NEUTRINO COSMOLOGY

ASI/COSMOS meeting on “LambdaCDM”
29 May 2019

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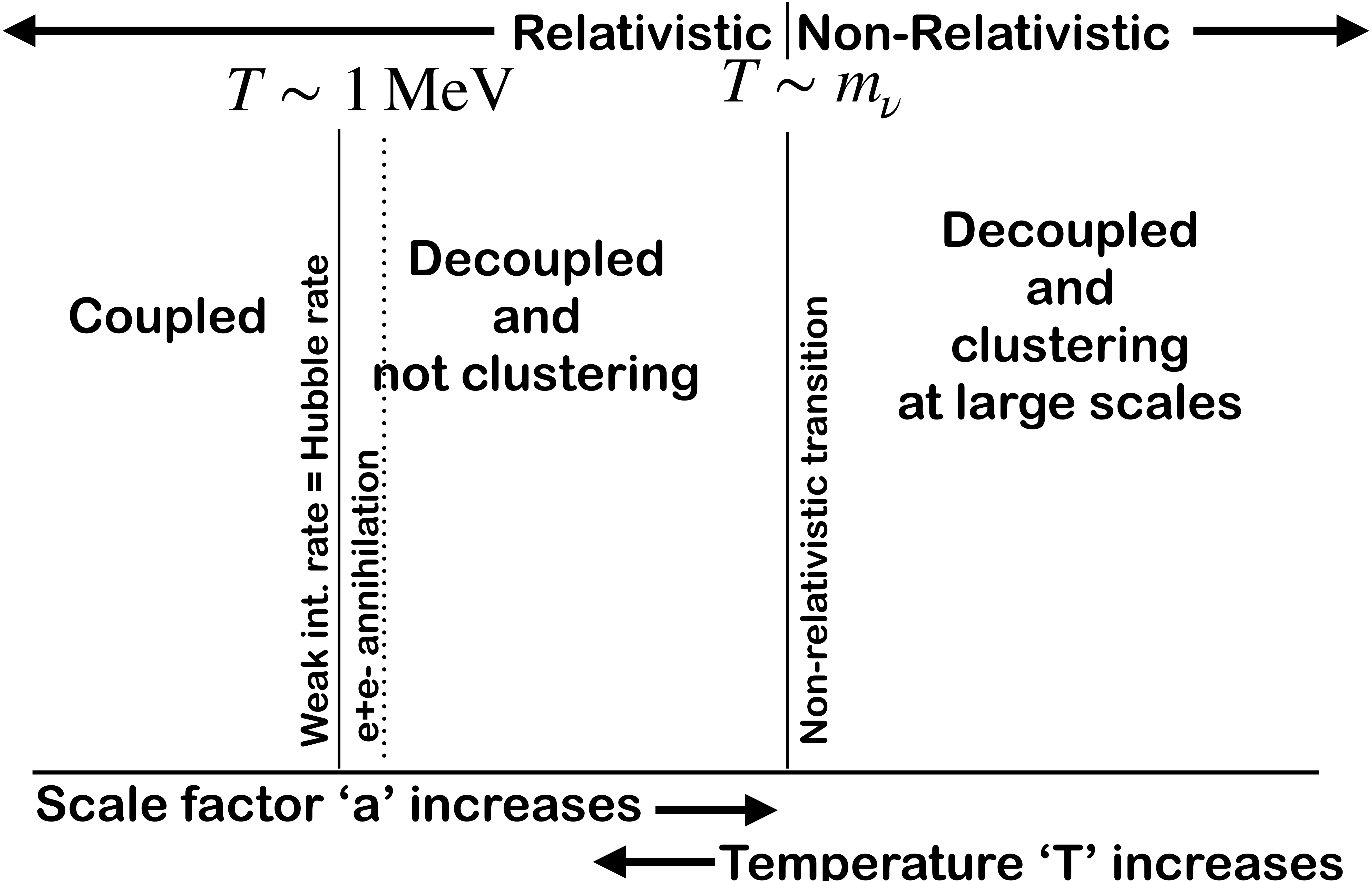
Credit:C.Pucci

Basics of neutrino cosmology



Basics of neutrino cosmology

← Relativistic Non-Relativistic →

T~m_nu

$$\rho_\nu \propto N_{\rm eff}$$

$$\rho_\nu \propto \Sigma m_\nu$$

$$N_{\rm eff} = \frac{\rho_{\rm rad} - \rho_\gamma}{\rho_\nu^{\rm st}} = 3.045$$

$$\Sigma m_\nu = \Sigma_{i=1,2,3} m_{\nu,i}$$

Distorsions due to non-inst decoupling
radiative corrections,
flavour oscillations
Dolgov, 1997, Mangano+,2005
deSalas&Pastor,2016

$$N_{\rm eff,\nu} \equiv \frac{\sum_i \rho_{\nu,i}}{\rho_{\nu,0}} = \frac{g/(2\pi)^3 \sum_i \int p_i^3 f_\nu(p_i, T_i) dp}{7/120\pi^2 T_\nu^4}$$

← **Temperature 'T' increases**

Effects on background quantities

Expansion rate

$$H(z)^2 = H_0^2 \left[(\Omega_c + \Omega_b)(1+z)^3 + \Omega_\gamma (1+z)^4 + \Omega_\Lambda + \frac{\rho_\nu(z)}{\rho_{\mathrm{crit},0}} \right]$$

modifies the angular size of the sound horizon at recombination $\theta_s = r_s / D_A$

modifies the angular scale of the Silk damping $\theta_d = \frac{r_d}{D_A} \propto \frac{1/\sqrt{H}}{1/H}$

$$1 + z_{\mathrm{eq}} = \frac{\Omega_c + \Omega_b}{\Omega_\gamma \left[1 + \frac{7}{8} \left(\frac{4}{11} \right)^{4/3} N_{\mathrm{eff}} \right]}$$

Matter-radiation equality

Perturbation effects

$$k_{\rm fs} \simeq 0.018\,\Omega_m^{1/2}\left(\frac{m_\nu}{1\,{\rm eV}}\right)^{1/2} h{\rm Mpc}^{-1}$$

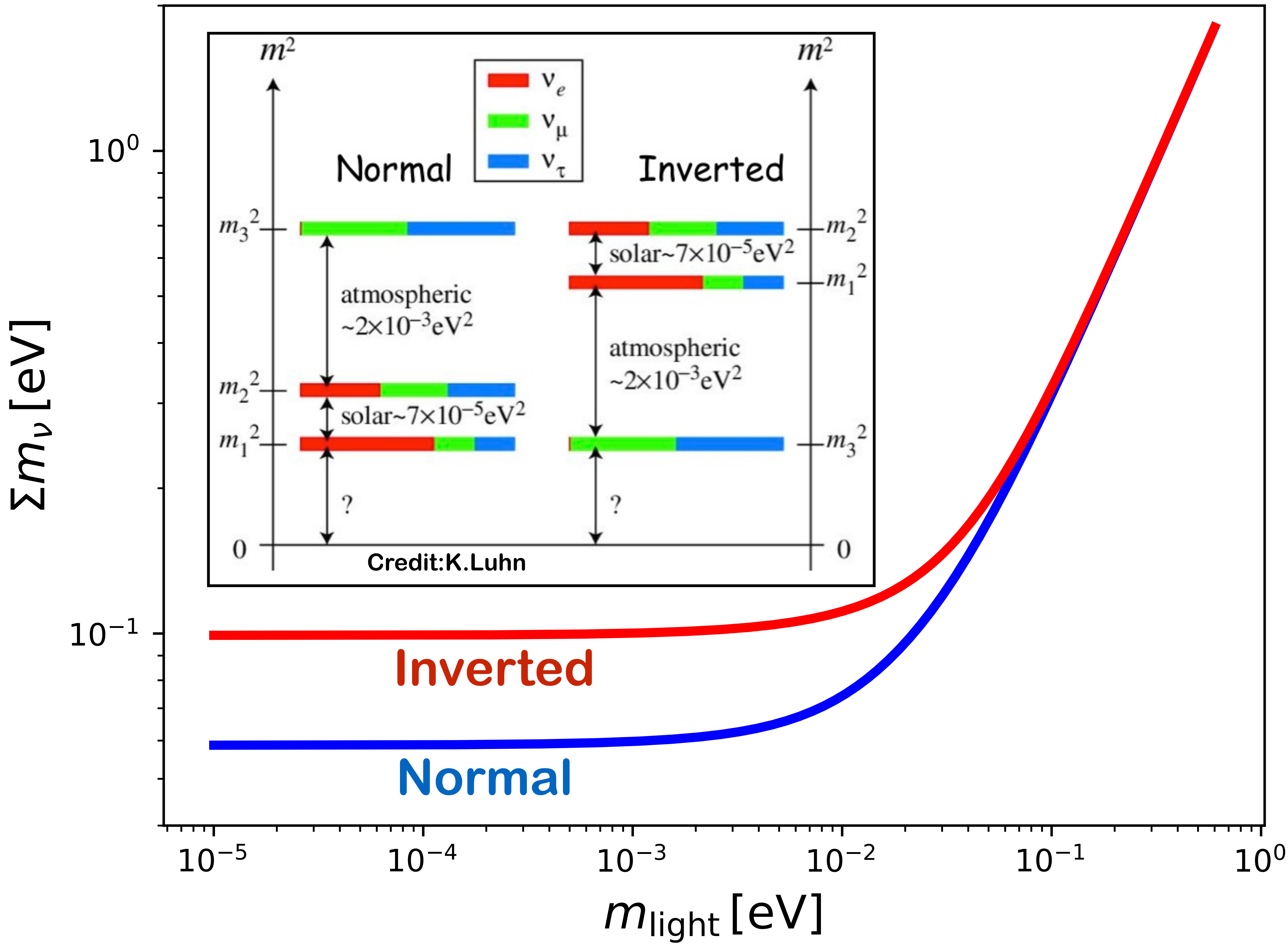
Free streaming scale

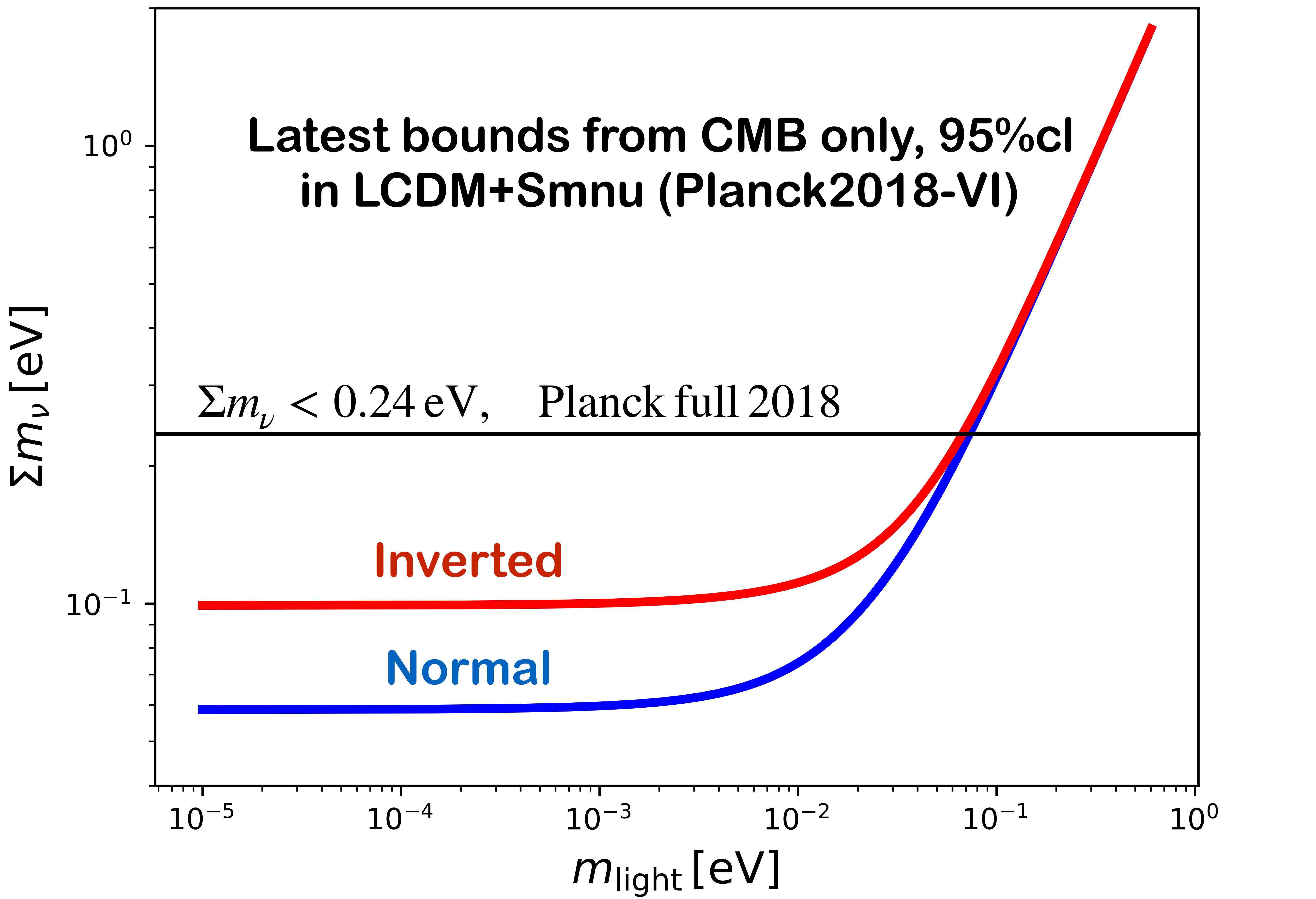
$$\delta_m(k \gg k_{\rm fs}) \propto a^{1-(3/5)\Omega_\nu/\Omega_m}$$

Suppressed growth

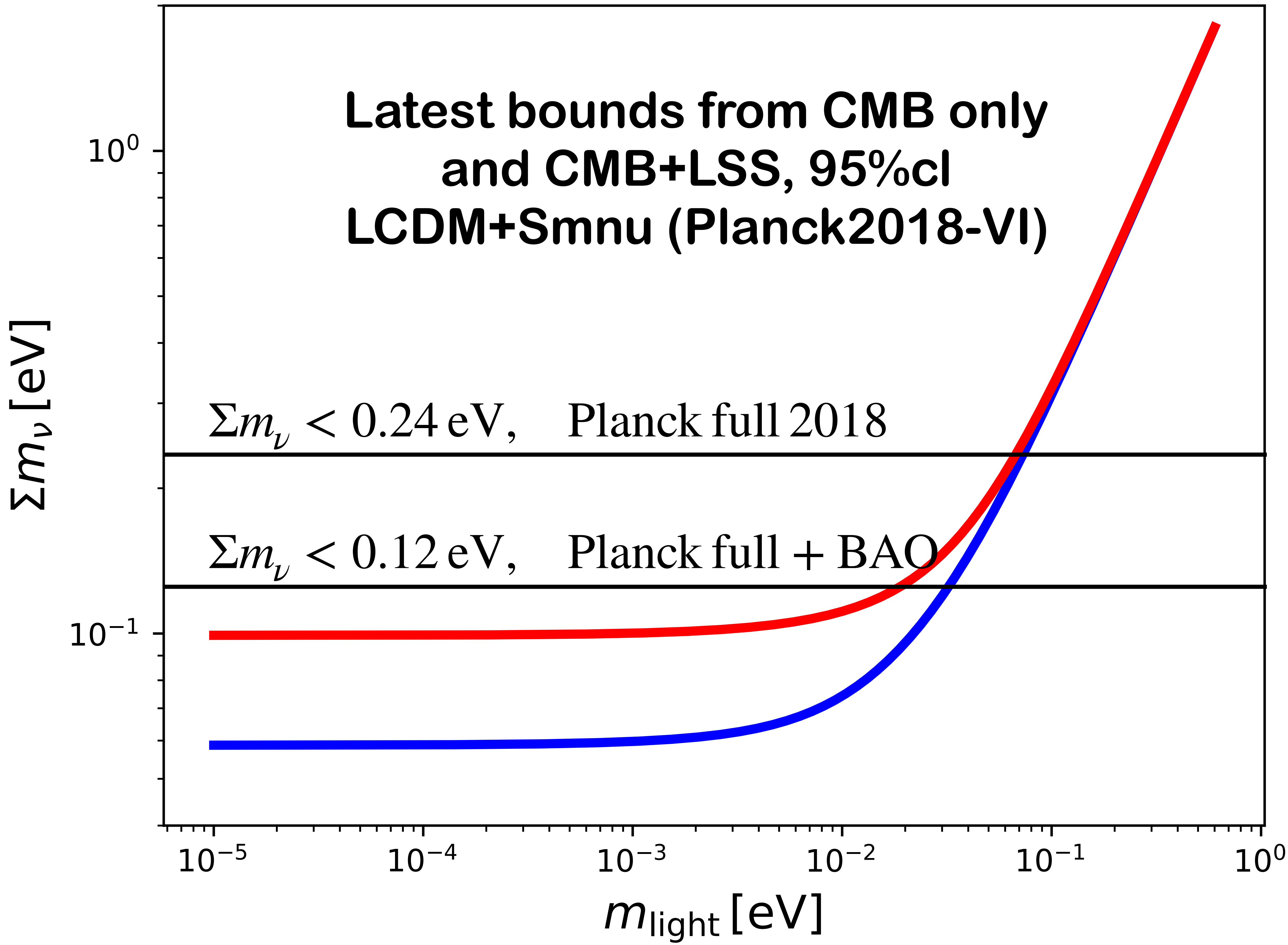
$$k_p r_s + \phi = p\pi$$

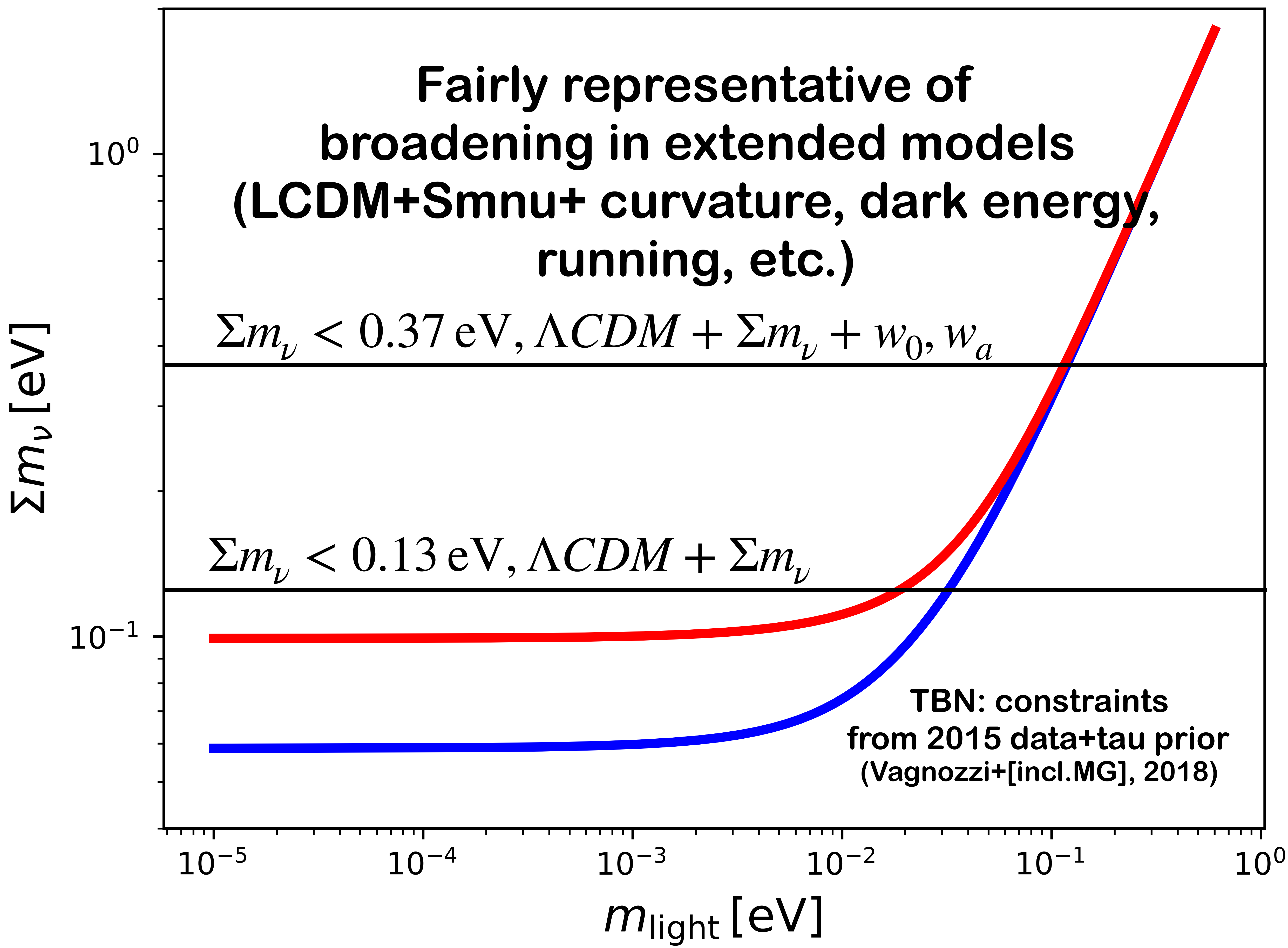
Acoustic phase shift


m^2
m^2
ν_e
ν_μ
ν_τ
Normal
Inverted
m_3^2
m_2^2
m_1^2
m_2^2
m_1^2
m_3^2
atmospheric ~2×10^{-3}eV^2
solar~7×10^{-5}eV^2
solar~7×10^{-5}eV^2
atmospheric ~2×10^{-3}eV^2
?
?
0
0
Credit:K.Luhn
10^0
10^{-1}
Inverted
Normal
10^{-5}
10^{-4}
10^{-3}
10^{-2}
10^{-1}
10^0
$m_{\rm light}$ [eV]
Σm_ν [eV]



Latest bounds from CMB only, 95%cl
in LCDM+Smnu (Planck2018-VI)
$\Sigma m_\nu < 0.24\,\mathrm{eV}$, Planck full 2018
Inverted
Normal
10^0
10^{-1}
Σm_ν [eV]
10^{-5}
10^{-4}
10^{-3}
10^{-2}
10^{-1}
10^0
$m_{\rm light}$ [eV]

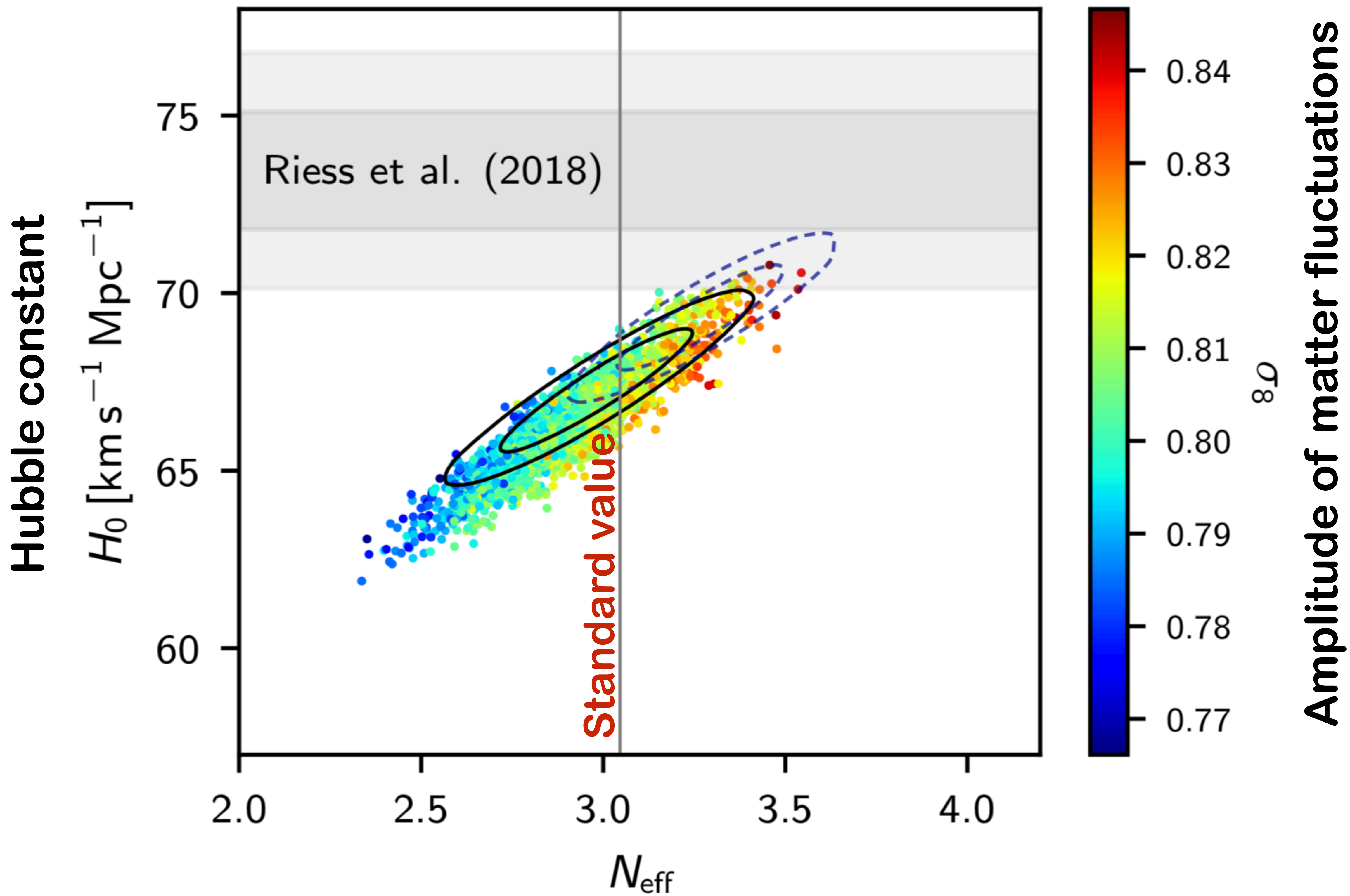

Latest bounds from CMB only
and CMB+LSS, 95%cl
LCDM+Smnu (Planck2018-VI)
$\Sigma m_\nu < 0.24$ eV, Planck full 2018
$\Sigma m_\nu < 0.12$ eV, Planck full + BAO
Σm_ν [eV]
10^0
10^{-1}
10^{-5}
10^{-4}
10^{-3}
10^{-2}
10^{-1}
10^0
$m_{\rm light}$ [eV]


Fairly representative of
broadening in extended models
(LCDM+Smnu+ curvature, dark energy,
running, etc.)
$\Sigma m_\nu < 0.37\,\mathrm{eV}, \Lambda CDM + \Sigma m_\nu + w_0, w_a$
$\Sigma m_\nu < 0.13\,\mathrm{eV}, \Lambda CDM + \Sigma m_\nu$
TBN: constraints
from 2015 data+tau prior
(Vagnozzi+[incl.MG], 2018)
10^0
10^{-1}
Σm_ν [eV]
10^{-5}
10^{-4}
10^{-3}
10^{-2}
10^{-1}
10^0
$m_{\rm light}$ [eV]

Current limits on Neff

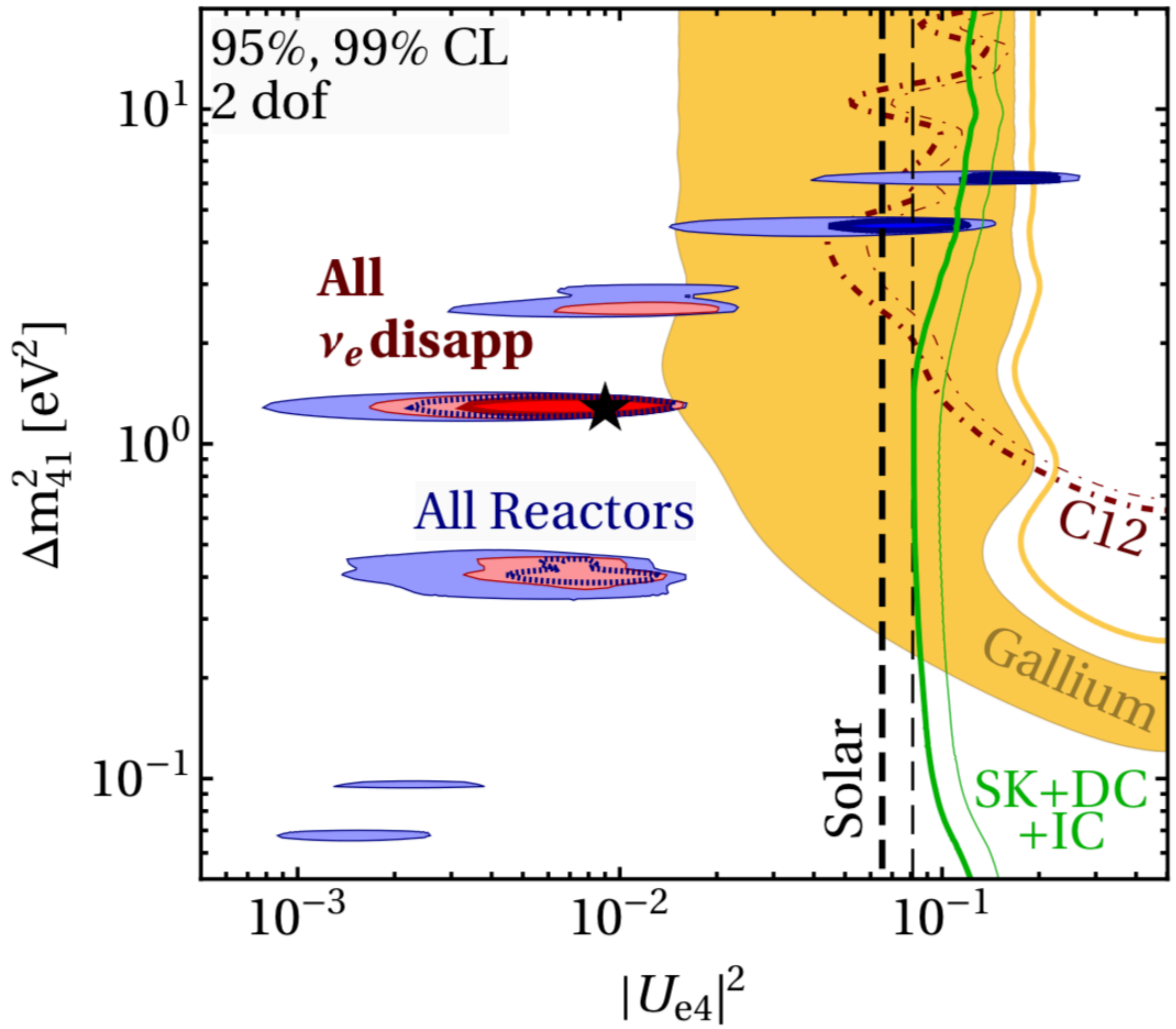
Planck collaboration, VI 2018

$$N_{\rm eff} = 2.99^{+0.34}_{-0.33}, 95\,\%\,c.l., \text{Planck2018 + BAO}$$



Light sterile in cosmology

Dentler+,2018



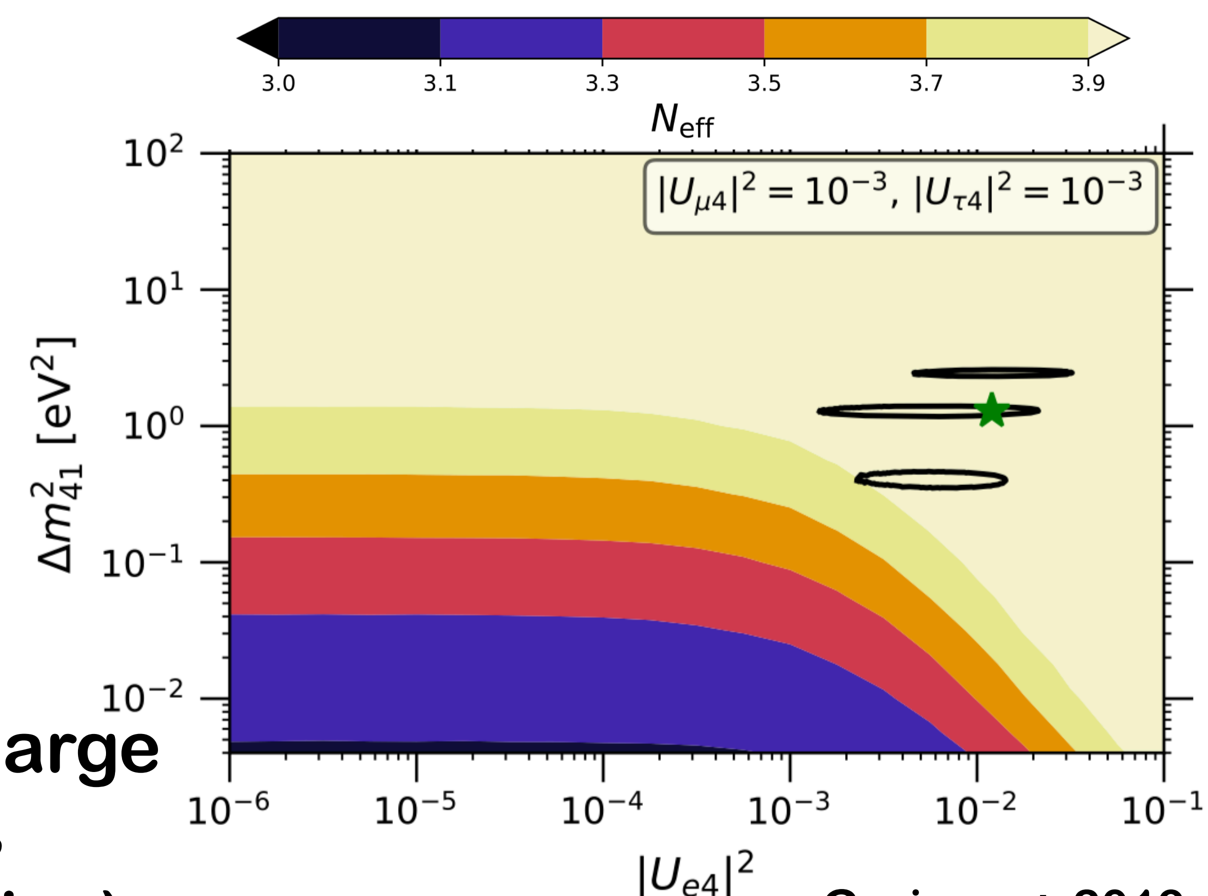
Anomalies in oscillations would require light sterile with large mixing angle.

If they exist, oscillations in the early Universe would create a population of sterile

$$\theta \qquad \Delta m^2$$

$$\rightarrow$$

$$\Delta N_{\rm eff} \qquad \Omega_{\nu,s} h^2 = \frac{m_{\rm eff}}{94\,{\rm eV}}$$

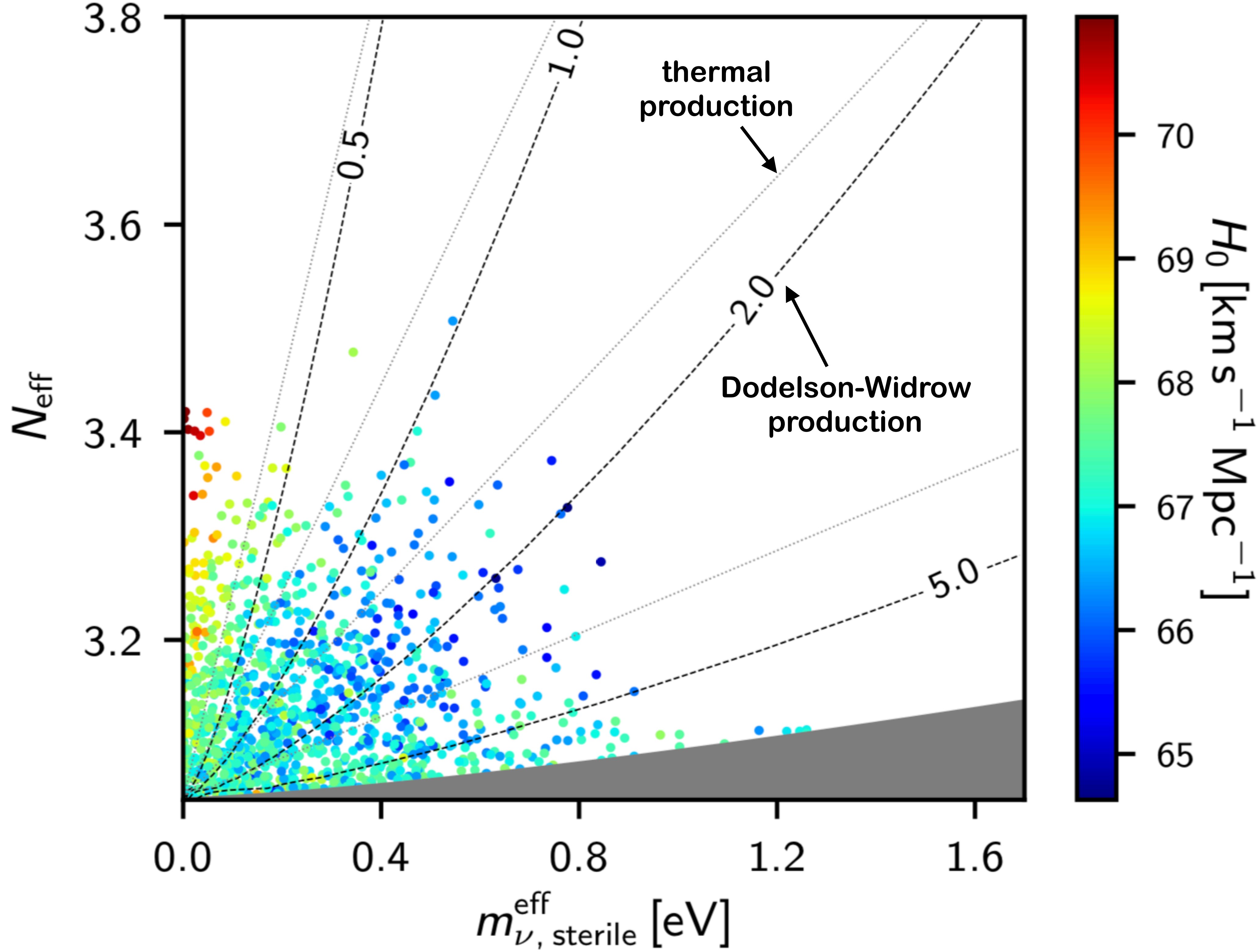


Gariazzo+,2019

Lab best fit is at odds with cosmology: too large contribution to Neff for large mixing angles (quick thermalisation of the sterile with active)

Current limits on Neff-meff

Planck collaboration, VI 2018



~eV thermalised sterile neutrino excluded at 7sigma

Non-standard models needed to make SBL compatible with cosmology

What next in neutrino cosmology

A new generation of ultimate cosmological surveys is approaching: Simons Observatory, Euclid, LiteBIRD, CMB-S4, DESI, LSST, SKA …

Does it mean that we are moving:

1) **Towards the first detection of the neutrino mass scale?**

2) **Towards the first probe of the physics of neutrino decoupling, and of BSM content at very early times?**

How to achieve robustness

1) **External redundancy and complementarity: Individual probes of very high and comparable sensitivity, cross-correlations,…**
2) **Internal redundancy and complementarity: Individual channels (e.g. temperature and polarisation; shear and galaxy; …) of high and comparable sensitivity**
3) **Know your instruments: extensive work on ‘mock’ calibration, sensitivity and systematics in preparation of the real analysis**

Route to robust neutrino mass bounds

- CMB lensing from SO combined with DESI BAO

$$\sigma(\Sigma m_\nu) = 0.04\,\mathrm{eV}\ [0.03\,\mathrm{eV}]$$

- Sunyaev-Zeldovich cluster counts from SO calibrated with LSST weak lensing

$$\sigma(\Sigma m_\nu) = 0.04\,\mathrm{eV}\ [0.03\,\mathrm{eV}]$$

- thermal SZ distortion maps from SO combined with DESI BAO

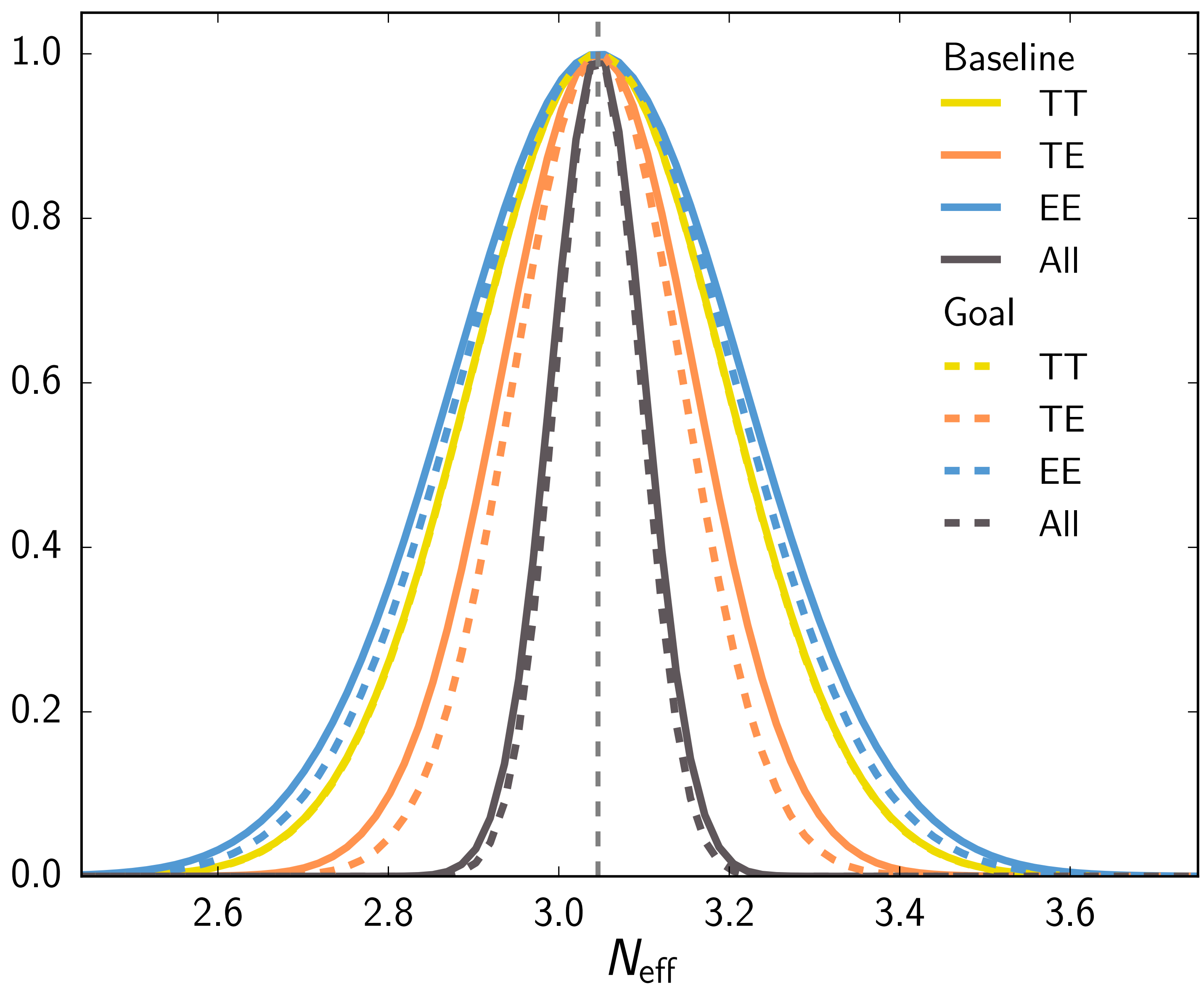
$$\sigma(\Sigma m_\nu) = 0.05\,\mathrm{eV}\ [0.04\,\mathrm{eV}]$$

- legacy SO dataset combined with cosmic-variance-limited measurement of reionization optical depth τ

$$\sigma(\Sigma m_\nu) = 0.02\,\mathrm{eV}$$

SO collaboration, 2018

Route to improved bounds on Neff



Primary CMB temperature and polarization power spectra from SO

$$\sigma(N_{\rm eff}) = 0.07\ [0.05]$$

SO collaboration, 2018

CONCLUSIONS

Determine CnB properties from neutrino peculiar effects on cosmological observables

Strong and robust constraints from cosmology

Neutrino masses: getting closer to cornering inverting hierarchy

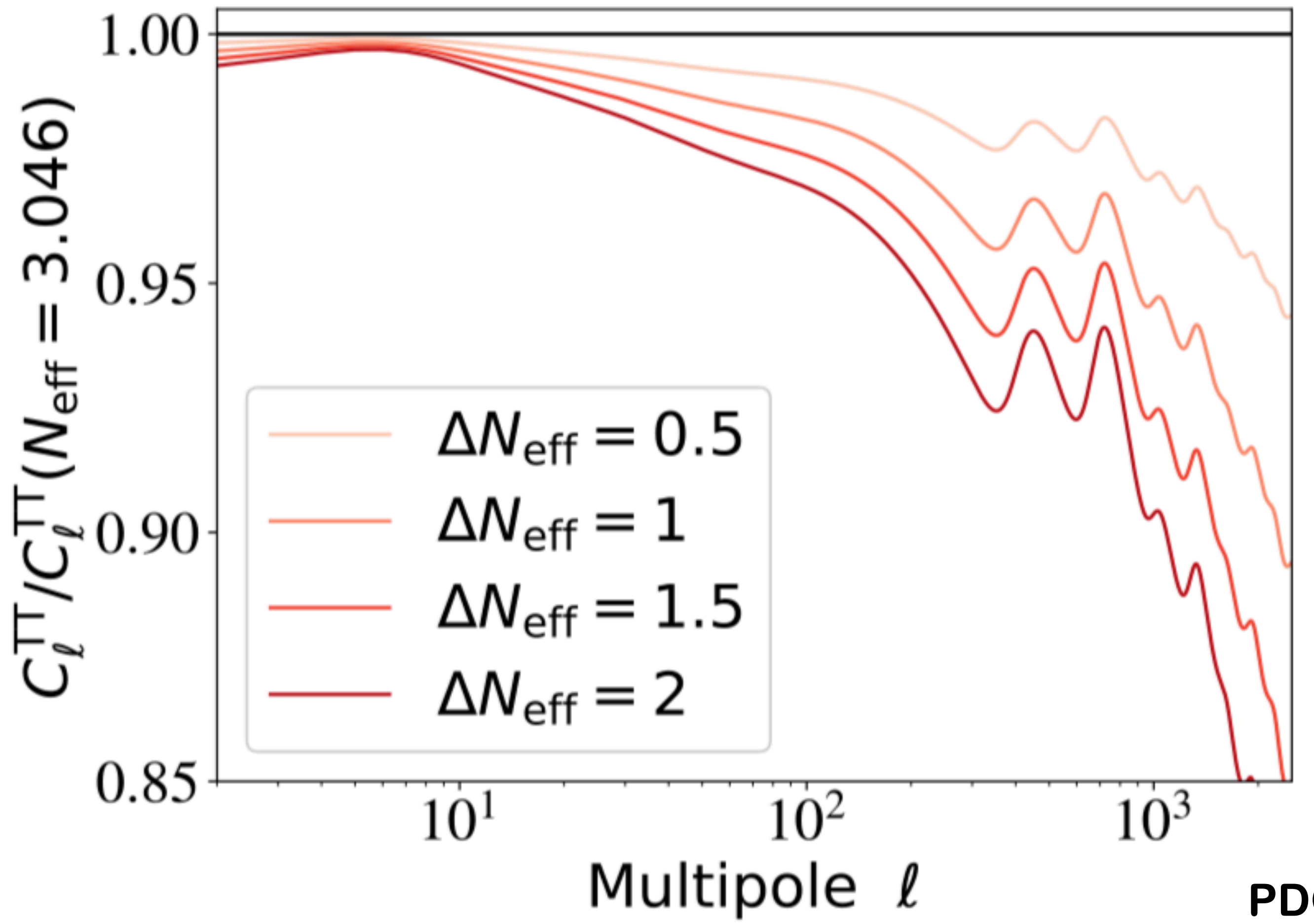
Neff: no preference for an additional thermalised species

Next generation surveys would probe the physics of non-instantaneous decoupling and detect the neutrino mass scale with high statistical significance

They would probe BSM scenarios and be complementary to lab and astro searches

BACKUP SLIDES

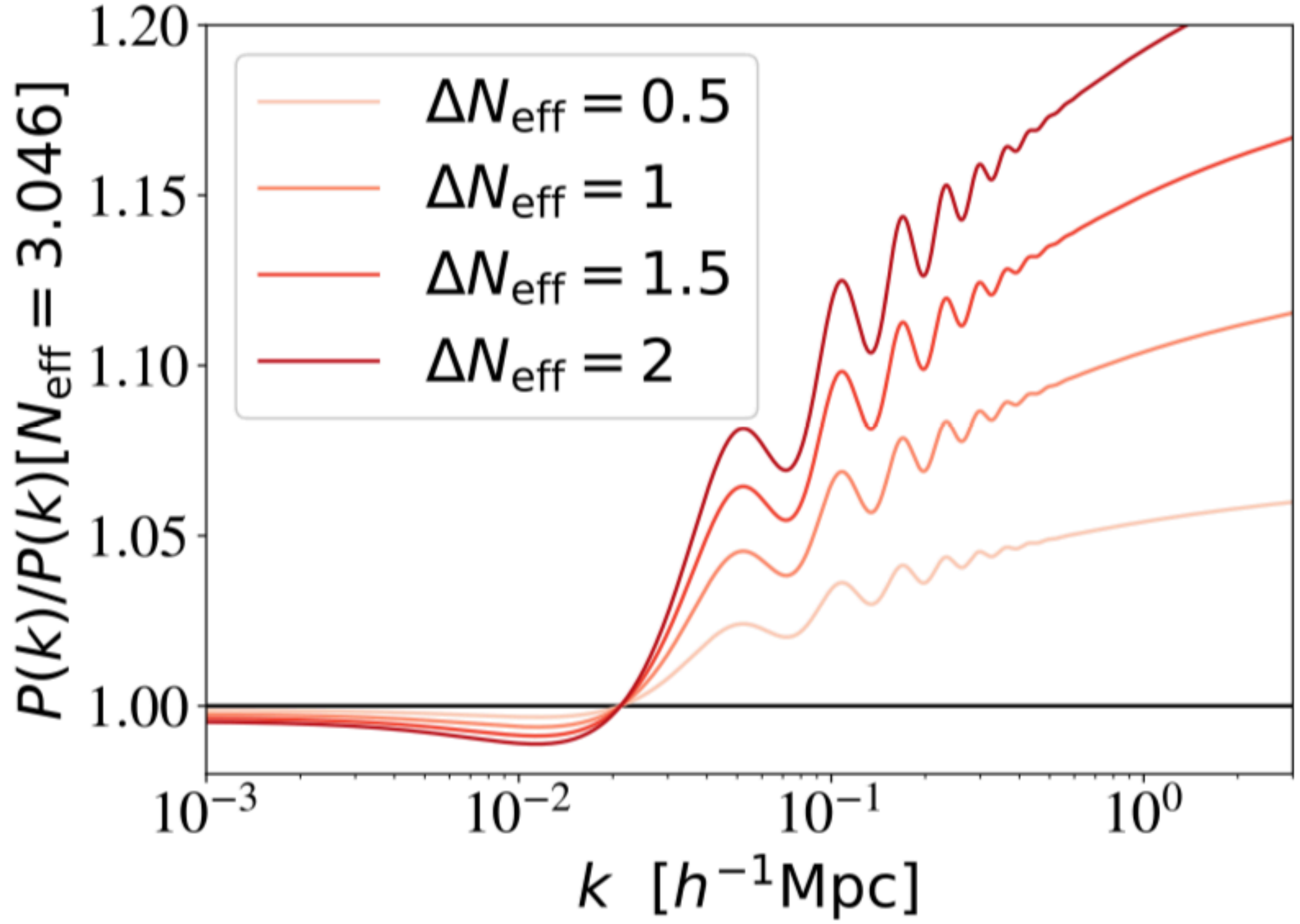
CMB power spectrum



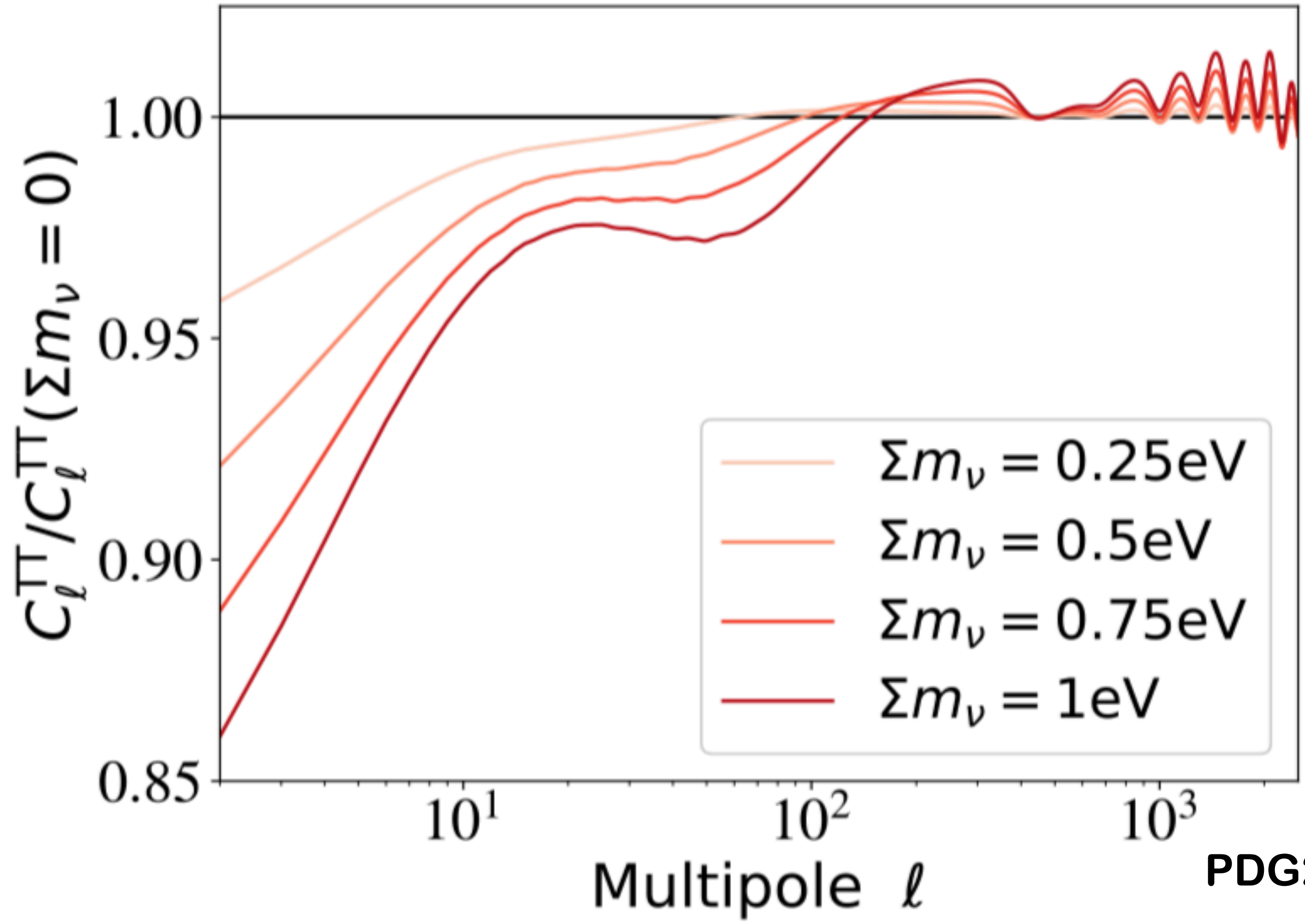
PDG2018

Figures obtained with fixed z_eq,z_L, obh2,tau

Matter power spectrum



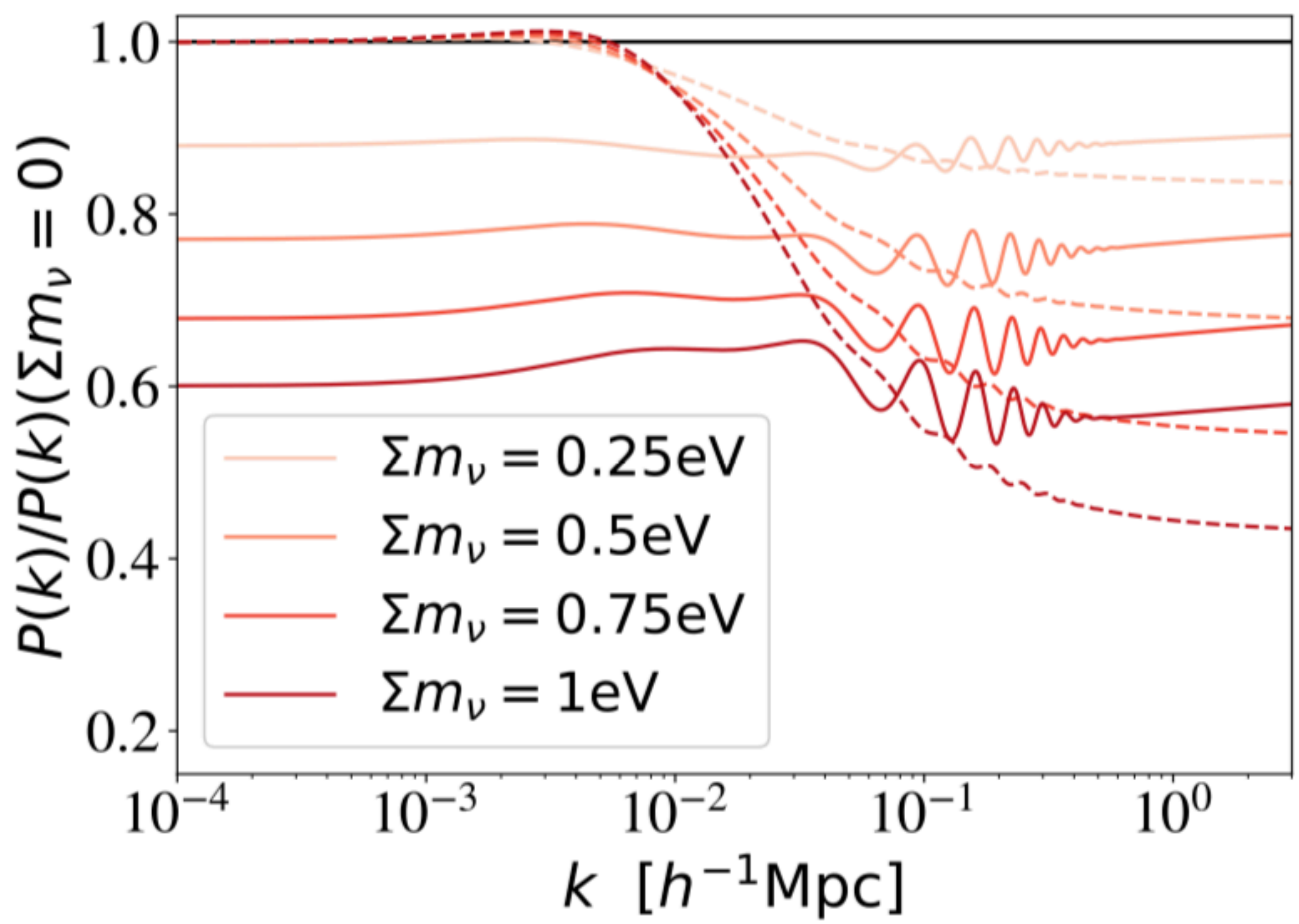
CMB power spectrum



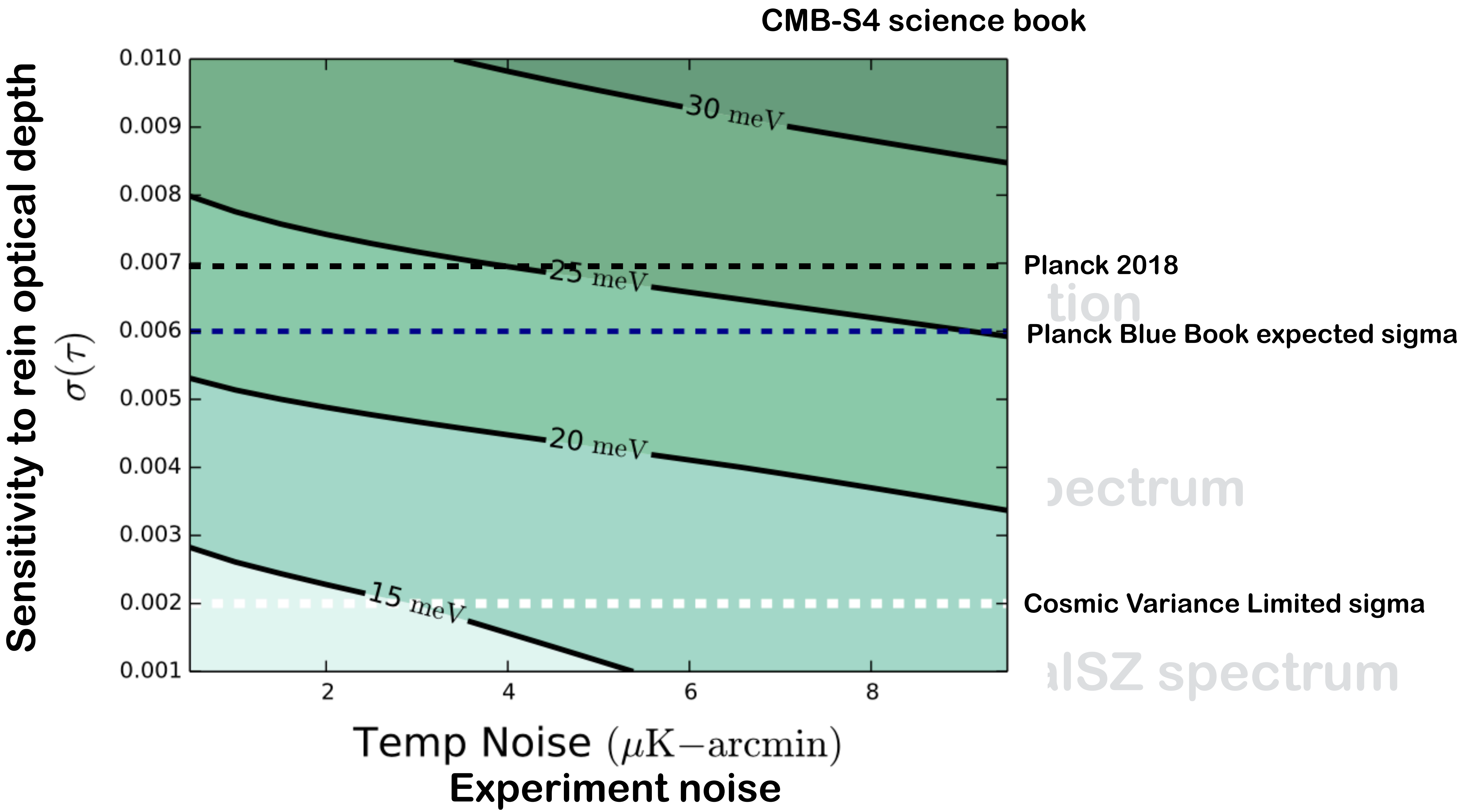
PDG2018

Figures obtained with fixed och2, obh2,tau,theta (solid); fixed omh2,obh2, Olambda (dashed)

Matter power spectrum



Route to robust neutrino mass bounds



- Cosmic-variance-limited measurements of tau

Neff as a proxy for extra light species

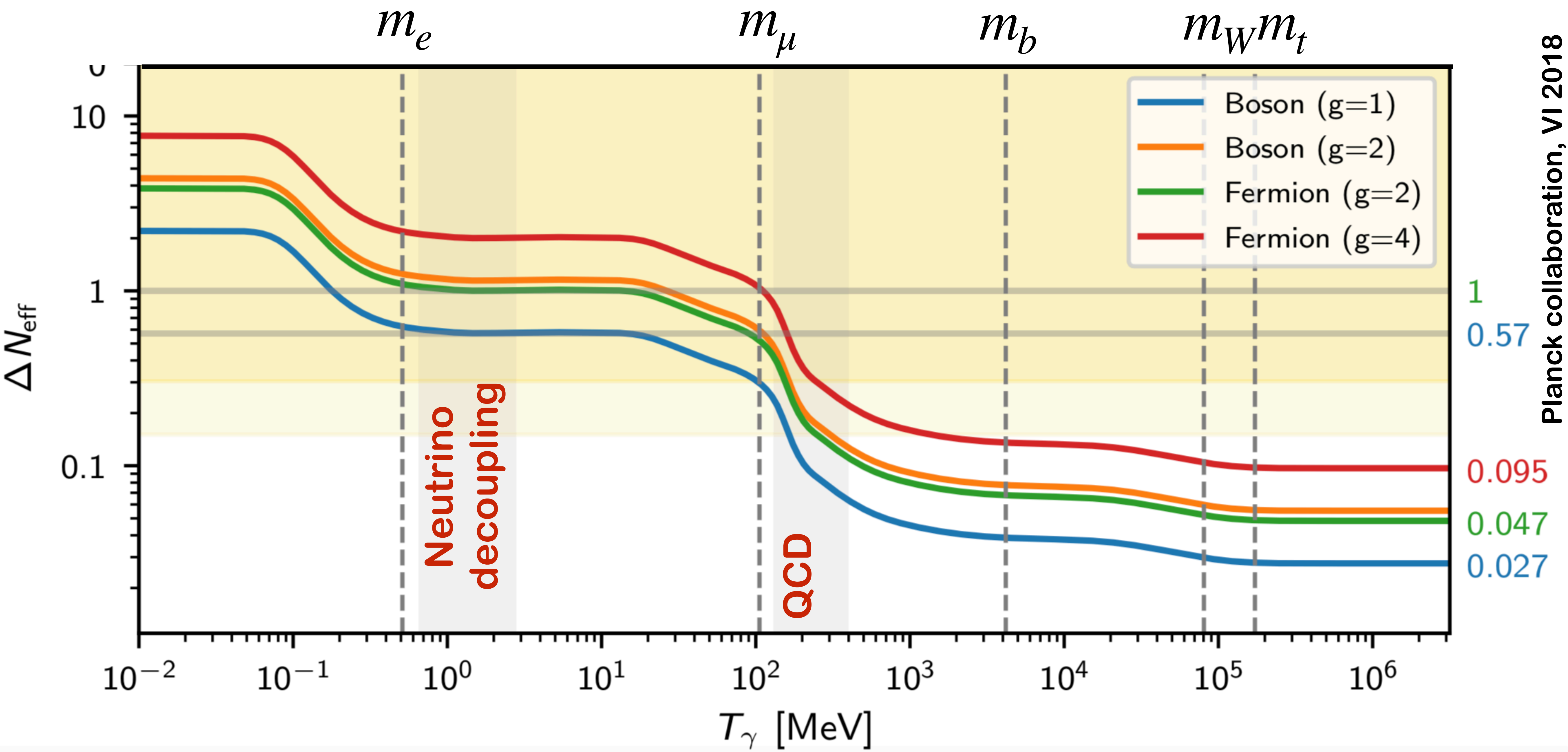
$$\rho_{\rm rad} = \rho_\gamma \left(1 + \frac{7}{8}\left(\frac{4}{11}\right)^{\frac{4}{3}} \left(N_{{\rm eff},\nu} + \Delta N_{\rm eff}\right)\right)$$

Light relics decoupling when thermalised with the primordial plasma

$$\Delta N_{\rm eff} = g \left(\frac{43}{4 g_{*,s}}\right)^{\frac{4}{3}} \times \begin{cases} 1/2, & \text{fermion} \\ 4/7, & \text{boson} \end{cases}$$

Effective dof at decoupling

Neff as a proxy for extra light species

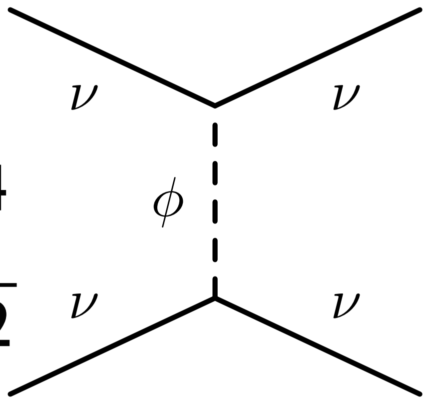


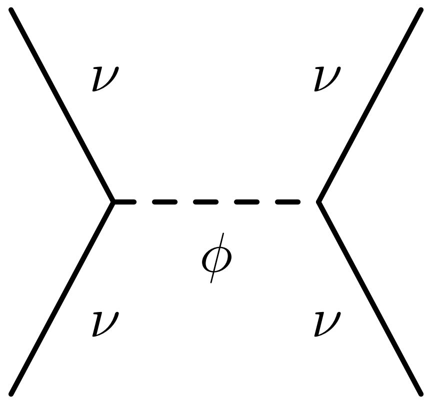
Presence of additional fully thermalised species decoupling after QCD phase transition excluded at 95%c.l.

Scalar-mediated neutrino interactions

$$\mathcal{L} \supset h_{ij}\bar{\nu}_i\nu_j\phi + g_{ij}\bar{\nu}_i\gamma_5\nu_j\phi + h.c.\,,$$

$$\Gamma_{\nu\nu} = \langle\sigma v\rangle n_{\rm eq} \propto g^4 T \qquad \sigma \sim \frac{g^4}{s} \sim \frac{g^4}{T^2}$$





Collisional processes can suppress stress and affect the perturbation evolution of cosmological neutrinos.

Parameterizing the interaction strength through:

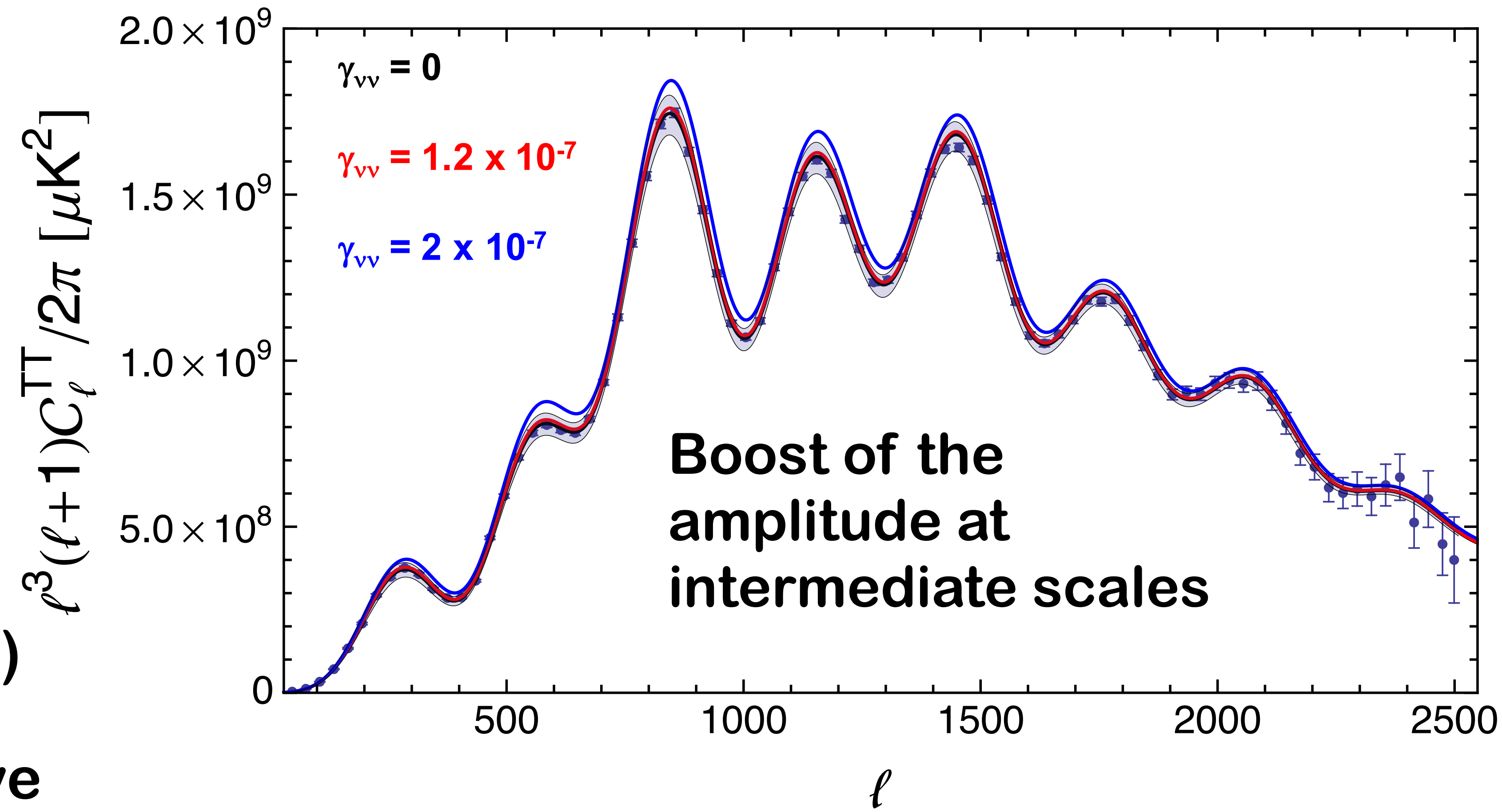
$$\Gamma_{\nu\nu} \equiv \gamma_{\nu\nu}^4 T$$

Planck constraints (95% CL):

$$\gamma_{\nu\nu}^4 < (0.3 \div 0.5) \times 10^{-27}$$

i.e $\gamma_{\nu\nu} \lesssim 10^{-7}$ **(Forastieri+2018)**

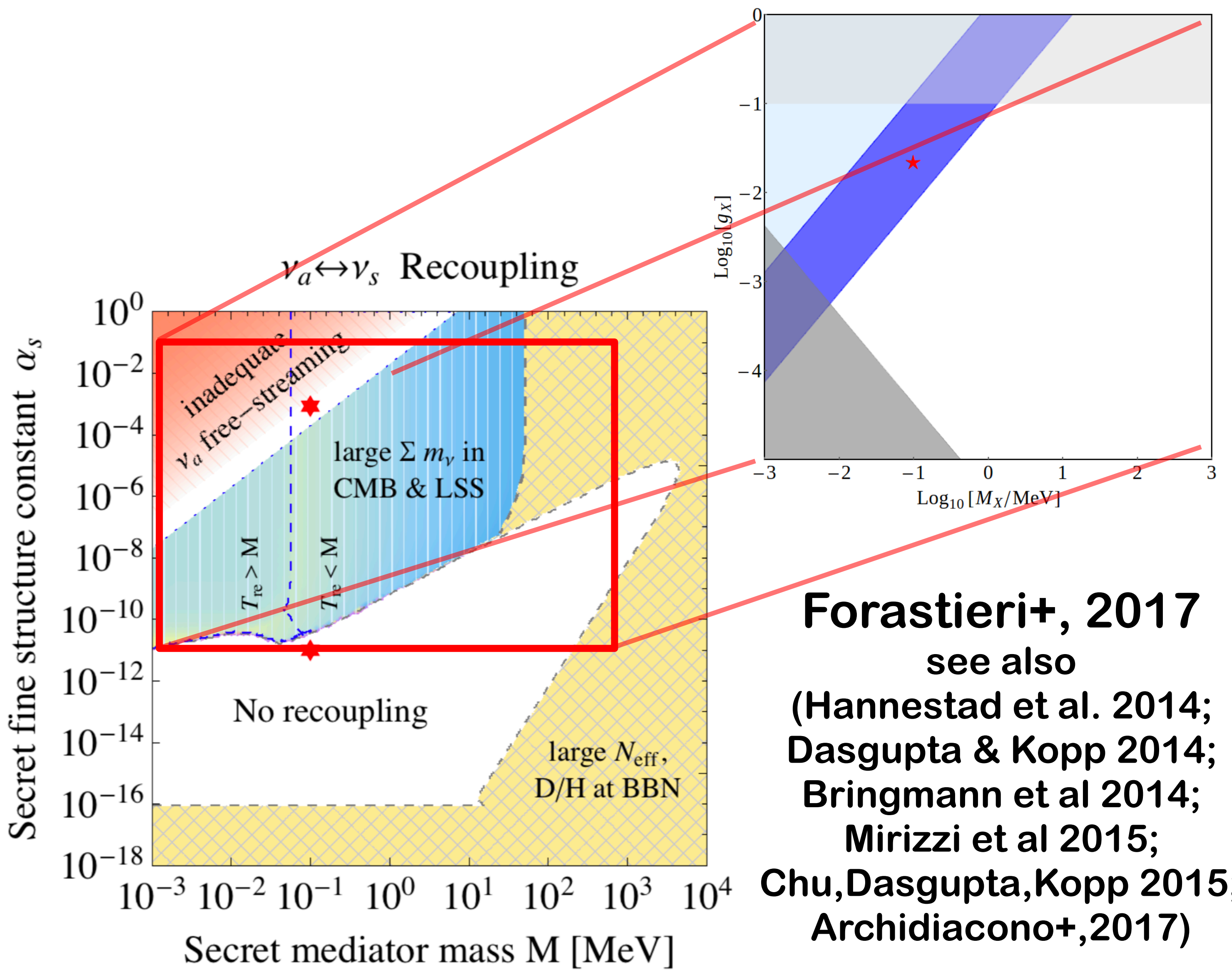
Future CMB bounds will improve by ~ 1 order of magnitude in $\gamma_{\nu\nu}^4$



Forastieri+ 2015; Archidiacono & Hannestad 2013; Cyr-Racine & Sigurdson 2013

Sterile neutrino interactions and SBL anomalies

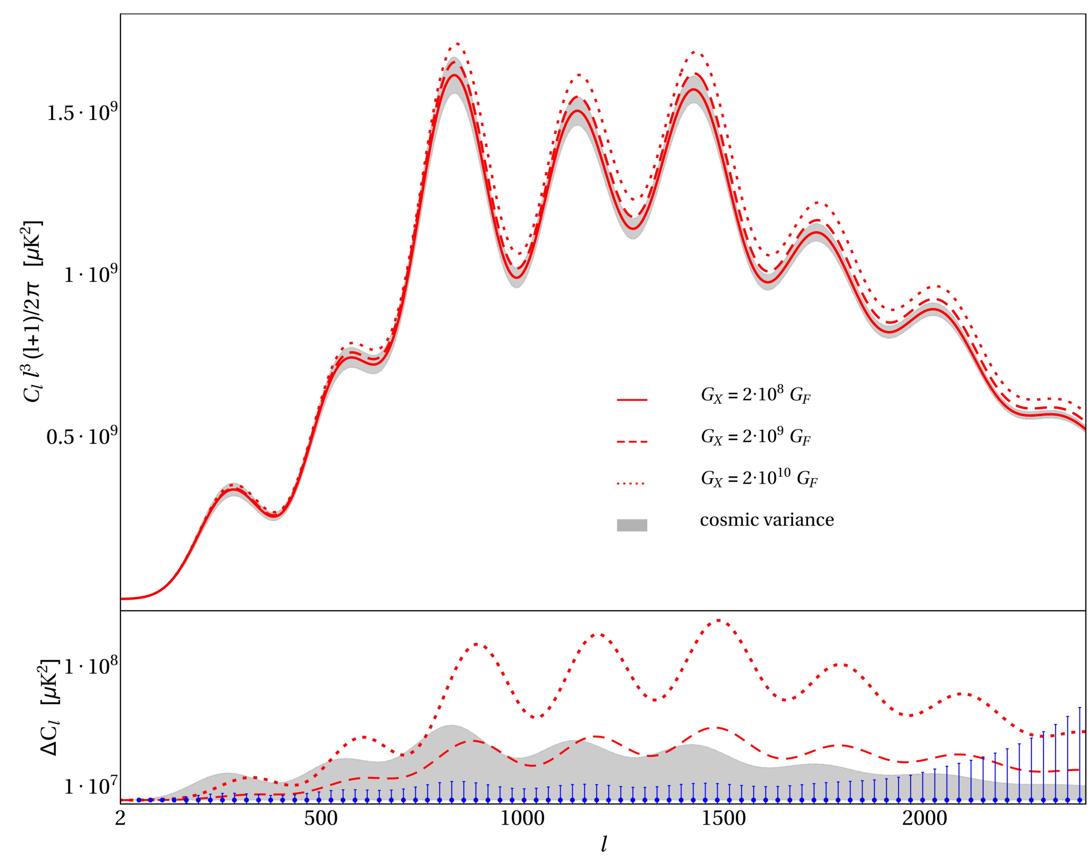
Sterile neutrino interpretation of SBL is in disagreement with cosmology (implies ΔN_{eff}=1)
Are "Secret" interactions in the sterile sector a way out?
Production of sterile neutrinos is delayed, but large values of G_X will leave an observational signature on the CMB spectrum.



$$\mathcal{L}_s = g_X \bar{\nu}_s \gamma_\mu \frac{1}{2}(1-\gamma_5)\, \nu_s X^\mu$$

$$G_X \sim g_X^2/M_X^2$$

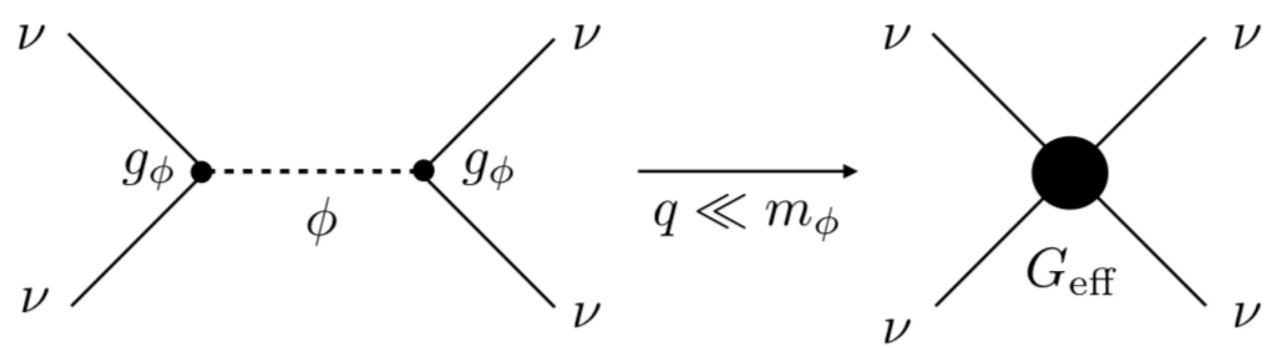
Forastieri+, 2017
see also
(Hannestad et al. 2014; Dasgupta & Kopp 2014; Bringmann et al 2014; Mirizzi et al 2015; Chu,Dasgupta,Kopp 2015; Archidiacono+,2017)



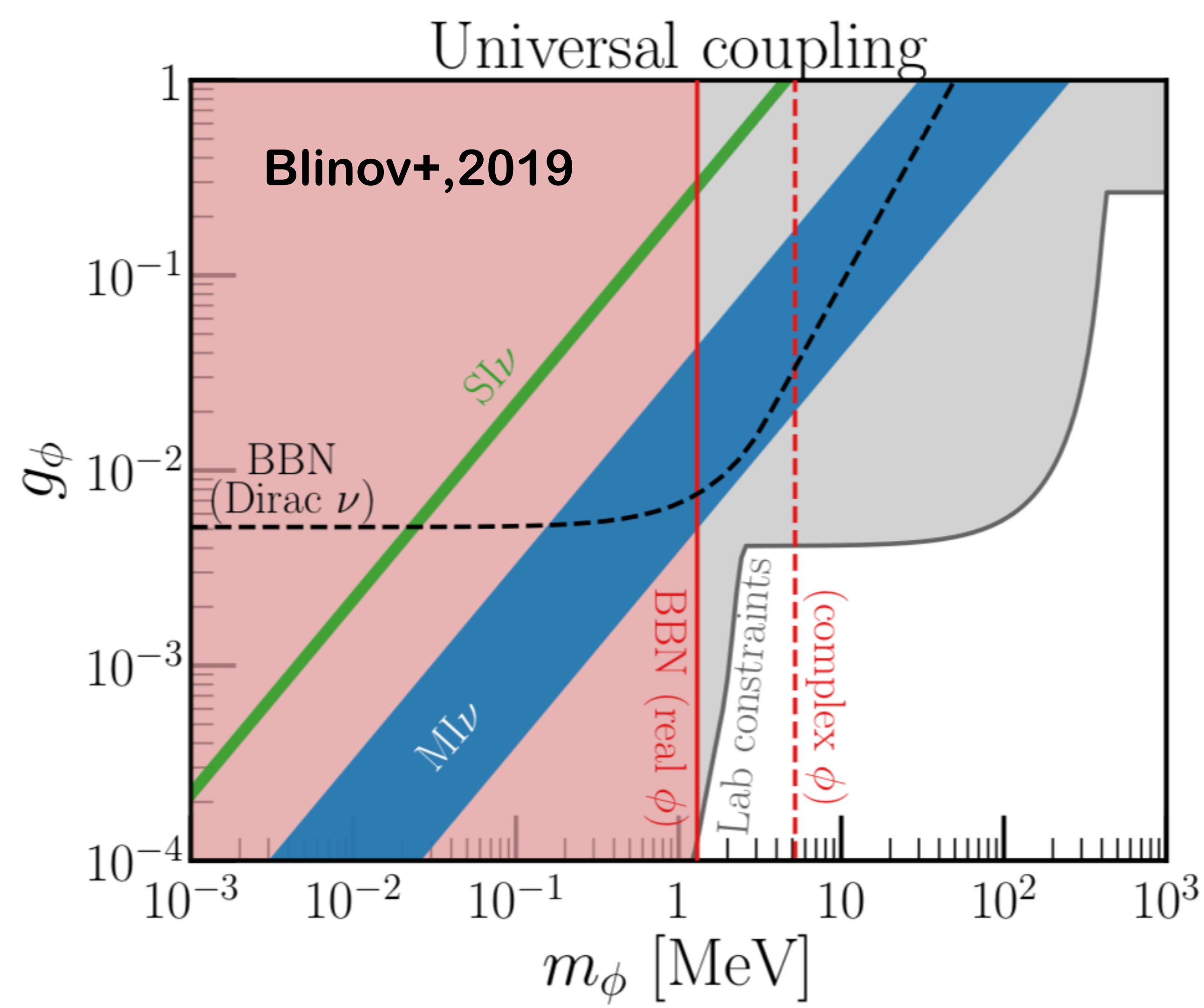
Scalar-mediated neutrino interactions

HEAVY MEDIATOR (m_phi>keV)

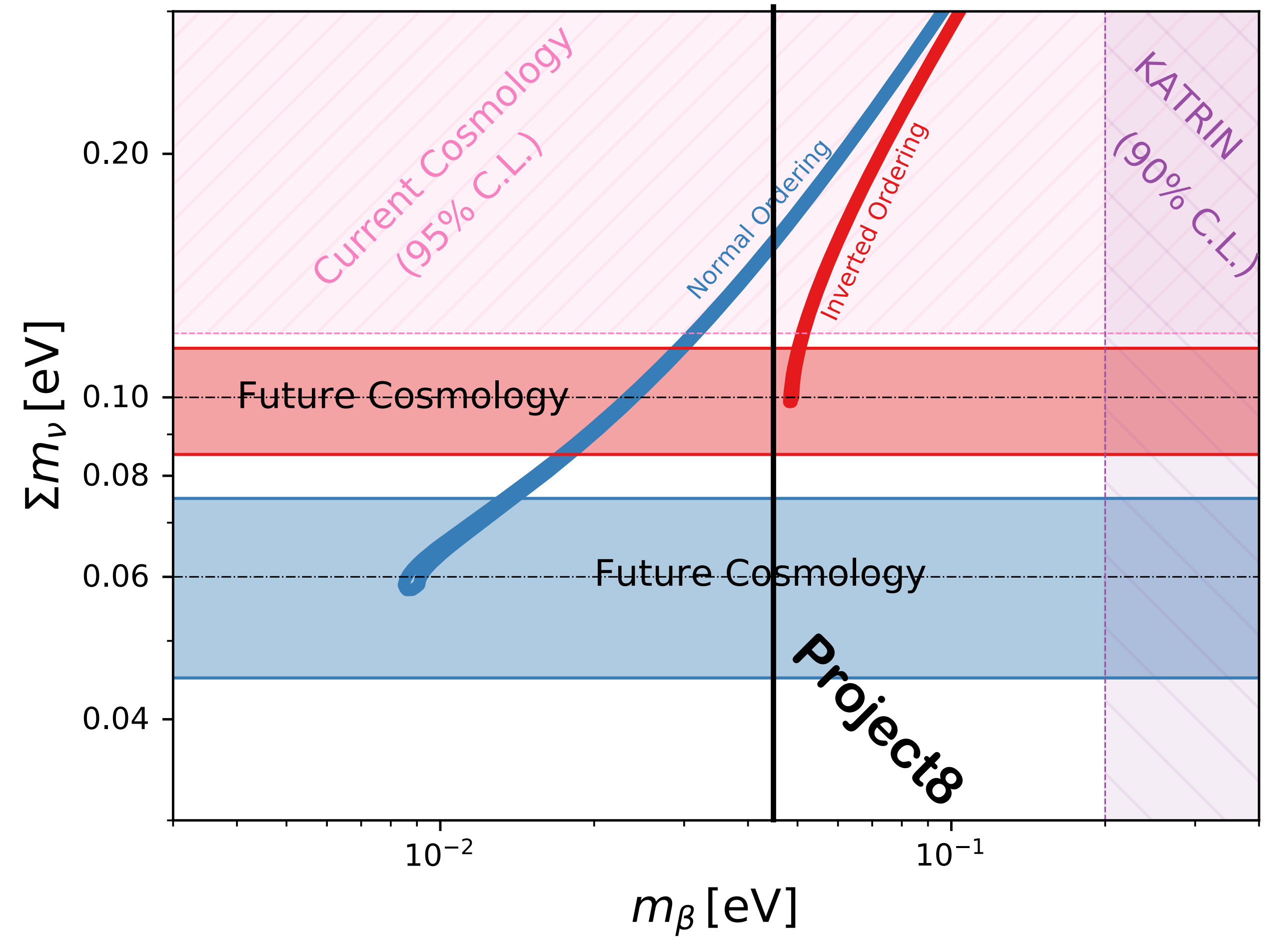
Parameterizing the universal coupling g_ij=g_nu delta_ij as Geff=g_nu^2/m_phi^2



Mildly interacting -> Geff<10^-4 MeV^-2
Strongly interacting -> 10^-2 MeV^-2<Geff<10^-1 MeV^-2

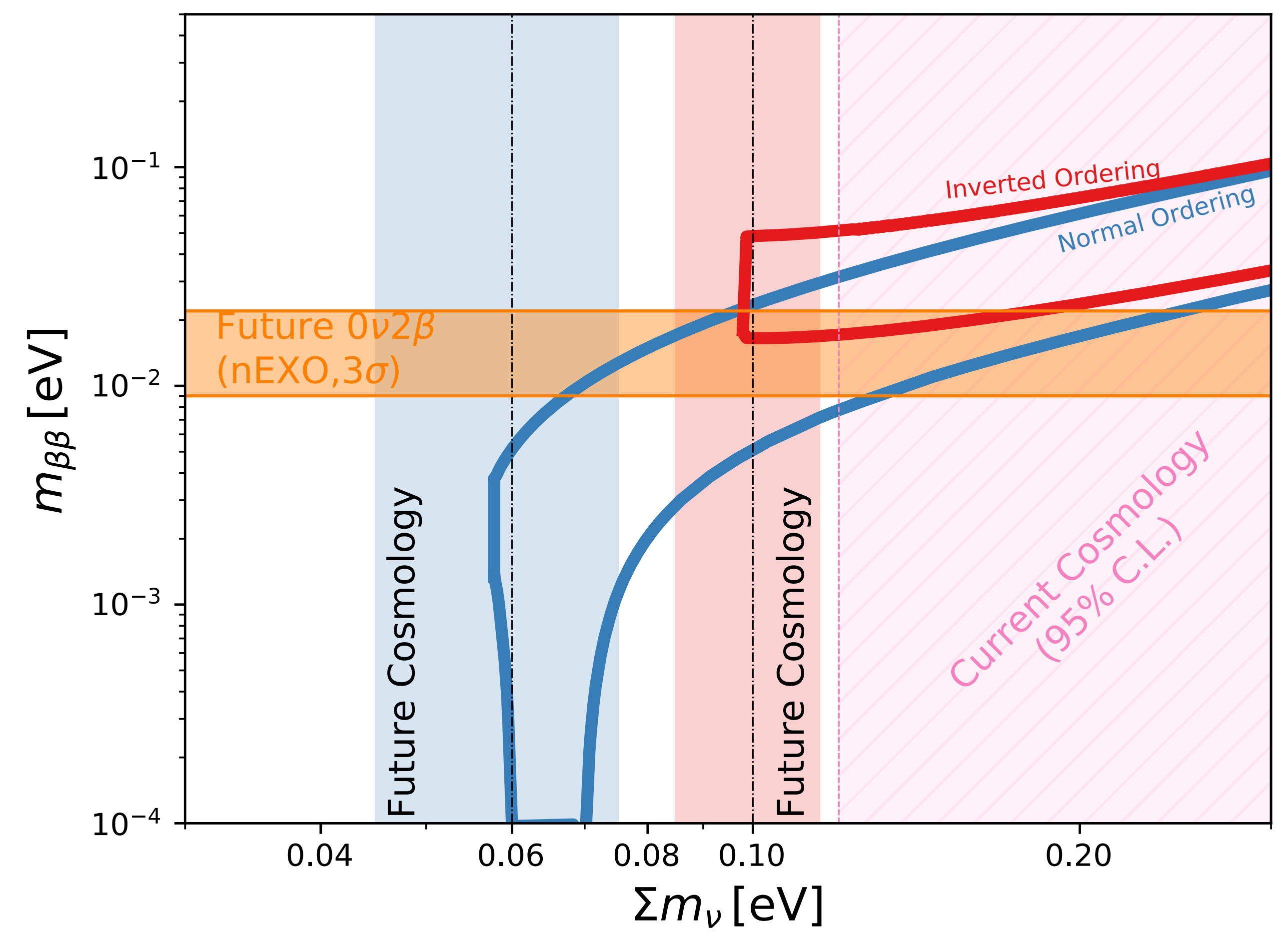


Expected sensitivity from S4-like surveys
~15 meV

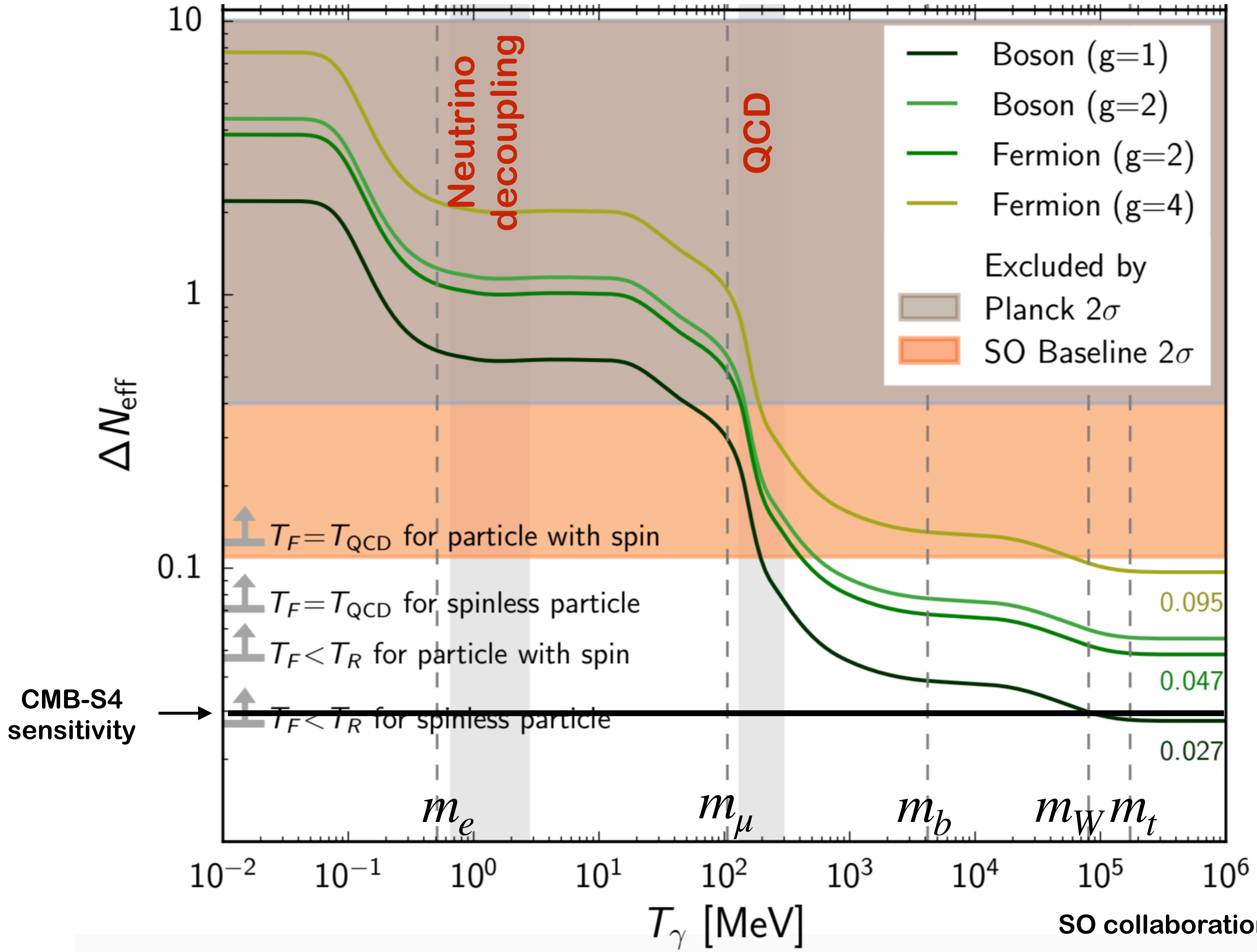


Complementarity with laboratory searches:

1) **independent cross-checks**
2) **interesting scenarios if in disagreement**



Route to improved bounds on Neff

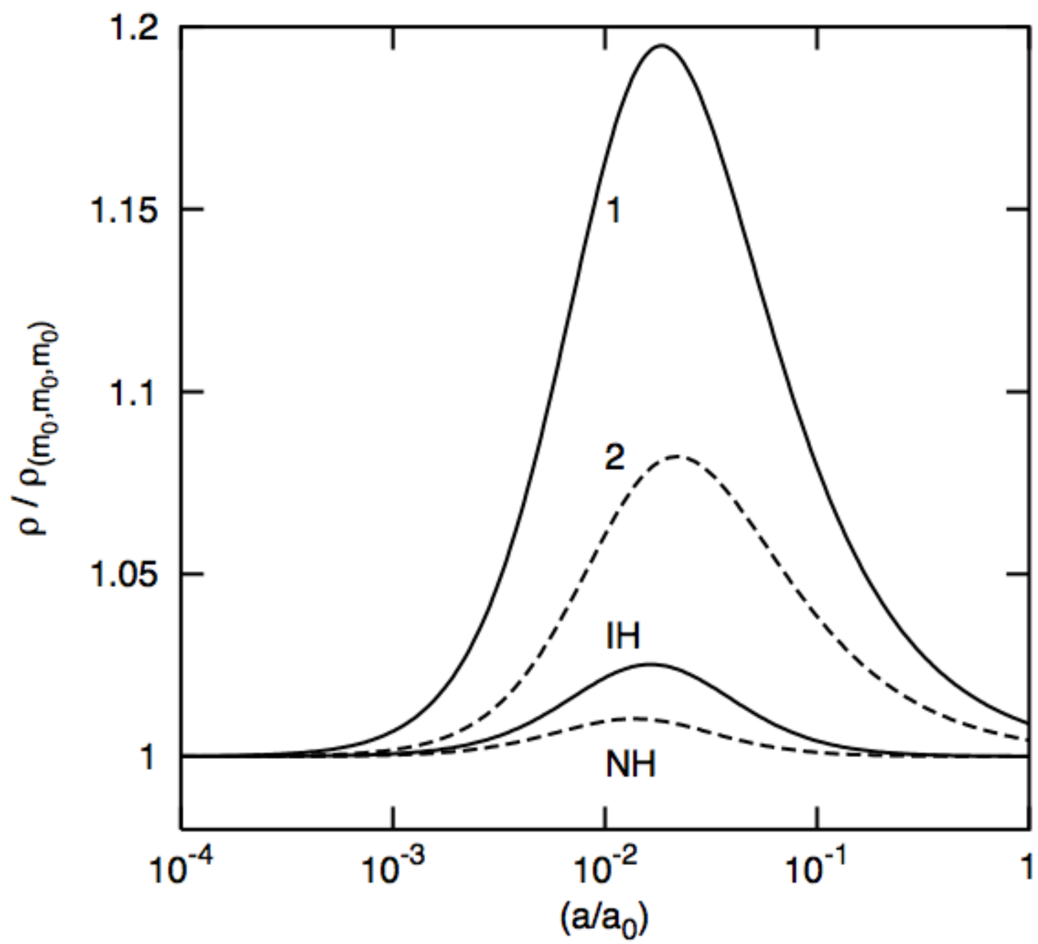


SO collaboration, 2018

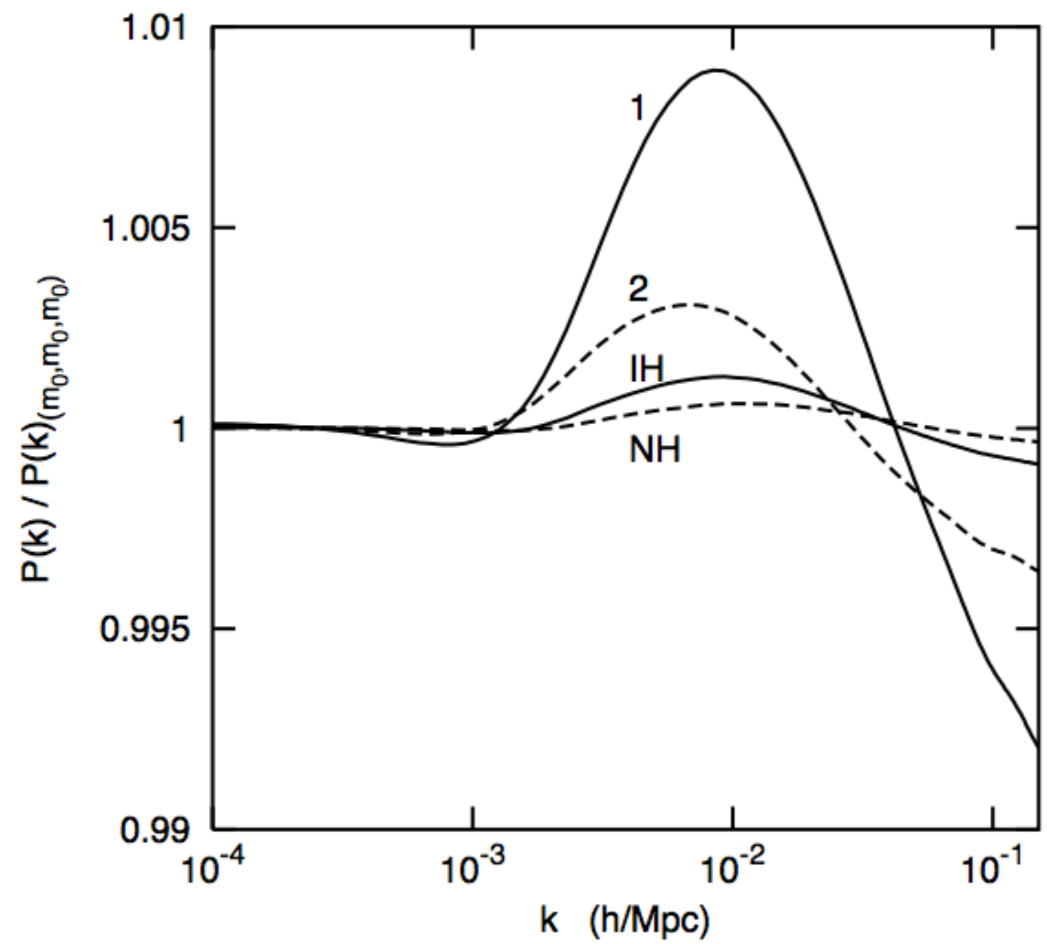
Sensitivity to the hierarchy

Physical effects due to different distribution of the sum of the masses for the 2 hierarchies

Total nu energy density



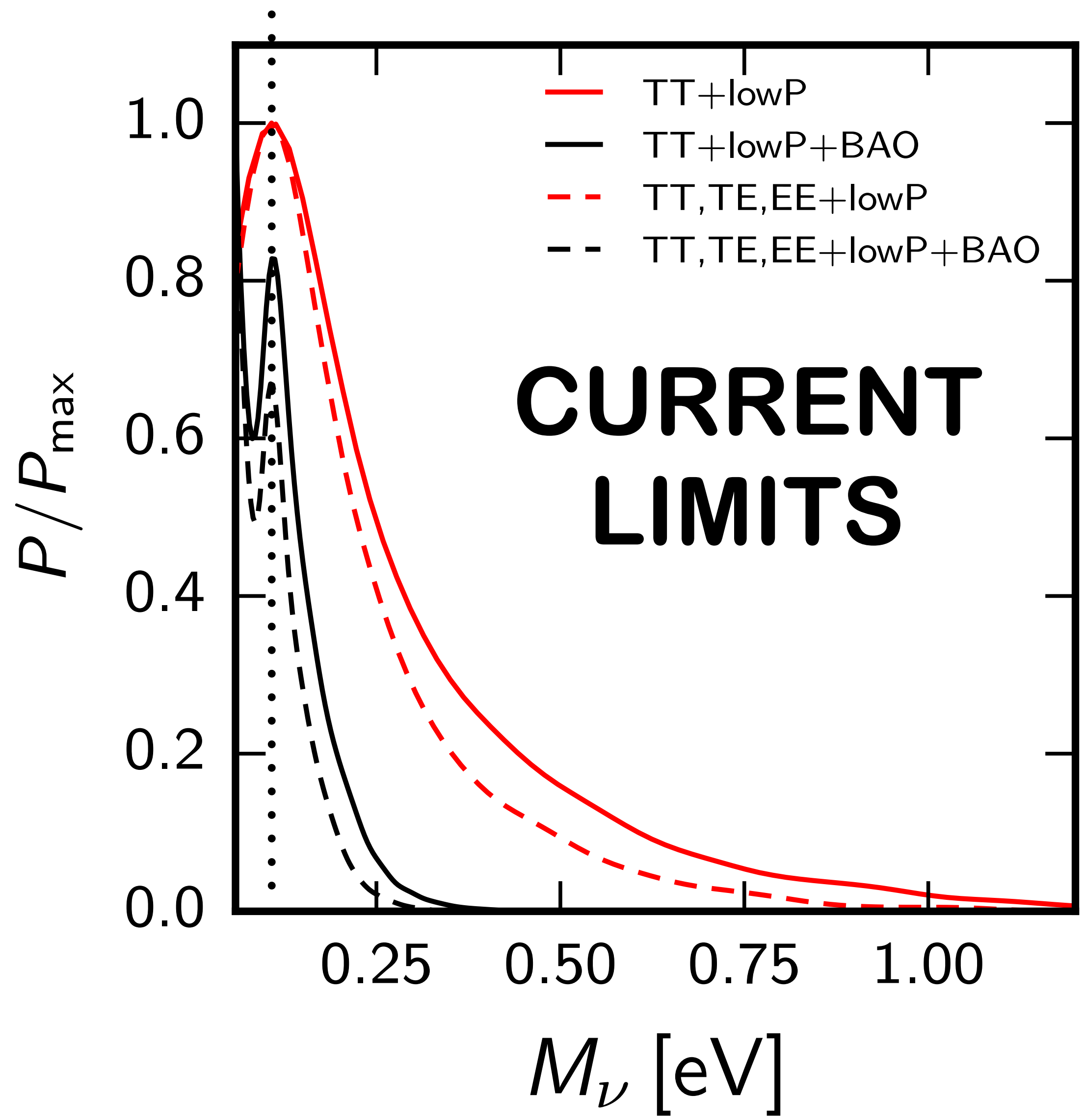
Matter power spectrum

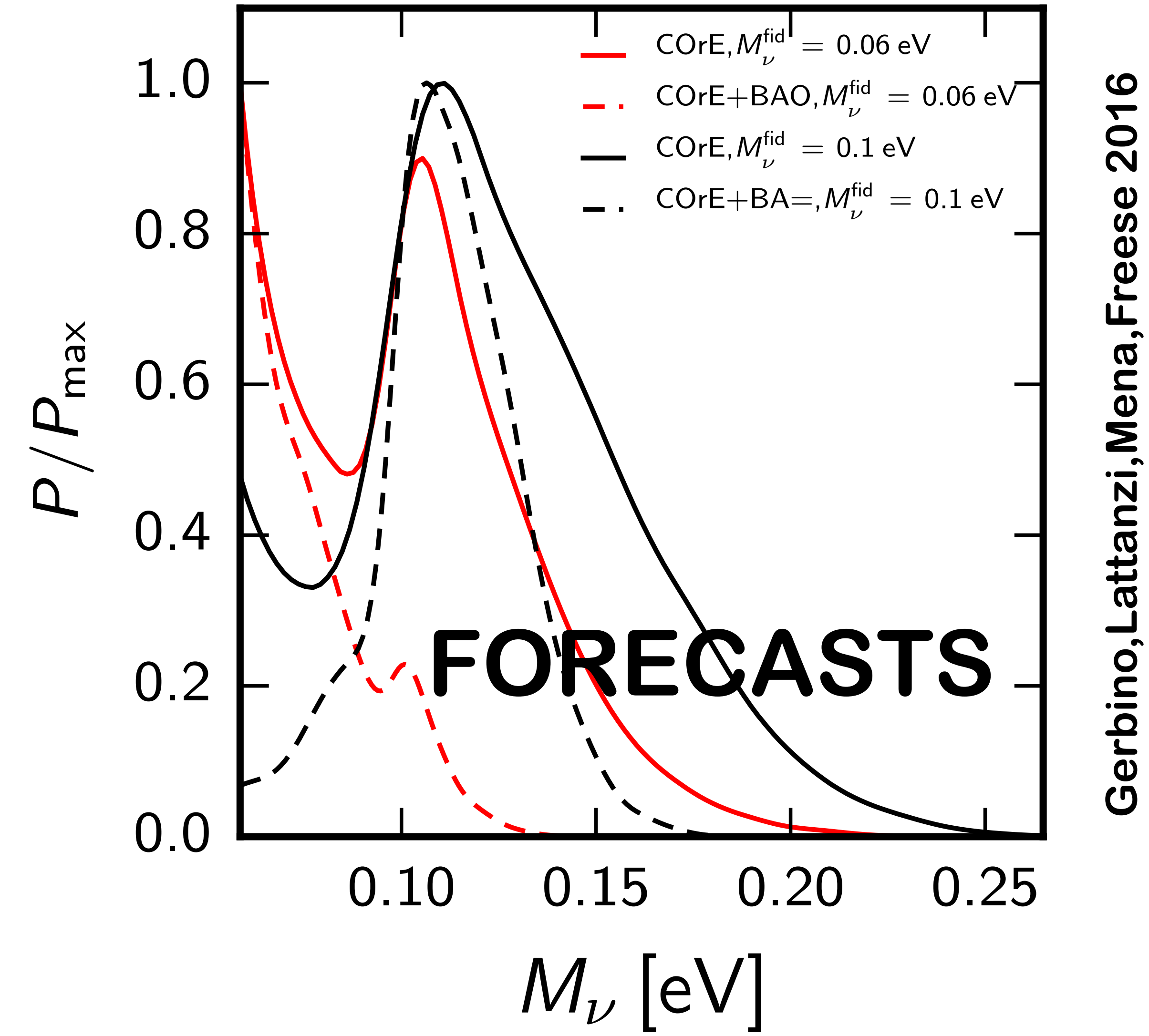


Lesgourgues&Pastor, 2006

Are current (and future) data sensitive to these effects? How much?

Sensitivity to the hierarchy





$$\mathcal{P}(h = NH) : \mathcal{P}(h = IH)$$

$$\mathcal{P}(h = NH) : \mathcal{P}(h = IH)$$

.... **3:2**

.... **0.06eV mass -> 9:1**

.... **0.1eV mass -> 1:1**

See also Hannestad&Schwetz,2016