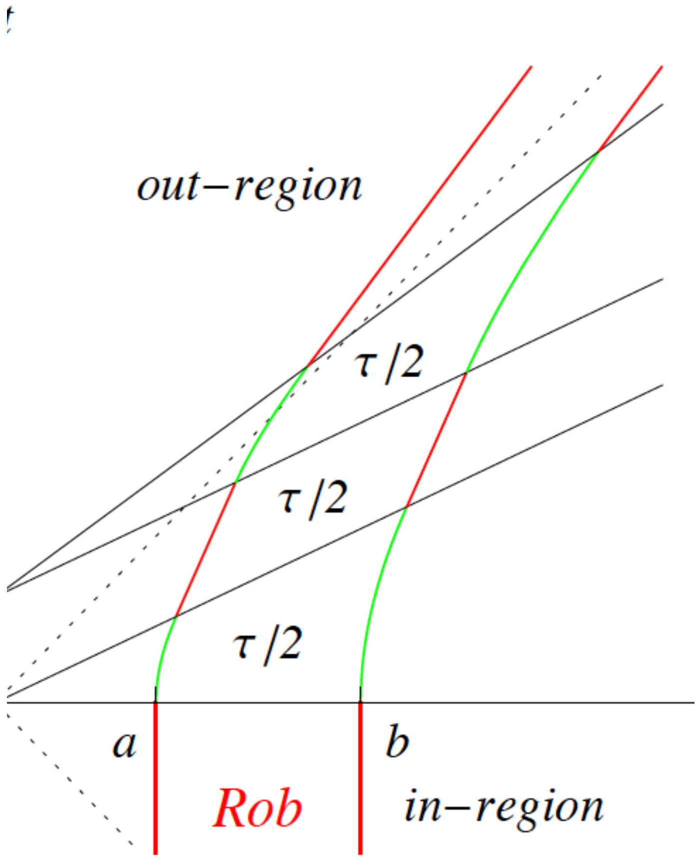
Friis, Bruschi, Louko, Fuentes PRD (R) 2012

results

non-uniform motion creates entanglement

gravity creates entanglement

quantum gates



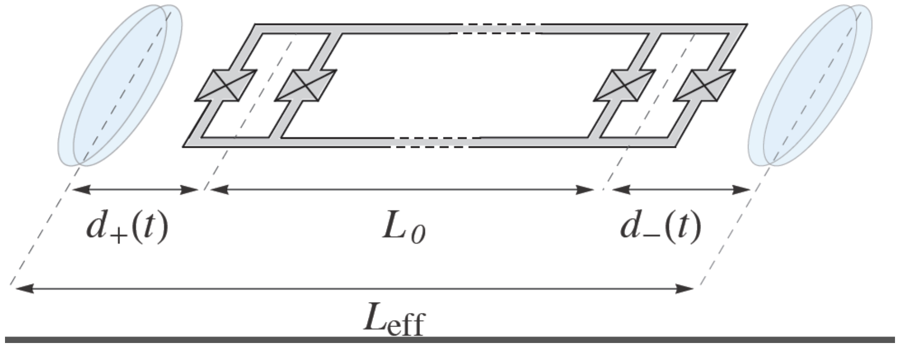
Friis, Huber, Fuentes, Bruschi PRD 2012
Bruschi, Lee, Dragan, Fuentes, Louko PRL 2012
Bruschi, Louko, Faccio & Fuentes NJP 2013
Bruschi, Sabin, Kok, Johansson, Delsing & Fuentes SR 2016

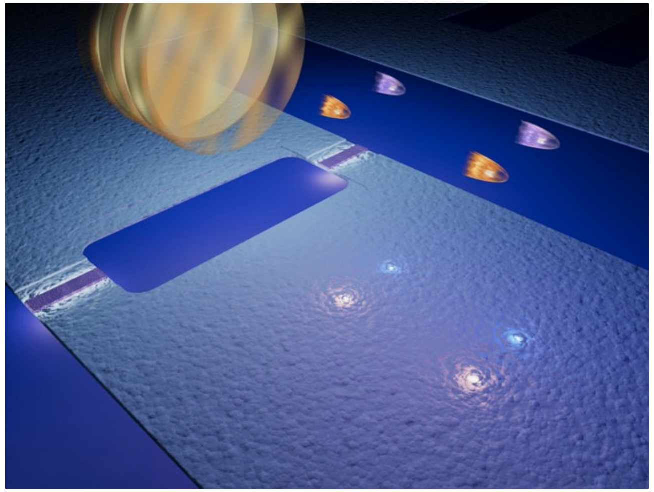
the relativistic motion of quantum systems can be used to produce quantum gates

two-mode squeezer
beam splitter

multi-qubit gates:
Dicke states
Multi-mode squeezer
Cluster states

Cavity moving at relativistic speeds using superconducting circuits



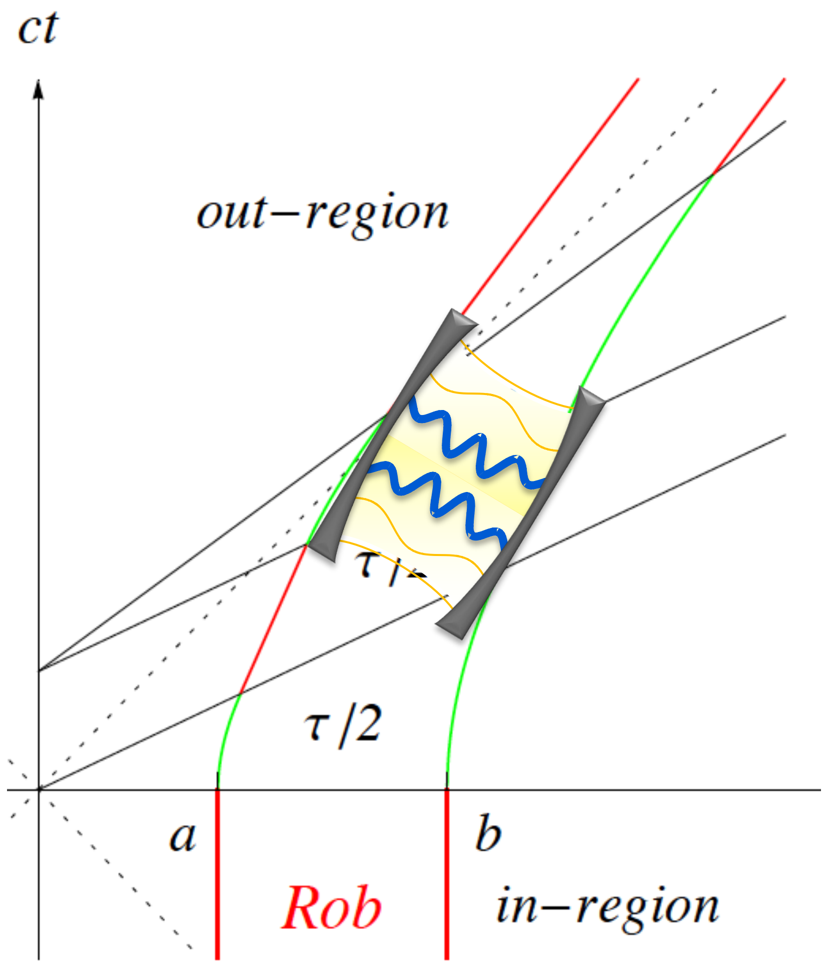


Generating Multimode Entangled Microwaves with a Superconducting Parametric Cavity

C. W. Sandbo Chang, M. Simoen, José Aumentado, Carlos Sabín, P. Forn-Díaz, A. M. Vadiraj, Fernando Quijandría, G. Johansson, I. Fuentes, and C. M. Wilson

Phys. Rev. Applied **10**, 044019 – Published 8 October 2018

Cavities moving in curved spacetime



PAPER • OPEN ACCESS

Dynamical Casimir effect in curved spacetime

Maximilian P E Lock[1,2] and Ivette Fuentes[4,2,3]

Published 7 July 2017 • © 2017 IOP Publishing Ltd and Deutsche Physikalische Gesellschaft

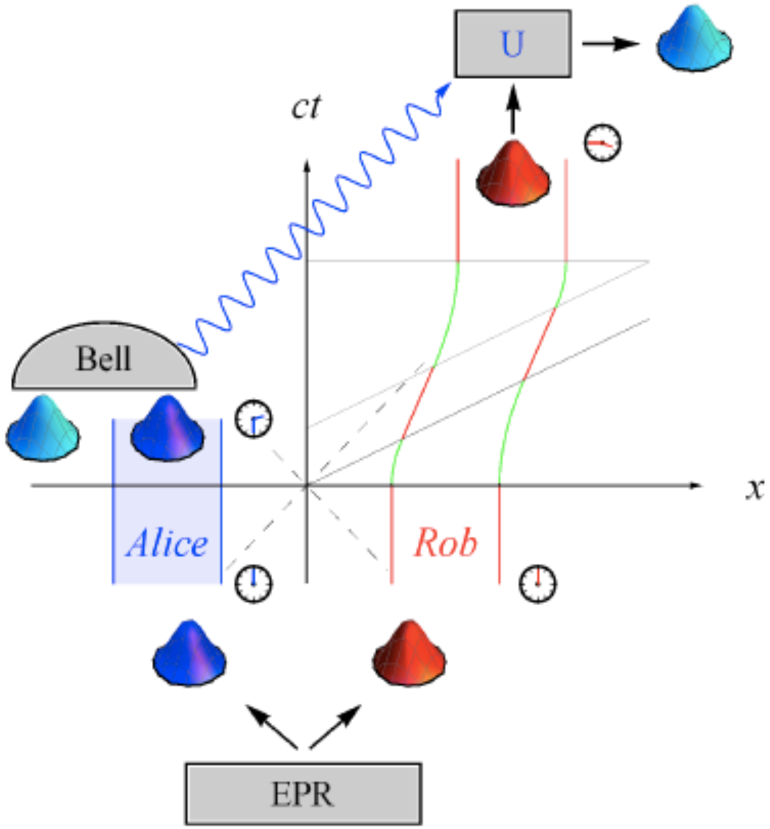
New Journal of Physics, Volume 19, July 2017

Citation Maximilian P E Lock and Ivette Fuentes 2017 *New J. Phys.* **19** 073005

We find corrections due to curvature effects

Teleporation with an accelerated partner

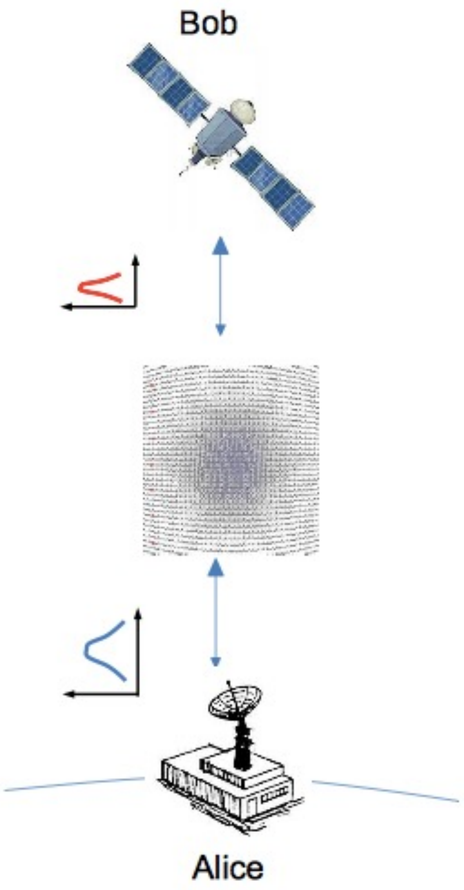
N. Friis, A. R. Lee, K. Truong, C. Sabín, E. Solano, G. Johansson, and I. Fuentes
Phys. Rev. Lett. **110**, 113602 – Published 12 March 2013



The fidelity of teleportation is affected by motion
It is possible to correct by local rotations
and trip planning

Spacetime effects on satellite-based quantum communications

David Edward Bruschi, Timothy C. Ralph, Ivette Fuentes, Thomas Jennewein, and Mohsen Razavi
Phys. Rev. D **90**, 045041 – Published 28 August 2014



Observable effects in satellite-based quantum cryptography

$$\Phi = \int_0^{+\infty} d\omega \left[\phi_\omega^{(u)} a_\omega + \phi_\omega^{(v)} b_\omega + \text{h.c.}\right]$$

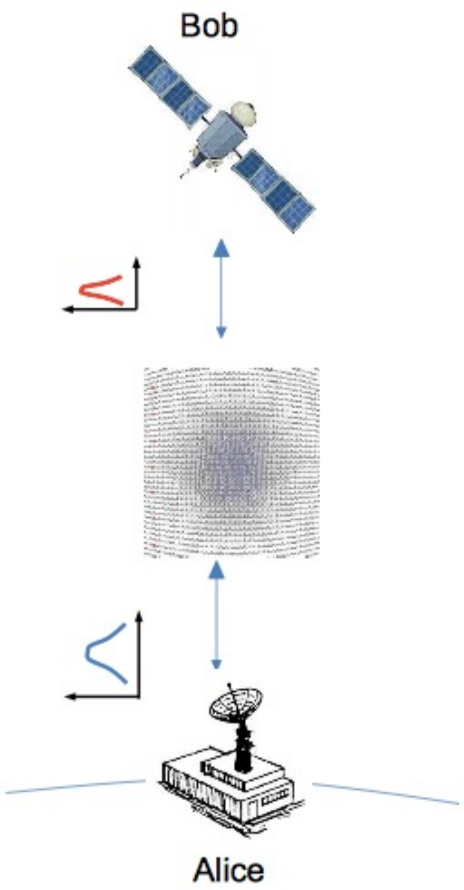
$$a_{\omega_0}(t) = \int_0^{+\infty} d\omega \, e^{-i\omega t} F_{\omega_0}(\omega) \, a_\omega .$$ Traveling wave-packet

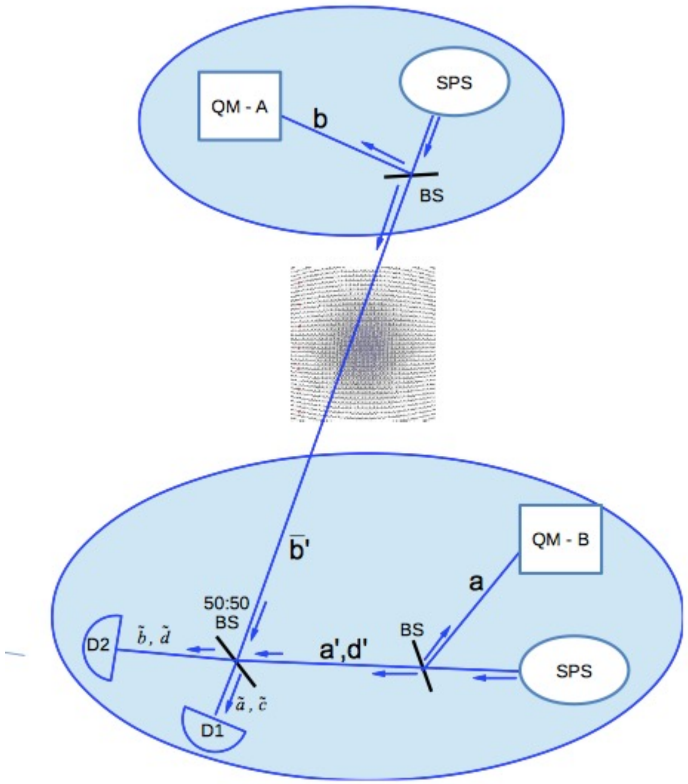
$$S = \begin{pmatrix} \Theta \mathbb{1}_2 & \sqrt{1-\Theta^2}\mathbb{1}_2 \\ -\sqrt{1-\Theta^2}\mathbb{1}_2 & \Theta \mathbb{1}_2 \end{pmatrix}$$

Spacetime acts as a beam-splitter

Spacetime effects on satellite-based quantum cryptography

Bruschi, Ralph, Fuentes, Jennewein, Razavi PRD 2014



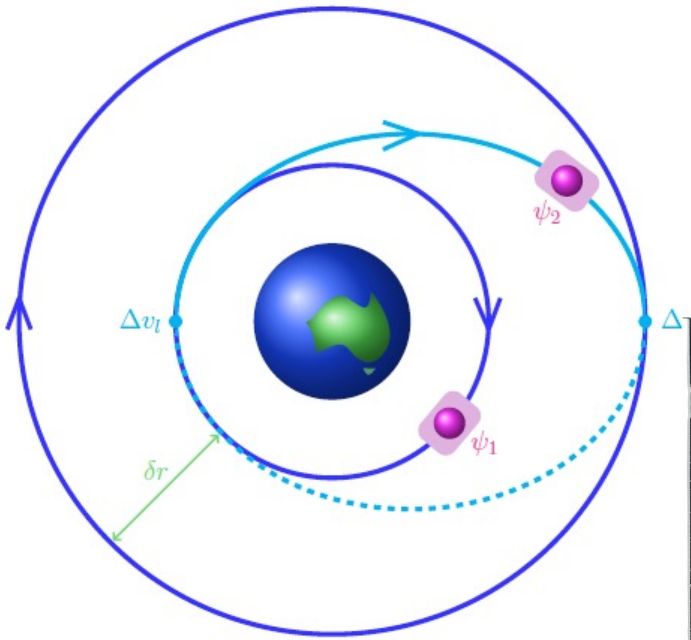


Testing the effects of gravity and motion on quantum entanglement in space-based experiments

David Edward Bruschi[1], Carlos Sabín[2], Angela White[3,6], Valentina Baccetti[4], Daniel K L Oi[5] and Ivette Fuentes[2]

Published 21 May 2014 •

New Journal of Physics, Volume 16, May 2014



Effects of gravity and motion on entanglement

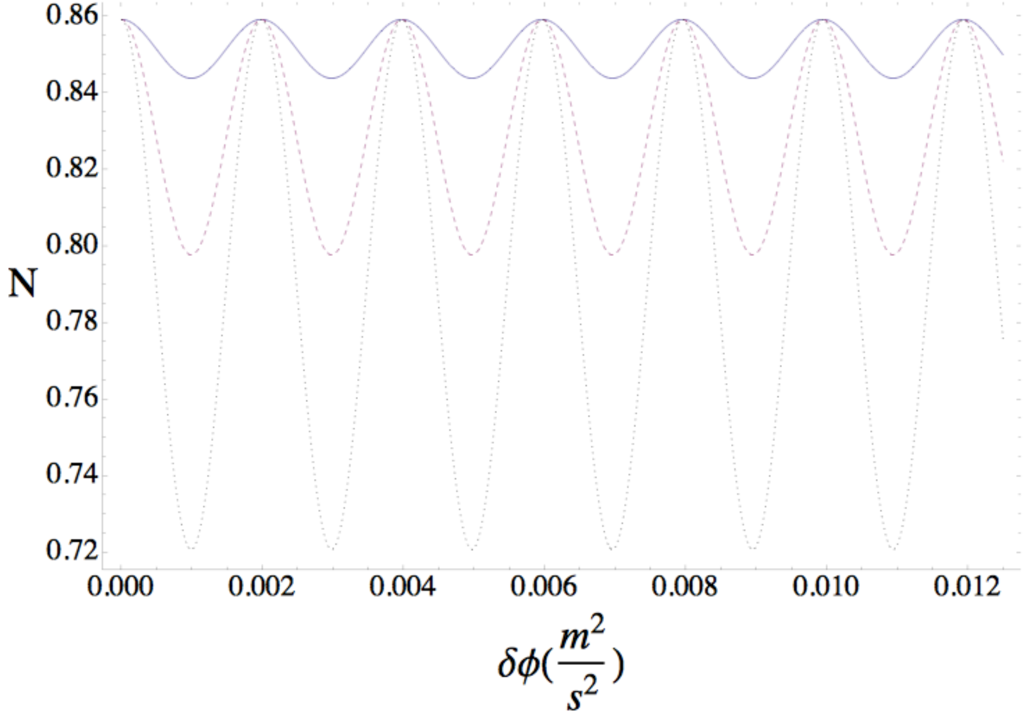
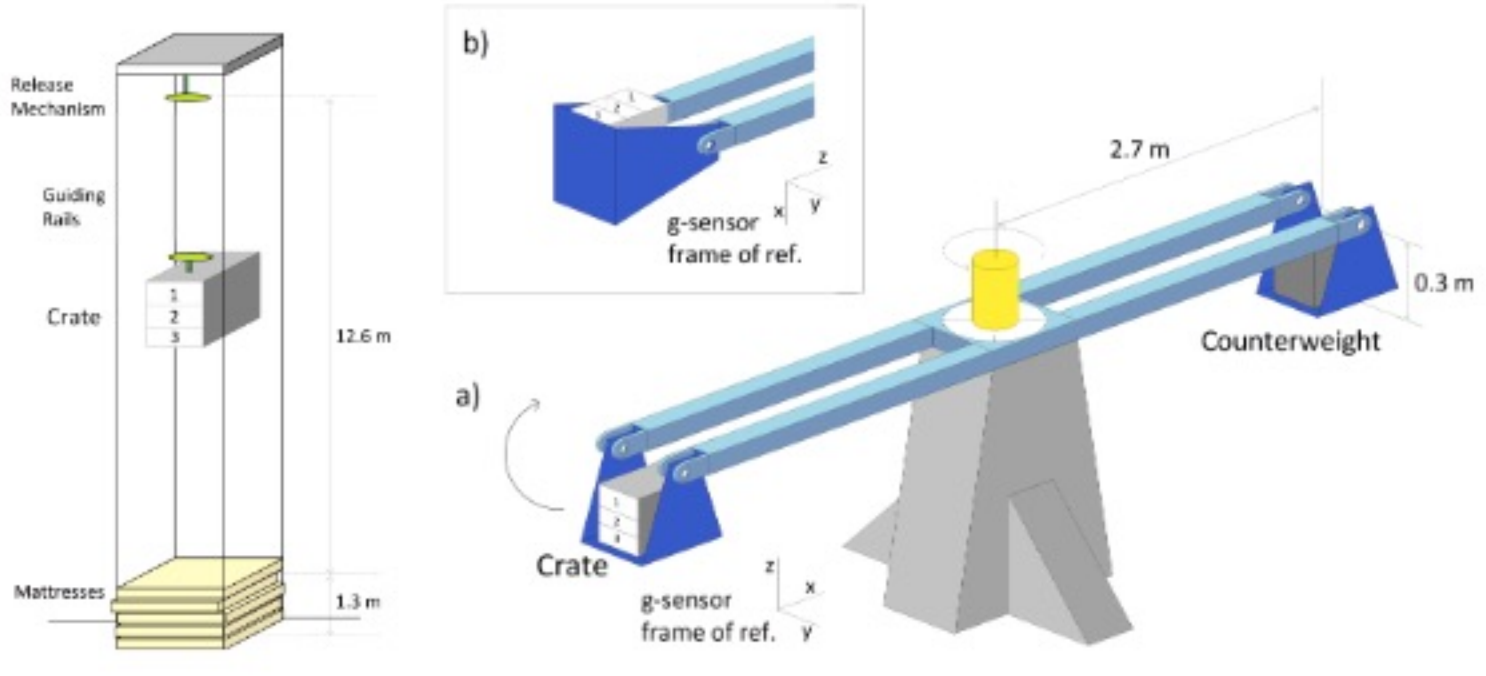


Figure 2. Negativity N vs. difference in gravitational field strength between initial and final orbits $\delta\phi$, after the first change in velocity Δv_l. The acceleration of the satellite is $a = 10^{-3}\,\mathrm{m/s^2}$ (solid, blue), $a = 2\cdot10^{-3}\,\mathrm{m/s^2}$ (red,dashed), $a = 3\cdot10^{-3}\,\mathrm{m/s^2}$ (black, dotted) while $L = 100\,\mu\mathrm{m}$, $c_s = 1\,\mathrm{mm/s}$, giving rise to $h^2 \simeq 0.05$ and $\Omega_1 = 2\pi \times 50$ Hz. The initial squeezing is $r = 1/2$.

Experimental test of photonic entanglement in accelerated reference frames

Matthias Fink, Ana Rodriguez-Aramendia, Johannes Handsteiner, Abdul Ziarkash, Fabian Steinlechner, Thomas Scheidl, Ivette Fuentes, Jacques Pienaar, Timothy C. Ralph & Rupert Ursin

Nature Communications **8**, Article number: 15304 (2017) | Cite this article

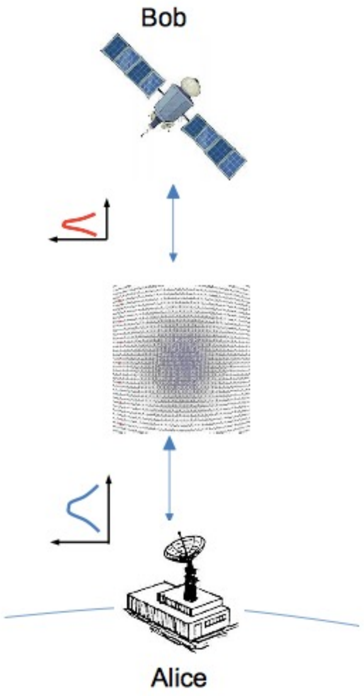


entanglement under uniform acceleration is conserved

Future experiments: non-uniform acceleration and Satellite-based experiments

Quantum estimation of the Schwarzschild spacetime parameters of the Earth

David Edward Bruschi, Animesh Datta, Rupert Ursin, Timothy C. Ralph, and Ivette Fuentes
Phys. Rev. D **90**, 124001 – Published 1 December 2014



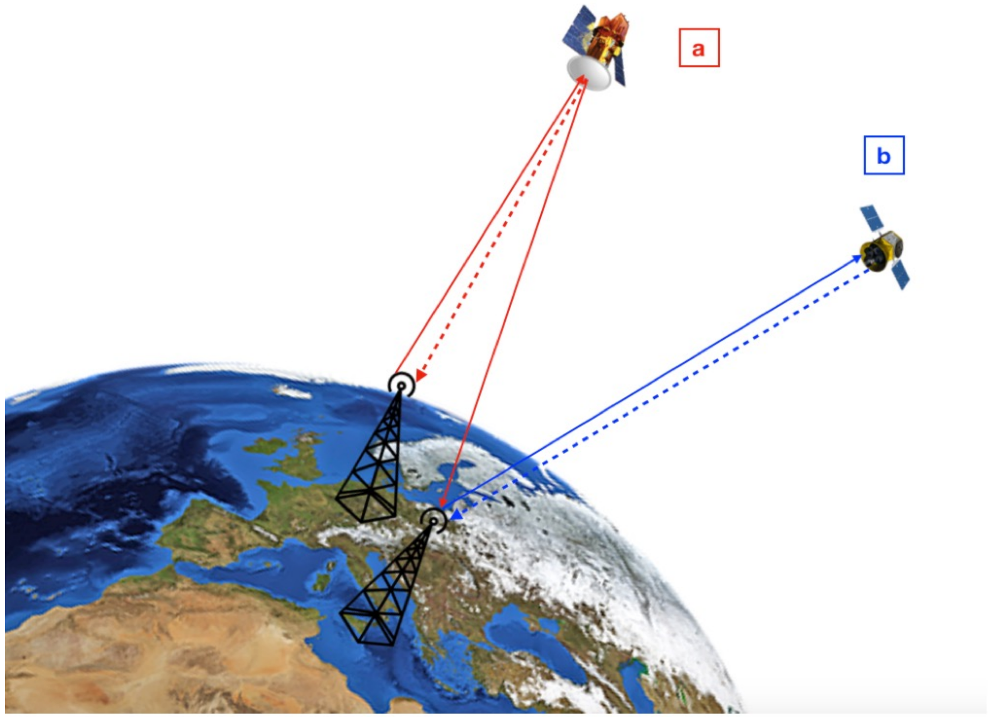
$$\boldsymbol{g} = \mathrm{diag}\left(-\left(1-\frac{2M}{r}\right), \frac{1}{1-\frac{2M}{r}}, r^2, r^2 \sin^2\theta\right)$$

Estimating the Schwarzschild radius of the Earth
Using traveling wave-packets

$$\frac{\Delta r_S}{r_S} \geq \frac{8\,\sigma\, r_A\,(r_A+L)}{\sqrt{N(\Omega_1^2+\Omega_2^2)}\, r_S\, L\, \sinh r} \sim 5.8\times 10^{-7}$$

Quantum-metrology estimation of spacetime parameters of the Earth outperforming classical precision

Jan Kohlrus, David Edward Bruschi, and Ivette Fuentes
Phys. Rev. A **99**, 032350 – Published 29 March 2019



Orbit	Light's trajectory	$\Delta r_S/r_S$	$\Delta R_E/R_E$	$\Delta h/h$
LEO	Lab1 $\leftrightarrows$ Lab2	8.50×10^{-10}	1.12×10^{-9}	3.56×10^{-9}
	Lab1 $\rightarrow$ Lab1	8.87×10^{-10}	1.17×10^{-9}	3.71×10^{-9}
	Lab2 $\rightarrow$ Lab2	8.15×10^{-10}	1.07×10^{-9}	3.41×10^{-9}
	Sat $\rightarrow$ Lab1	1.77×10^{-9}	2.33×10^{-9}	7.43×10^{-9}
	Sat $\rightarrow$ Lab2	1.63×10^{-9}	2.14×10^{-9}	6.83×10^{-9}
VLEO	Lab1 $\rightarrow$ Lab1	6.06×10^{-10}	6.30×10^{-10}	1.57×10^{-8}
	Sat $\rightarrow$ Lab 1	1.21×10^{-9}	1.26×10^{-9}	3.15×10^{-8}

TABLE II: Precision bounds obtained through the quantum metrology scheme described above, for the different possible configurations of the reflecting and downlink schemes. Results for uplinks are very similar to those in downlinks.

Join my group

Available theoretical and experimental PhD studentships

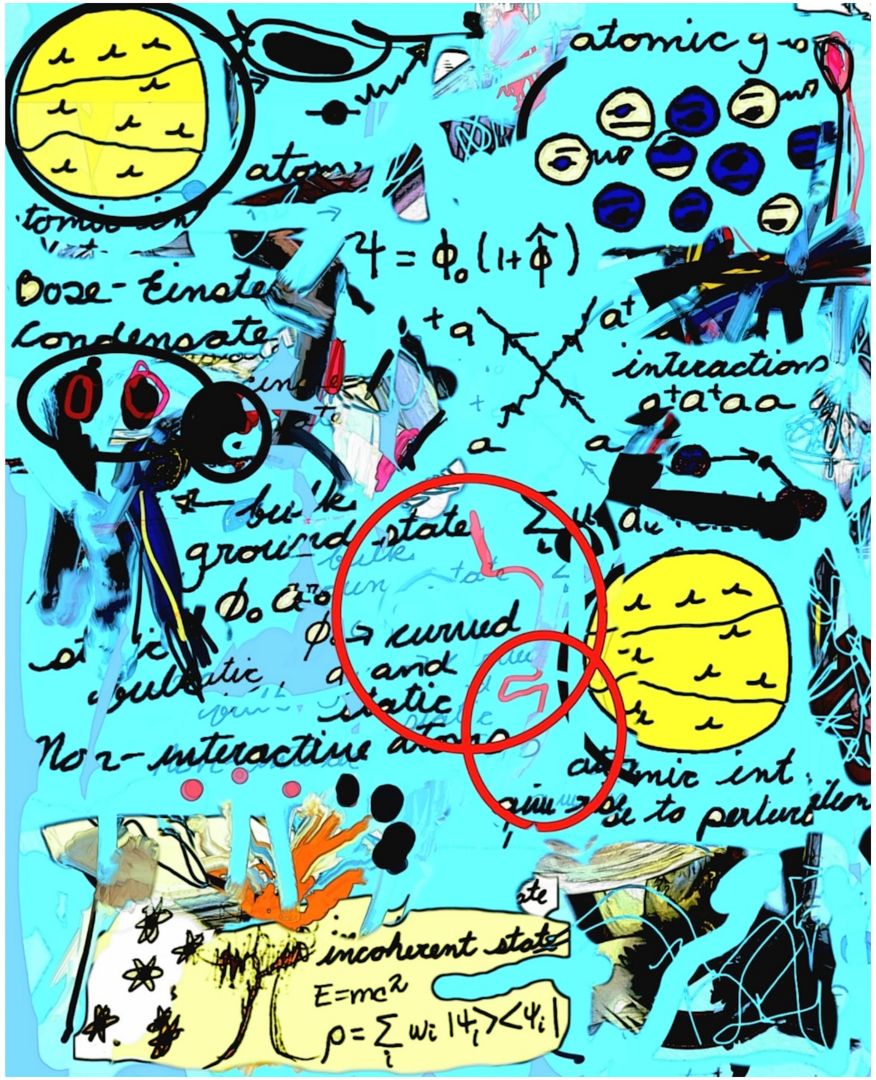
Home UK students at University of Southampton (start this or next year)
International students, Emmy Studentship at Keble College, Oxford (start next year)

Available theoretical and experimental postdoctoral positions

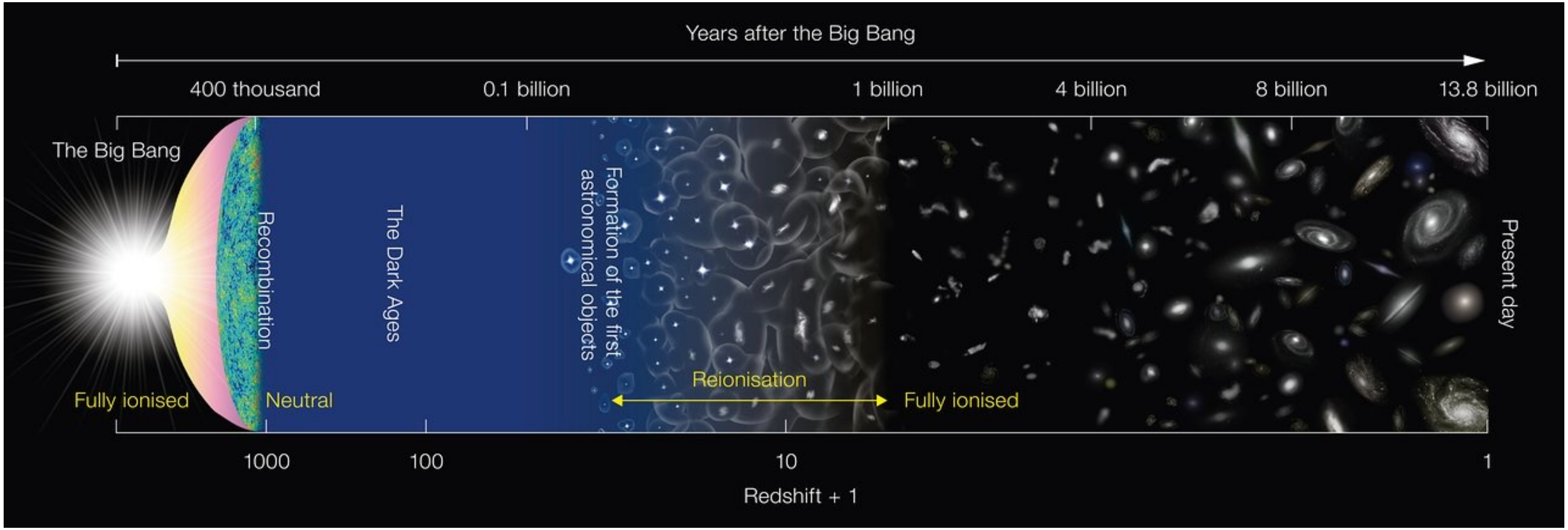
Informal inquires: i.fuentes-guridi@soton.ac.uk

look at my website for more information about my group
https://ivettefuentes.weebly.com

Relativistic Quantum Metrology
and
Analogue experiments



Many fundamental questions unanswered



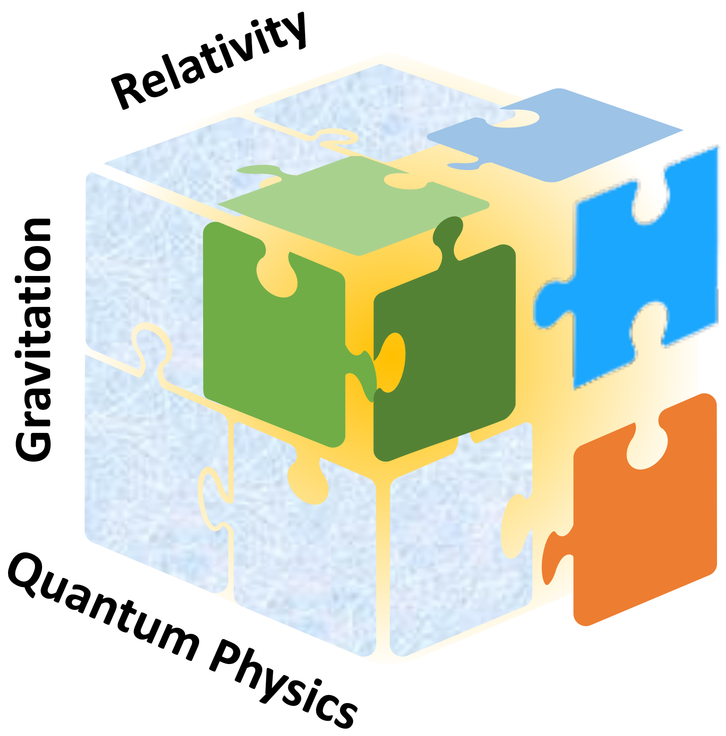
What is the nature of dark matter?

Is dark energy driving the accelerated expansion of the Universe?

What physics dominated the Universe at early times?

Does the equivalence principle holds for quantum systems?

Underpinning our difficulty to find answers



Quantum physics and General Relativity
are incompatible

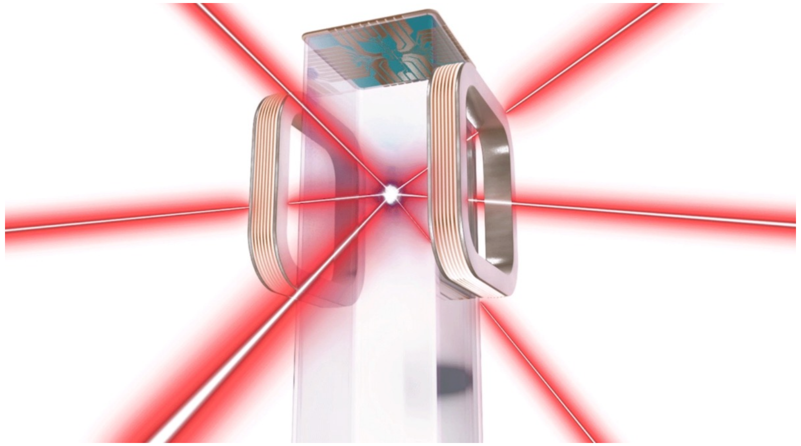
Technology and instruments must be developed first

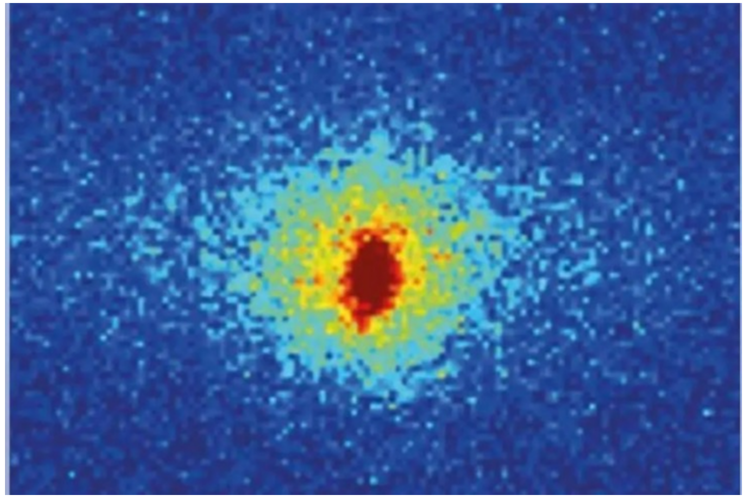
The interplay of theory and experiment

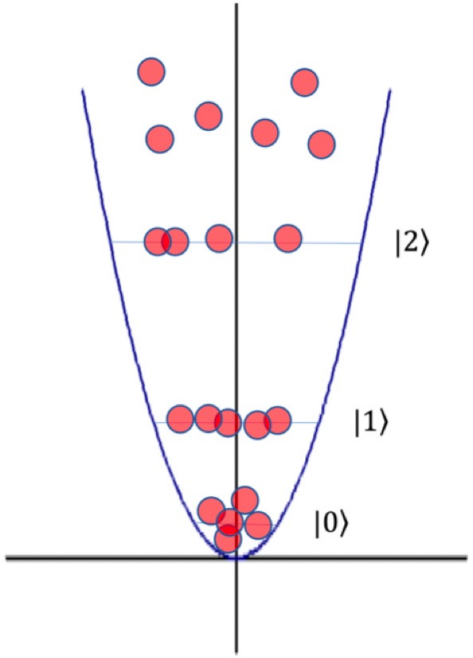
Main aim: Develop new instruments
window to scales that have not been explored before
New Physics!

Quantum technologies can help us put the pieces of the puzzle together

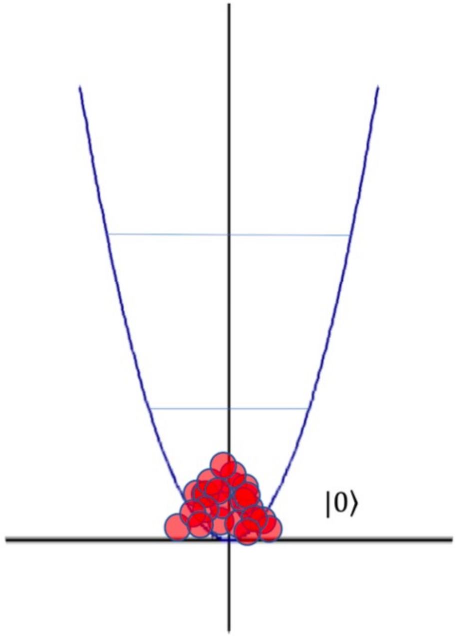
Bose-Einstein Condensate (BEC)







Cold atoms in a thermal distribution

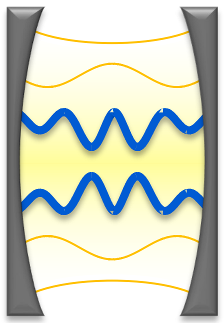


Atoms in a BEC form a lump

BEC in spacetime

A covariant formalism is available

Phonons are a relativistic quantum field



$$\mathcal{L} = -\sqrt{-g}\left\{g^{\mu\nu}\partial_\mu\hat{\Phi}^\dagger\partial_\nu\hat{\Phi} + \left(\frac{m^2c^2}{\hbar^2} + V\right)\hat{\Phi}^\dagger\hat{\Phi} + U\left(\hat{\Phi}^\dagger\hat{\Phi}, \lambda_i\right)\right\}$$

$$U\left(\hat{\Phi}^\dagger\hat{\Phi}, \lambda_i\right) \approx \frac{\lambda}{2}\hat{\Phi}^\dagger\hat{\Phi}^\dagger\hat{\Phi}\hat{\Phi}$$

$$\widehat{\Phi} = \phi(1 + \widehat{\psi})$$

For uniform density

$$\Box\,\psi = \frac{1}{\sqrt{-\mathfrak{g}}}\partial_\mu(\sqrt{-\mathfrak{g}}\,\mathfrak{g}^{\mu\nu}\partial_\mu)\psi$$

effective metric

$$\mathfrak{g}_{ab} = \rho\frac{c}{c_s}\left[g_{ab} + \left(1 - \frac{c_s^2}{c^2}v_a v_b\right)\right]$$

spacetime metric

analogue metric

Demonstrate QFT in CS

Fagnocchi et. al NJP 2010 (quantum field flat spacetime)
Visser & Molina-Paris NJP 2010 (classical field in curved spacetime)
D. E. Bruschi et. al., NJP 16 5:053041, 2014 (quantum field in curved spatime)
D. Hartley. et. al., PRD 89 025011, 2018

Analogue spacetimes and simulations

Analogue Systems

moving fluid
superfluid helium
Bose–Einstein condensate
gravity waves in water
electromagnetic waves in a dielectric medium

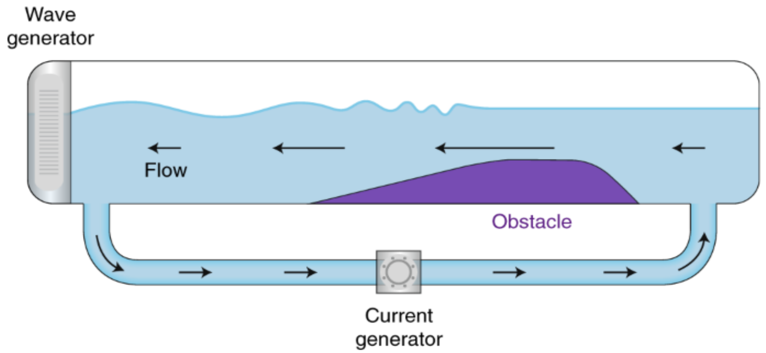
Published: 12 October 2014

Observation of self-amplifying Hawking radiation in an analogue black-hole laser

Jeff Steinhauer

Nature Physics **10**, 864–869 (2014) | Cite this article

13k Accesses | **227** Citations | **388** Altmetric | Metrics



Published: 12 June 2017

Rotational superradiant scattering in a vortex flow

Theo Torres, Sam Patrick, Antonin Coutant, Maurício Richartz, Edmund W. Tedford & Silke Weinfurtner

Nature Physics **13**, 833–836 (2017) | Cite this article

11k Accesses | **141** Citations | **240** Altmetric | Metrics

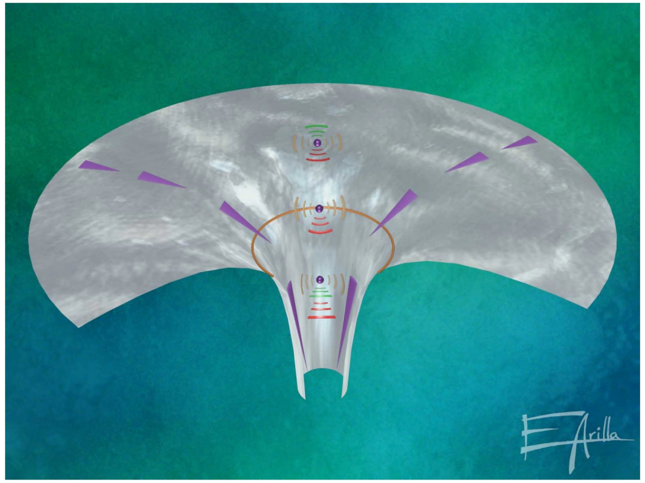
Quantum analogue of a Kerr black hole and the Penrose effect in a Bose-Einstein condensate

D. D. Solnyshkov, C. Leblanc, S. V. Koniakhin, O. Bleu, and G. Malpuech
Phys. Rev. B **99**, 214511 – Published 24 June 2019

Covariant formulation

$$\mathcal{L} = -\sqrt{-g}\left\{g^{\mu\nu}\partial_\mu\hat{\Phi}^\dagger\partial_\nu\hat{\Phi} + \left(\frac{m^2c^2}{\hbar^2} + V\right)\hat{\Phi}^\dagger\hat{\Phi} + U\left(\hat{\Phi}^\dagger\hat{\Phi}, \lambda_i\right)\right\}$$

BEC: barotropic, irrotational and inviscid fluid, in a covariant formalism.



i) simulating a spacetime metric
ii) simulating the effects of spacetime dynamics on a phononic field.

1. Superfluid regime (barotropic and inviscid)

Bogoliubov approximation so $\left\langle\hat{\Phi}\right\rangle \approx \phi$ and $\left\langle\hat{\psi}^\dagger\hat{\psi}\right\rangle \ll |\phi|^2$.

$$\hat{\Phi} = \phi\left(1 + \hat{\psi}\right) \qquad \phi = \sqrt{\rho}e^{i\theta}$$

2. In the long wavelength regime $|k| \ll \frac{\sqrt{2}}{\xi}\left(1+\frac{\hbar^2}{2m^2\xi^2u_0^2}\right)\min\left[1, \frac{mu_0\xi}{\sqrt{2}\hbar}\right],$

Velocity flows $u_\mu = \frac{\hbar}{m}\partial_\mu\theta$

Scattering length

Healing length $\xi = \frac{1}{\sqrt{\lambda\rho}}.$

$\lambda = 8\pi a.$

phonons should have wavelengths far longer than the healing length

$U\left(\phi^\dagger\phi, \lambda\right) = \frac{1}{2}\lambda\left|\phi^\dagger\phi\right|^2 + \cdots$

Strength of the interaction

The fluctuations are massless phonons

Plugging the Bogoliubov approximation in the equations of motion for the Lagrangian yields

$$\frac{1}{\sqrt{-G}}\partial_\mu\sqrt{-G}G^{\mu\nu}\partial_\nu\hat{\psi} = 0$$

$v^\mu = \frac{c}{|u|}u^\mu$ Normalized velocity flows

$$G_{\mu\nu} = \frac{\rho c}{c_s}\left[g_{\mu\nu} + \left(1-\frac{c_s^2}{c^2}\right)\frac{v_\mu v_\nu}{c^2}\right].$$

$c_s^2 = \frac{c^2c_0^2}{c_0^2+|u|^2}$

Effective speed of sound

$$c_0^2 = \frac{\hbar^2}{2m^2}\rho\partial_\rho^2U\left(\rho,\lambda\right) = \frac{\hbar^2}{2m^2}\lambda\rho.$$

The equations of motion of the mean field ϕ can be split into real and imaginary components resulting in a continuity equation

$$\nabla_\mu \left(\rho u^\mu\right) = 0$$

PHYSICAL REVIEW D

covering particles, fields, gravitation, and cosmology

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Quantum simulation of dark energy candidates

Daniel Hartley, Christian Käding, Richard Howl, and Ivette Fuentes

Phys. Rev. D **99**, 105002 – Published 10 May 2019

and an equation directly relating bulk field properties with potentials:

$$|u|^2 = c^2 + \frac{\hbar^2}{m^2}\left\{V + \partial_\rho U\left(\rho, \lambda\right) - \frac{\nabla_\mu \nabla^\mu \sqrt{\rho}}{\sqrt{\rho}}\right\}.$$

$$G_{\mu\nu} \propto \begin{pmatrix} -\left(c_s^2 - v^2\right)/c^2 & -v_i/c \\ -v_j/c & \delta_{ij} \end{pmatrix}$$

Acoustic metric commonly used. We can obtain it in our covariant formulation in the following way.

Flat spacetime

We define a "flat" acoustic metric for the phonons in the absence of any simulated fields as the acoustic metric in coordinates with the following conditions:

1. The background spacetime is flat, i.e. $g_{\mu\nu} = \eta_{\mu\nu} = \mathrm{diag}\,(-1, 1, 1, 1)$,
2. No flows, i.e. $v_i = 0$,
3. Static density, i.e. $\rho_0 = \rho_0\,(\boldsymbol{x}) \Leftrightarrow \partial_t \rho_0 = 0$, and
4. Unperturbed interaction strength λ, so $\partial_t \lambda = \partial_x \lambda = 0$.

The acoustic metric in $n+1$ dimensions then has the form

$$G^0_{\mu\nu} = \frac{\rho_0 c}{c_{s0}} \begin{pmatrix} -c_{s0}^2/c^2 & \\ & \mathbb{I}_n \end{pmatrix}$$

$$c_{s0} = \frac{\hbar}{m}\sqrt{4\pi\rho\, a}.$$

Simplest simulation

Gravitational wave in 1+1 dimension

$$g_{\mu\nu} = \eta_{\mu\nu} + h_{\mu\nu}$$

A gravitational wave moving in the z-direction

$$\eta_{\mu\nu} = \begin{pmatrix} -c^2 & 0 & 0 & 0 \\ 0 & 1 & 0 & 0 \\ 0 & 0 & 1 & 0 \\ 0 & 0 & 0 & 1 \end{pmatrix}$$

$$h_{\mu\nu} = \begin{pmatrix} 0 & 0 & 0 & 0 \\ 0 & h_+(t) & h_\times(t) & 0 \\ 0 & h_\times(t) & -h_+(t) & 0 \\ 0 & 0 & 0 & 0 \end{pmatrix}$$

time-dependent perturbations in two different polarisations.

In one dimension

$$ds^2 = -c^2\,dt^2 + (1 + h_+(t))\,dx^2.$$

BEC in flat spacetime
in the absence of spatial flows

$v^t = c$ and $v^x = 0$, then the line element is conformal to:

$$ds^2 = -c_s^2\,dt^2 + dx^2$$

$$\mathfrak{g}_{ab} = \frac{\rho\,c_s}{c} \begin{pmatrix} -c_s^2 & 0 & 0 & 0 \\ 0 & 1 + h_+(t) & h_\times(t) & 0 \\ 0 & h_\times(t) & 1 - h_+(t) & 0 \\ 0 & 0 & 0 & 1 \end{pmatrix}$$

Look FB resonance and add tupac's references

Now, if we consider that the speed of sound can depend on t:

$$c_s(t) = c_{s0}\, f(t),$$

the corresponding line element is conformal to:

$$c_{s0} = \frac{\hbar}{m}\sqrt{4\pi\rho\, a}.$$

$$ds^2 = -c_{s0}^2\, dt^2 + \frac{1}{(f(t))^2}\, dx^2.$$

Therefore, if the speed of sound varies in time such that

$$f(t) = (1 - \frac{A_+ \sin\Omega\, t}{2}),$$

we find, up to the first order in A_+:

$$ds^2 = -c_{s0}^2\, dt^2 + (1 + A_+ \sin\Omega\, t)\, dx^2.$$

So the experimental task is to modulate the speed of sound as:

$$c_s(t) = c_{s0}(1 - \frac{A_+ \sin\Omega\, t}{2}).$$

Vary scattering length
using magnetic fields

background scattering length

$$a = a_{bg}(1 - \frac{\omega}{B - B_0})$$

Feshbach resonances

Recent experiments

Open Access

Thermalization by a synthetic horizon

Lotte Mertens, Ali G. Moghaddam, Dmitry Chernyavsky, Corentin Morice, Jeroen van den Brink, and Jasper van Wezel

Phys. Rev. Research **4**, 043084 – Published 8 November 2022

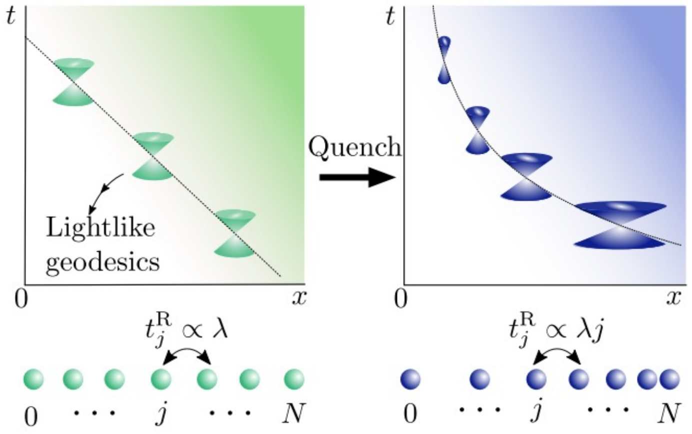


FIG. 1. Schematic overview of the quench set-up. In the top part of the figure a lightlike geodesic is drawn in a spacetime diagram for flat spacetime (left) and curved spacetime with a horizon at $x = 0$ (right). In the lower part the corresponding tight binding models are shown with constant hopping (left) and position dependent hopping (right).

28 November 2022 / Evrim Yazgin

Black hole simulation began glowing, providing hope for unifying gravity and quantum mechanics

The experiment may prove the existence of Hawking radiation.

Article | Published: 09 November 2022

Quantum field simulator for dynamics in curved spacetime

Celia Viermann, Marius Sparn, Nikolas Liebster, Maurus Hans, Elinor Kath, Álvaro Parra-López, Mireia Tolosa-Simeón, Natalia Sánchez-Kuntz, Tobias Haas, Helmut Strobel, Stefan Floerchinger & Markus K. Oberthaler

Nature **611**, 260–264 (2022) | Cite this article

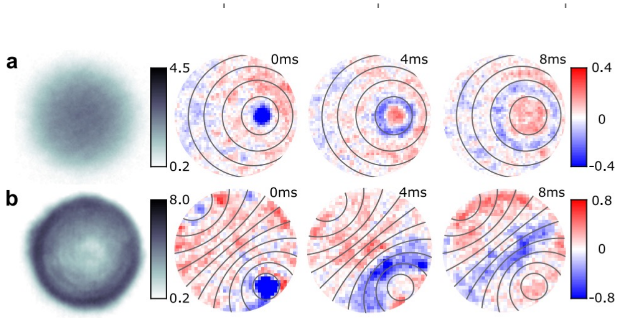


FIG. 2. **Configurable density distribution for hyperbolic and spherical geometry. a)** Density distribution of the condensate for hyperbolic geometry and expansion of a phononic wave packet depicted in th sate (blue under-, rec two-by-two region of p els. **b)** Density distrib the two-dimensional c propagation of the way ing confirming the suc positive curvature. E the initial perturbatio respectively.

RESEARCH ARTICLE | PHYSICS

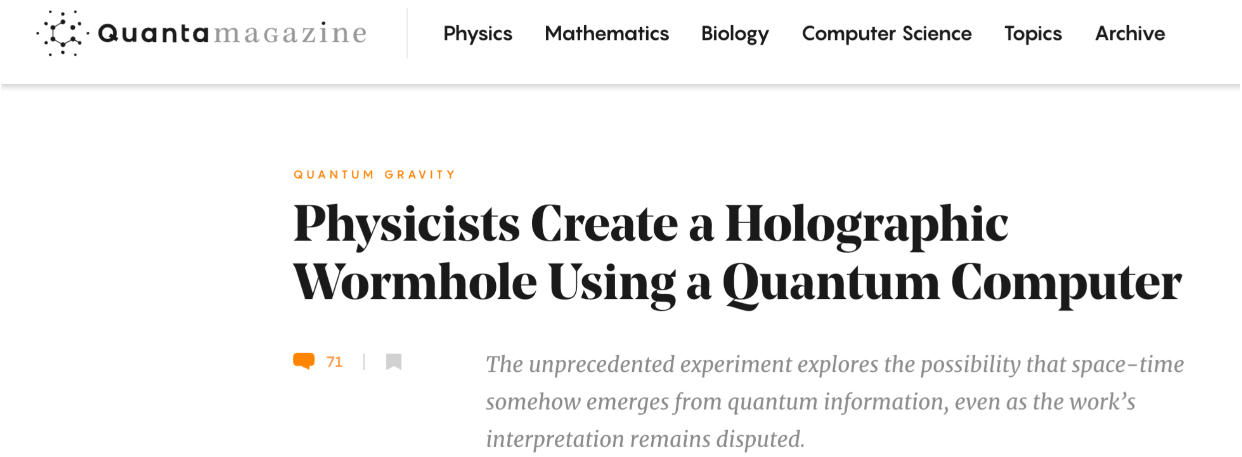
Experimental observation of curved light-cones in a quantum field simulator

Mohammadamin Tajik, Marek Gluza, Nicolas Sebe, +9, and Jörg Schmiedmayer Authors Info & Affiliations

Edited by Angel Rubio, Max-Planck-Institut fur Struktur und Dynamik der Materie, Hamburg, Germany; received January 23, 2023; accepted March 24, 2023

May 15, 2023 | 120 (21) e2301287120 | https://doi.org/10.1073/pnas.2301287120

Uses and Limitations


Quantamagazine
Physics Mathematics Biology Computer Science Topics Archive

QUANTUM GRAVITY

Physicists Create a Holographic Wormhole Using a Quantum Computer

71

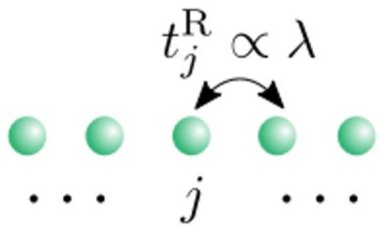
The unprecedented experiment explores the possibility that space-time somehow emerges from quantum information, even as the work's interpretation remains disputed.

-self-consistency checks
-useful to study the behaviour of matter under the influence of fields
with non-linear behaviour since their non-linearities can make computer simulations challenging.

A good analogy can be an invaluable tool in studying a complex or inaccessible system but…

Analogue experiments can't falsify or verify a physical theory
(as much as a computer simulation can't do this either)

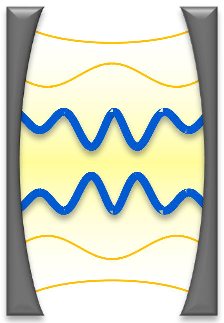
Attention!
New phenomena can be misinterpreted
Analogues cannot truly discover new effects



BEC in spacetime

A covariant formalism is available

Phonons are a relativistic quantum field



$$\mathcal{L} = -\sqrt{-g}\left\{g^{\mu\nu}\partial_\mu\hat{\Phi}^\dagger\partial_\nu\hat{\Phi} + \left(\frac{m^2c^2}{\hbar^2} + V\right)\hat{\Phi}^\dagger\hat{\Phi} + U\left(\hat{\Phi}^\dagger\hat{\Phi}, \lambda_i\right)\right\}$$

$$U\left(\hat{\Phi}^\dagger\hat{\Phi}, \lambda_i\right) \approx \frac{\lambda}{2}\hat{\Phi}^\dagger\hat{\Phi}^\dagger\hat{\Phi}\hat{\Phi}$$

$$\widehat{\Phi} = \phi(1 + \widehat{\psi})$$

For uniform density

$$\Box\,\psi = \frac{1}{\sqrt{-\mathfrak{g}}}\partial_\mu(\sqrt{-\mathfrak{g}}\mathfrak{g}^{\mu\nu}\partial_\mu)\psi$$

effective metric

$$\mathfrak{g}_{ab} = \rho\frac{c}{c_s}\left[g_{ab} + \left(1 - \frac{c_s^2}{c^2}v_a v_b\right)\right]$$

spacetime metric

analogue metric

Demonstrate QFT in CS

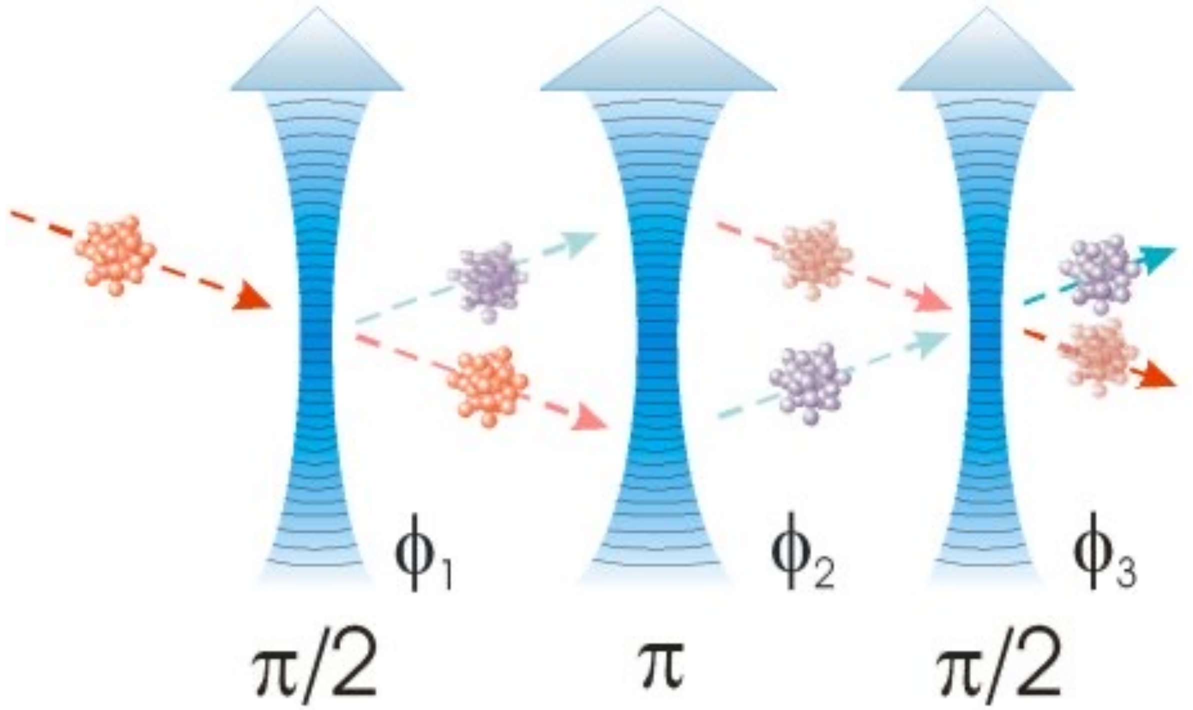
Fagnocchi et. al NJP 2010 (quantum field flat spacetime)

Visser & Molina-Paris NJP 2010 (classical field in curved spacetime)

D. E. Bruschi et. al., NJP 16 5:053041, 2014 (quantum field in curved spatime)

D. Hartley. et. al., PRD 89 025011, 2018

Atom interferometer: quantum spatial interferometery



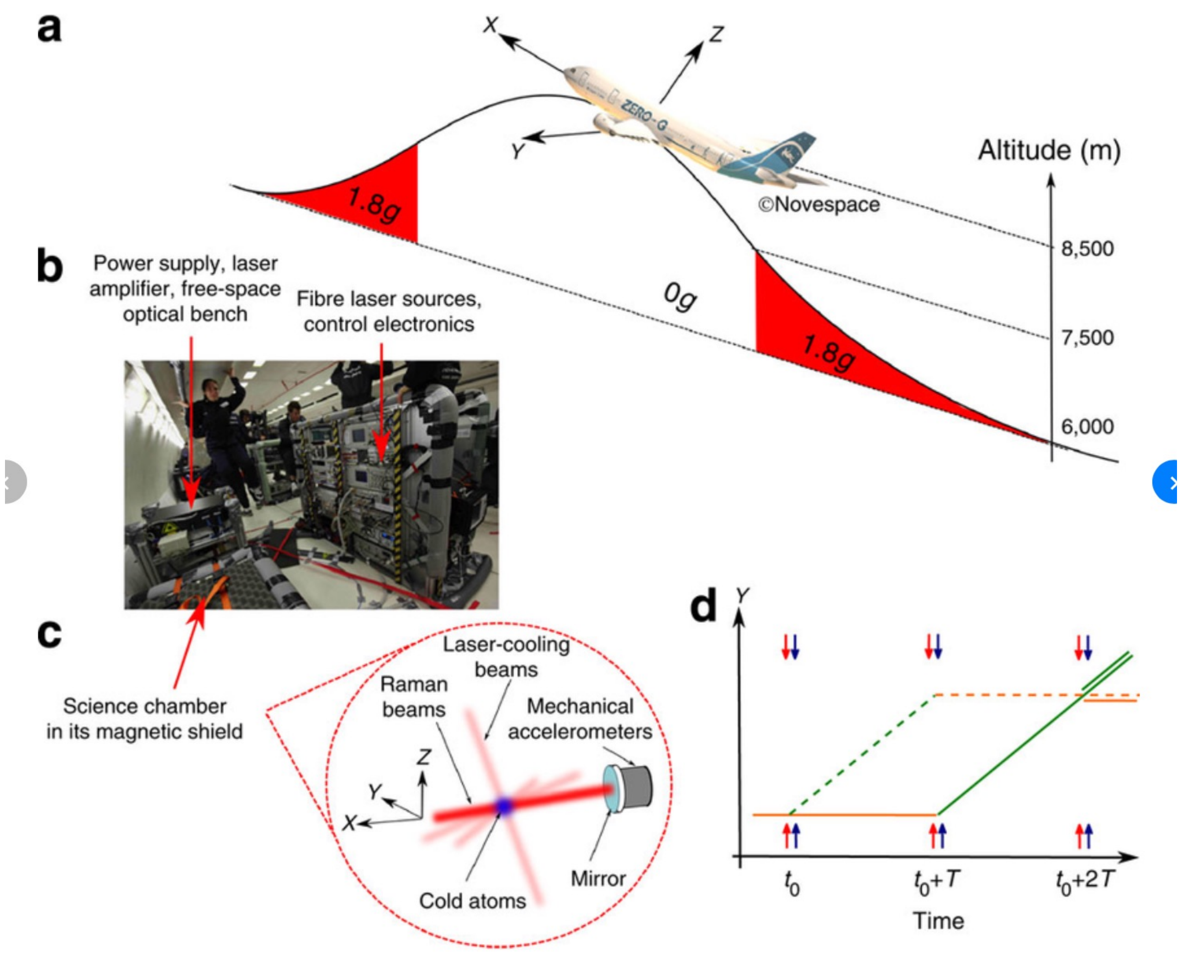
$$\delta a = 1 \Big/ \left(\sqrt{N} k T^2 \right)$$

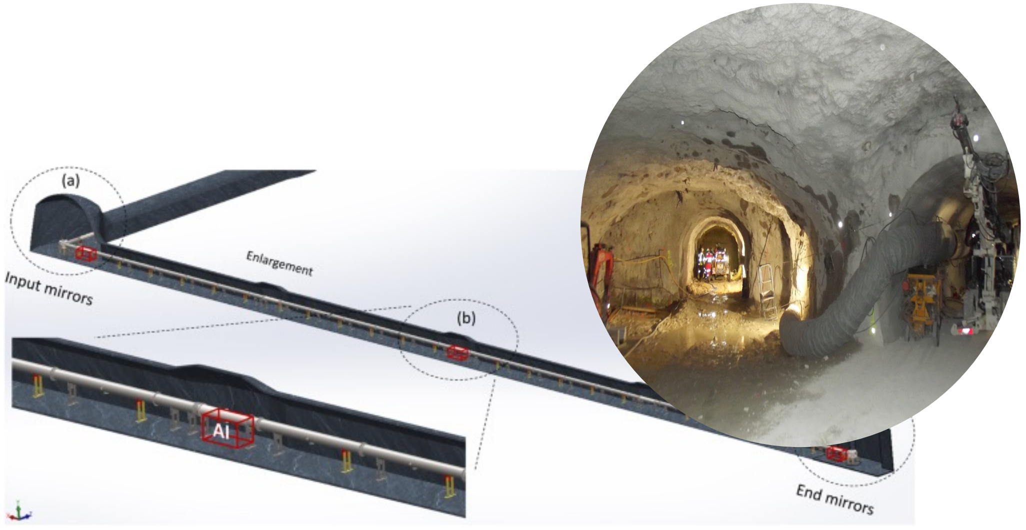
Single particle detector, local

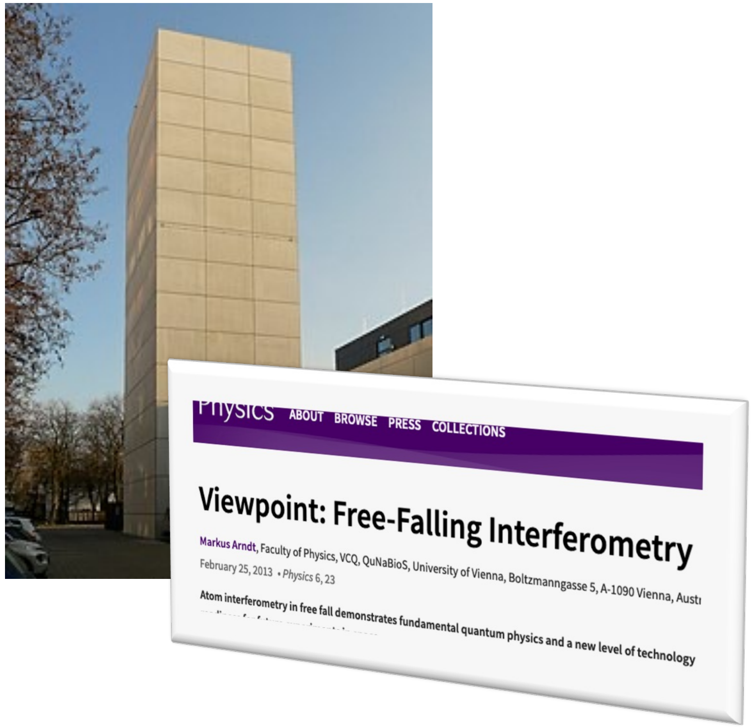
Interferometry in the <u>spatial domain</u>: limited by time of flight

Compatible with Newtonian physics

Gravimeters are going Big







But…I want one in my phone

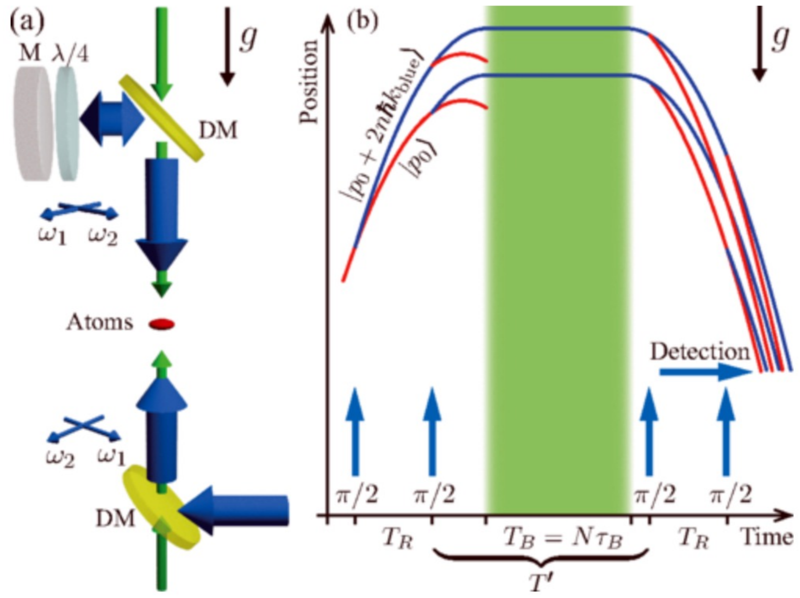
Change of paradigm!

PHYSICAL REVIEW A
covering atomic, molecular, and optical physics and quantum information

Trapped-atom interferometer with ultracold Sr atoms

Xian Zhang, Ruben Pablo del Aguila, Tommaso Mazzoni, Nicola Poli, and Guglielmo M. Tino
Phys. Rev. A **94**, 043608 – Published 4 October 2016

Images



MINIATURIZED ATOM-CHIP GRAVIMETER

Matterwave interferometers based on cold atoms are commonly used as gravimeters. They reach accuracies of up to 10^{-9}g and are nowadays even commercially available.

We have demonstrated a compact quantum gravimeter, which employs an atom chip for the rapid and efficient creation of Bose-Einstein condensates (BEC). At the same time, the atom chip serves for complete state preparation of the atomic cloud and as a retroreflector for the laser beam to create an optical lattice. With the lattice, we split, redirect, and recombine the BEC to form a Mach-Zehnder interferometer and measure the local gravitational acceleration.

To extend the interferometer time and increase the device's sensitivity, we employ the optical lattice for an innovative launch mechanism. In this way, we acquire an intrinsic sensitivity of $\Delta g/g = 10^{-7}$, while keeping all atom-optical operations in a volume of less than a one-centimeter cube.

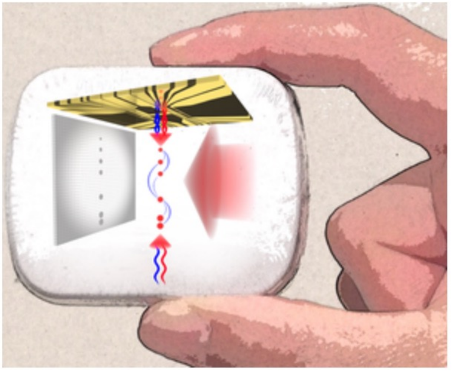


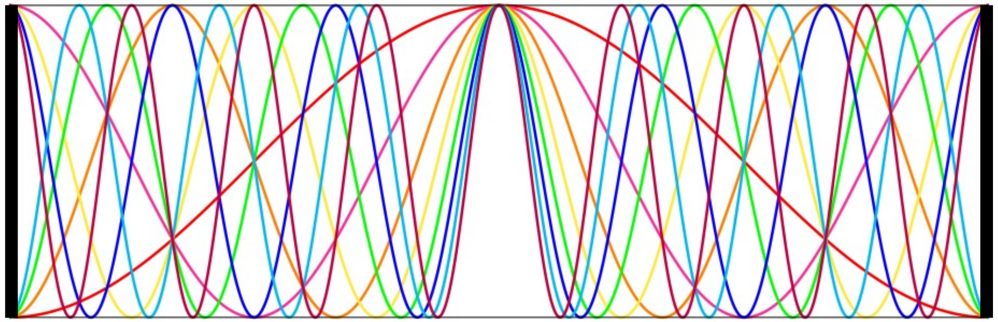
Image by S. Abend and E. Rasel/Leibniz Univ. of Hannover (Physics 9, 131, 2016)

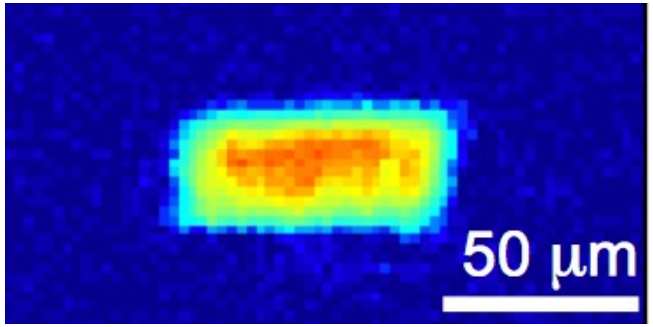
RELATED PUBLICATIONS

S. Abend et. al. *Atom-Chip Fountain Gravimeter* Phys. Rev. Lett. **117**, 203003 (2016)

Quantum frequency interferometry

Howl & Fuentes arXiv:1902.09883





Uses interactions: collective excitations, entanglement between atoms

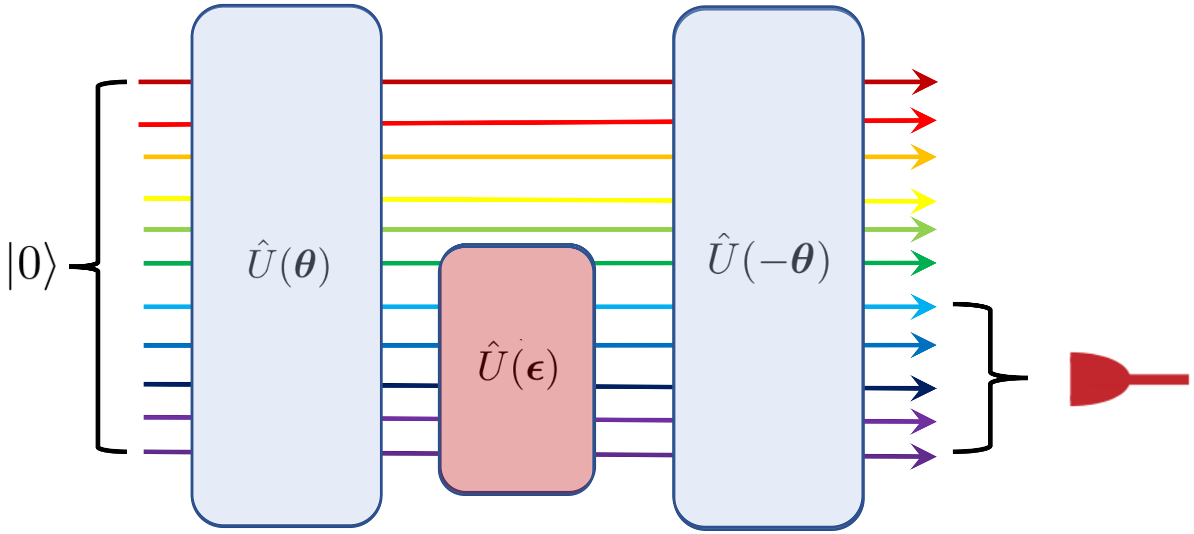
Implementation of frequency modes: phonons in a BEC (massless quantum field)

Interferometry in the frequency (time) domain, non-local

We use parametric amplification produced by the non-linearity introduced by atomic collisions

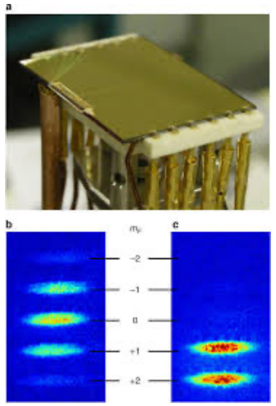
Compatible with General Relativity: underpinned by QFT in curved spacetime

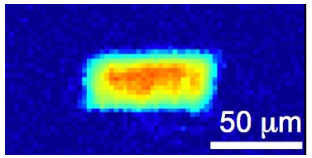
BECs and quantum frequency interferometry



- Detector can be miniaturized
- High sensitivity
- High resilience to noise

Howl & Fuentes arXiv:1902.09883



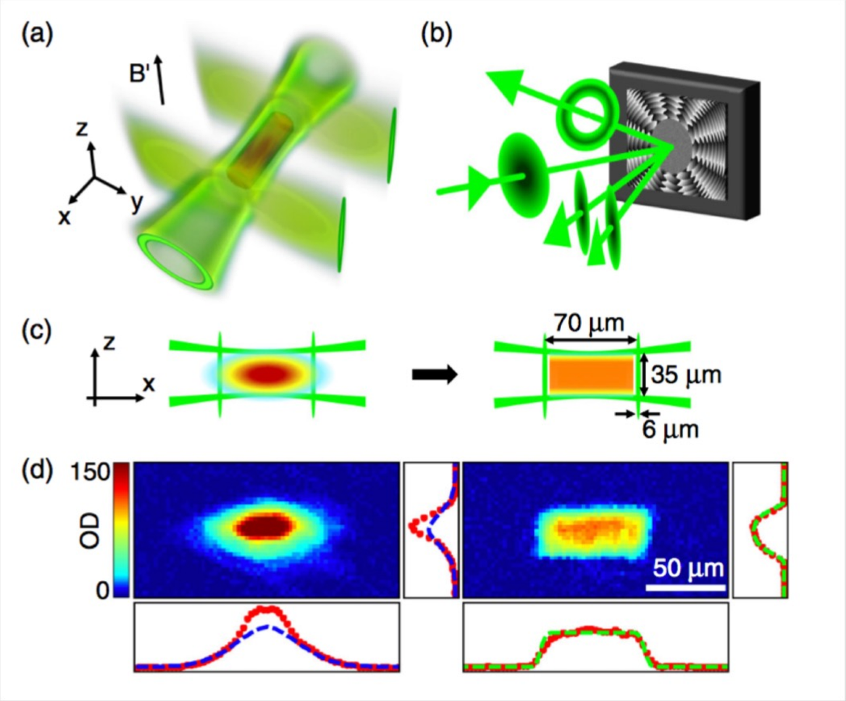


Quantum sensors underpinned by QFTCS

- Continuous source gravitational wave detector
- Quantum relativistic clocks
- Dark energy
- Proper acceleration
- Local gravitational fields (UK patent No.1908538.0)
- Gravitational gradient (UK patent No. 2000112.9)
- Curvature
- Spacetime parameters
- Dark Matter!

Demonstrate particle creation by spacetime dynamics!

Bose Einstein Condensate in a box



mean field (ground state)

$$\hat{\Phi} = \phi(1 + \widehat{\psi})$$

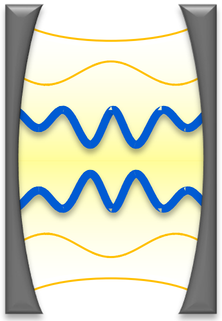
phonons
density fluctuations due to interactions

quasi-uniform density

$$\rho = \phi^{\dagger}\phi$$

Gaunt *et. al.* PRL 110 200406 (2013)

BEC in flat spacetime



$$\mathfrak{g}_{ab} = \left(\frac{n_0^2\, c_s^{-1}}{\rho_0 + p_0}\right) \begin{pmatrix} -c_s^2 & 0 & 0 & 0 \\ 0 & 1 & 0 & 0 \\ 0 & 0 & 1 & 0 \\ 0 & 0 & 0 & 1 \end{pmatrix}$$

Minkowski space but with speed of sound

$$\tau = (c/c_s)t \quad \Rightarrow \quad ds^2 = -cdt^2 + dx^2$$

phonons in a cavity-type 1-dimensional trap

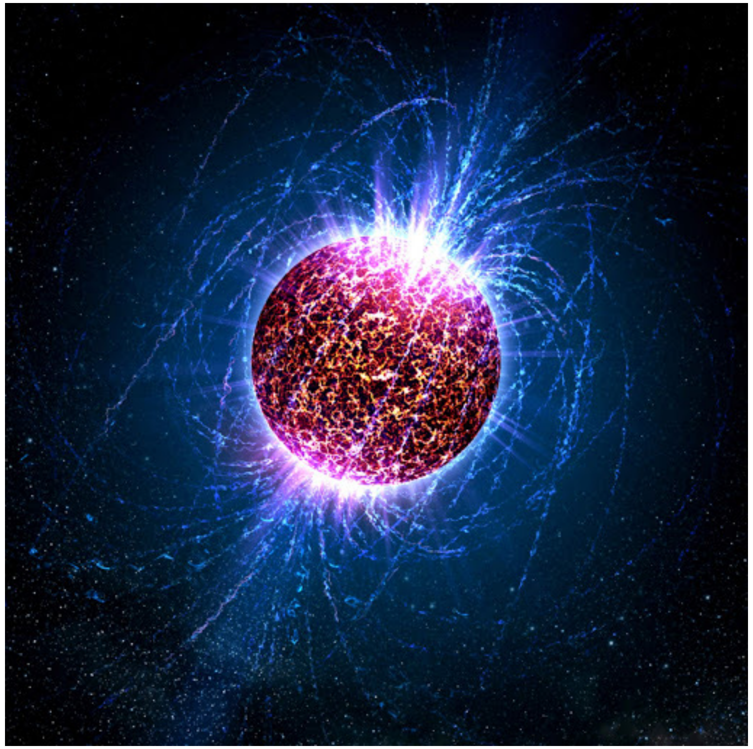
$$\omega_n = \frac{n\,\pi\,c_s}{L}$$ spectrum

$$\Box\phi(t,x) = 0$$

$$\phi_n = \frac{1}{\sqrt{n\,\pi}} \sin\frac{n\pi(x - x_L)}{L}\, e^{-i\,\omega_n\,t}$$

Solutions to the K-G equation

Gravitational wave spacetime



$$g_{\mu\nu} = \eta_{\mu\nu} + h_{\mu\nu}$$

$$h_{\mu\nu} = \begin{pmatrix} 0 & 0 & 0 & 0 \\ 0 & h_+(t) & h_\times(t) & 0 \\ 0 & h_\times(t) & -h_+(t) & 0 \\ 0 & 0 & 0 & 0 \end{pmatrix}$$

$$\mathfrak{g}_{ab} = \left(\frac{n_0^2\, c_s^{-1}}{\rho_0 + p_0}\right) \begin{pmatrix} -c_s^2 & 0 & 0 & 0 \\ 0 & 1 + h_+(t) & h_\times(t) & 0 \\ 0 & h_\times(t) & 1 - h_+(t) & 0 \\ 0 & 0 & 0 & 1 \end{pmatrix}$$

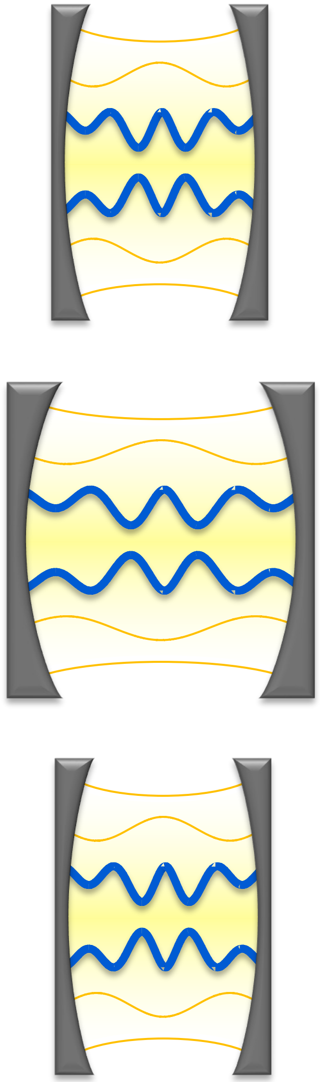
In a one-dimensional trap

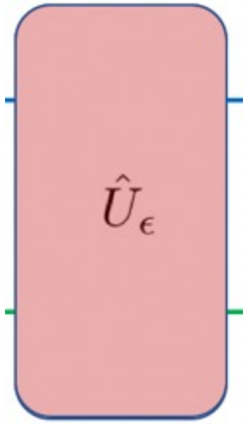
$$ds^2 = -c_s^2\, dt^2 + (1 + h_+(t))\, dx^2.$$

$$h_+(t) = \epsilon \sin \Omega t$$ Continuous sources

$$\omega_n = \frac{n\, \pi\, c_s}{L}$$ Resonance!

Field transformations





Bogoliubov transformations

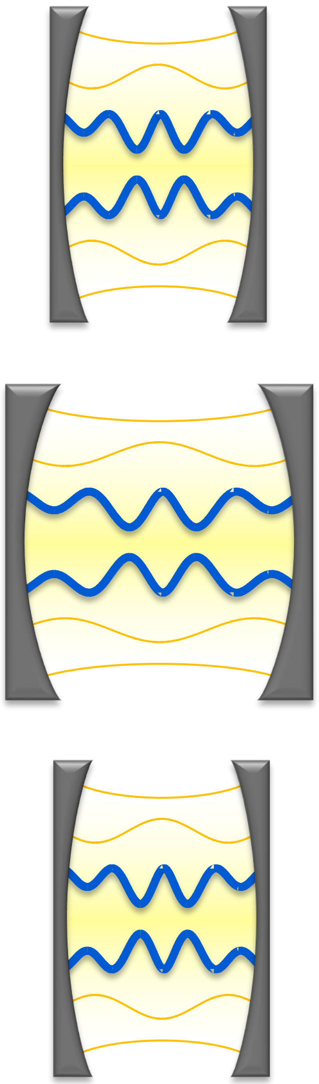
$$\tilde{a}_m = \sum_n \left(\alpha^*_{mn} a_n - \beta^*_{mn} a^\dagger_n \right)$$

$$\alpha_{mn} = (\tilde{\phi}_m, \phi_n) \text{ and } \beta_{mn} = -(\tilde{\phi}_n, \phi^*_m)$$

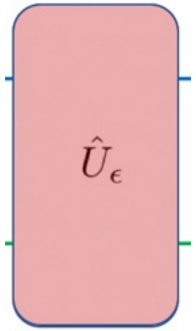
$$\begin{aligned} \beta_{jk}(t) &= -\frac{\epsilon}{2}\sqrt{\frac{n}{m}}\,\omega_m\, t\,[-x_L + (-1)^{m+n}(L + x_L)]\delta_{jm}\,\delta_{kn} + \mathcal{O}(\epsilon^2) \\ \alpha_{jk}(t) &= 0 + \mathcal{O}(\epsilon^2), \end{aligned}$$

Bruschi, Fuentes & Louko PRD (R) 2011

Field transformations



The dynamics of spacetime produces a Bogoliubov transformation on the field modes



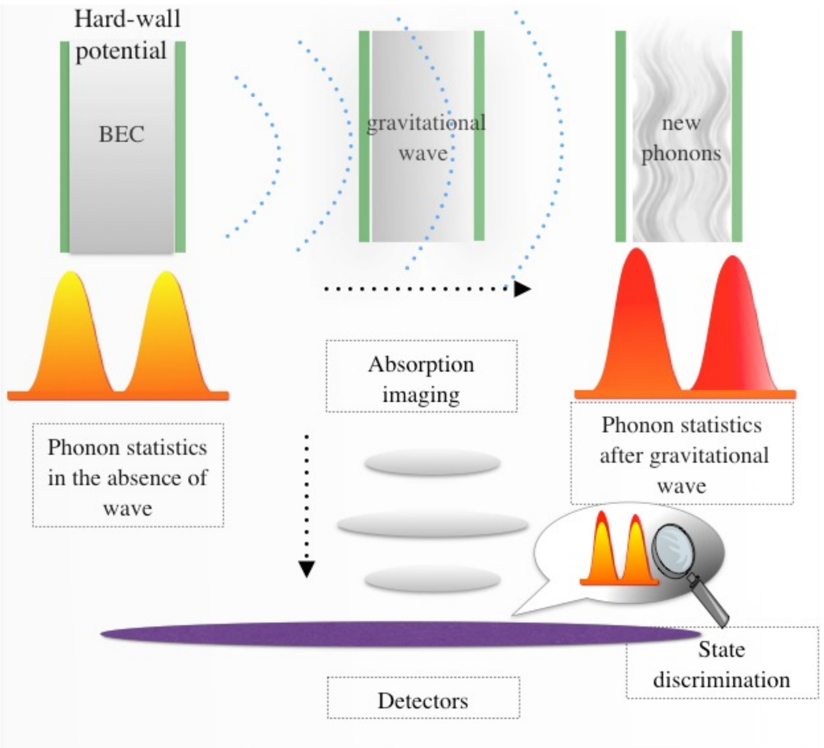
$$\tilde{a}_m = \sum_n \left(\alpha^*_{mn}\, a_n - \beta^*_{mn}\, a^\dagger_n \right)$$

$$\alpha_{mn} = (\tilde{\phi}_m, \phi_n) \text{ and } \beta_{mn} = -(\tilde{\phi}_n, \phi^*_m)$$

Bogoliubov transformation for monocromatic gravitational wave

$$\begin{aligned} \beta_{jk}(t) &= -\frac{\epsilon}{2}\sqrt{\frac{n}{m}}\,\omega_m\, t\,[-x_L + (-1)^{m+n}(L + x_L)]\delta_{jm}\,\delta_{kn} + \mathcal{O}(\epsilon^2) \\ \alpha_{jk}(t) &= 0 + \mathcal{O}(\epsilon^2), \end{aligned}$$

Application: gravitational wave detector



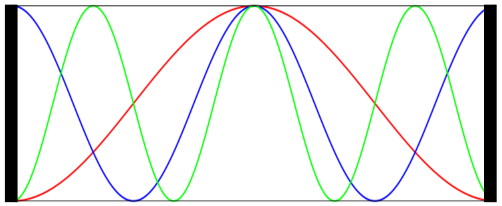
Phonon creation by gravitational waves

Carlos Sabín[1], David Edward Bruschi[2], Mehdi Ahmadi[1] and Ivette Fuentes[1]

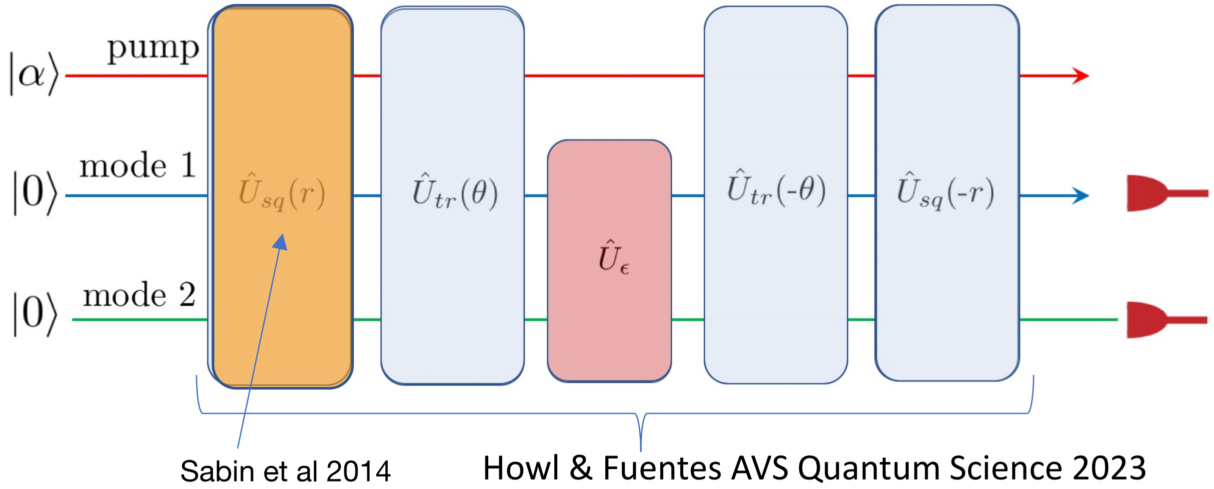
Published 7 August 2014 • © 2014 IOP Publishing Ltd and Deutsche Physikalische Gesellschaft
New Journal of Physics, Volume 16, August 2014

Requires high phonon numbers ~$\sqrt{N_0}$
and long phonon lifetimes ~10s

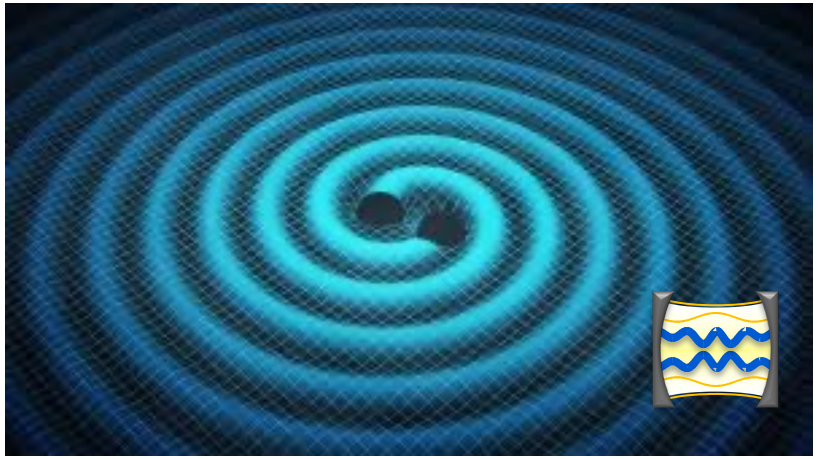
Three mode application



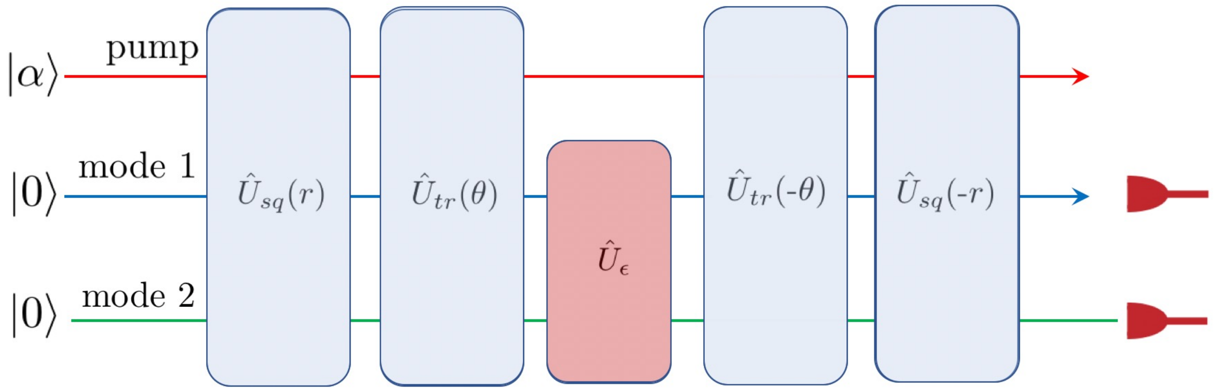
Circuit representation



Improves the sensitivity by several orders of magnitude.
Squeezing can be much smaller than assumed previously and the system can suffer from short phononic lifetimes.



Interferometric transformations



Two-mode squeezing operation

$$\boldsymbol{S}_s = \begin{pmatrix} \mathbf{1} & \mathbf{0} & \mathbf{0} \\ \mathbf{0} & \cosh r \mathbf{1} & \sinh r (\cos\vartheta_{sq}\boldsymbol{\sigma_z} + \sin\vartheta_{sq}\boldsymbol{\sigma_x}) \\ \mathbf{0} & \sinh r (\cos\vartheta_{sq}\boldsymbol{\sigma_z} + \sin\vartheta_{sq}\boldsymbol{\sigma_x}) & \cosh r \mathbf{1} \end{pmatrix},$$

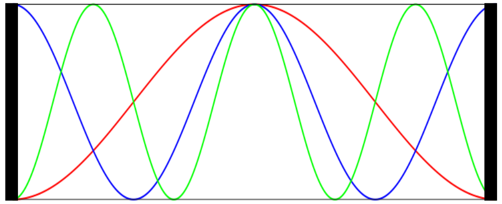
The tritter transformation is

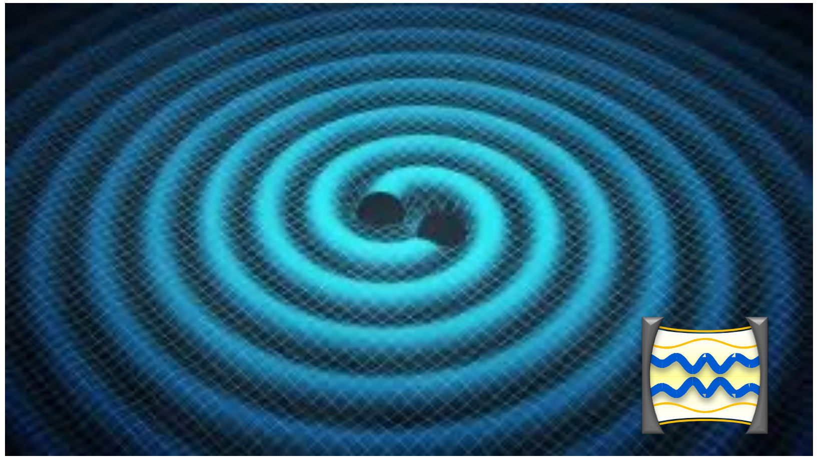
$$\boldsymbol{S}_{tr} = \begin{pmatrix} \cos\theta\mathbf{1} & \frac{1}{\sqrt{2}}\sin\theta(\sin\vartheta\mathbf{1} + i\cos\vartheta\boldsymbol{\sigma_y}) & \frac{1}{\sqrt{2}}\sin\theta(\sin\vartheta\mathbf{1} + i\cos\vartheta\boldsymbol{\sigma_y}) \\ -\frac{1}{\sqrt{2}}\sin\theta(\sin\vartheta\boldsymbol{\sigma_z} - i\cos\vartheta\boldsymbol{\sigma_y}) & \cos^2(\frac{\theta}{2})\mathbf{1} & -\sin^2(\frac{\theta}{2})\mathbf{1} \\ -\frac{1}{\sqrt{2}}\sin\theta(\sin\vartheta\boldsymbol{\sigma_z} - i\cos\vartheta\boldsymbol{\sigma_y}) & -\sin^2(\frac{\theta}{2})\mathbf{1} & \cos^2(\frac{\theta}{2})\mathbf{1} \end{pmatrix}.$$

Howl & Fuentes arXiv:1902.09883

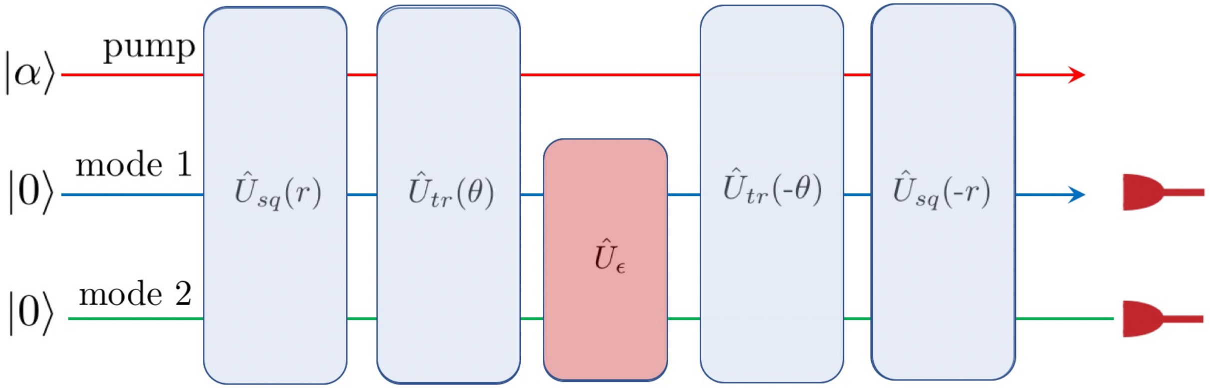
Application: gravitational wave detector

Three mode application





Interferometric transformations



Two-mode squeezing operation

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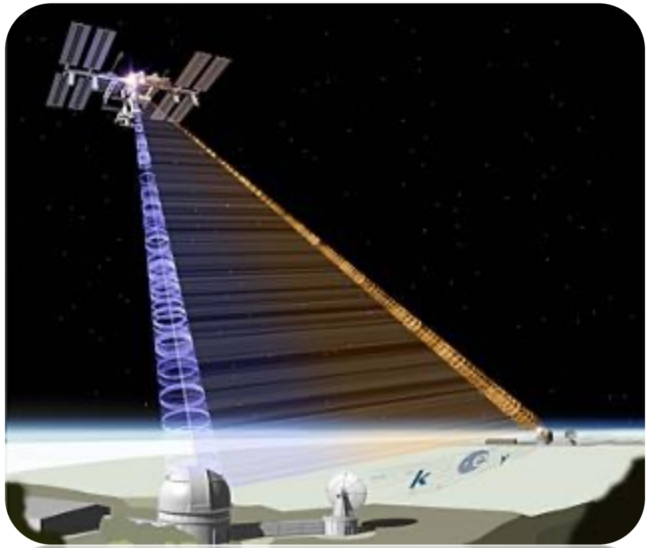
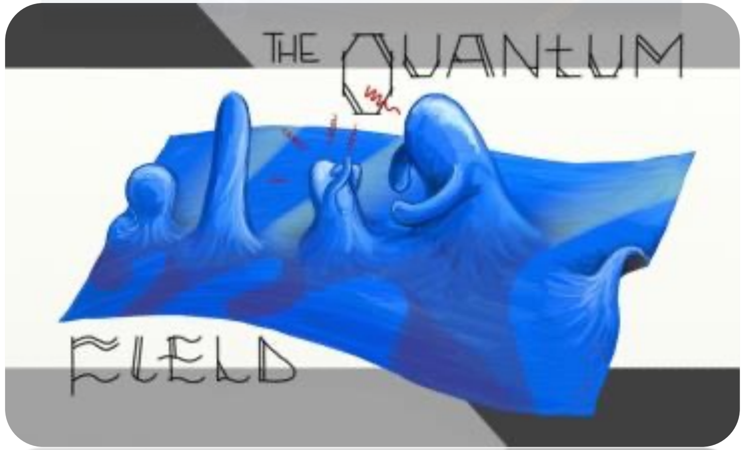
The tritter transformation is

$$\boldsymbol{S}_{tr} = \begin{pmatrix} \cos\theta\mathbf{1} & \frac{1}{\sqrt{2}}\sin\theta(\sin\vartheta\mathbf{1} + i\cos\vartheta\boldsymbol{\sigma_y}) & \frac{1}{\sqrt{2}}\sin\theta(\sin\vartheta\mathbf{1} + i\cos\vartheta\boldsymbol{\sigma_y}) \\ -\frac{1}{\sqrt{2}}\sin\theta(\sin\vartheta\boldsymbol{\sigma_z} - i\cos\vartheta\boldsymbol{\sigma_y}) & \cos^2(\frac{\theta}{2})\mathbf{1} & -\sin^2(\frac{\theta}{2})\mathbf{1} \\ -\frac{1}{\sqrt{2}}\sin\theta(\sin\vartheta\boldsymbol{\sigma_z} - i\cos\vartheta\boldsymbol{\sigma_y}) & -\sin^2(\frac{\theta}{2})\mathbf{1} & \cos^2(\frac{\theta}{2})\mathbf{1} \end{pmatrix}.$$

Howl & Fuentes Howl & Fuentes AVS Quantum Science 2023

Quantum metrology



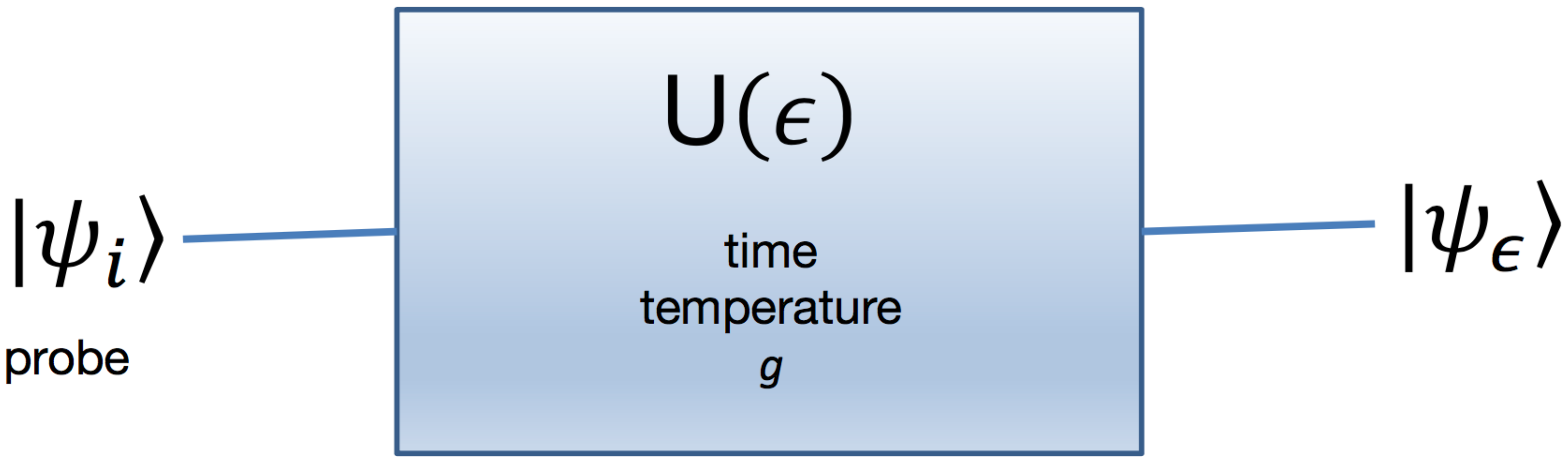


- Enables ultrasensitive devices for measuring fields, frequencies, time
- Quantum clocks and sensors are being sent to space… relativity cannot be ignored

Used to measure gravitational parameters…

gravitational field strengths
accelerations

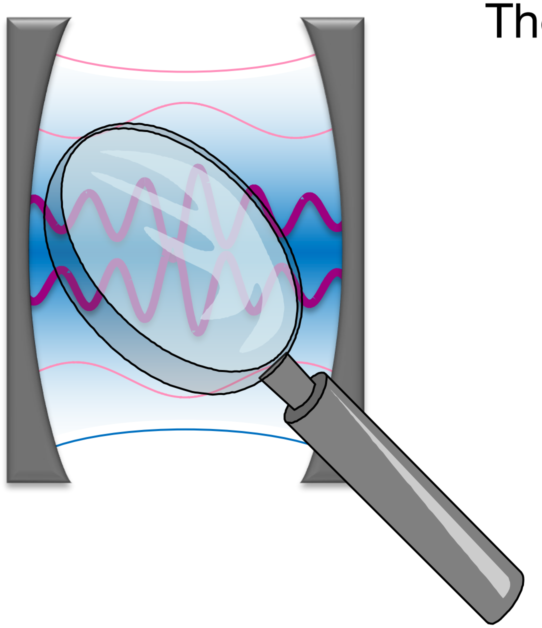
Quantum Metrology



$$\langle\psi_\epsilon|\psi_{\epsilon+d\epsilon}\rangle \ll 1$$

Exploit quantum properties of the probe state to estimate with high precision parameters in the theory (Hamiltonian)

Quantum Metrology



There is an optimal measurement such that

parameter

$$\langle(\Delta\hat{\epsilon})^2\rangle \geq \frac{1}{M H_\epsilon}$$

Error

Quantum Fisher information
M number of measurements

$$H_\epsilon = \frac{8\left(1 - \sqrt{\mathcal{F}(\sigma_\epsilon, \sigma_{\epsilon+d\epsilon})}\right)}{d\epsilon^2}$$

Fidelity

state σ_ϵ

$\langle\psi_\epsilon|\psi_{\epsilon+d\epsilon}\rangle \ll 1$

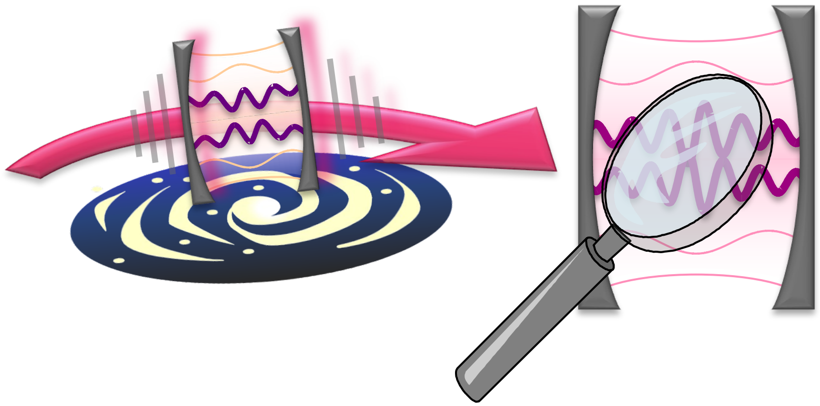
General framework for RQM

Ahmadi, Bruschi, Sabin, Adesso, Fuentes, Nature Sci. Rep. 2014
Ahmadi, Bruschi, Fuentes PRD 2014

Fisher information in QFT:
Analytical formulas in terms of
general Bogoliubov coefficients

for small parameters

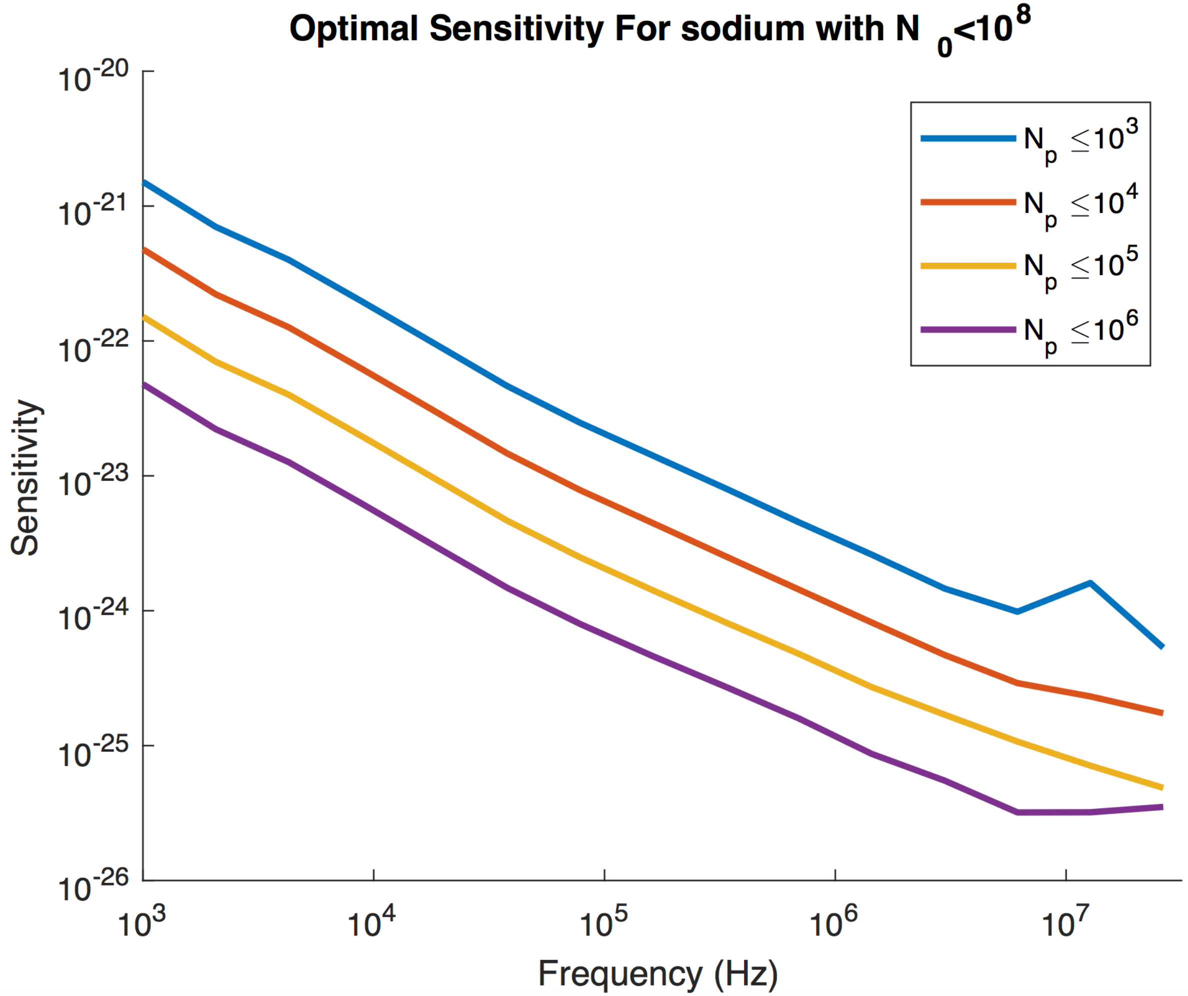
$$
\begin{aligned}
H &= \epsilon^{-2}\Re\Big[4\cosh r(f_\alpha^n + f_\beta^n + f_\alpha^m + f_\beta^m)\\
&+ 4\cosh^2 r(2|\beta_{nm}(t)|^2 - f_\alpha^n + f_\beta^n - f_\alpha^m + f_\beta^m)\\
&- 4\sinh^2 r(-f_\alpha^n + f_\beta^n - f_\alpha^m + f_\beta^m + 2\beta_{nm}(t)^2 - 2\alpha_{nm}(t)^2)\\
&+ 4\sinh r\,\Re[\mathcal{G}^{\alpha\beta}_{nm} + \mathcal{G}^{\alpha\beta}_{nm}] - 4\cosh^4 r|\beta_{nm}(t)|^2\\
&- \frac{1}{2}\sinh^2 2r(2|\alpha_{nm}(t)|^2 - 3|\beta_{nm}(t)|^2 - \beta_{nm}(t)^2\Big].
\end{aligned}
$$



$$
\begin{aligned}
f_\alpha^i &= \frac{1}{2}\sum_{n\neq k,k'}|\alpha_{ni}|^2\\
f_\beta^i &= \frac{1}{2}\sum_{n\neq k,k'}|\beta_{ni}|^2\\
\mathcal{G}^{\alpha\beta}_{ij} &= \sum_{n\neq k,k'}\alpha_{ni}\beta^*_{nj}
\end{aligned}
$$

Single-mode
Two-mode channels

Detector sensitivity



$$\Delta\epsilon \geq \frac{1}{\sqrt{MF_Q}}$$

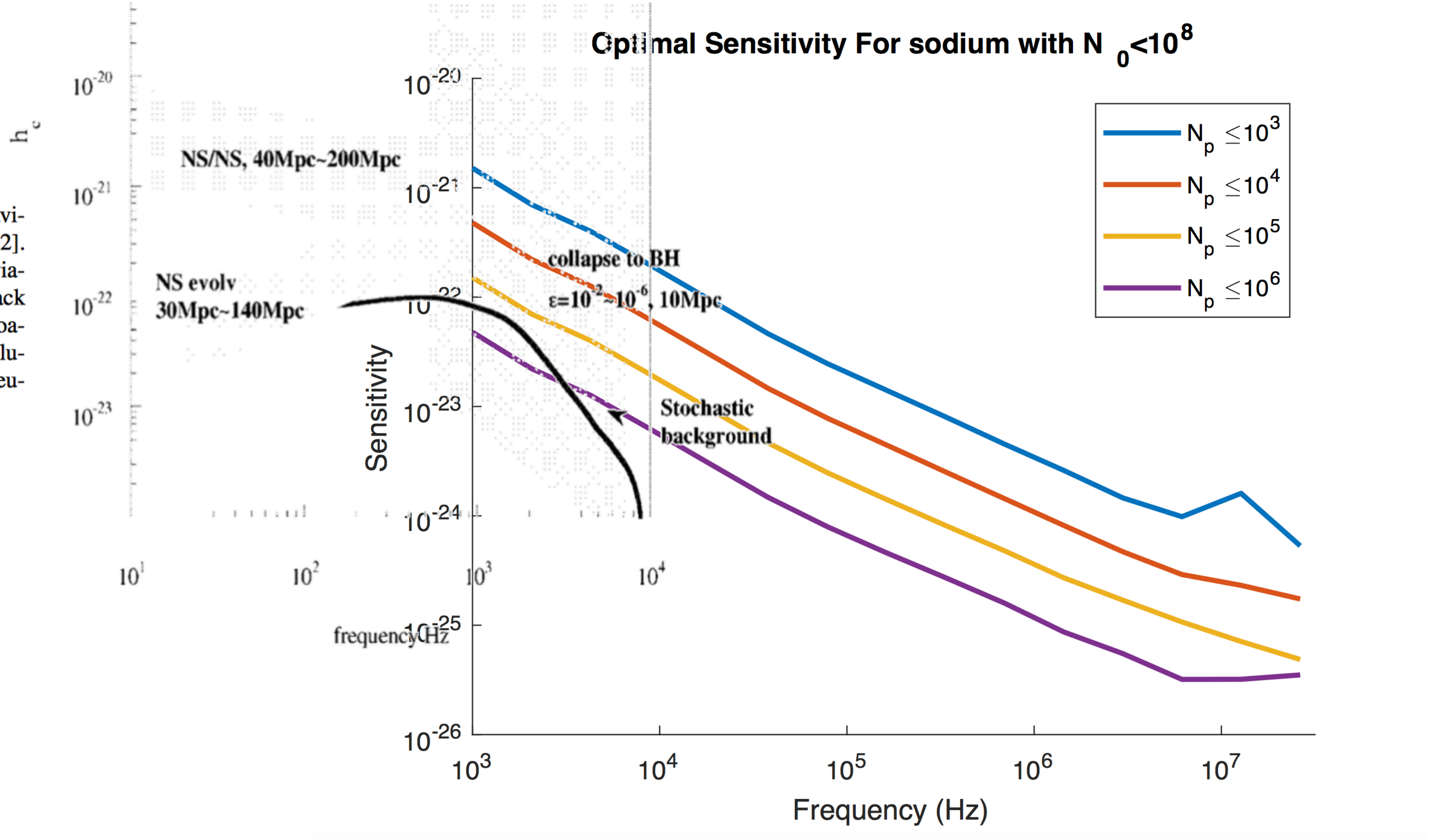
$$\Delta\epsilon \geq \frac{m}{\sqrt{2\pi}\hbar}\frac{\alpha^3}{\theta N_0^2\sqrt{N_p\tau t}}\sqrt{\frac{L^7}{a^3}}\frac{\sqrt{nl}\,(l-n)^2}{(l^2+n^2)}$$

m atomic mass
θ tritter angle
N_0 number of atoms in the ground state
N_p number of phonons in the two-mode squeezed state
$\alpha = \sqrt{A}/L$, A area, L length
a scattering length
$M = \tau/t$ number of measurements
τ integration time
t interaction time (lifetime of the phonons)
$\Omega = \omega_n + \omega_l$ the frequency of the gw
$\omega_j = j\pi c_s/L$ phonon frequency
n, l mode numbers

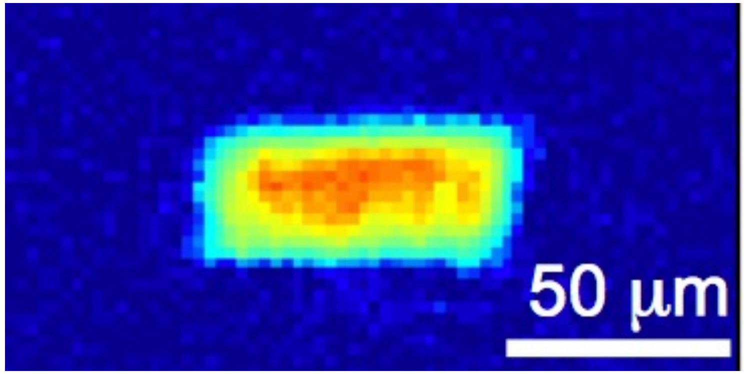
Howl & Fuentes arXiv:1902.09883

Detector sensitivity

Figure 8. Spectrum of gravitational wave sources [18, 22]. In this figure, the abbreviations are: BH, collapse to black hole; NS/NS, neutron star coalescence; NS evol, secular evolution of a nonaxisymmetric neutron star.



Constraints



	$(\gamma^{La}_{k_BT\ll\mu})^{-1}$ (1.56)	$(\gamma^{La}_{k_BT\gg\mu})^{-1}$ (1.56)	$(\gamma^{Be,0})^{-1}$ (1.55)	$t_{1/2}$ (3.3)
$t \approx$	$\frac{640}{3\pi}\frac{\hbar^7\beta^4 L n_0^3 a_s^2}{l m^3}$	$\frac{8}{\pi}\frac{\beta L}{a_s l}$	$\frac{640}{3\pi^6}\frac{m L^5 n_0}{\hbar l^5}$	$\frac{1}{50}\frac{m}{\hbar n_0^2 a_s^4}$
$r \propto$	$\frac{\hbar^8\beta^4 L^2 n_0^{9/2} a_s^{7/2}}{l^2 m^4}$	$\frac{\hbar^2 L^2 \beta n_0^{3/2} a_s^{1/2}}{l^2 m}$	$\frac{L^6 n_0^{5/2} a_s^{3/2}}{l^6}$	$\frac{L}{l n_0^{1/2} a_s^{5/2}}$
$\beta\omega_l$	$2\sqrt{\pi}\frac{\hbar l\sqrt{n_0 a_s}}{Lm}$		weak interactions	$\lvert a_s\rvert n_0^{1/3} \ll 1$
phonon regime	$\sqrt{\frac{\pi l^2}{L^2 n_0 a_s}} \ll 1$		ultracold regime	$\frac{1}{4\pi}\frac{m}{\hbar^2\beta\, n_0\, a_s} \ll 1$

Table 8.2: Approximate values for the limitations to the phonon and condensate life times and resulting proportionalities of the maximum two-mode squeeze factor from SECTION 8.1 in the case where the respective damping or particle losses become dominant.
The two bottom rows show a measure $\beta\omega_l$ for the thermal occupation of the initial state, which should be minimal, and restrictions from the diluteness condition and the assumptions of the phononic and the ultracold regime.
To see where one could optimize and which boundaries will be encountered, all quantities above are given in terms of the parameters characterizing an individual experiment.

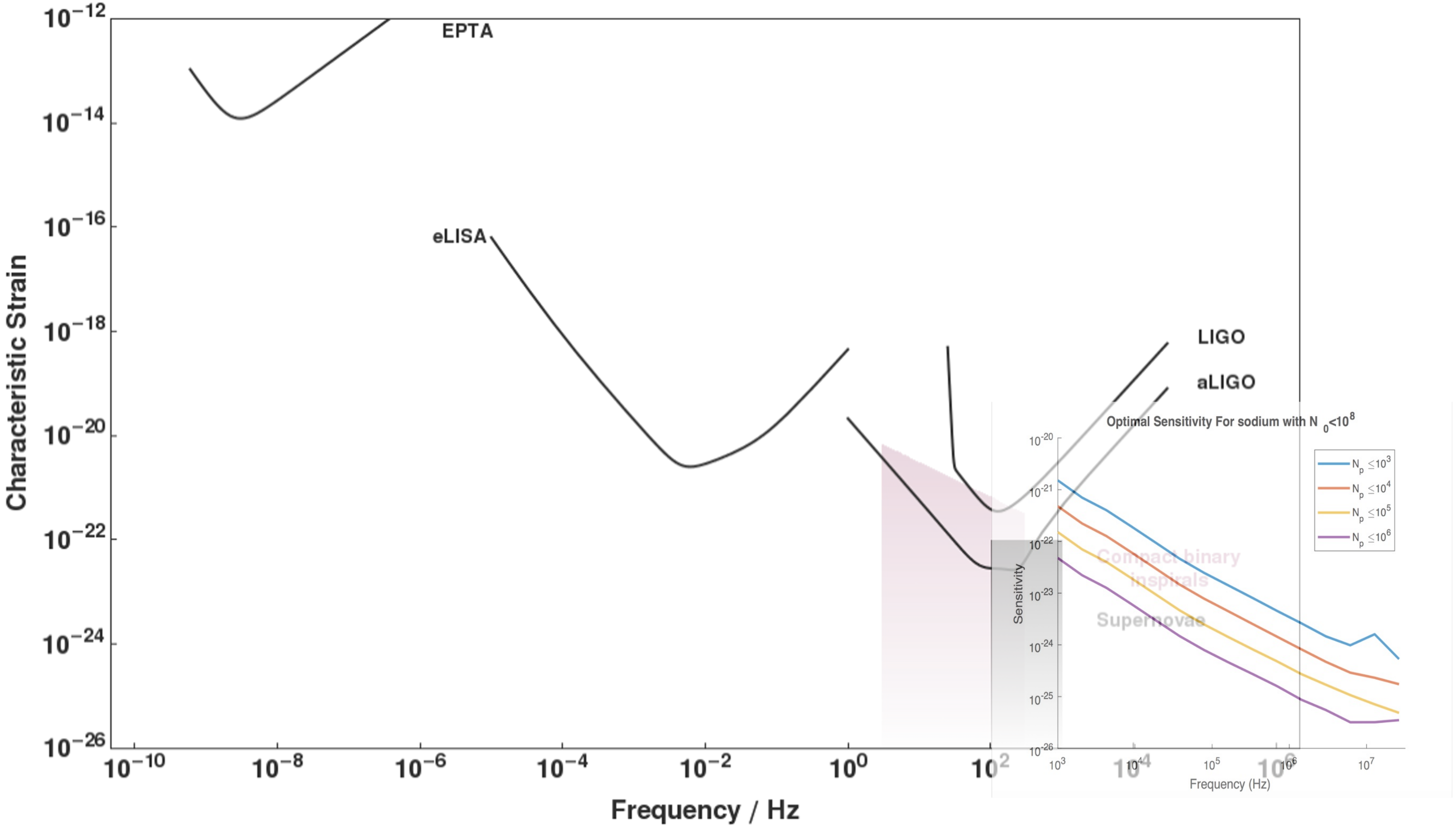
Quantum decoherence of phonons in Bose–Einstein condensates

Richard Howl[1], Carlos Sabín[2], Lucia Hackermüller[3] and Ivette Fuentes[1,4]

Published 29 November 2017 • © 2017 IOP Publishing Ltd
Journal of Physics B: Atomic, Molecular and Optical Physics, Volume 51, Number 1
Citation Richard Howl *et al* 2018 *J. Phys. B: At. Mol. Opt. Phys.* **51** 015303

P. Juschitz, Two-mode Phonon Squeezing in Bose-Einstein Condensates for Gravitational Wave Detection arXiv:2101.05051

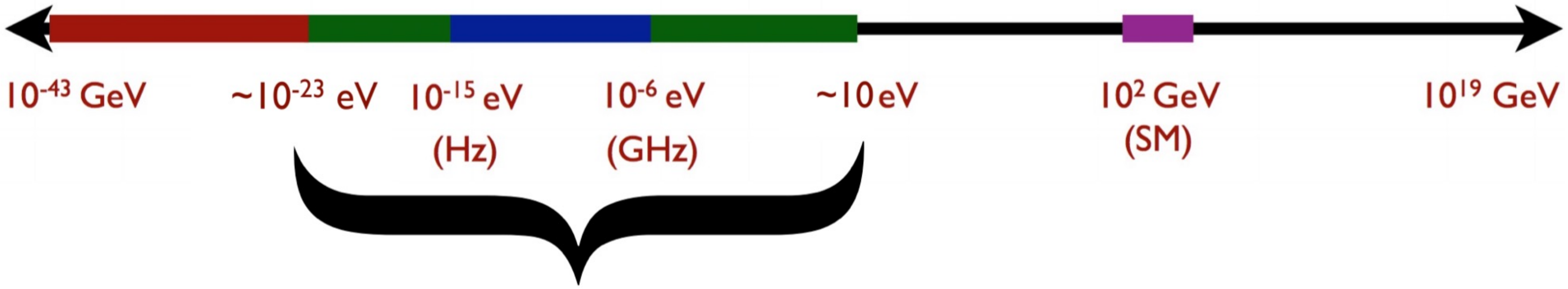
Search for yet unknow sources



Know sources: 10 kHz

Exotic sources:
primordial black holes
boson stars
Early Universe Cosmology
Phase transitions
Preheating after inflation,
Cosmic strings
Dark matter
Ultralight 10^{-8}- 10^{14} Hz
Decay: Penrose CCC

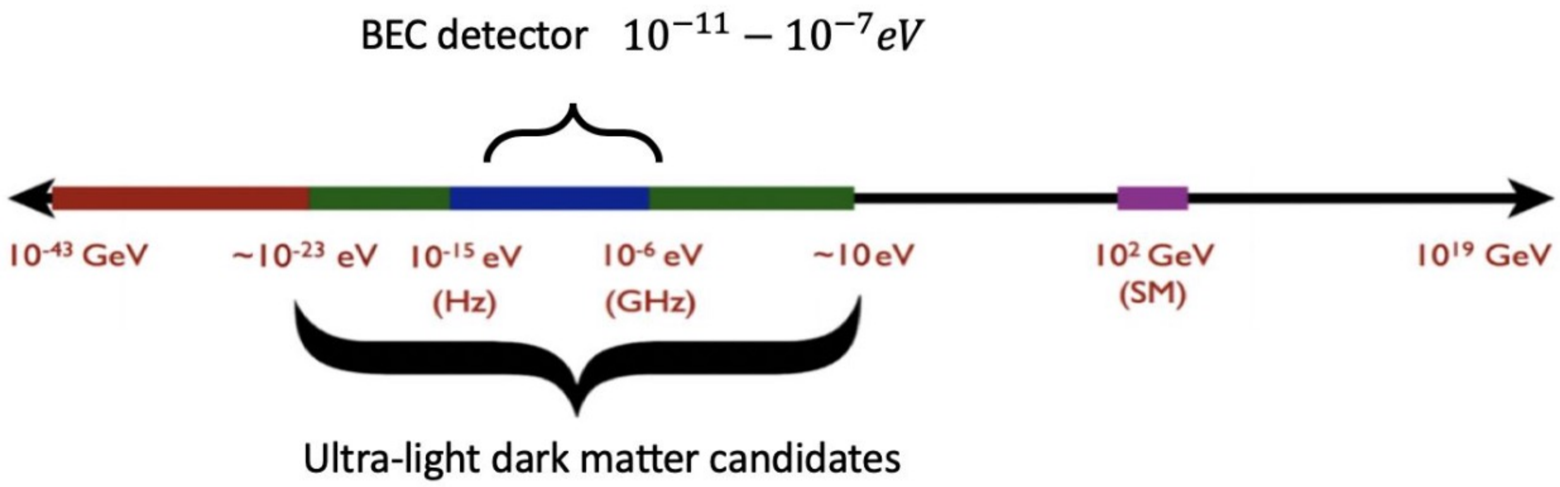
Ultra-Light Dark Matter (bosonic)



→ search for *coherent effects of the entire field,* not single hard particle scatterings

Generic Candidates: Light pseudo-Nambu-Goldstones (axions and "axion like particles" — ALPs); Massive hidden vector bosons (aka "dark photons"); Light scalars (moduli/dilatons…)

Slide by John March-Russel


BEC detector $10^{-11} - 10^{-7} eV$
10^{-43} GeV
$\sim 10^{-23}$ eV
10^{-15} eV
(Hz)
10^{-6} eV
(GHz)
~ 10 eV
10^{2} GeV
(SM)
10^{19} GeV
Ultra-light dark matter candidates

Different principle

Interferometer arm length L

$$\frac{\Delta L}{L}$$

Resonance

$$\Omega = \omega_n + \omega_m$$

wave

quantum excitations

Quantum Weber Bar

Temperature

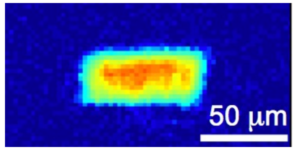
Weber bar

$T \sim 4$ K

BEC

$T \sim 5 x 10^{-10} K$

Initial quantum states
Squeezing
Parametric amplification



How can it work if its so small?

$$\omega \sim \frac{C_s}{L}$$

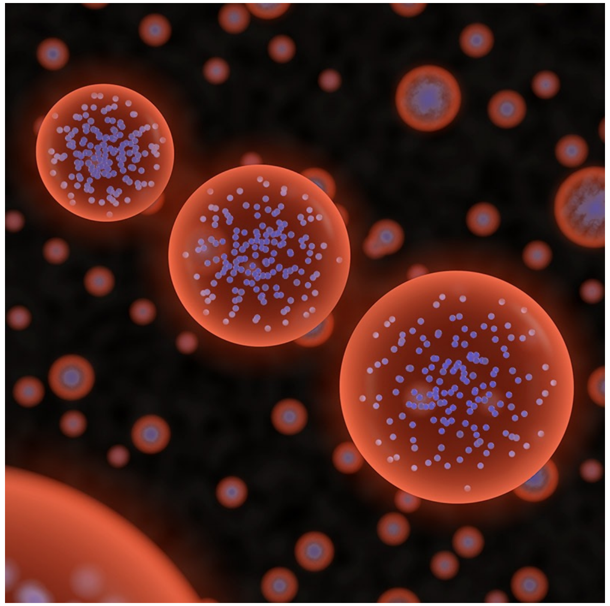
$$\omega \sim \frac{C}{L}$$

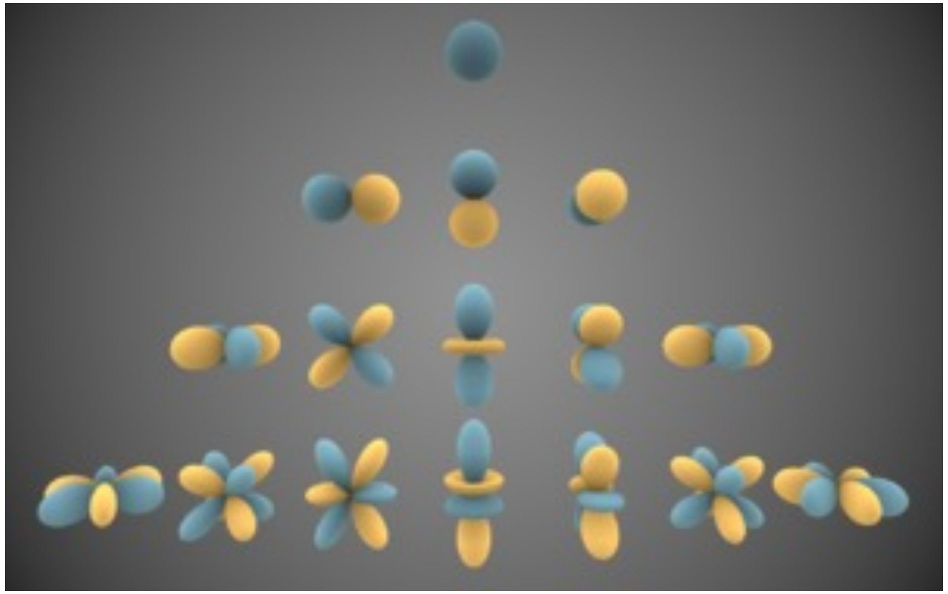
Speed of sound: $C_s = 10\text{mm/s}$
$L = 10^{-1}$-10^{-3}mm

Speed of light: C= $2.99\text{x}10^{11}$mm/s
L =2.99Km-2990Km

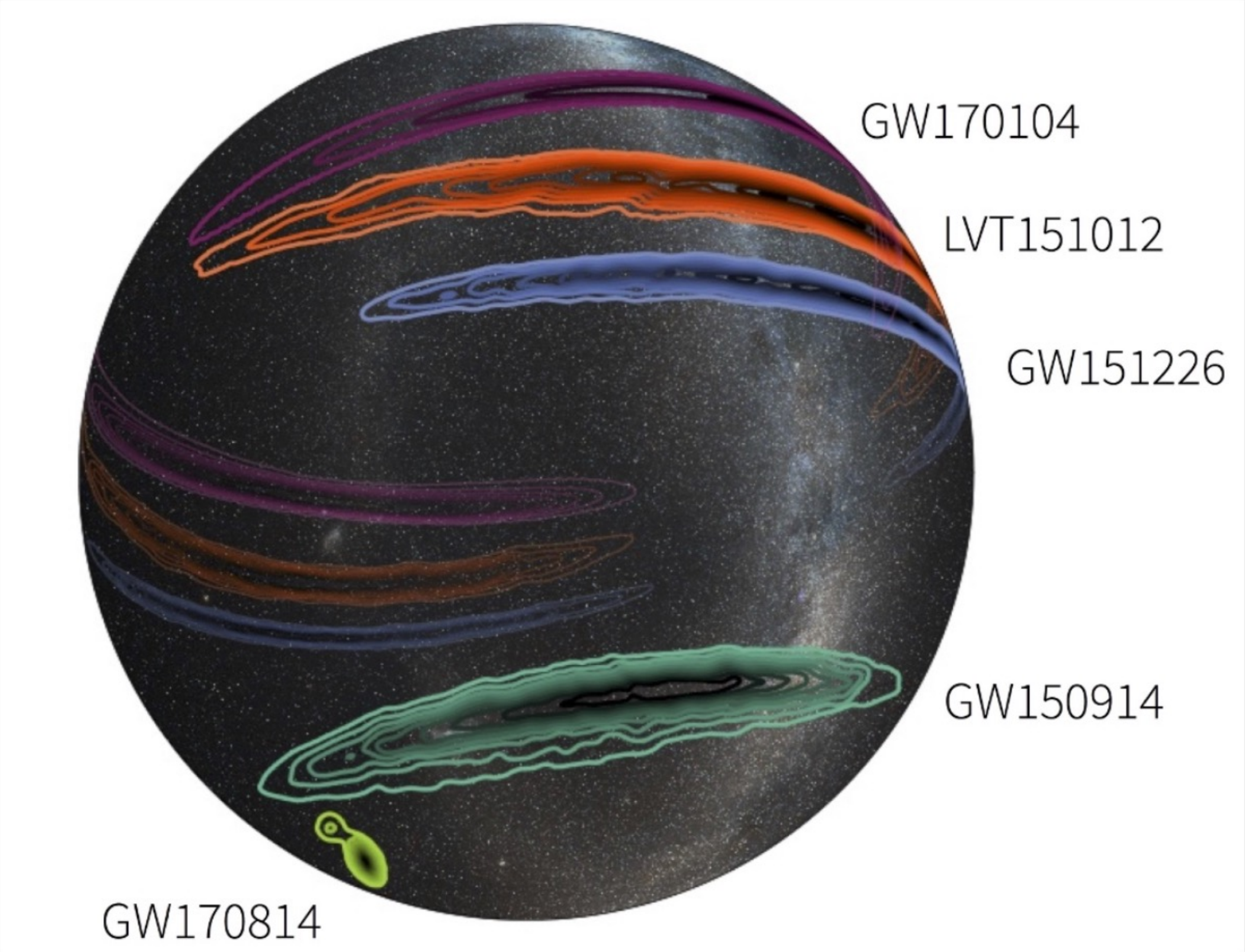
Broadband: 3D

Spherical BEC





We are studying different geometries

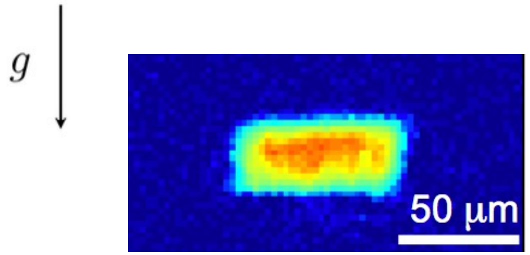


This three-dimensional projection of the Milky Way galaxy onto a transparent globe shows the probable locations of the three confirmed black-hole merger events observed by the two LIGO detectors—GW150914 (dark green), GW151226 (blue), GW170104 (magenta)—and a fourth confirmed detection (GW170814, light green, lower-left) that was observed by Virgo and the LIGO detectors. Also shown (in orange) is the lower significance event, LVT151012. Image credit: LIGO/Virgo/Caltech/MIT/Leo Singer (Milky Way image: Axel Mellinger).

Commercial applications

$$\boldsymbol{g} = \mathrm{diag}\left(-f(r), \frac{1}{f(r)}, r^2, r^2 \sin^2\vartheta\right)$$

$$f(r) = 1 - r_S/r$$



Phononic gravimeter:
Same sensitivity but much smaller system.

Phononic gradiometer:
Improves the state of the art by at least two orders of magnitude

1. WO2020249974 - QUANTUM GRAVIMETERS AND GRADIOMETERS

PCT Biblio. Data Description Claims Drawings ISR/WOSA/A17(2)(a) National Phase Notices Documents

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Part 1: 1 2 3 4 5 6 **Part 2:** A B C D E

PATENT COOPERATION TREATY
PCT
INTERNATIONAL SEARCH REPORT
(PCT Article 18 and Rules 43 and 44)

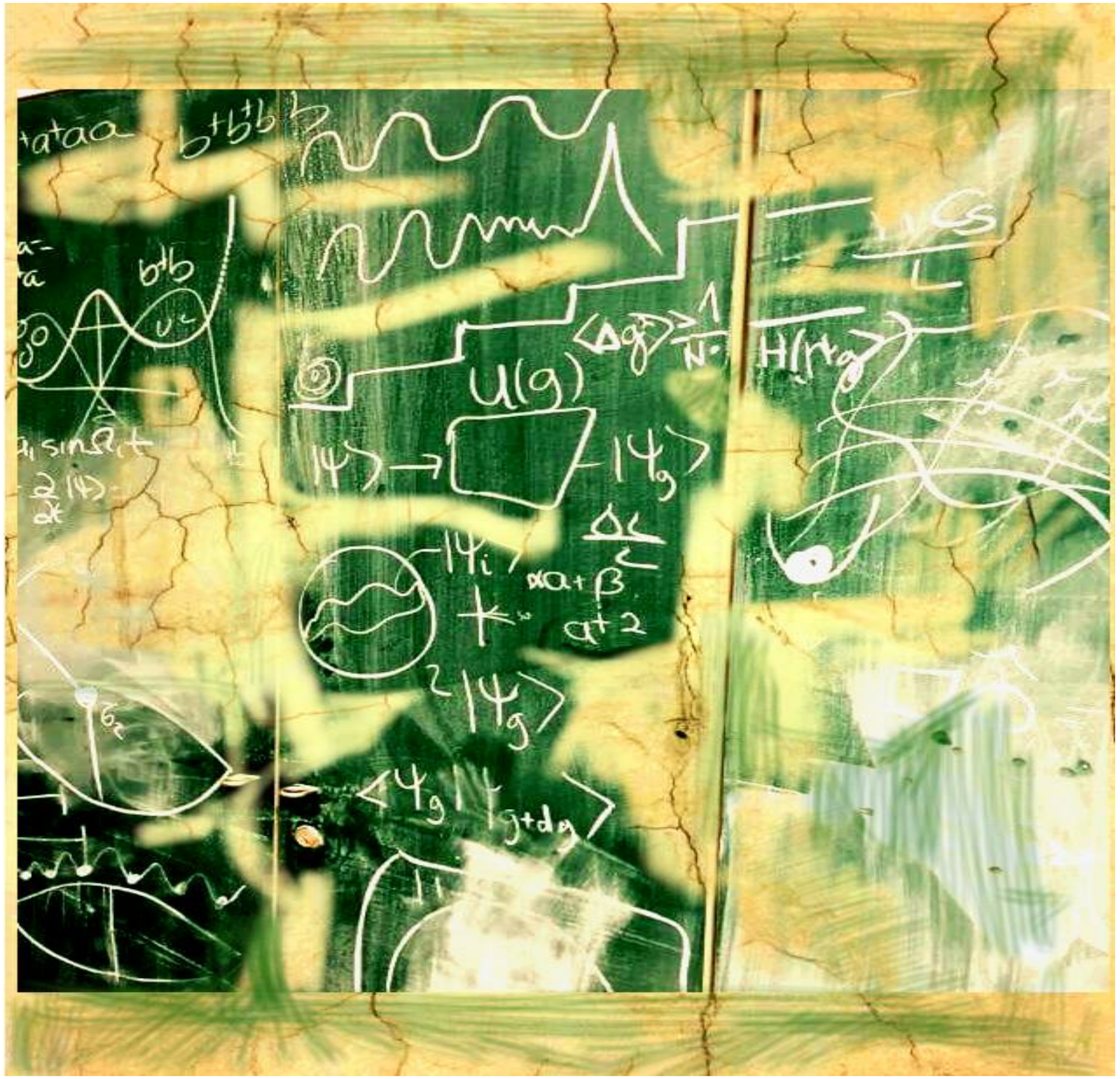
International application No. PCT/GB2020/051434	Applicant's or agent's file reference DJC96140P.WO
International filing date (day/month/year) 12 June 2020	(Earliest) Priority Date (day/month/year) 13 June 2019

Setup	Running time	Length	$\Delta r_S/r_S$
[14, 15, 17] Atom. Int.	$100\,\mathrm{s} - 8\,\mathrm{h}$	$0.2 - 2.5\,\mathrm{m}$	10^{-9}
[16] BEC-chip	$100\,\mathrm{s}$	$10^{-2}\,\mathrm{m}$	10^{-10}
Phononic MZI	$6\,\mathrm{s}$	$10^{-4}\,\mathrm{m}$	10^{-8}

Quantum technologies could be useful for Fundamental Physics

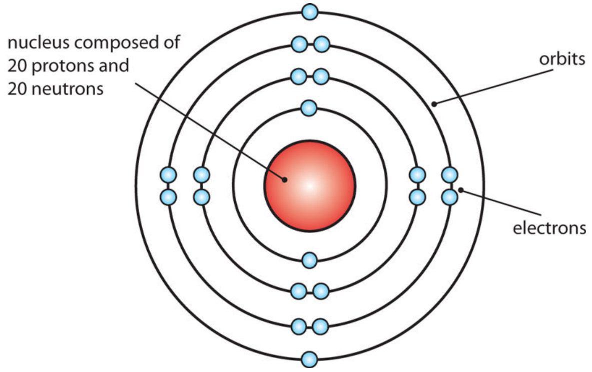
Main aim: Develop new instruments
window to scales that have not been explored before
New Physics!

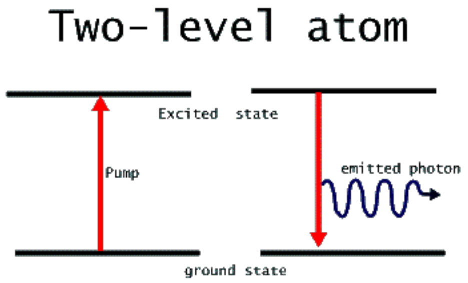
Quantum technologies can help us put the pieces of the puzzle together



Relativistic effects
on
quantum clocks

Quantum clocks (quantum 1.0)





Optical clocks routinely achieving 10^{-17} - 10^{-18} systematic uncertainty

Time

Quantum mechanics

Time is absolute (Galilean trans.).

Space and time are different.

- Time is a parameter.
- Space is an operator.

Particles can be in a superposition of positions at once.

Relativity

Space and time are not different.

Time is observer dependent

Time flows at different rates in different points in space (Lorentz trans.)

Proper time

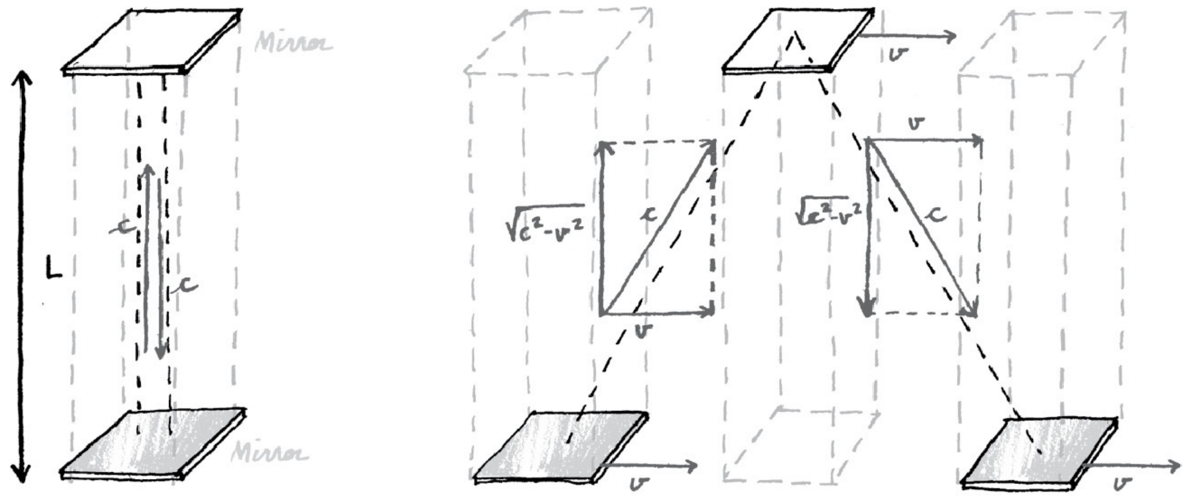
$$\Delta\tau = \int_P d\tau = \int_P \frac{1}{c}\sqrt{g_{\mu\nu}\, dx^\mu\, dx^\nu}$$

metric

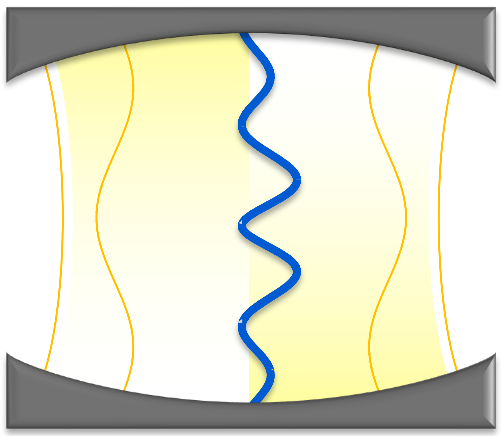
Does quantum theory imposes fundamental limitations on how well we can measure proper time?

Einstein's light clock

Classical light



Quantum: photons



Lindkvist, Sabin, Fuentes, Dragan, Svensson, Delsing, Johansson PRA (2014)

Relativistic quantum clock model

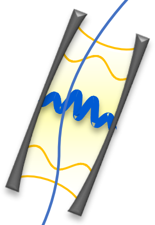
Open Access | Published: 19 May 2015

Motion and gravity effects in the precision of quantum clocks

Joel Lindkvist, Carlos Sabín, Göran Johansson & Ivette Fuentes

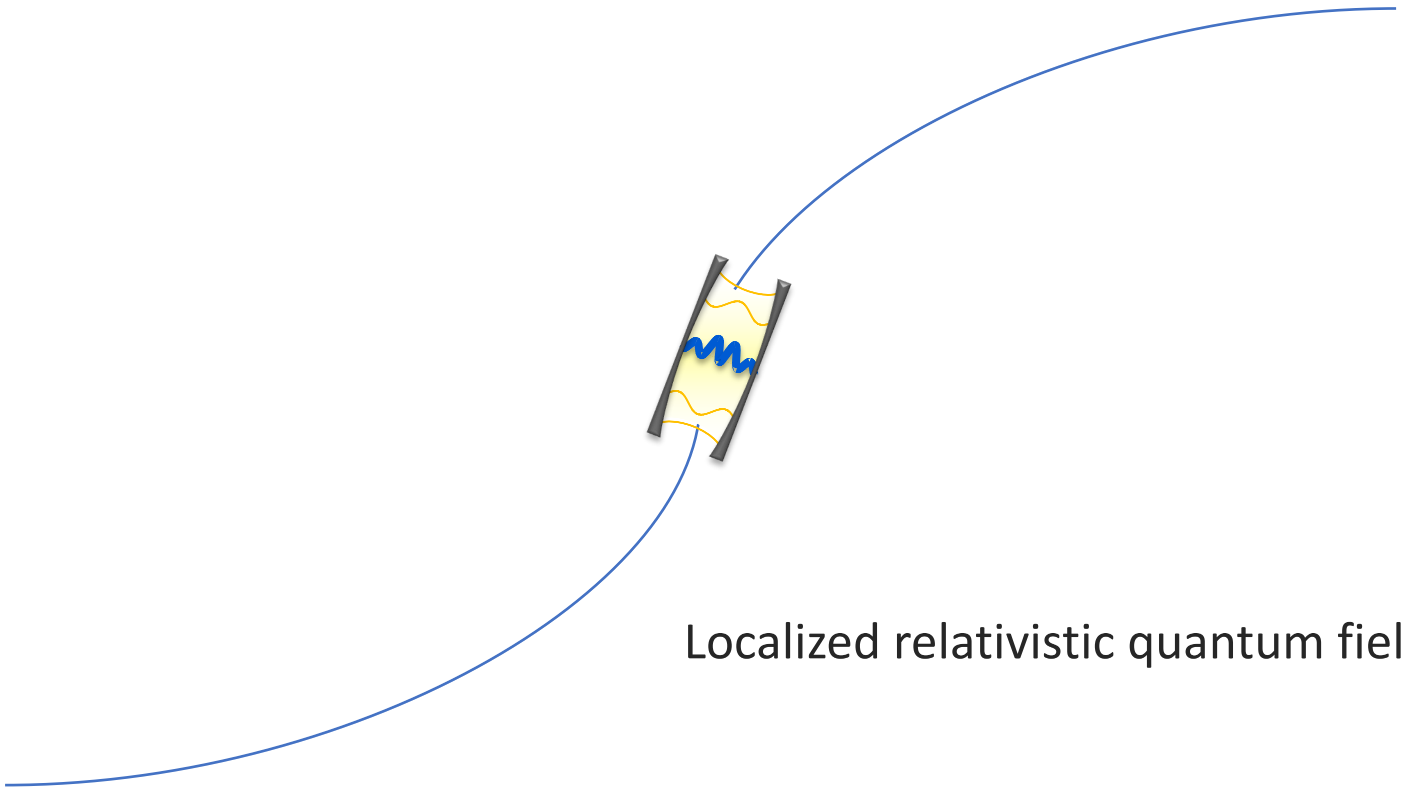
Scientific Reports **5**, Article number: 10070 (2015) | Cite this article

Clock: one mode of the field
in a coherent or squeezed state



How does motion, curvature and gravity affect the clock ticks and the clock's precision?

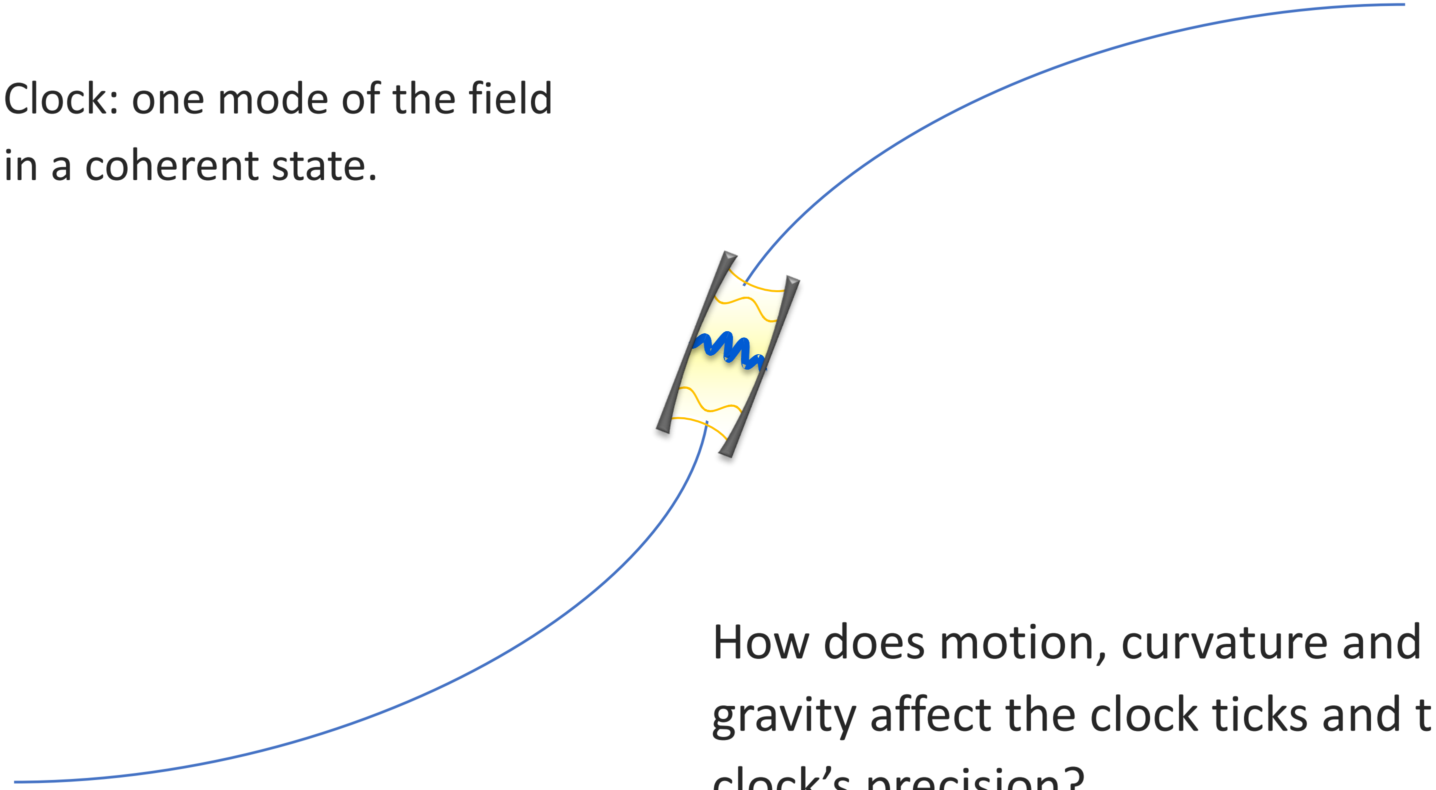
Relativistic quantum clock model



Localized relativistic quantum field

Relativistic quantum clock model

Clock: one mode of the field
in a coherent state.



How does motion, curvature and gravity affect the clock ticks and the clock's precision?

Precision: Quantum Fisher information

Lindkvist, Sabin, Johansson and Fuentes, Sci. Rep. (2015)

$$H_\theta = 4\alpha^2 P\left[\cosh(2r) + \sinh(2r)\cos\phi\right] + \frac{4\sinh^2(2r)}{1+P^2} \qquad P = \frac{1}{4\sqrt{\det\sigma}}$$

Squeezing

Coherent state amplitude

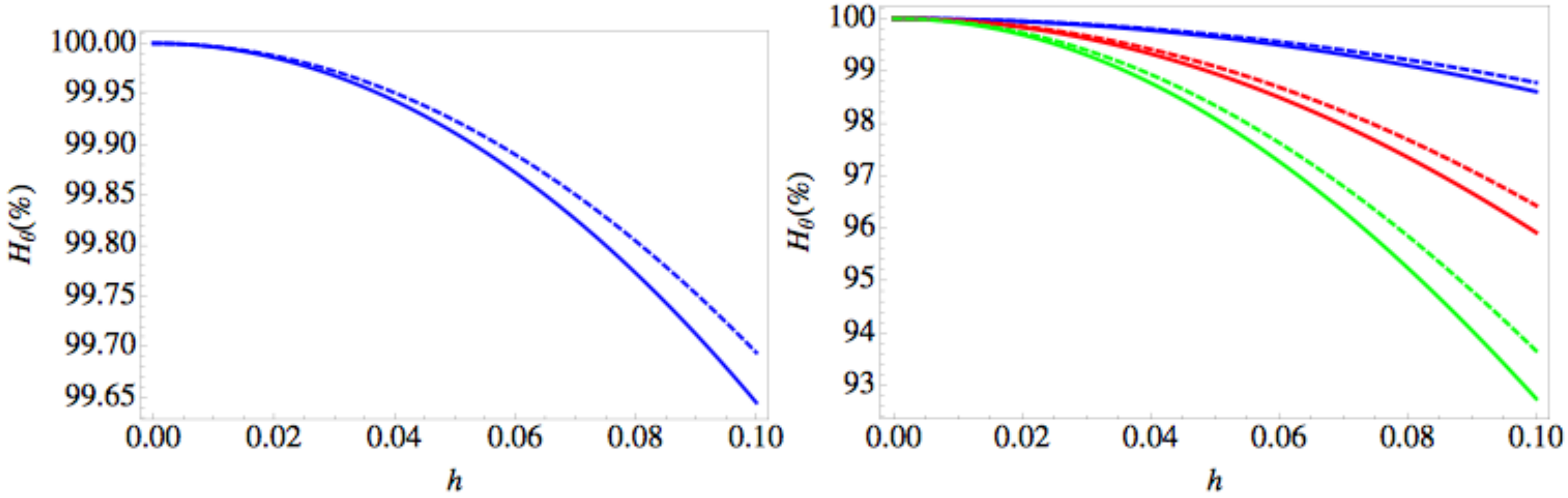
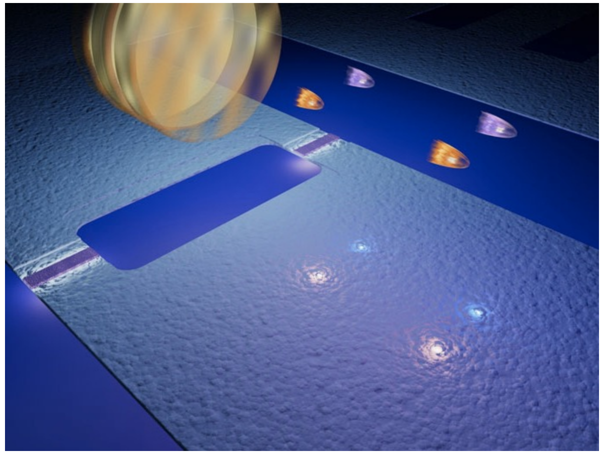
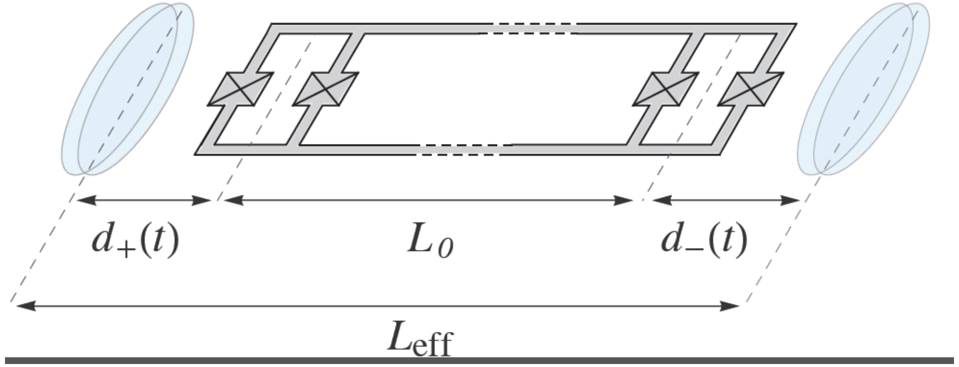


FIG. 2: a) Ratio of the transformed and original QFI for an initial coherent state as a function of h, for $\theta_a = \pi$ and initial photon numbers $N > 1$. The solid(dashed) curve is for $\theta_0 = 0(\pi/2)$). b) Ratio of the transformed and original QFI for an initially squeezed vacuum as a function of h, for $\theta_a = \pi$ and initial photon numbers $N = 1$ (blue), $N = 5$ (red) and $N = 10$ (green). The solid(dashed) curves are for $\theta_0 = 0(\pi/2)$).

Implementing the twin-paradox

Lindkvist, Sabin, Fuentes, Dragan, Svensson, Delsing, Johansson PRA (2014)





What new did we learn:

Quantum particle creation makes clock tick slower

Precision of quantised light clocks in a Schwarzschild spacetime

Tupac Bravo,[1,*] Dennis Rätzel,[2,†] David Edward Bruschi,[3,4,1,‡] and Ivette Fuentes[5,§]

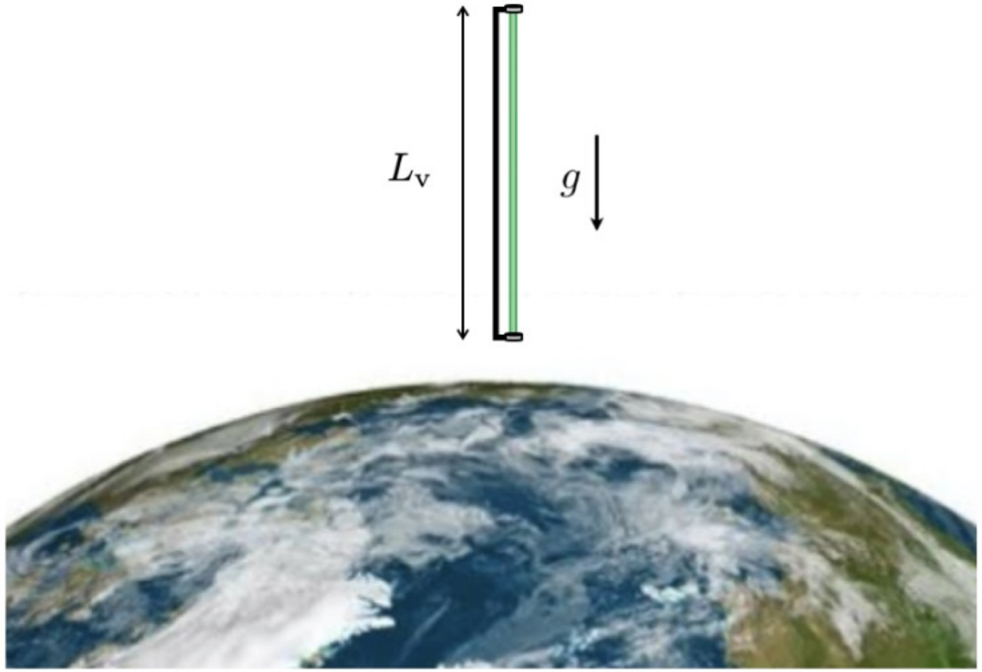


FIG. 1. Vertical light clock in the gravitational field of the Earth. The mirrors are held together by a $L_v = 20cm$ rigid rod.

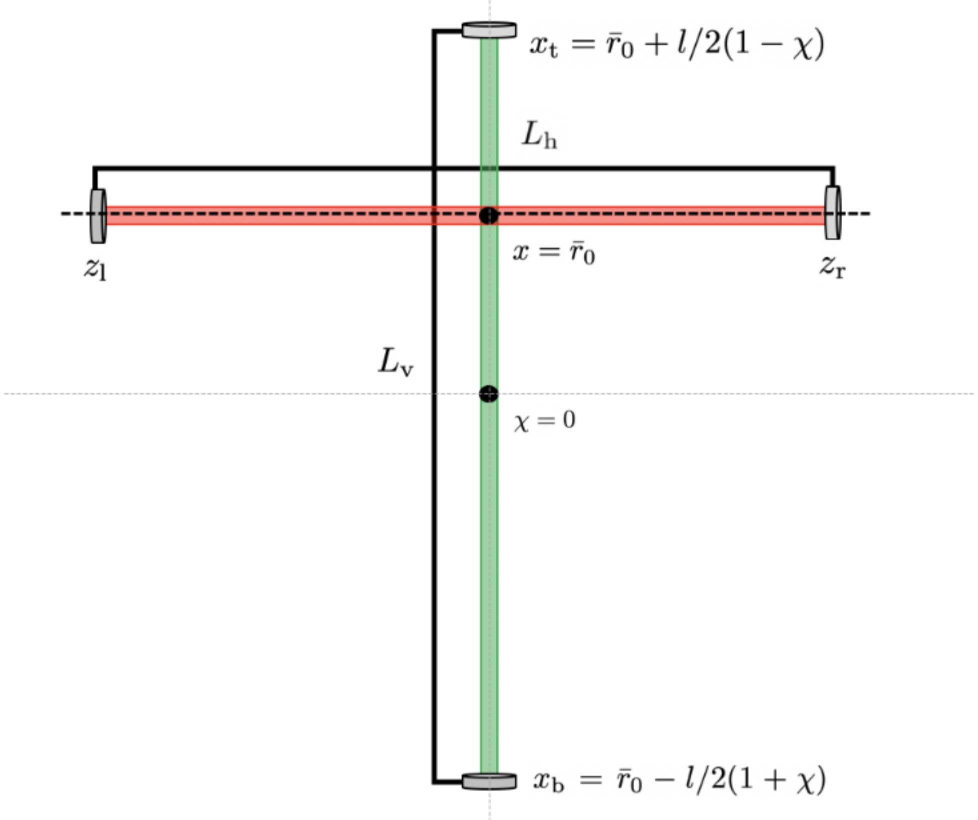


FIG. 4. The horizontal and vertical cavities intersect at $x = \bar{r}_0$. This figure shows where the mirrors are located.

[Submitted on 16 Apr 2022]

Gravitational time dilation in extended quantum systems: the case of light clocks in Schwarzschild spacetime

Tupac Bravo, Dennis Rätzel, Ivette Fuentes

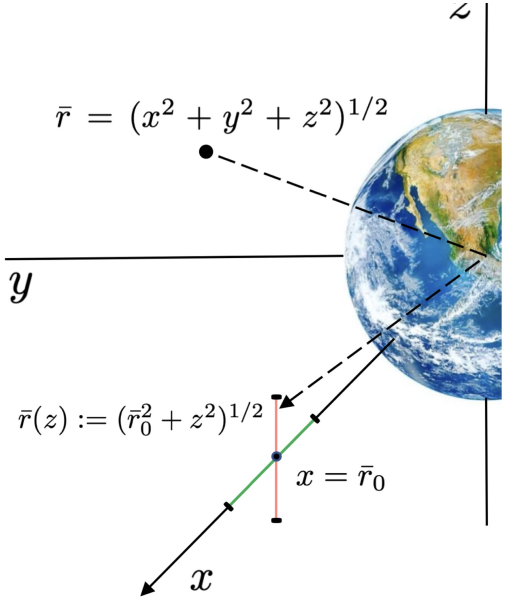
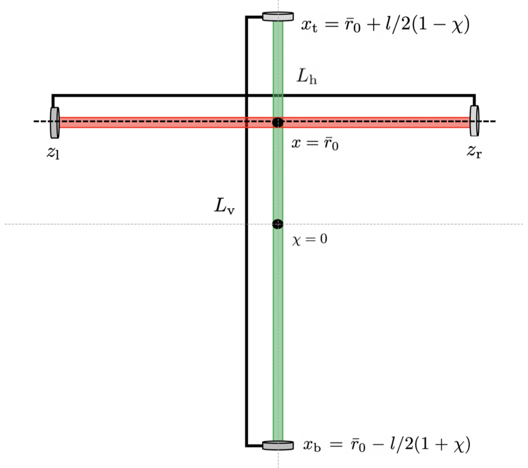
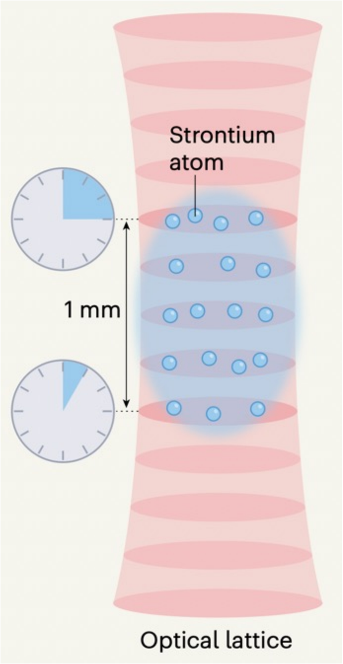


FIG. 1. Coordinate system used to quantize the field in horizontal (red) and vertical (green) cavities.





And for quantum clocks with squeezed vacuum input states

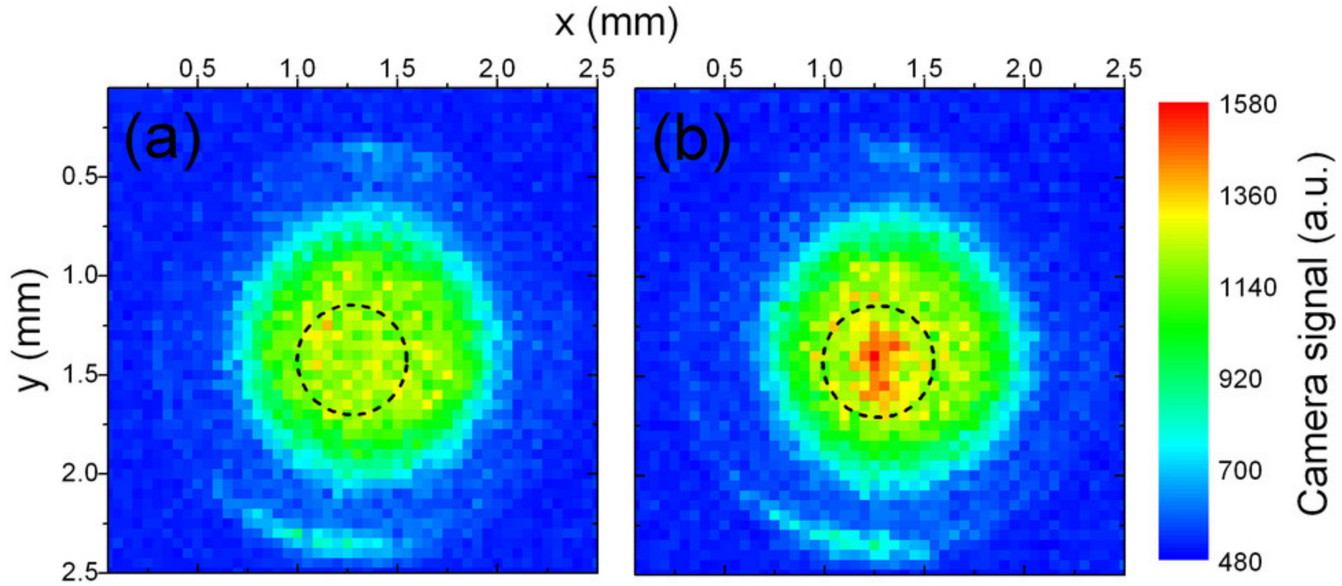
$$\Delta_{\mathrm{h},k}(\tau_0) = \frac{1}{2\sqrt{\mathcal{M}}} \frac{1}{\sqrt{N_p(N_p+1)}} \frac{L}{\pi\, c\, k}$$

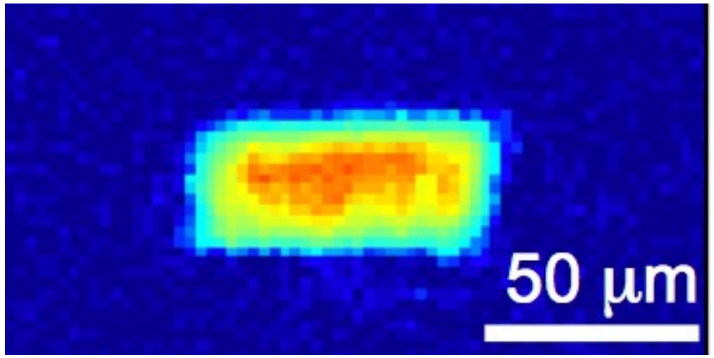
$$\Delta_{\mathrm{v},k}(\tau_0) = \Delta_{\mathrm{h},k}(\tau_0) \left(1 + \frac{r_{\mathrm{S}} L}{4\bar{r}_0^2}\chi\right)$$

$$\boldsymbol{g} = \boldsymbol{\eta} + \frac{r_{\mathrm{S}}}{\bar{r}} \mathbb{1} + \frac{r_{\mathrm{S}}^2}{2\bar{r}^2} \mathrm{diag}\left(-1, \frac{3}{4}, \frac{3}{4}, \frac{3}{4}\right)$$

Clock made of atoms in a trap

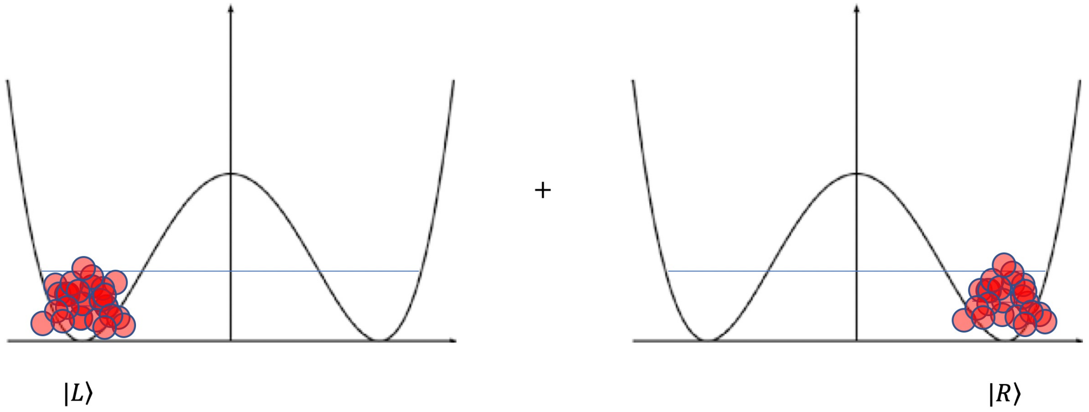
Notion of time?





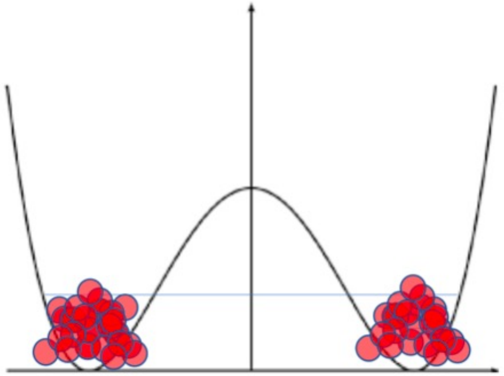
Use collective modes of vibration

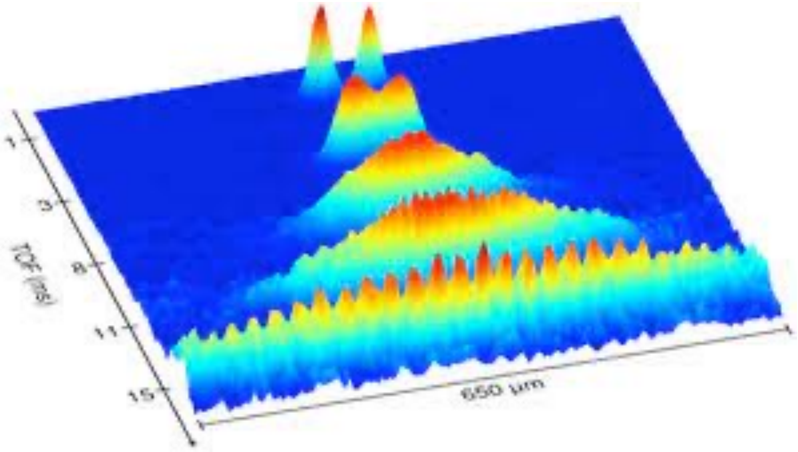
Testing gravitational induced collapse



Large mass experiments

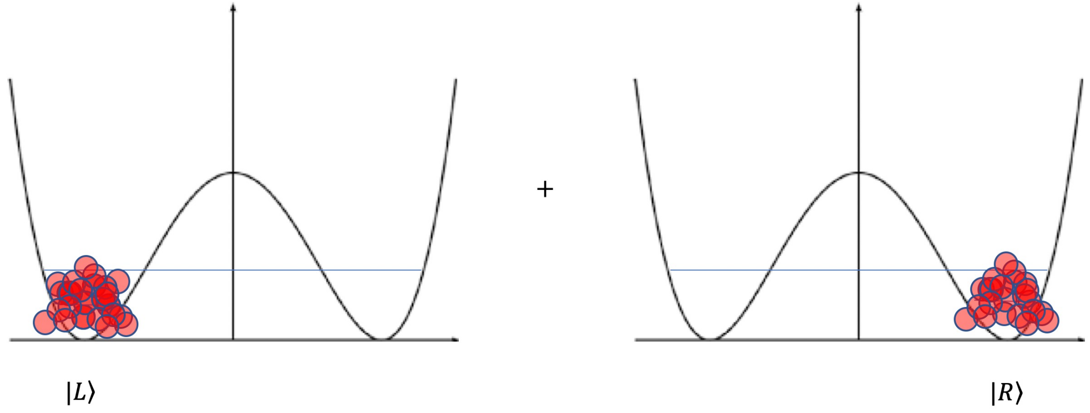
Spatial superpositions in a double-well potential





Schrödinger Cat state

$$|\psi\rangle = \frac{1}{\sqrt{2}}\left(|N_L 0_R\rangle + |0_L N_R\rangle\right)$$



Test Penrose’s gravitational induced collapse

Decoherence time must be longer than the collapse time

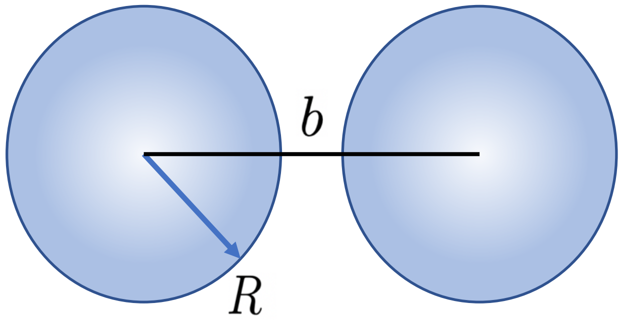
$$\tau \sim \frac{\hbar}{E_G} \qquad E_G = \frac{13Gm^2N^2}{14R}.$$

Non-uniform spherical just touching

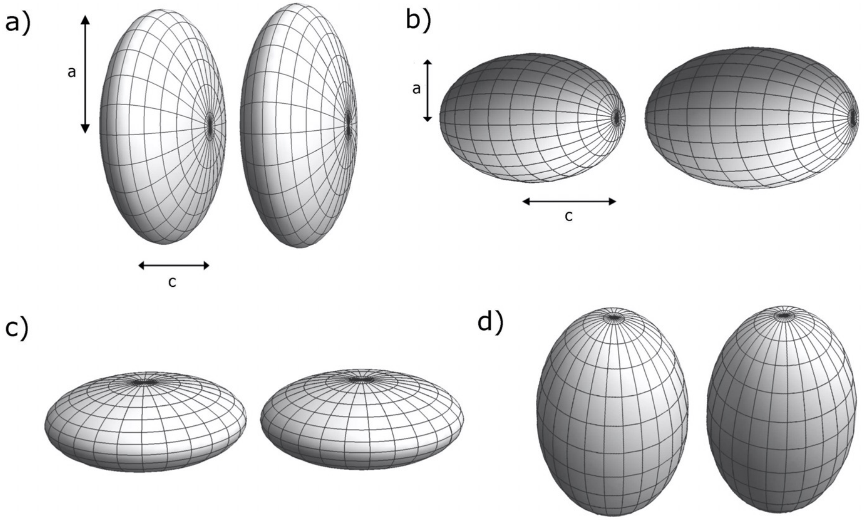
Gravitational self-energy

For a uniform sphere

$$E_G = \begin{cases} \frac{6GM^2}{5R}\left(\frac{5}{3}\lambda^2 - \frac{5}{4}\lambda^3 + \frac{1}{6}\lambda^5\right) & \text{if } 0 \leq \lambda \leq 1, \\ \frac{6GM^2}{5R}\left(1 - \frac{5}{12\lambda}\right) & \text{if } \lambda \geq 1, \end{cases}$$



where $\lambda := b/(2R)$



Gravitational collapse depends strongly in the mass geometry: signature

a) and b) can collapse at shorter times depending on the ellipticity

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Decoherence and noise

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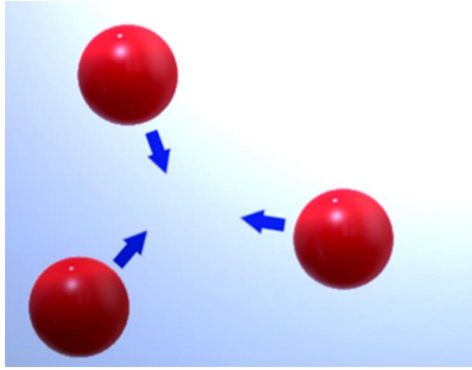
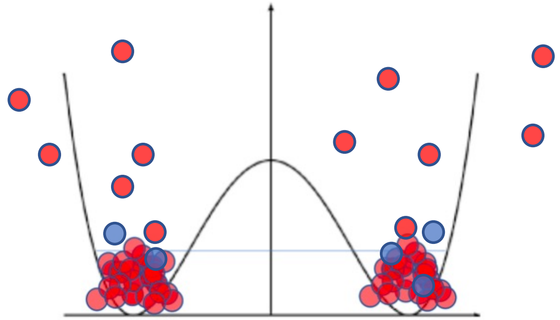
- Three-body recombination
- Two-body losses
- Thermal cloud interactions
- Foreign atom interactions
- Decoherence from the trapping potential





Advantage: atom-atom interactions can be controlled in a BEC

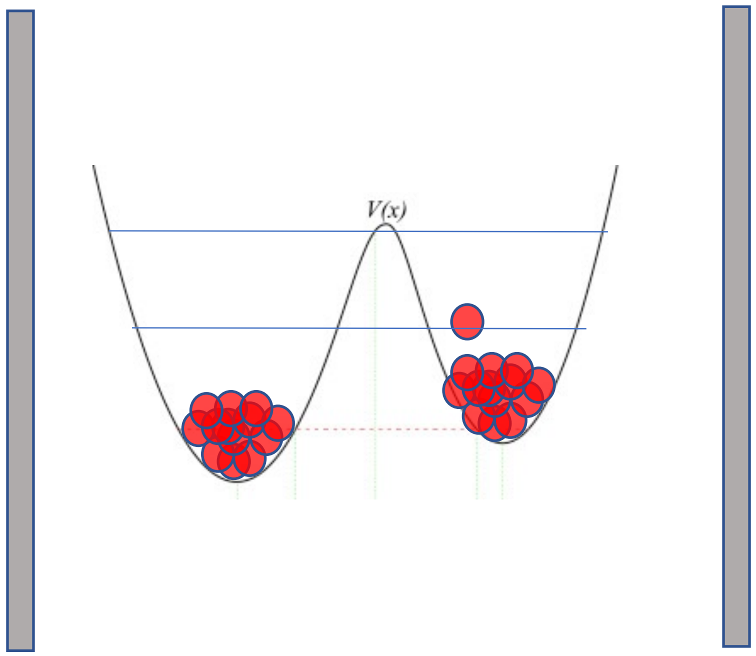
It would be convenient to supressed them after the superposition has been created.

$\gamma = 1/(8\pi)$ A ^{133}Cs BEC with 4×10^9 atoms and R=1 μm would collapse in approximately 2 seconds

$\gamma = 8\pi$ $2\,\mathrm{s}$ would occur when $N \approx 10^9$ and $R = 0.1\,\mathrm{mm}$ or $N \approx 10^8$ and $R = 1\,\mu\mathrm{m}$.

Confidential: shaking of the building

- Work in progress with Roger, Richard and Devang



$$\Delta E_G \Delta t = \frac{\hbar}{2}$$

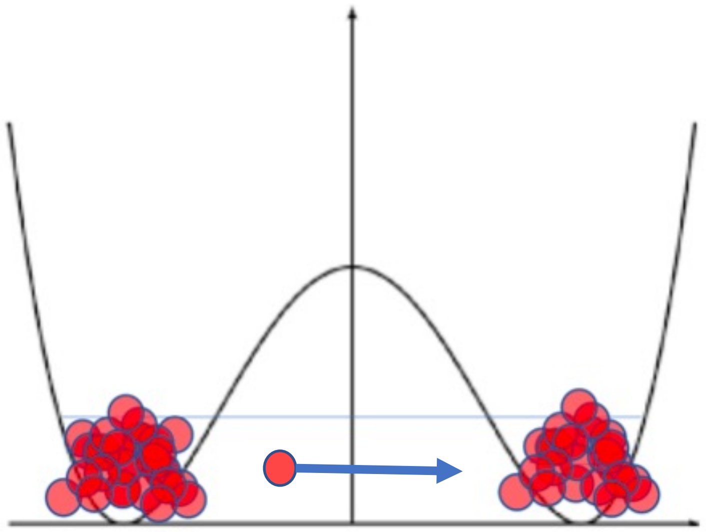
So far, the community has worked on making the gravitational self-energy larger by increasing the mass to get times that are short enough to see collapse. For this N $> 10^9$

For much smaller masses collapse times are very long. We propose to measure a very small energy uncertainty due to gravity: shaking of the building

A very small energy difference between wells can excite and atom to the next energy level. Experiments can detect single atoms in a higher state!

Other BEC states

- Work in progress with Penrose & Westbrook



In a BEC atoms are not bound together and can tunnel between left and right wells.

We are studying gravitational self-energy for a variety of BEC quantum states

Gravitational self-energy for spheroidals

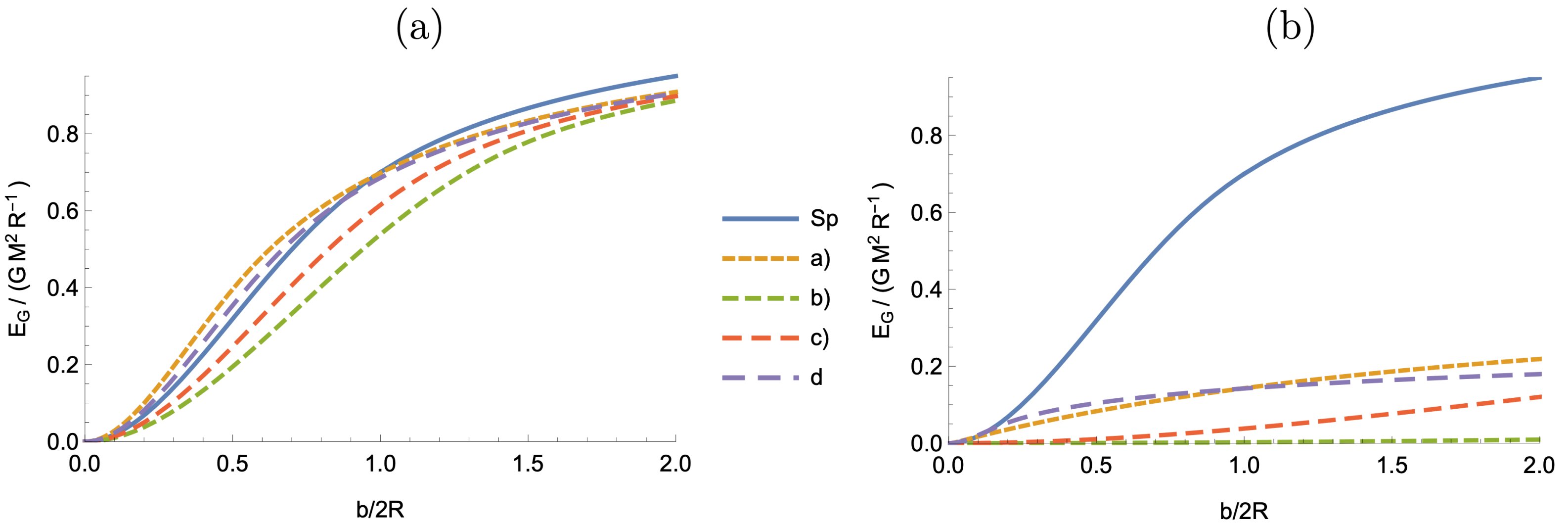


Figure 4. Both plots are of the gravitational self-energy of the difference between displaced uniform spherical and spheroidal mass distributions, E_G, against $b/(2R)$, where R is the radius of the sphere and b is the distance between the centres of the states. All mass distributions have the same total mass and volume. The solid line is for the spherical case, and the various dashed and dotted lines are for the a), b), c) and d) spheroidal configurations illustrated in Figure 3. The left plot is for $\epsilon = 0.5$ (ellipticity $e = 0.87$), and the right plot is for $\epsilon = 0.01$ (ellipticity $e = 0.99995$).

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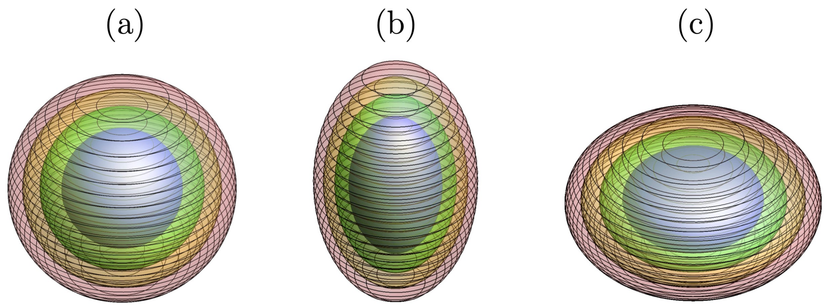
BEC geometries and density

BEC densities are not uniform
quadratic and Gaussian densities



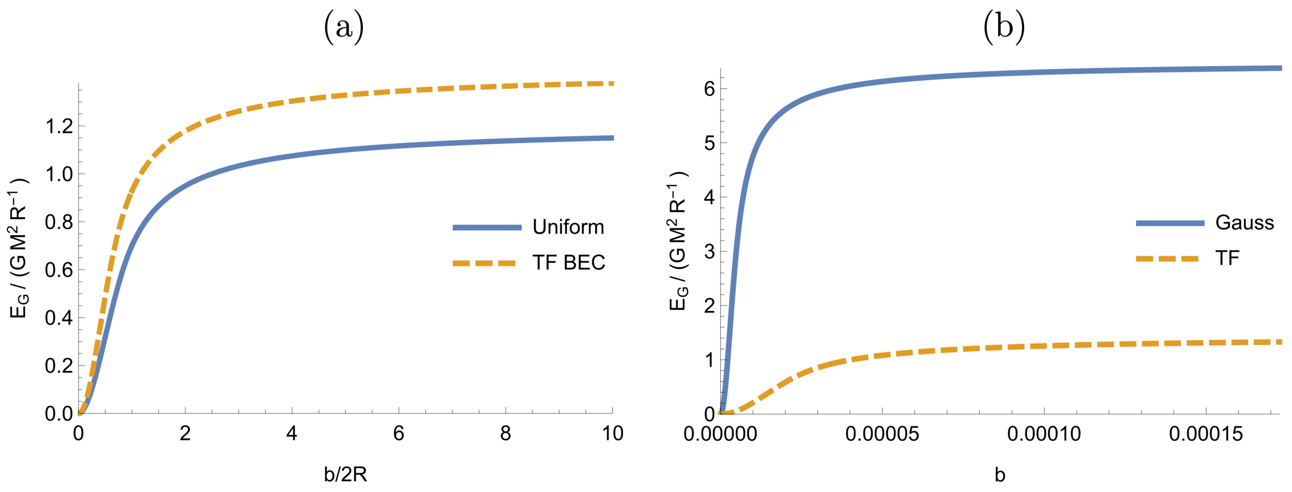


Figure 7. On the left is the gravitational self-energy of the difference between displaced spherical BECs (in the TF approximation) and displaced uniform spheres, $E_G/(GM^2R^{-1})$, against $b/(2R)$ where R is the radius of the spheres, M is their mass and b is the distance between the centres of the sphere states. On the right is $E_G/(GM^2R^{-1})$ against b of spherical ^{133}Cs BECs in the TF and Gaussian approximations with 10^6 atoms, the same trapping frequency $\omega_0 = 100\,\mathrm{Hz}$, and with the standard scattering length in the former regime, but with zero scattering in the latter so that we have an ideal BEC in that case.

Result: non-uniform densities produce shorter collapse times

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Using Quantum technologies we hope to understand better physics at the interplay of quantum and gravity.