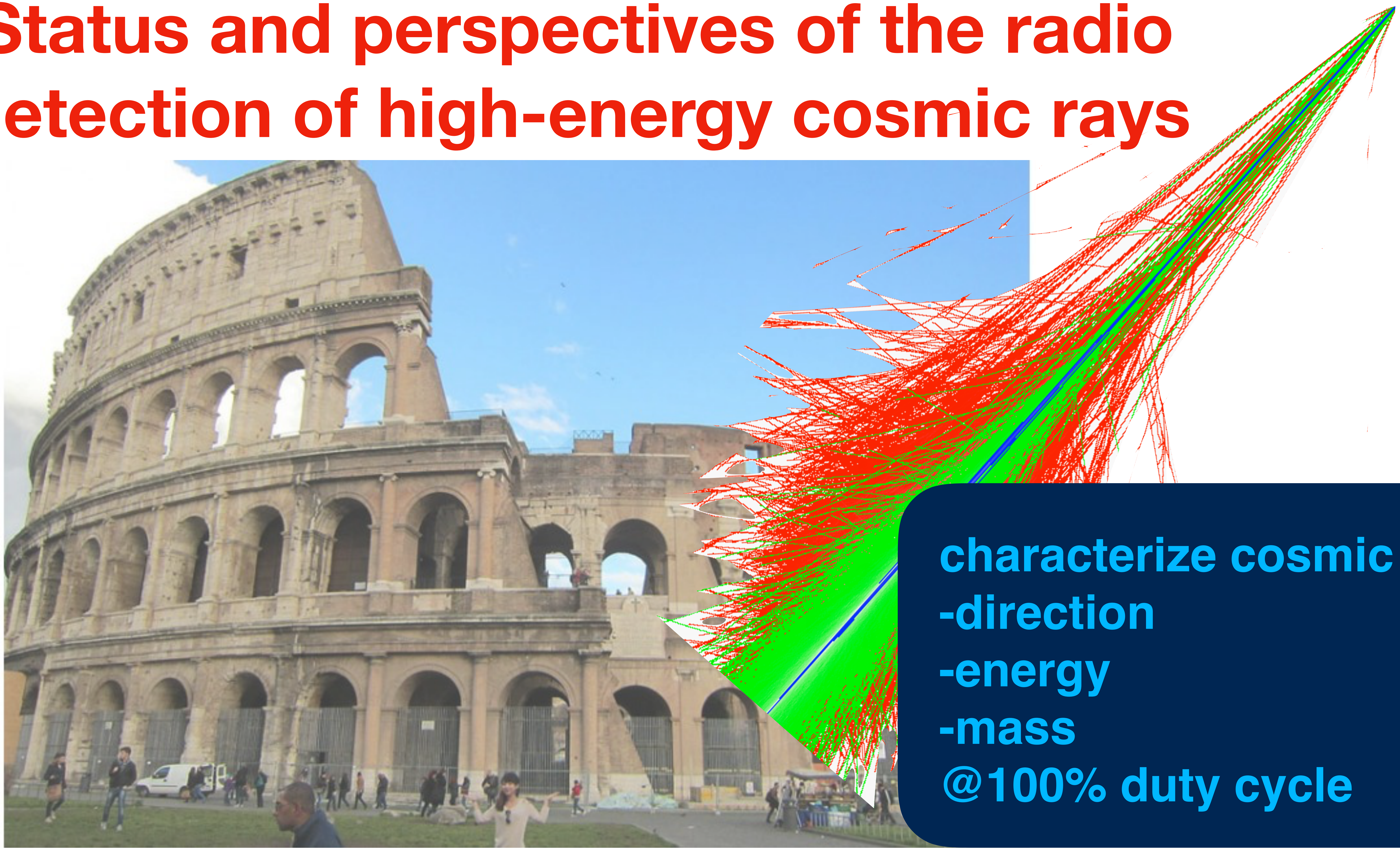
7th Roma International Conference on AstroParticle Physics

RICAP 18

Status and perspectives of the radio detection of high-energy cosmic rays



characterize cosmic rays:
-direction
-energy
-mass
@100% duty cycle

Nikhef
VUB
erc
LOFAR
PIERRE AUGER OBSERVATORY

taskleader radio at Pierre Auger Observatory

PI LOFAR key science project Cosmic Rays

Jörg R. Hörandel

Radboud University Nijmegen, Nikhef, VU Brussels

http://particle.astro.ru.nl

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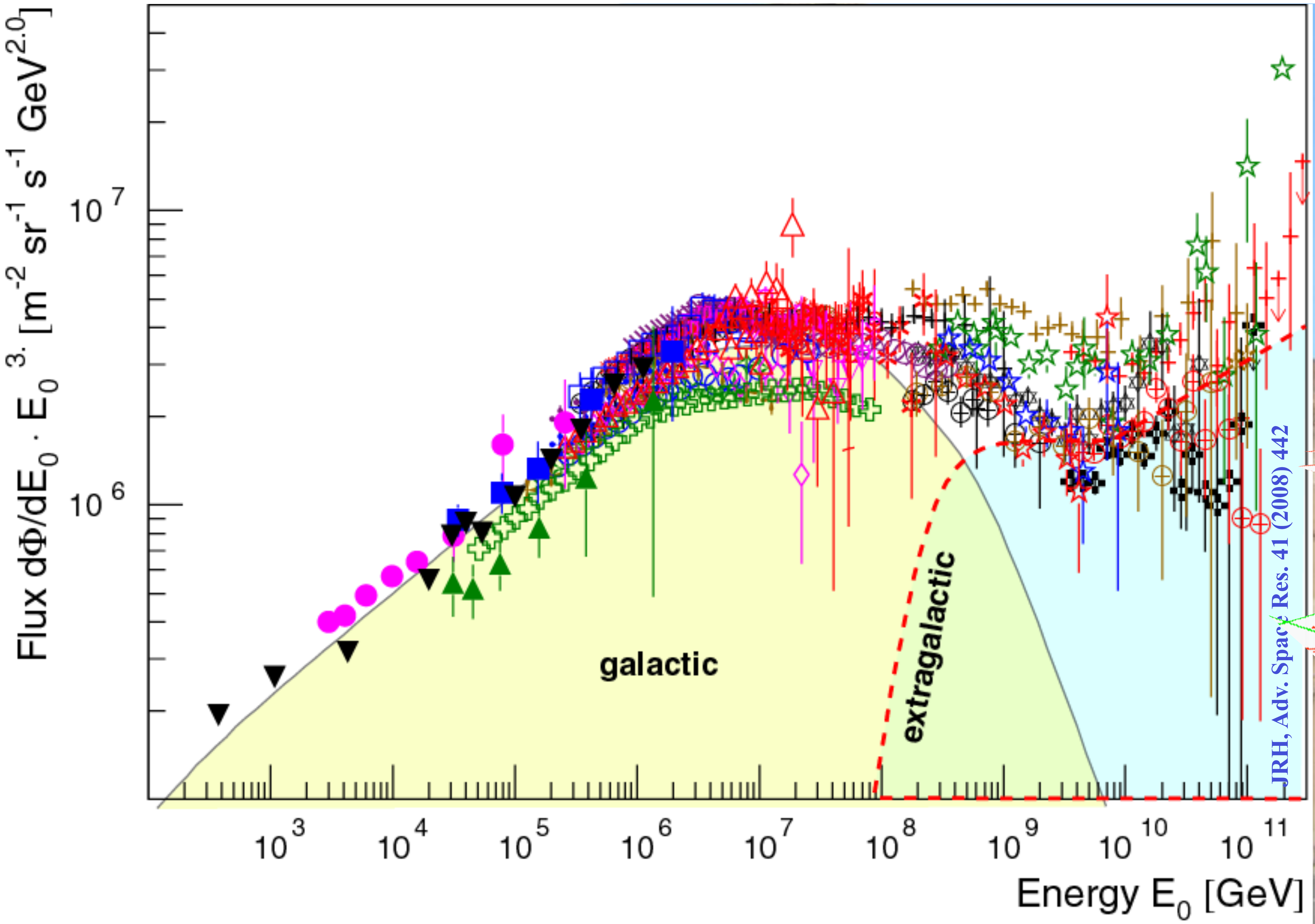

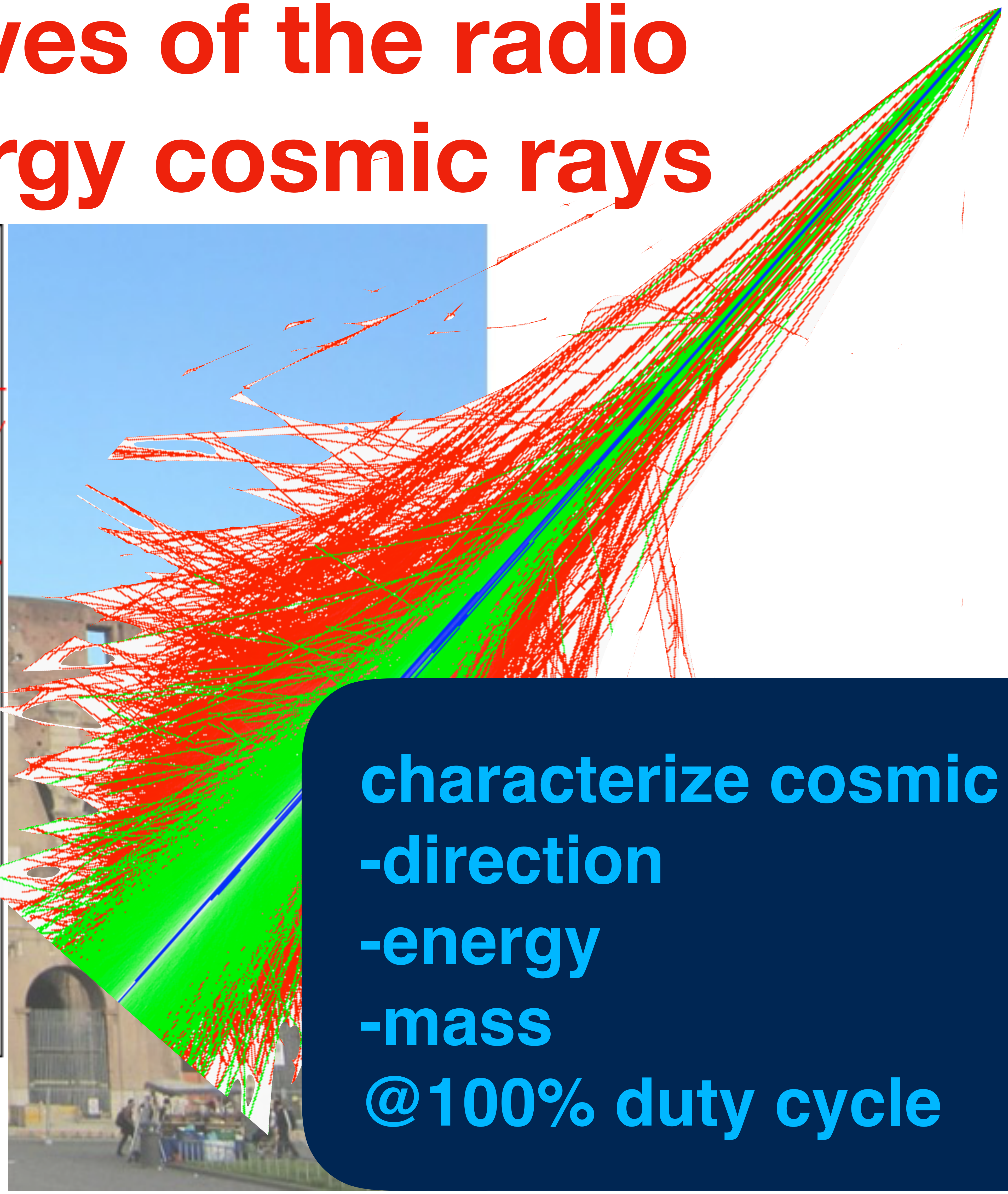

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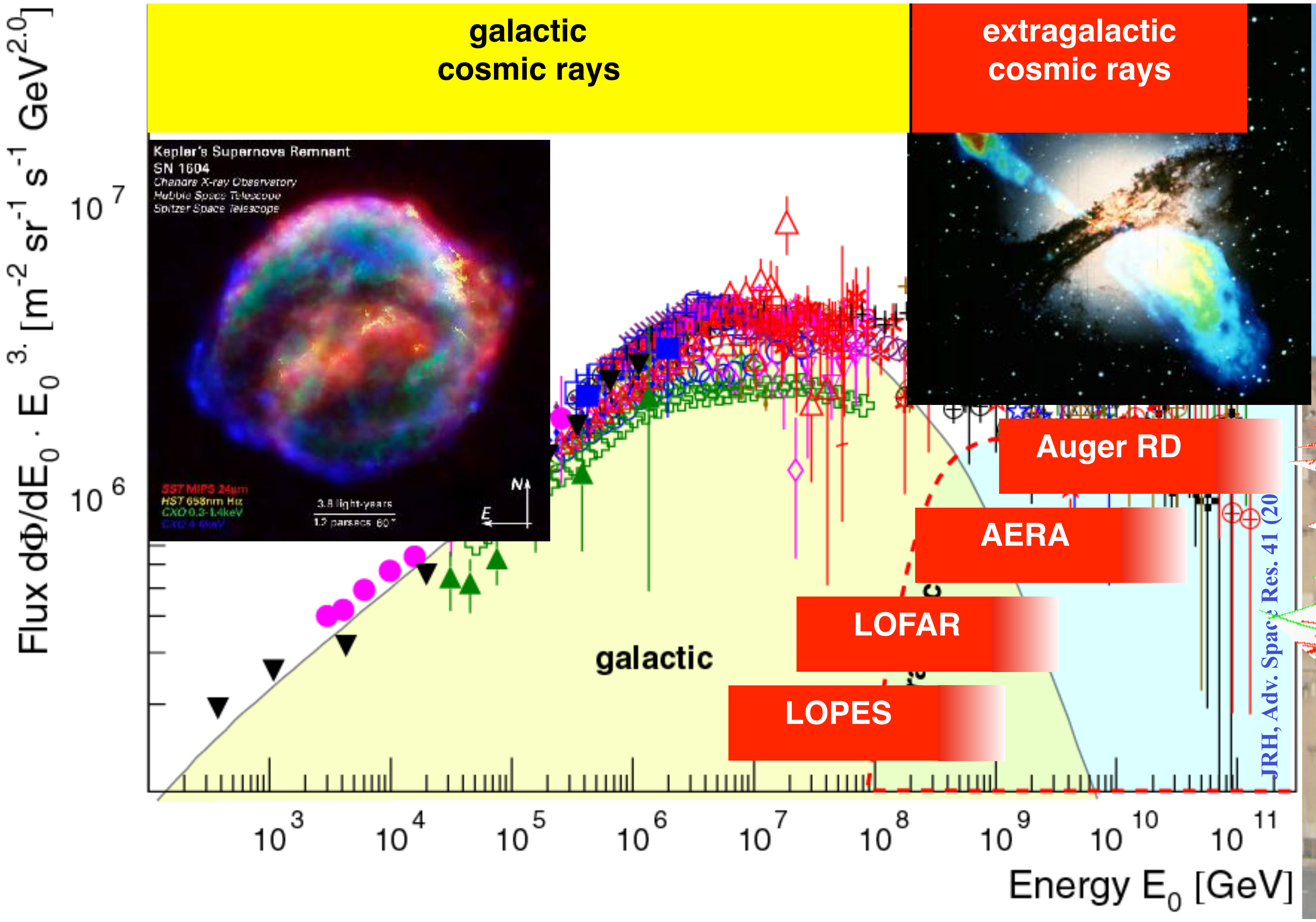

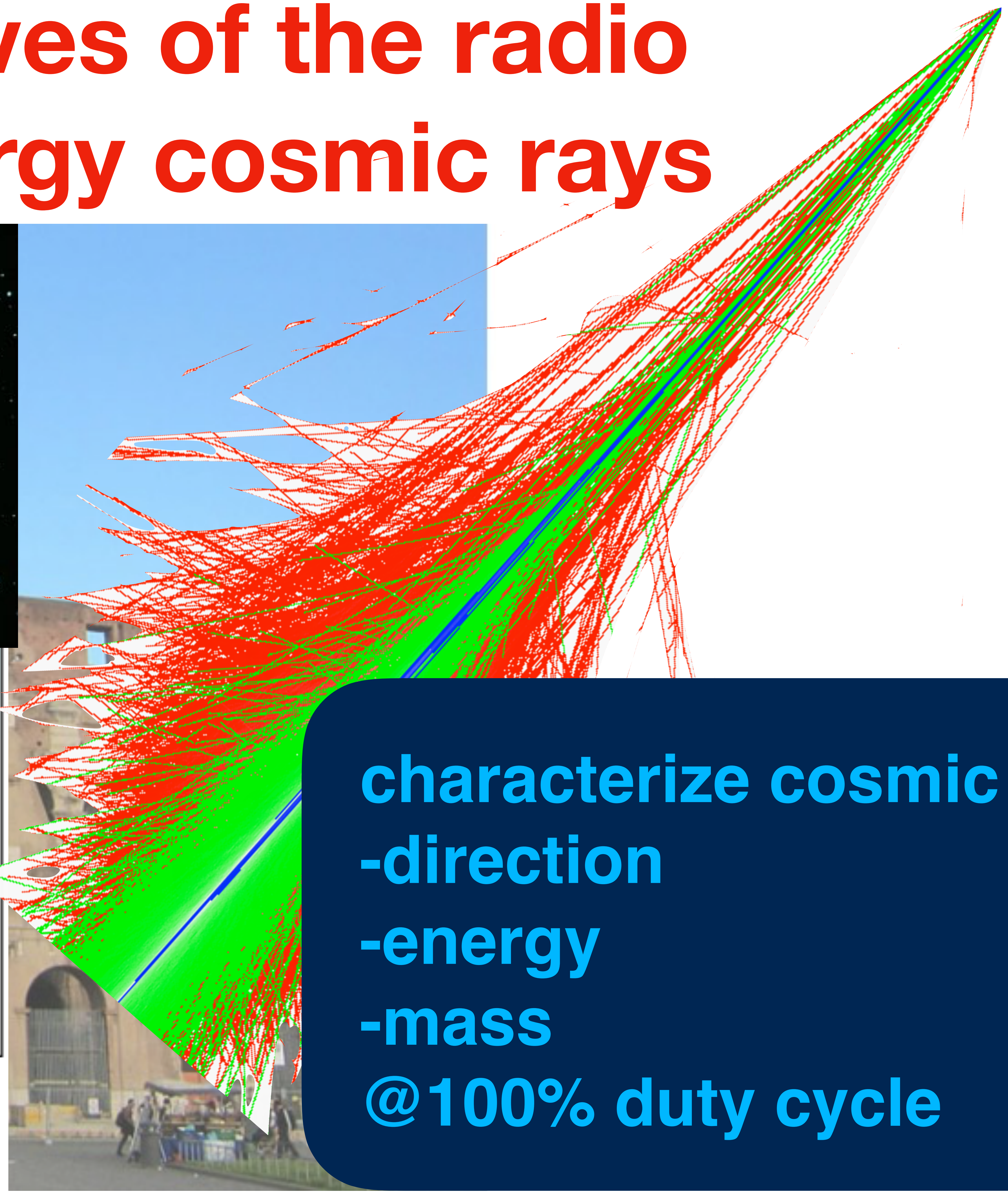

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Radio Emission in Air Showers

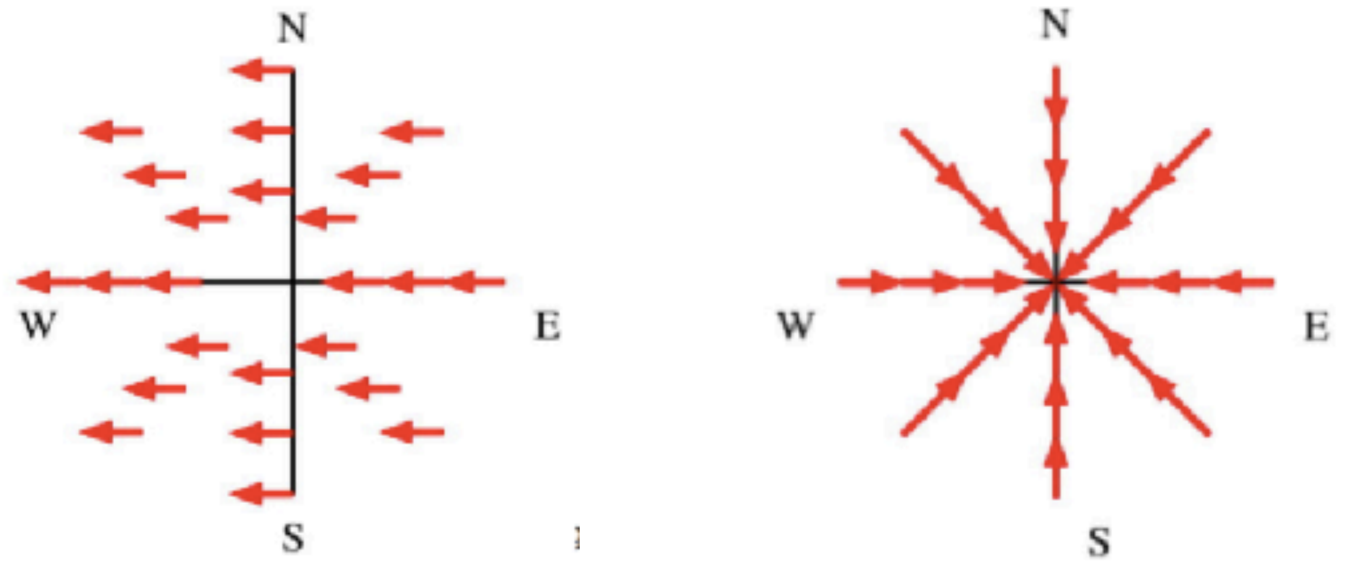
Mainly: Charge separation in **geomagnetic field**

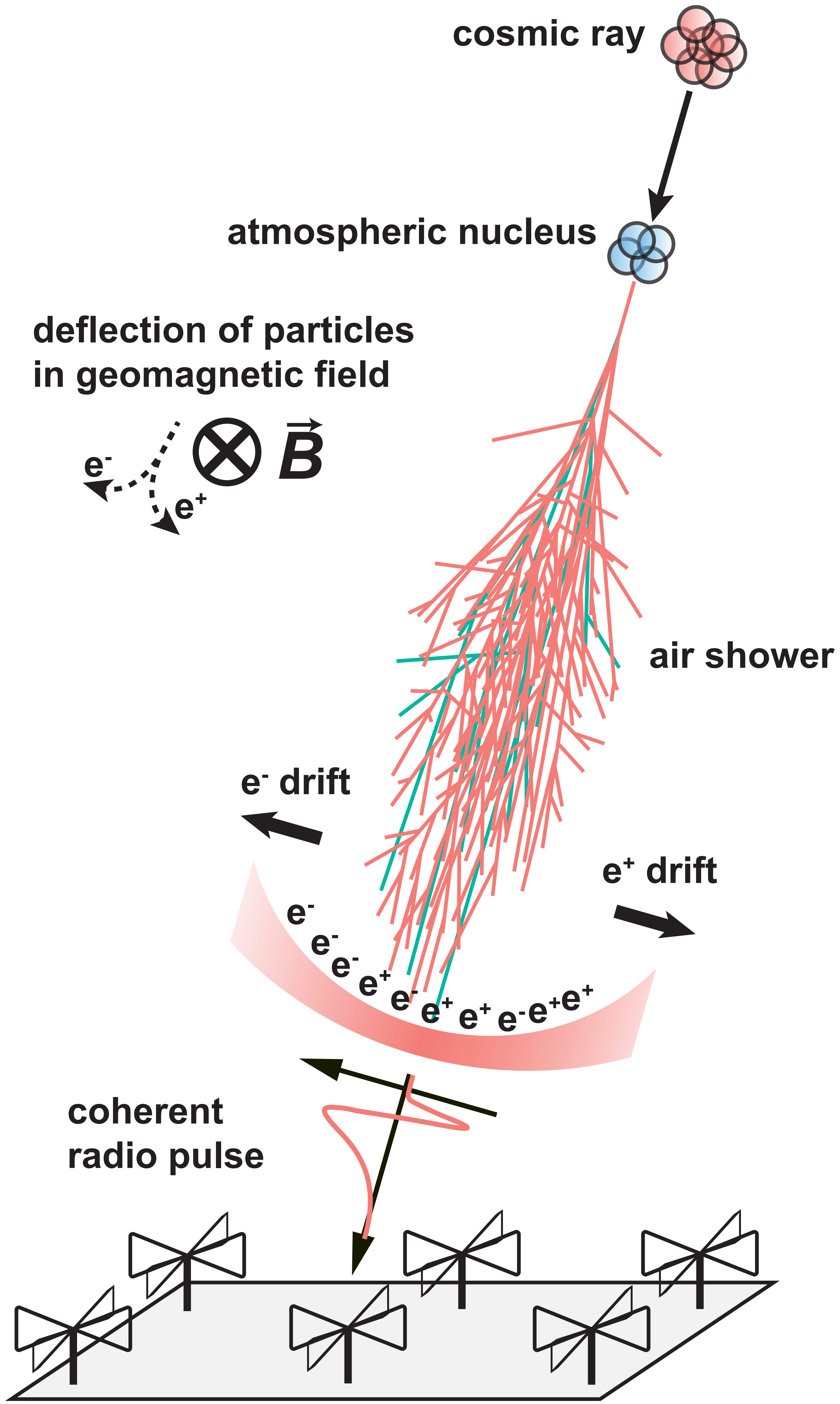
$$\vec{E} \propto \vec{v} \times \vec{B}$$

Theory predicts additional mechanisms:

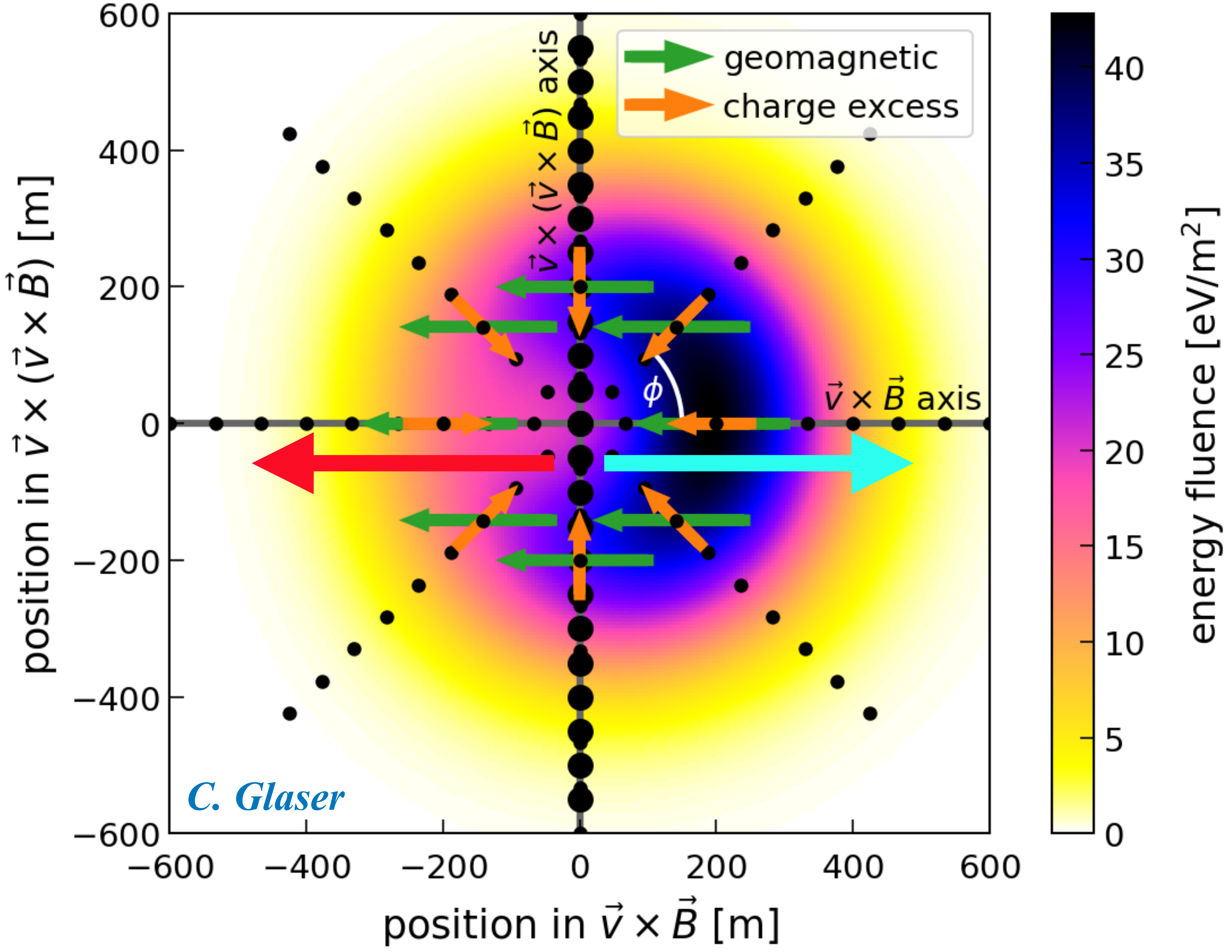
- excess of electrons in shower: **charge excess**
- superposition of emission due to **Cherenkov** effects in atmosphere

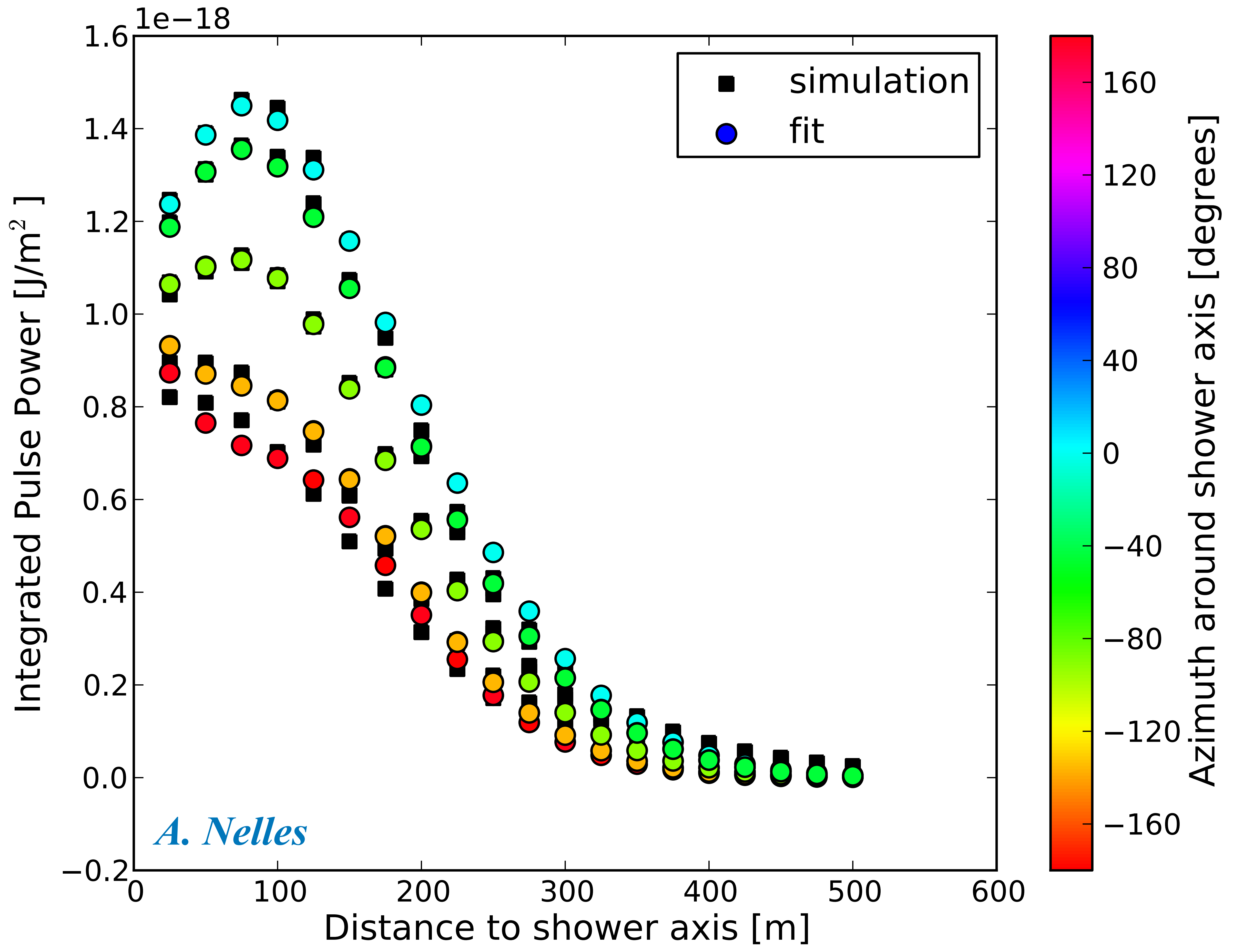
polarization of radio signal





Footprint of radio emission on the ground





33RD INTERNATIONAL COSMIC RAY CONFERENCE, RIO DE JANEIRO 2013
THE ASTROPARTICLE PHYSICS CONFERENCE

ICRC 2013

The renaissance of radio detection of cosmic rays

TIM HUEGE[1]

2018: beyond capabilities of standard installations

2016: radio technique mature: properties of cosmic rays

2014: understanding the emission processes

2005: understanding the signal

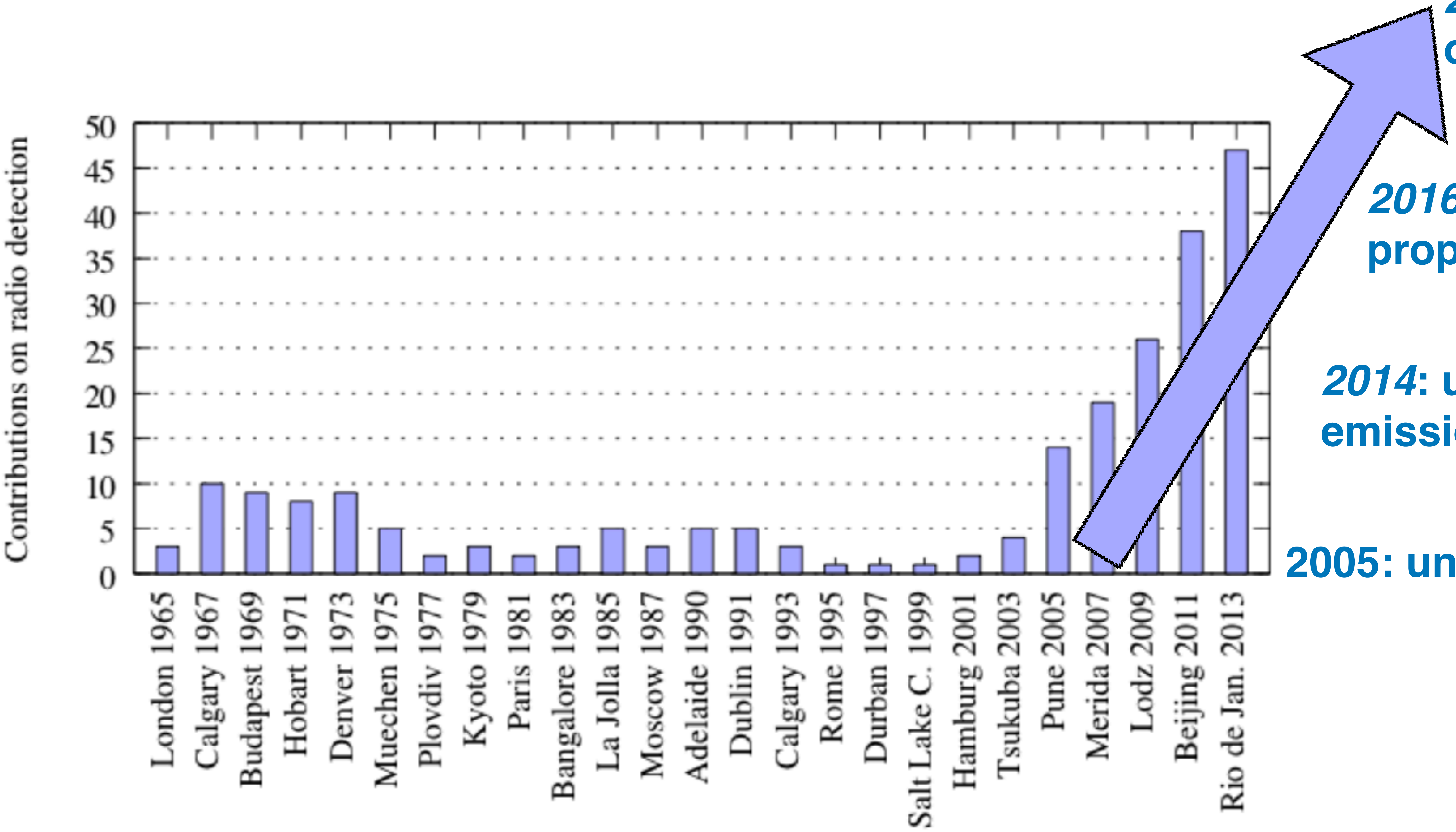


Figure 1: Number of contributions related to radio detection of cosmic rays or neutrinos to the ICRCs since 1965. The field has grown very impressively since the modern activities started around 2003. Data up to 2007 were taken from [11].

Radio detection of extensive air showers around the world

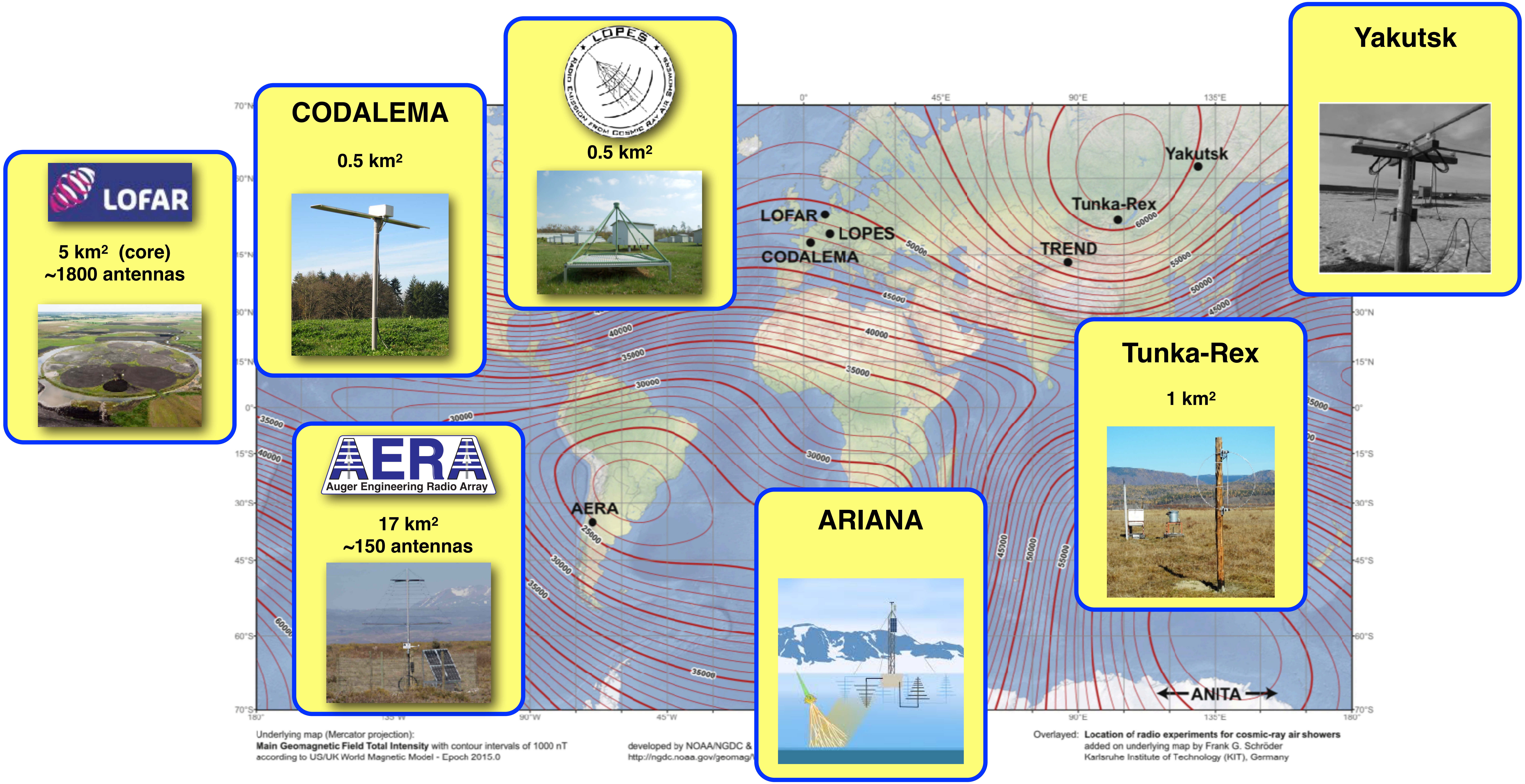
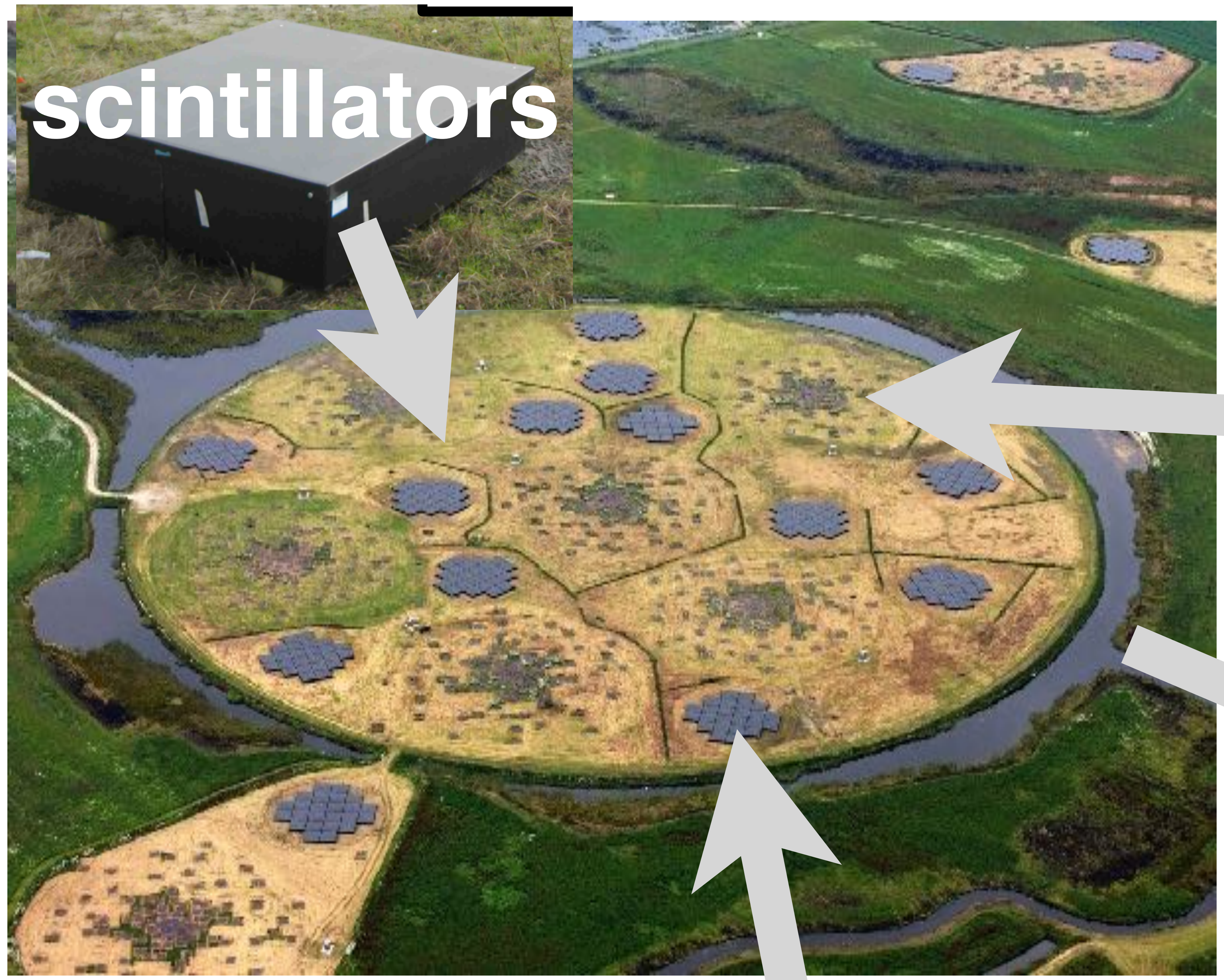
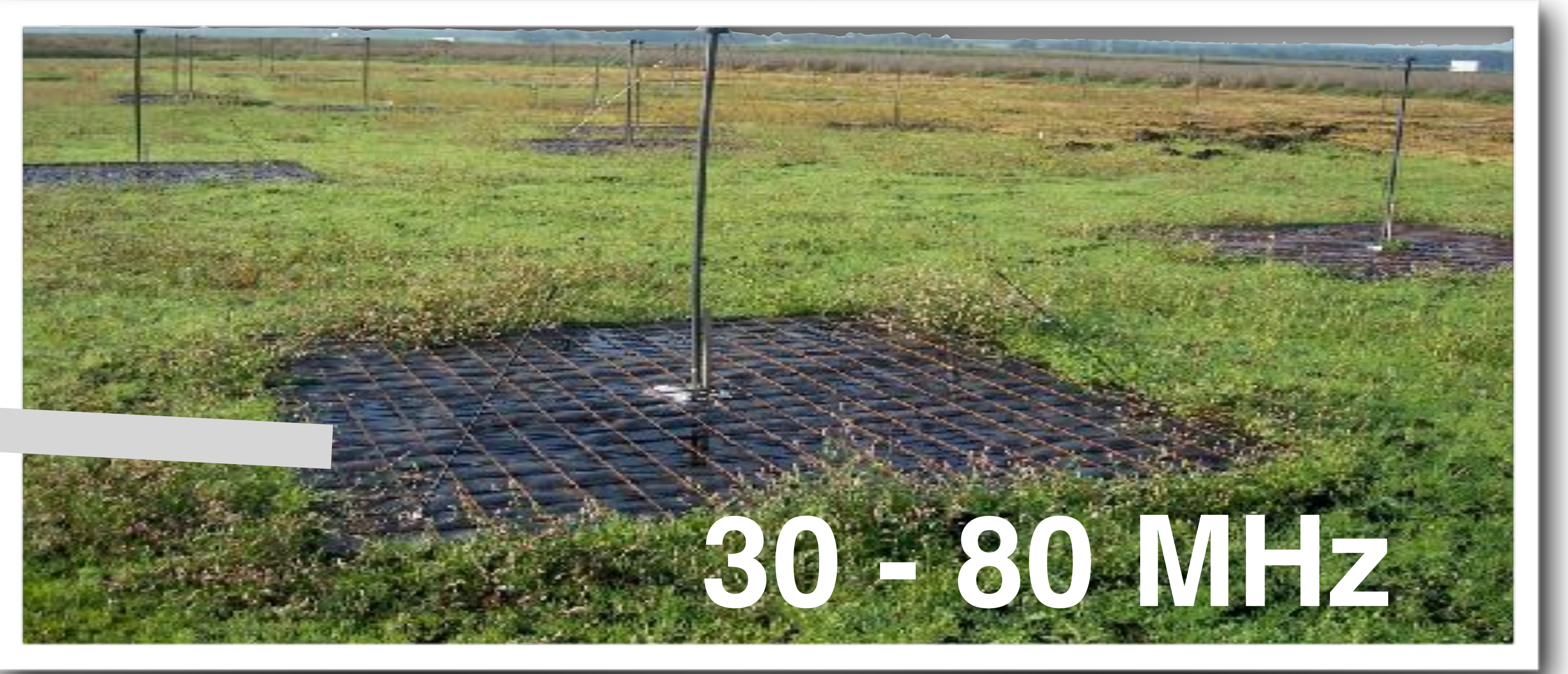
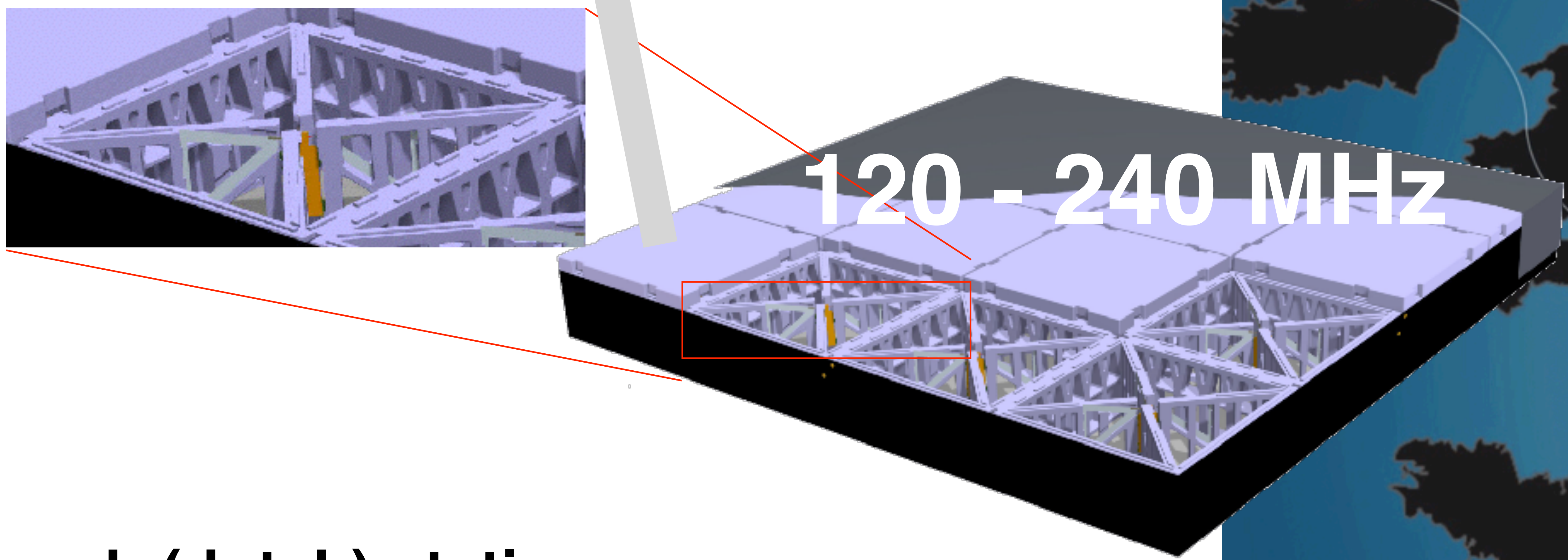


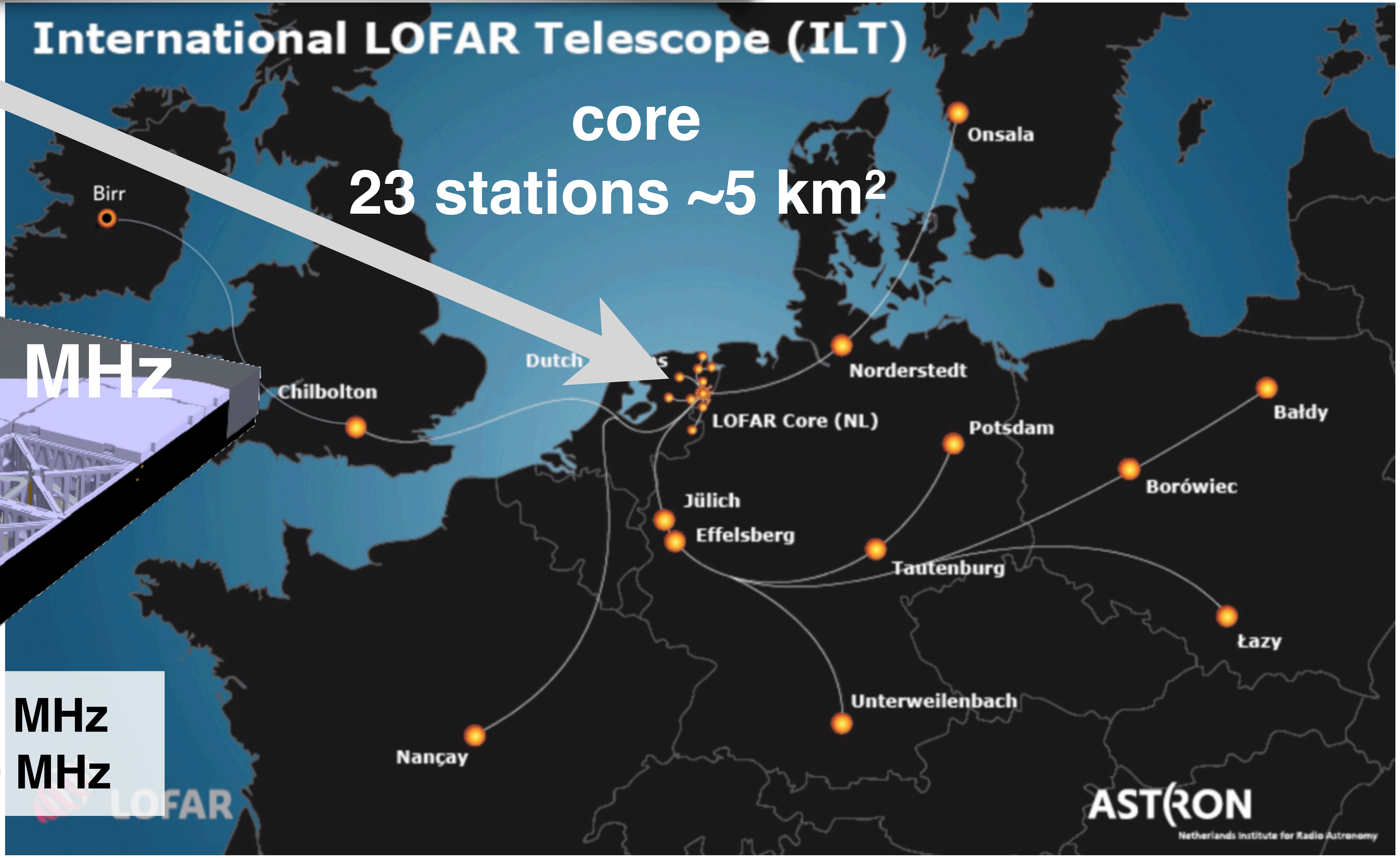
Fig. 21. Map of the total geomagnetic field strengths (world magnetic model [207]) and the location of various radio experiments detecting cosmic-ray air showers.

F.G. Schröder, Prog. Part. Nucl. Phys. 92 (17) 1









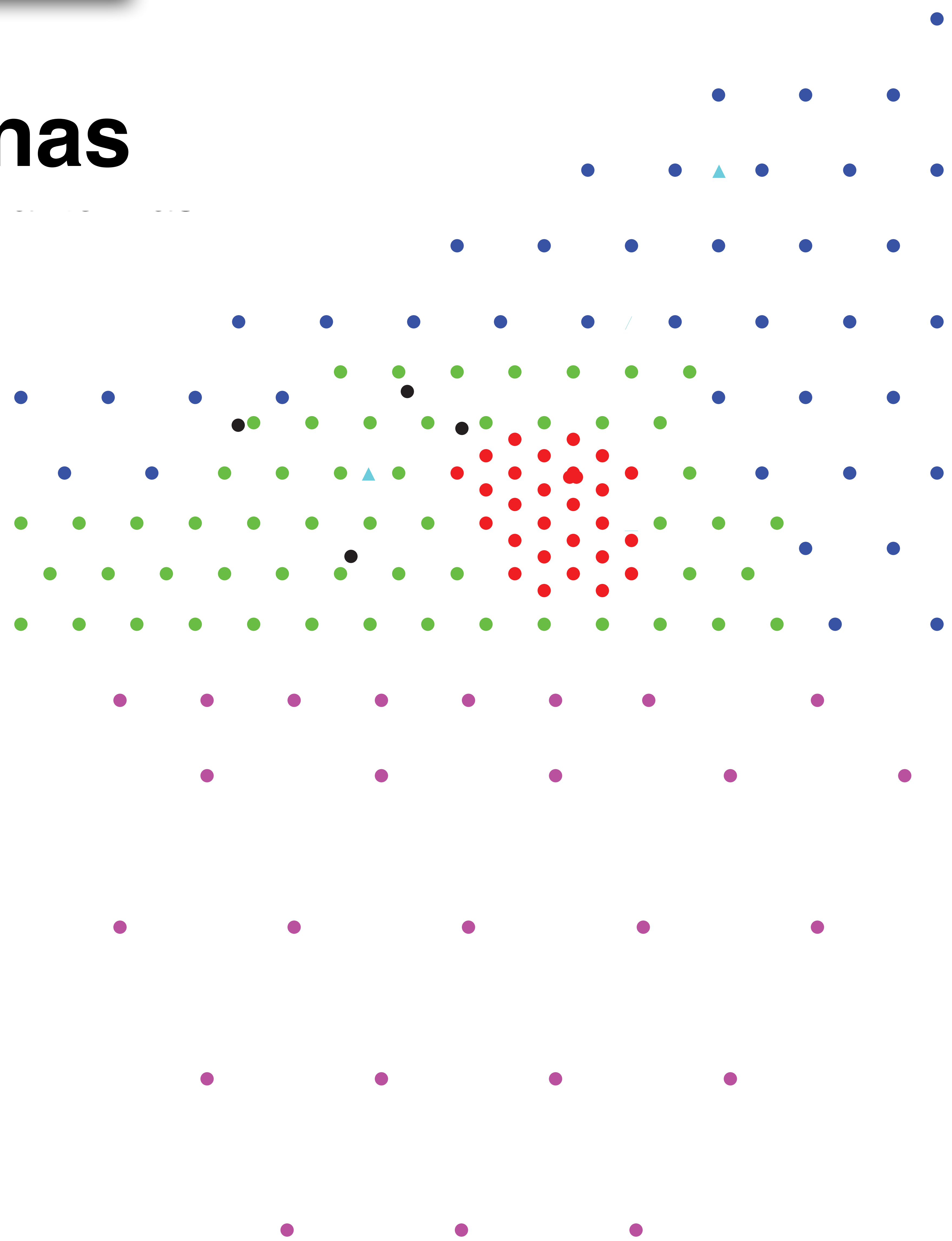
each (dutch) station:
- 96 low-band antennas 30- 80 MHz
- high-band antennas (2x24 tiles) 120-240 MHz

M. van Haarlem et al., A&A 556 (2013) A2

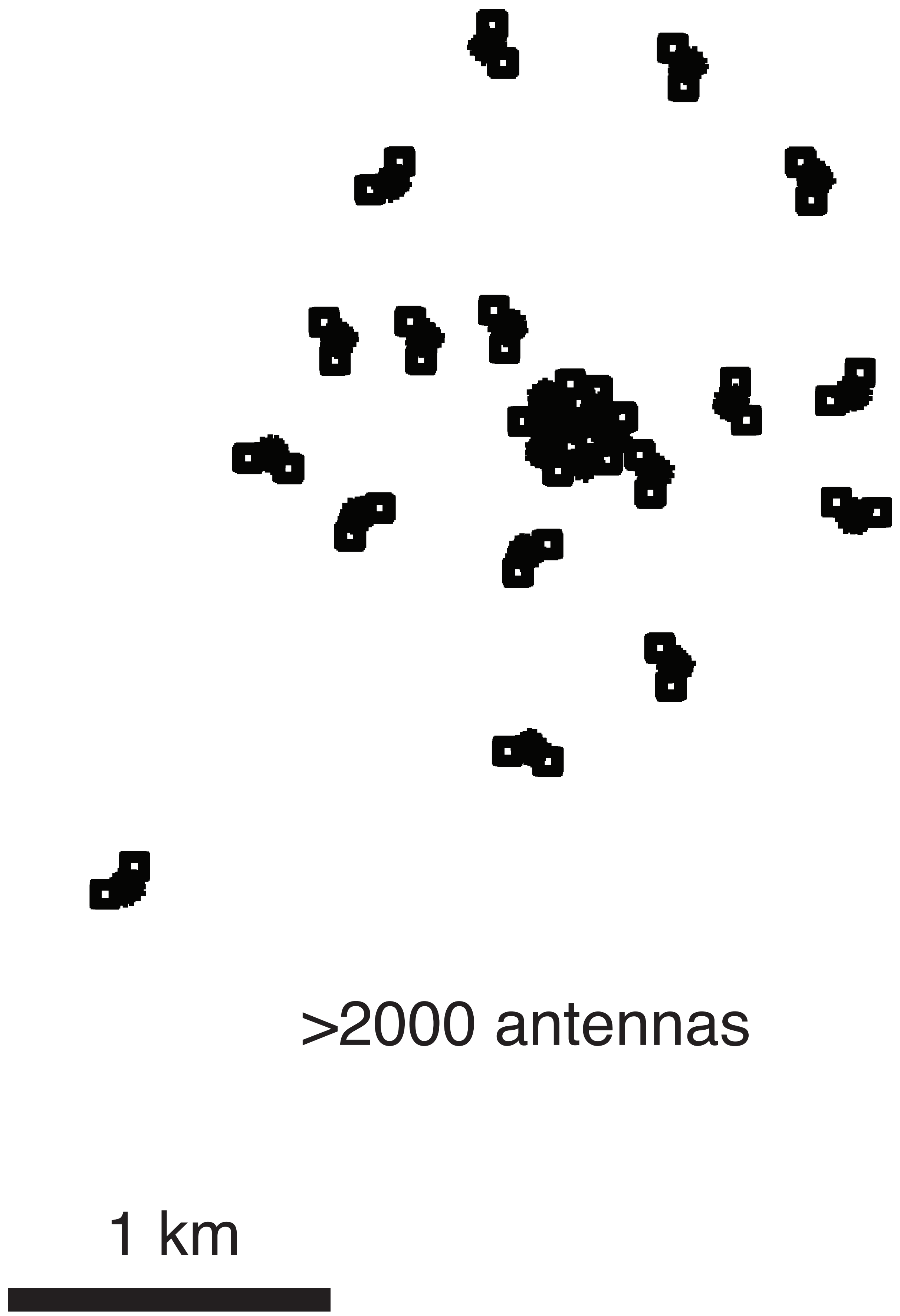
S. Thoudam et al., Nucl. Instr. Meth. A 767 (2014) 339

AERA Auger Engineering Radio Array

~150 antennas
~17 km²
30-80 MHz



LOFAR core
23 stations ~5 km²



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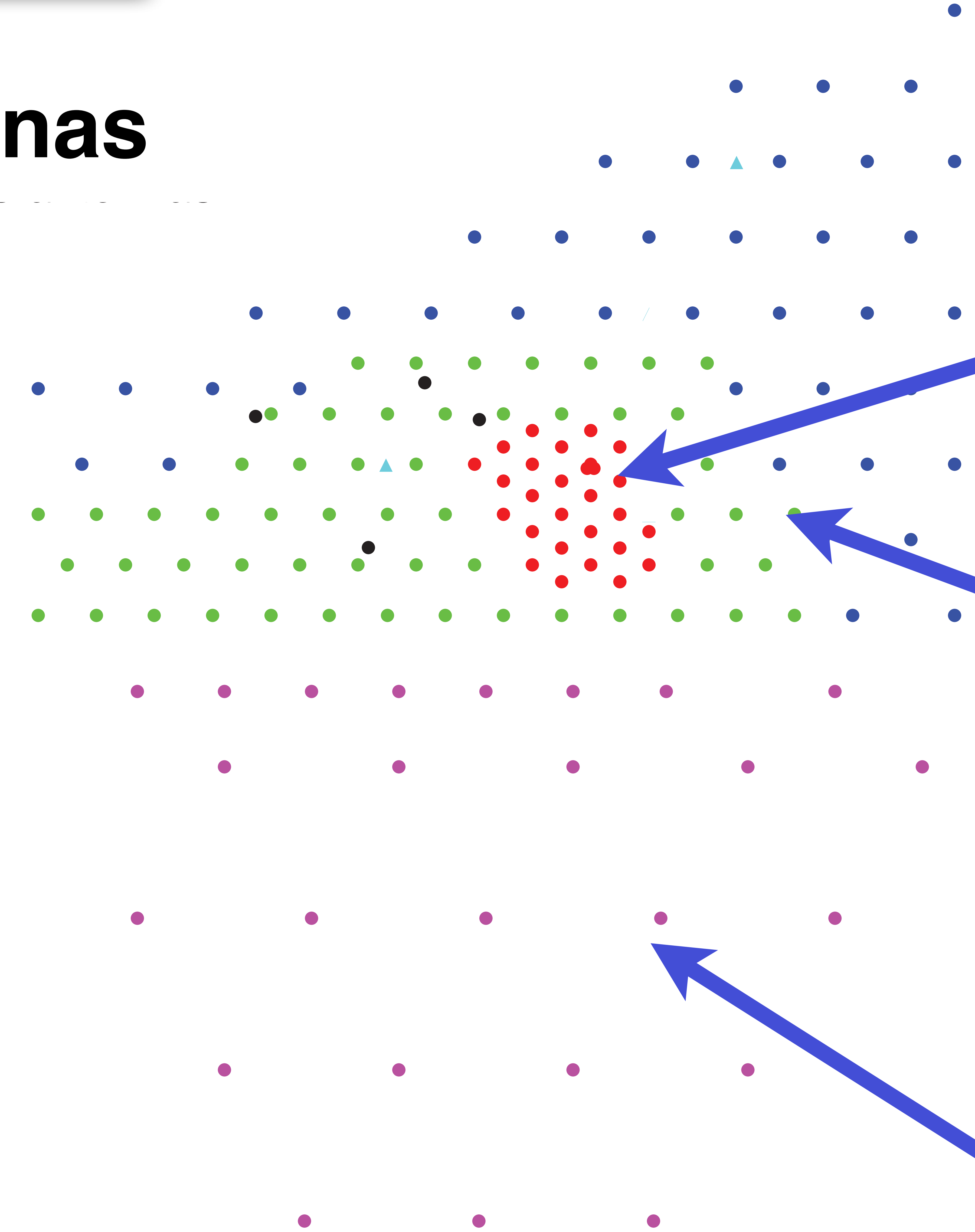

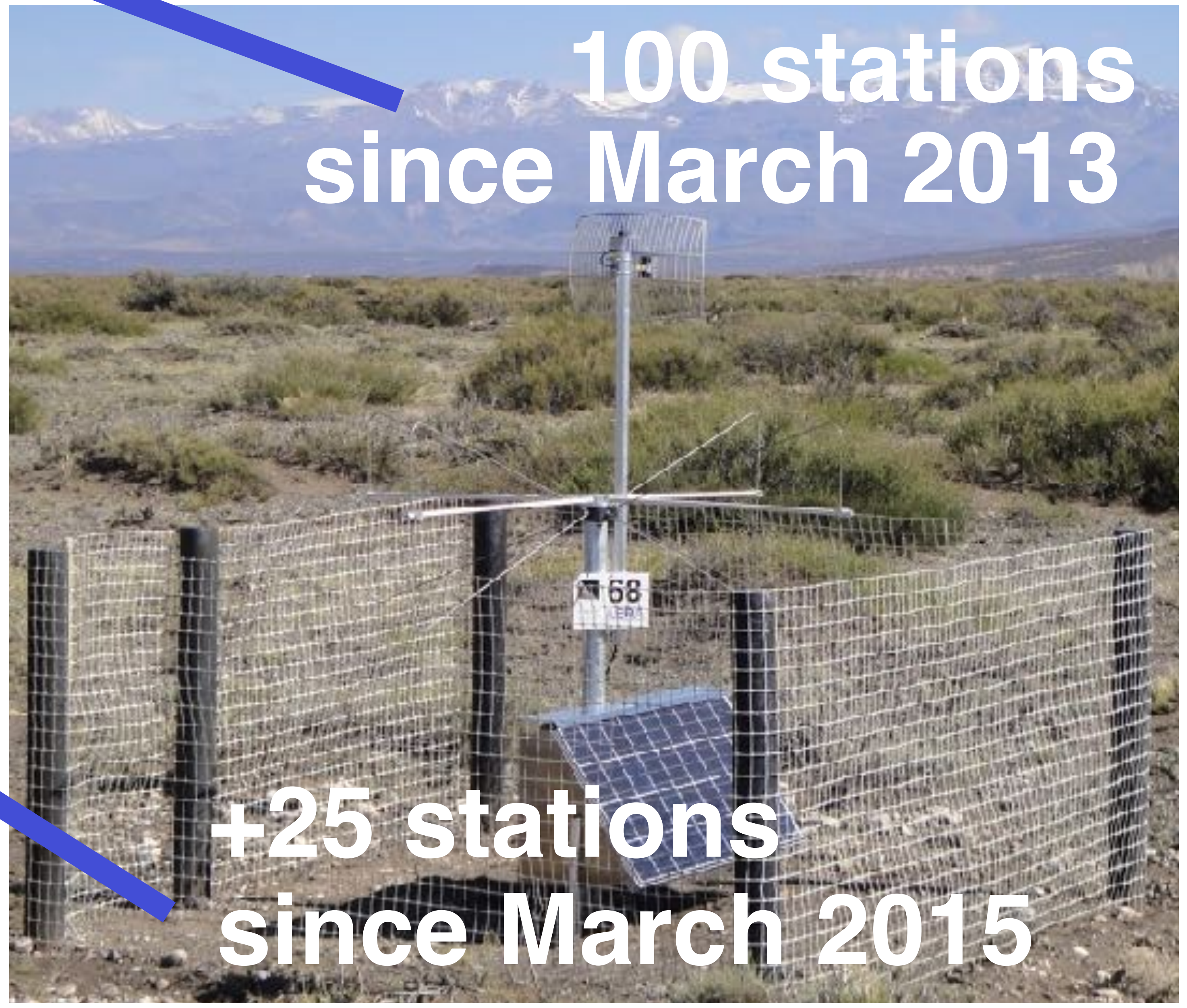

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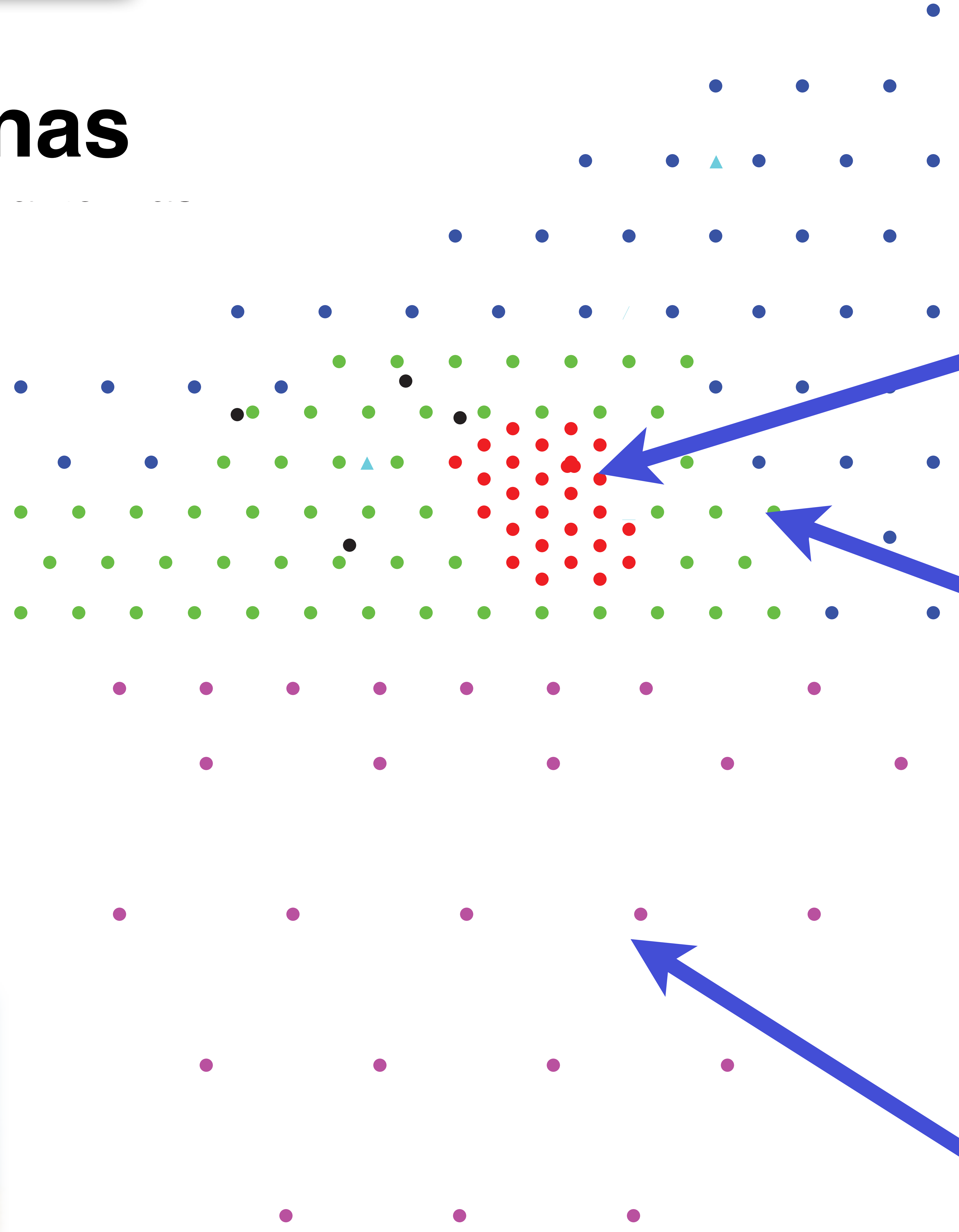
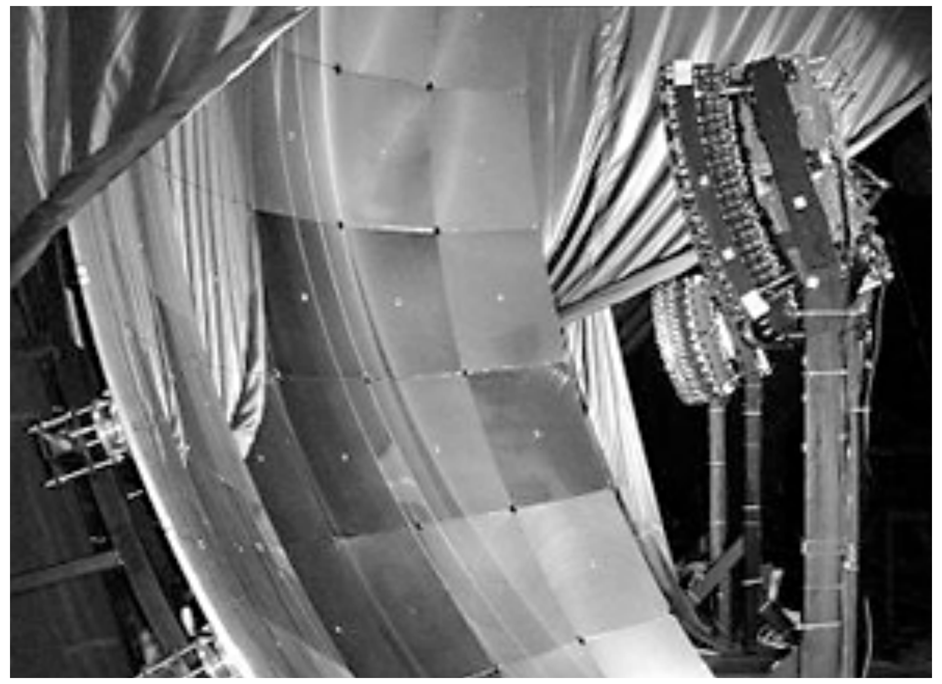
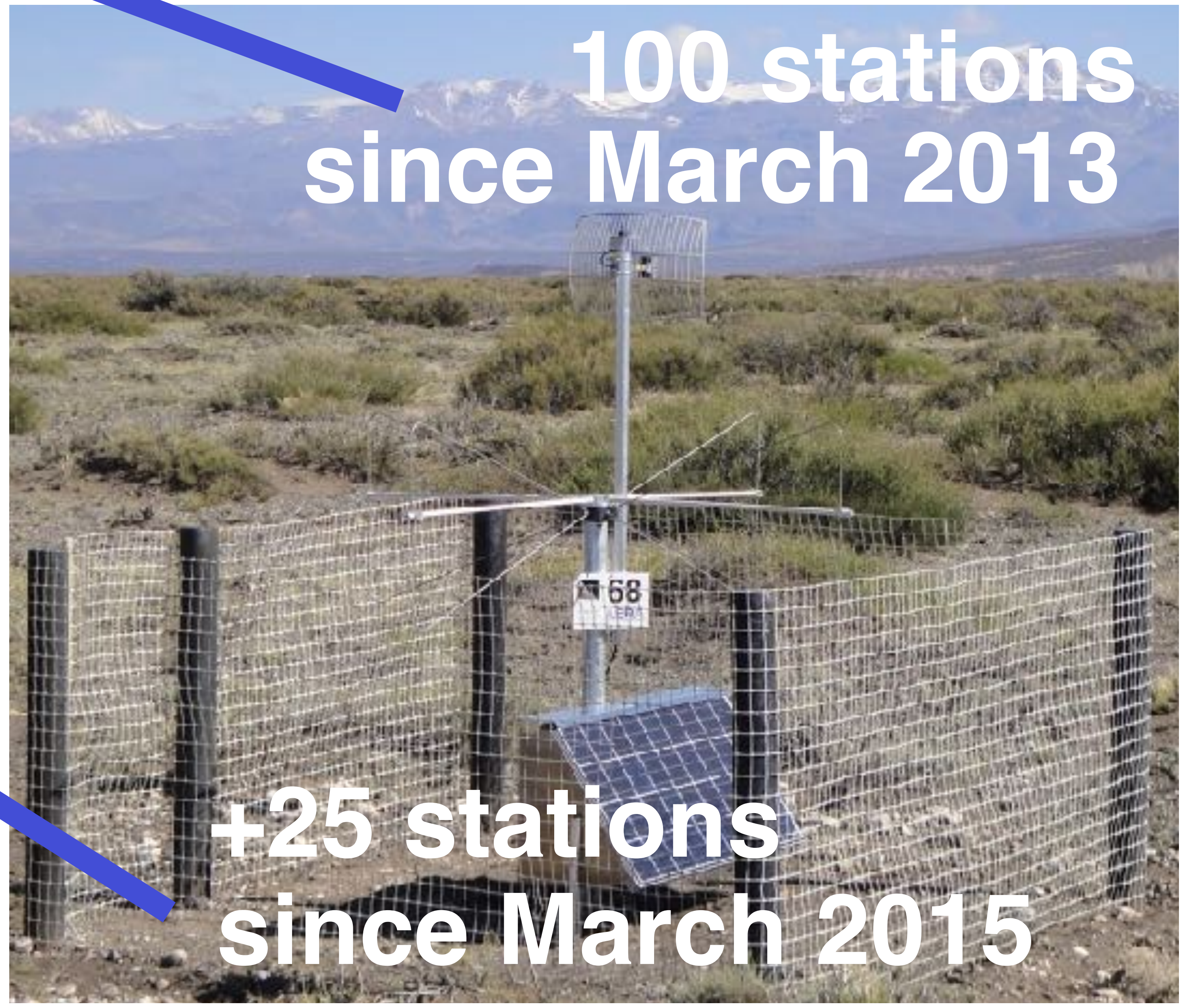




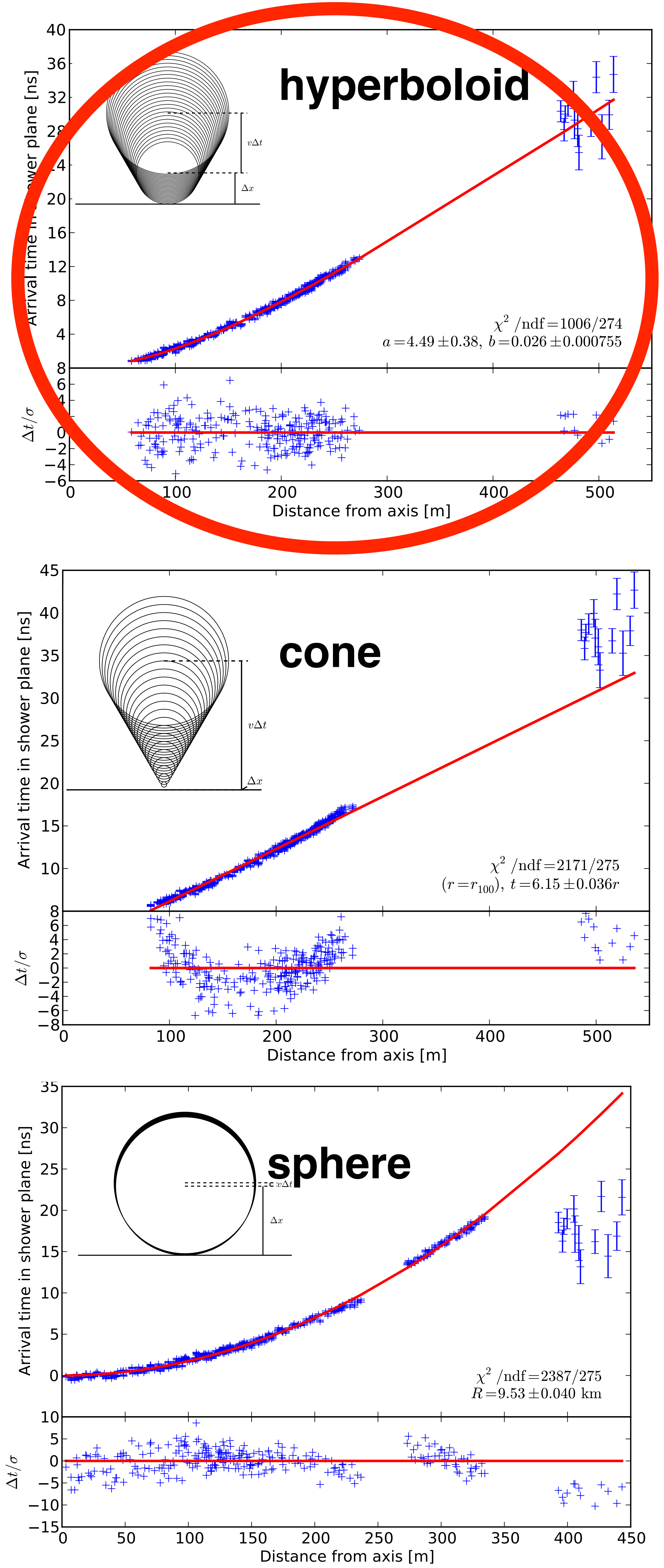


Properties of incoming cosmic ray

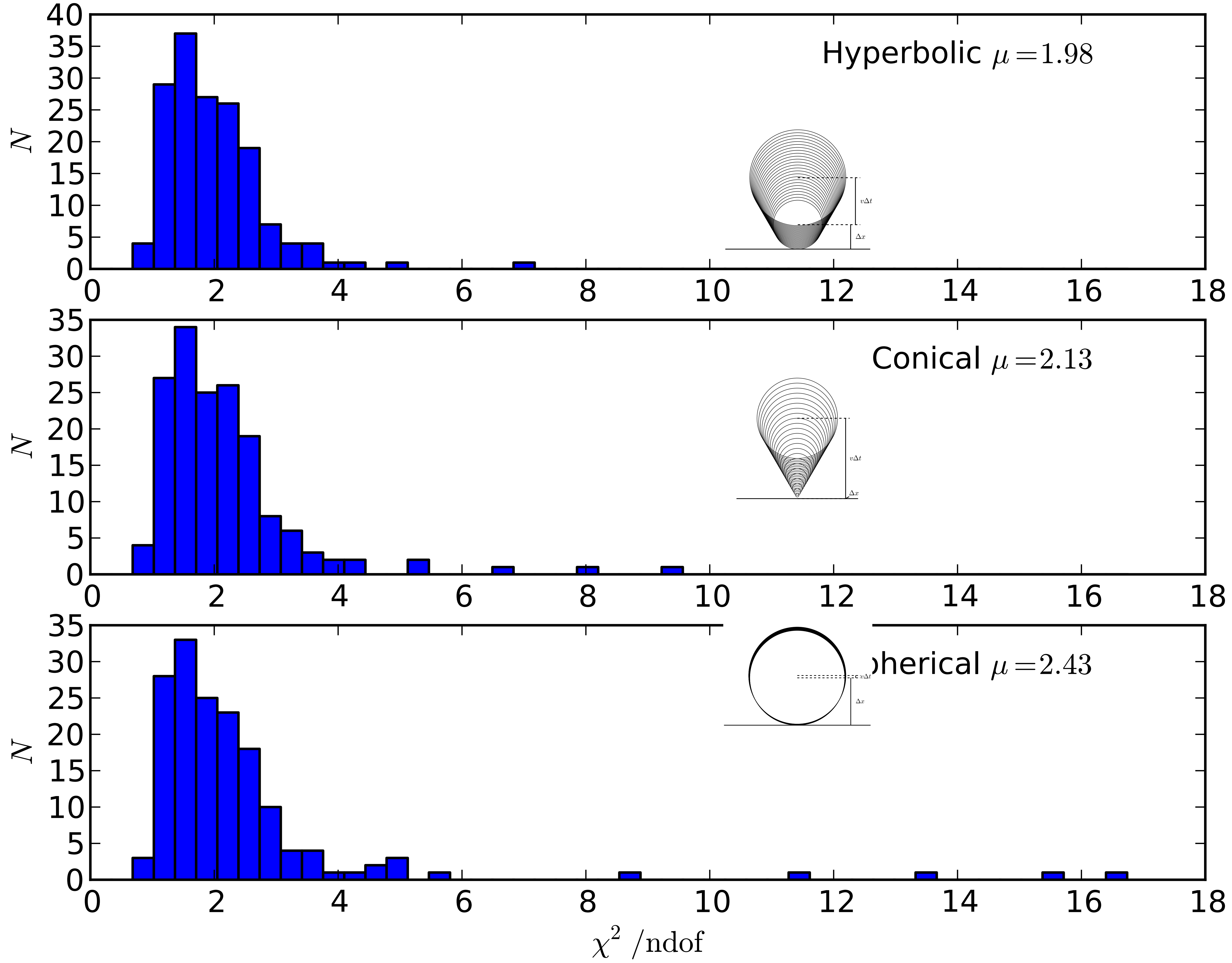
- direction
- energy
- type

Direction

Shape of Shower Front

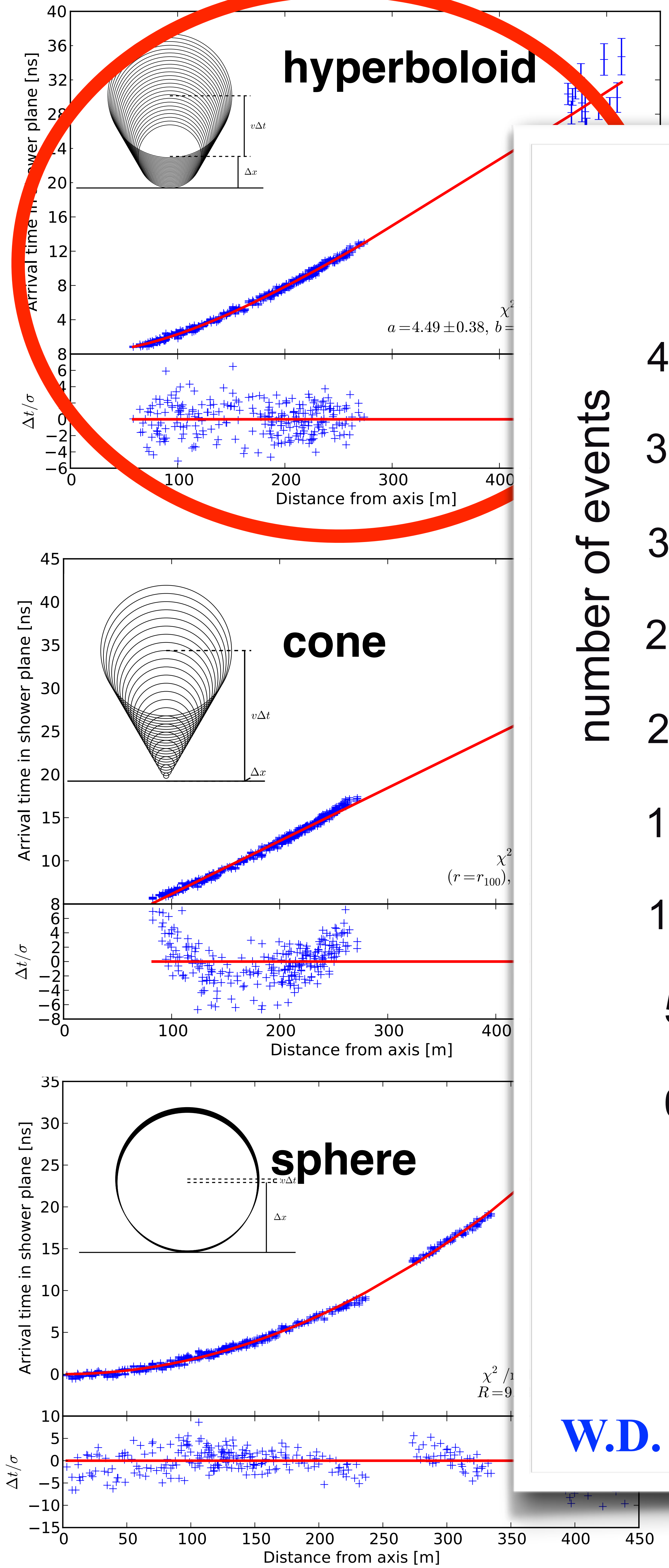


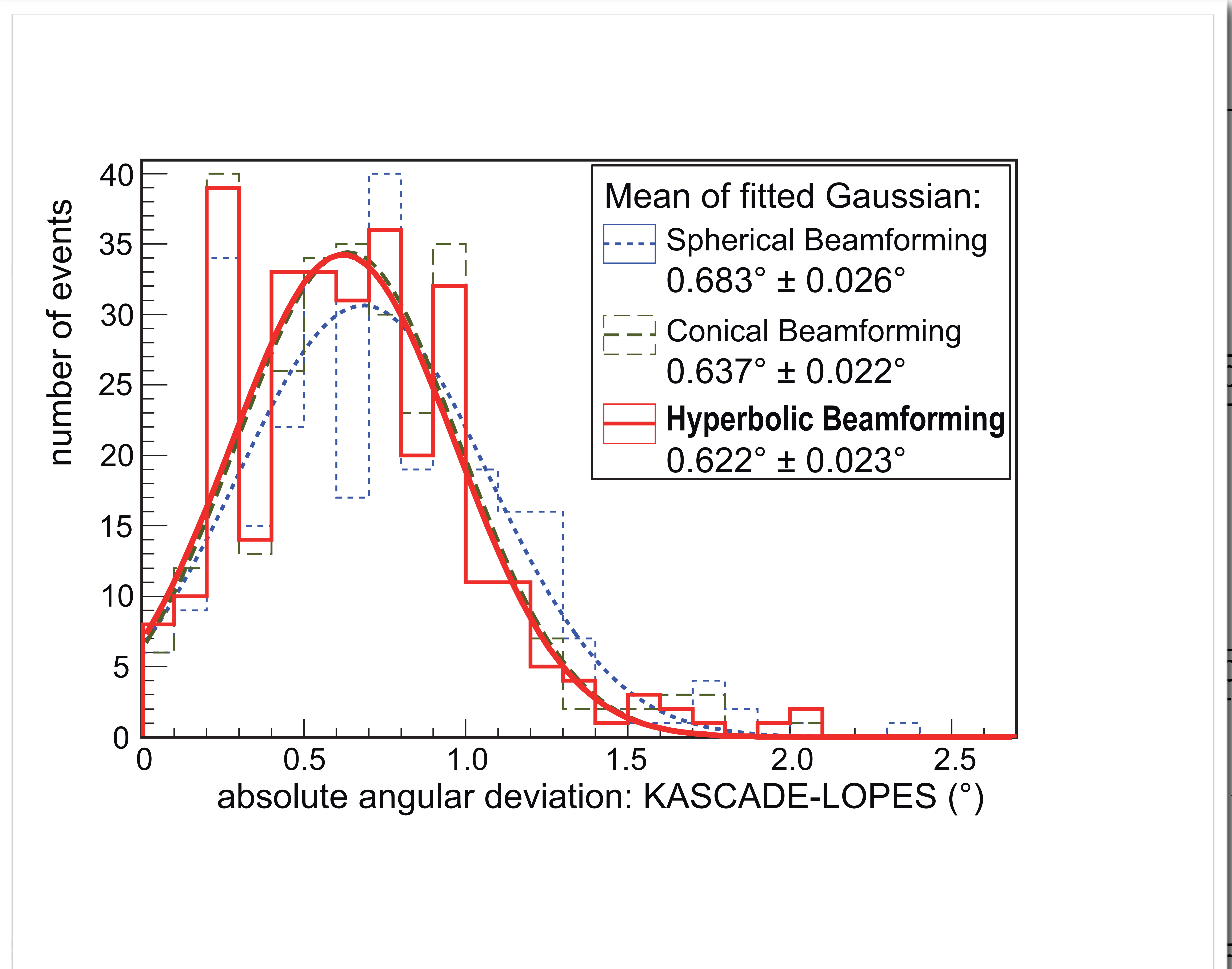
fit quality



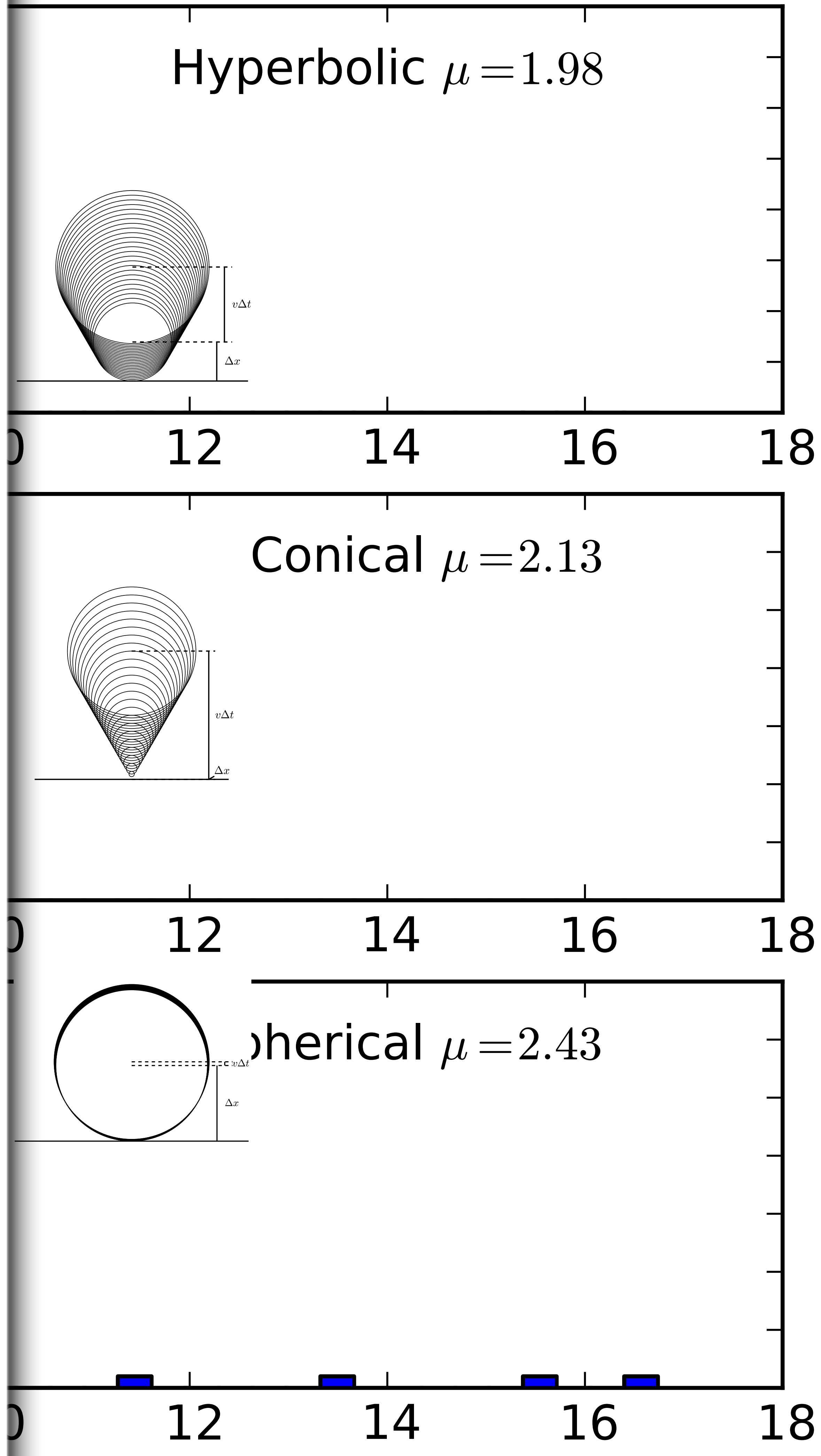
A. Corstanje et al., Astropart. Phys. 61 (2015) 22

Shape of Shower Front





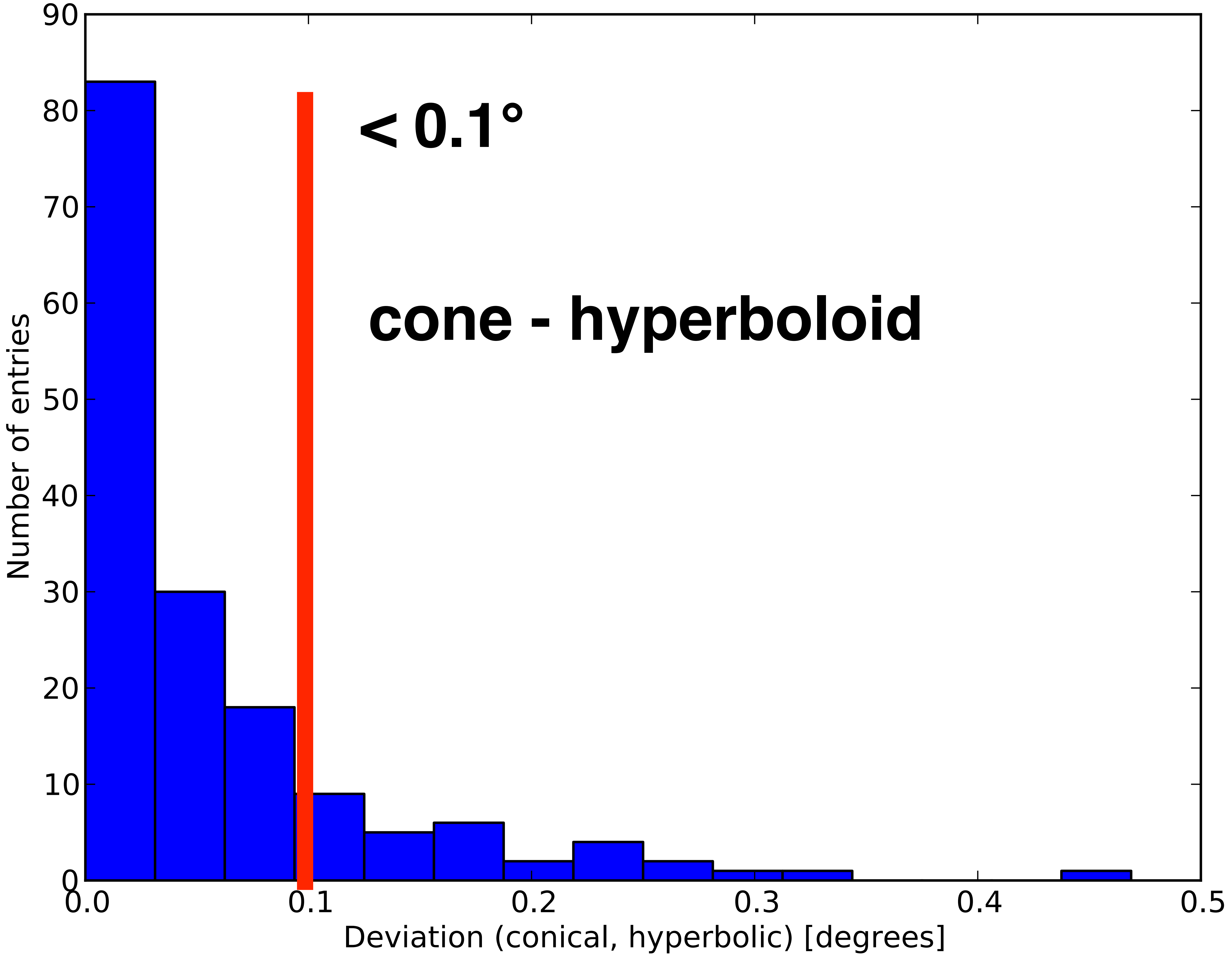
W.D. Apel et al., JCAP 1409 (2014) no.09, 025

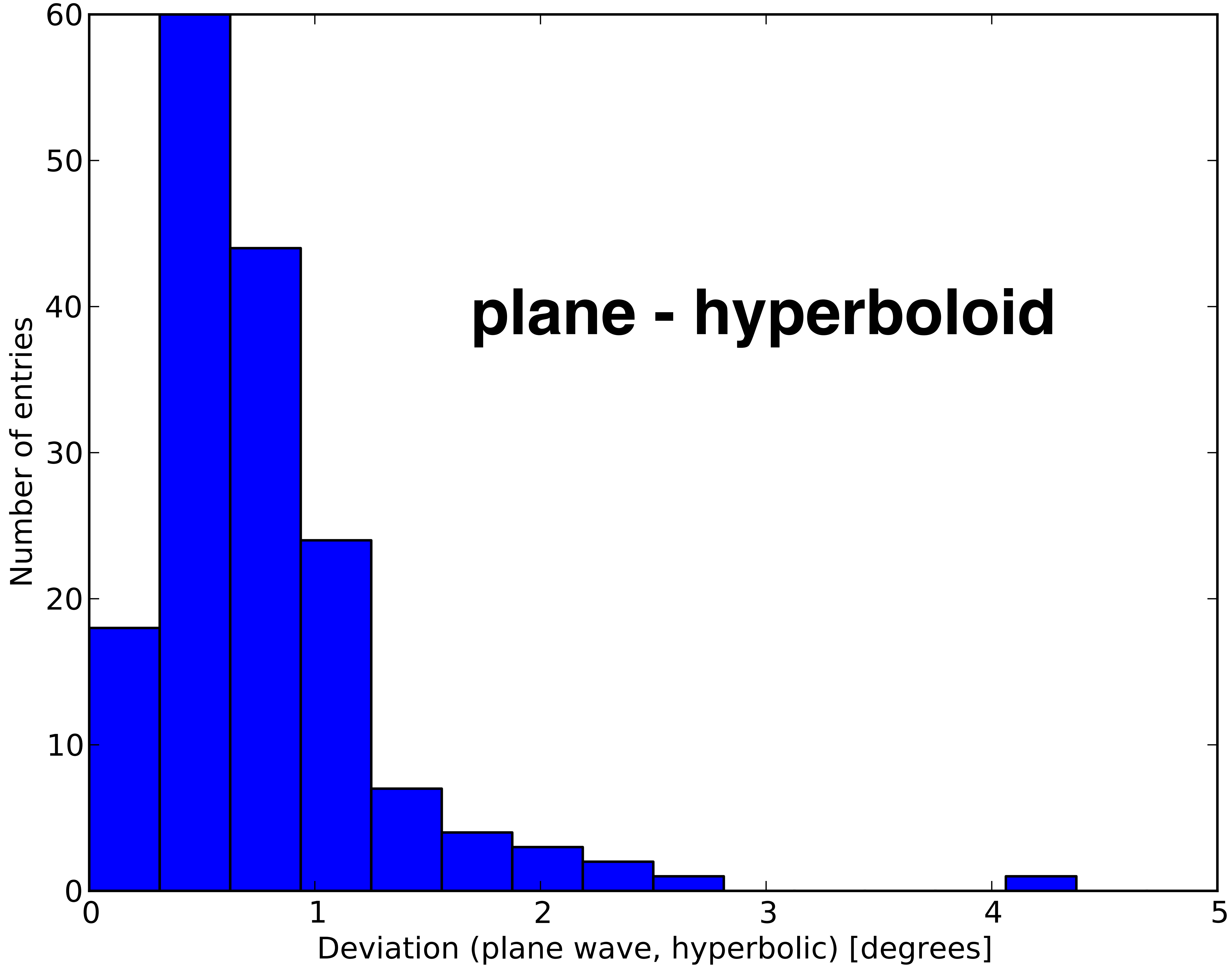


A. Corstanje et al., Astropart. Phys. 61 (2015) 22

Accuracy of Shower Direction

angular difference between..

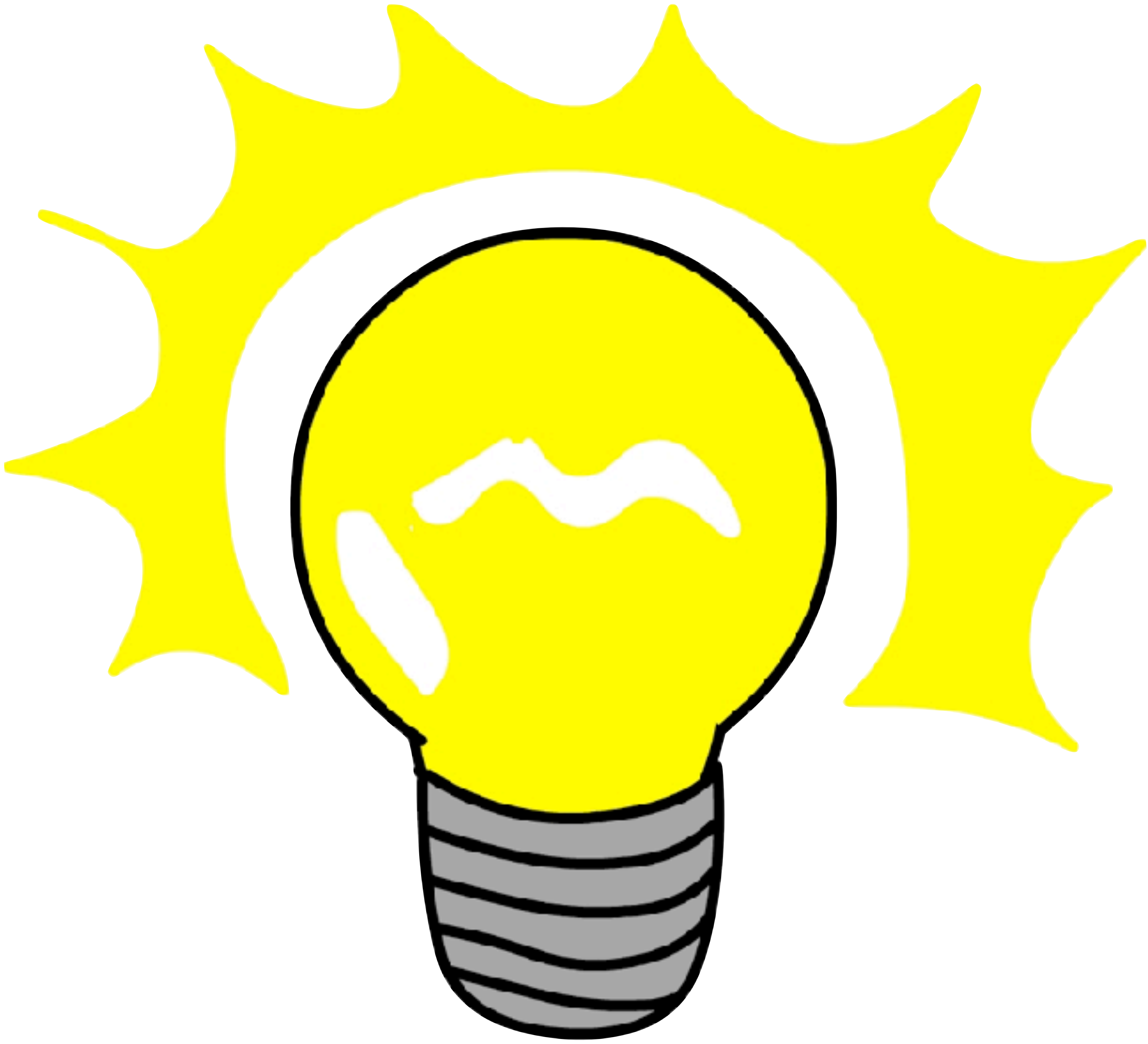




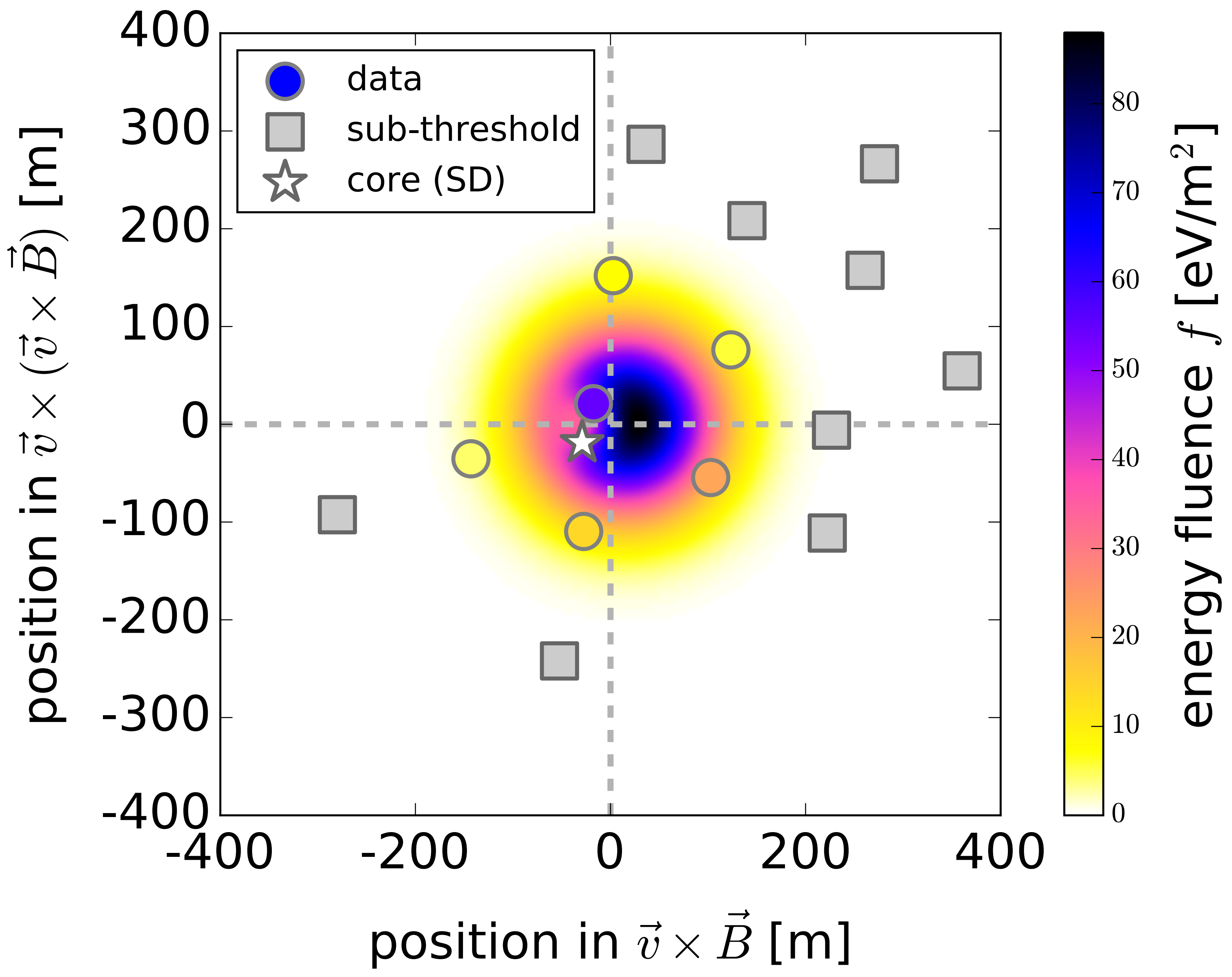
A. Corstanje et al., Astropart. Phys. 61 (2015) 22

LOFAR

Energy

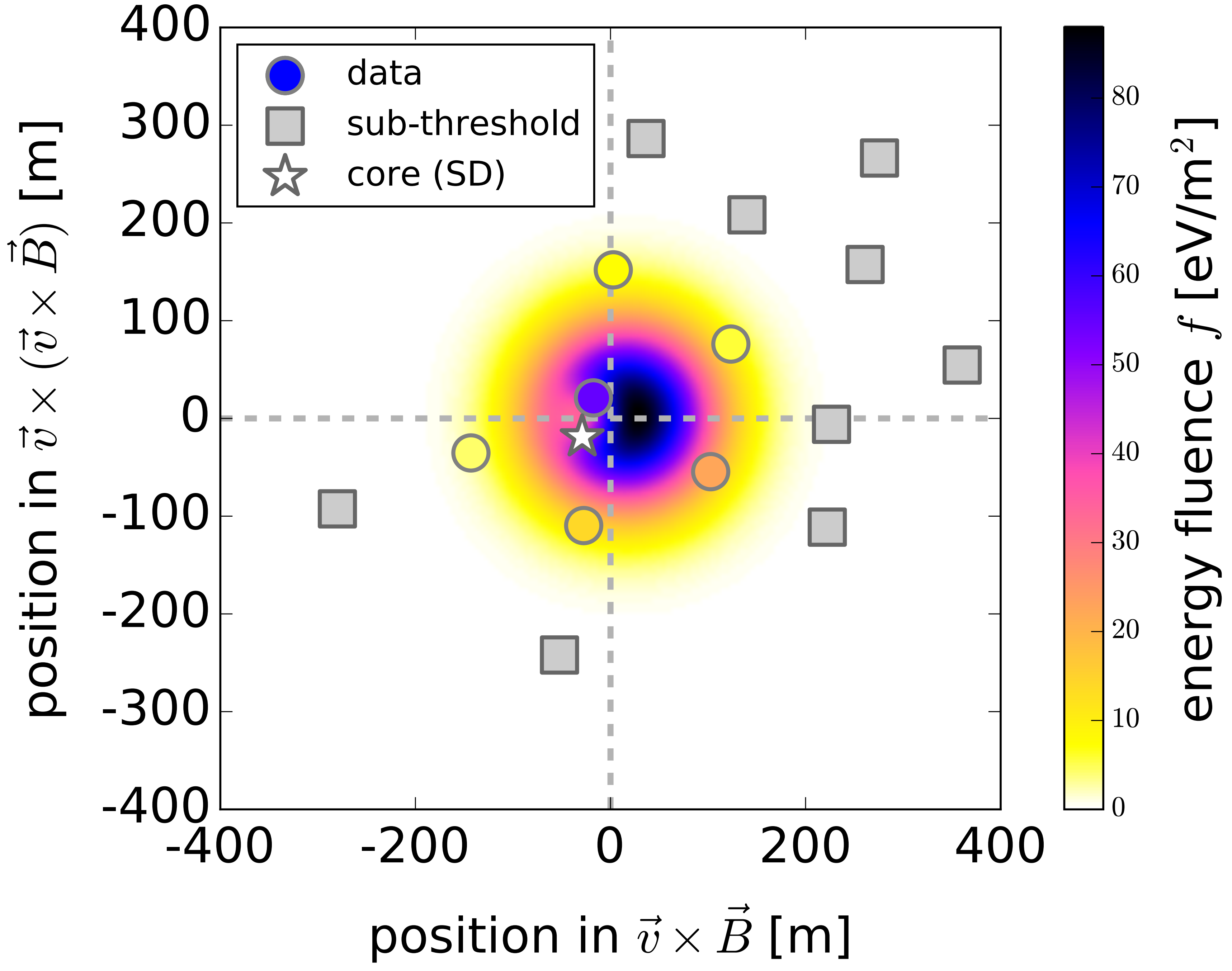


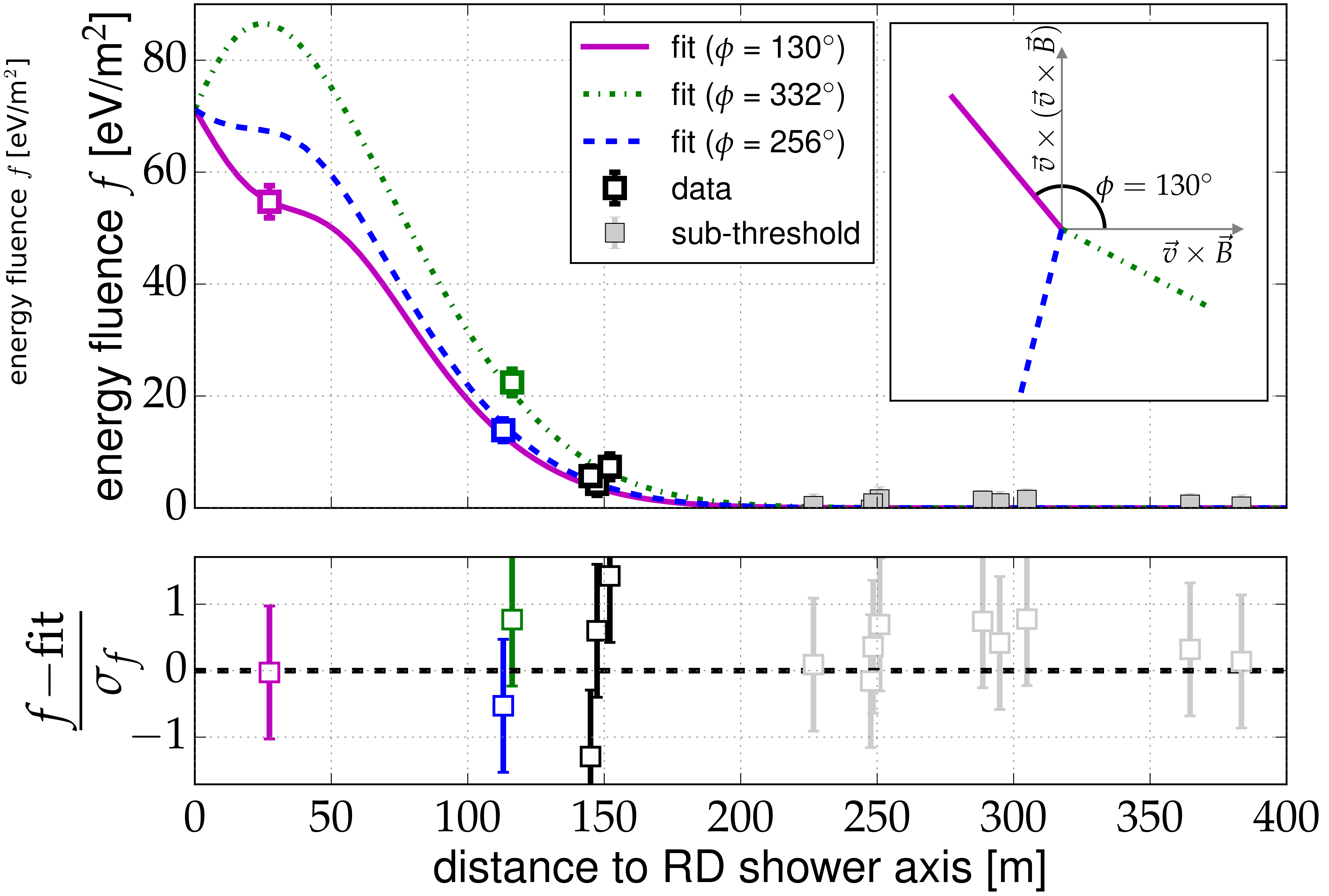
Measurement of the Radiation Energy in the Radio Signal of Extensive Air Showers as a Universal Estimator of Cosmic-Ray Energy



A. Aab et al., PRL 116 (2016) no.24, 241101

Measurement of the Radiation Energy in the Radio Signal of Extensive Air Showers as a Universal Estimator of Cosmic-Ray Energy





PIERRE AUGER OBSERVATORY

AERA Auger Engineering Radio Array

A. Aab et al., PRL 116 (2016) no.24, 241101

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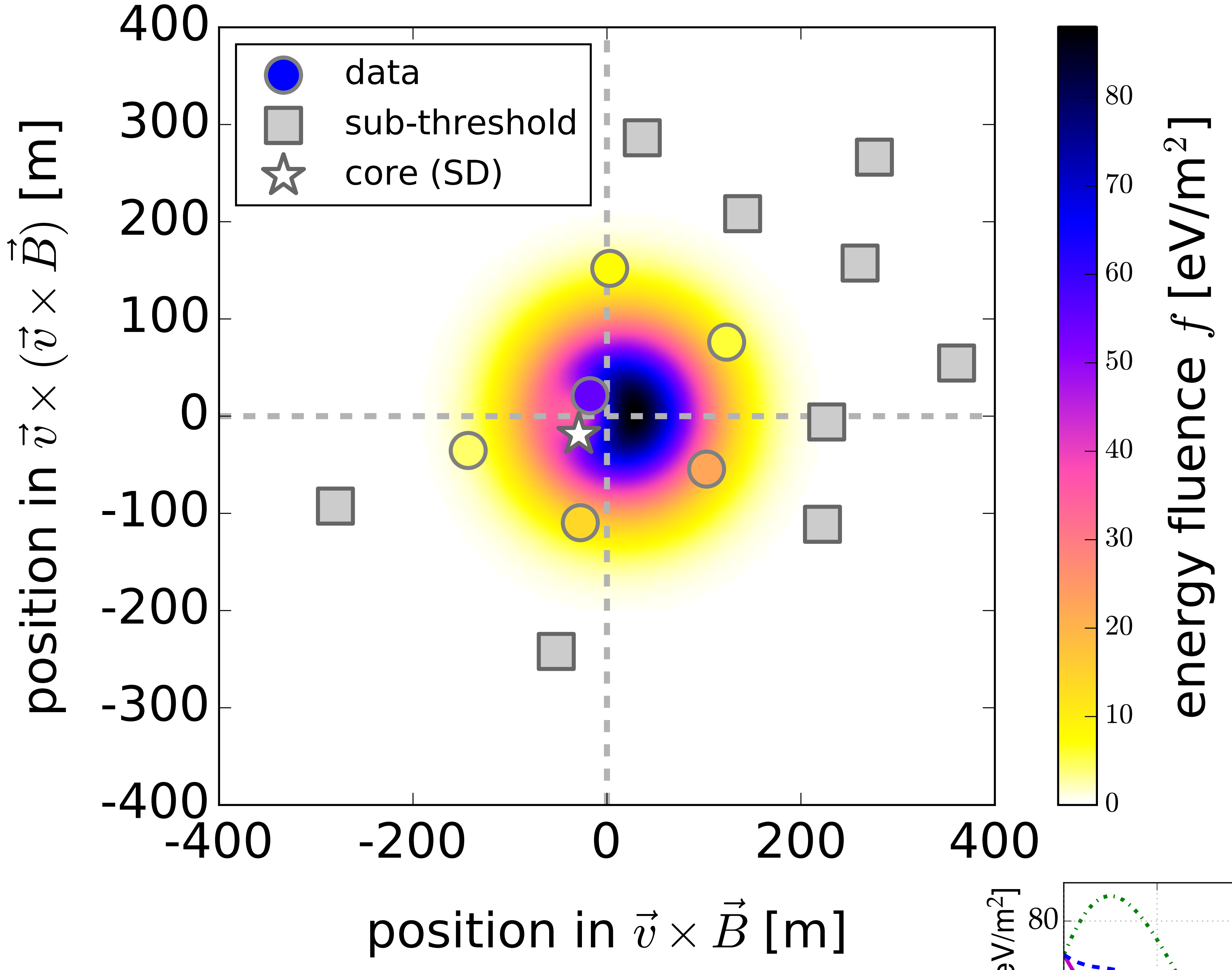
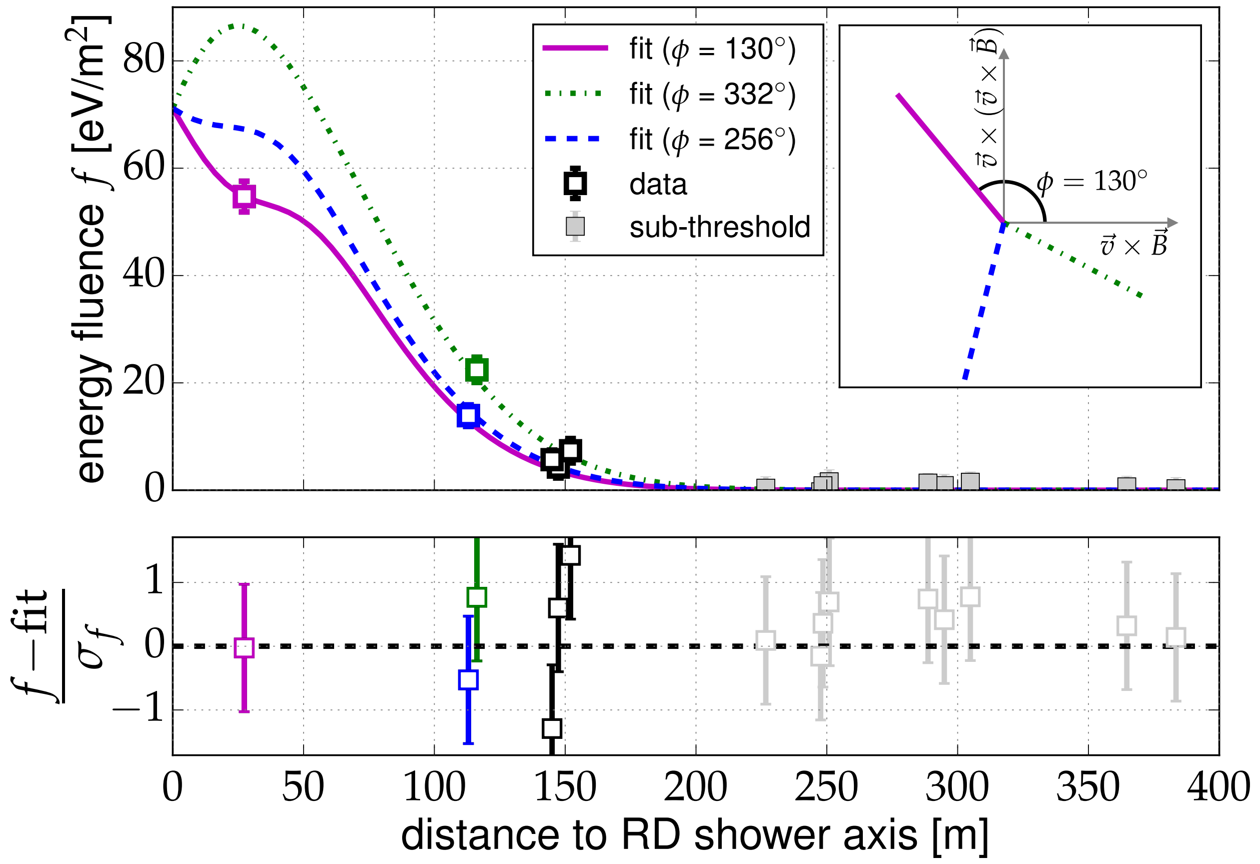
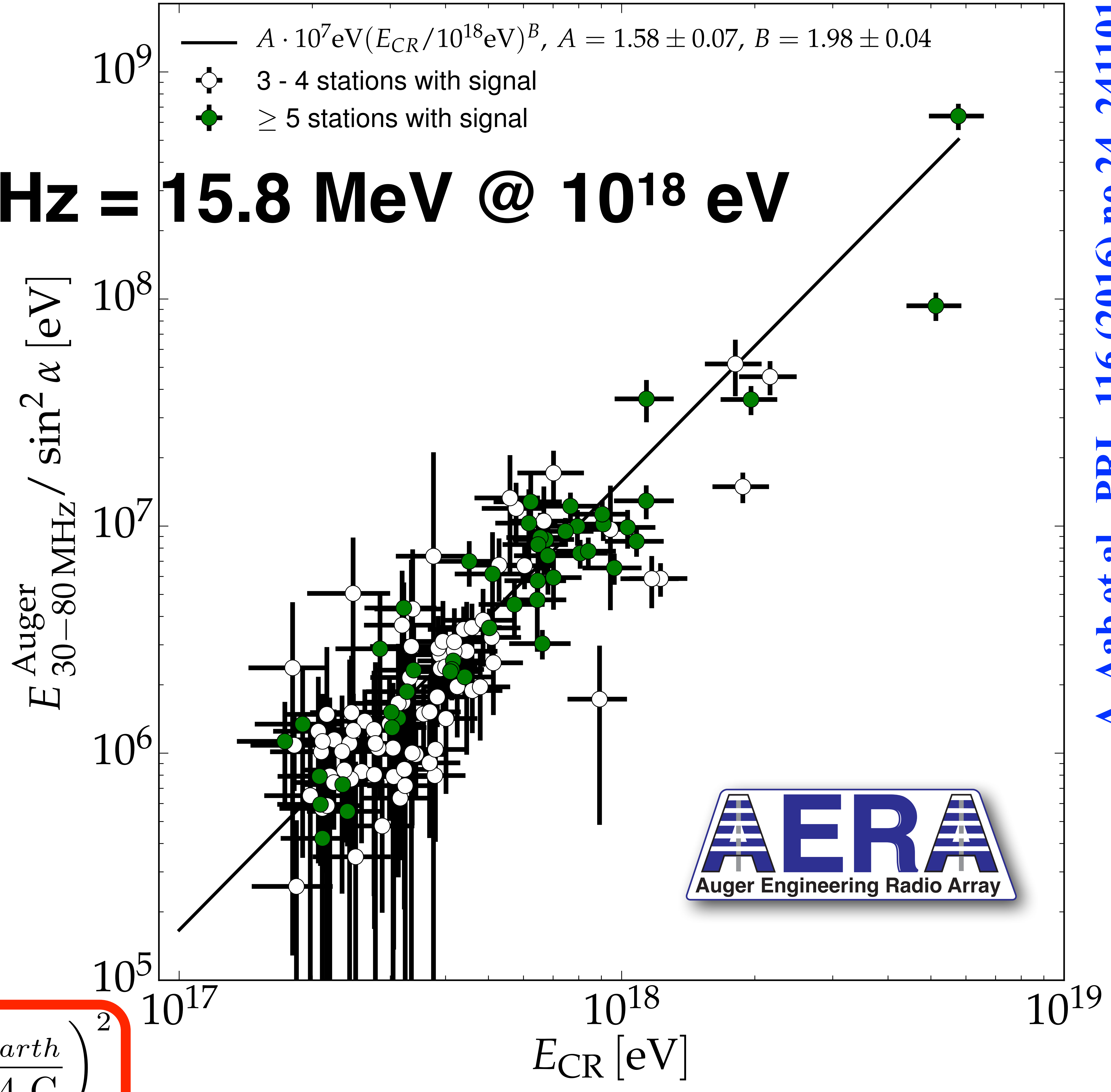




$E_{30\text{-}80}$ MHz = 15.8 MeV @ 10^{18} eV



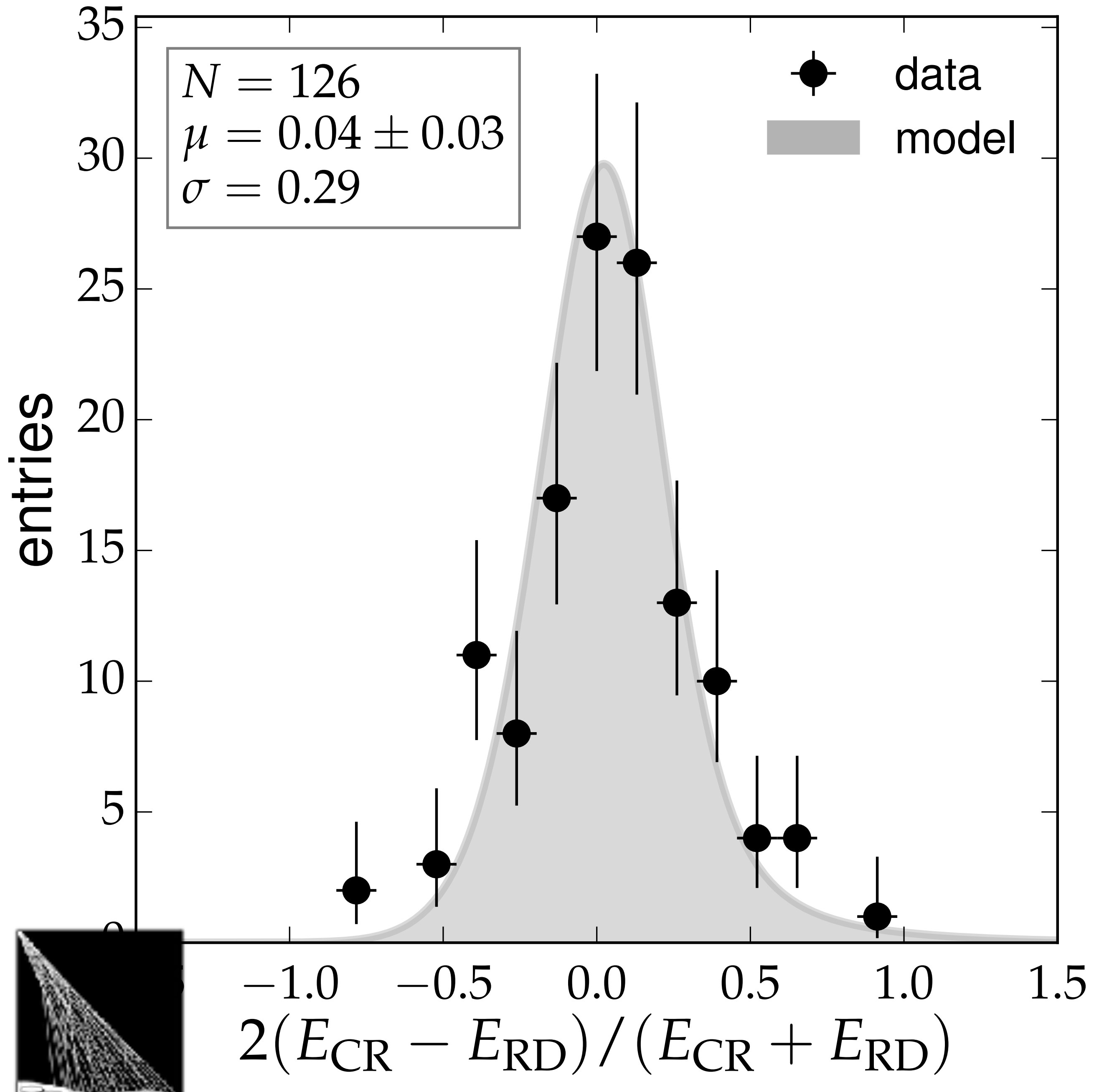
PIERRE AUGER OBSERVATORY

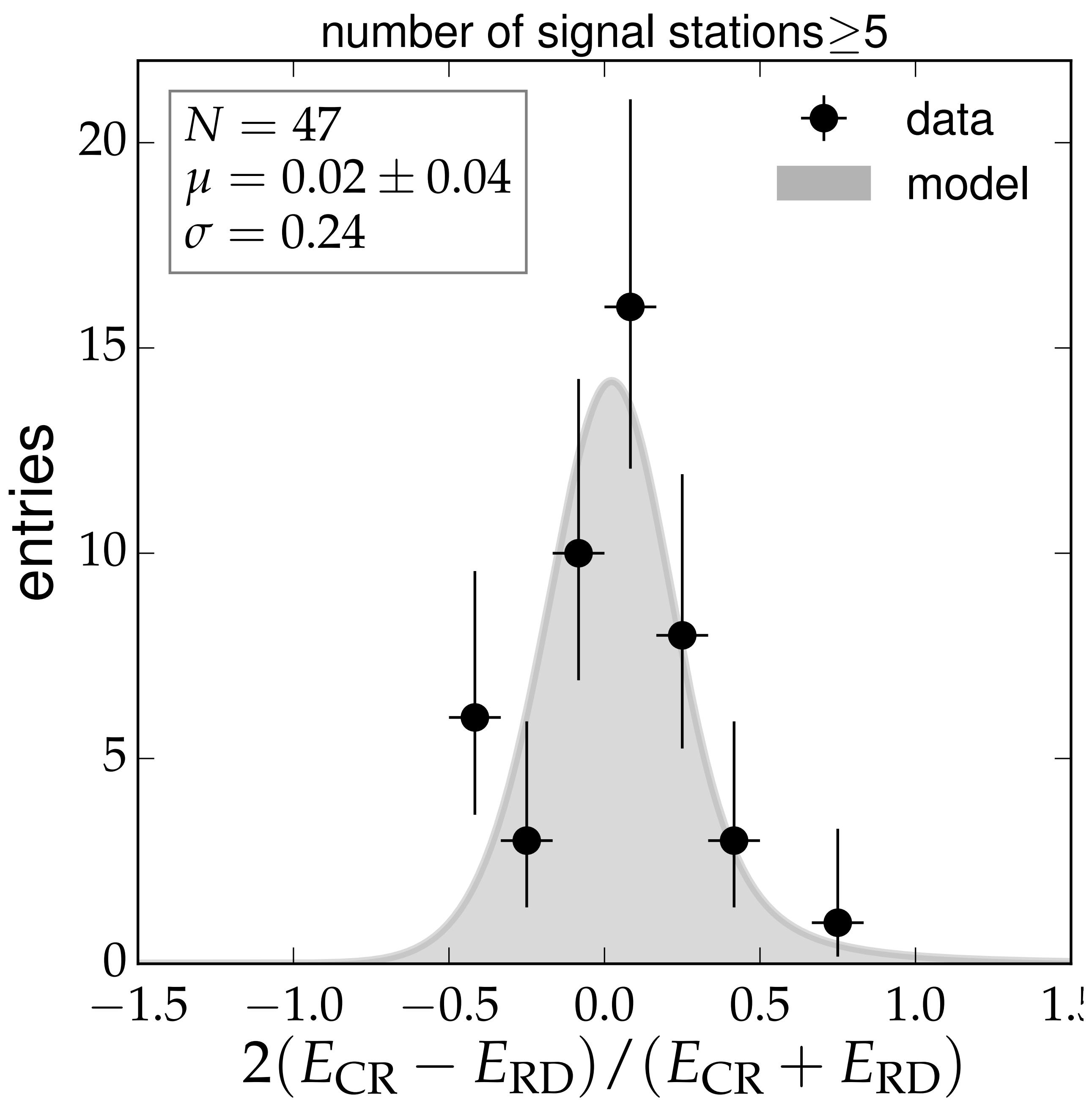
$$E_{30\text{-}80\ \mathrm{MHz}} = (15.8 \pm 0.7(\mathrm{stat}) \pm 6.7(\mathrm{syst})\ \mathrm{MeV}) \times \left(\sin\alpha \frac{E_{CR}}{10^{18}\ \mathrm{eV}} \frac{B_{Earth}}{0.24\ \mathrm{G}}\right)^2$$

A. Aab et al., PRL 116 (2016) no.24, 241101

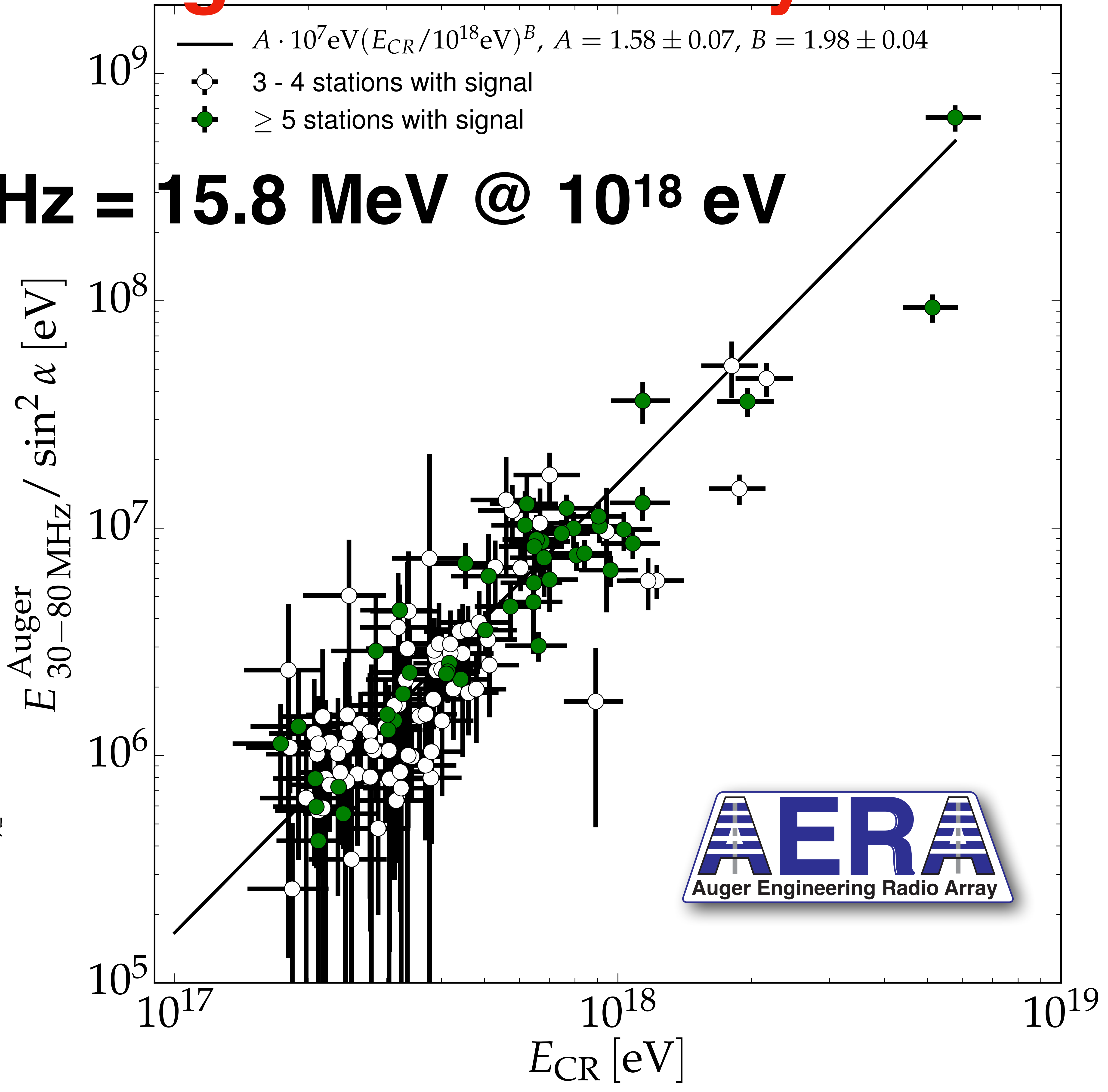
Energy Estimation of Cosmic Rays with the Engineering Radio Array of the Pierre Auger Observatory

$E_{30\text{-}80}$ MHz = 15.8 MeV @ 10^{18} eV





$\sigma \approx 24\%$

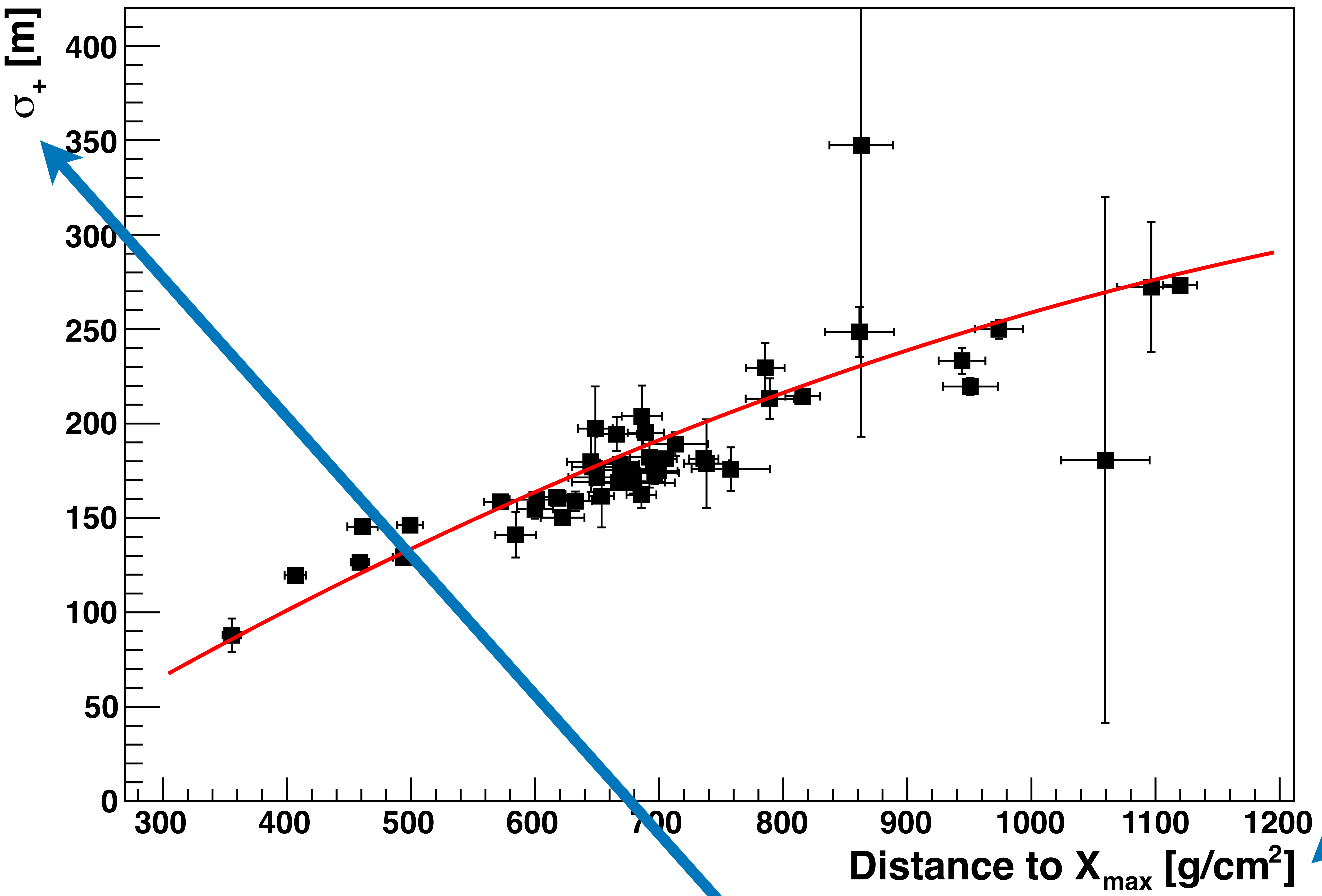


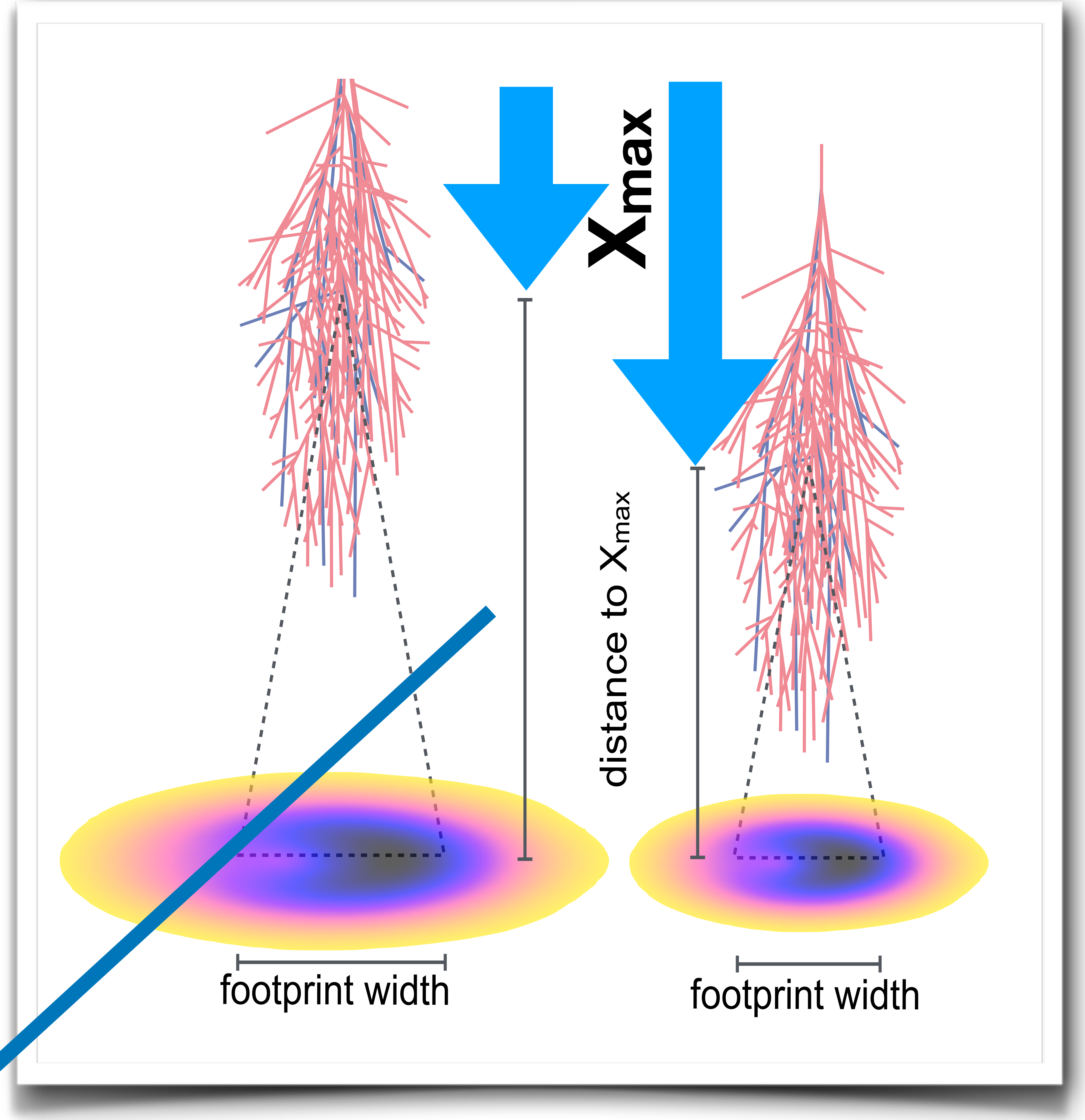
A. Aab et al., PRL 116 (2016) no.24, 241101

A. Aab et al., PRD 93 (2016) no.12, 122005

Particle type
Mass

Distance to Xmax

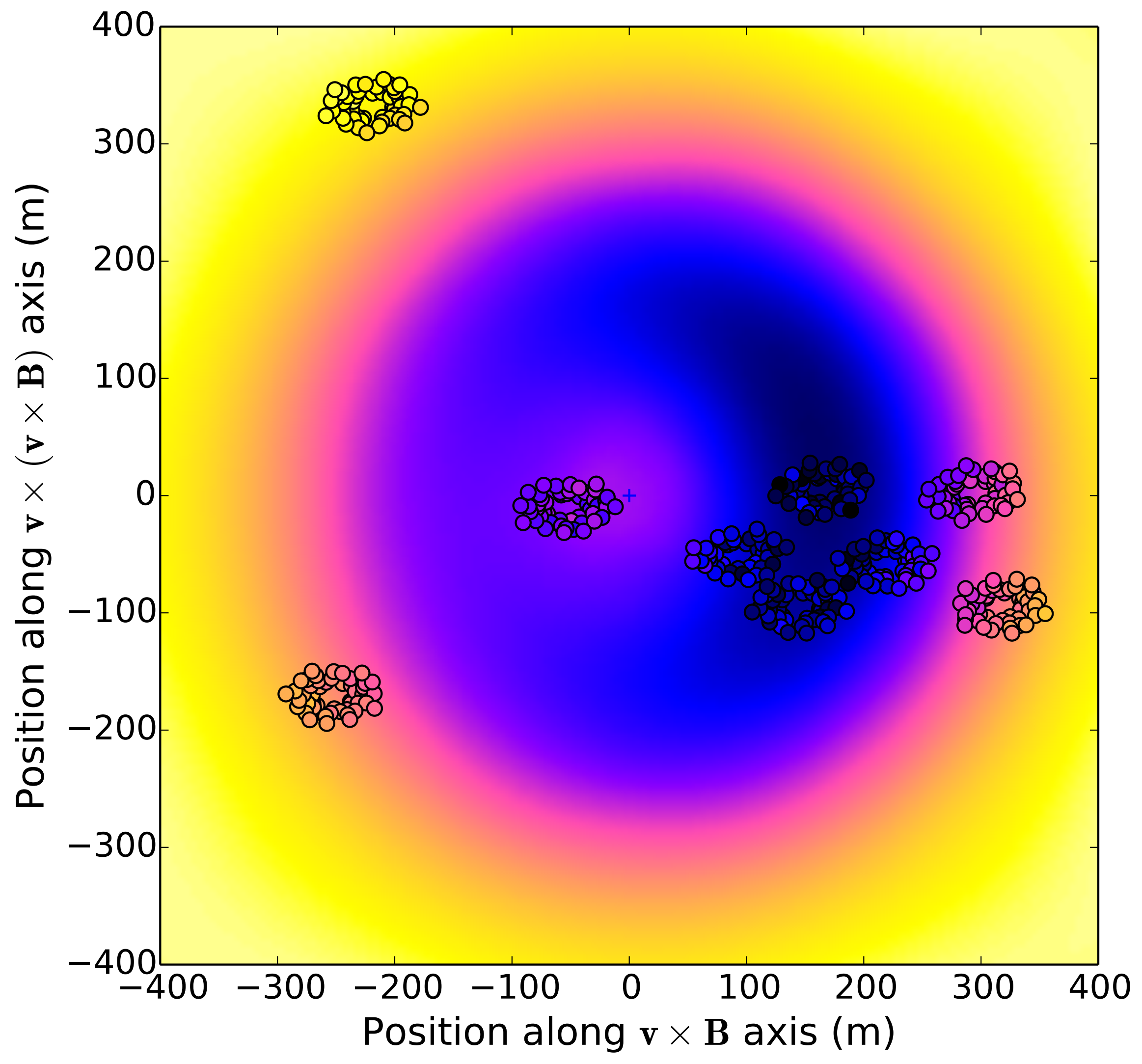


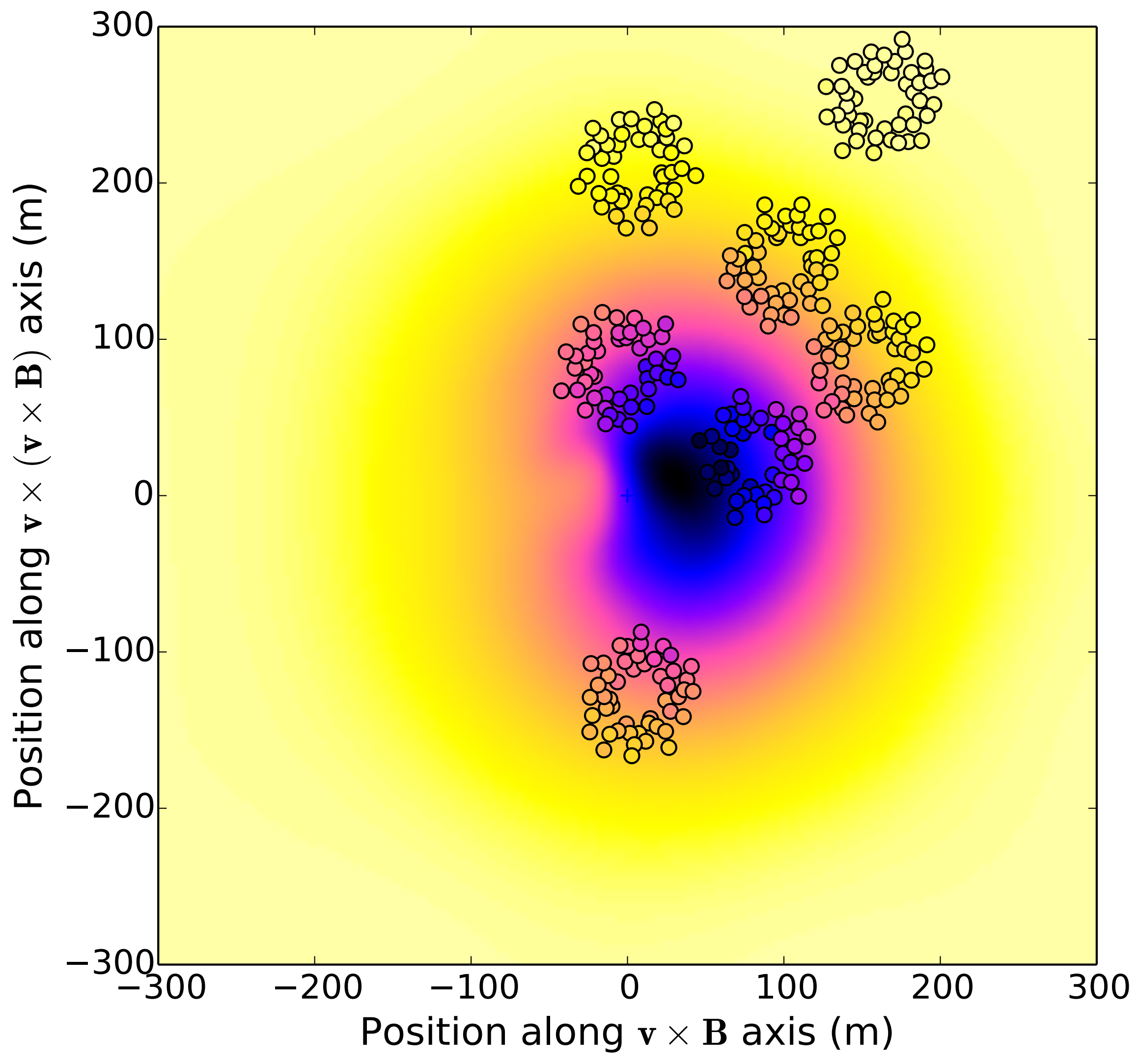


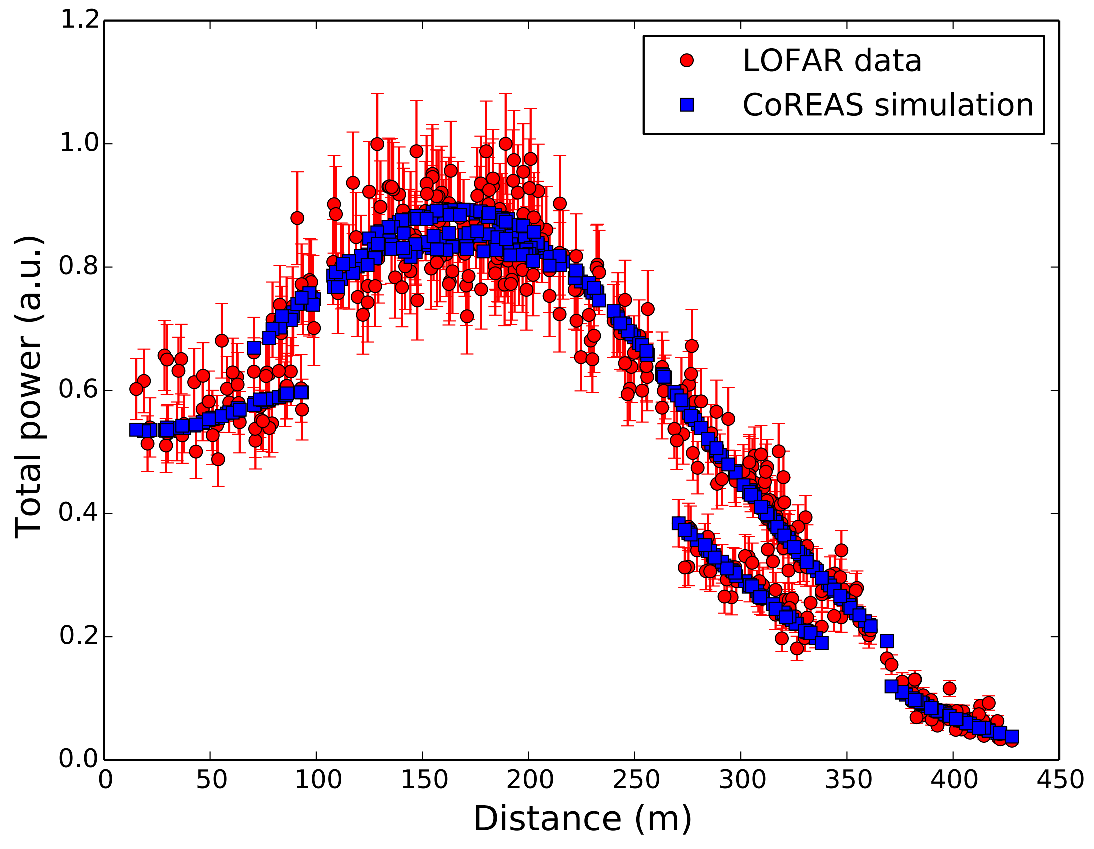
$$P(x',y') = A_+ \cdot \exp\left(\frac{-[(x'-X_+)^2+(y'-Y_+)^2]}{\sigma_+^2}\right) - A_- \cdot \exp\left(\frac{-[(x'-X_-)^2+(y'-Y_-)^2]}{\sigma_-^2}\right) + O$$

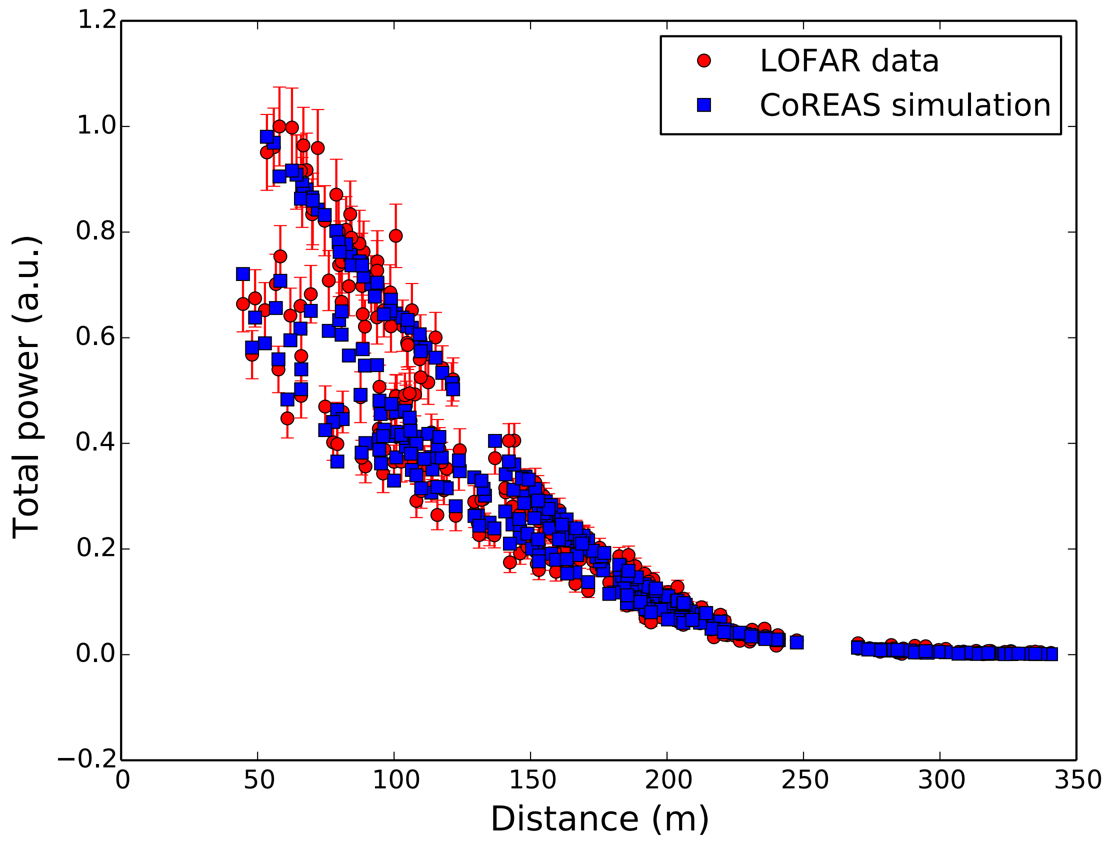
A. Nelles et al., JCAP 05 (2015) 018

Measurement of particle mass



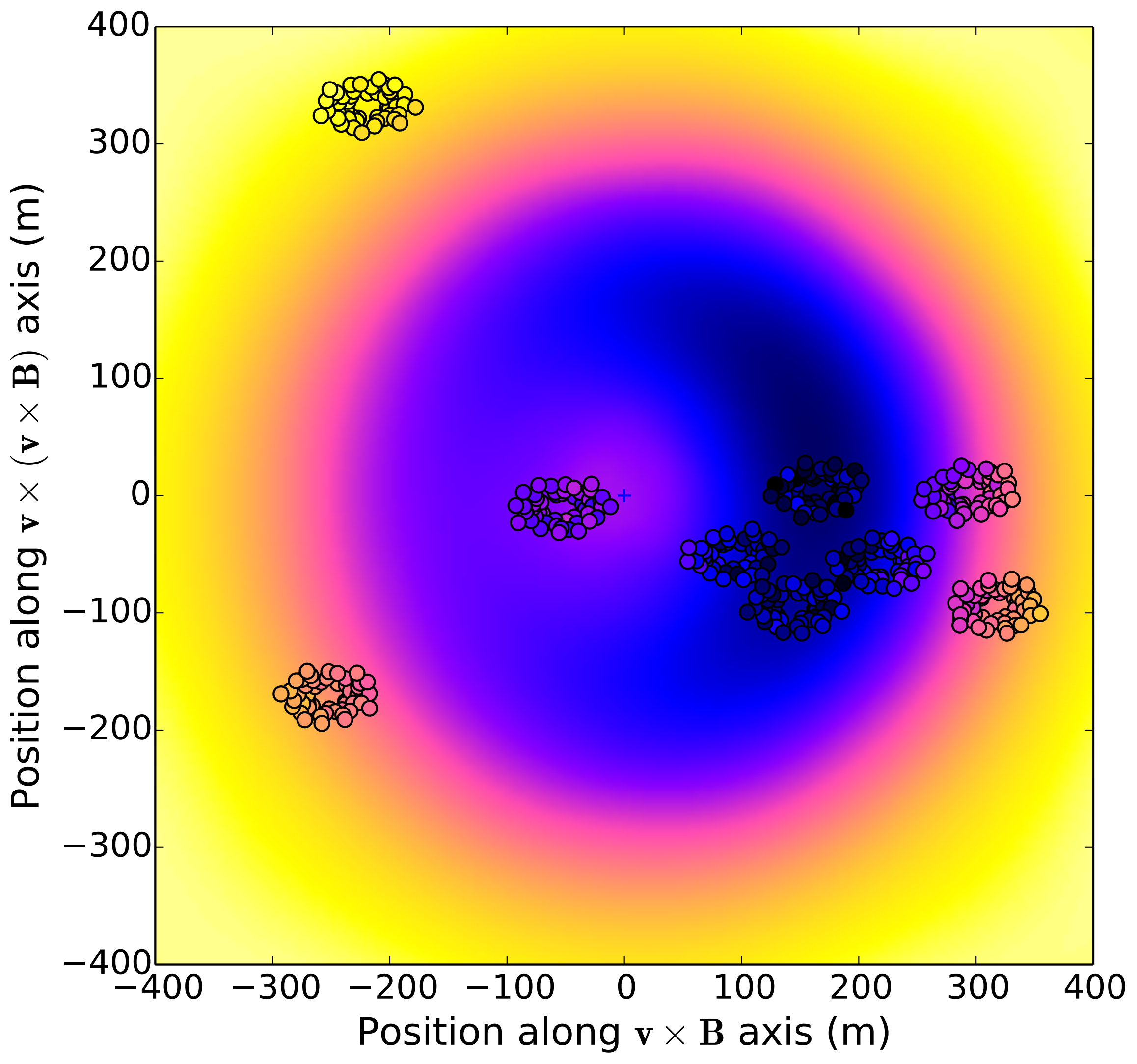


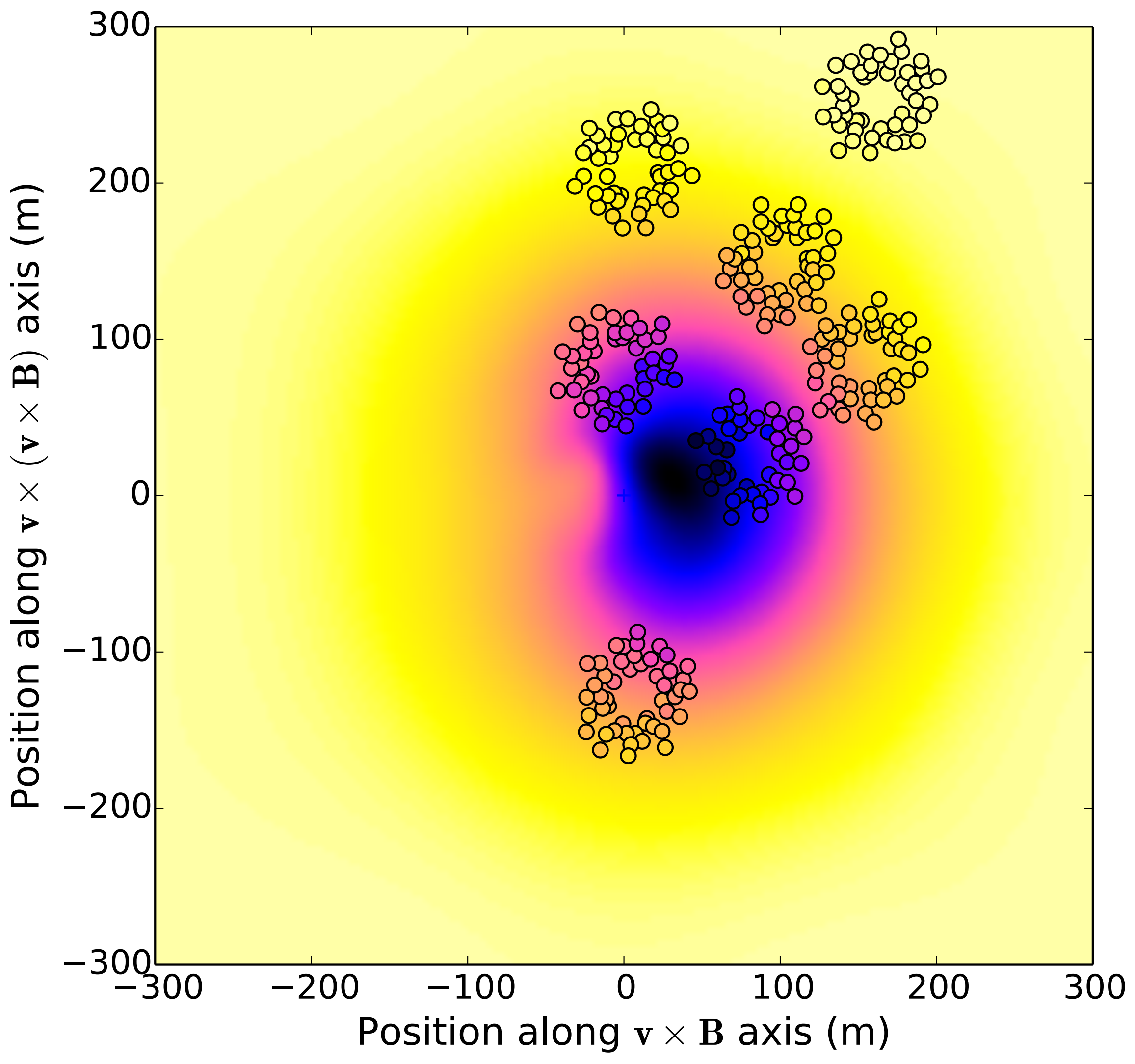


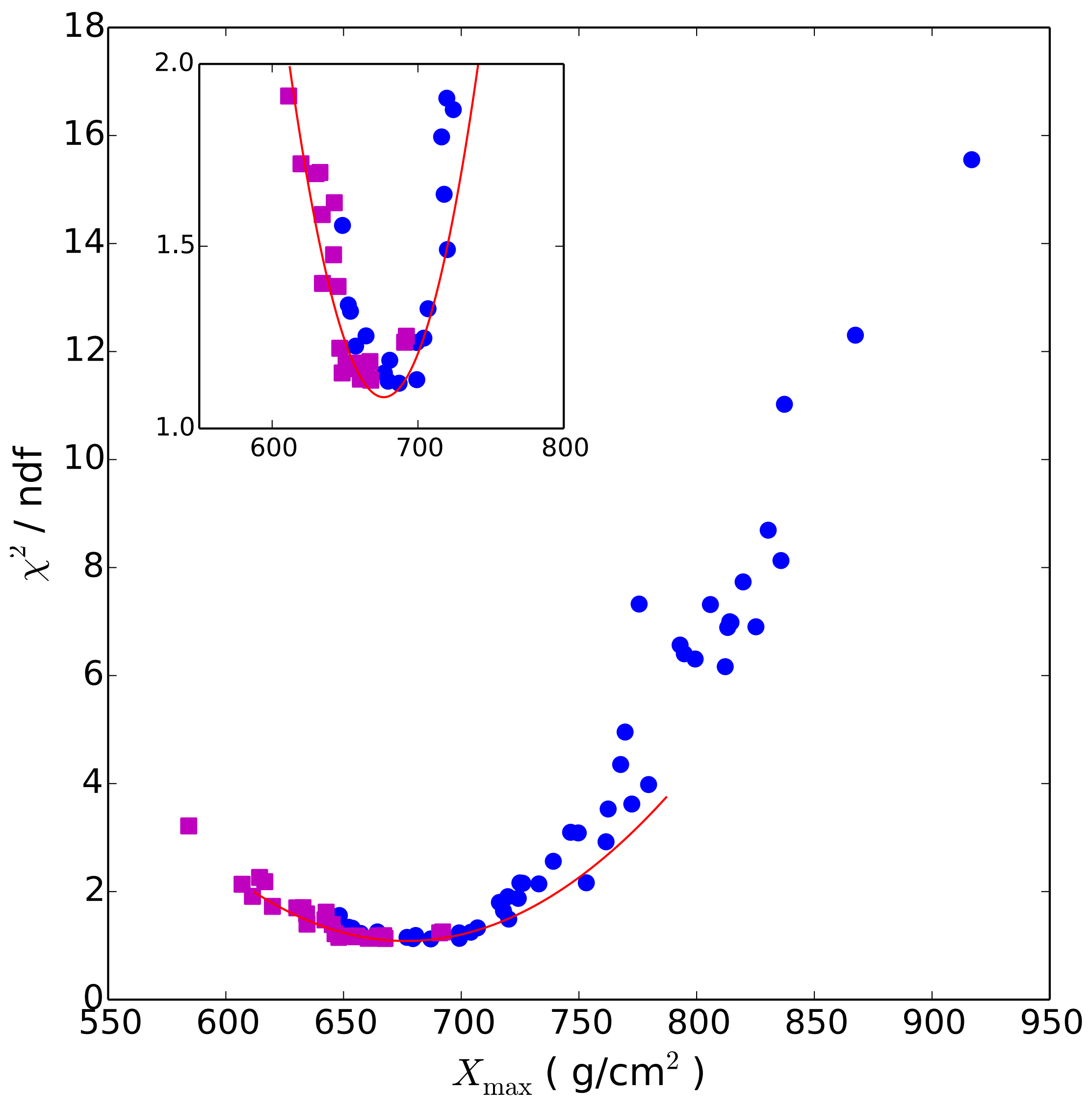


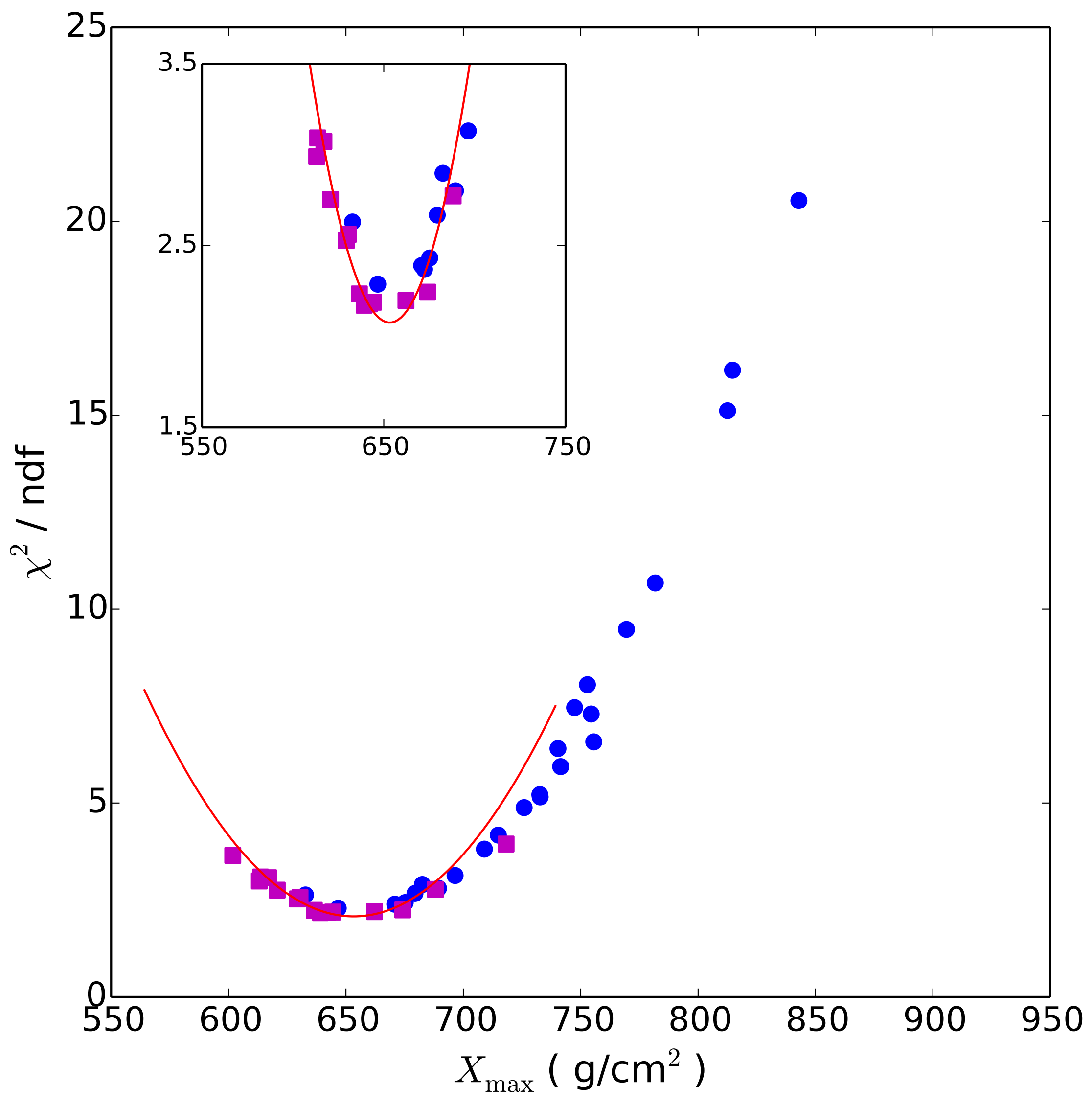
S. Buitink et al., PRD 90 (2014) 082003

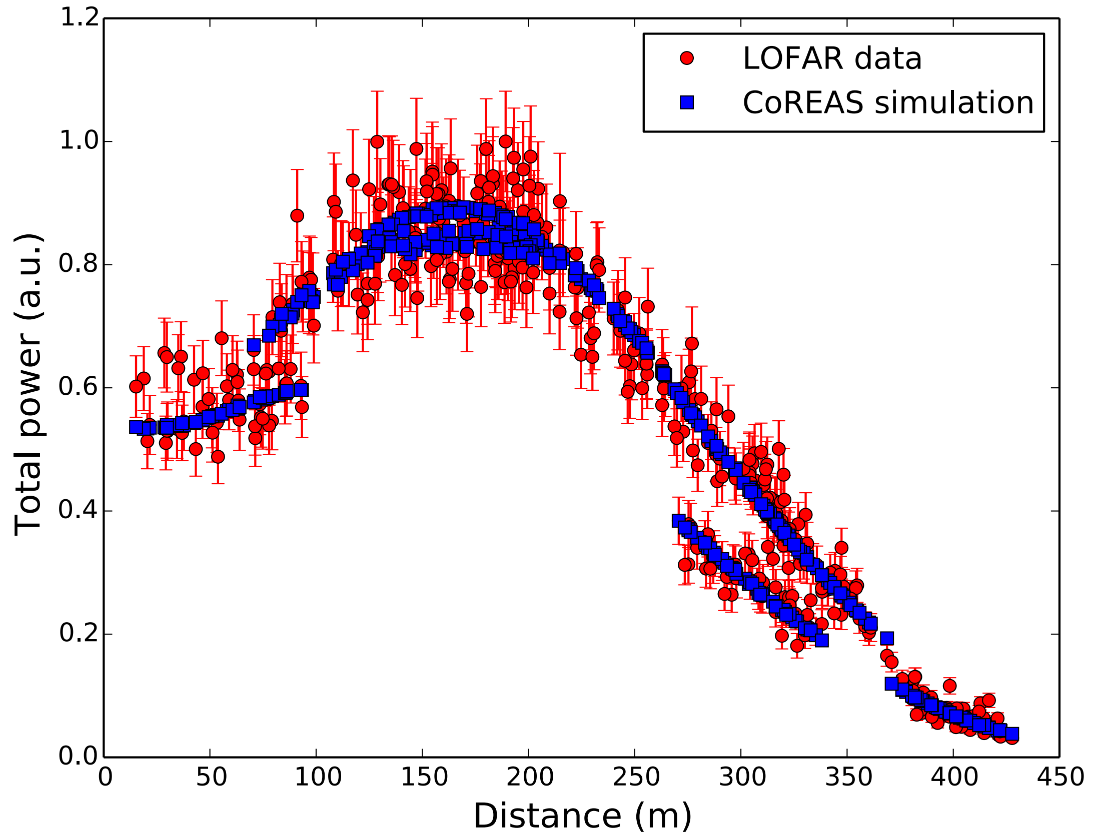
Measurement of particle mass

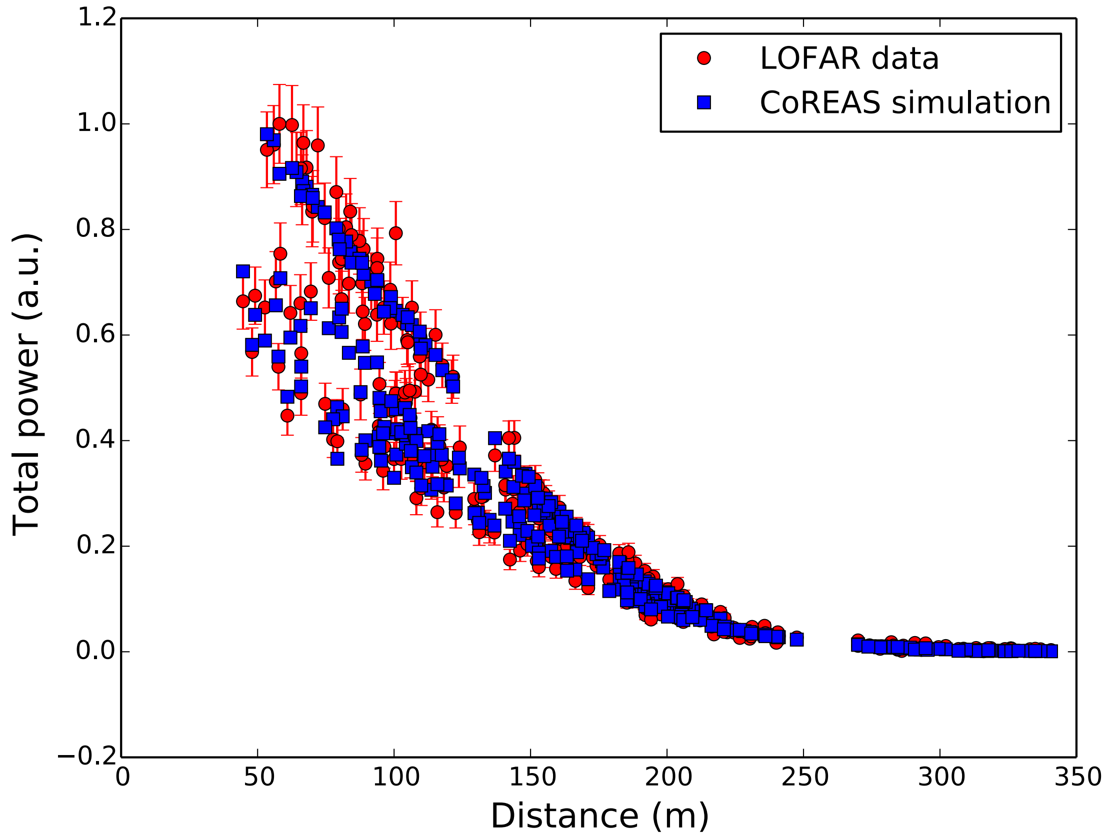












S. Buitink et al., PRD 90 (2014) 082003

Measurement of particle mass

$\sigma_E \approx 32\%$

$\sigma_{X_{max}} \approx 17\ \mathrm{g/cm^2}$

S. Buitink et al., PRD 90 (2014) 082003

Depth of the shower maximum

LETTER nature
doi:10.1038/nature16976

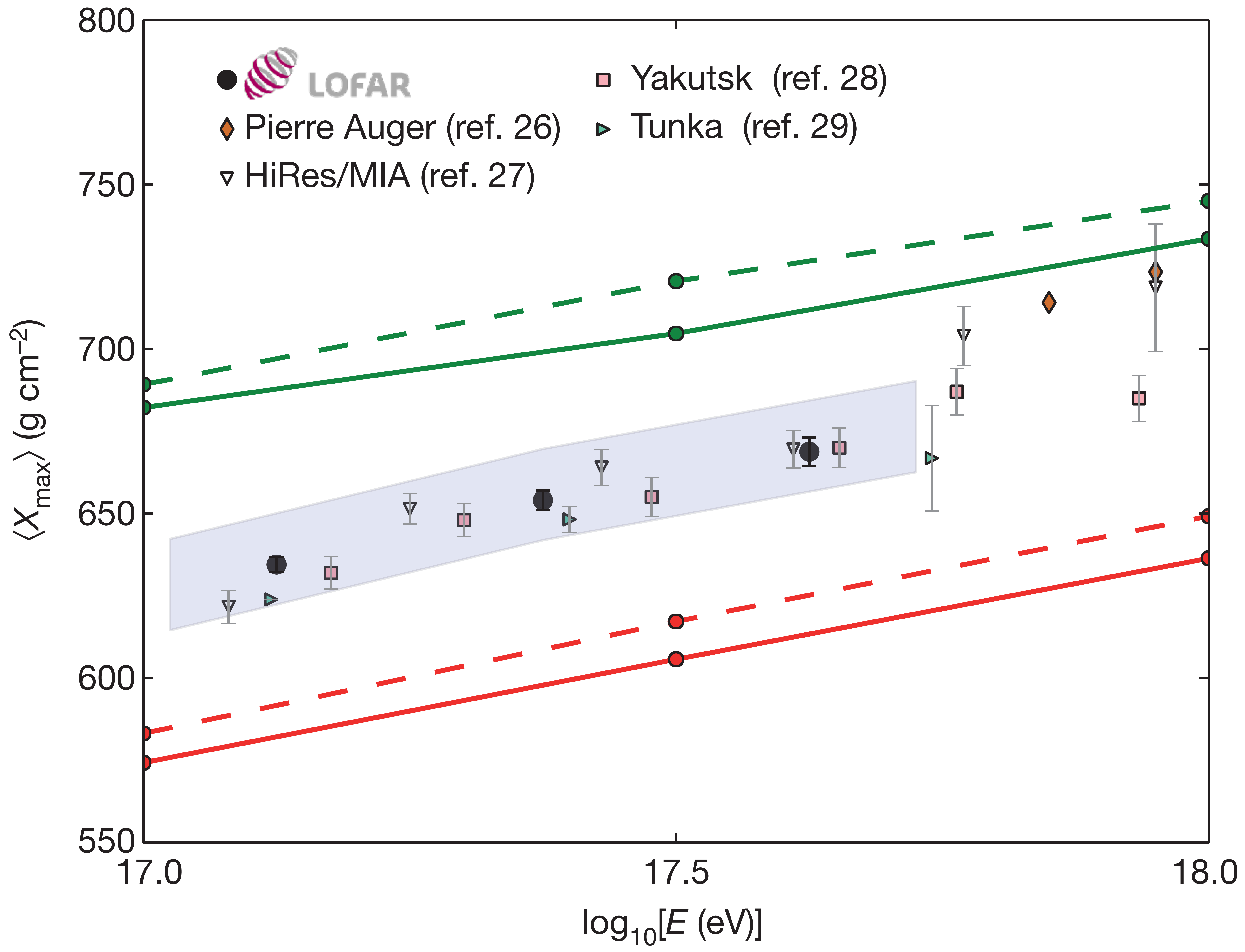
A large light-mass component of cosmic rays at 10^{17}–$10^{17.5}$ electronvolts from radio observations

S. Buitink[1,2], A. Corstanje[2], H. Falcke[2,3,4,5], J. R. Hörandel[2,4], T. Huege[6], A. Nelles[2,7], J. P. Rachen[2], L. Rossetto[2], P. Schellart[2], O. Scholten[8,9], S. ter Veen[3], S. Thoudam[2], T. N. G. Trinh[8], J. Anderson[10], A. Asgekar[3,11], I. M. Avruch[12,13], M. E. Bell[14], M. J. Bentum[3,15], G. Bernardi[16,17], P. Best[18], A. Bonafede[19], F. Breitling[20], J. W. Broderick[21], W. N. Brouw[3,13], M. Brüggen[19], H. R. Butcher[22], D. Carbone[23], B. Ciardi[24], J. E. Conway[25], F. de Gasperin[19], E. de Geus[3,26], A. Deller[3], R.-J. Dettmar[27], G. van Diepen[3], S. Duscha[3], J. Eislöffel[28], D. Engels[29], J. E. Enriquez[3], R. A. Fallows[3], R. Fender[30], C. Ferrari[31], W. Frieswijk[3], M. A. Garrett[3,32], J. M. Grießmeier[33,34], A. W. Gunst[3], M. P. van Haarlem[3], T. E. Hassall[21], G. Heald[3,13], J. W. T. Hessels[3,23], M. Hoeft[28], A. Horneffer[5], M. Iacobelli[3], H. Intema[32,35], E. Juette[27], A. Karastergiou[30], V. I. Kondratiev[3,36], M. Kramer[5,37], M. Kuniyoshi[38], G. Kuper[3], J. van Leeuwen[3,23], G. M. Loose[3], P. Maat[3], G. Mann[20], S. Markoff[23], R. McFadden[3], D. McKay-Bukowski[39,40], J. P. McKean[3,13], M. Mevius[3,13], D. D. Mulcahy[21], H. Munk[3], M. J. Norden[3], E. Orru[3], H. Paas[41], M. Pandey-Pommier[42], V. N. Pandey[3], M. Pietka[30], R. Pizzo[3], A. G. Polatidis[3], W. Reich[5], H. J. A. Röttgering[32], A. M. M. Scaife[21], D. J. Schwarz[43], M. Serylak[30], J. Sluman[3], O. Smirnov[17,44], B. W. Stappers[37], M. Steinmetz[20], A. Stewart[30], J. Swinbank[23,45], M. Tagger[33], Y. Tang[3], C. Tasse[44,46], M. C. Toribio[3,32], R. Vermeulen[3], C. Vocks[20], C. Vogt[3], R. J. van Weeren[16], R. A. M. J. Wijers[23], S. J. Wijnholds[3], M. W. Wise[3,23], O. Wucknitz[5], S. Yatawatta[3], P. Zarka[47] & J. A. Zensus[5]

Cosmic rays are the highest-energy particles found in nature. Measurements of the mass composition of cosmic rays with energies of 10^{17}–10^{18} electronvolts are essential to understanding whether they have galactic or extragalactic sources. It has also been proposed that the astrophysical neutrino signal[1] comes from accelerators capable of producing cosmic rays of these energies[2]. Cosmic rays initiate air showers—cascades of secondary particles in the atmosphere—and their masses can be inferred from measurements of the atmospheric depth of the shower maximum[3] (X_{max}; the depth of the air shower when it contains the most particles) or of the composition of shower particles reaching the ground[4]. Current measurements[5] have either high uncertainty, or a low duty cycle and a high energy threshold. Radio detection of cosmic rays[6–8] is a rapidly developing technique[9] for determining X_{max} (refs 10, 11) with a duty cycle of, in principle, nearly 100 per cent. The radiation is generated by the separation of relativistic electrons and positrons in the geomagnetic field and a negative charge excess in the shower front[6,12]. Here we report radio measurements of X_{max} with a mean uncertainty of 16 grams per square centimetre for air showers initiated by cosmic rays with energies of 10^{17}–$10^{17.5}$ electronvolts. This high resolution in X_{max} enables us to determine the mass spectrum of the cosmic rays: we find a mixed composition, with a light-mass fraction (protons and helium nuclei) of about 80 per cent. Unless, contrary to current expectations, the extragalactic component of cosmic rays contributes substantially to the total flux below $10^{17.5}$ electronvolts, our measurements indicate the existence of an additional galactic component, to account for the light composition that we measured in the 10^{17}–$10^{17.5}$ electronvolt range.

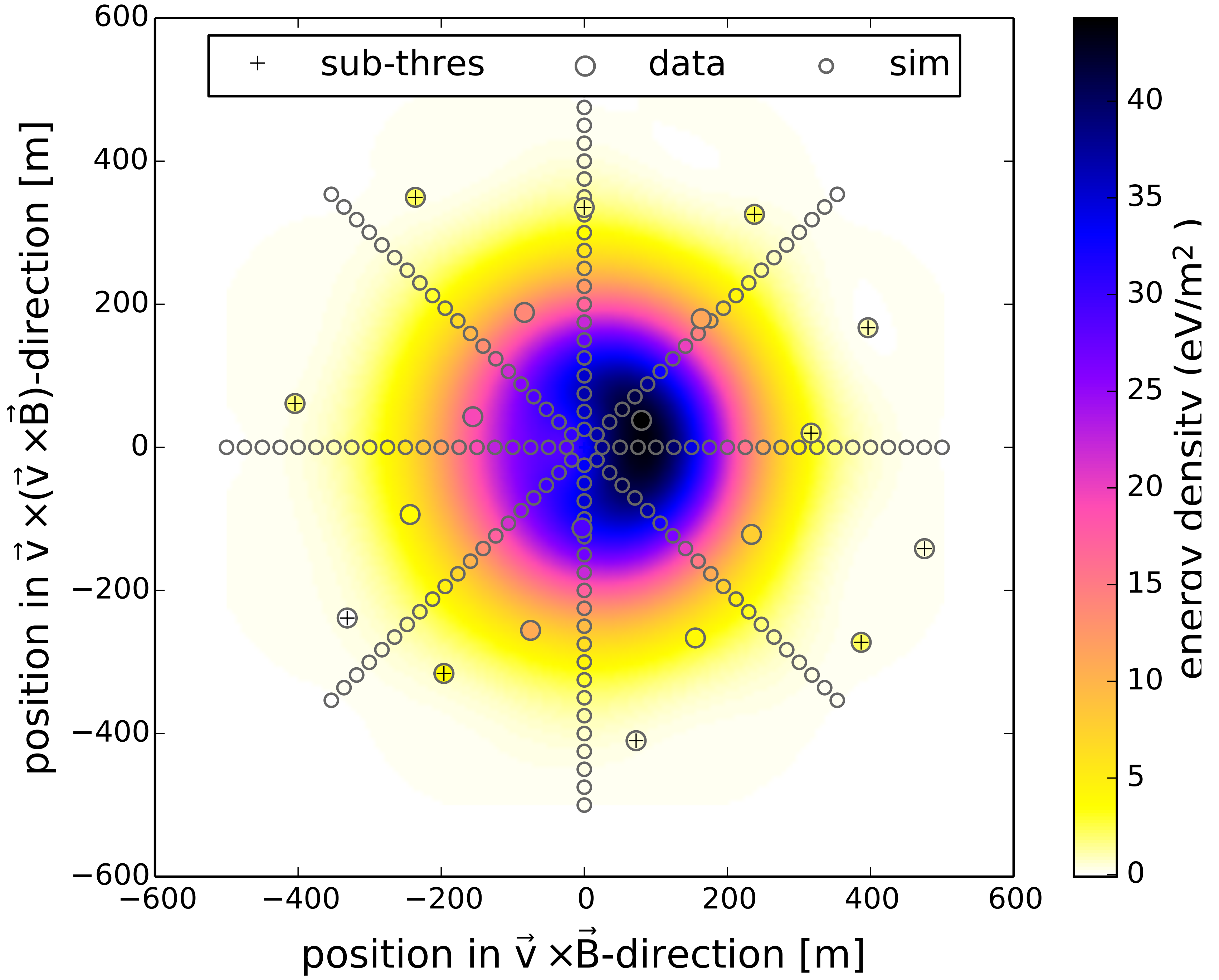
Observations were made with the Low Frequency Array (LOFAR[13]), a radio telescope consisting of thousands of crossed dipoles with built-in air-shower-detection capability[14]. LOFAR continuously records the radio signals from air showers, while simultaneously running astronomical observations. It comprises a scintillator array (LORA) that triggers the read-out of buffers, storing the full waveforms received by all antennas.

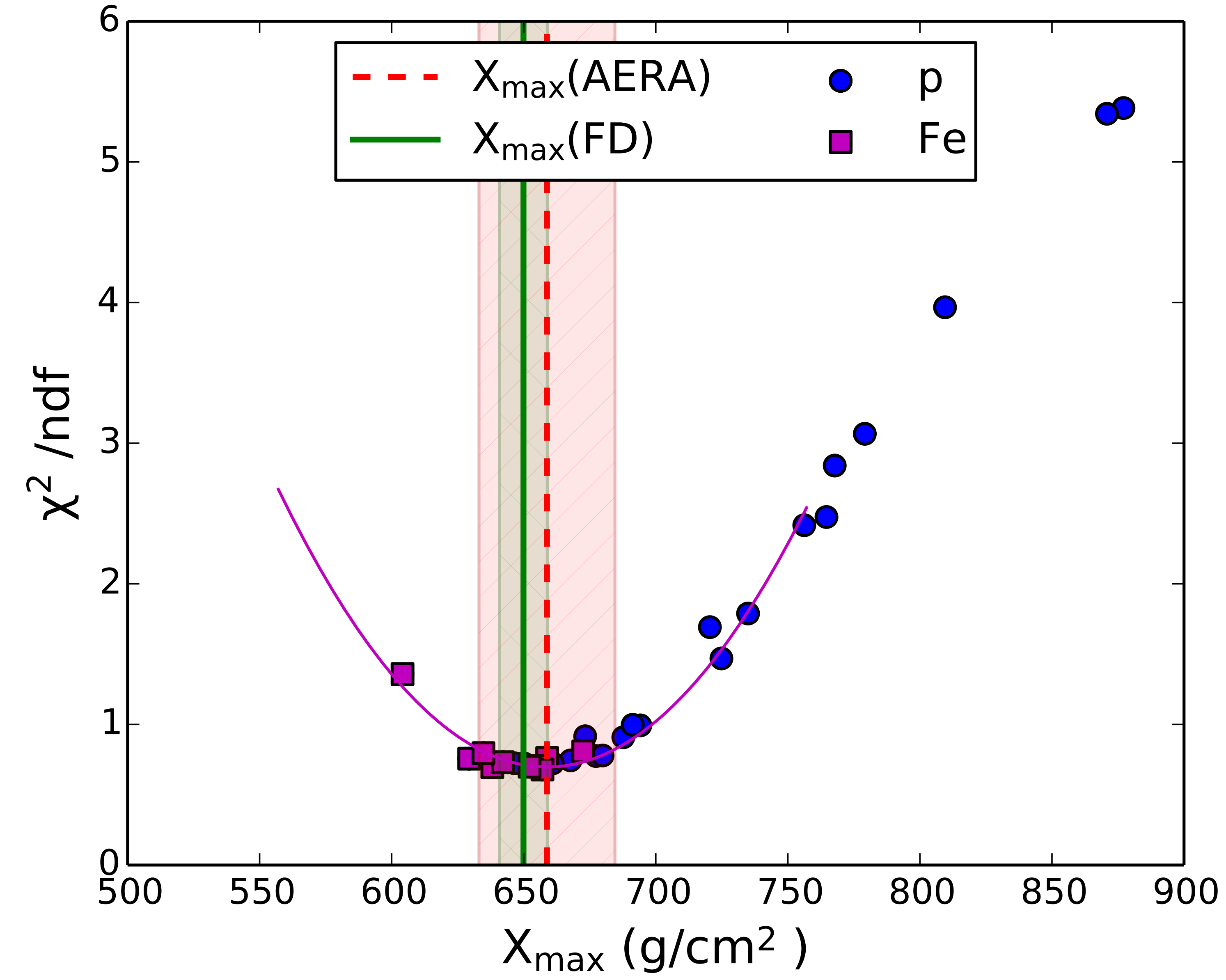
We selected air showers from the period June 2011 to January 2015 with radio pulses detected in at least 192 antennas. The total uptime was about 150 days, limited by construction and commissioning of the

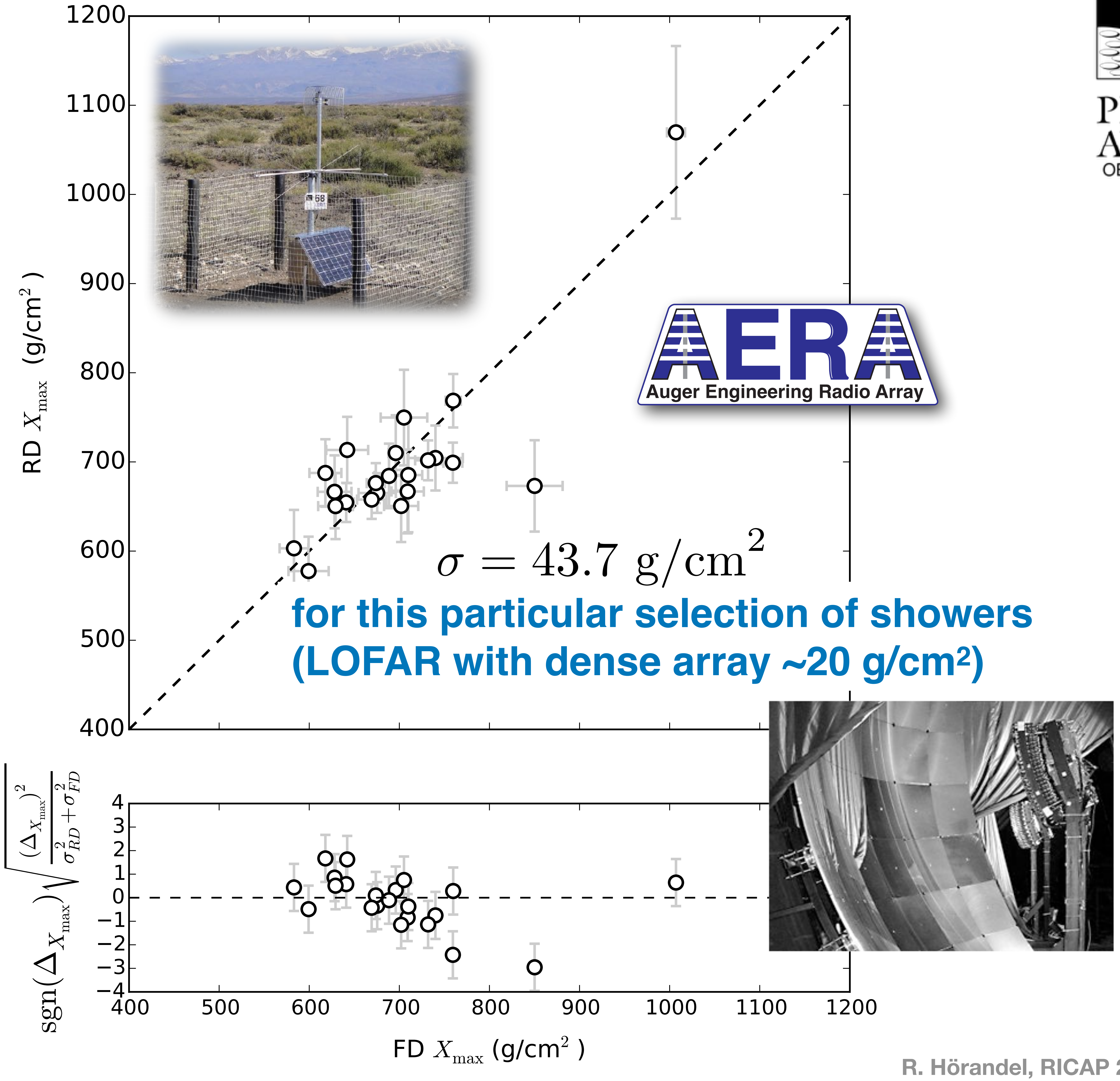


S. Buitink et al., Nature 531 (2016) 70

Xmax radio vs fluorescence







for this particular selection of showers
(LOFAR with dense array ~20 g/cm²)

J. Schulz, PhD thesis RU Nijmegen (2016)

Determine the properties of the incoming particle with the radio technique

- direction ~ 0.1° - 0.5°
- energy ~ 20% - 30%
- type (X_{max}) ~ 20 - 40 g/cm²

(depending on detector spacing)

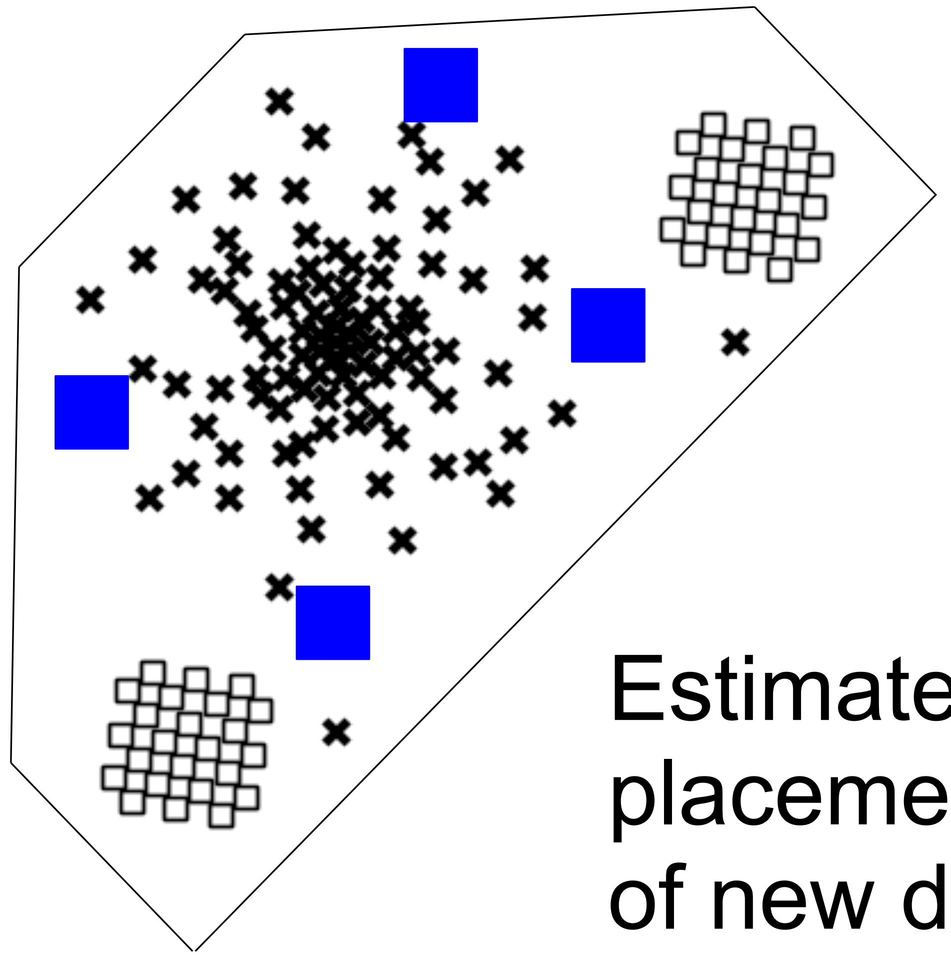
—> radio technique is routinely used to measure properties of cosmic rays

ongoing and future work

Extension of scintillator array (LORA)

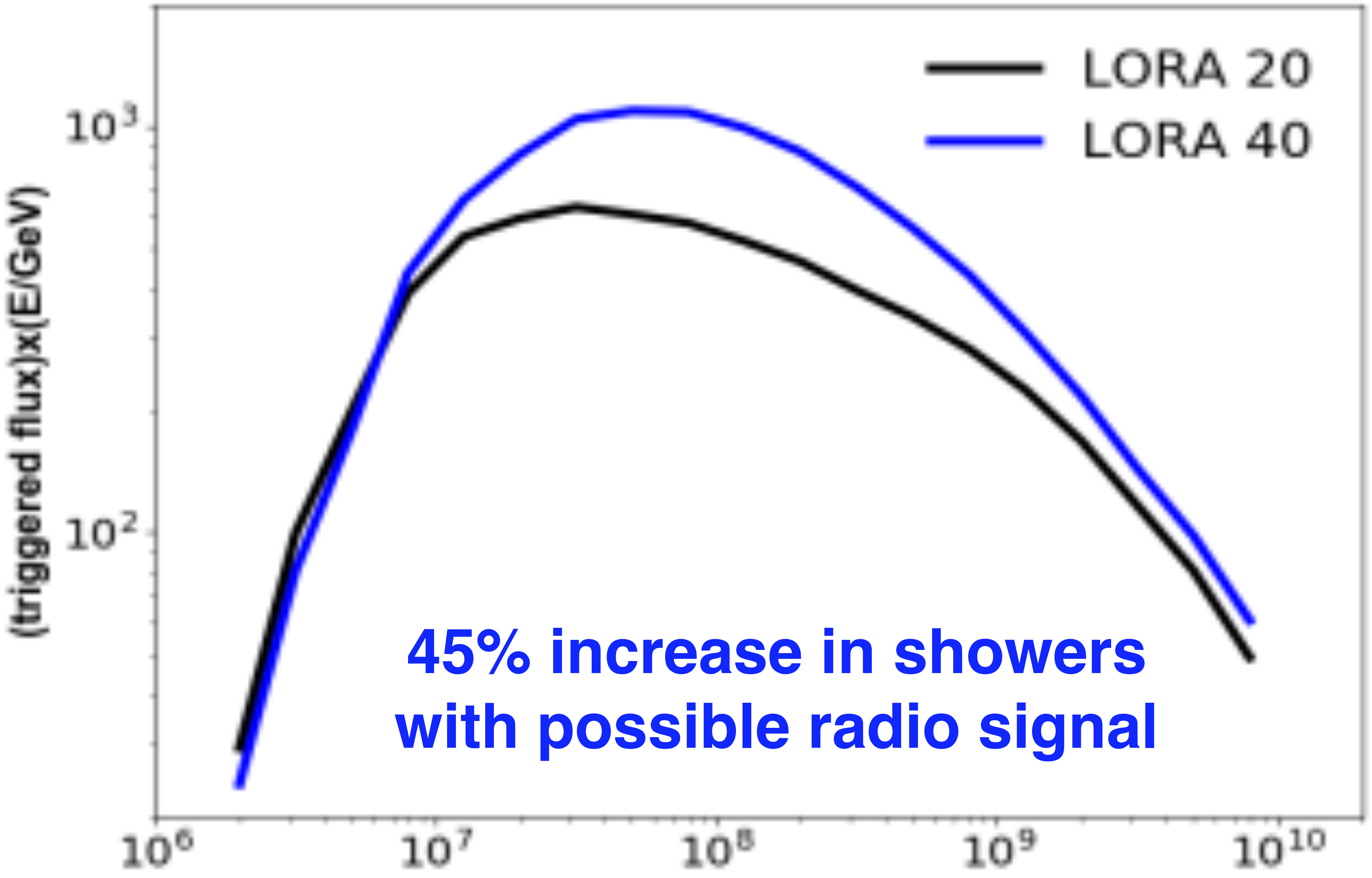
2.5 km

Existing station
New station



Estimated placement of new detectors

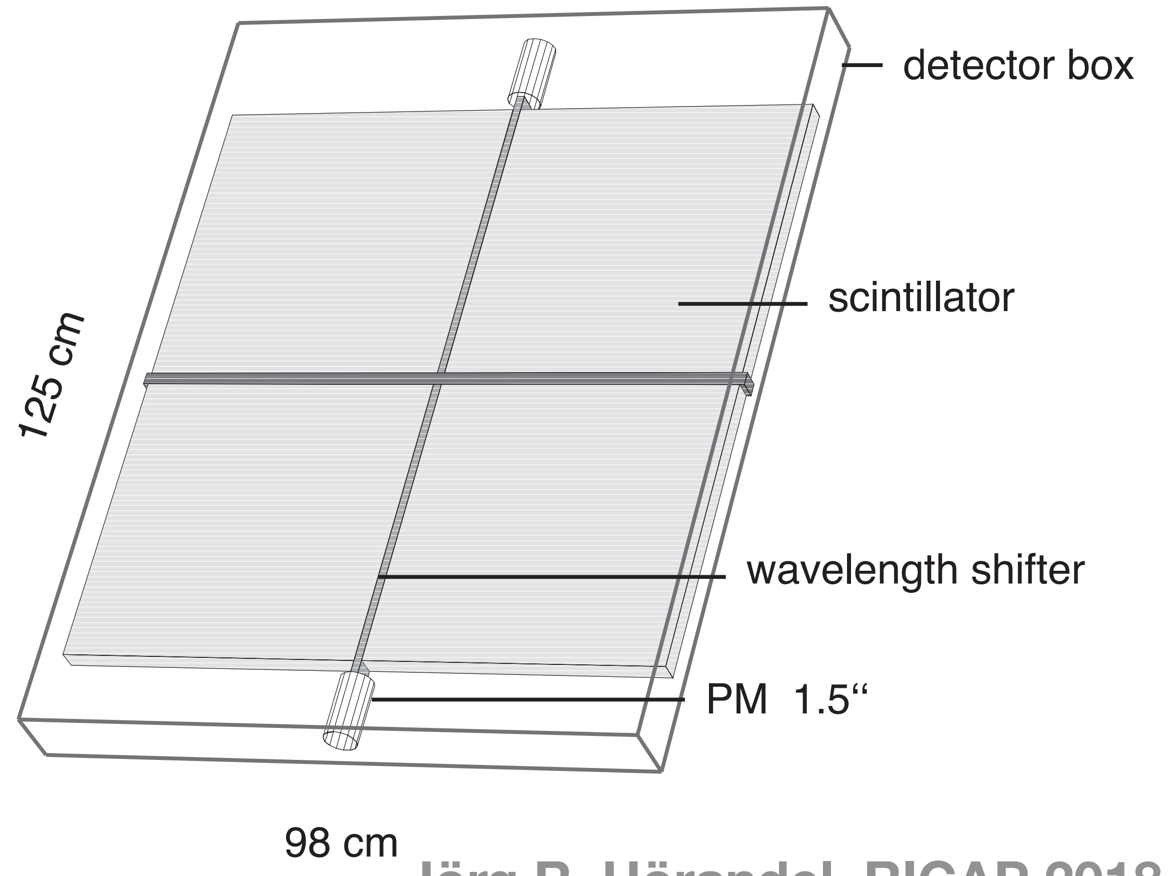
S. Buitink
K. Mulrey



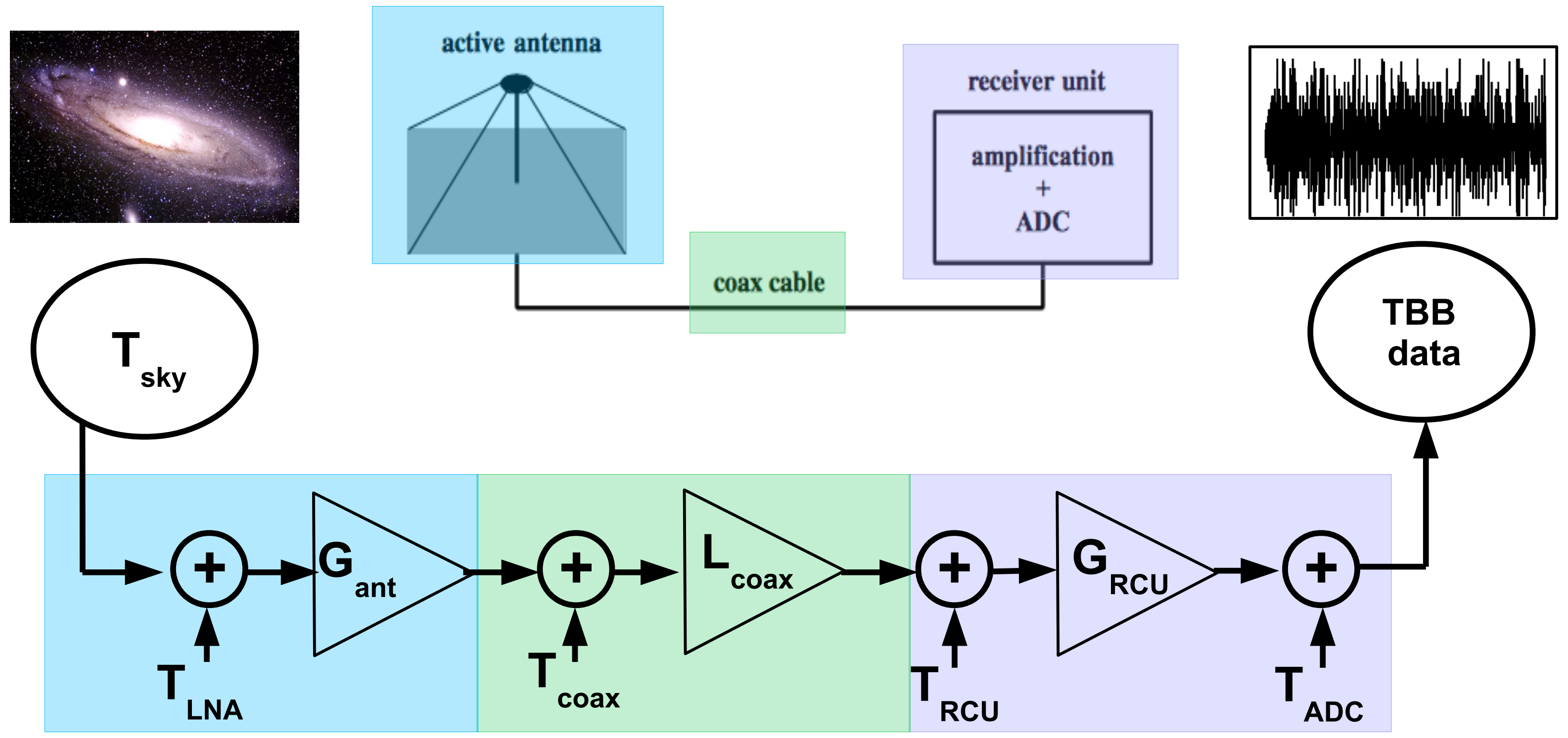
45% increase in showers with possible radio signal

10^{16} eV **10^{18} eV**

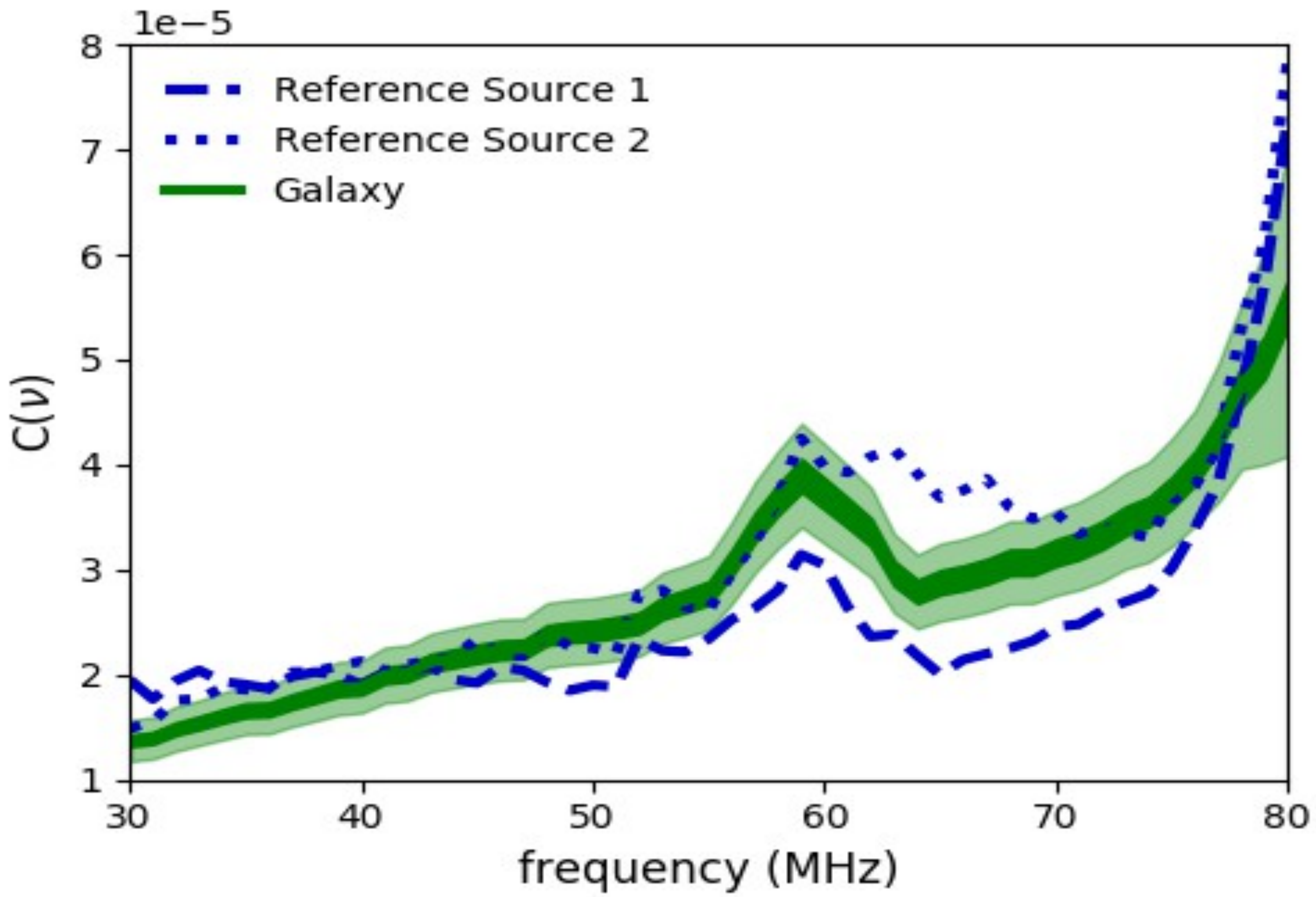
adding 20 scintillator stations in 2018

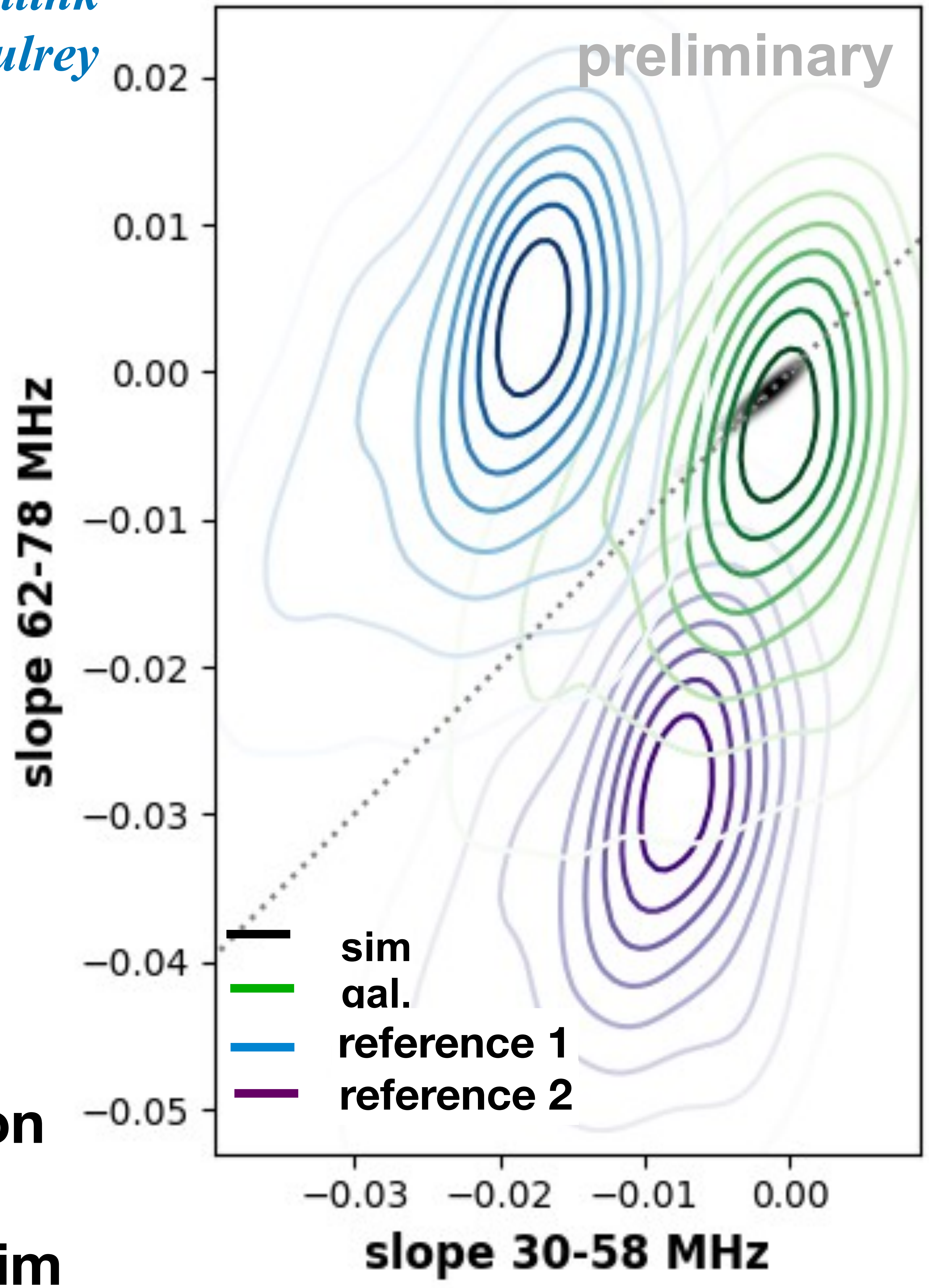


Improved calibration

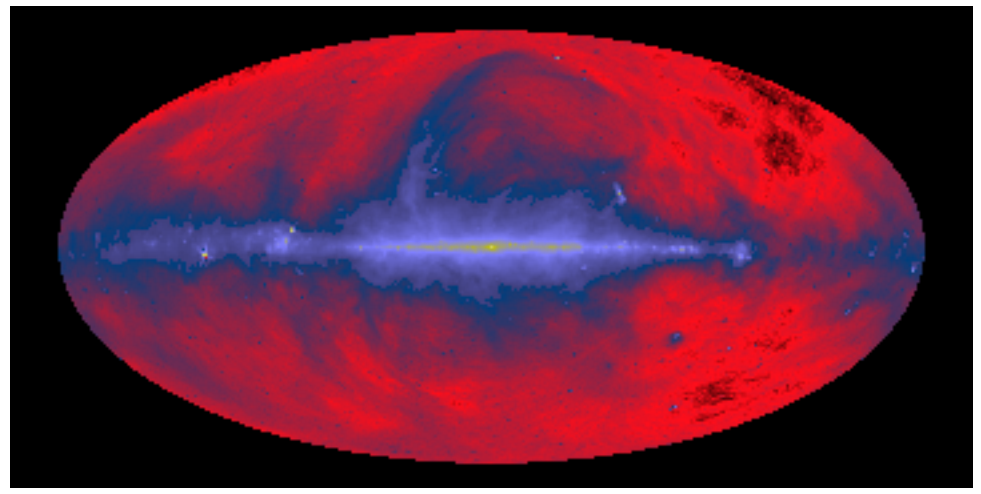


S. Buitink
K. Mulrey





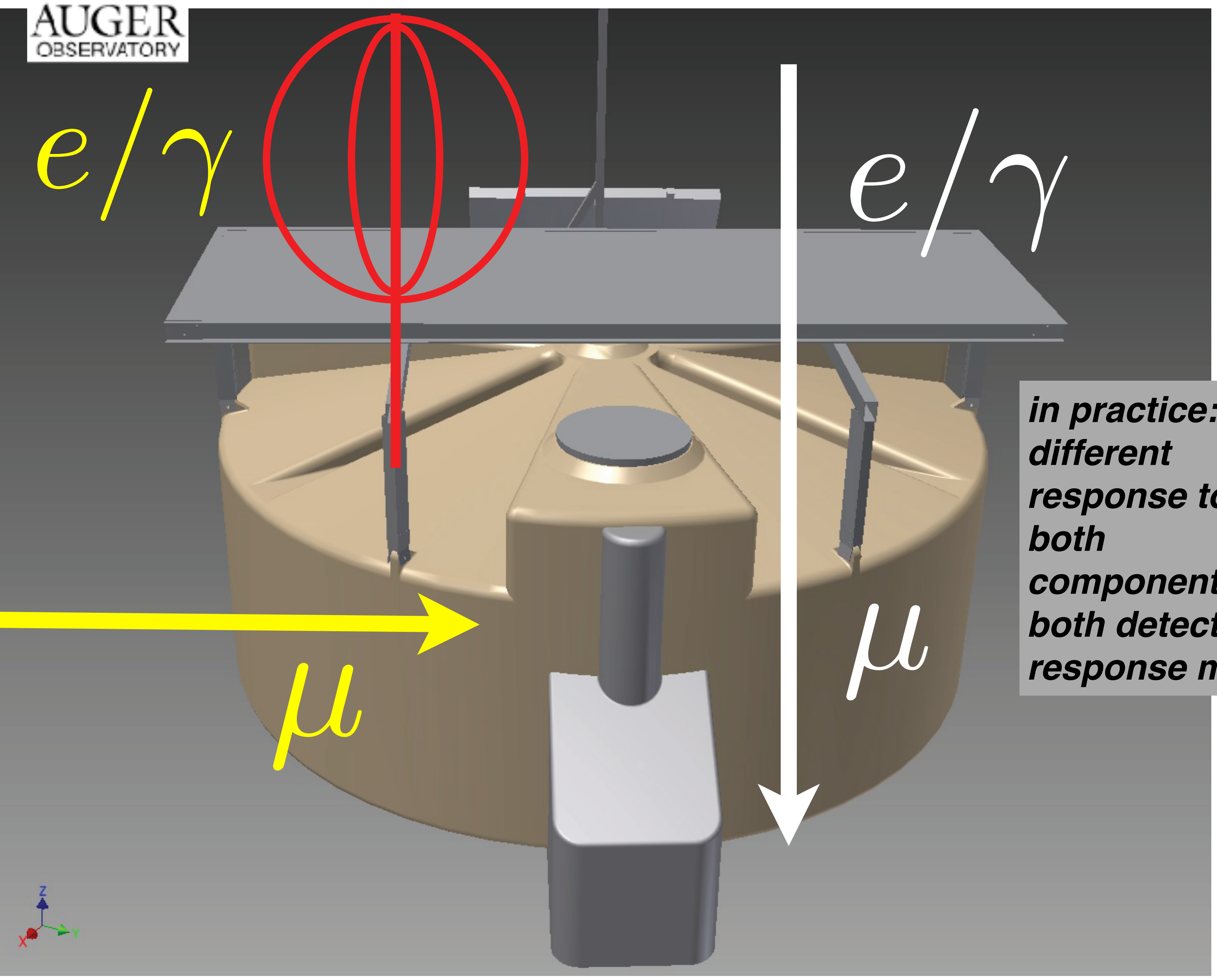
Galactic emission (LFmap)



reference source

- complete signal chain calibration
- systematic uncertainty of Galactic calibration <14% (<77 MHz)
- spectral slope in agreement with CoREAS sim

Upgrade of the Pierre Auger Observatory
(astro-)physics of the highest-energy particles in nature



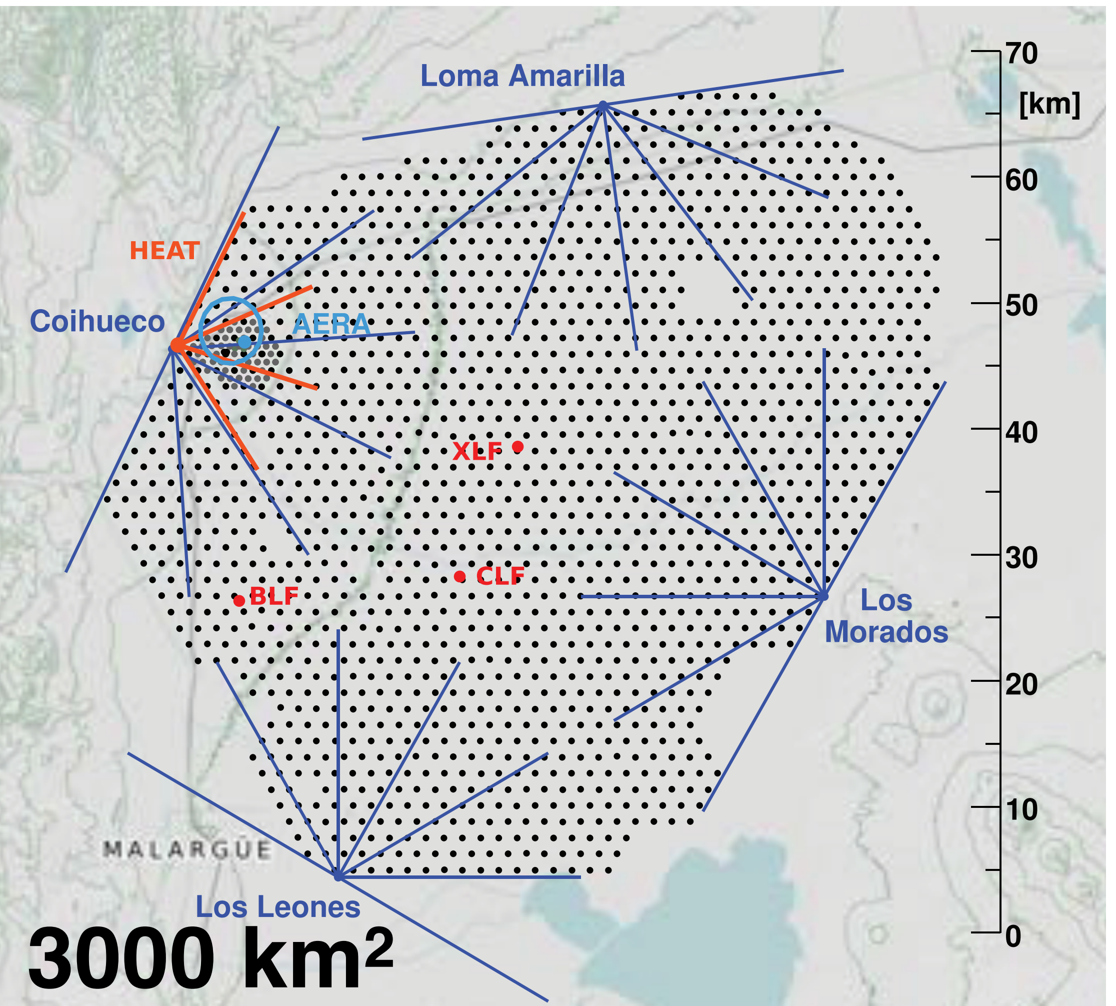
upgrade PAO

- electronics
- scintillator layer
- radio detector

in practice: different response to both components in both detectors: response matrix

Key science questions

- What are the **sources** and **acceleration** mechanisms of ultra-high-energy cosmic rays (UHECRs)?
- Do we understand **particle** acceleration and **physics** at energies well beyond the LHC (Large Hadron Collider) scale?
- What is the fraction of **protons**, **photons**, and **neutrinos** in cosmic rays at the highest energies?



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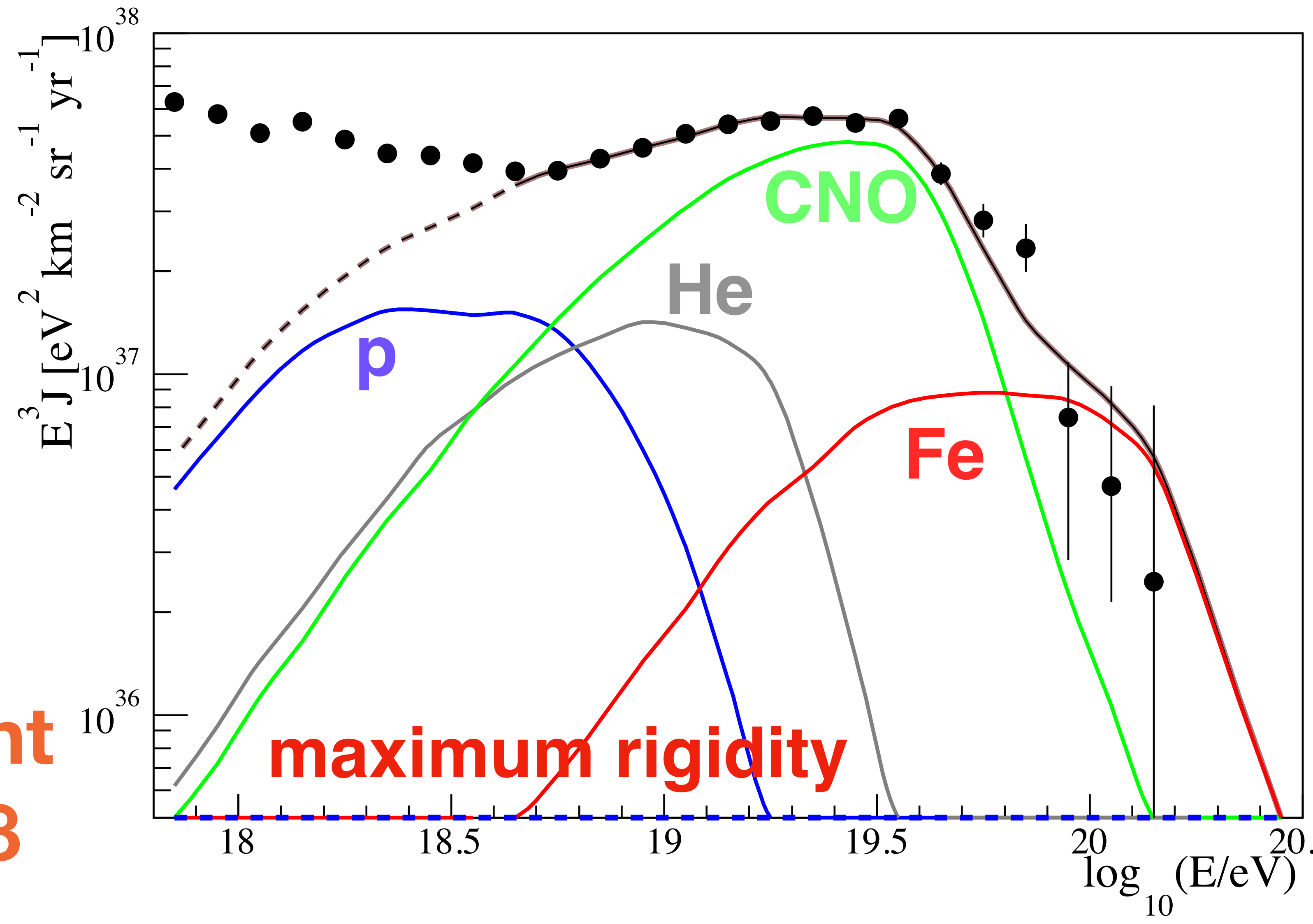
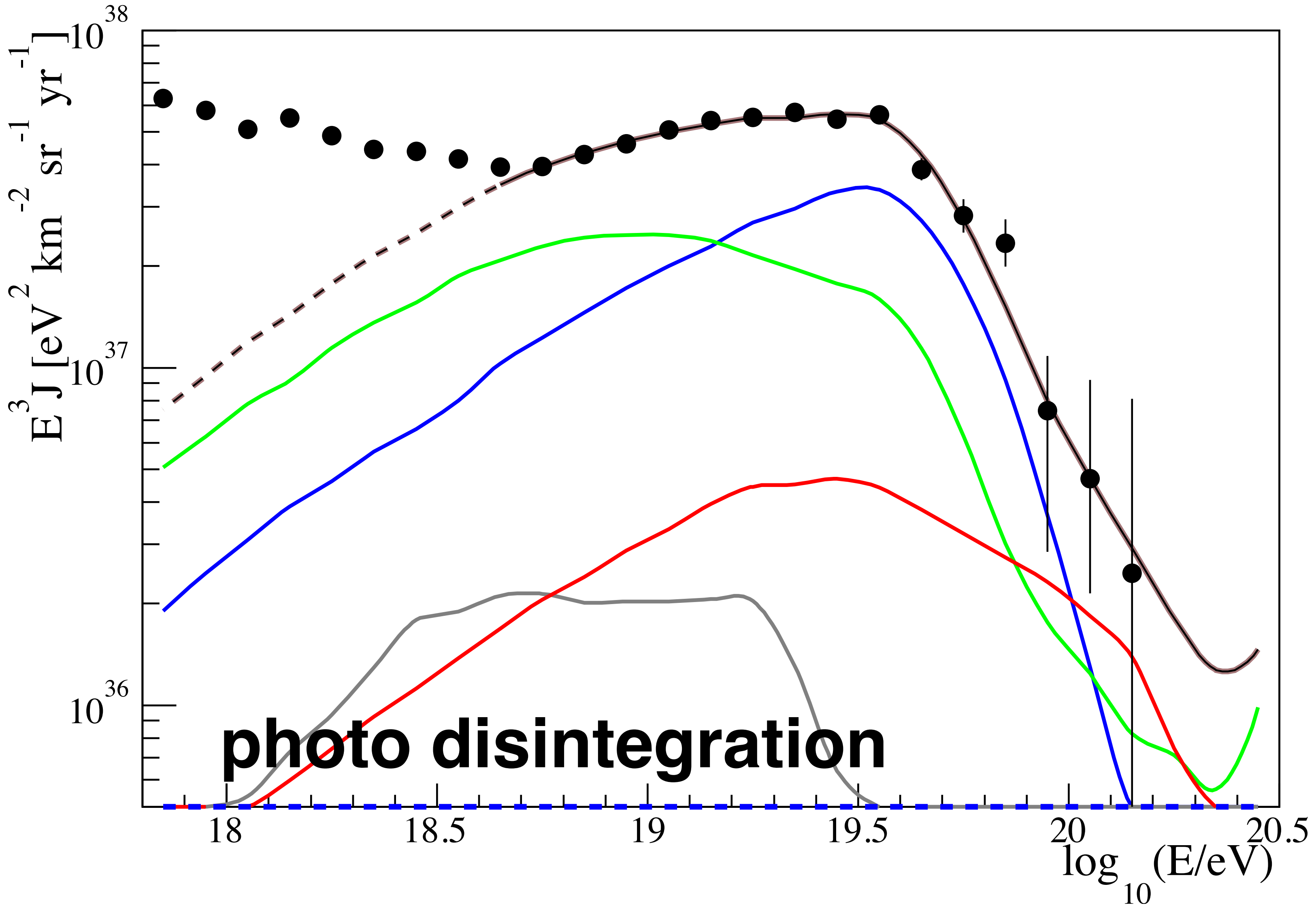
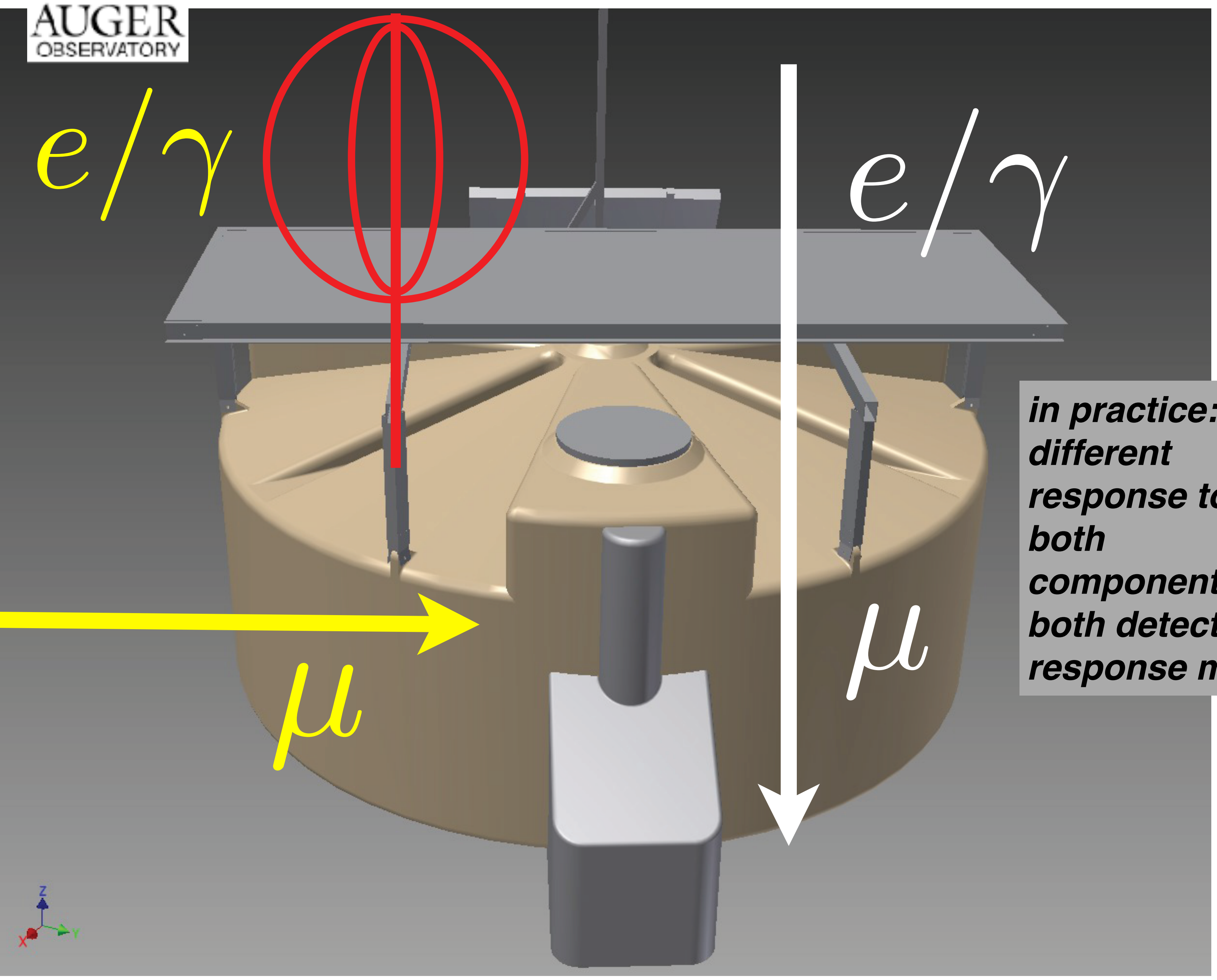
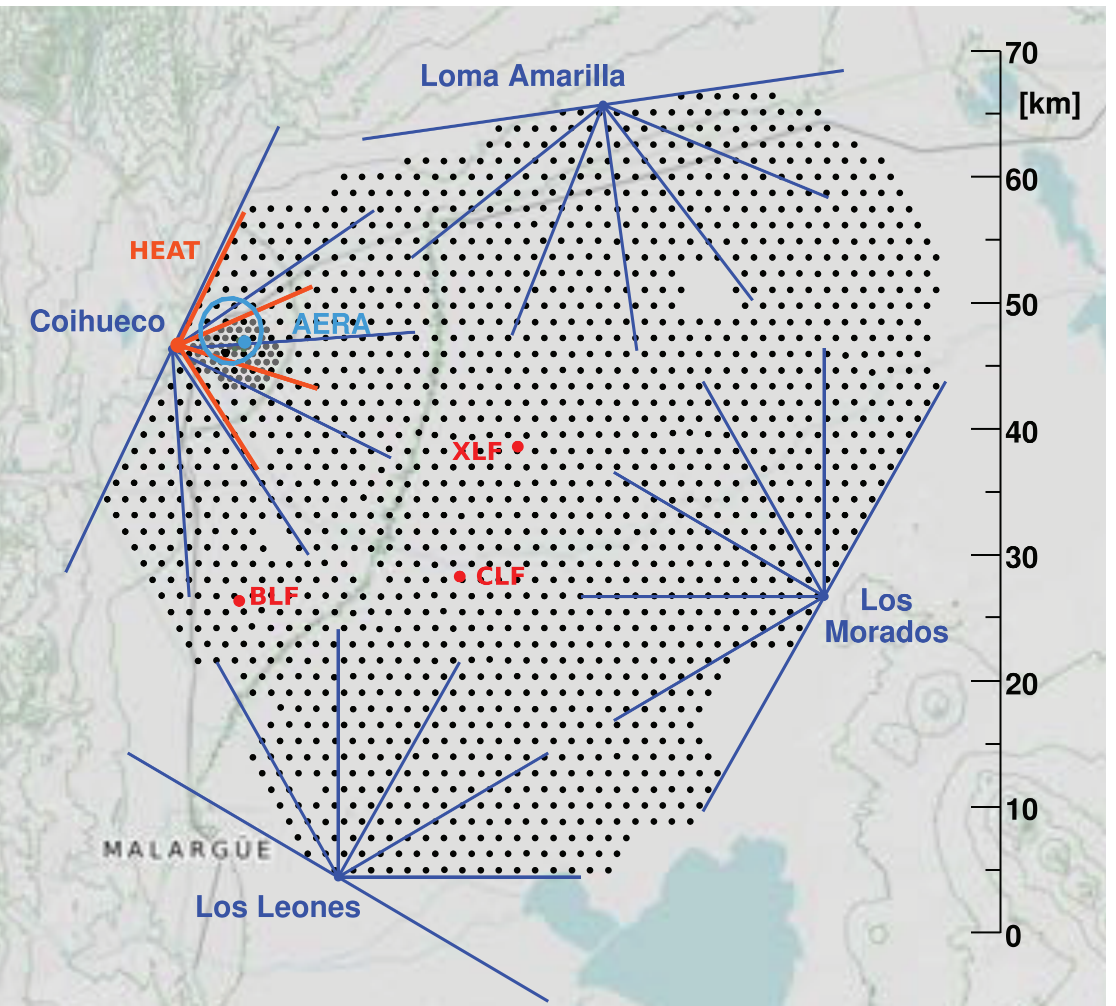




A large radio array the Pierre Auger Observatory



in practice: different response to both components in both detectors: response matrix



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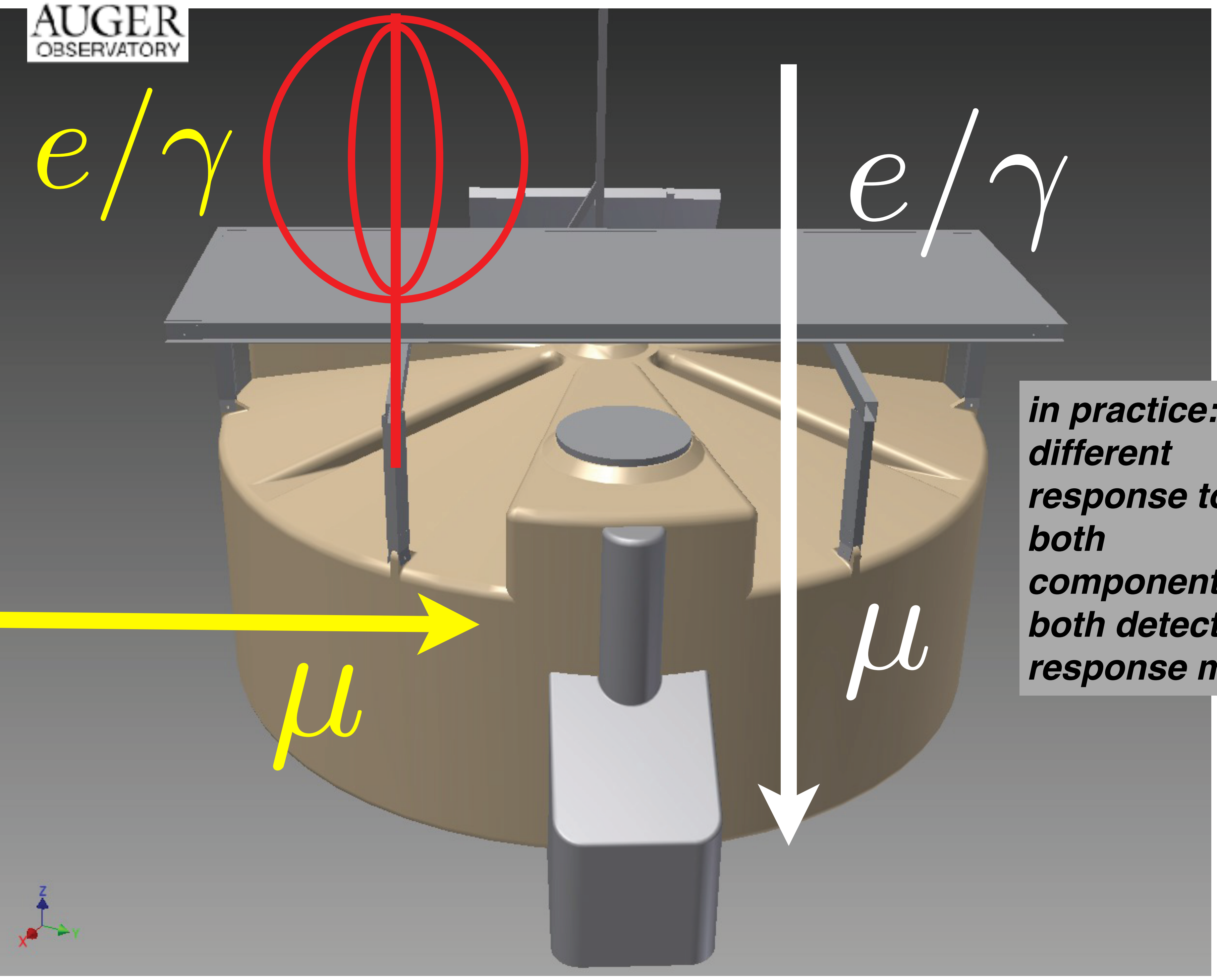
objective

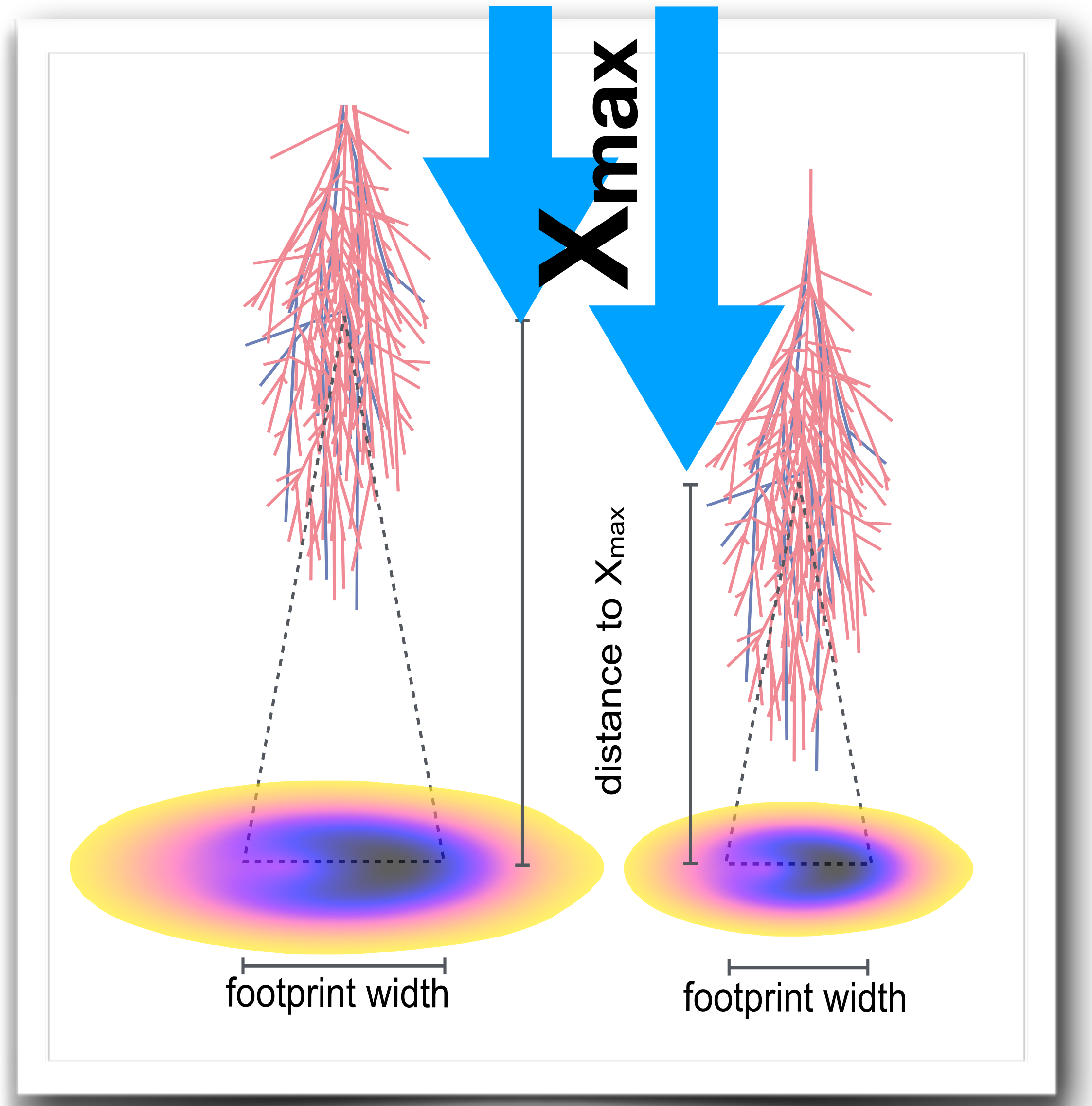
- origin of cosmic rays
- type of particle up to highest energies
- isolate protons, photons, neutrinos
- extend e/m-muon separation to high zenith angles
  --> horizontal air showers
  (i.e. increase exposure of SSD analyses)
- increase the sky coverage/overlap with TA
- absolute energy calibration from 1st principles
- independent mass scale
- clean e/m measurement
  --> shower physics

A large radio array the Pierre Auger Observatory



in practice: different response to both components in both detectors: response matrix

***attention*:**
type of particle determined

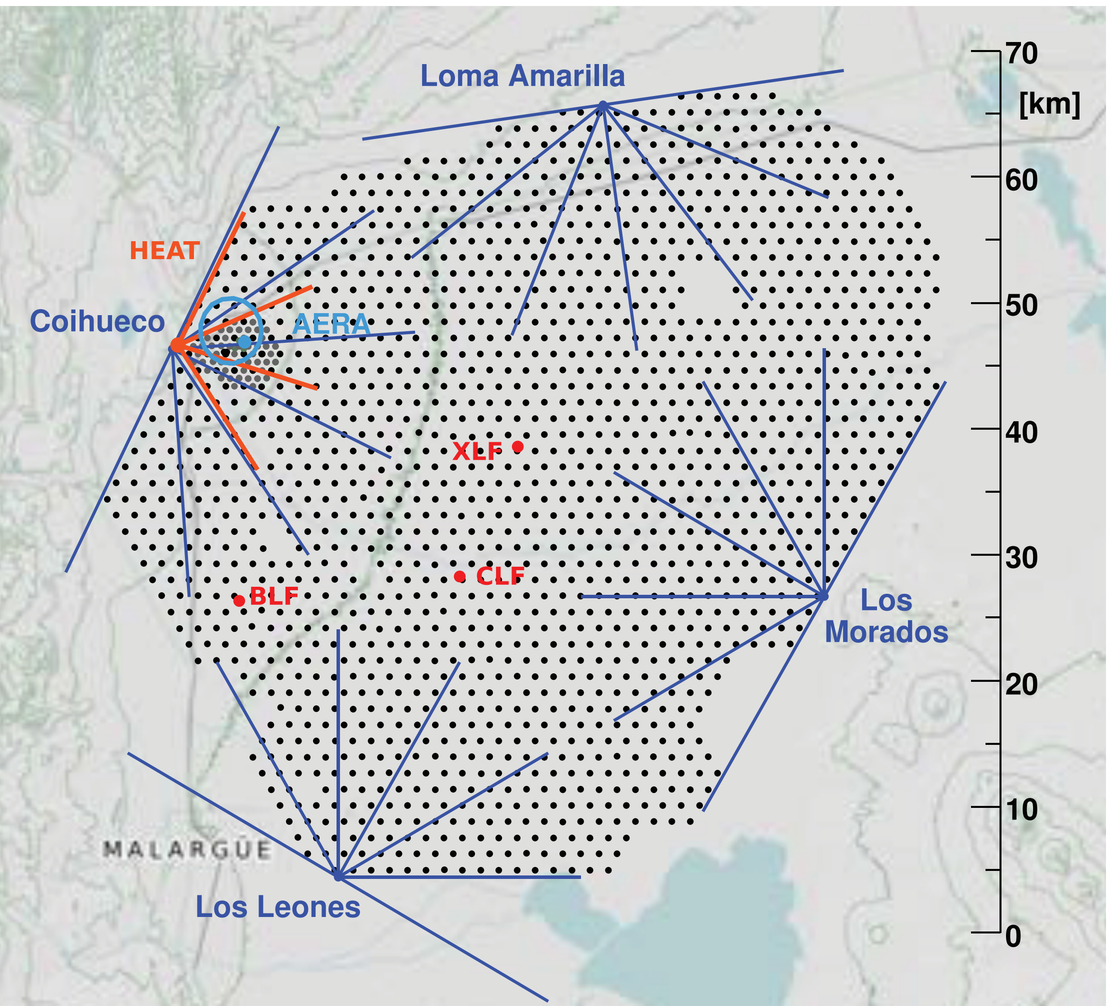


for vertical showers:

size of footprint
geometrical measurement

for horizontal showers:

electron/muon ratio
important: radio emission not absorbed in atmosphere



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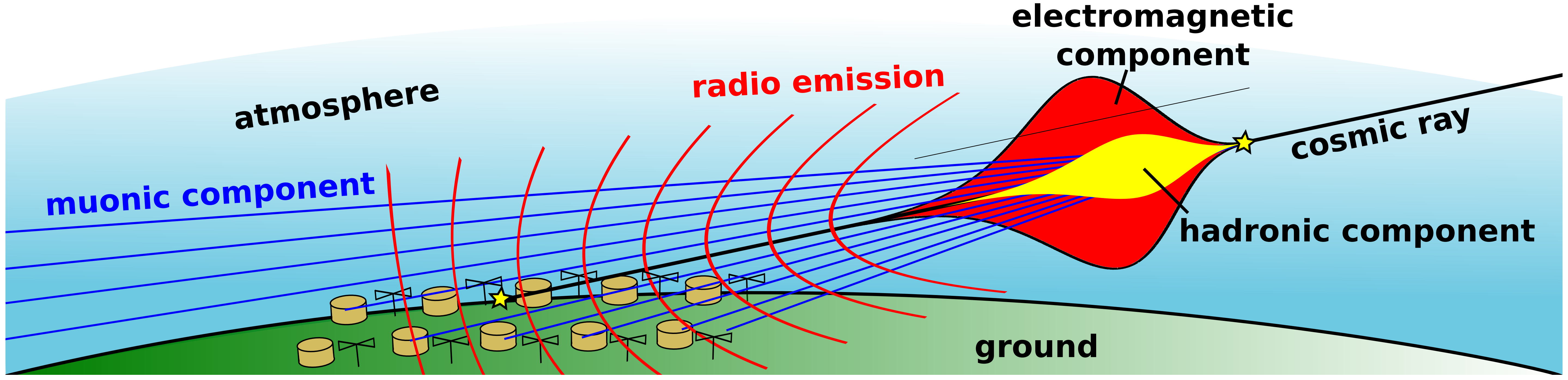


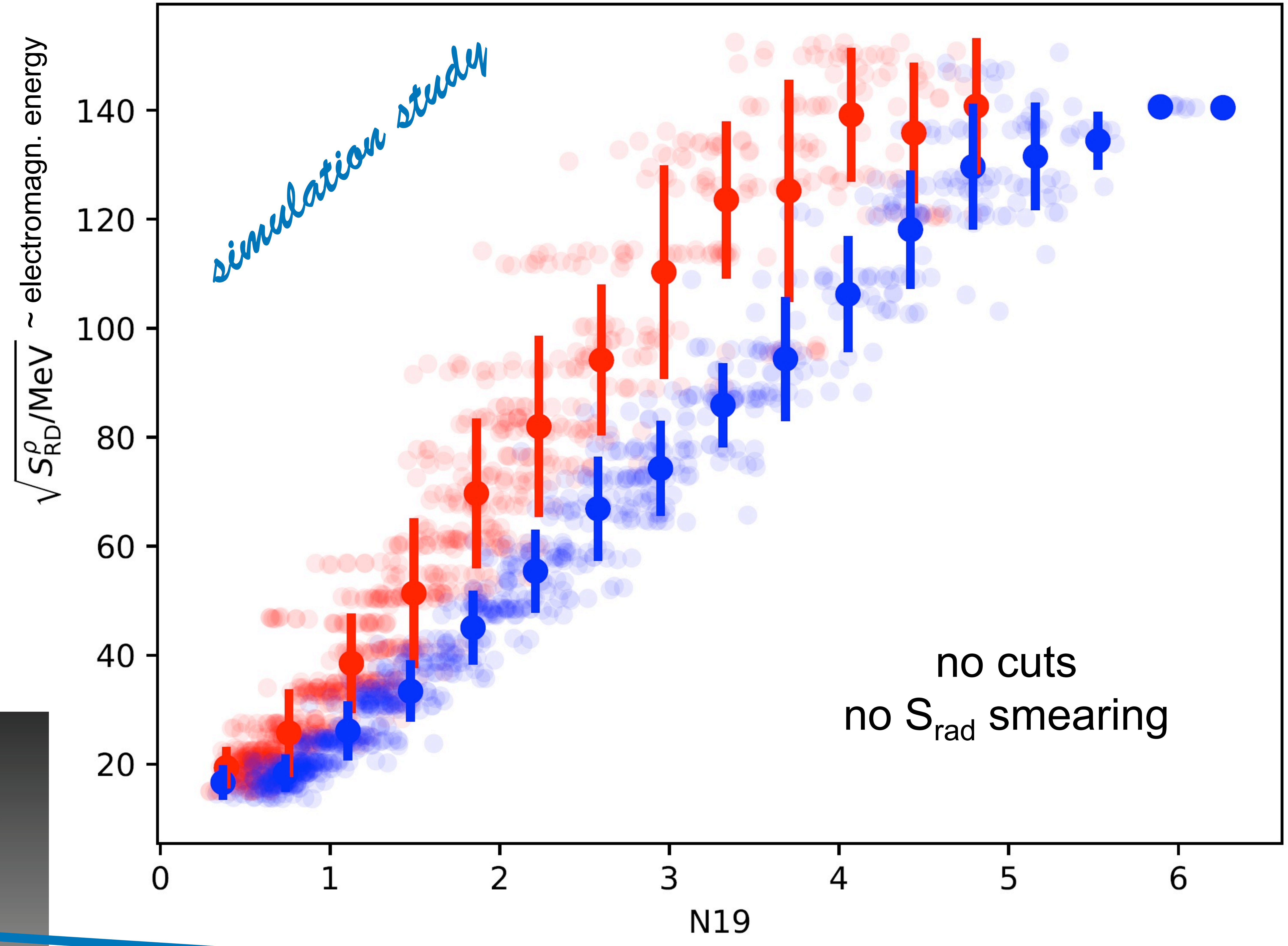
Radio detector provides good mass separation

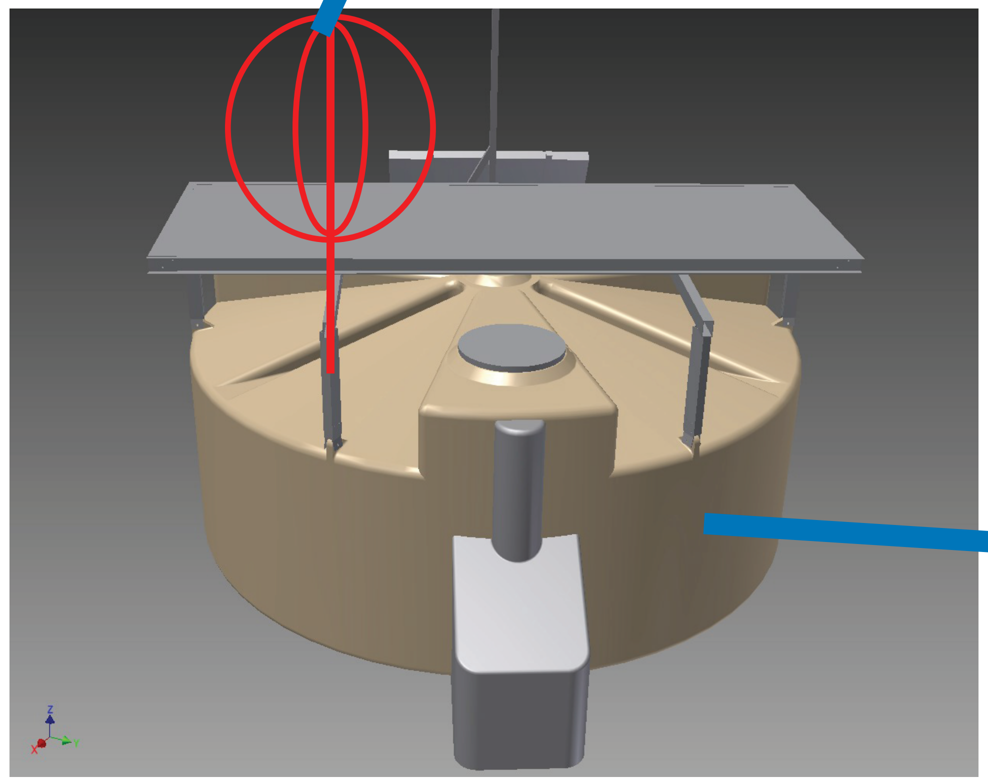
PIERRE AUGER OBSERVATORY

ARENA

J.R. Hörandel

radio detector (e/m)





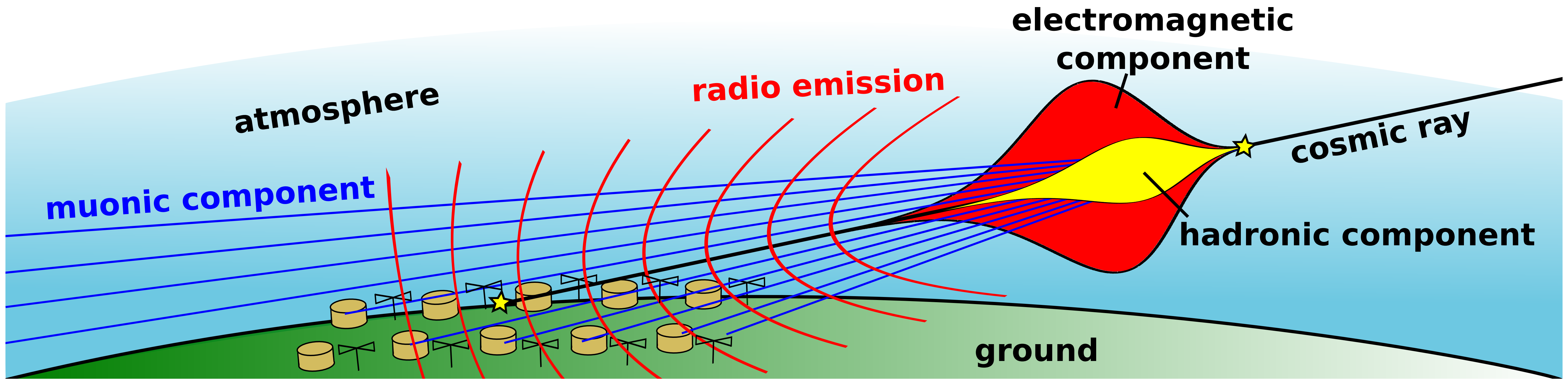
water Cherenkov detector (μ)

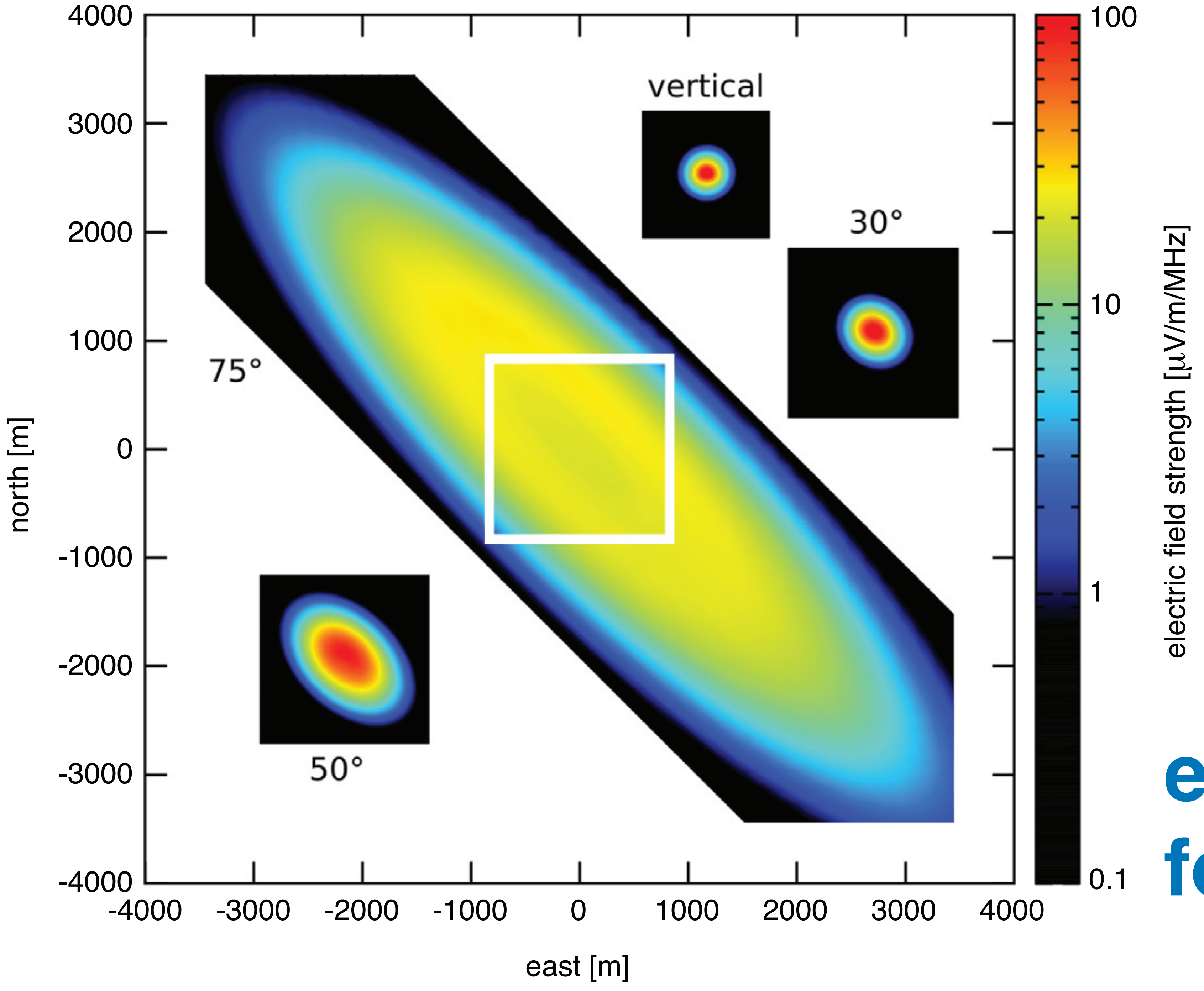
- can separate species with S_{rad} and N19
- separation increases with energy
- scaling at highest energies probably artifact of maximum simulated energy

A large radio array at the Pierre Auger Observatory

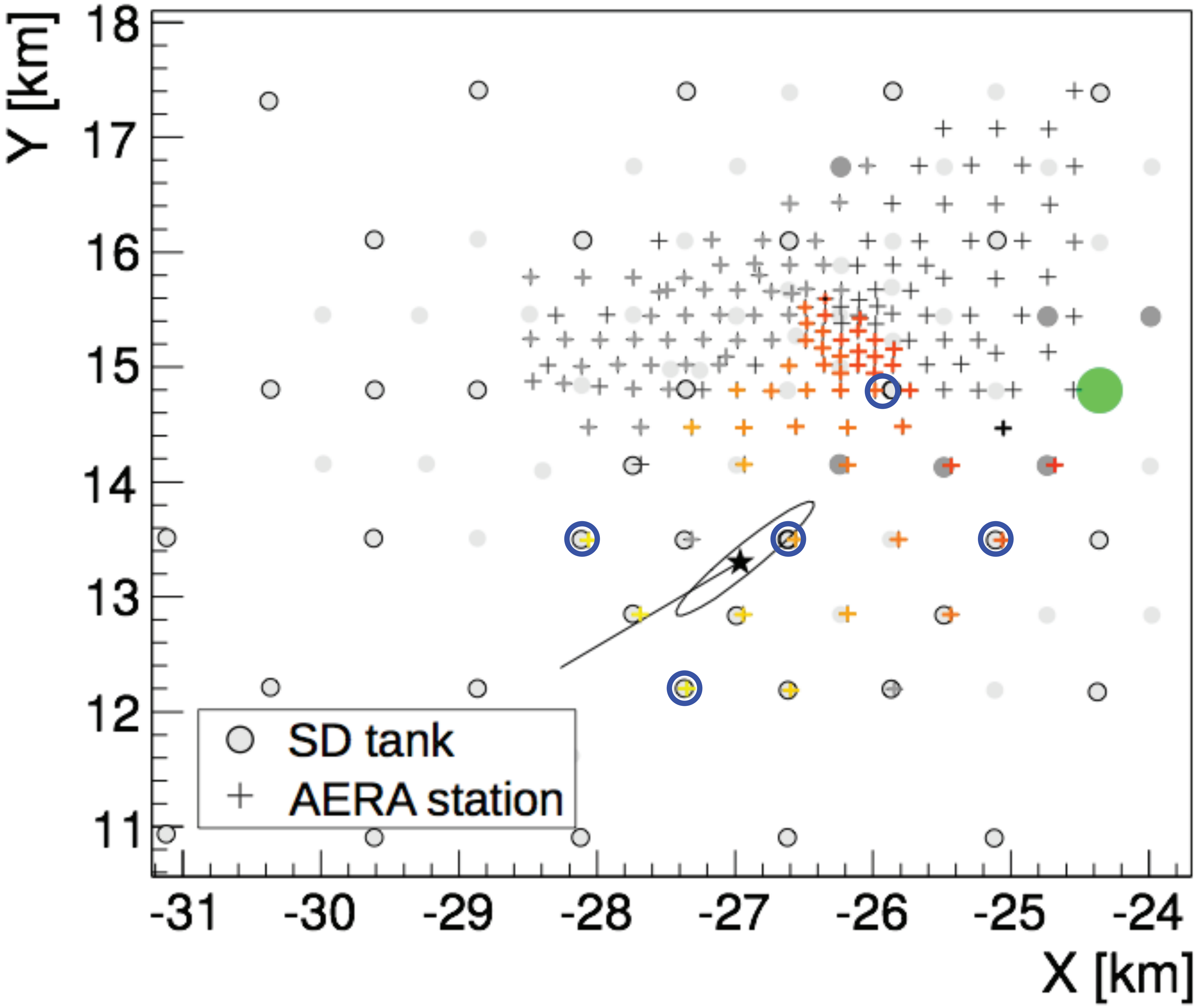
preparatory work & feasibility

AERA 17 km² --> 3000 km²





expect large radio footprint from simulations

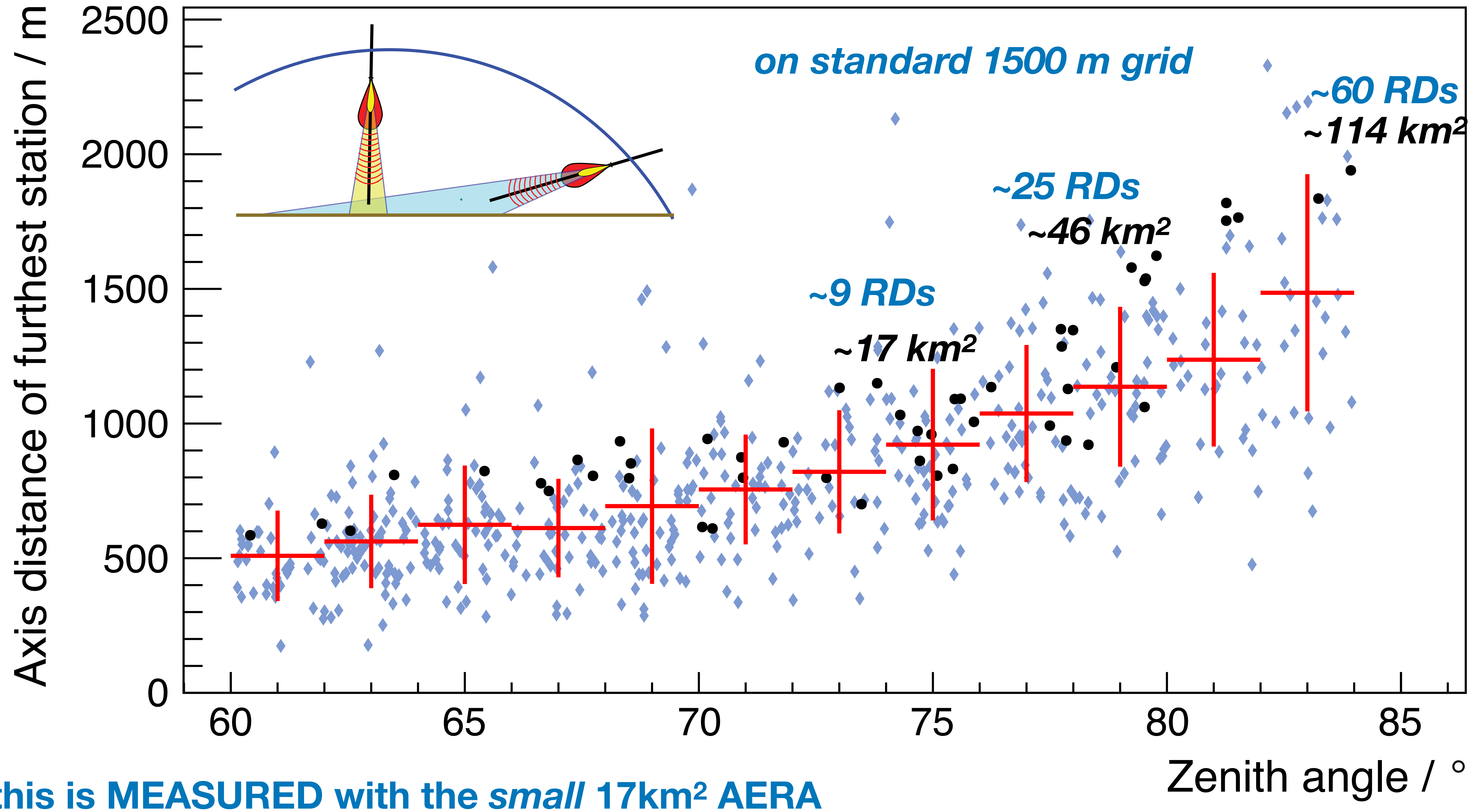


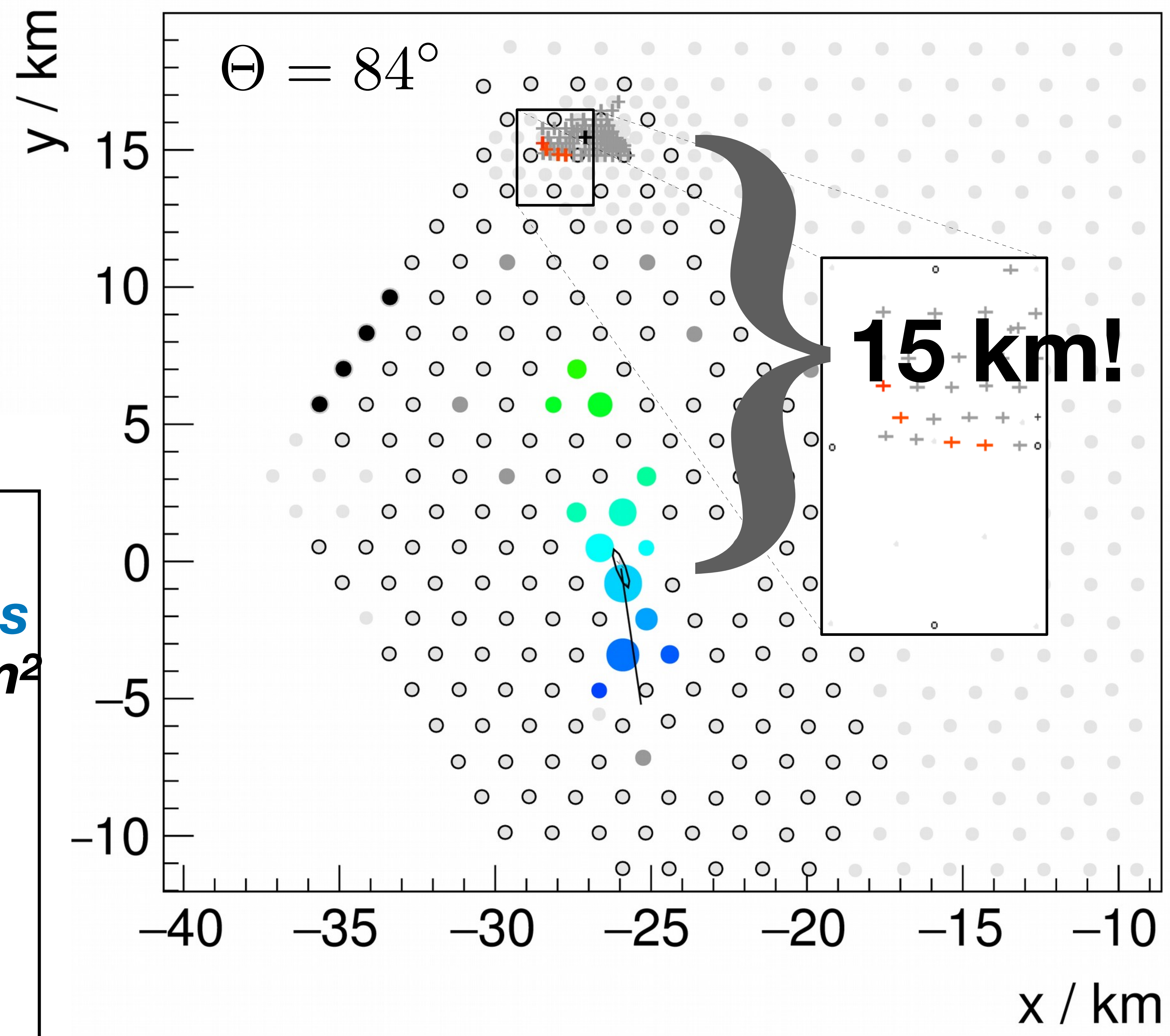
horizontal air showers registered and reconstructed with existing AERA

see e.g. T. Huege, Phys. Rep. 620 (2016) 1

Horizontal air showers have large footprints in radio emission

M. Gottowik

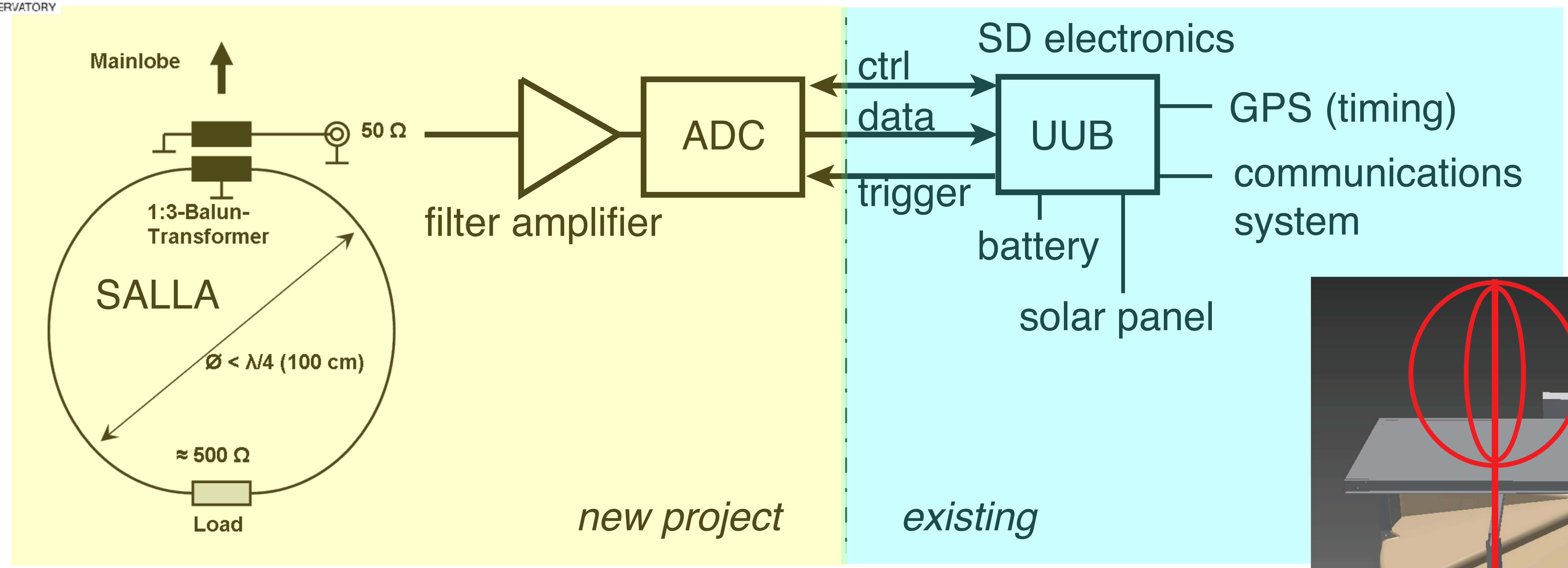


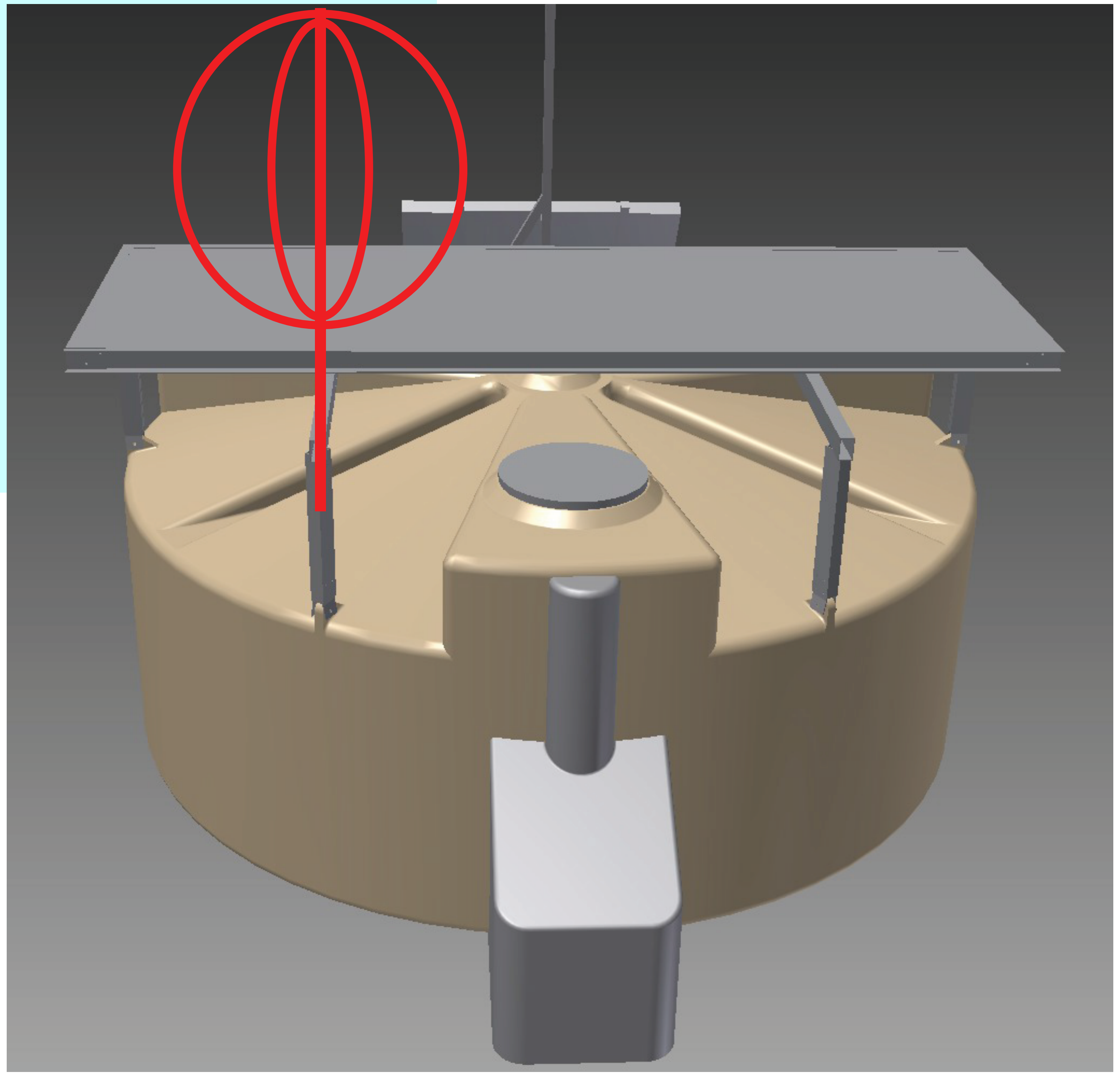


Pierre Auger coll. submitted (2018)
arXiv: 1806.05386

this is MEASURED with the *small* 17km² AERA

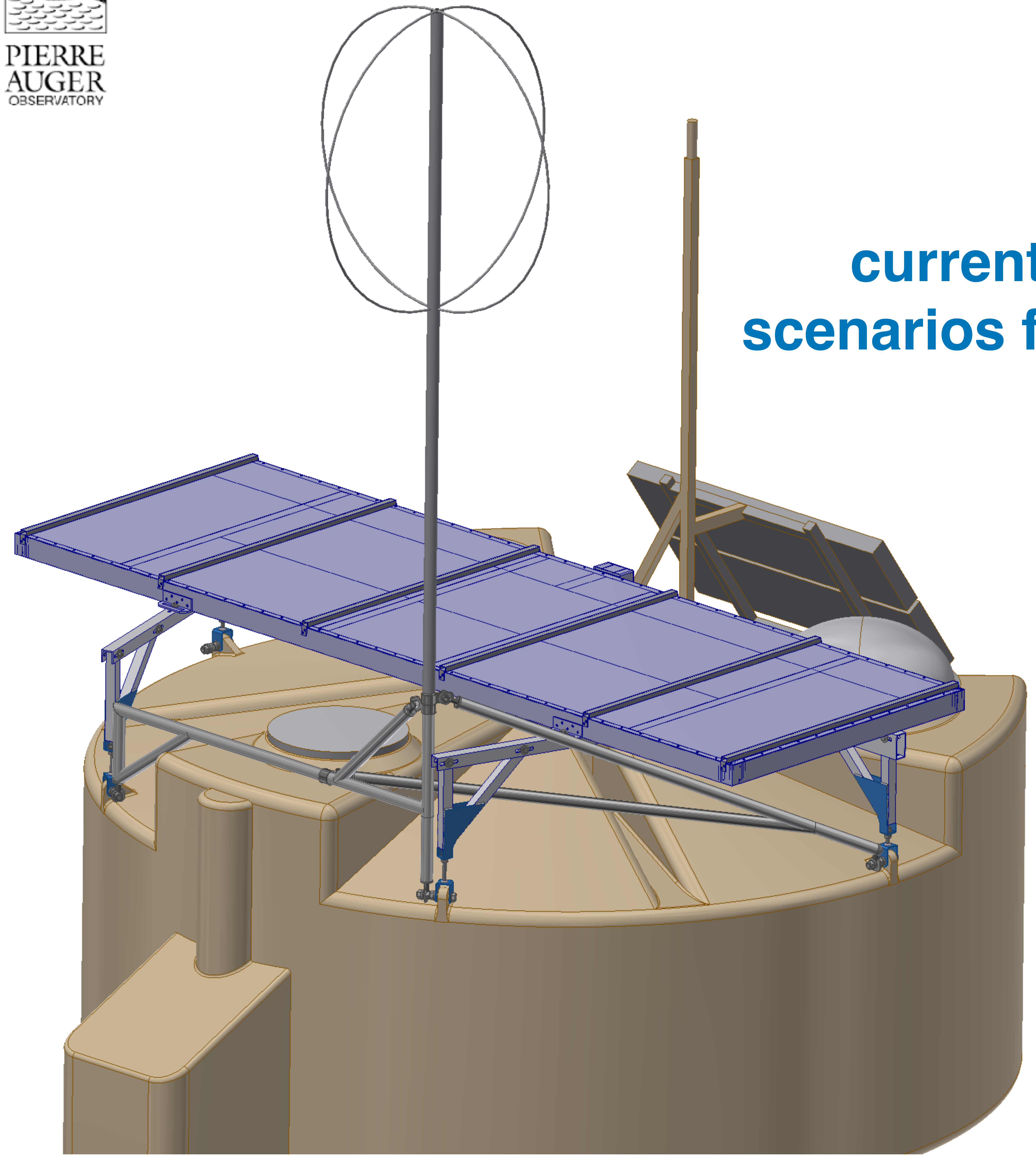
Integration of radio upgrade (RD), scintillator upgrade (SSD), and water Cherenkov detector in *ONE* unit



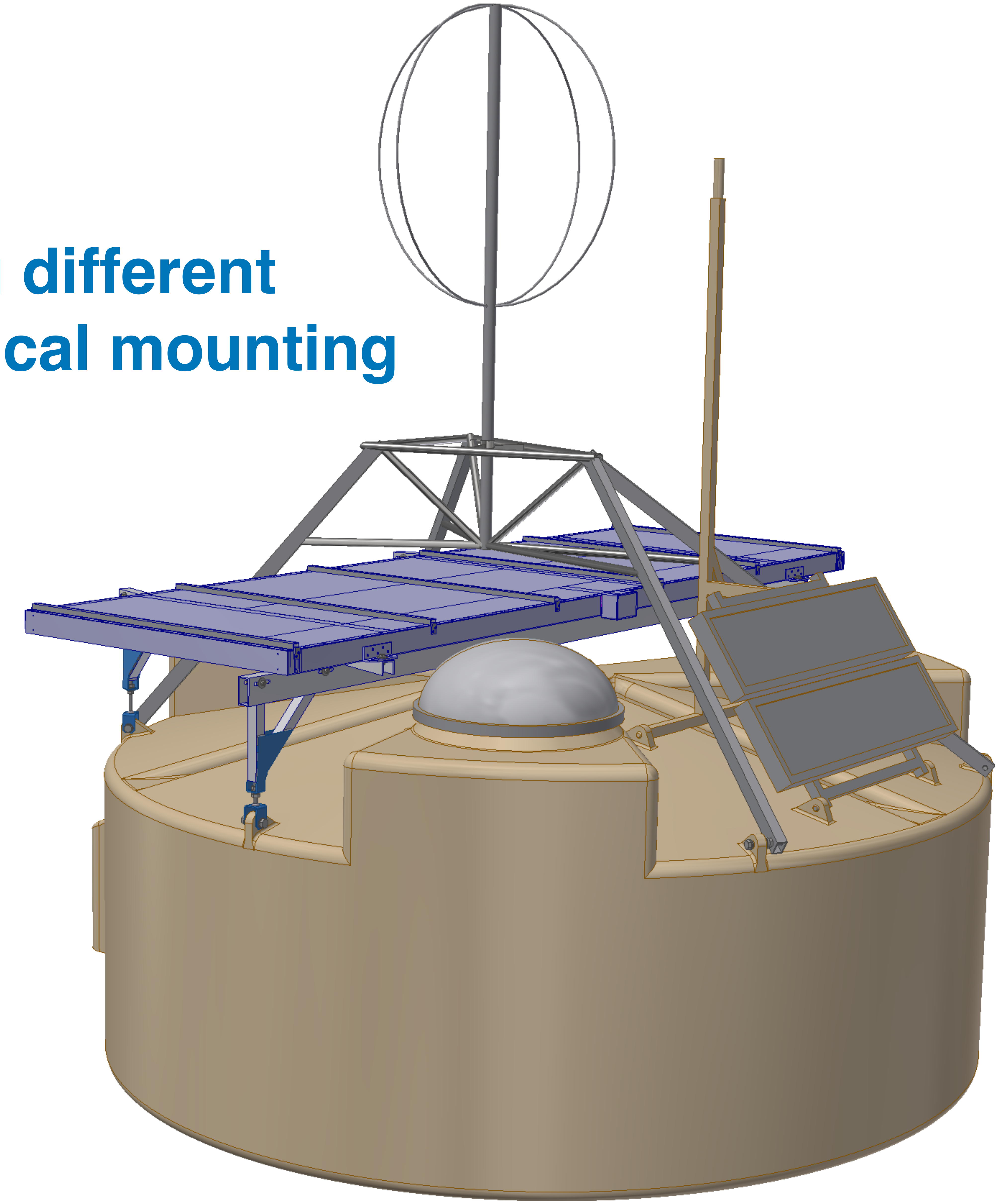


Shared infrastructure **(solar power, battery, GPS timing, communications system) and integrated data acquisition**

Antenna mounting



currently studying different
scenarios for mechanical mounting

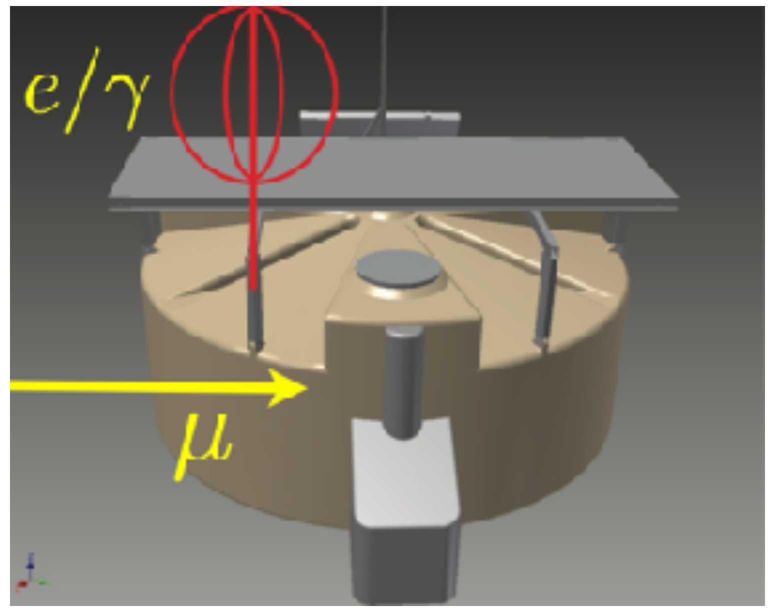


7th Roma International Conference on AstroParticle Physics

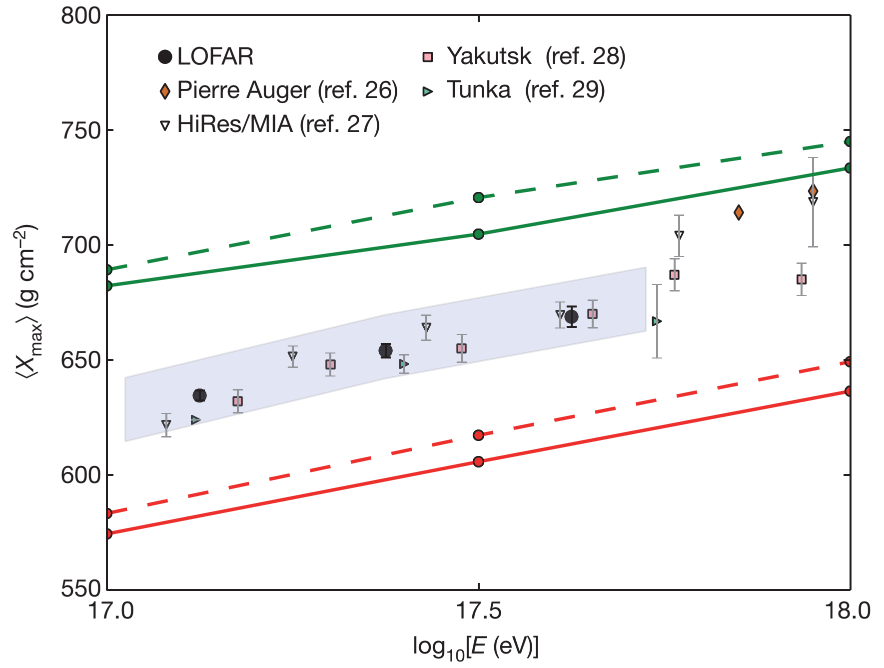
Status and perspectives of the radio detection of high-energy cosmic rays

huge progress in last decade

2018: beyond capabilities of standard installations

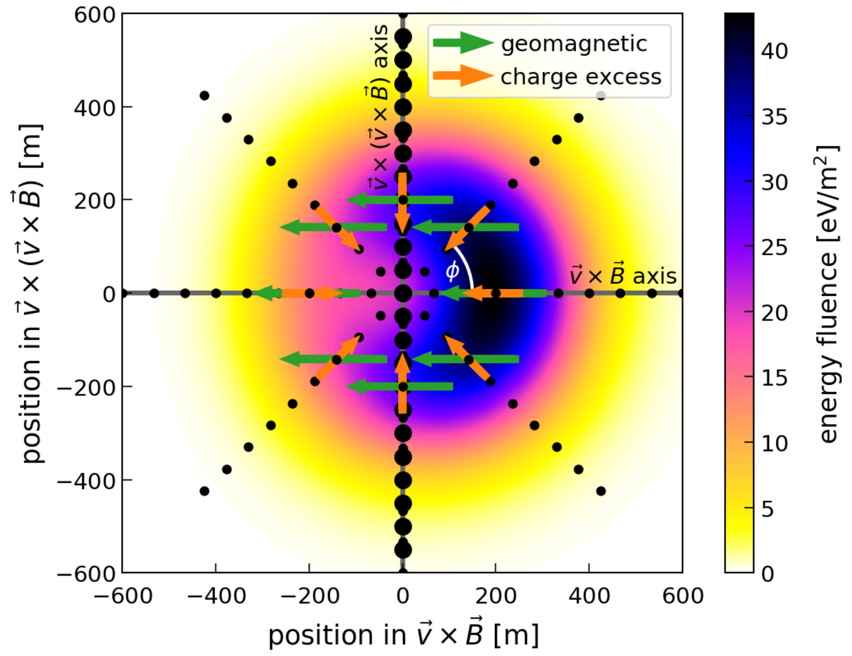


2016: radio technique mature: properties of cosmic rays



2014: understanding the emission processes

2013: CoREAS radio simulation in CORSIKA



2011: endpoint formalism

2005: understanding the radio signal

taskleader radio at Pierre Auger Observatory
PI LOFAR key science project Cosmic Rays
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Radboud University Nijmegen, Nikhef, VU Brussels
http://particle.astro.ru.nl