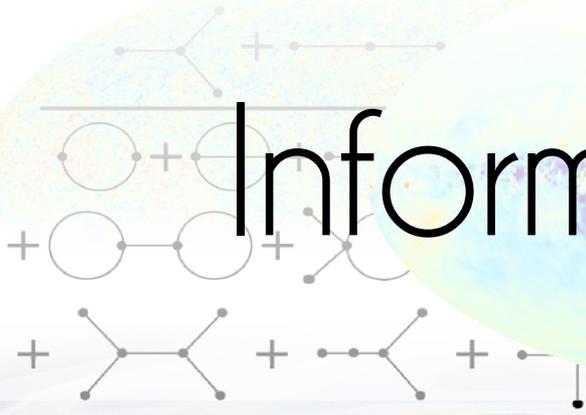
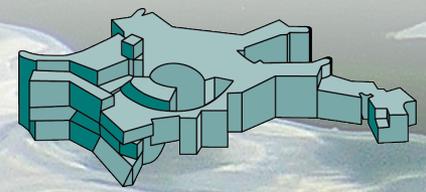


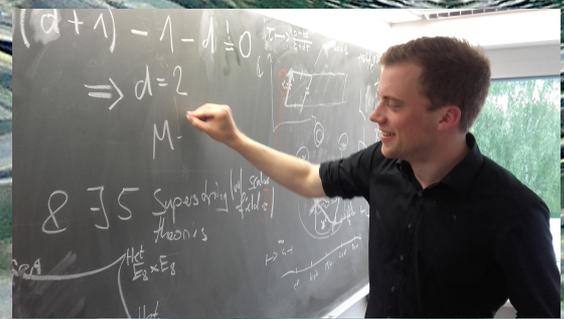
Information field theory



Torsten Enßlin
MPI for Astrophysics
Ludwig Maximilian University Munich



IFT Team: Philipp Arras, Michael Bell, Vanessa Böhm, Sebastian Dorn, Philipp Frank, Maksim Greiner, Sebastian Hutschenreuter, Henrik Junklewitz, Jakob Knollmüller, Reimar Leike, Marco Selig, Theo Steininger, Niels Oppermann, Natalia Porqueres, Daniel Pumpe, Martin Reinecke, Valentina Vacca, ...many more



Information Field Theory

signal field: ∞ degrees of freedom
data set: finite
→ additional information needed

information:
physical laws, symmetries, continuity,
statistical homogeneity/isotropy, ...

most important:
space, time, & fields
are continuous

combining concrete (data) &
abstract (knowledge) information
→ **information theory for fields**

Information Field Theory

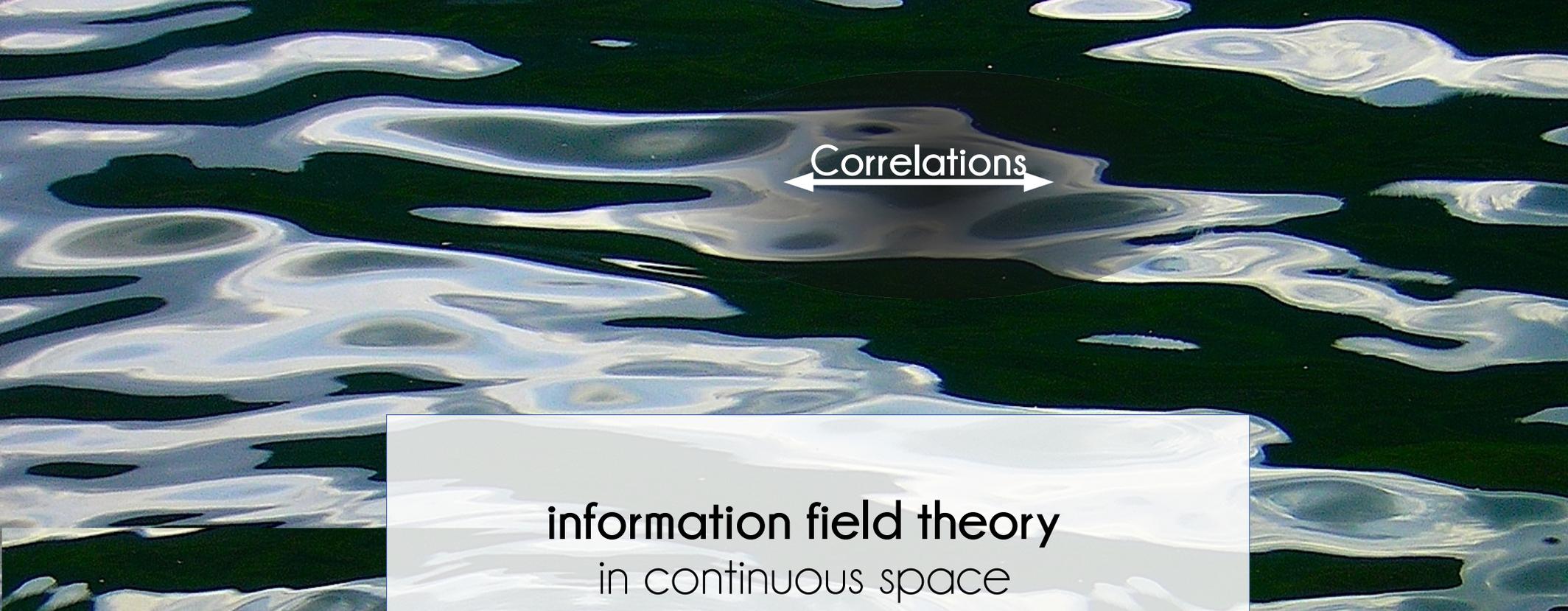
$s =$ signal

$d =$ data

$$\mathcal{P}(s|d) = \frac{\mathcal{P}(d|s) \mathcal{P}(s)}{\mathcal{P}(d)}$$

$$H(d, s) = -\log \mathcal{P}(d, s)$$

$$Z(d) = \int \mathcal{D}s \mathcal{P}(d, s)$$

A close-up photograph of water ripples on a dark surface, creating a complex, wavy pattern of light and dark green and blue. A white double-headed arrow is superimposed over the center of the image, with the word "Correlations" written above it.

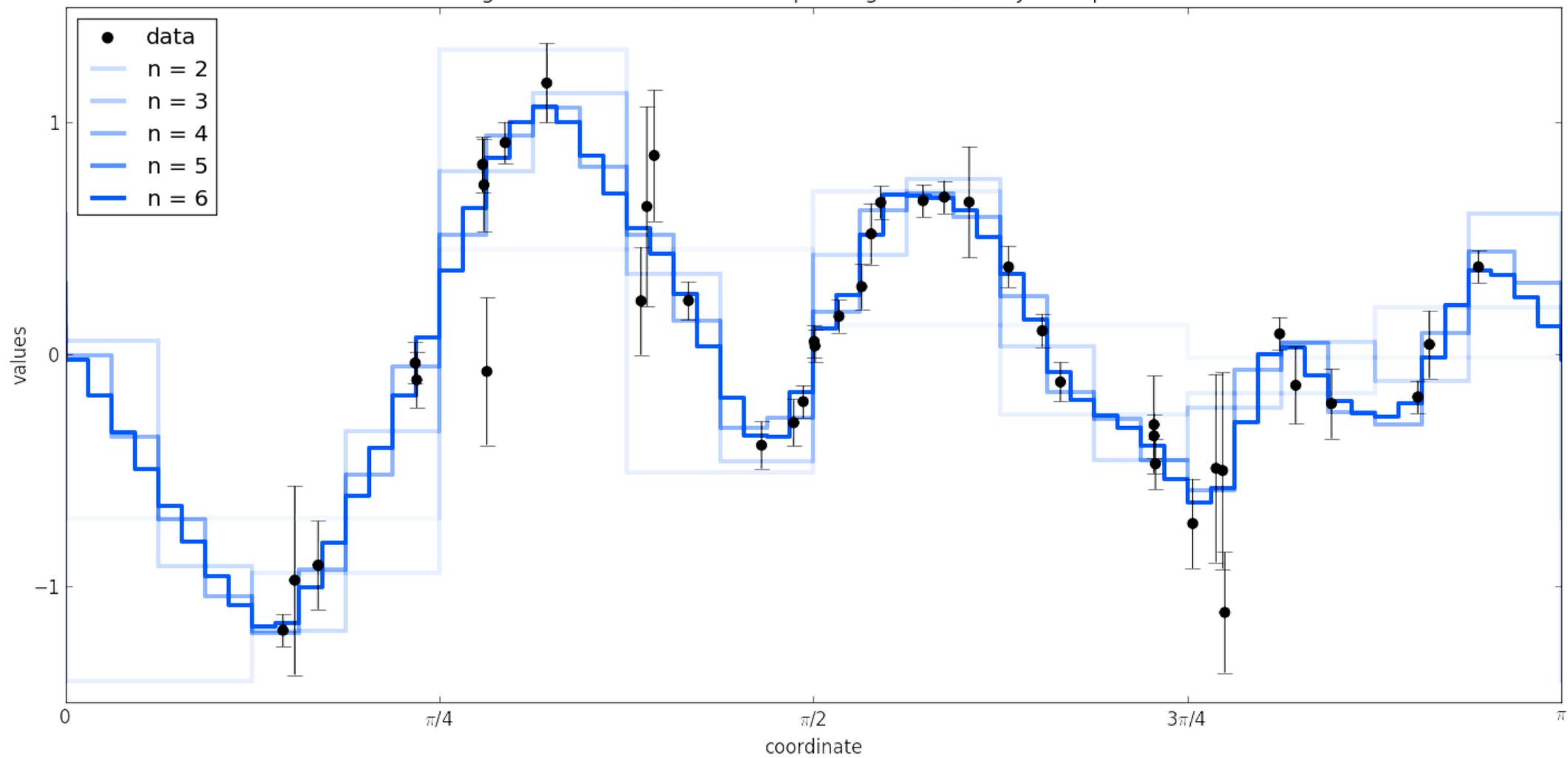
Correlations

A semi-transparent white rectangular box containing text, centered over the water ripples background.

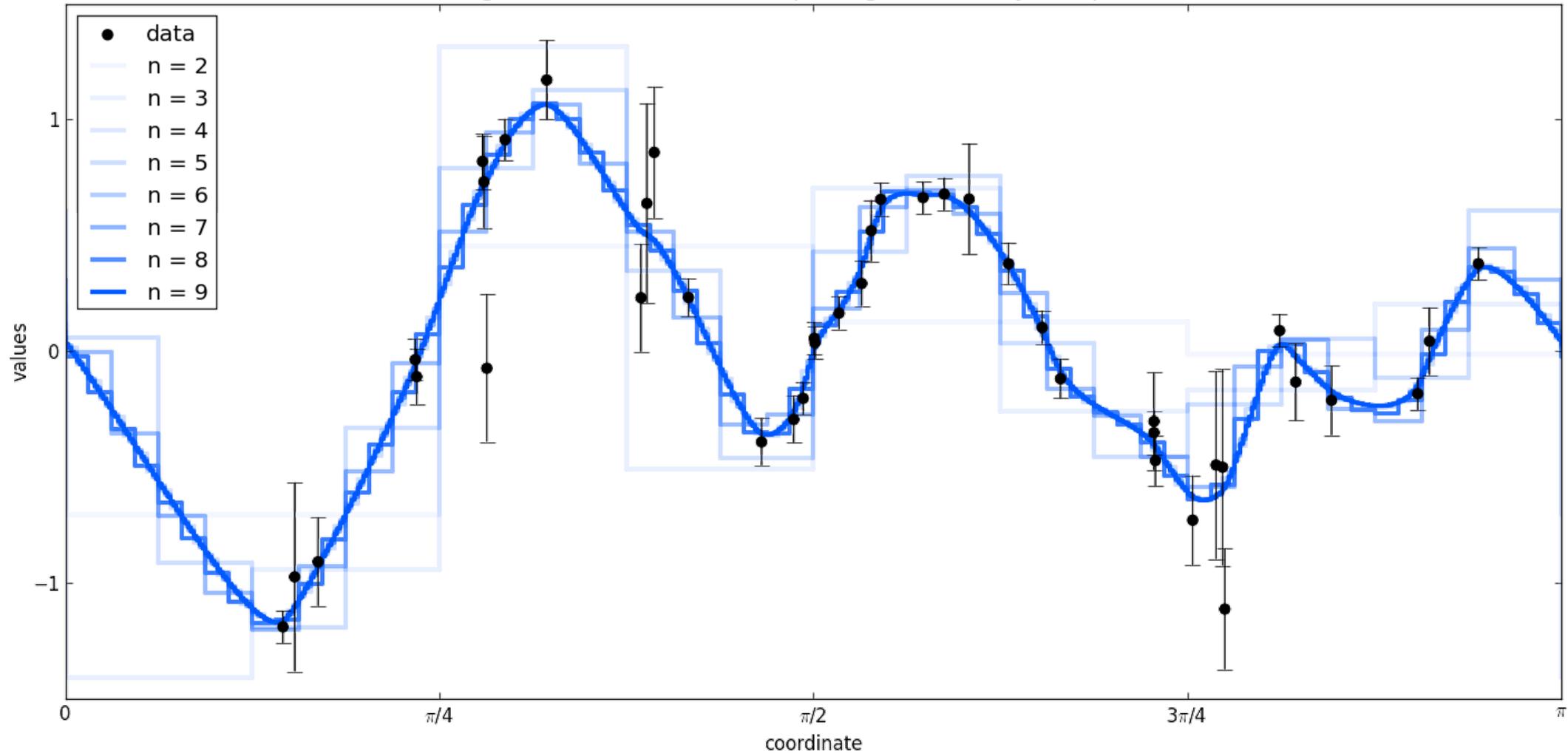
information field theory
in continuous space
calculations in discretized spaces



signal reconstruction with 2^n pixels given 42 noisy data points



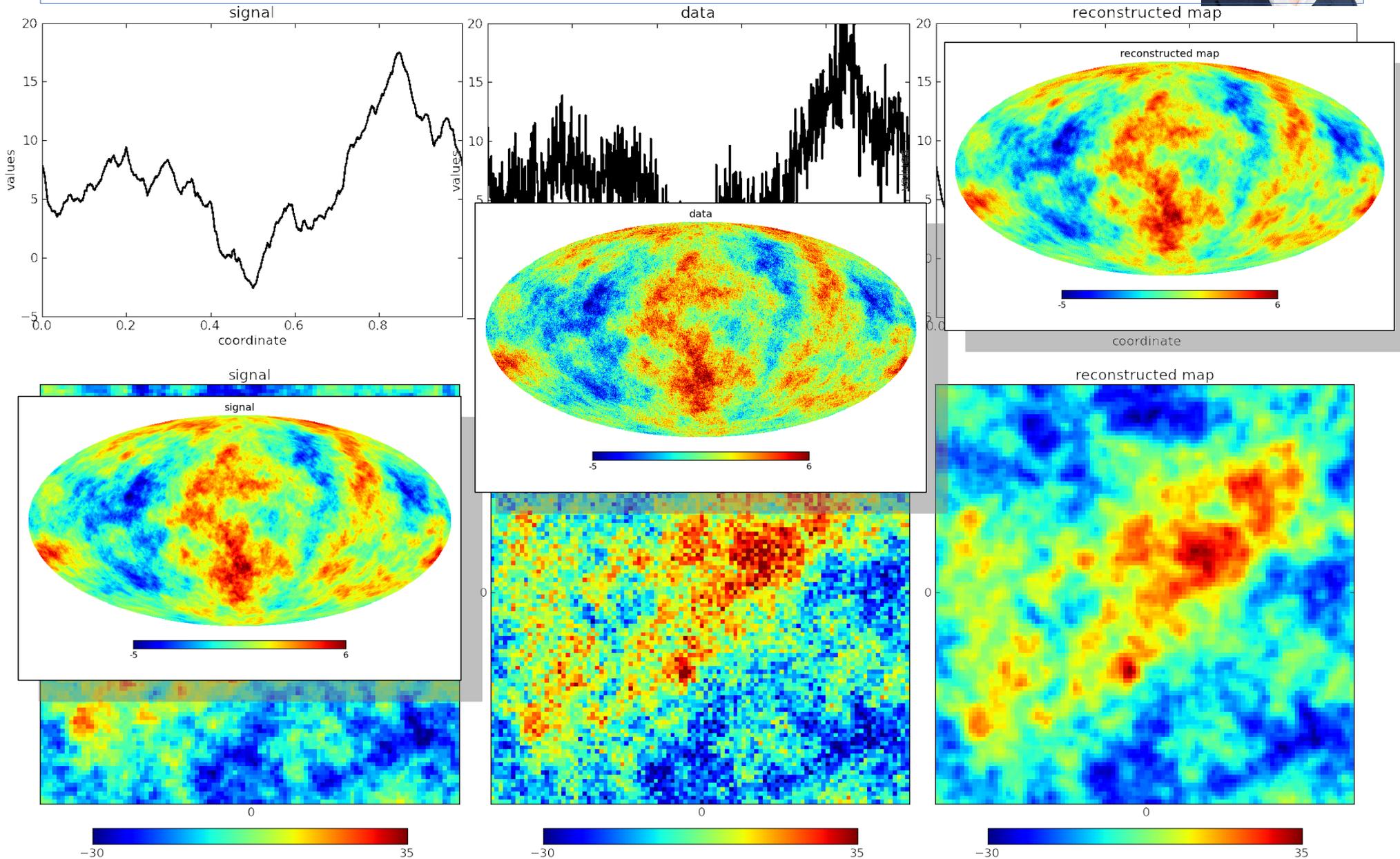
signal reconstruction with 2^n pixels given 42 noisy data points





NIFTY - Numerical Information Field Theory

Selig et al. (2013), Steininger et al. (sub.)
Code @ <https://gitlab.mpcdf.mpg.de/ift/NIFTy>



Extended VLA

LOFAR

RESOLVE

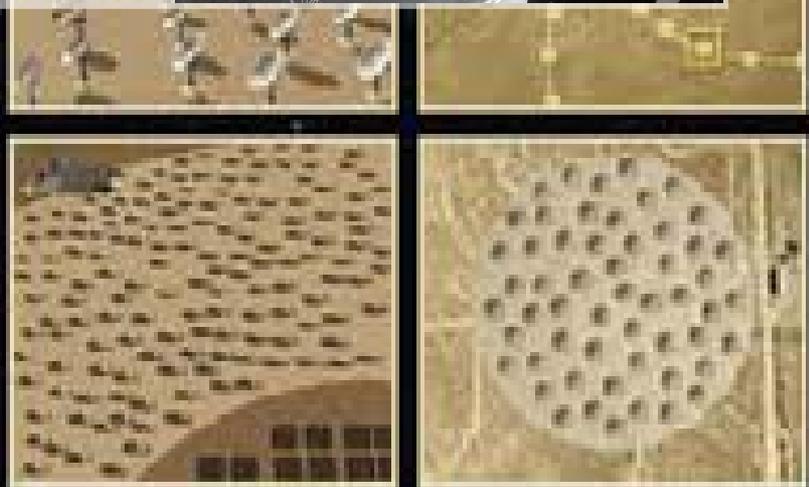
ALMA

Junklewitz et al.
Arras et al.

APERTIF

Meerkat

ASKAP



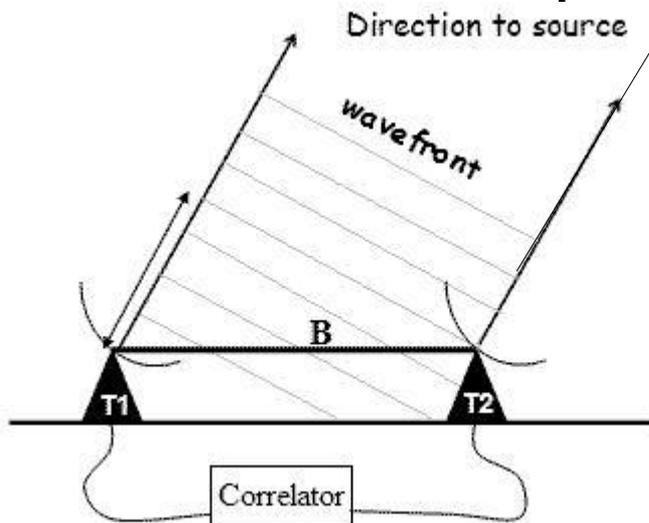
Data model

$$d = R e^s + n$$

known

unknown

known response



Gaussian priors:

$$\mathcal{P}(s) = \mathcal{G}(s, S) \quad \text{unknown}$$

$$\mathcal{P}(n) = \mathcal{G}(n, N) \quad \text{known}$$

Information Hamiltonian

$$\mathcal{H}(\mathbf{d}, \mathbf{s}, \boldsymbol{\tau}) = -\log \mathcal{P}(\mathbf{d}, \mathbf{s}, \boldsymbol{\tau})$$

likelihood

$$= \frac{1}{2} (\mathbf{d} - \mathbf{R}e^{\mathbf{s}})^{\dagger} \mathbf{N}^{-1} (\mathbf{d} - \mathbf{R}e^{\mathbf{s}}) + \frac{1}{2} \log (\det [\mathbf{N}])$$

prior /
regularization

$$+ \frac{1}{2} \mathbf{s}^{\dagger} \mathbf{S}^{-1} \mathbf{s} + \frac{1}{2} \log (\det [\mathbf{S}])$$

hyper-prior /
intelligence

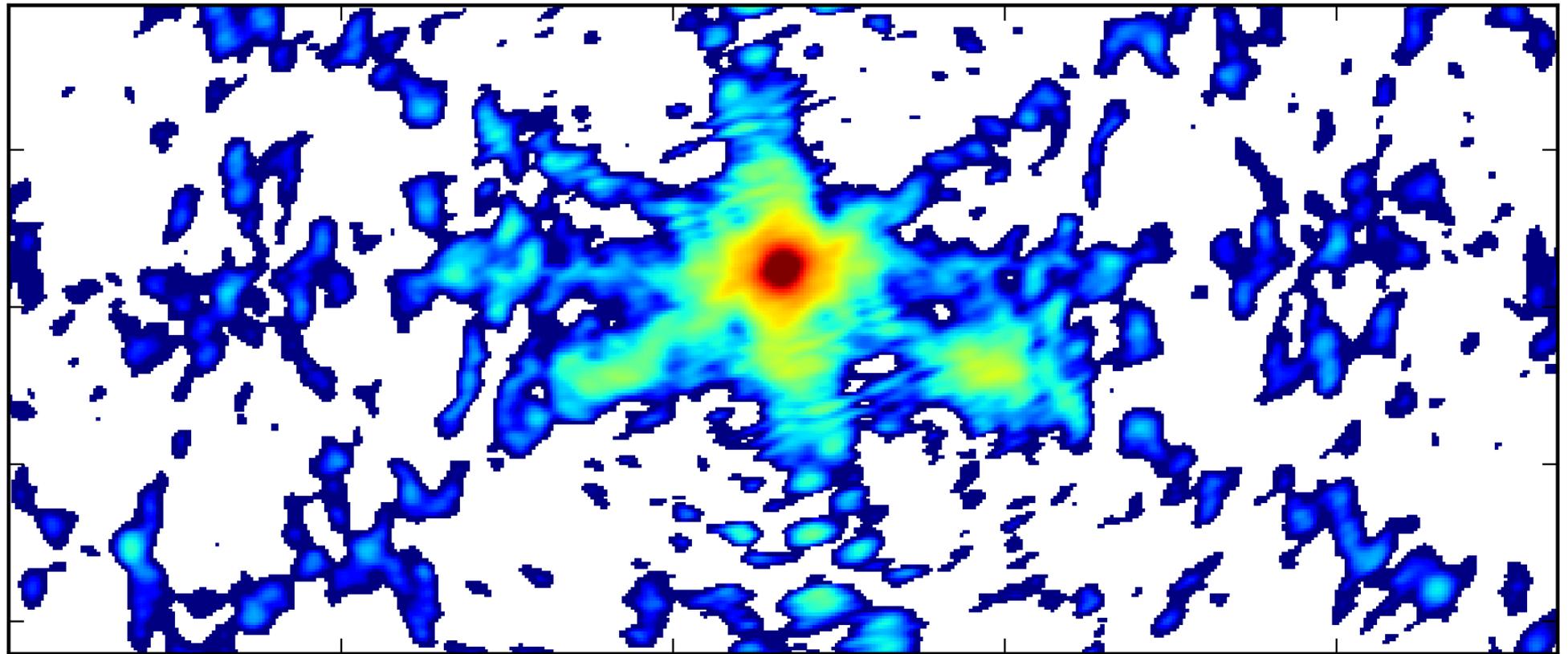
$$+ (\boldsymbol{\alpha} - \mathbf{1})^{\dagger} \boldsymbol{\tau} + \mathbf{q}^{\dagger} e^{-\boldsymbol{\tau}} + \frac{1}{2} \boldsymbol{\tau}^{\dagger} \mathbf{T} \boldsymbol{\tau}$$

correlation
structure

$$\mathbf{S} = \sum_k e^{\tau_k} \mathbf{S}_k$$

Abell 2219 @ 8415 MHz - data by Valentina Vacca

dirty 8415



16 28 42

16 28 39

16 28 36

16 28 33

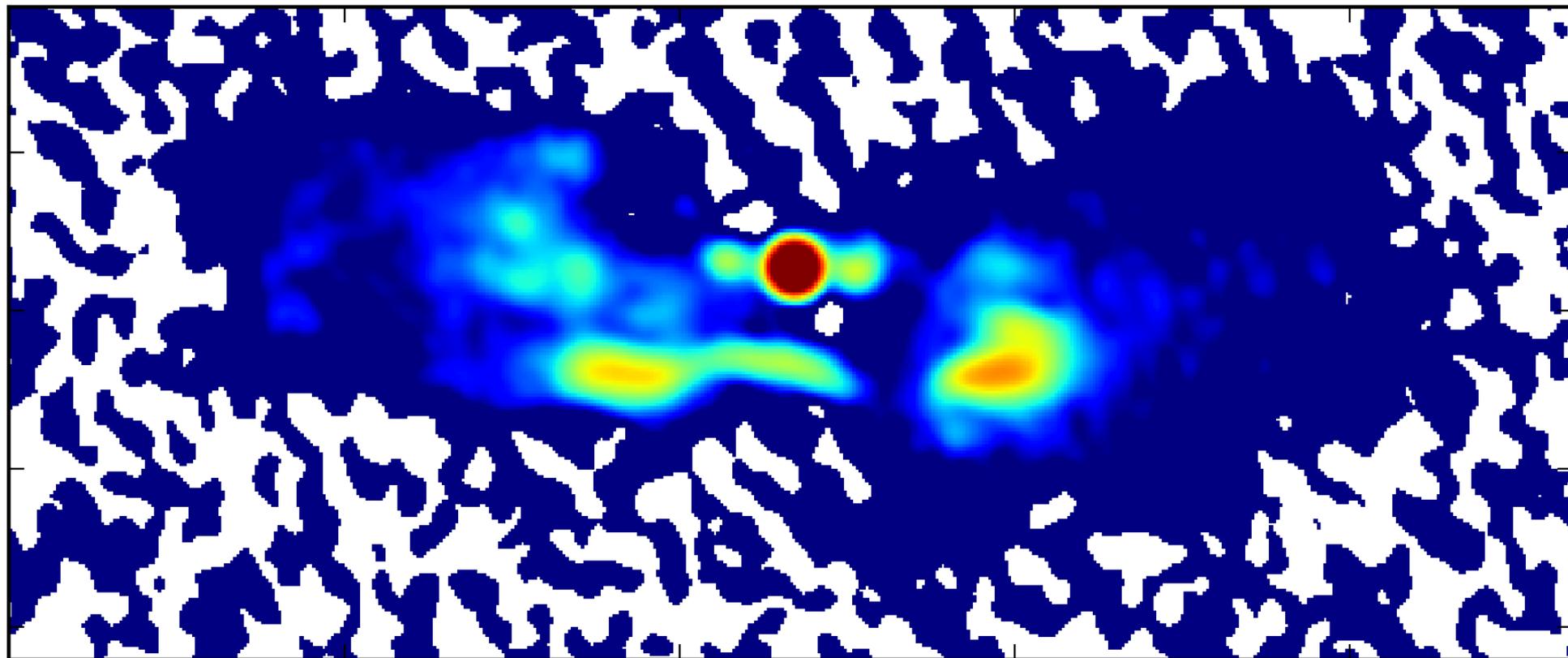


5e+05

5e+07

Abell 2219 @ 8415 MHz - CLEAN map by Valentina Vacca

CLEAN 8415



16 28 42

16 28 39

16 28 36

16 28 33

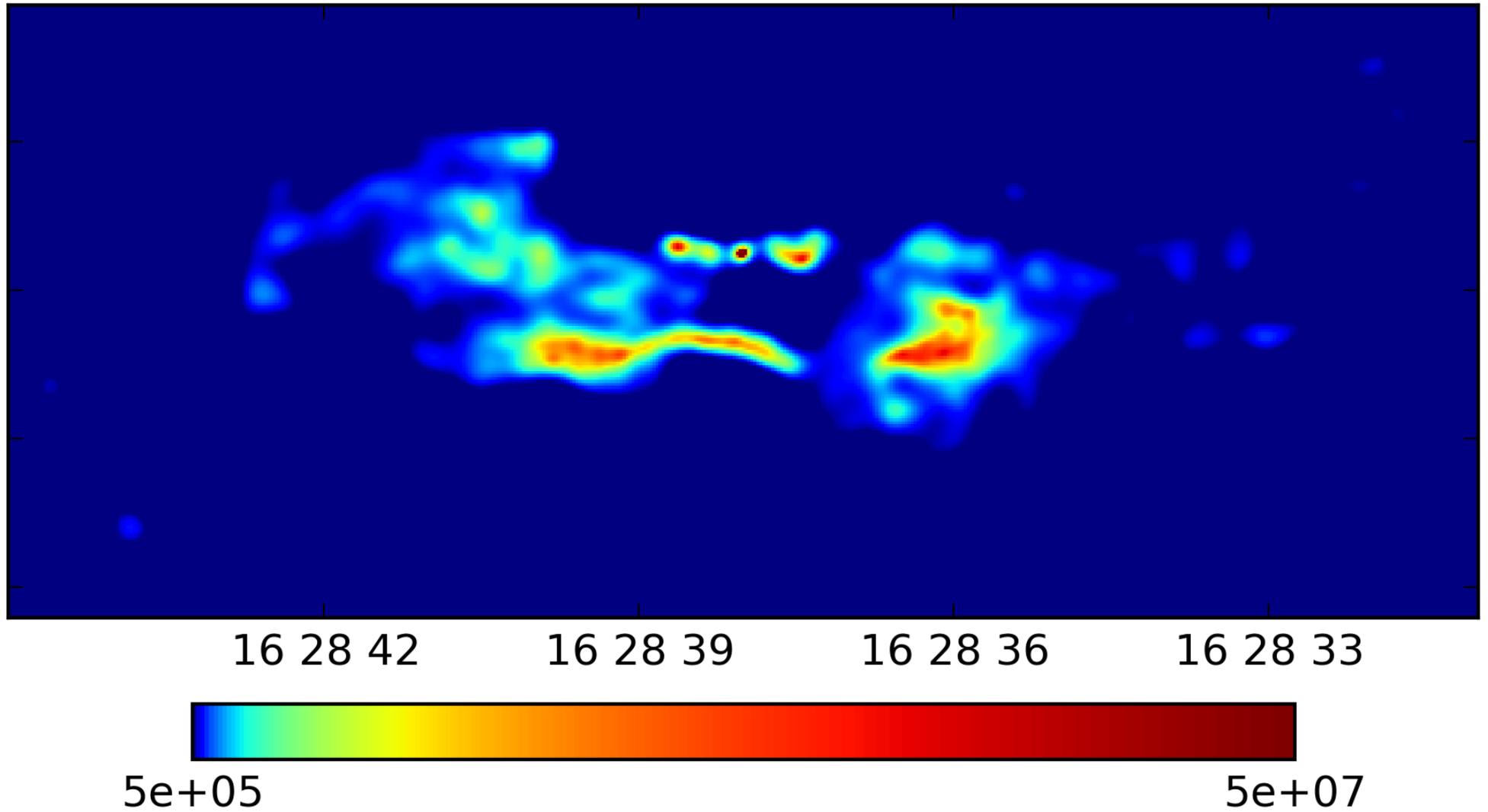


5e+05

5e+07

Abell 2219 @ 8415 MHz - **fast-RESOLVE** map by Maksim Greiner

RESOLVE 8415



IFT gang



Maksim Greiner
fastRESOLVE

Data fusion

single dish

interferometer

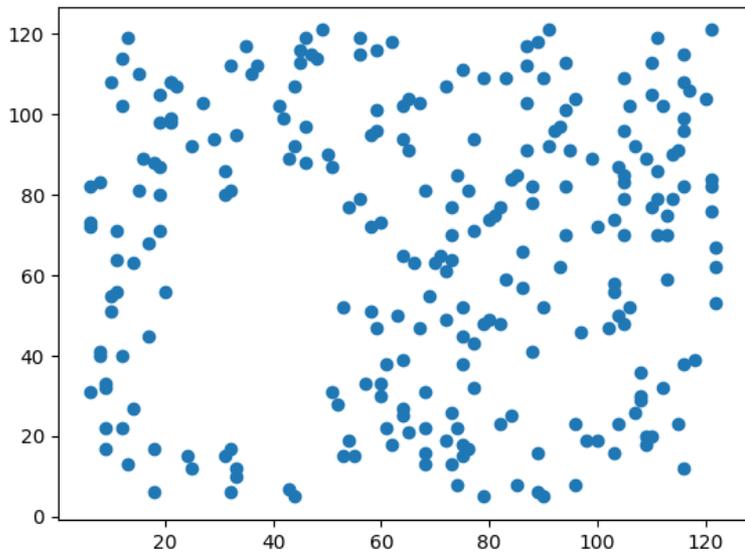


Data fusion

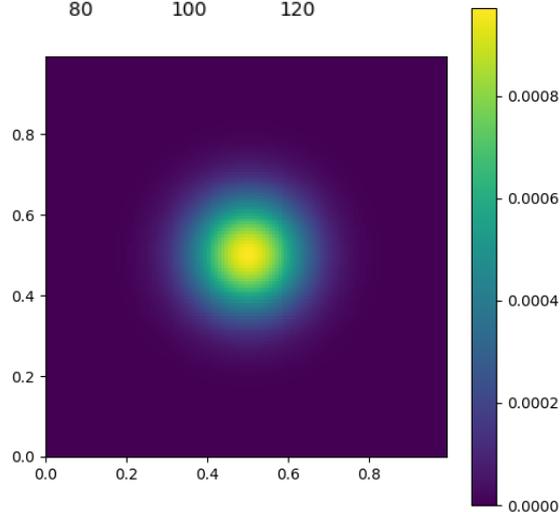
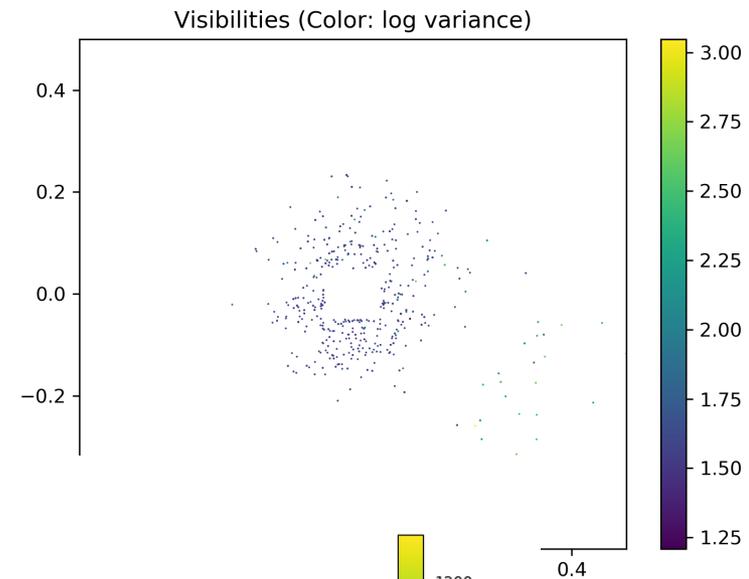
$$\mathcal{P}(d_1, d_2, s) = \mathcal{P}(d_1|s) \mathcal{P}(d_2|s) \mathcal{P}(s) \cdot \mathcal{H}(s)$$

single dish

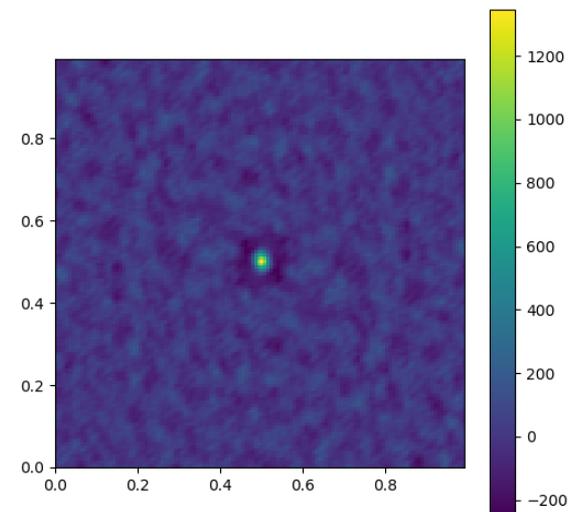
pointing



interferometer

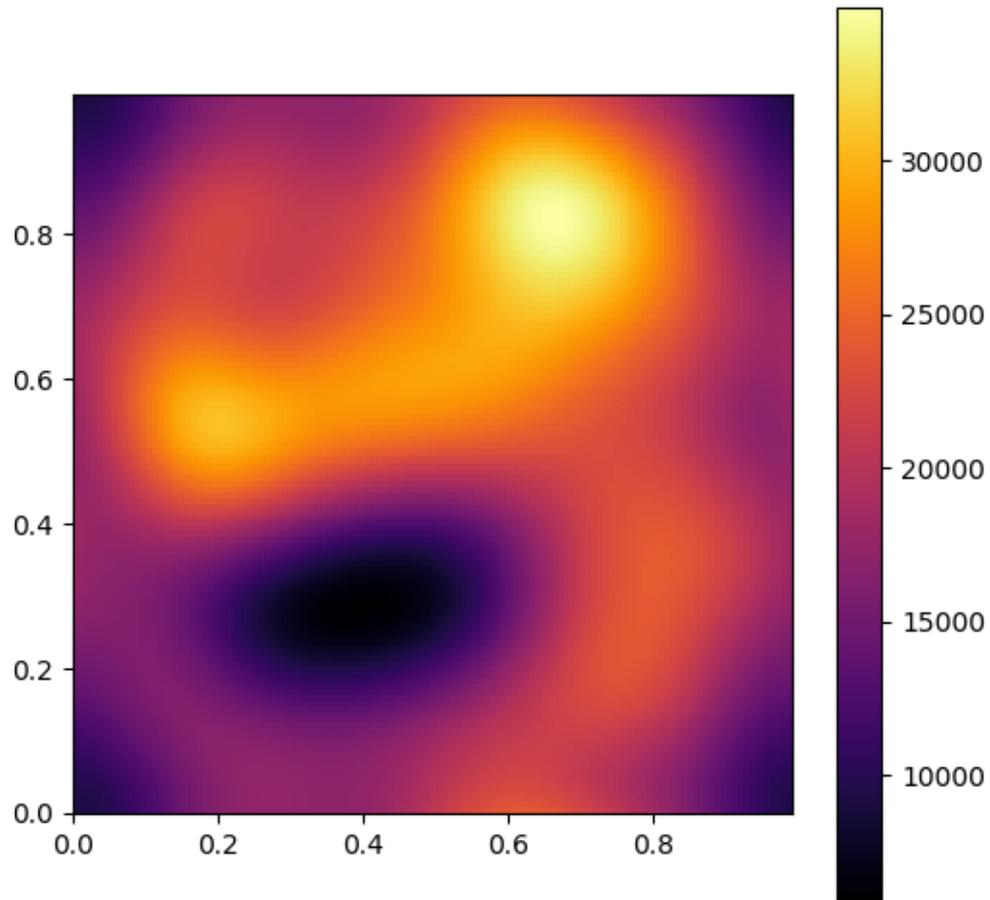


point
spread
function

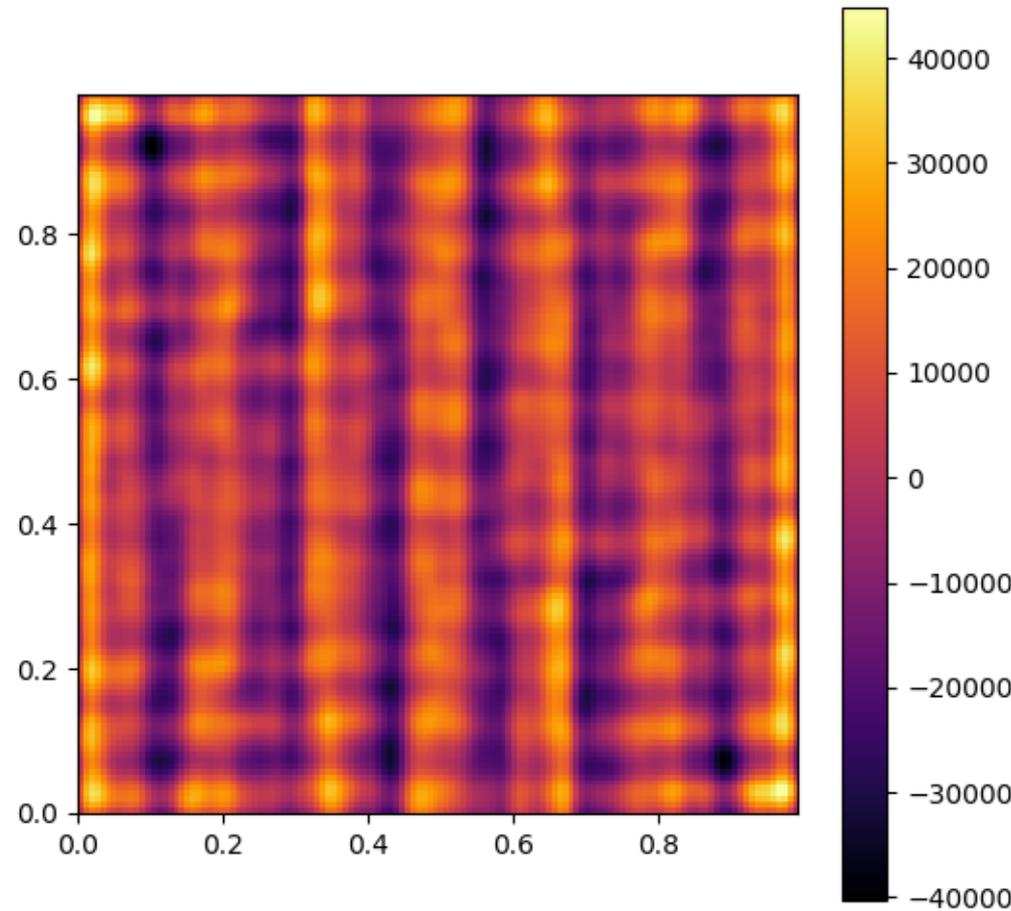


Dirty images

single dish

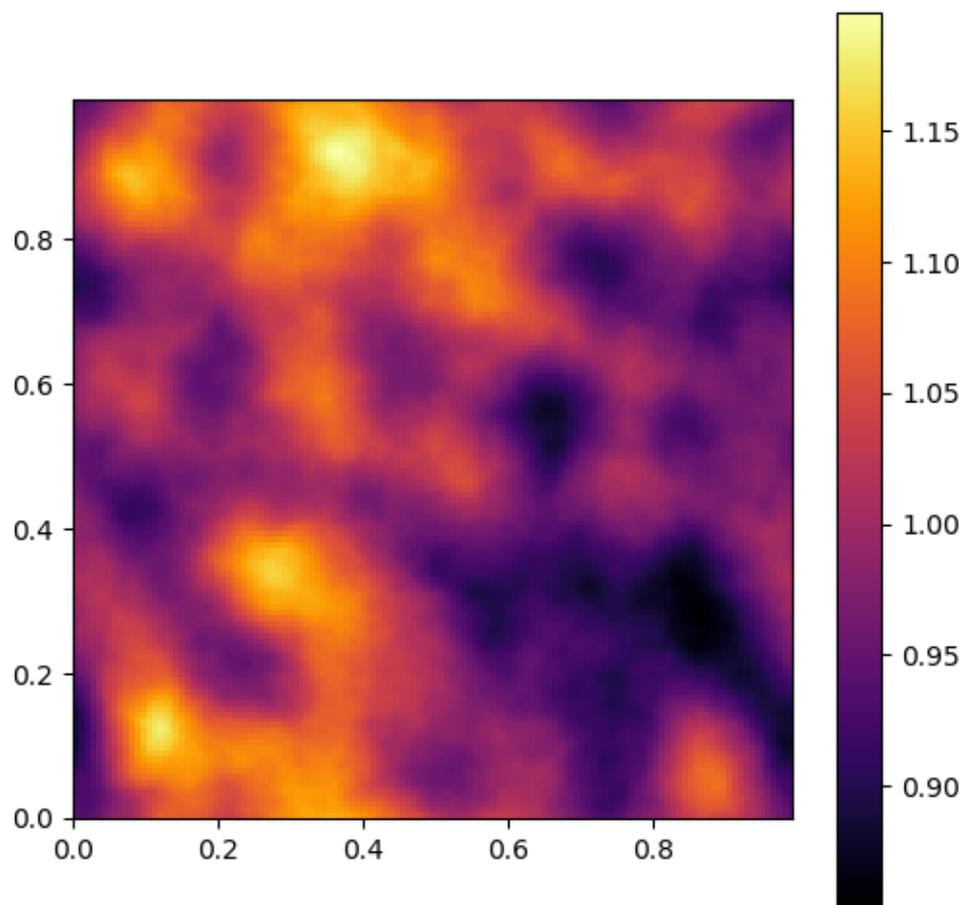


interferometer

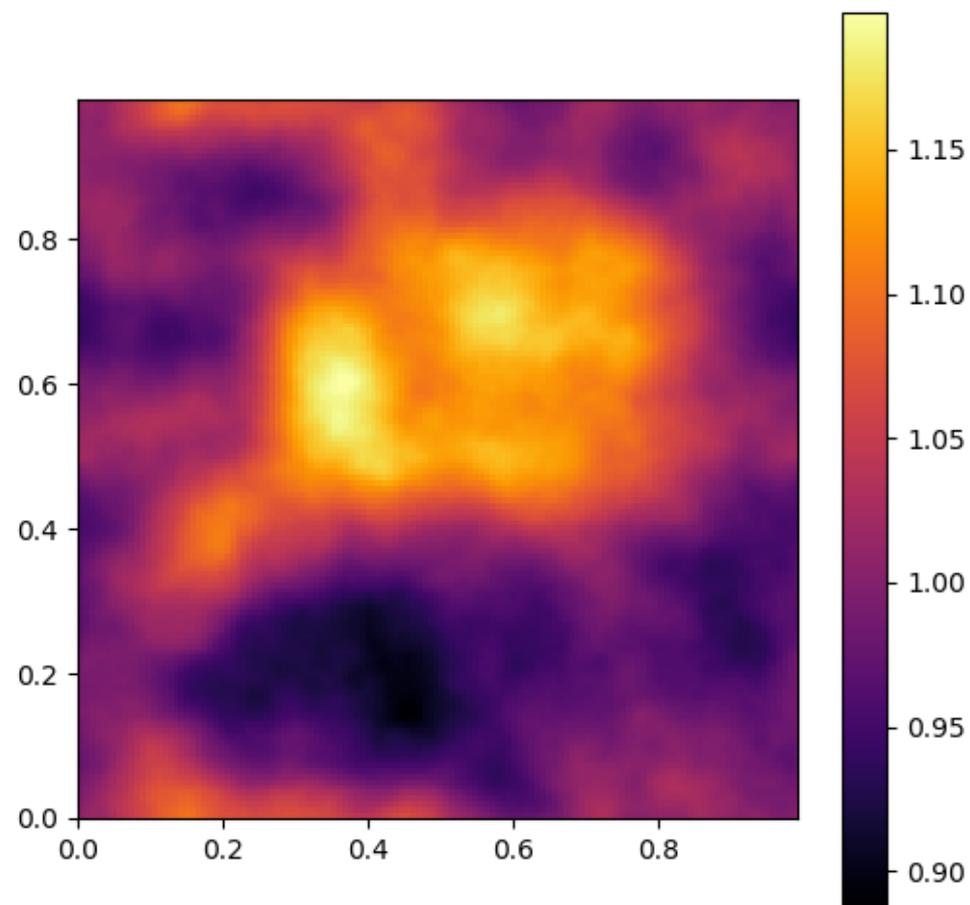


Images

single dish

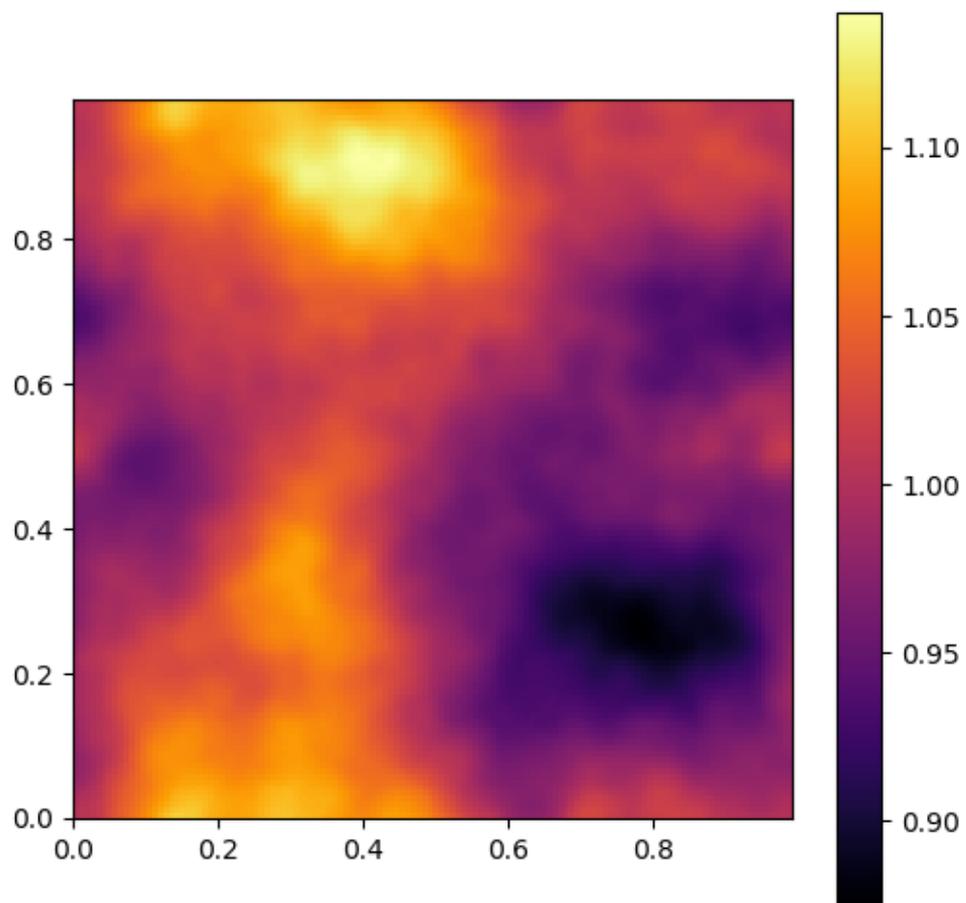


interferometer

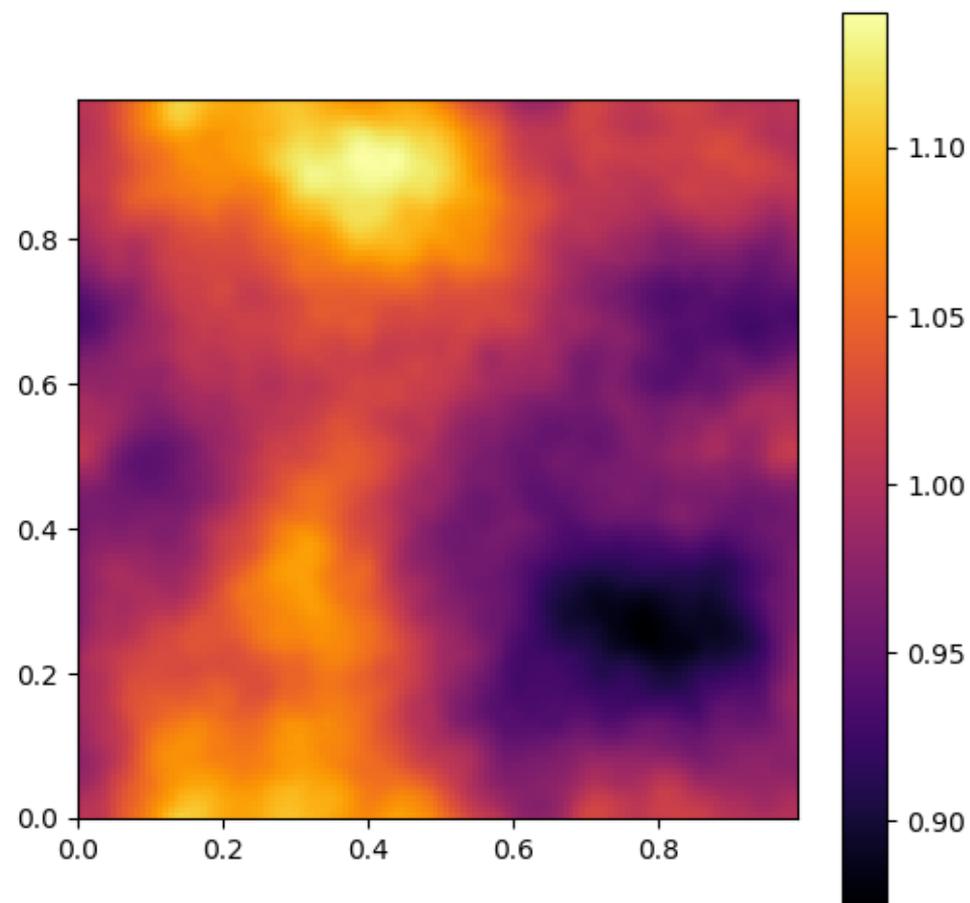


Image

single dish + interferometer

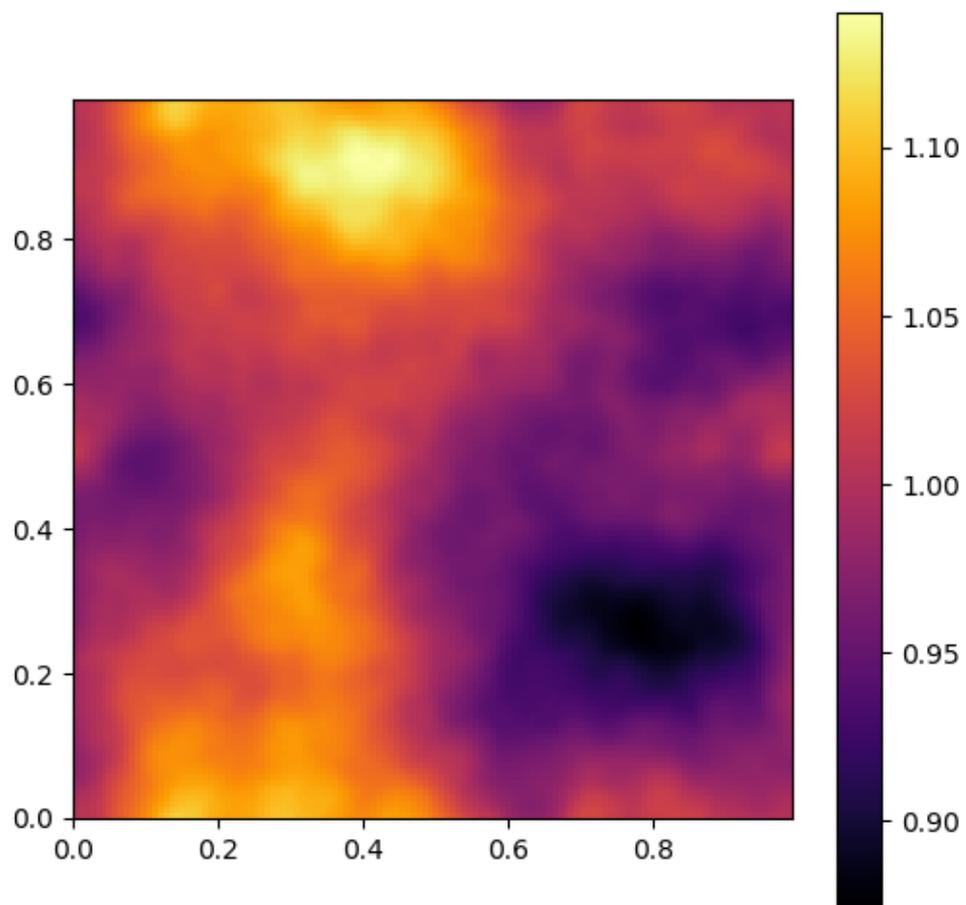


single dish + interferometer

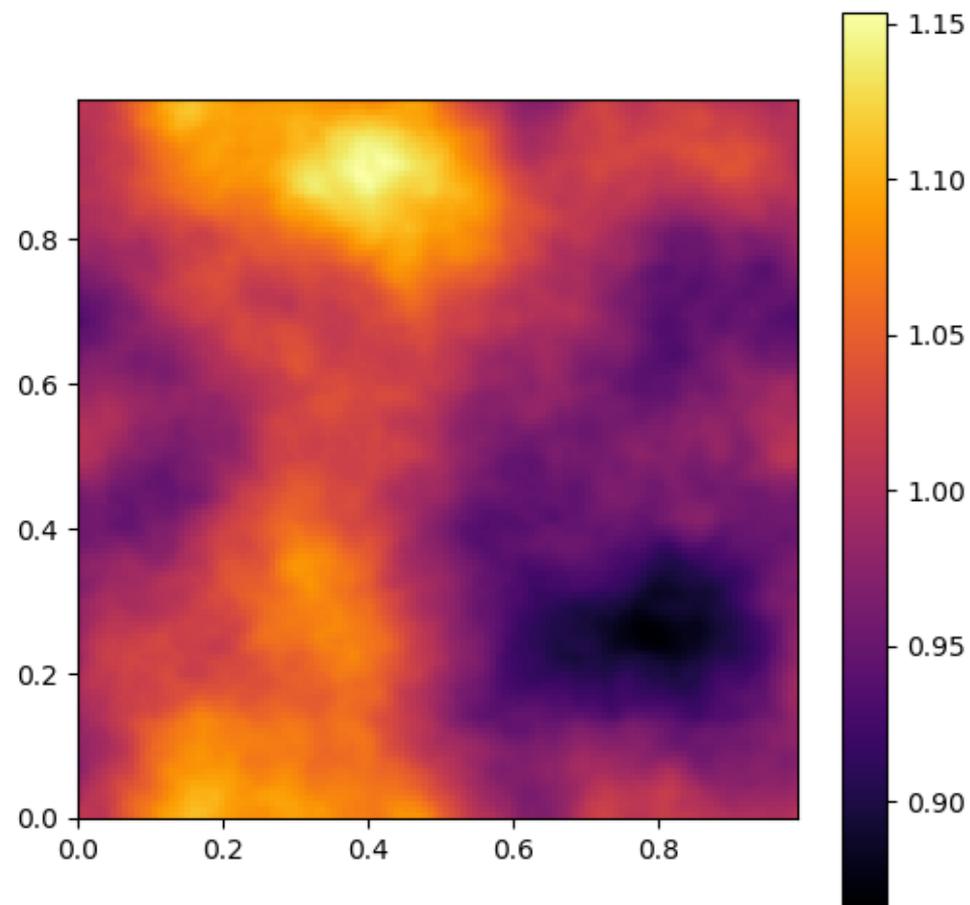


Image

single dish + interferometer



real sky

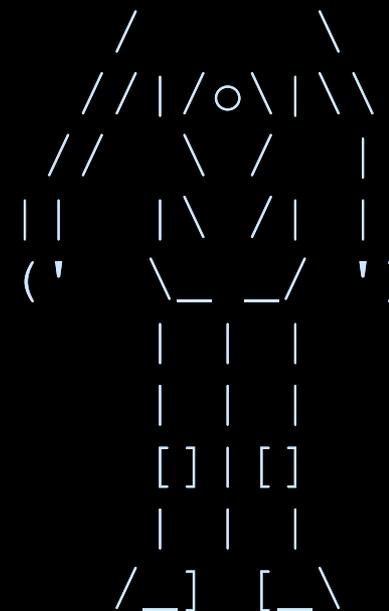
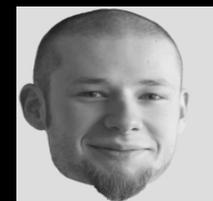


Denoising, Deconvolving, and Decomposing Photon Observations

Selig et al. (2014)

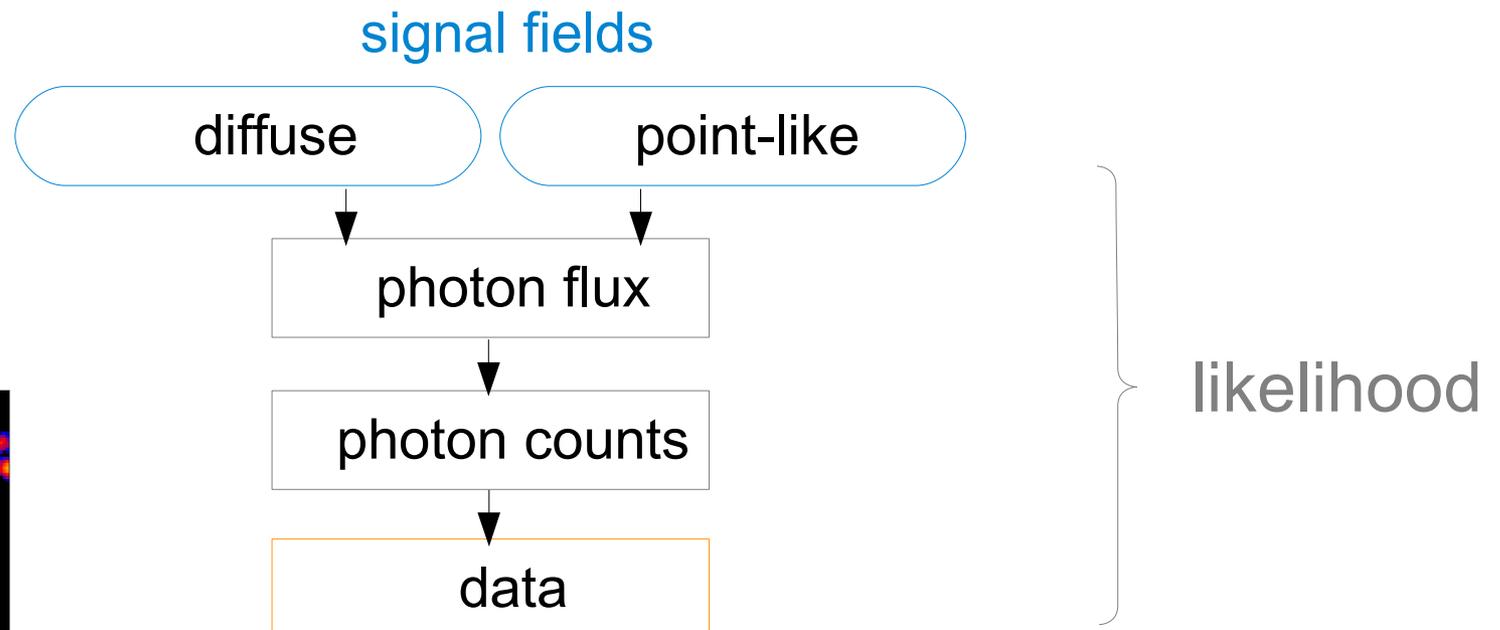
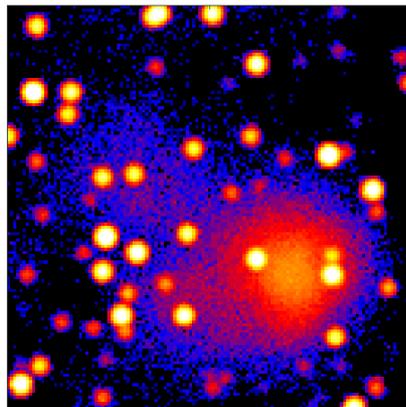
www.mpa-garching.mpg.de/ift/d3po

D³PO



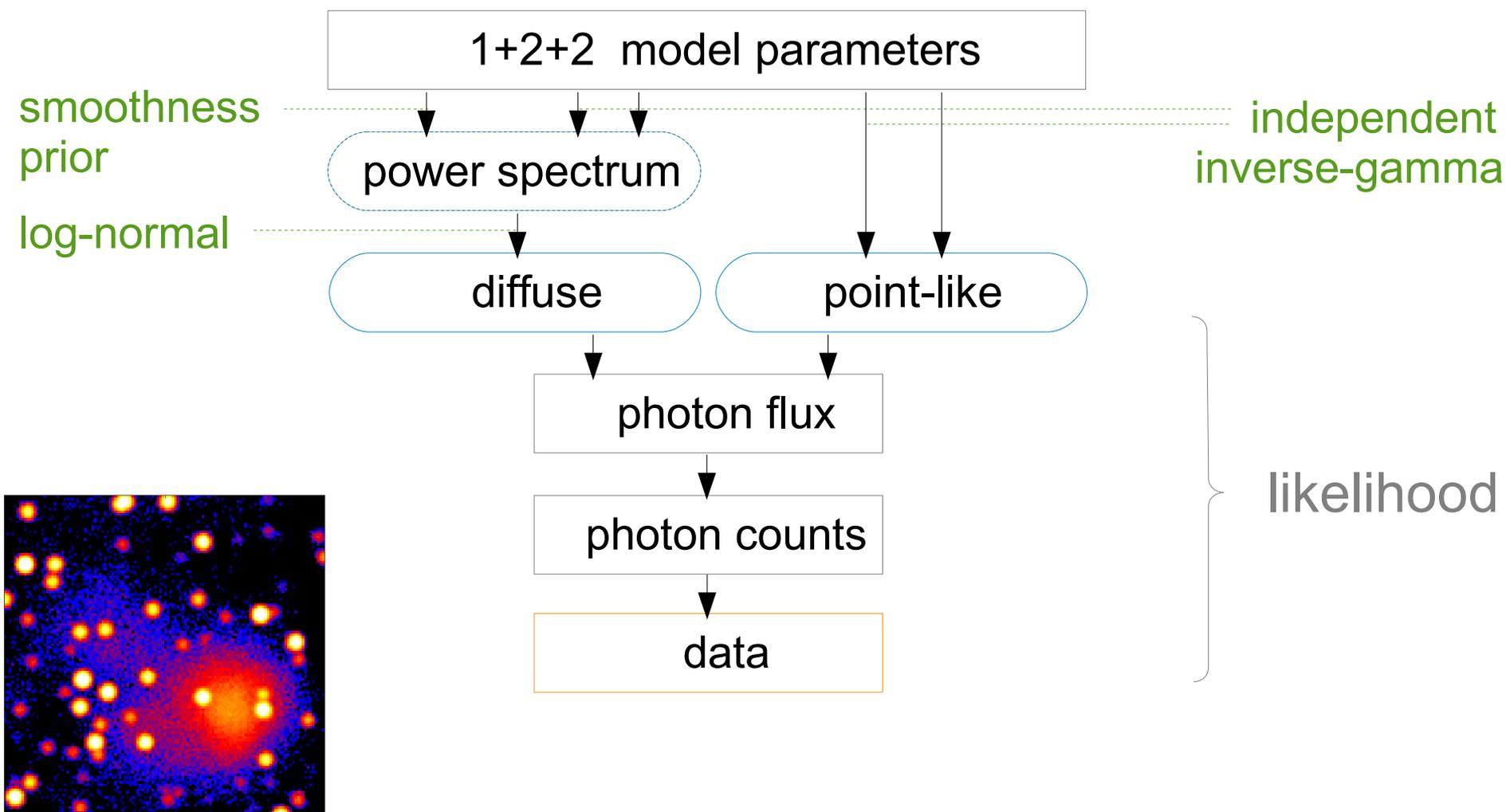
D³PO – challenges & assumptions

Selig & Enßlin
(2014)
arXiv: 1311.1888



D³PO – challenges & assumptions

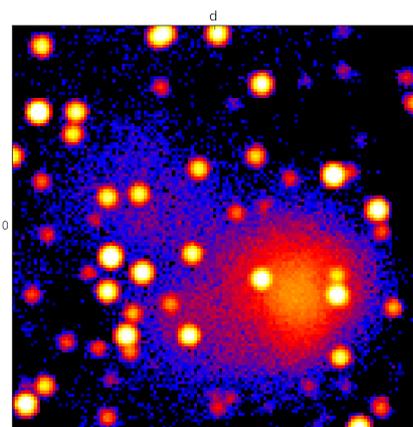
Selig & Enßlin
(2014)
arXiv: 1311.1888



D³PO – challenges & assumptions

Selig & Enßlin
(2014)
arXiv: 1311.1888

$$\begin{aligned}\mathcal{H}(\mathbf{s}, \boldsymbol{\tau}, \mathbf{u} | \mathbf{d}) &= -\log \mathcal{P}(\mathbf{s}, \boldsymbol{\tau}, \mathbf{u} | \mathbf{d}) \\ &= H_0 + \mathbf{1}^\dagger \mathbf{R} (e^{\mathbf{s}} + e^{\mathbf{u}}) - \mathbf{d}^\dagger \log (\mathbf{R} (e^{\mathbf{s}} + e^{\mathbf{u}})) \\ &\quad + \frac{1}{2} \log (\det [\mathbf{S}]) + \frac{1}{2} \mathbf{s}^\dagger \mathbf{S}^{-1} \mathbf{s} \\ &\quad + (\boldsymbol{\alpha} - \mathbf{1})^\dagger \boldsymbol{\tau} + \mathbf{q}^\dagger e^{-\boldsymbol{\tau}} + \frac{1}{2} \boldsymbol{\tau}^\dagger \mathbf{T} \boldsymbol{\tau} \\ &\quad + (\boldsymbol{\beta} - \mathbf{1})^\dagger \mathbf{u} + \boldsymbol{\eta}^\dagger e^{-\mathbf{u}}\end{aligned}$$

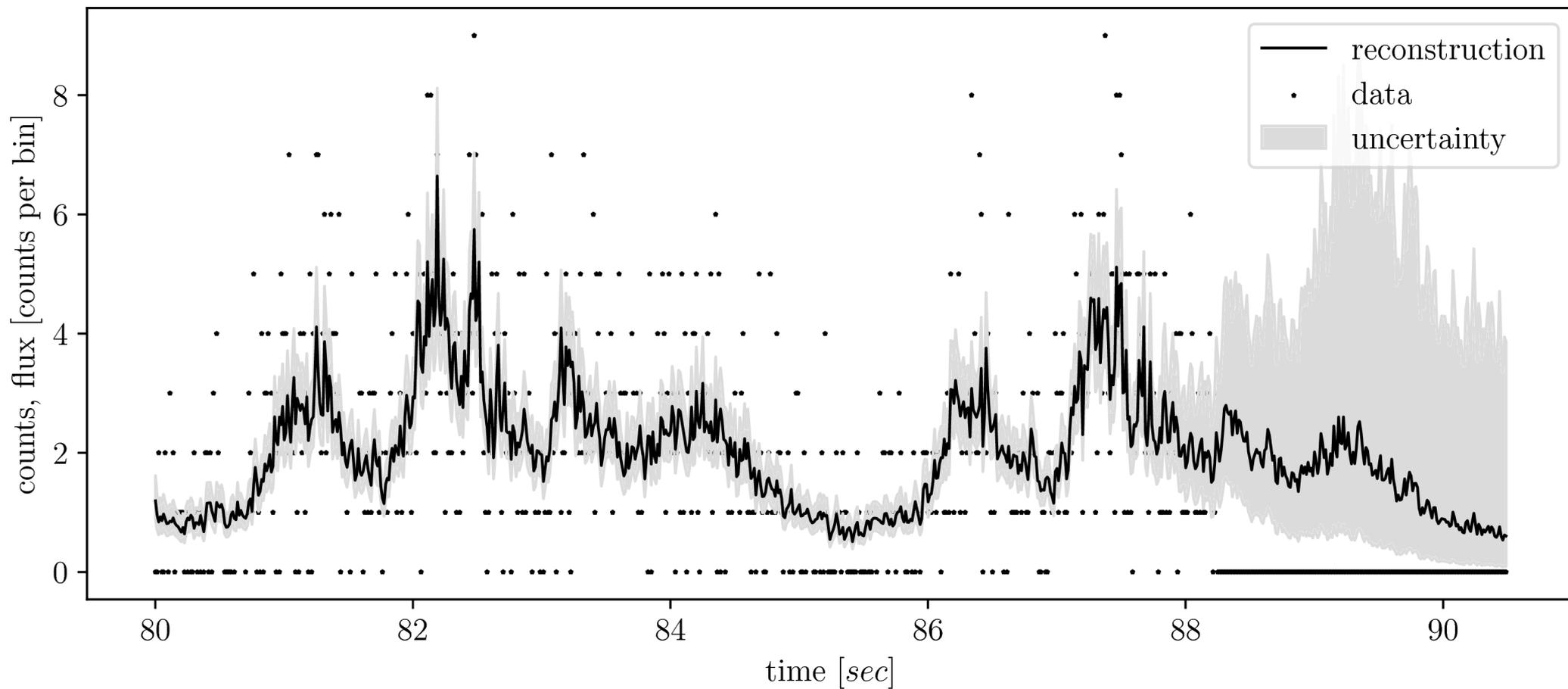


$$\mathbf{S} = \sum_k e^{\tau_k} \mathbf{S}_k$$

D³PO in 1D & QPOs

Magnetar flare SGR 1900+14

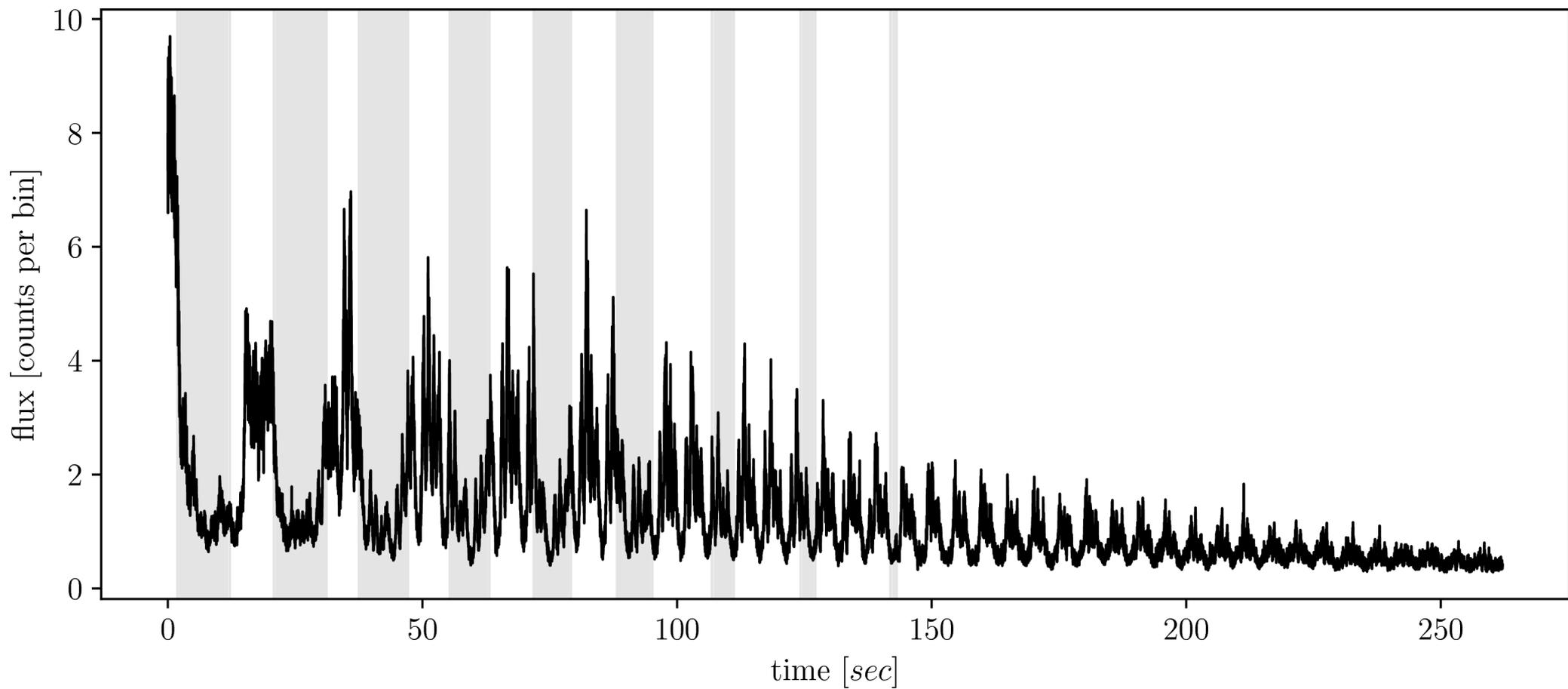
Pumpe et al. arXiv:1708.05702



D³PO in 1D & QPOs

Magnetar flare SGR 1900+14

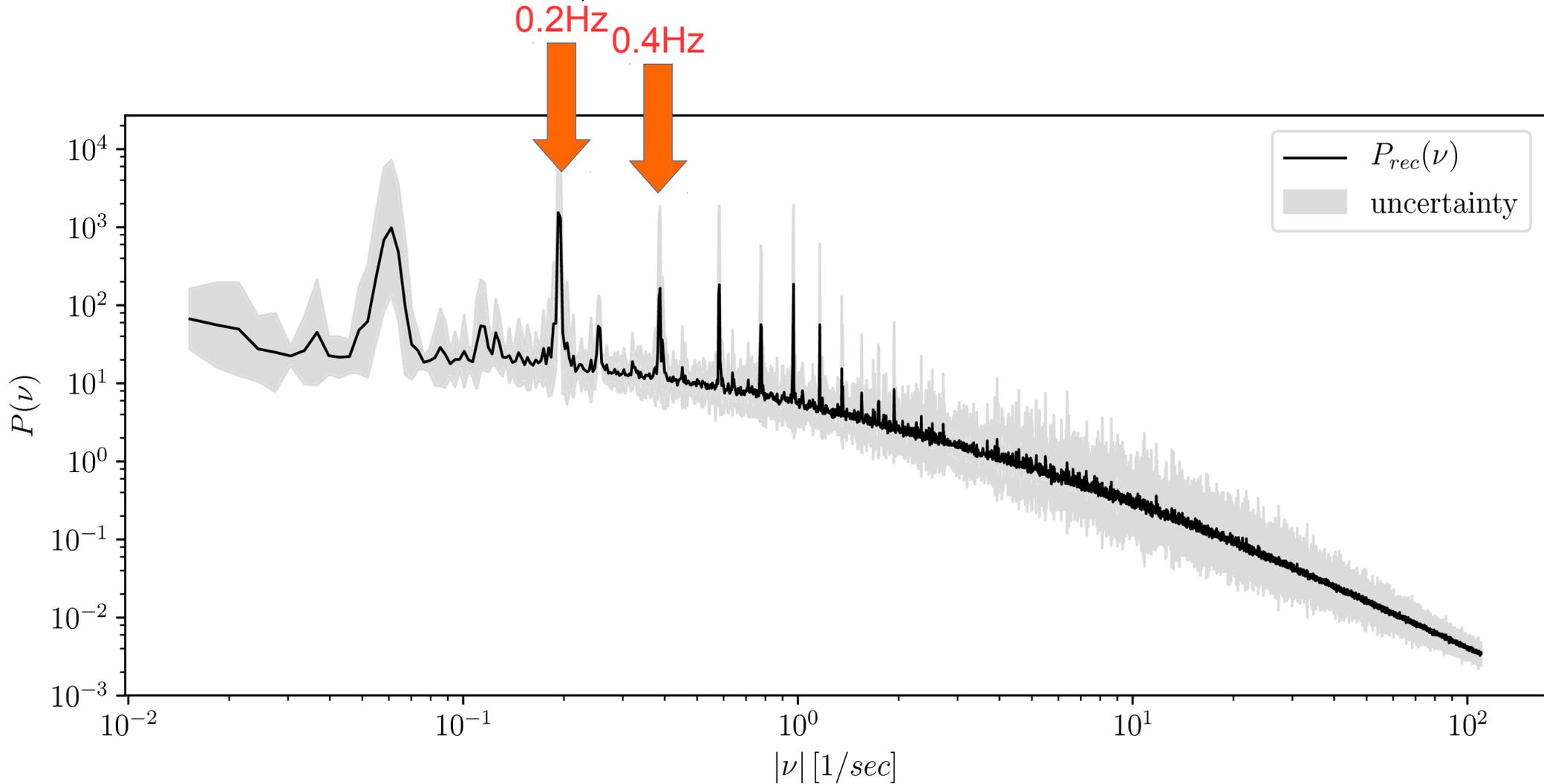
Pumpe et al. arXiv:1708.05702

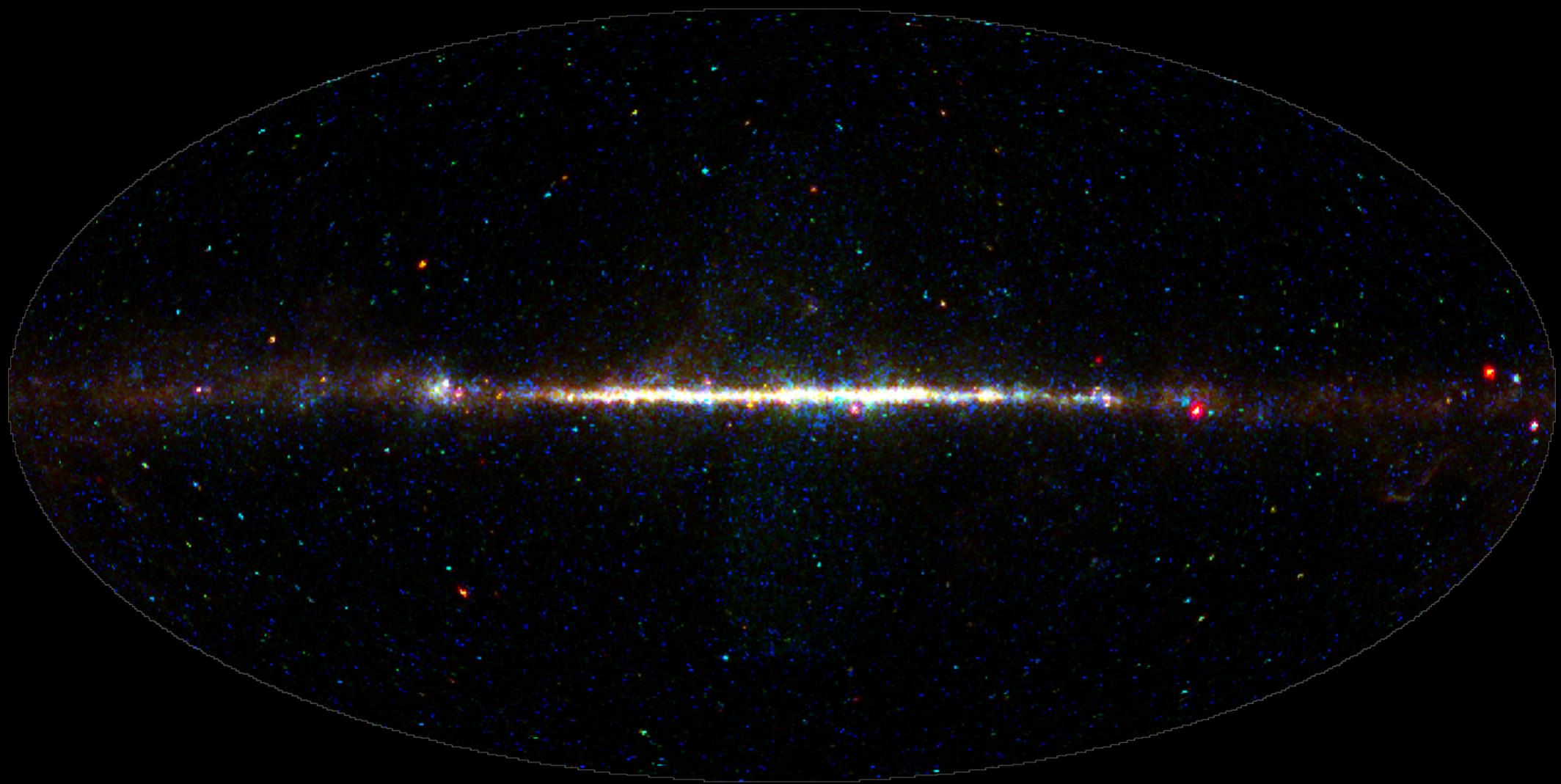


D³PO in 1D & QPOs

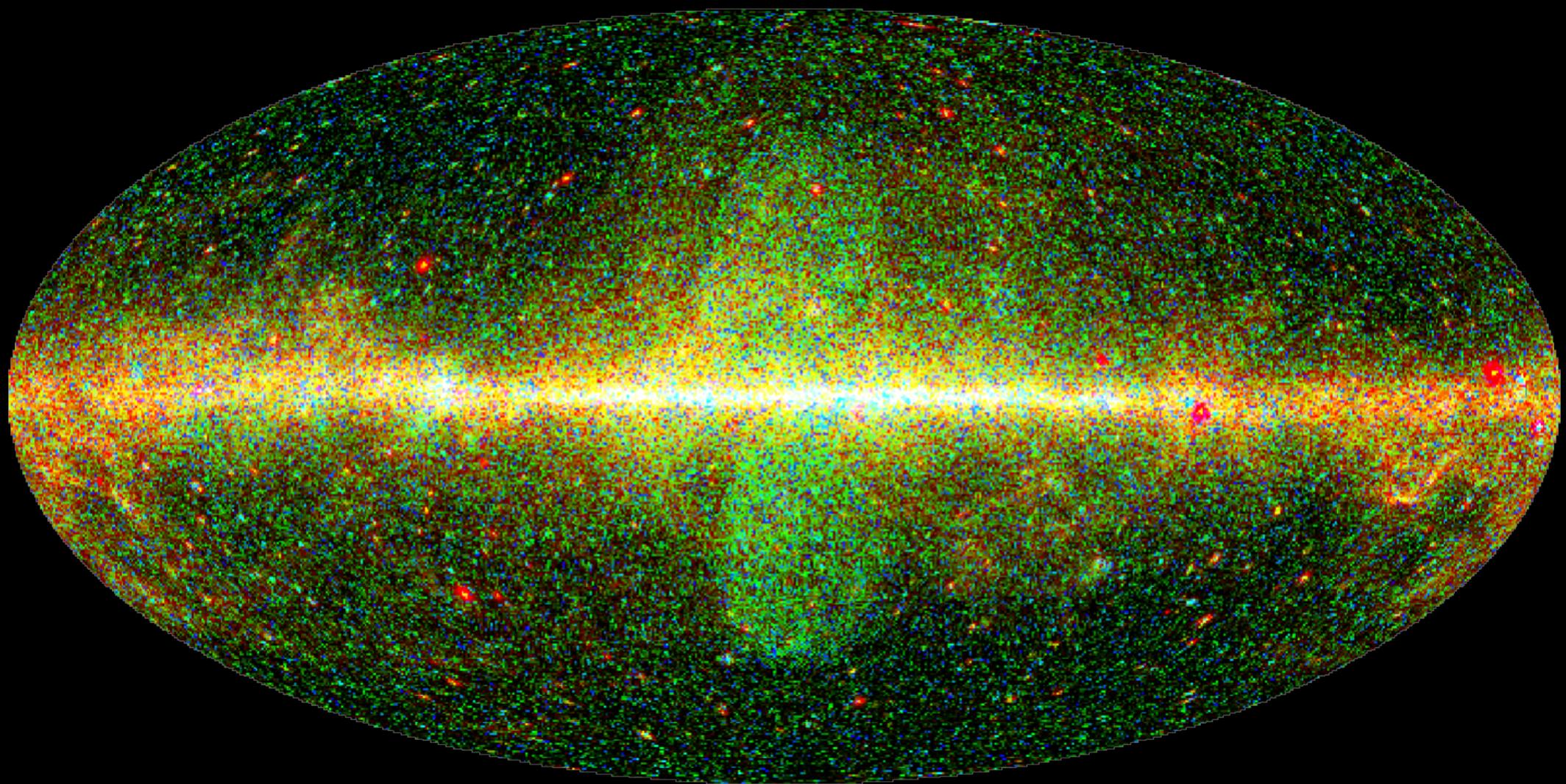
Magnetar flare SGR 1900+14

Pumpe et al. arXiv:1708.05702

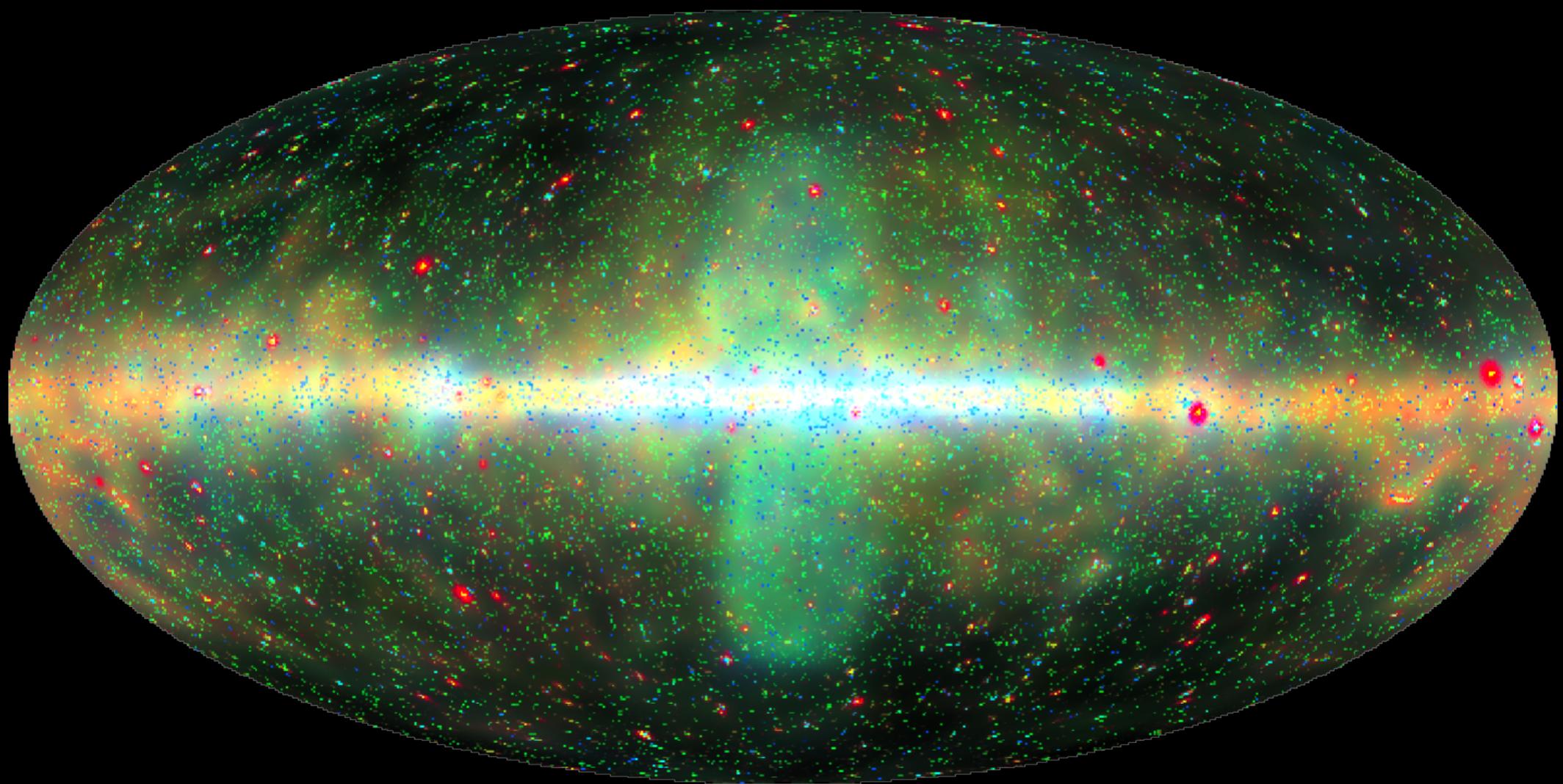




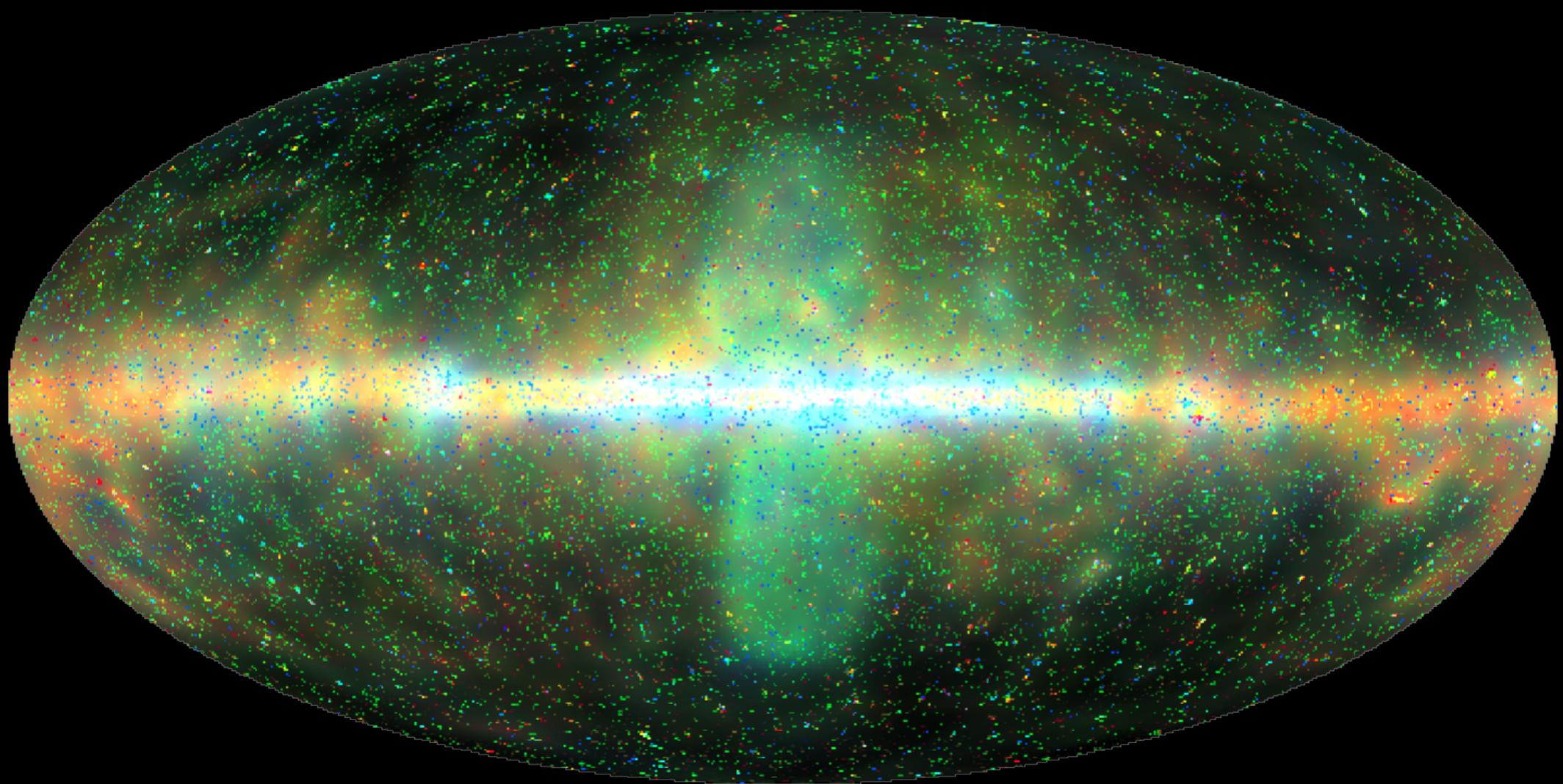
data



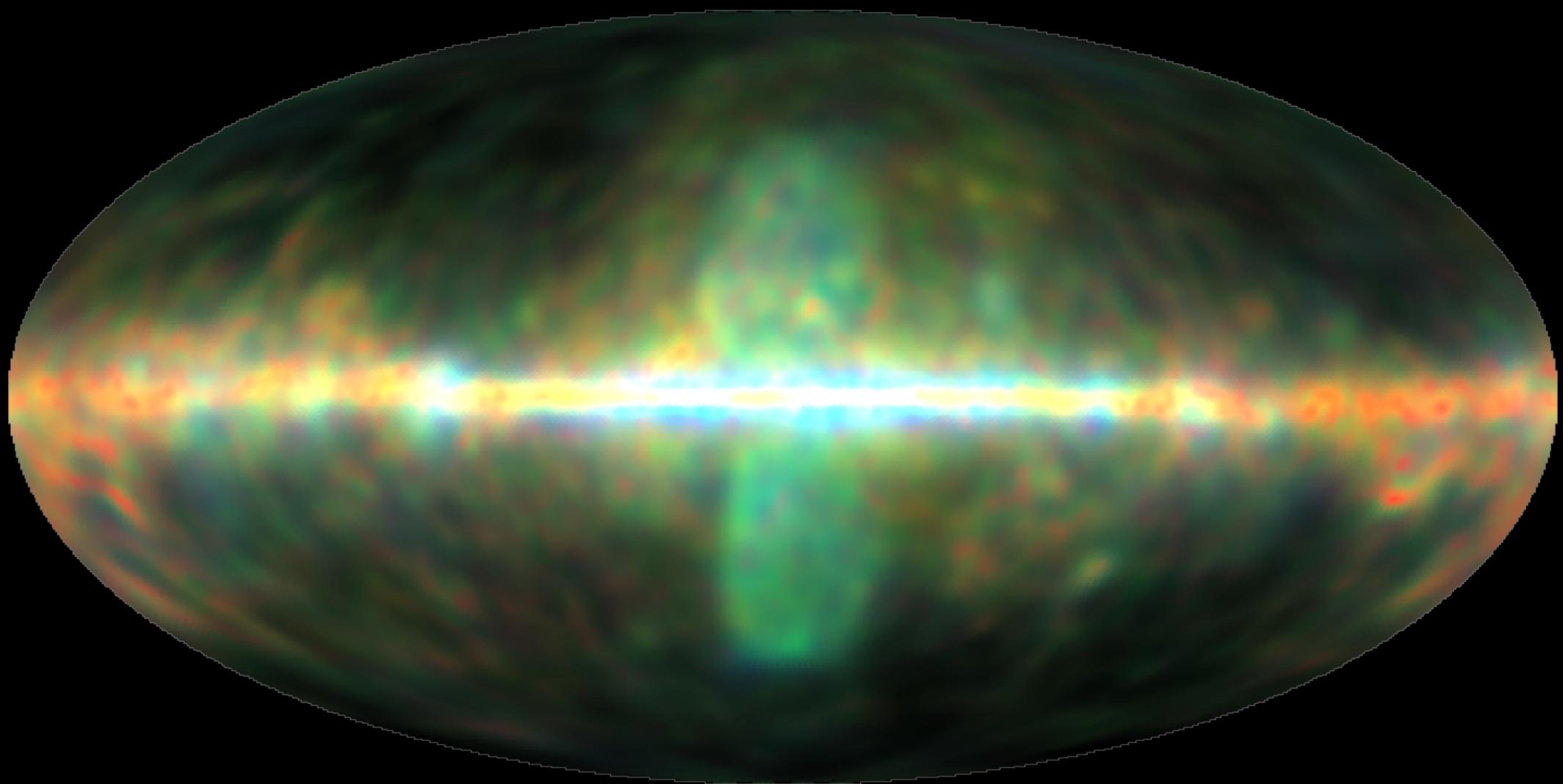
log-data



log-data ... denoised

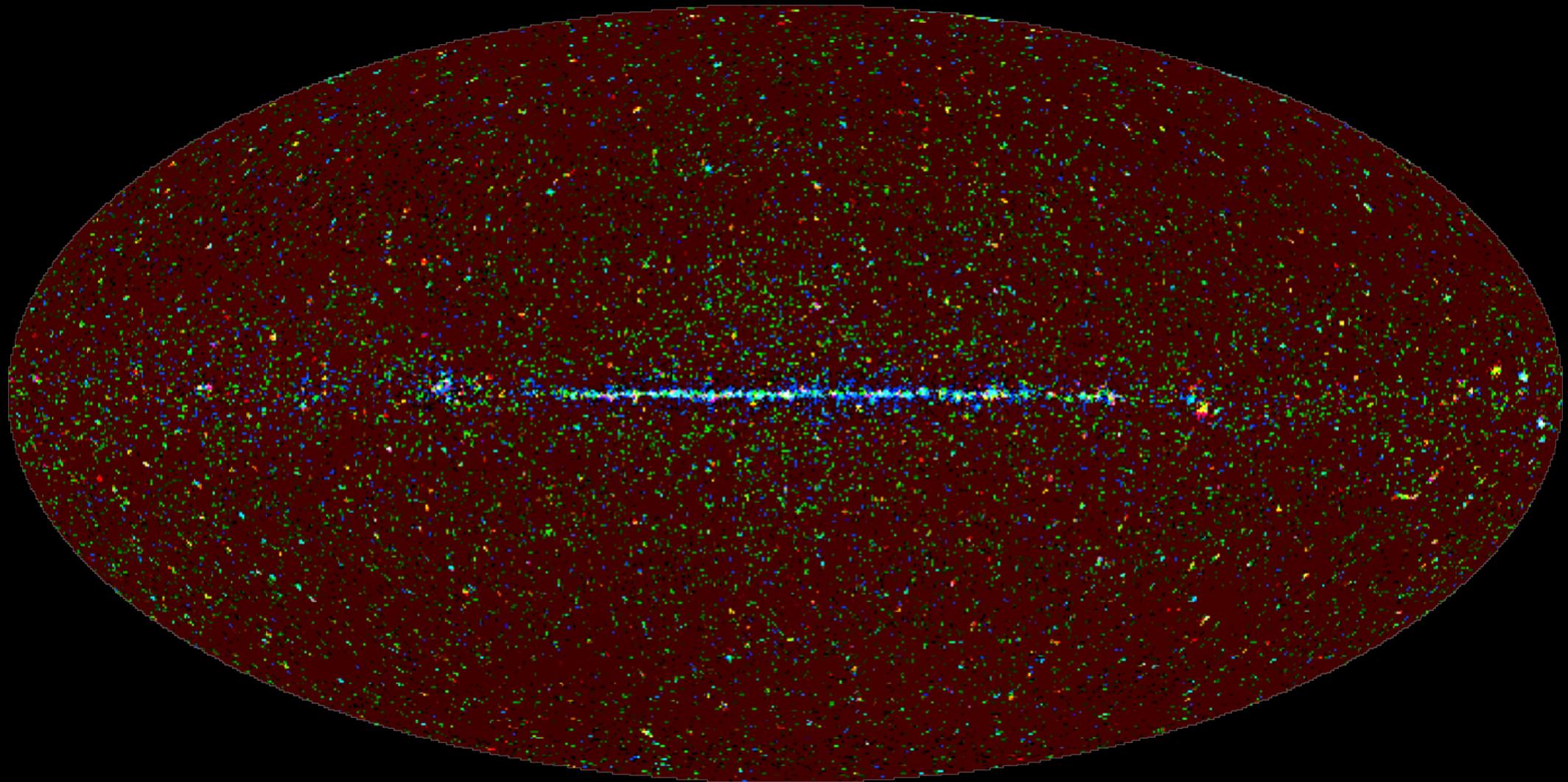


log-data ... denoised ... deconvolved



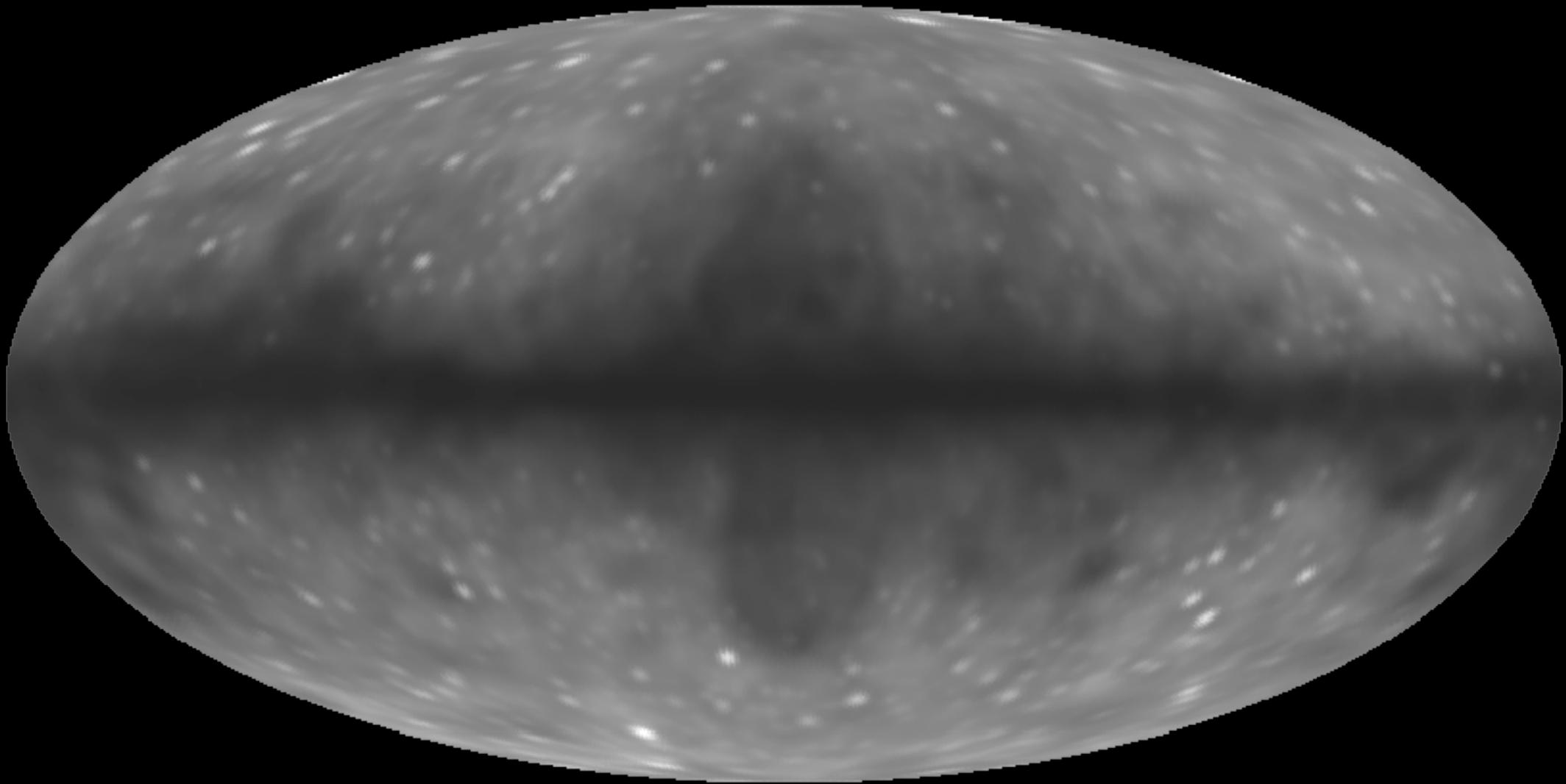
log-data ... denoised ... deconvolved ... decomposed

Selig, Vacca, Oppermann, Enßlin (2015)



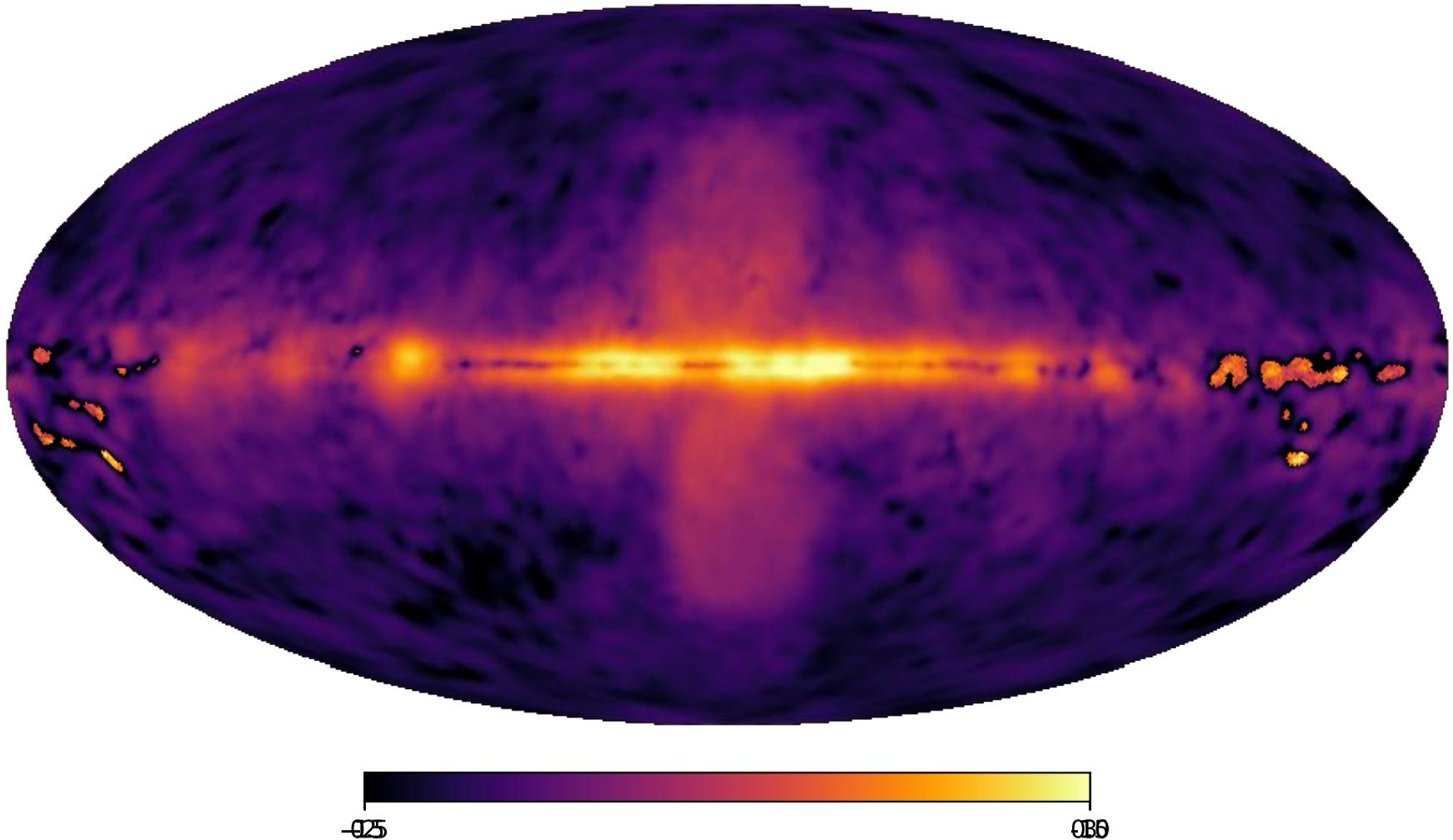
log-data ... denoised ... deconvolved ... decomposed

Selig, Vacca, Oppermann, Enßlin (2015)

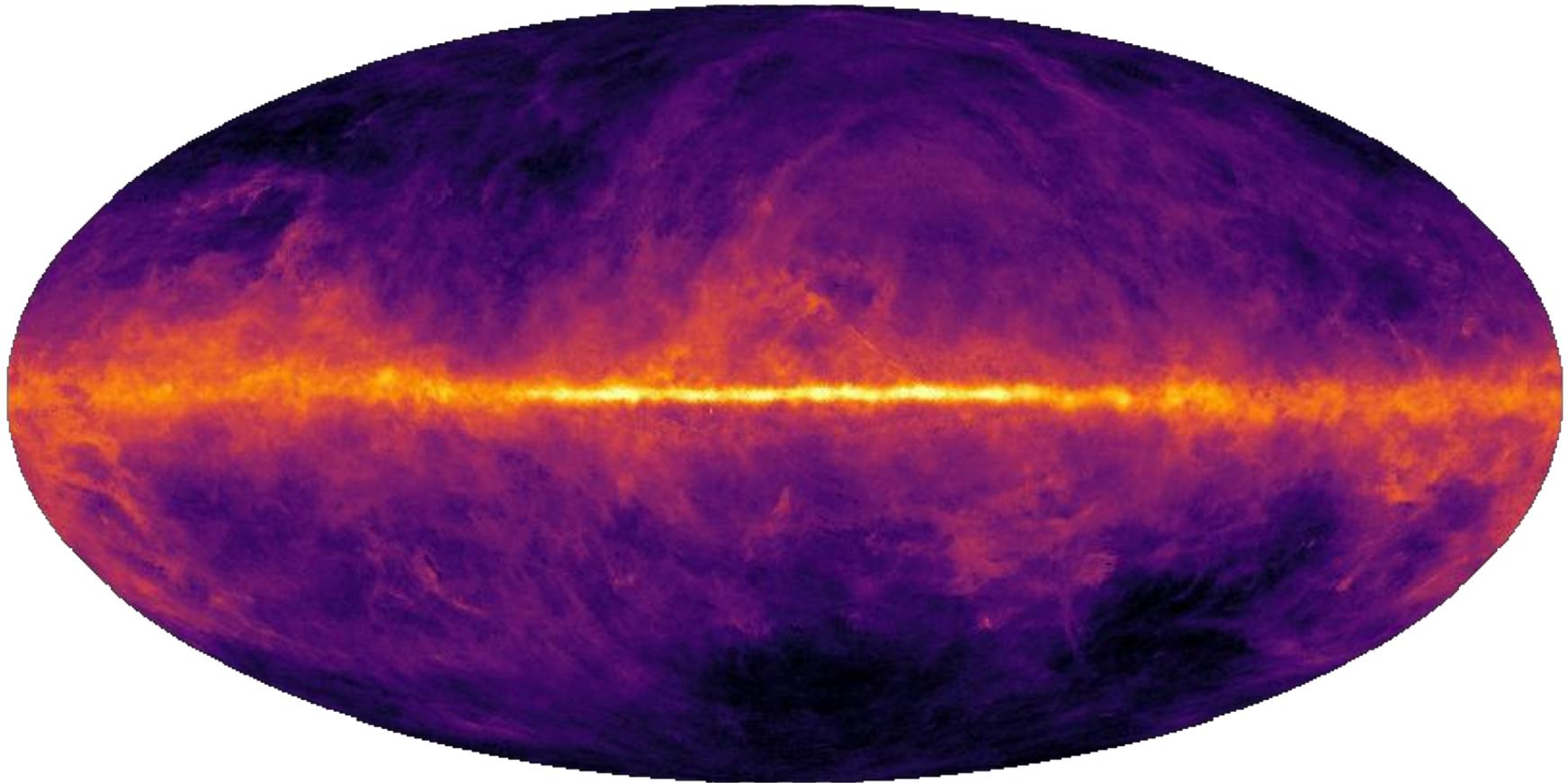


relative uncertainty of diffuse emission

Sharpening up the gamma-ray sky



Restoring the hadronic sky

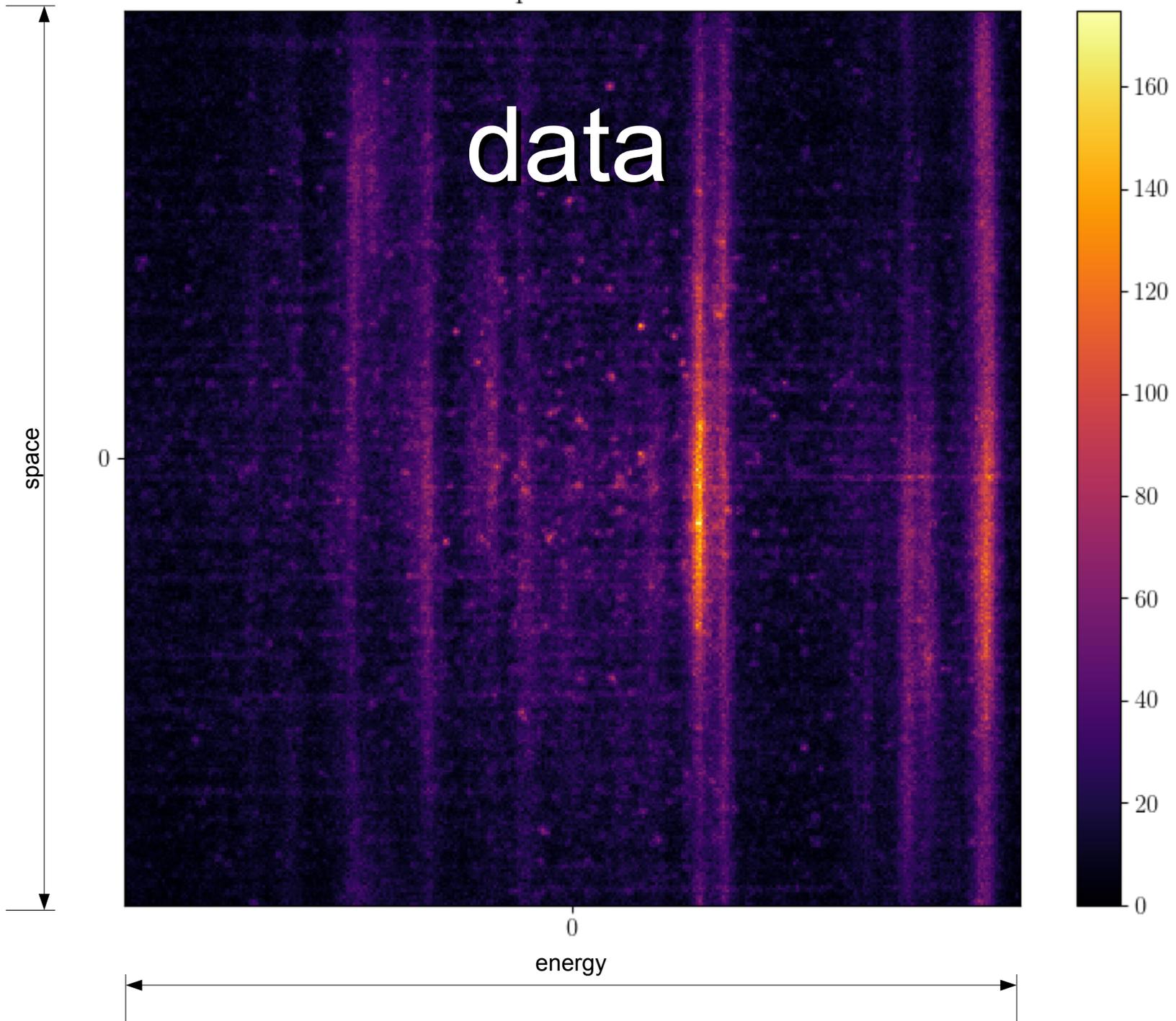


Glimpses into the future



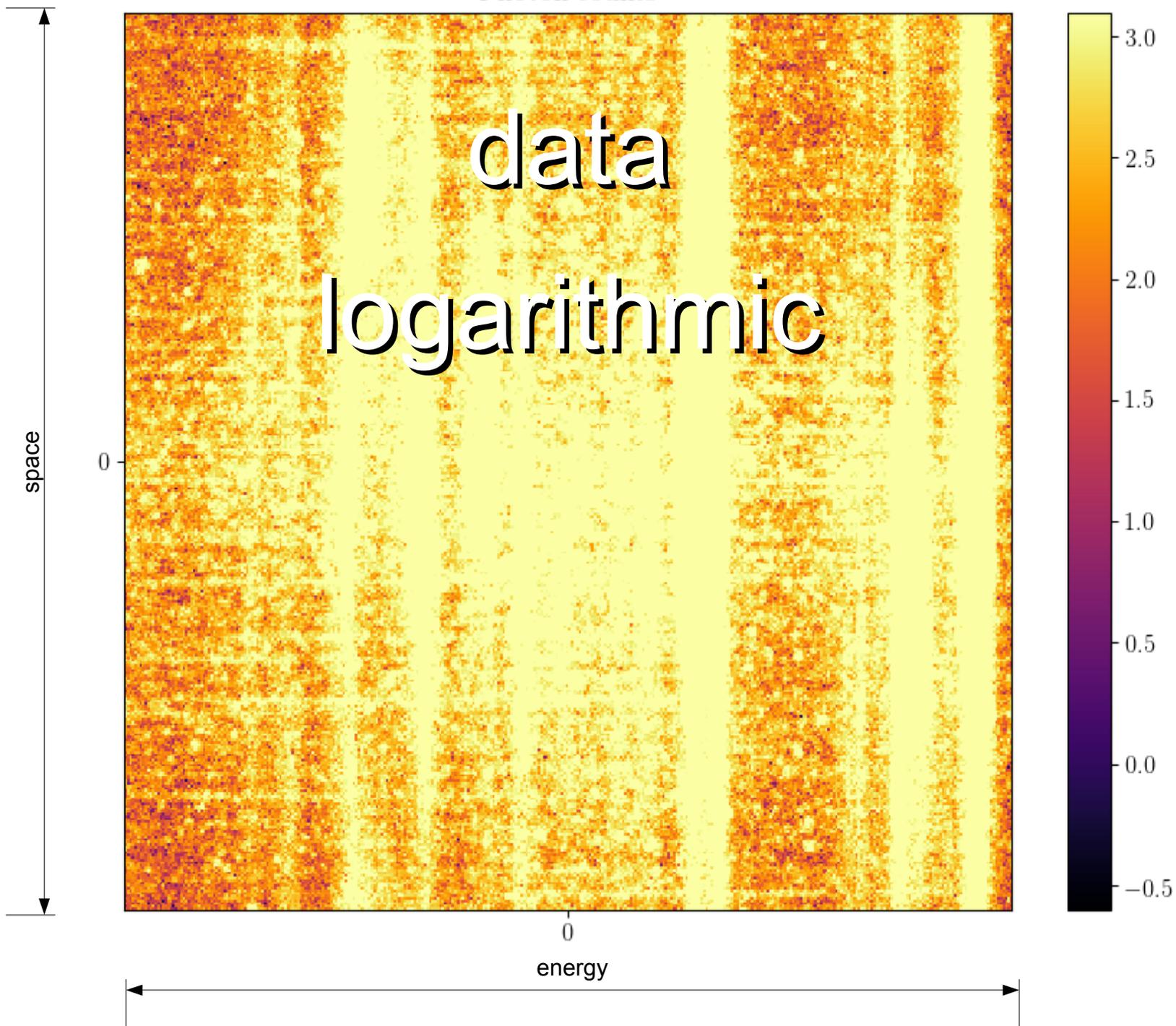
Daniel Pumpe: D⁴PO
arXiv:1802.020135

Raw photon counts

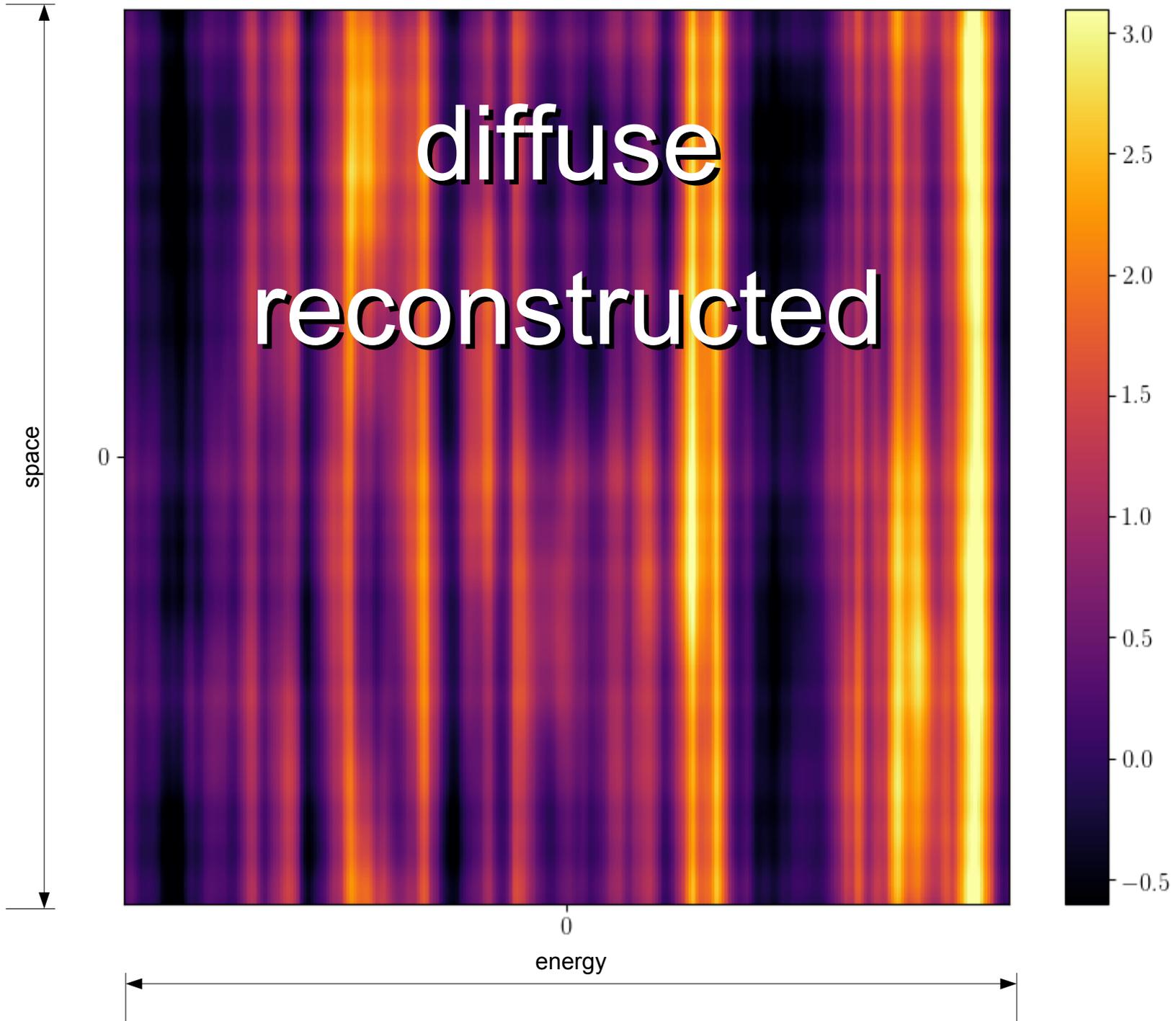


2018

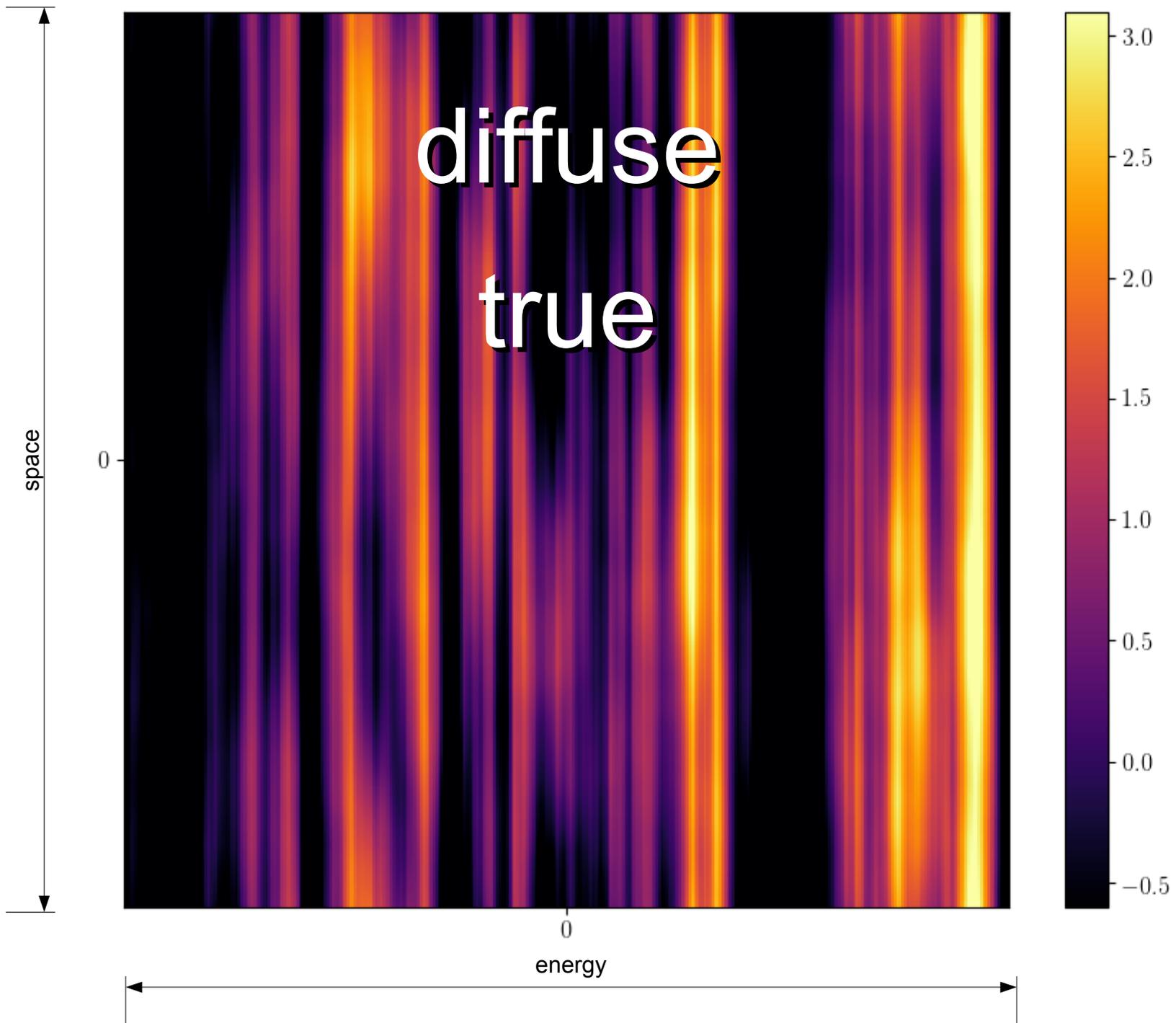
Photon counts



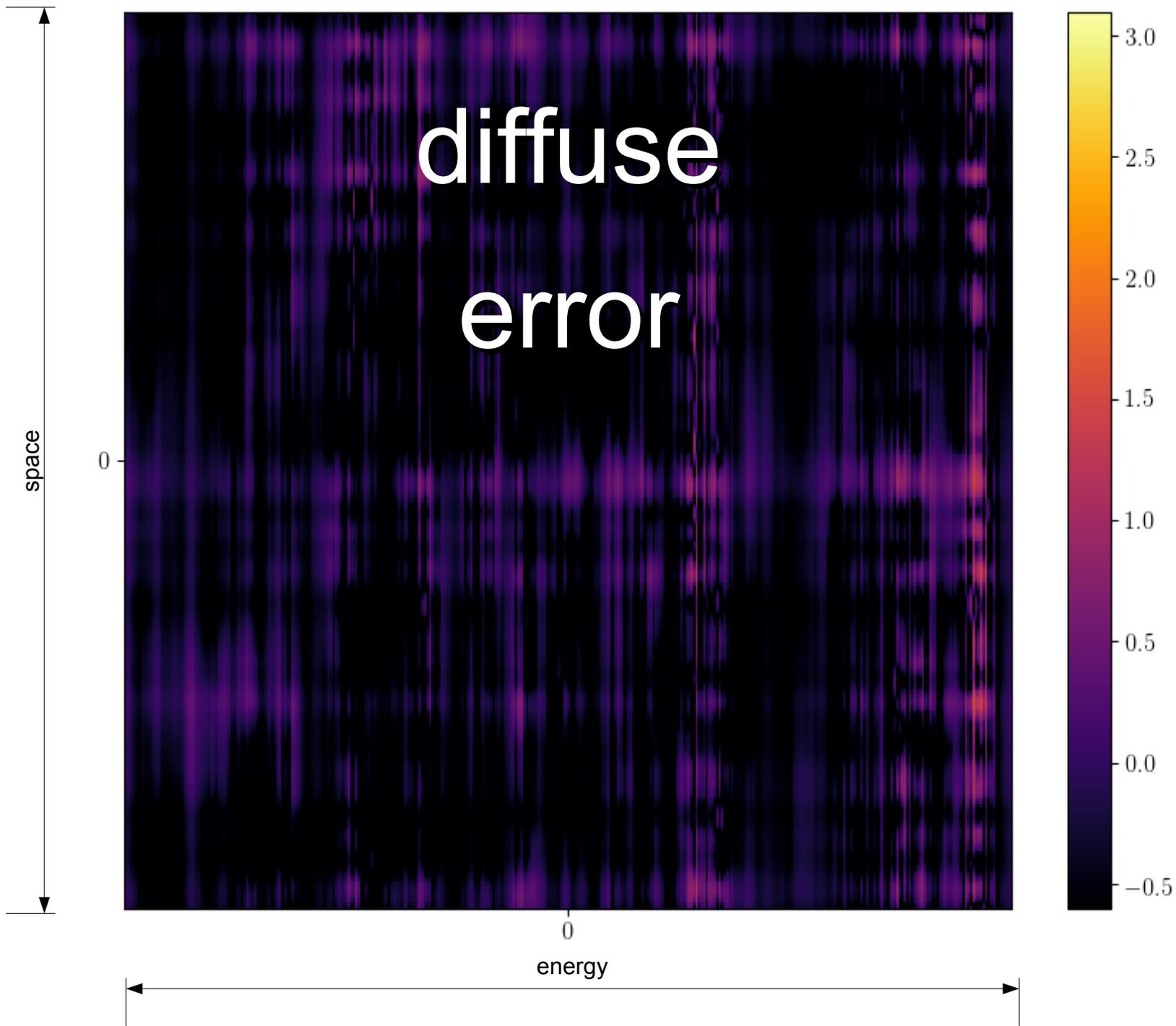
Reconstructed diffuse flux



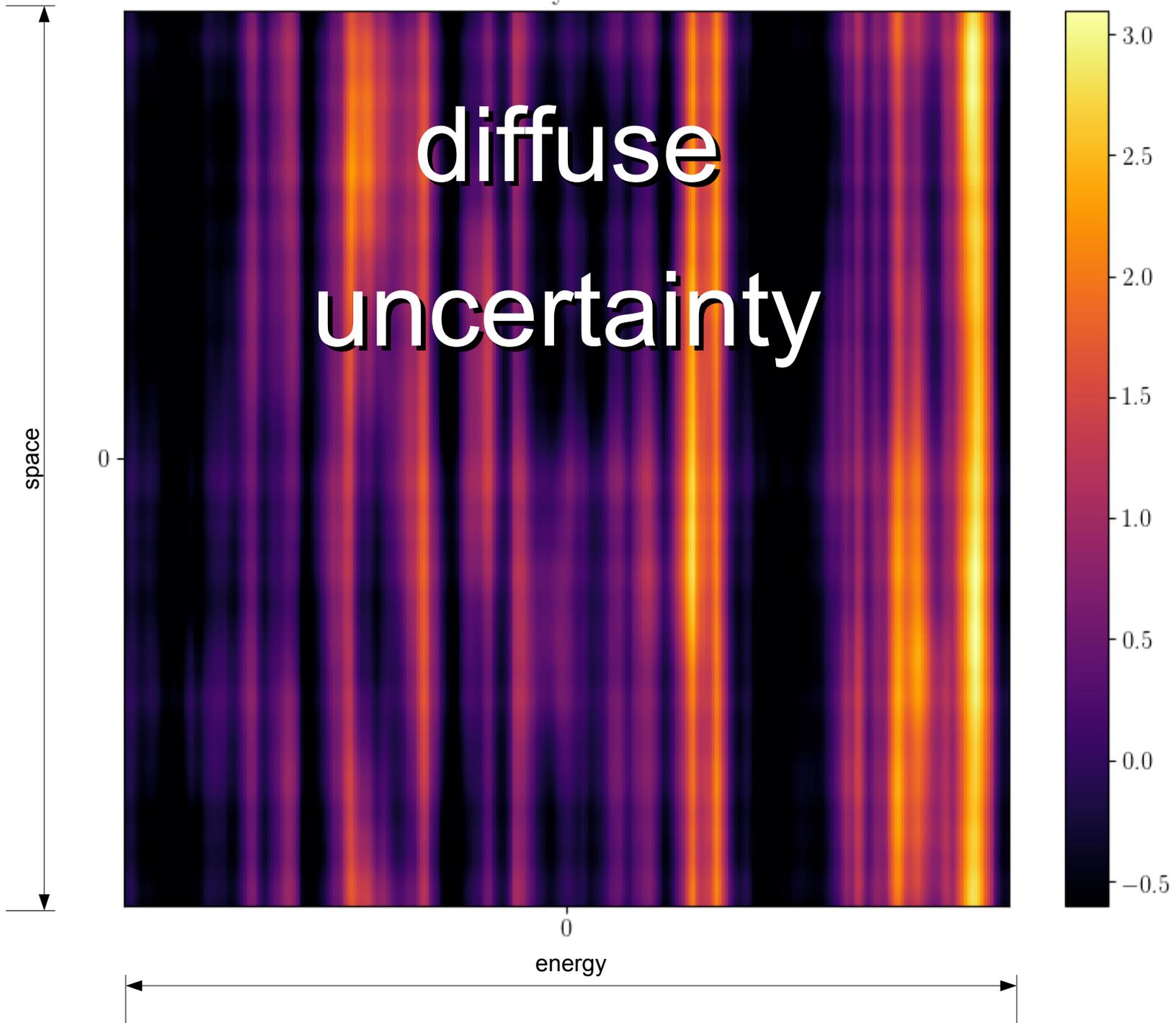
Diffuse flux



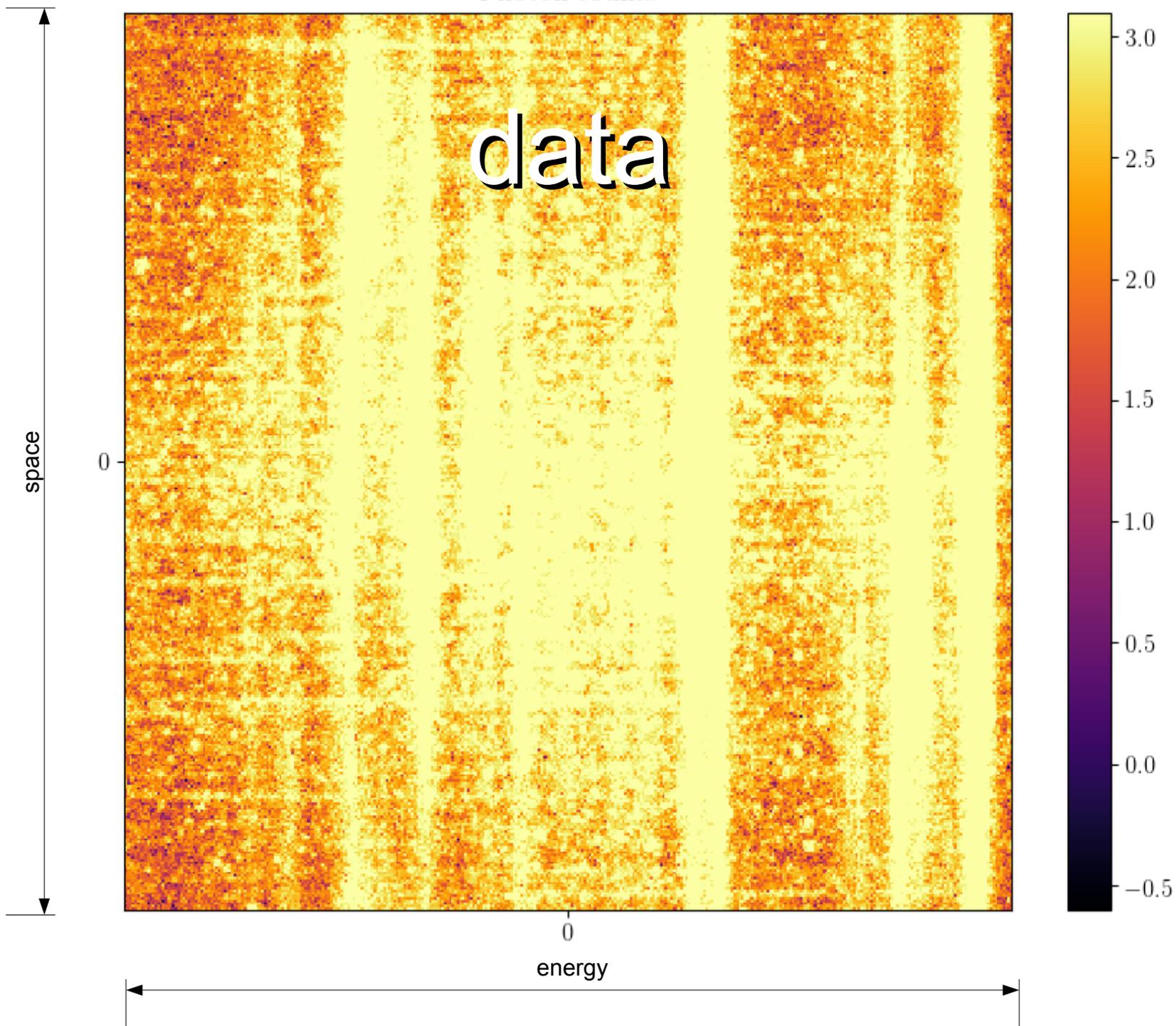
Error diffuse flux



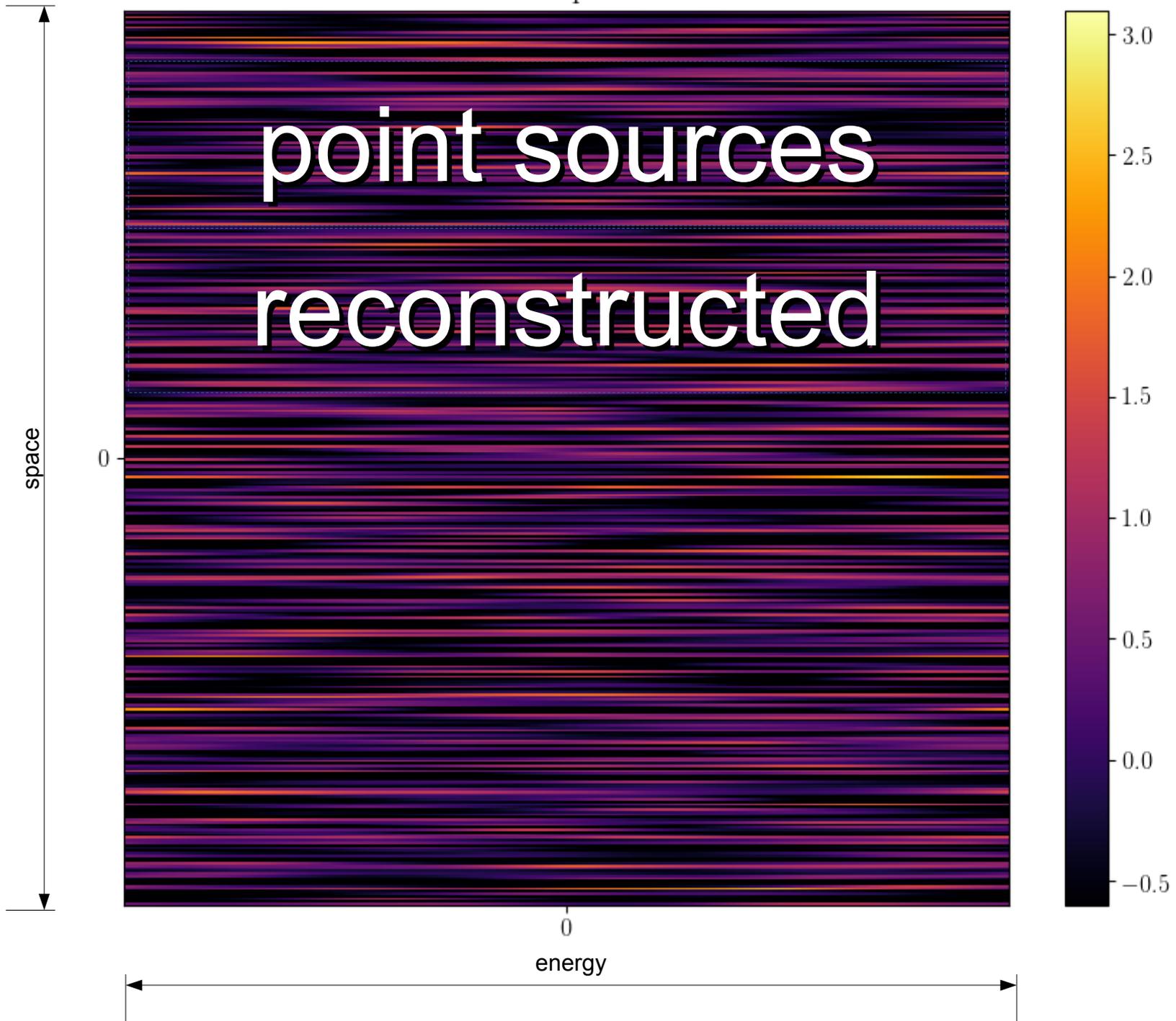
Uncertainty diffuse flux



Photon counts

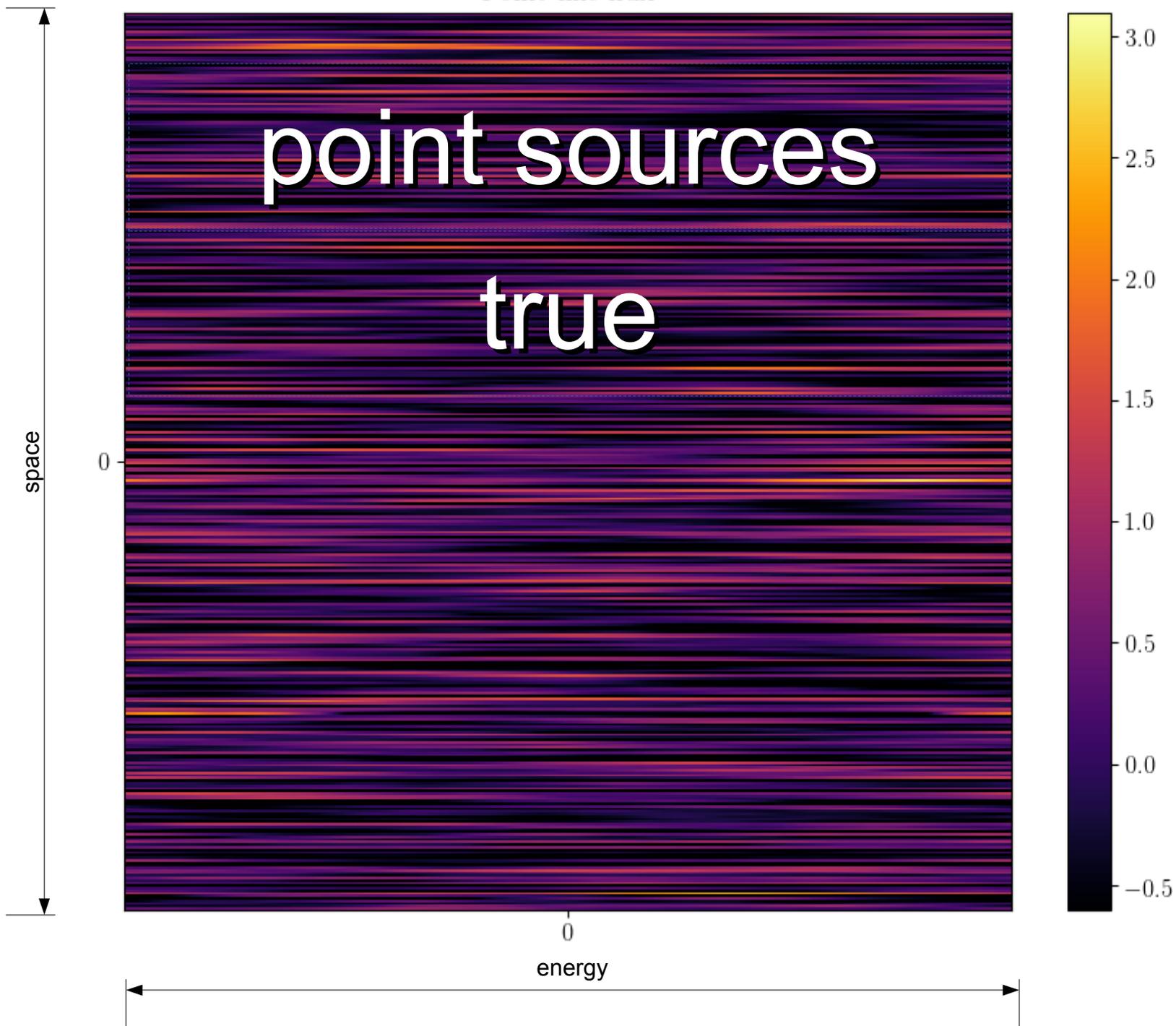


Reconstructed point-like flux



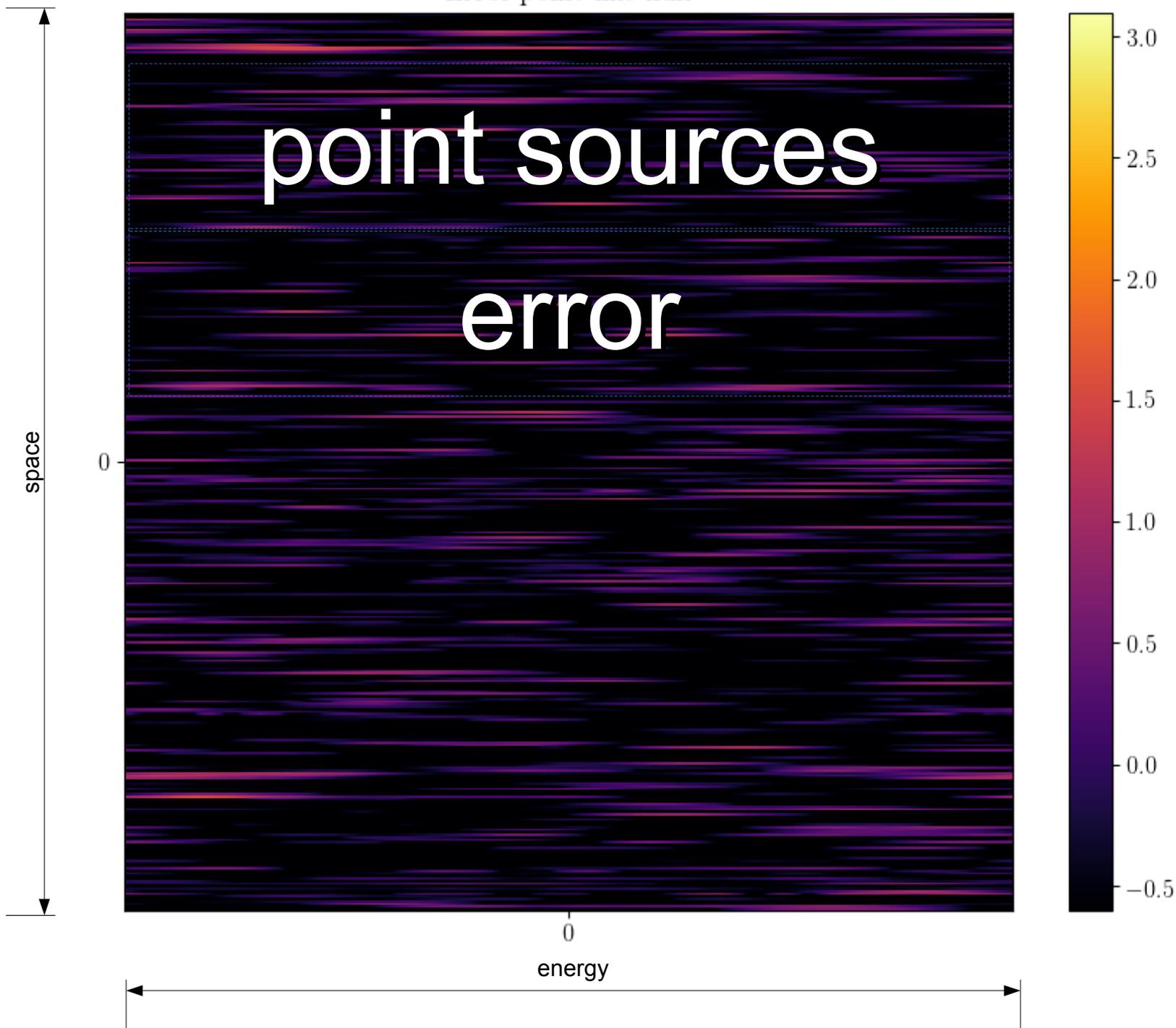
2018

Point-like flux



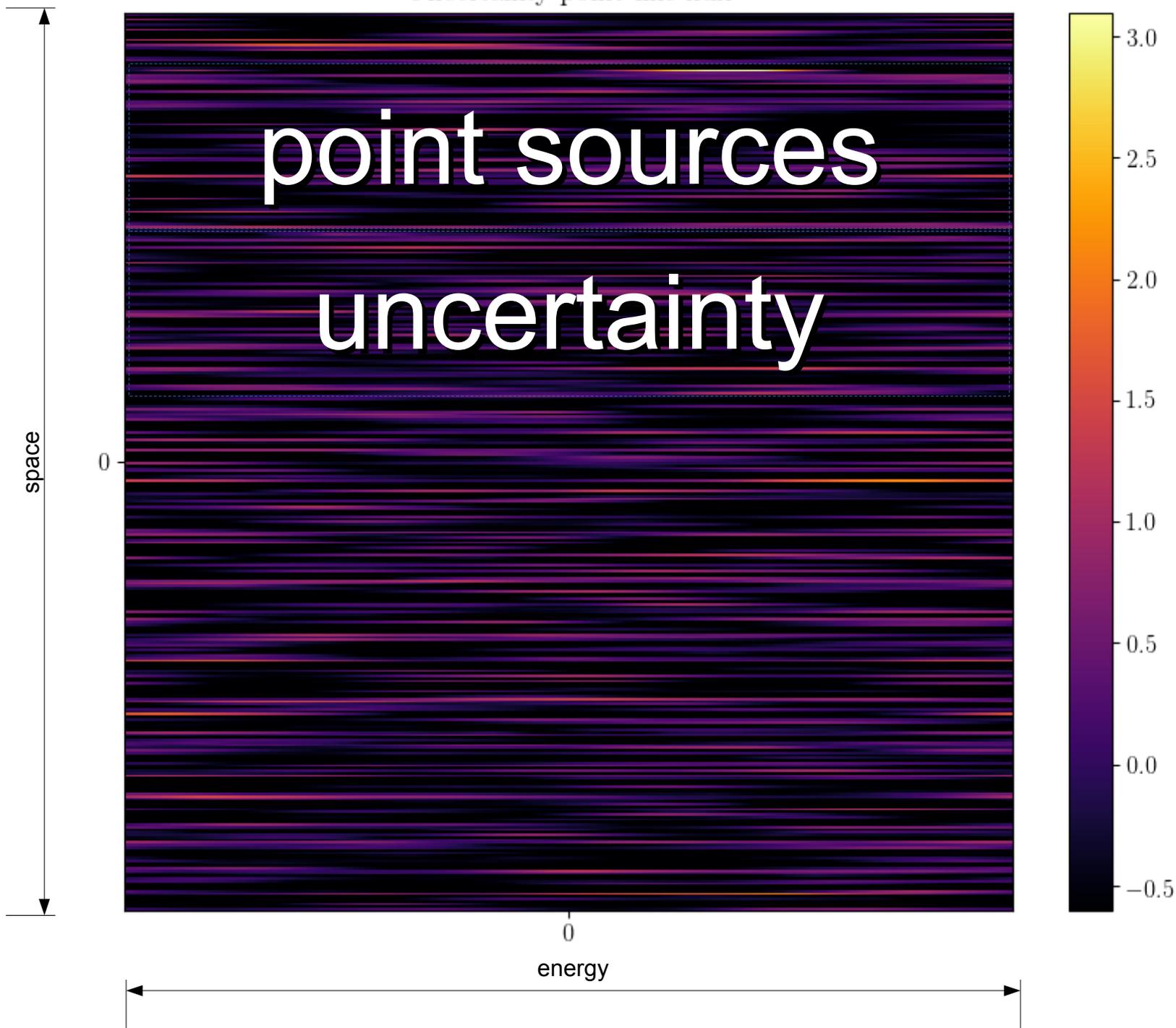
2018

Error point-like flux



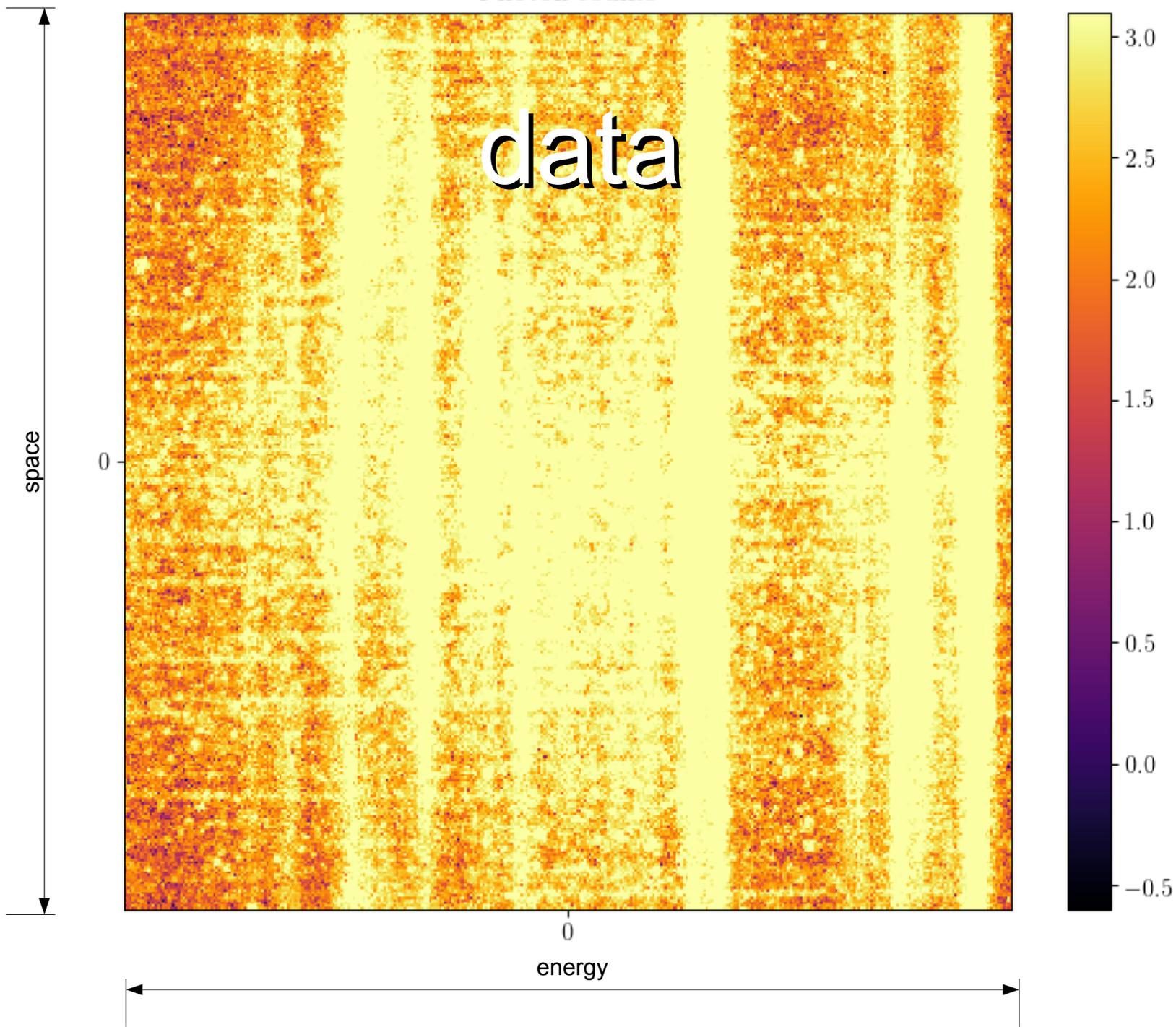
2018

Uncertainty point-like flux

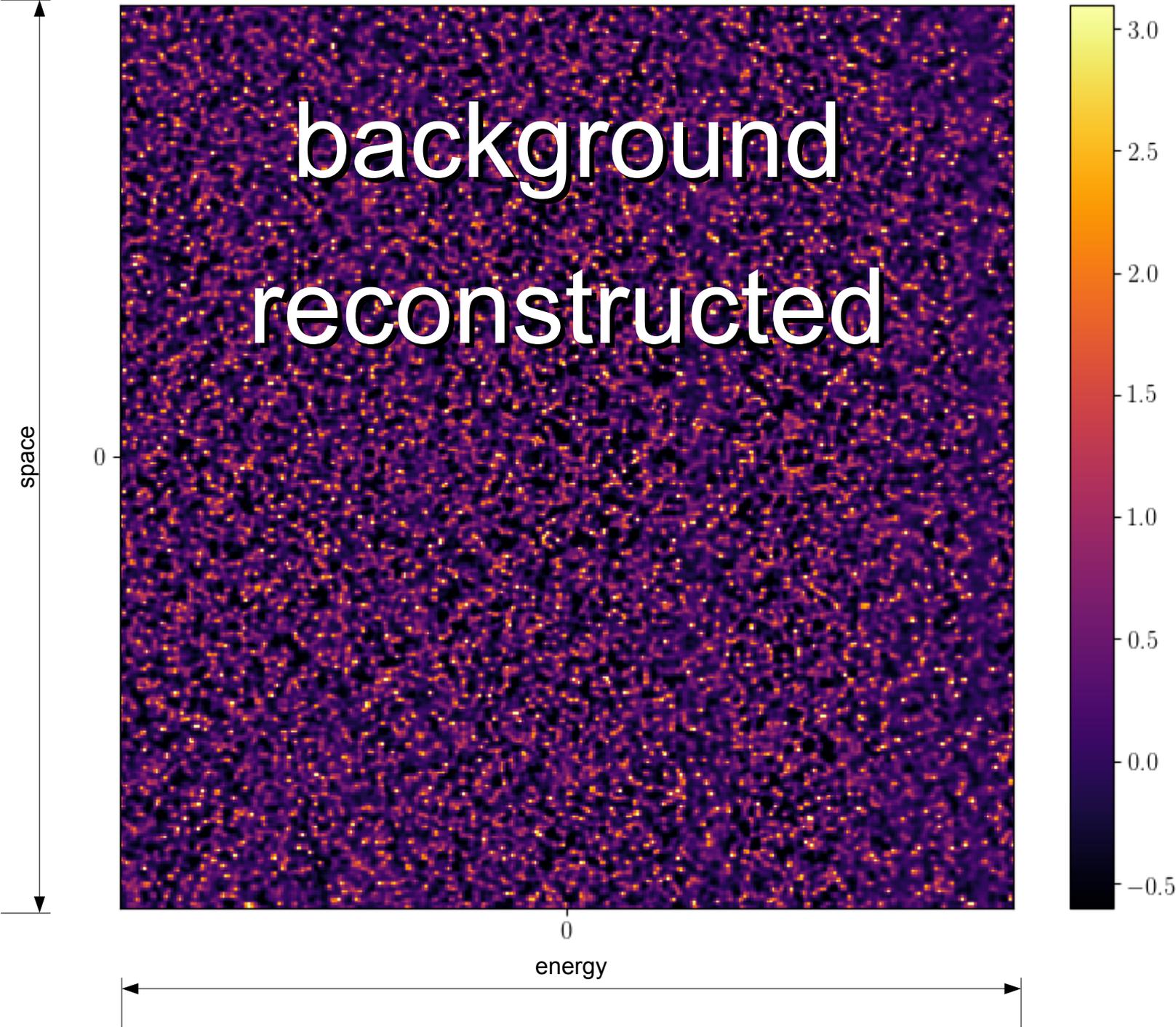


2018

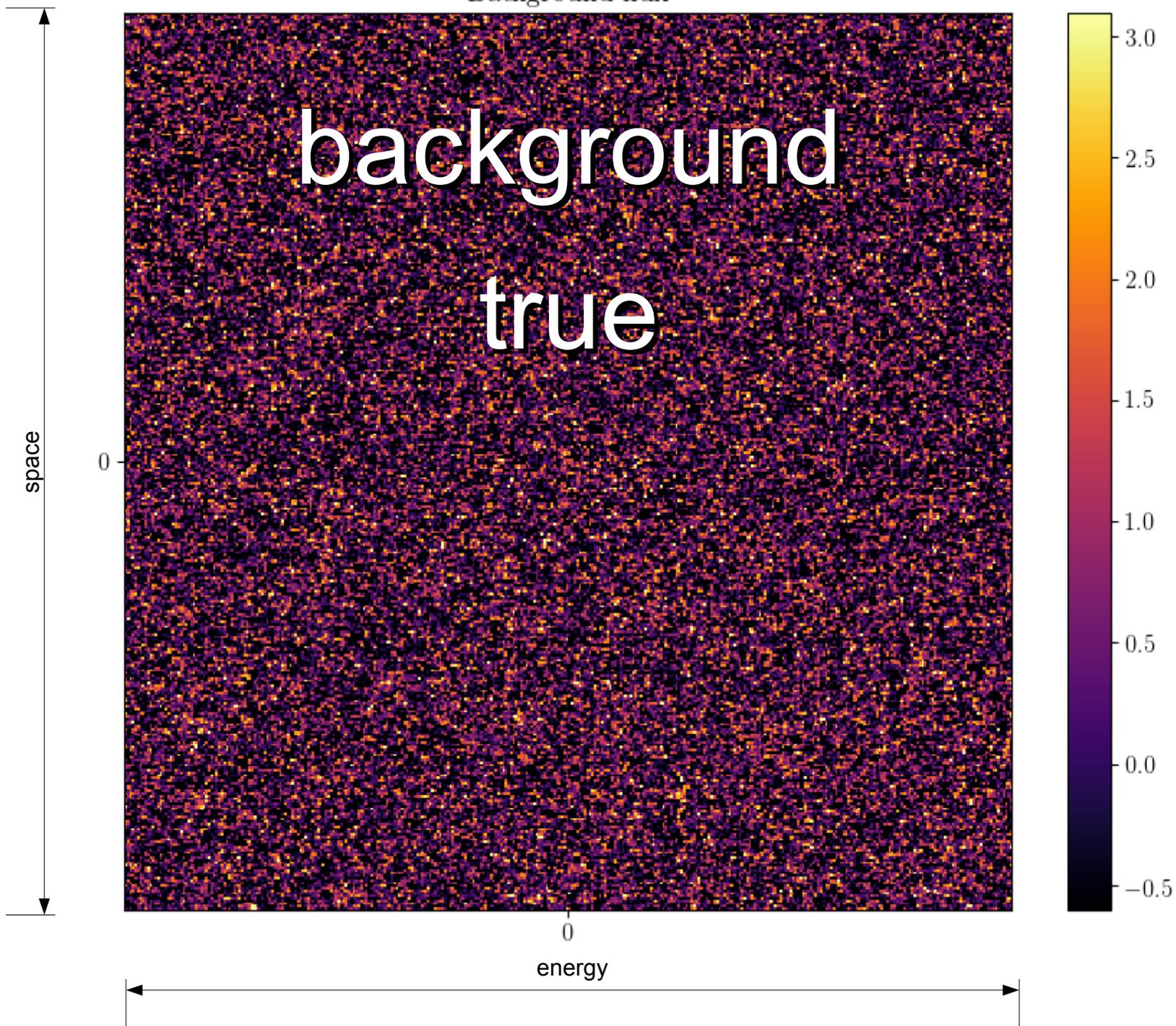
Photon counts



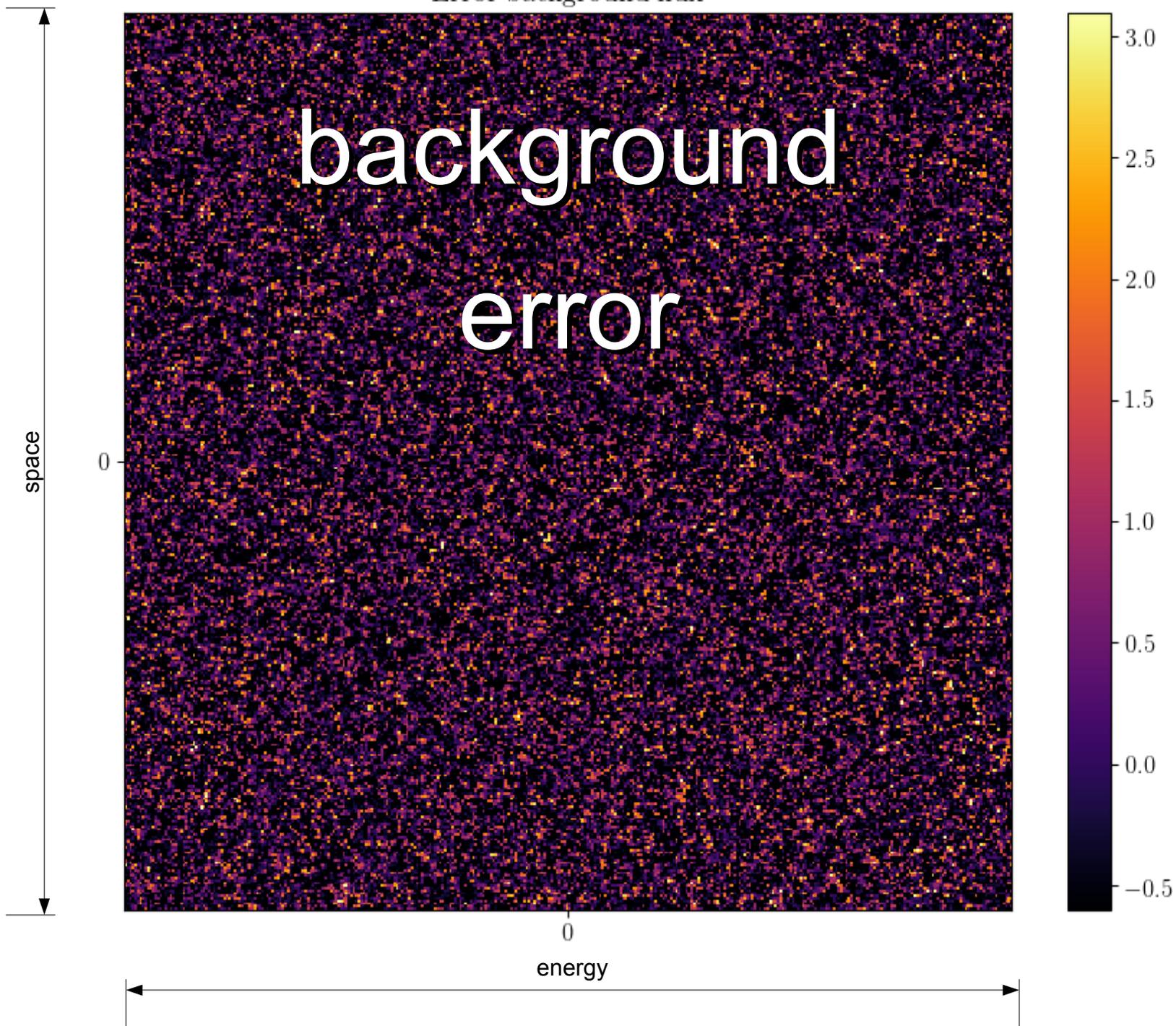
Reconstructed background flux



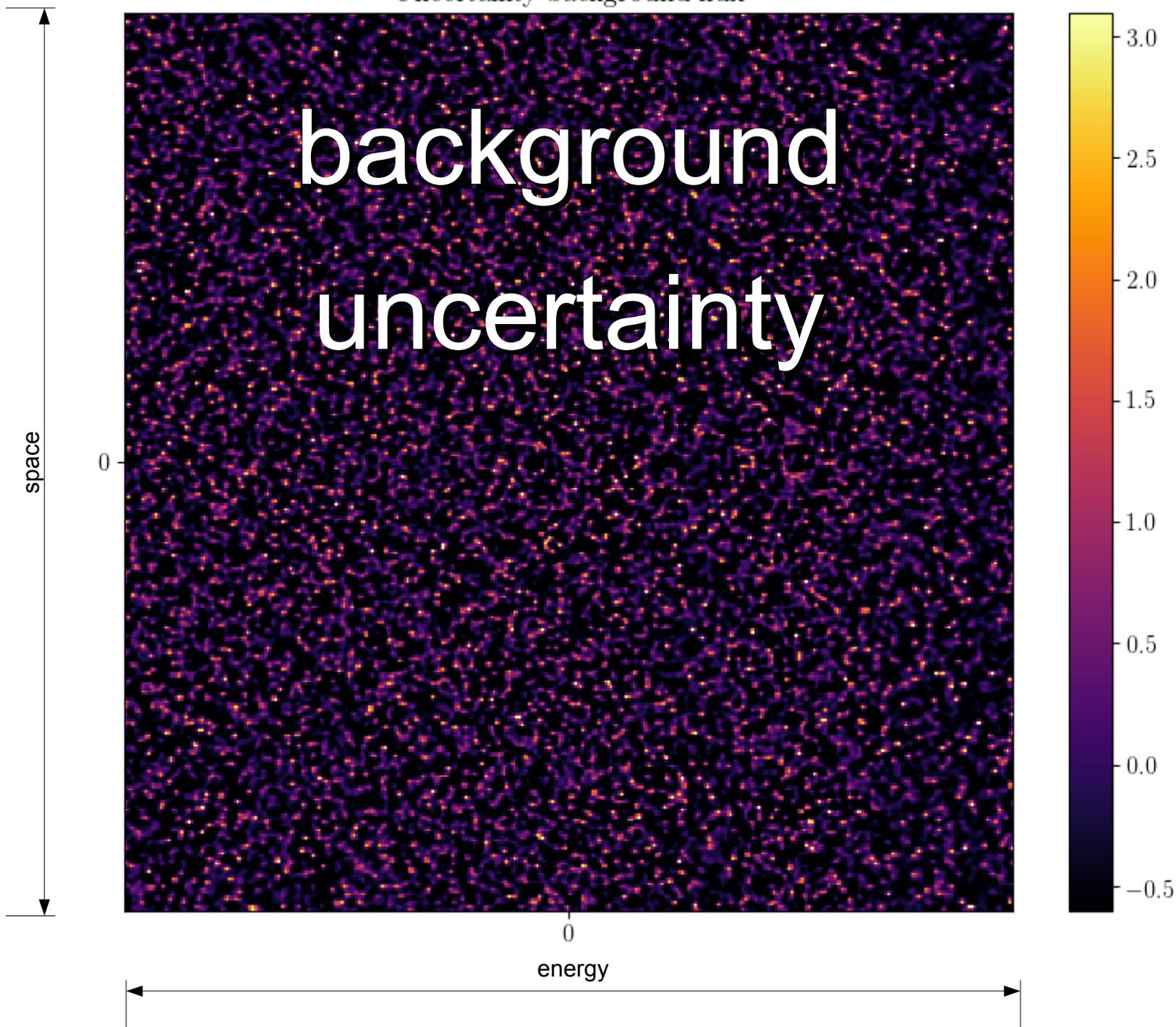
Background flux



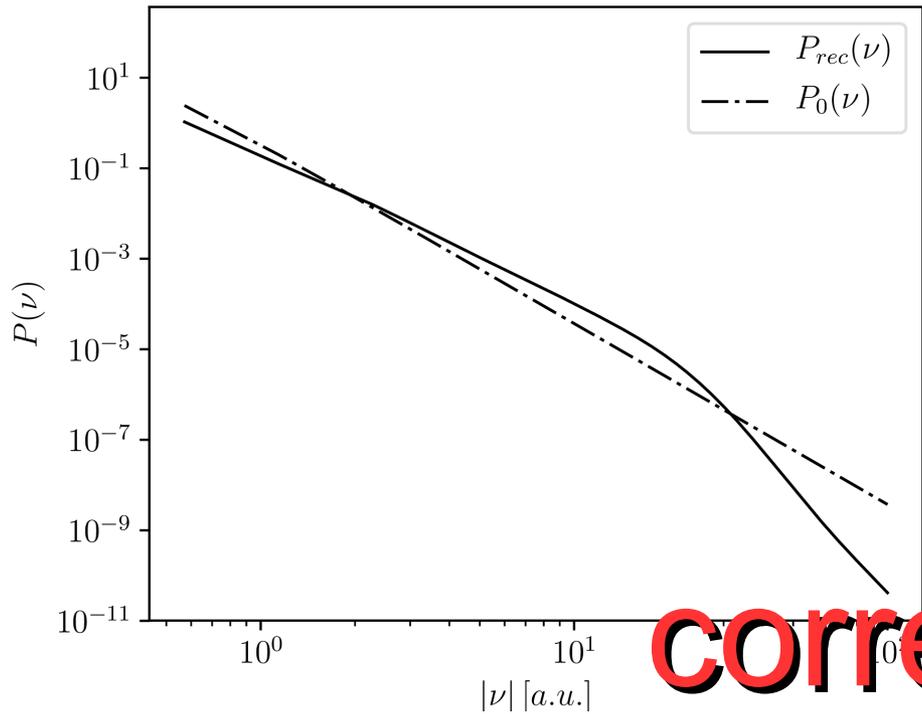
Error background flux



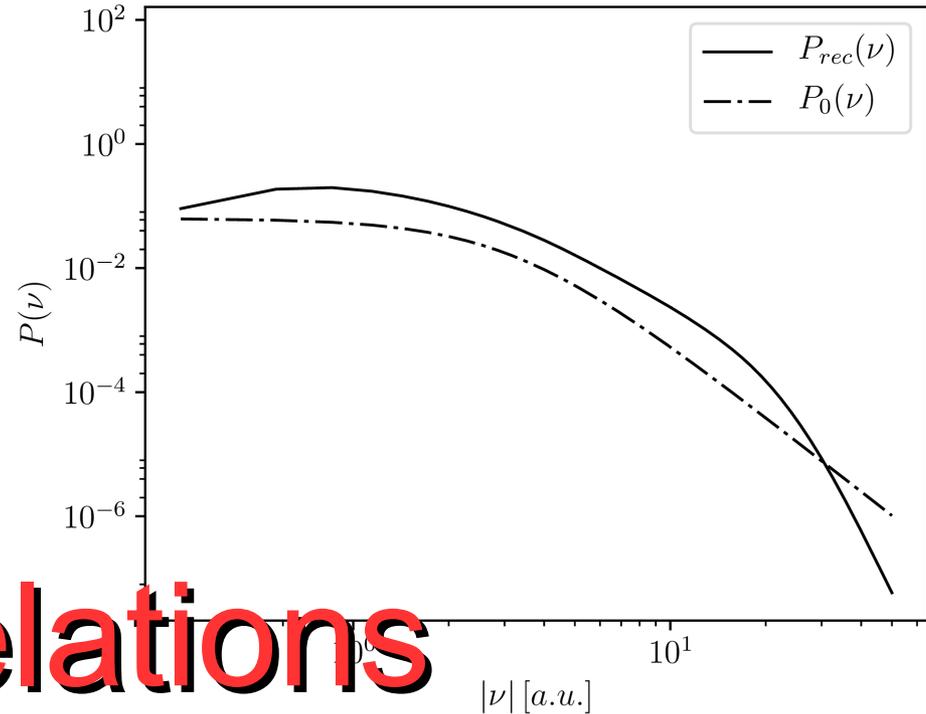
Uncertainty background flux



Spatial correlation of diffuse flux

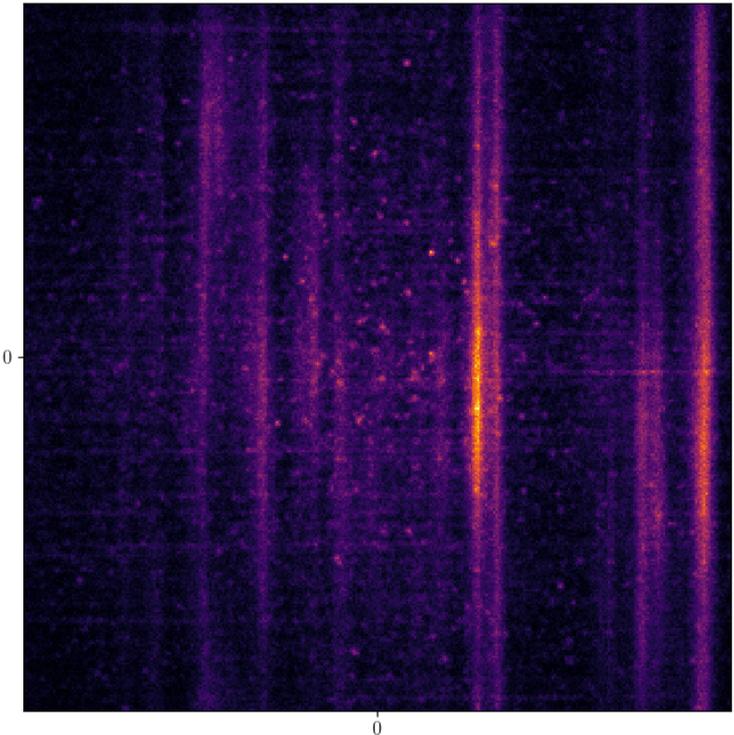


Energy correlation of diffuse flux

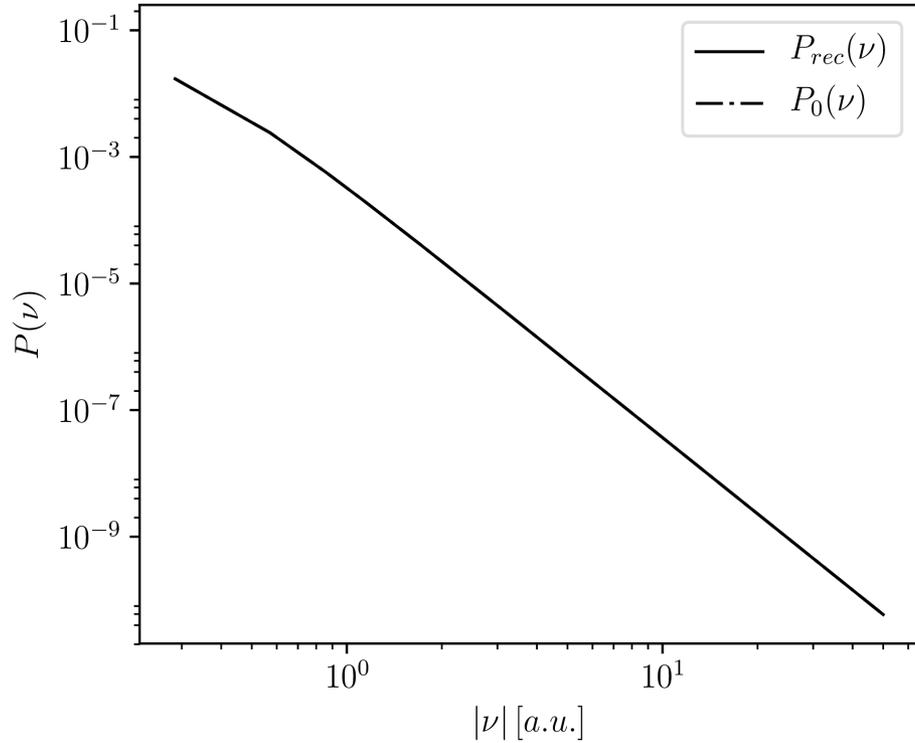


correlations

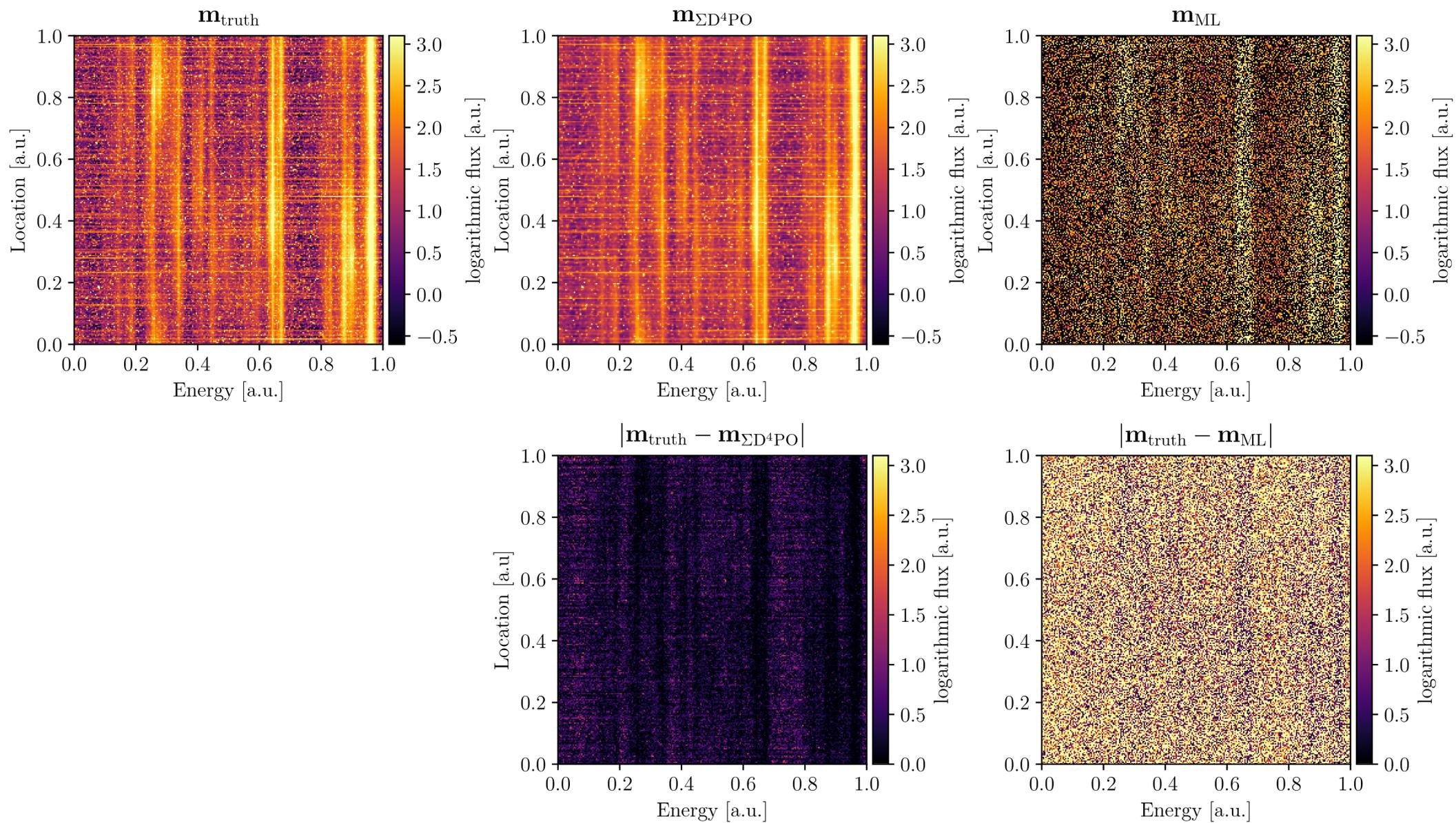
Raw photon counts



Energy correlation of point flux



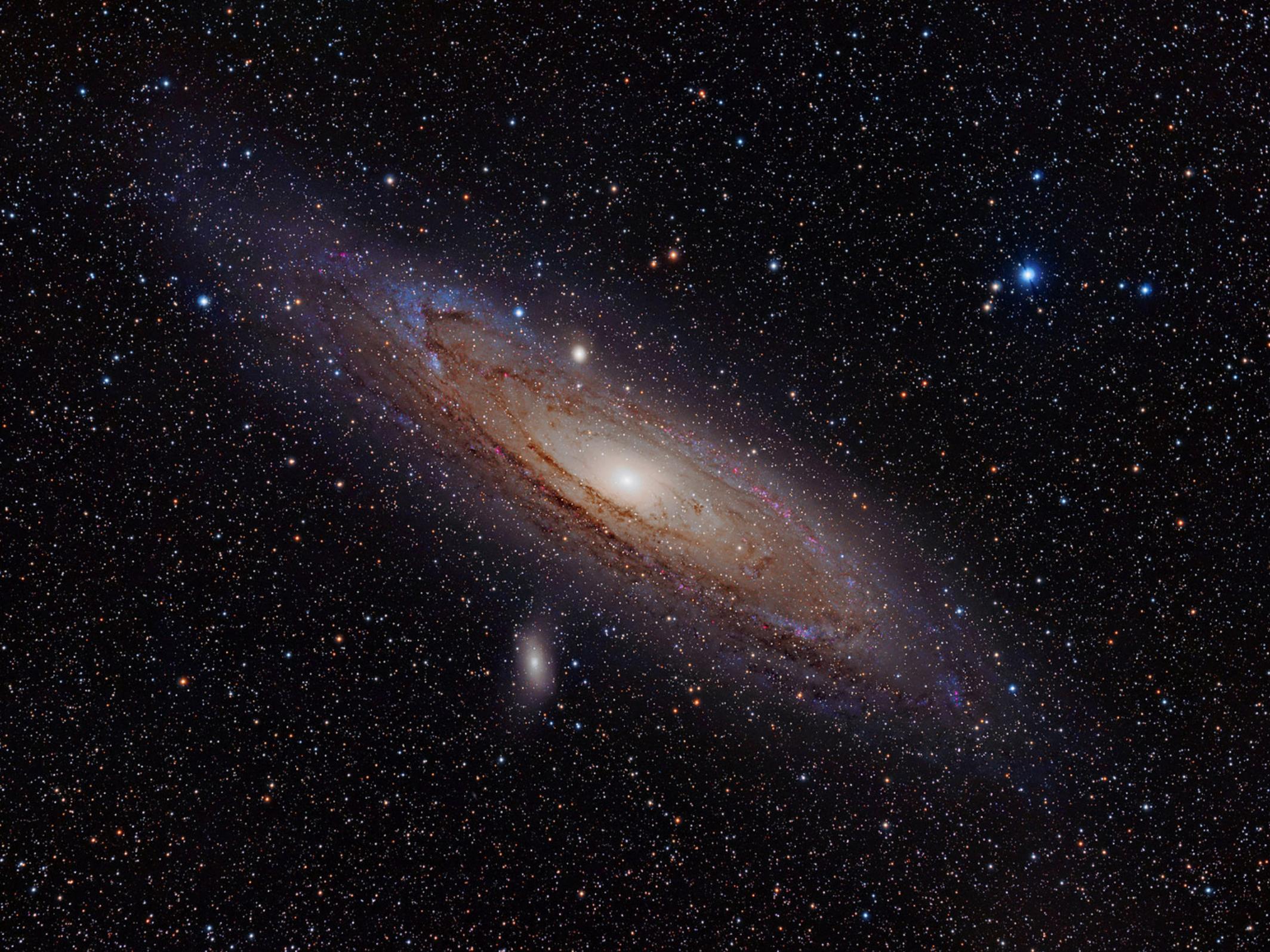
D⁴PO vs ML



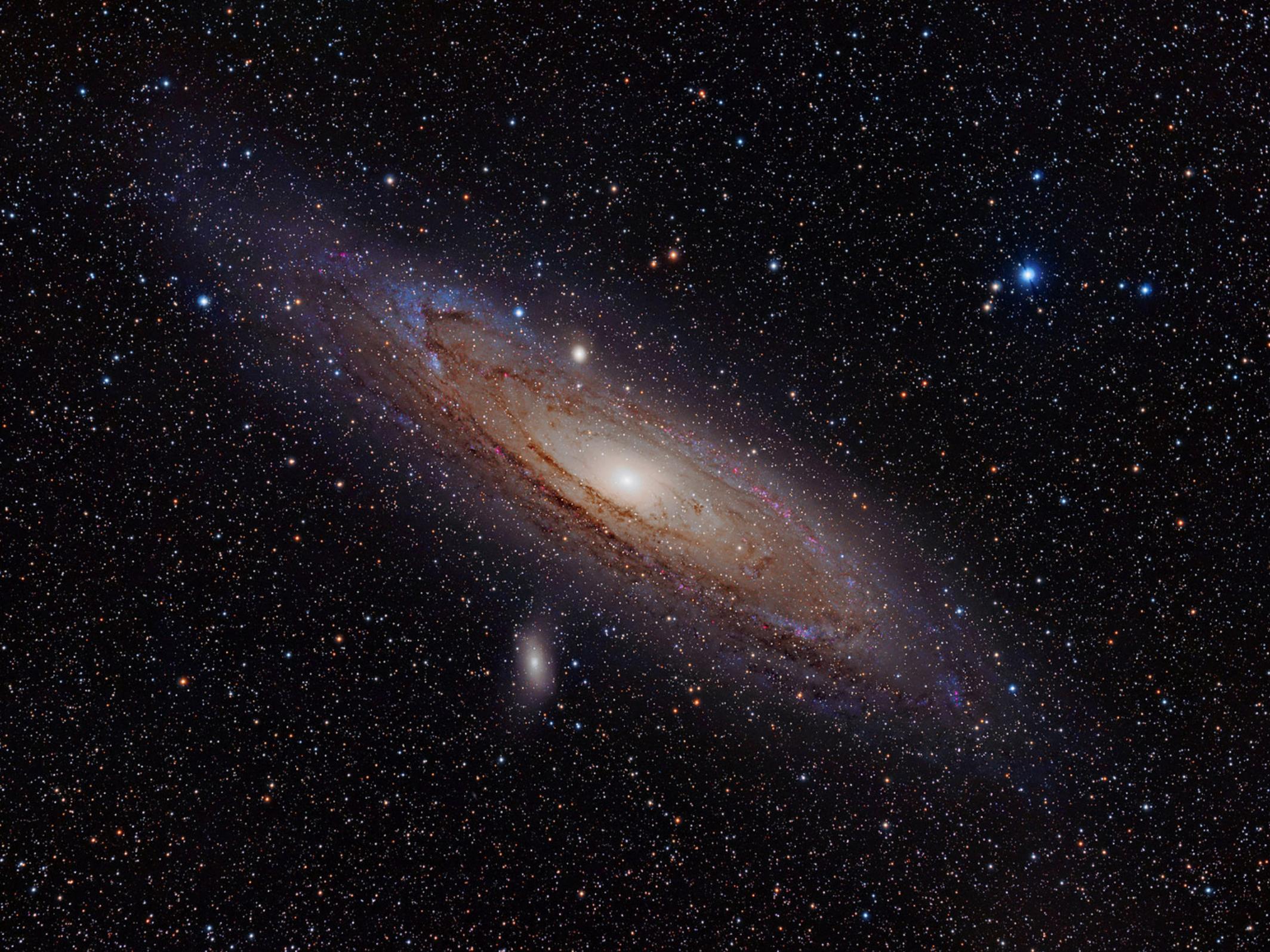
Glimpses into the future



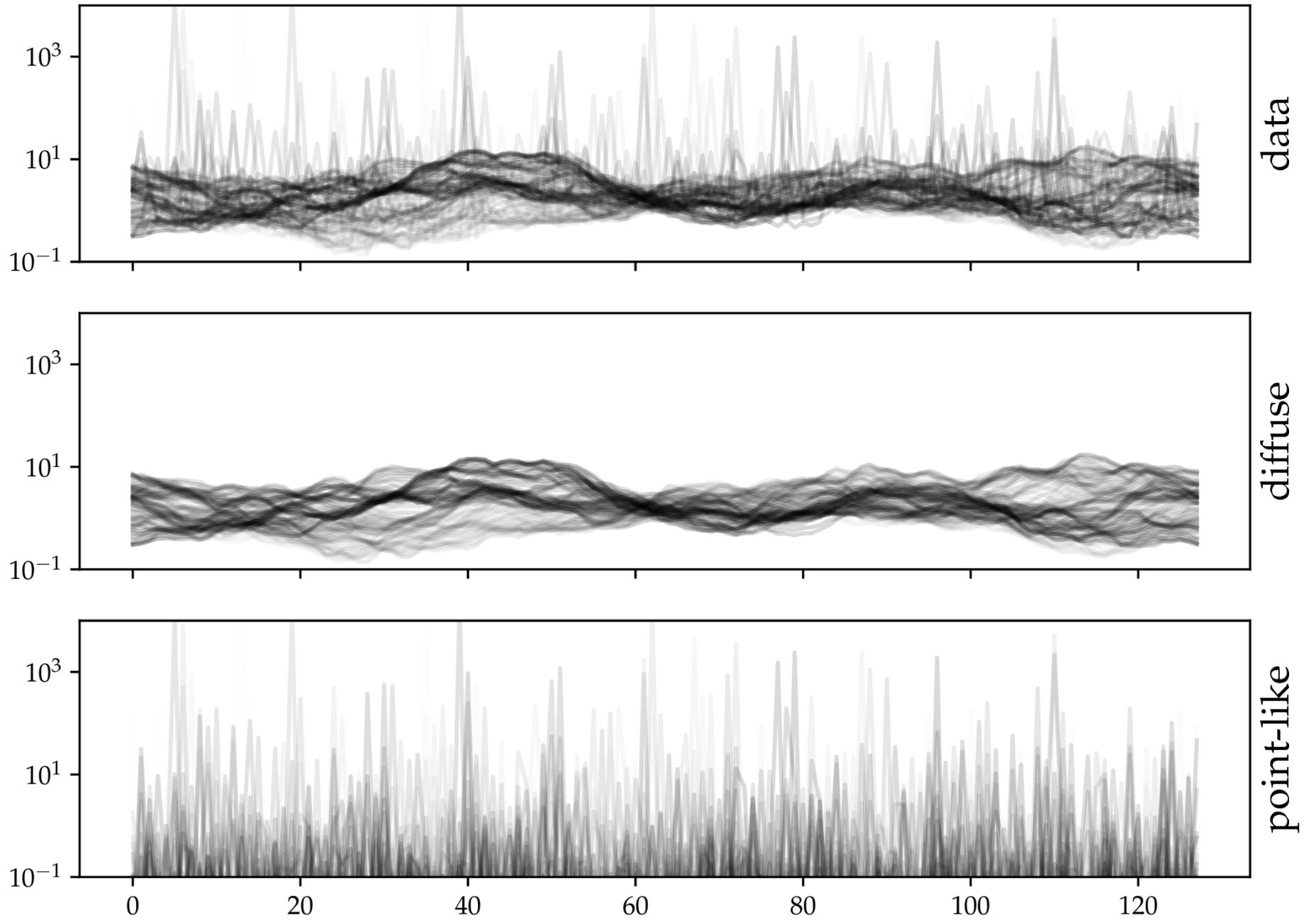
Jakob Knollmüller
starblade



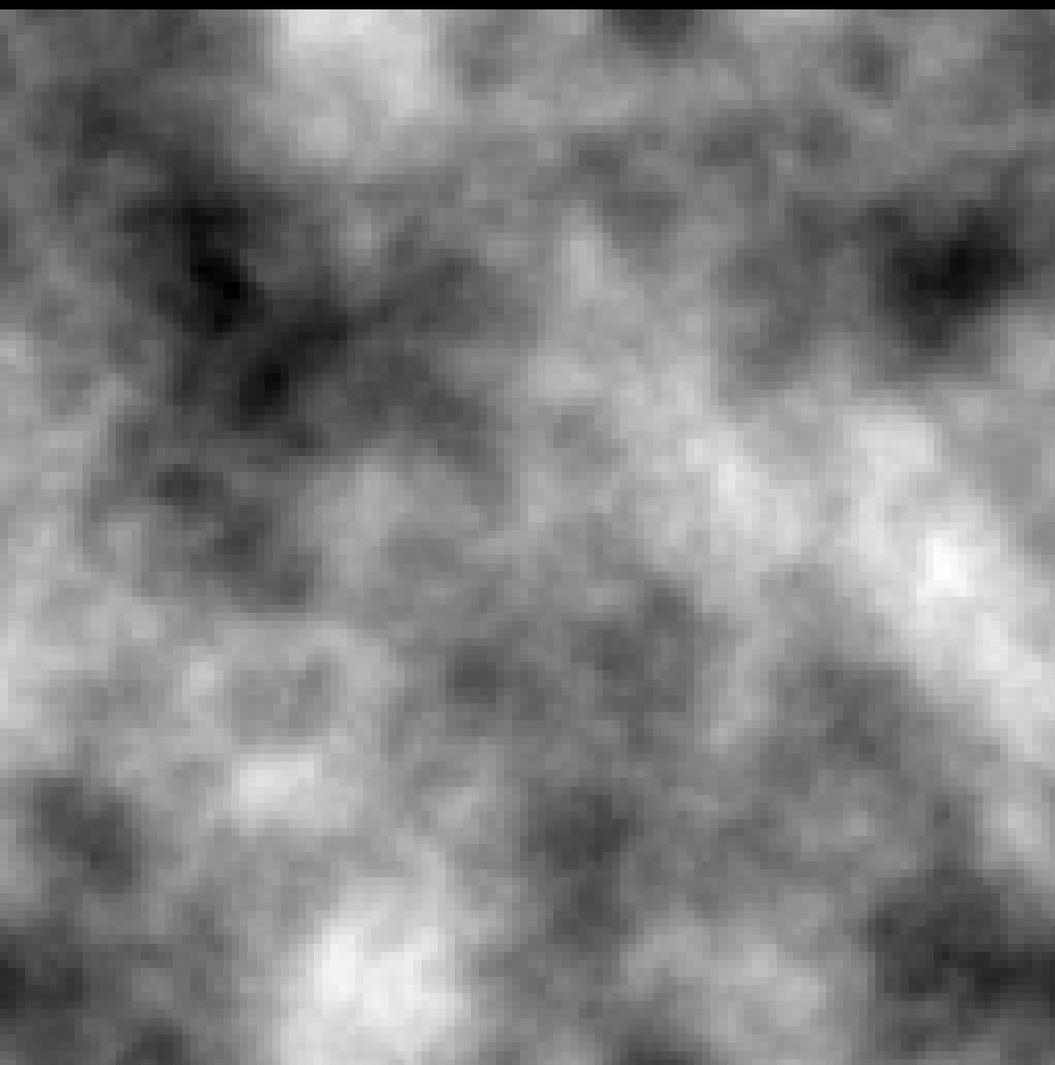




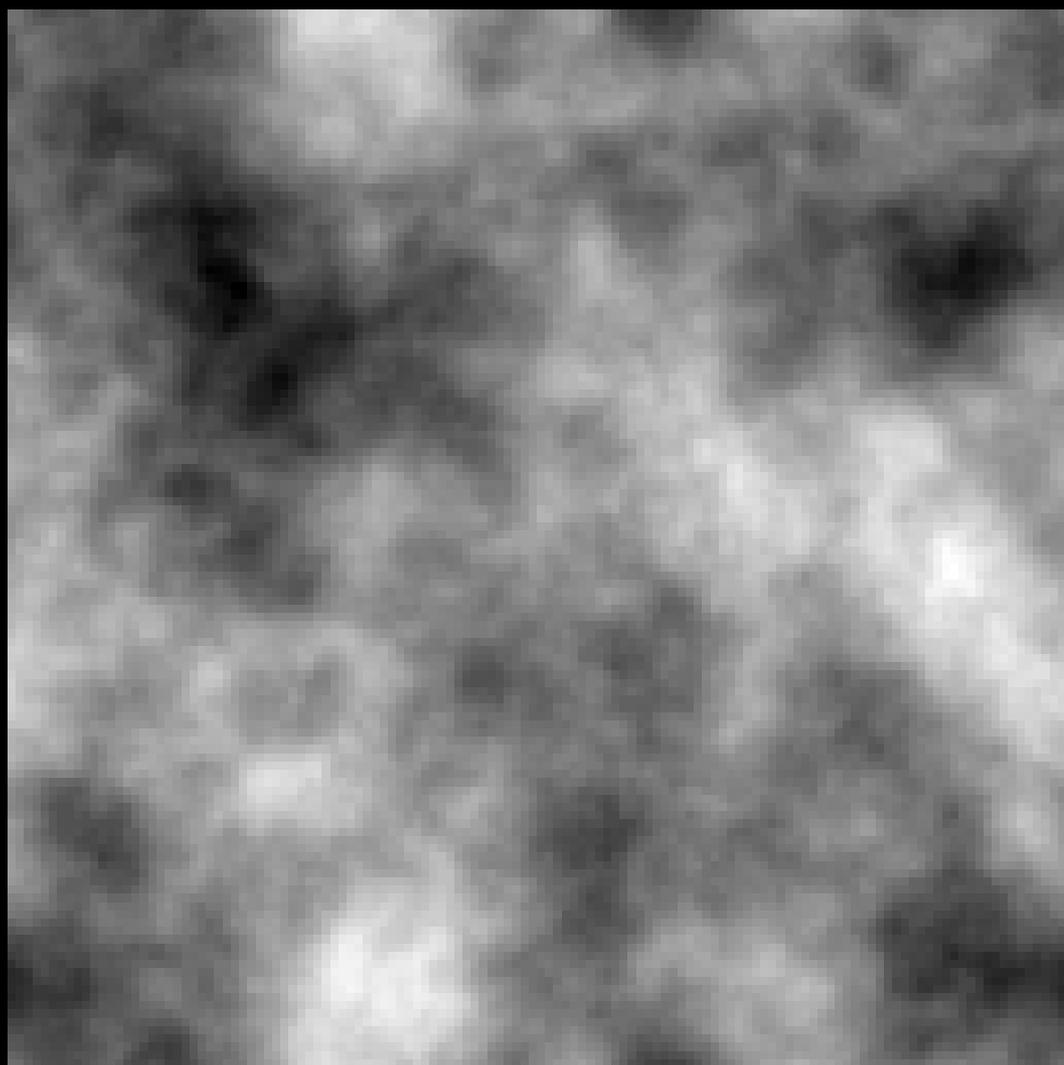
data and true components



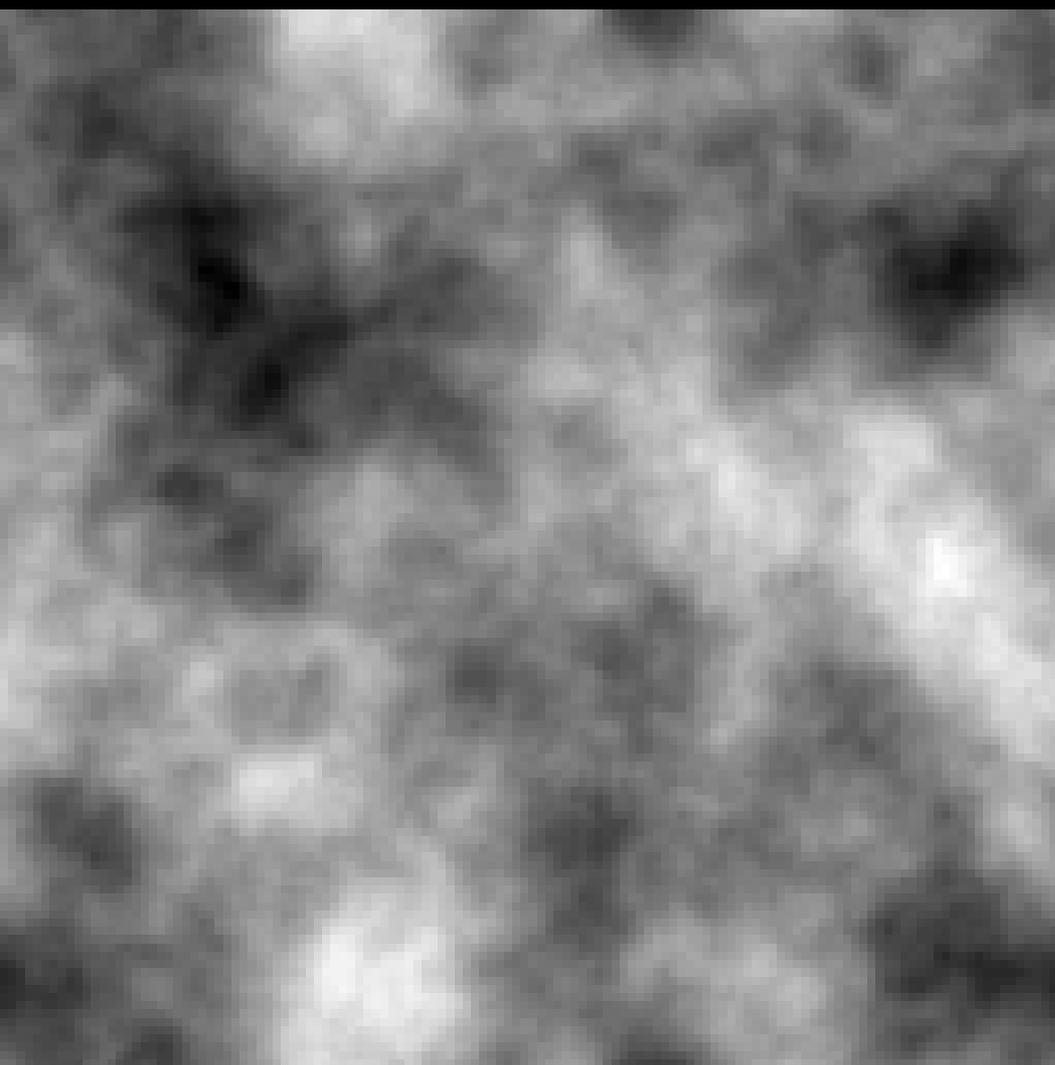
ground truth / starblade



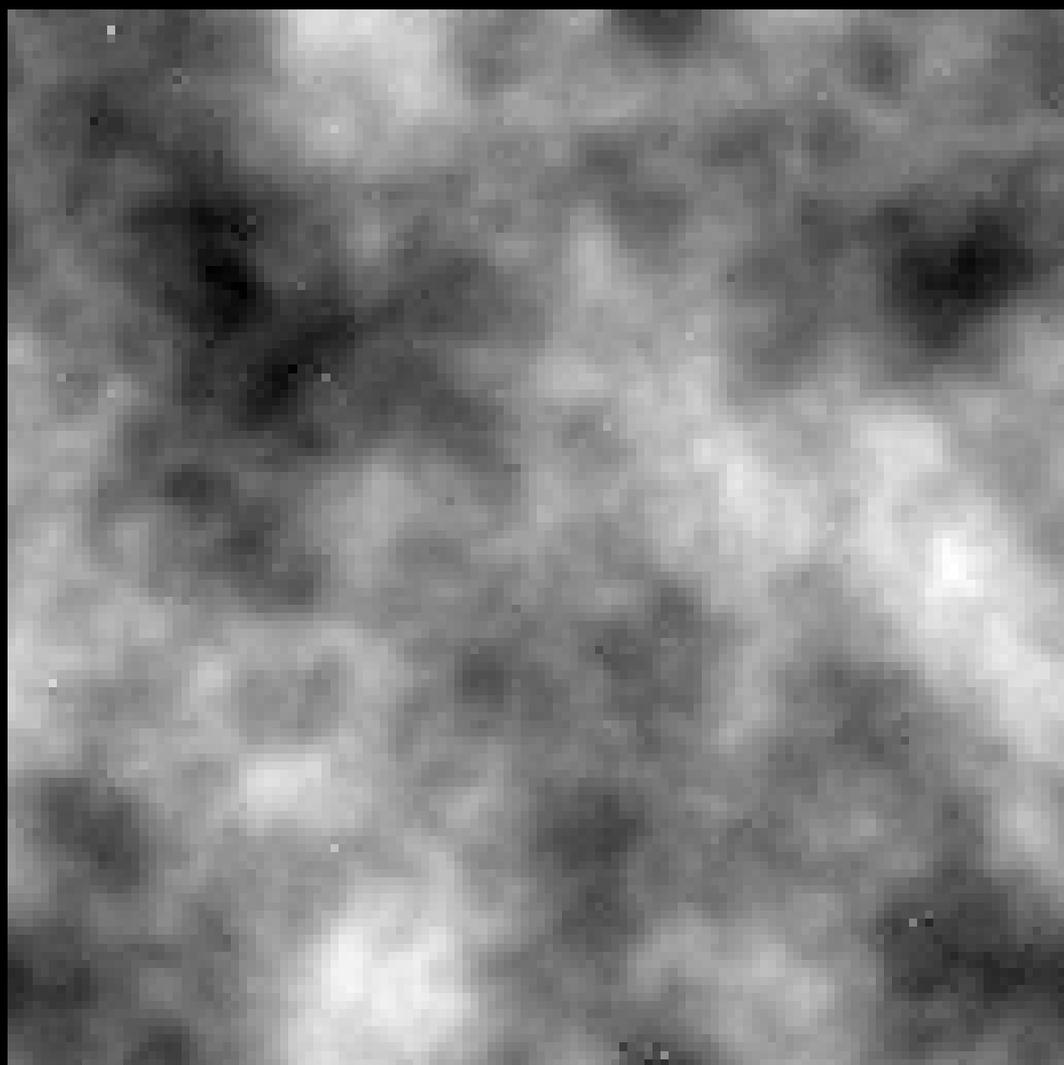
ground truth / autoencoder



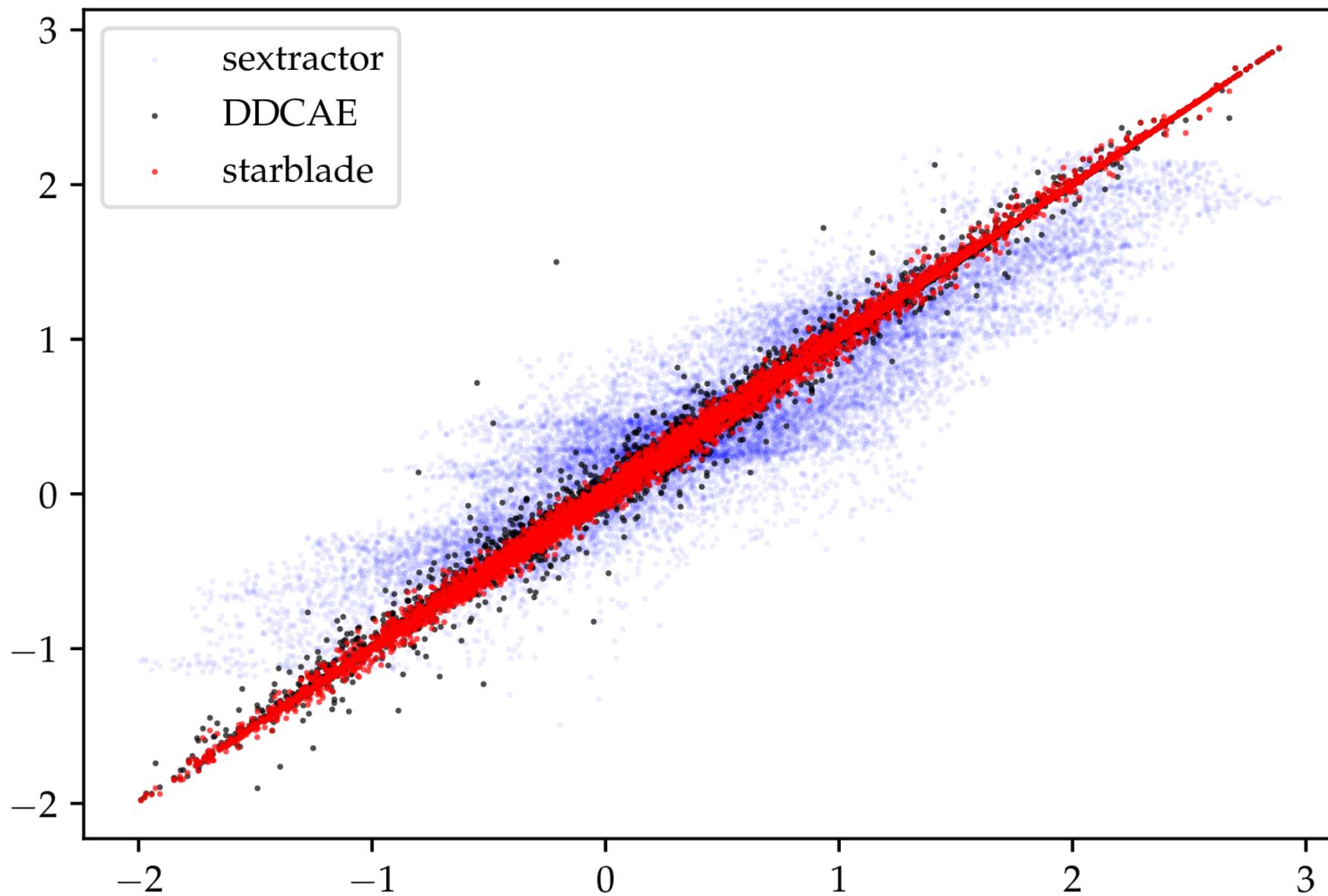
ground truth / starblade

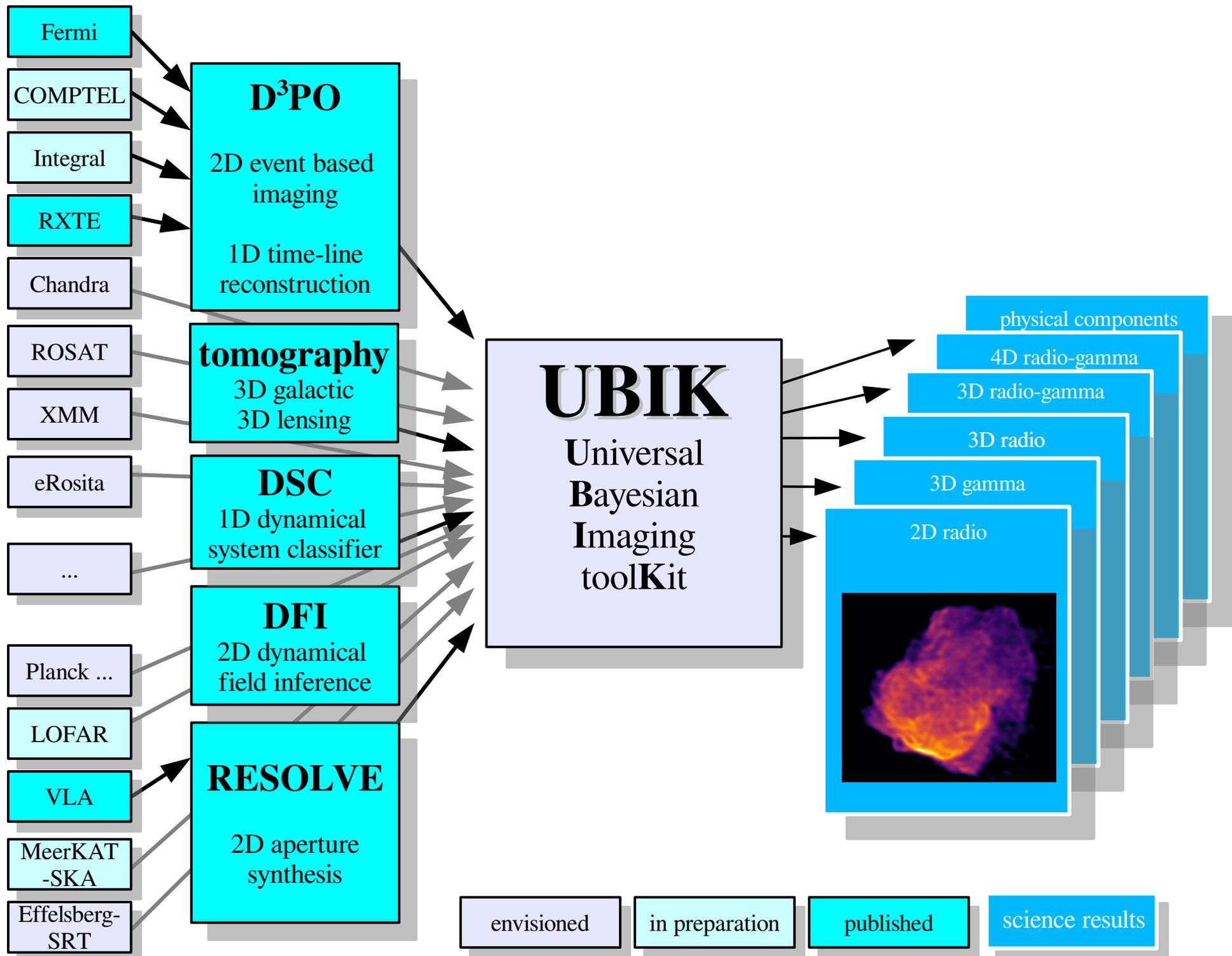


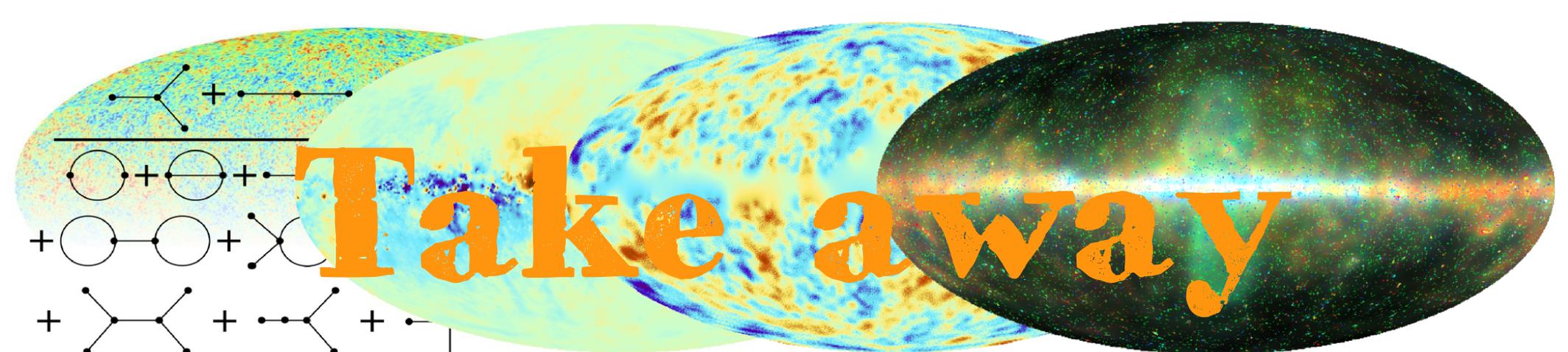
ground truth / autoencoder



diffuse truth vs DDCAE and starblade







Take away

Imaging goes inference

IFT

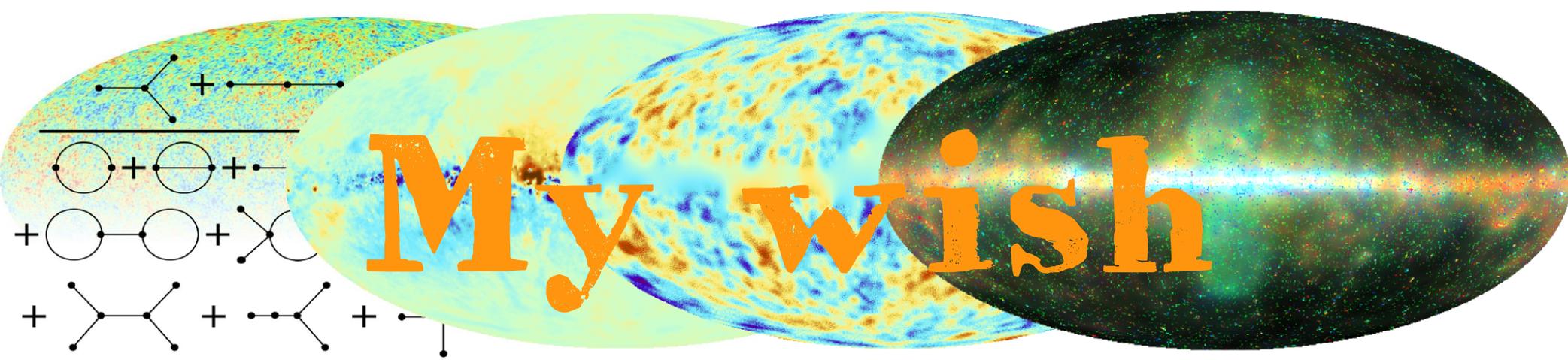
- information field theory

NIFTy

- numerical IFT in Python

UBIK

- multi-instrument & -dimension
imaging & tomography



My wish

If you build an instrument please provide

- **a fast response code**

R : sky \rightarrow data

- **its derivative**

$R' = dR/dsky$: (sky, dsky) \rightarrow ddata

- **its adjoint**

R'^{\dagger} : (sky, ddata) \rightarrow dsky

and UBIK will be your friend.

Thank you!

