systematics uncertainties in the determination of the local dark matter density
(+ complementarity of different targets in direct detection)

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[1] the relevance of the local dark matter density

$$\rho_0 \equiv \rho_{dm}(R_0 \sim 8 \text{ kpc})$$

:: ρ_0 is a main astrophysical unknown for DM searches ::

key ingredient to compute DM signals and draw limits

uncertainties on ρ_0 are crucial in interpreting positive DM detections

scattering off nuclei

$$\frac{dR}{dE} \propto n_{dm} \int_{v_{min}}^{\infty} dv \, \frac{f(v)}{v} \propto \rho_0$$

signal: nuclei recoils

sensitive to $\langle \rho_0 \rangle_{mpc}$

capture in Sun/Earth

$$\frac{dN_{dm}}{dt} = C - 2\Gamma_{ann}$$

$$C \propto n_{dm} \int_0^{v_{max}} dv \, \frac{f(v)}{v} \propto \rho_0$$

signal: ν from Sun/Earth

sensitive to $\langle \rho_0 \rangle$

halo annihilation/decay

$$\frac{d\phi}{dE} \propto \langle \sigma_{ann} v \rangle n_{dm}^k \propto \rho_0^k$$

signals: γ, e^+, $\bar{p}$, ν

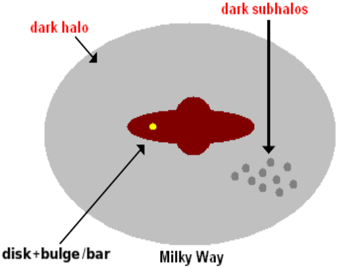
sensitive to $\langle \rho_0 \rangle$

[not the largest unknown]

[1] from dynamical observables to ρ_0

Milky Way mass model

bulge(+bar)	$\lesssim 3$ kpc	$\rho_b(x, y, z)$	x_b, y_b, z_b
disk	$\lesssim 10$ kpc	$\rho_d(r, z)$	Σ_d, r_d, z_d
dark halo	$\lesssim 200$ kpc	$\rho_{dm}(x, y, z) \propto \rho_0$	
+gas...			



a model fixes $M_i(R), \phi_i(R)$

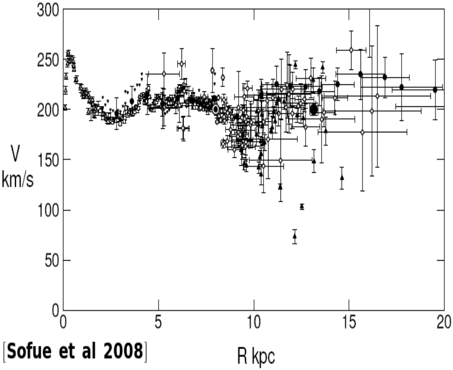
$$\sum_i \frac{d\phi}{dR}(R) \equiv \frac{G}{R^2} \sum_i M_i(< R) = \frac{v^2(R)}{R} \qquad v_0 \equiv v(R_0)$$

spherical average local density

$$\bar{\rho}_0 \simeq \frac{1}{4\pi R_0^2} \left(\frac{1}{G} \left. \frac{\partial \left(v^2 R\right)}{\partial R} \right|_{R_0} - \left. \frac{dM_d}{dR} \right|_{R_0} \right)$$

[1] from dynamical observables to ρ_0

observables



[Sofue et al 2008]

$R_0, \quad A - B = v_0/R_0, \quad A + B = -v_0'$
[fix v_0, v_0']

mass enclosed

$M(< 50 \text{ kpc}) \quad M(< 100 \text{ kpc})$

local surface density

$\Sigma_{|z|<1.1 \text{ kpc}} \quad \Sigma_*$

terminal velocities $\quad R < R_0$

$v(R) = v_T(l) + v_0 sin(l)$

velocity dispersions $\quad R \gtrsim R_0 \quad$ (tracer populations)

Jeans (sph., steady) $\frac{\partial(\nu\sigma_R^2)}{\partial R} + \frac{2\beta\sigma_R^2\nu}{R} = -\nu\sum_i \frac{d\phi_i}{dR} = -\frac{\nu G}{R^2}\sum_i M_i(< R)$

$\sigma_{los} \propto \sigma_R$

microlensing

$\tau_{LMC} \sim 10^{-7} \quad \tau_{bulge} \sim 10^{-6} \quad$ [constrain M_b]

[1] from dynamical observables to ρ_0

aim: use observables to constrain mass model parameters

selected references (different models/observables)

Caldwell & Ostriker '81	$\rho_0 = 0.23 \pm \times 2 \ \mathrm{GeV/cm^3}$
Gates, Gyuk & Turner '95	$\rho_0 = 0.30^{+0.12}_{-0.11} \ \mathrm{GeV/cm^3}$
Moore et al '01	$\rho_0 \simeq 0.18 - 0.30 \ \mathrm{GeV/cm^3}$
Belli et al '02	$\rho_0 \simeq 0.18 - 0.71 \ \mathrm{GeV/cm^3}$ (isoth.)
Strigari & Trotta '09	$\Delta\rho_0/\rho_0 = 20\%$ (projected; 2000 halo stars, v_{esc})
Catena & Ullio '09	$\rho_0 \simeq 0.39 \pm 0.03 \ \mathrm{GeV/cm^3}$ $\Delta\rho_0/\rho_0 = 7\%$!!
Salucci et al '10	$\rho_0 \simeq 0.43 \pm 0.21 \ \mathrm{GeV/cm^3}$

usual assumptions: $\rho_{dm} = \rho_{dm}(r)$, ρ_{dm} from DM-only simulations

[2] our numerical framework

difficult to obtain a MW-like galaxy at $z = 0$ with simulations

usually large bulges and small disks result (L problem)

recent sucessful attempt: Agertz, Teyssier & Moore 2010

dark matter + gas + stars

cosmological setup

WMAP 5yr cosmology
select DM-only halo
$M_{vir} \sim 10^{12}\ \mathrm{M}_\odot \quad R_{vir} \sim 205$ kpc
no major merger for $z < 1$

baryonic features

star formation (Schmidt law; ϵ_{ff}, n_0)

$\dot{\rho}_g = -\epsilon_{ff} \frac{\rho_g}{t_{ff}}$

stellar feedback (SNII, SNIa, wind)

numerical features

$m_{DM} = 2.5 \times 10^6\ \mathrm{M}_\odot$
$\Delta x = 340$ pc

main result

MW-like galaxy with $v_c \sim const$, $\quad B/D \sim 0.25$, $\quad r_d \sim 4 - 5$ kpc

[2] our numerical framework

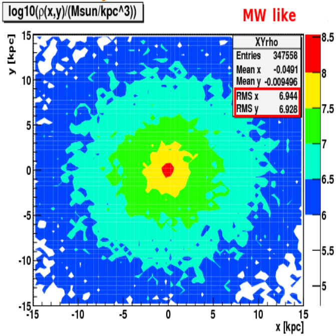
Run	ϵ_{ff}	Feedback	Star formation threshold, n_0
SR6-n01e1	1 %	SNII	$0.1\,cm^{-3}$
SR6-n01e2	2 %	SNII	$0.1\,cm^{-3}$
SR6-n01e5	5 %	SNII	$0.1\,cm^{-3}$
SR6-n01e1ML	1 %	SNII, SNIa, mass loss	$0.1\,cm^{-3}$ **MW like**
SR6-n01e2ML	2 %	SNII, SNIa, mass loss	$0.1\,cm^{-3}$
SR6-n01e5ML	5 %	SNII, SNIa, mass loss	$0.1\,cm^{-3}$ **baryon+**
SR6-n1e1	1 %	SNII	$1\,cm^{-3}$
SR6-n1e2	2 %	SNII	$1\,cm^{-3}$
SR6-n1e5	5 %	SNII	$1\,cm^{-3}$

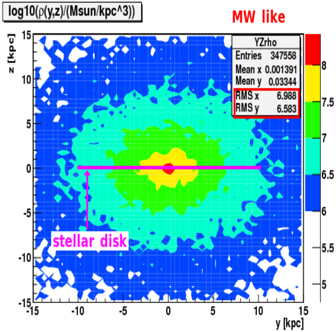
Run	$M_{disk,s}$	$M_{disk,g}$	$M_{bulge,s}$	r_d [kpc] (1)	f_{gas} (2)	B/D	B/T	j_{bar} (3)
SR6-n01e1	8.6	1.6	2.0	3.8	0.13	0.23	0.19	1920
SR6-n01e2	7.4	1.3	4.6	7.6	0.10	0.62	0.38	1655
SR6-n01e5	5.6	0.72	7.0	~ 15.0	0.05	1.25	0.56	1305
SR6-n01e1ML	8.0	2.3	2.2	5.0	0.18	0.27	0.21	1960
SR6-n01e2ML	8.1	1.6	3.8	5.0	0.12	0.47	0.32	1718
SR6-n01e5ML	5.5	0.93	7.2	~ 15.0	0.07	1.30	0.57	1464
SR6-n1e1	6.6	3.3	2.9	2.7	0.26	0.44	0.31	1594
SR6-n1e2	6.4	2.4	4.3	2.5	0.18	0.67	0.40	1804
SR6-n1e5	6.0	2.1	5.2	2.7	0.16	0.87	0.46	1643

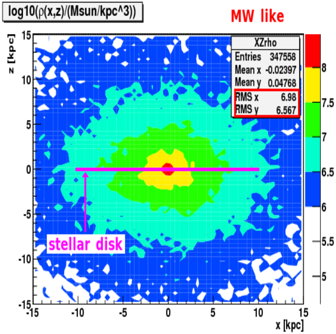
[Agertz, Teyssier & Moore 2010]

to bracket uncertainties we consider: DM only, MW like, baryon+

[3] halo shape: a first look







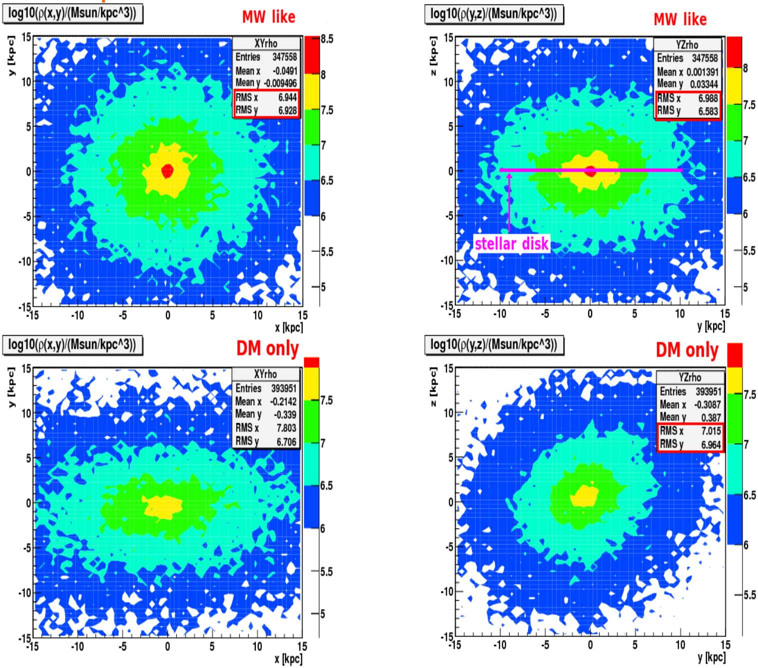
profiles of dark matter density

SR6-n01e1ML :: MW like

approximately axisymmetric halo

$$10^7\ \mathrm{M}_\odot/\mathrm{kpc}^3 \sim 0.38\ \mathrm{GeV}/\mathrm{cm}^3$$

[3] halo shape: a first look



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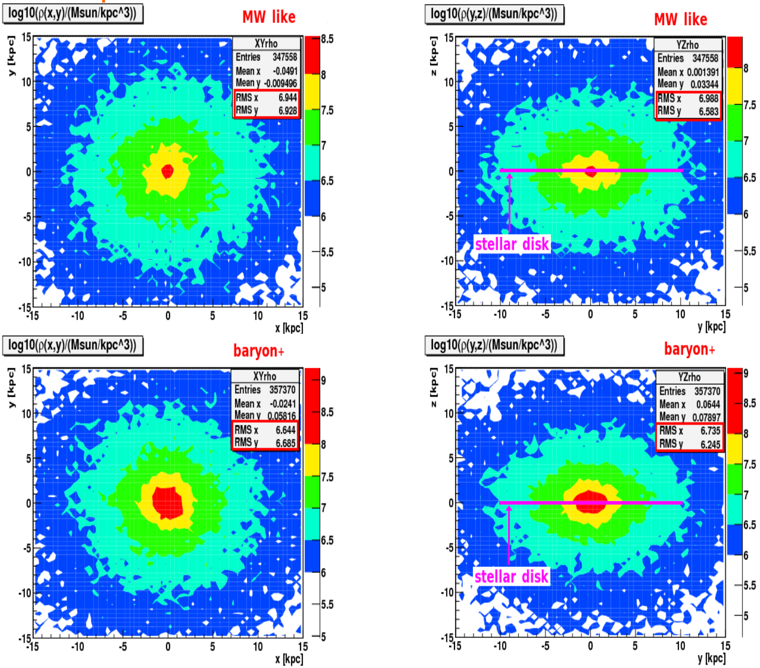
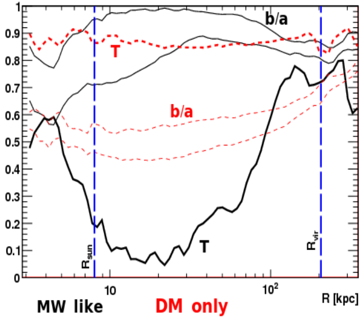
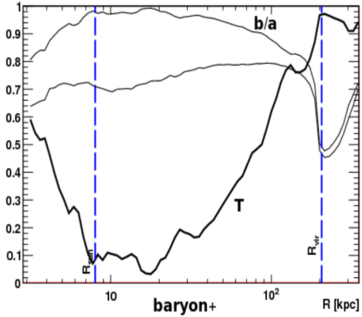
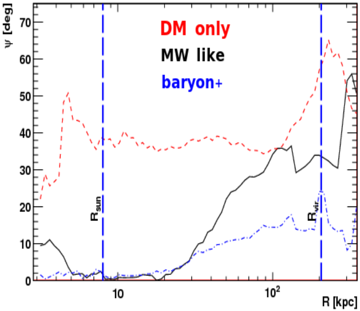


[3] halo shape: getting more quantitative







inclusion of baryons

prolate $\longrightarrow$ oblate halo shape

flattening aligned with stellar disk for $R \lesssim 20$ kpc

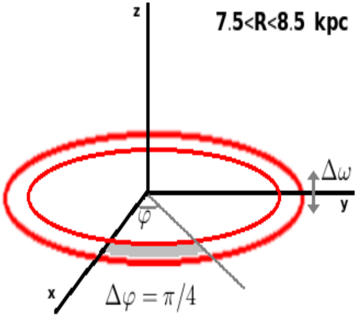
[MP, Agertz, Bertone, Moore & Teyssier '10]

[3] halo shape: consequences for ρ_0

- many studies assume a spherical halo [e.g. Catena & Ullio, Strigari & Trotta]
- data then constrains the spherical average local density $\bar{\rho}_0$:

$$\bar{\rho}_0 \simeq \frac{1}{4\pi R_0^2}\left(\frac{1}{G}\left.\frac{\partial(v^2R)}{\partial R}\right|_{R_0} - \left.\frac{dM_d}{dR}\right|_{R_0}\right)$$

- model triaxial halo is tricky (b/a, c/a not known nor constant)
- to estimate systematic uncertainty compare $\bar{\rho}_0 \leftrightarrow \rho_0$ in simulations



strategy

spherical shell $7.5 < R < 8.5$ kpc

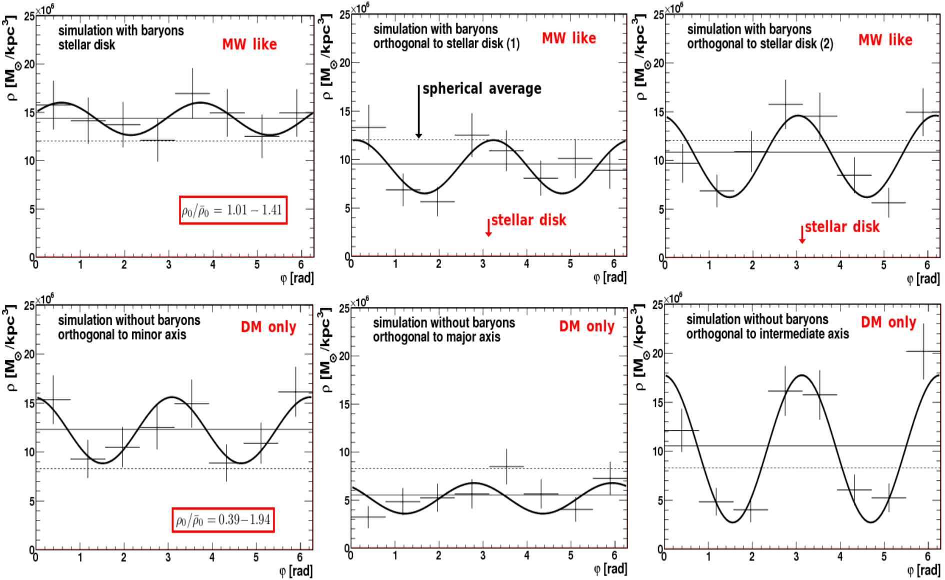
select particles in 3 orthogonal rings

divide rings into 8 portions $\Delta\varphi = \pi/4$

evaluate ρ along the ring, $\rho(\varphi)$

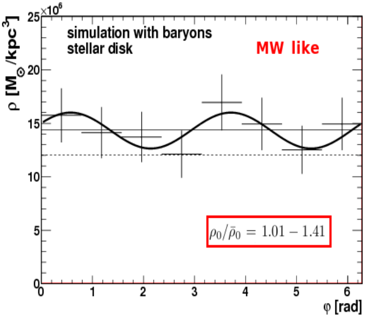
[3] halo shape: consequences for ρ_0

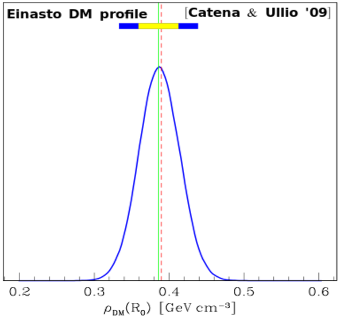
[MP, Agertz, Bertone, Moore & Teyssier '10]



[3] halo shape: consequences for ρ_0

just an exercise...



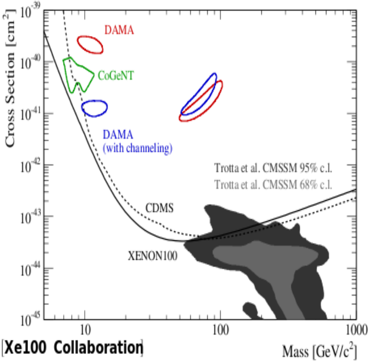


$$\rho_0 = 0.466 \pm 0.033(\mathrm{stat}) \pm 0.077(\mathrm{syst})\ \mathrm{GeV/cm^3}$$

:: syst > stat ::

future: bayesian study with triaxial halo

[4] astrophysical uncertainties: why do we care?



[Xe100 Collaboration]

direct detection

$$\frac{dR}{dE_R} = \frac{\rho_0}{m_\chi m_N} \int_{v_{min}}^{\infty} d^3\vec{v}\, v\, f(\vec{v} + \vec{v}_E; v_{esp}, v_0) \frac{d\sigma_{\chi N}}{dE_R}$$

standard assumptions

$$\rho_0 = 0.3\ \mathrm{GeV/cm^3}$$

$$f(v) \propto e^{-v^2/v_0^2},\ v_0 \simeq 220\ \mathrm{km/s},\ v_{esc} \simeq 600\ \mathrm{km/s}$$

exclusion limits are not rigid

astrophysics should really be treated as a nuisance in direct searches

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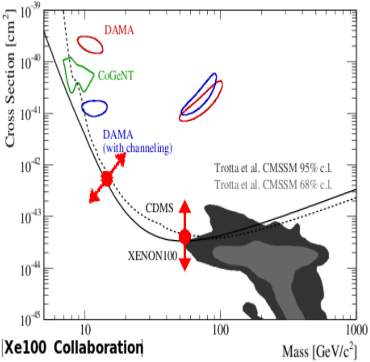

[Xe100 Collaboration]

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[5] an application: different targets in direct detection

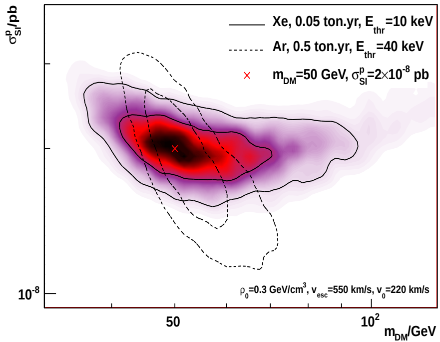
spin independent scattering rate

$$\frac{dR}{dE_R} = \frac{\rho_0 \sigma^p_{SI}}{m_\chi \mu^2_{p\chi}} \times \underbrace{\mathbf{A}^2 F^2(E_R, \mathbf{A})}_{\text{nuclear physics}} \times \underbrace{\mathcal{F}\left(v_{min}(E_R, \mathbf{A}, m_\chi)\right)}_{\text{astrophysics}}$$

$$v_{min} = \sqrt{m_N E_R/(2\mu^2_{N\chi})}$$

breaking $\rho_0 \sigma^p_{SI}/m_\chi$ degenaracy use several energy bins and/or targets

kinematic limit $m_\chi \gg m_N \Rightarrow v_{min} \simeq \sqrt{E_R/(2m_N)}$; $dR/dE_R \propto \rho_0 \sigma^p_{SI}/m_\chi$



work in progress

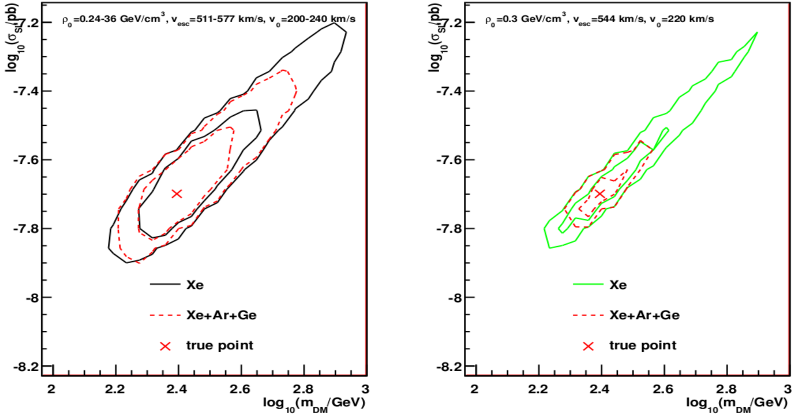
[with Bertone, Trotta, Baudis]

very preliminary!

[5] an application: different targets in direct detection

how much do astrophysical uncertainties blow up target complementarity?

with astrophysical uncertainties | fixed astrophysics



work in progress [with Bertone, Trotta, Baudis] **very preliminary!**

[..] conclusions

ρ_0 in light of recent N-body+hydro simulations

> baryons turn DM halo from prolate to oblate

> flattening is along the disk

> $\rho_0/\bar{\rho}_0 \simeq 1.21 \pm 0.20$

astrophysical uncertainties: not an academic question!

ultimately limit our ability to combine signals

example: complementarity between different targets in direct detection

upcoming direct detection experiments and results urge for accurate control over systematics of astrophysical parameters

[+] the role of baryons on dark matter halos

adiabatic contraction [Blumenthal et al 1986]

spherical mass distribution $M_i(< R_i)$: baryons + dark matter $\quad f_b \sim 0.17$
baryons cool and contract slowly $\rightarrow M_b(< R)$
circular orbits + $L = const$

$$R\,(M_b(< R) + M_{dm}(< R)) = R_i M_i(< R_i) = R_i M_{dm}(< R)/(1 - f_b)$$

$$\rho_{dm} \propto R^{-2}\,\frac{dM_{dm}}{dR}$$

final DM profile is significantly contracted
[+ Gnedin et al 2004, Gustafsson et al 2006]

halo shape

DM-only halos are prolate
+ baryons: more oblate halos (still triaxial)
in any case, $\rho_{dm} \neq \rho_{dm}(r)$

aim
address systematics on ρ_0 in light of recent N-body+hydro simulations
a realistic pdf on ρ_0 is needed if we are to convincingly identify WIMPs

[+] halo shape: getting more quantitative

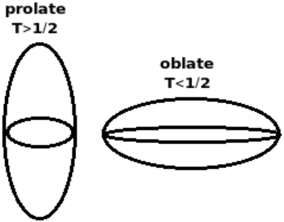
inertia calculations

for a set of N_p particles, $J_{ij} = \frac{\sum_{k=1}^{N_p} m_k x_{i,k} x_{j,k}}{\sum_{k=1}^{N_p} m_k}$

principle axes: eigenvectors $\vec{j}_a$ (major), $\vec{j}_b$ (intermediate), $\vec{j}_c$ (minor)

axis ratios: $b/a = \sqrt{J_b/J_a}$, $c/a = \sqrt{J_c/J_a}$

triaxiality: $T = \frac{1-b^2/a^2}{1-c^2/a^2}$

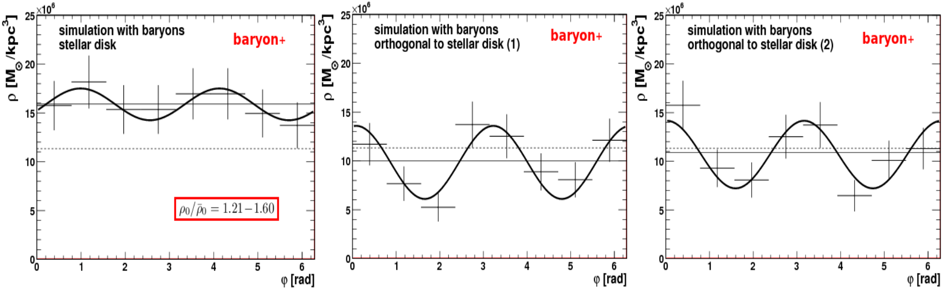


iterative procedure ['a la Katz et al '91]

$$r < R \rightarrow b/a,\ c/a,\ \vec{j}_{a,b,c} \rightarrow q = \sqrt{x^2 + \frac{y^2}{(b/a)^2} + \frac{z^2}{(c/a)^2}} < R \rightarrow \ldots$$

convergence criterium: 0.5% change in b/a, c/a

[+] halo shape: consequences for ρ_0



MW like	$\rho_0/\bar{\rho}_0 = 1.01 - 1.41$
baryon+	$\rho_0/\bar{\rho}_0 = 1.21 - 1.60$
DM only	$\rho_0/\bar{\rho}_0 = 0.39 - 1.94$

/ $\rho(\varphi) > \bar{\rho}_0$ because halo is flattened

/ halo-to-halo scatter can change normalisation

[MP, Agertz, Bertone, Moore & Teyssier '10]