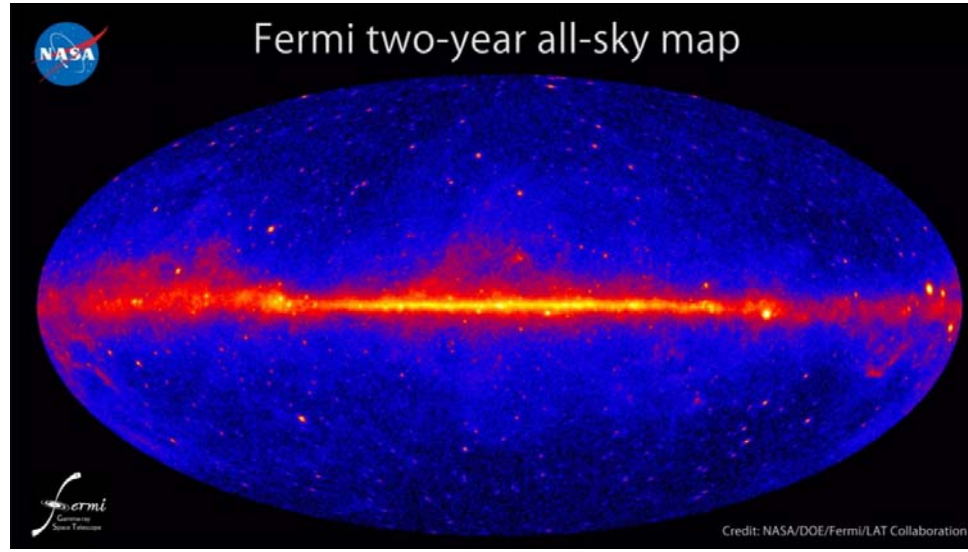
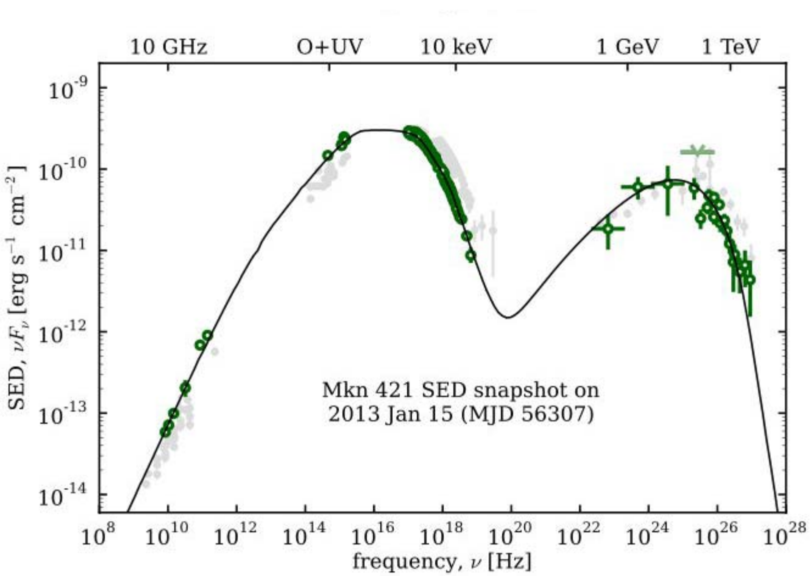


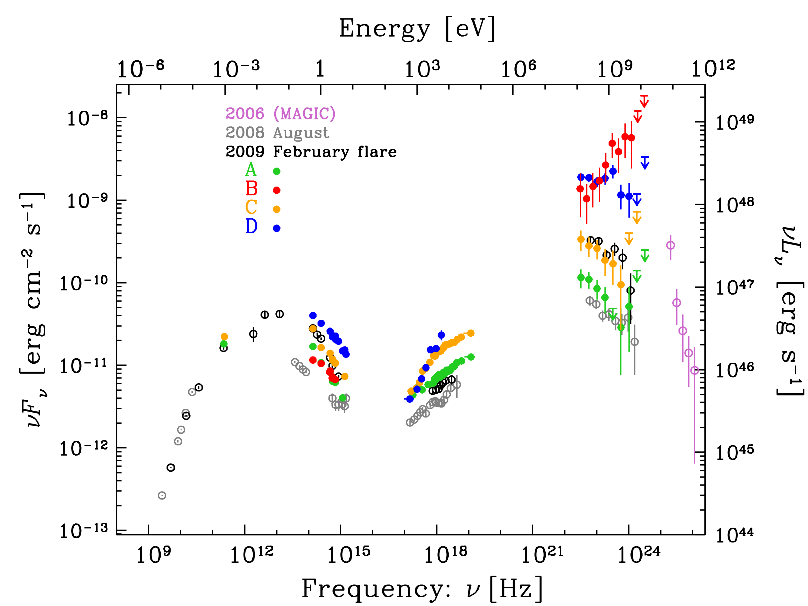
Joint Fermi / VHE / NuSTAR / Swift observations of blazars

Greg Madejski
SLAC/KIPAC
w/ David Sanchez, Krzysztof Nalewajko, David Paneque, Mislav Balokovic, Amy Furniss, Berrie Giebels, Meg Urry, Marek Sikora, and the members of the Fermi, Veritas, MAGIC, H.E.S.S., & NuSTAR teams

Small dollop of blazar phenomenology

- Blazar spectra show two broad “humps” one peaking in the far IR – to – soft X-rays, another peaking in the MeV – Gev γ-ray range, sometimes extends to the TeV VHE γ-ray regime
- The low-energy hump emission (radio, opt.) – synchrotron emission of <u>plasma consisting of relativistic particles</u> accelerated in the jet
- The high-energy peak - inverse Compton process, by the same electrons that produced the synchrotron hump
- Volume can be estimated from variability time scales



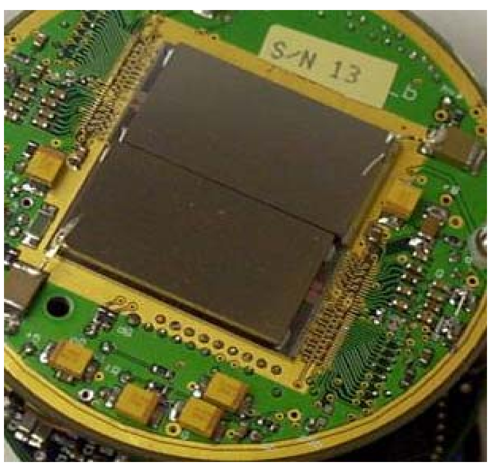


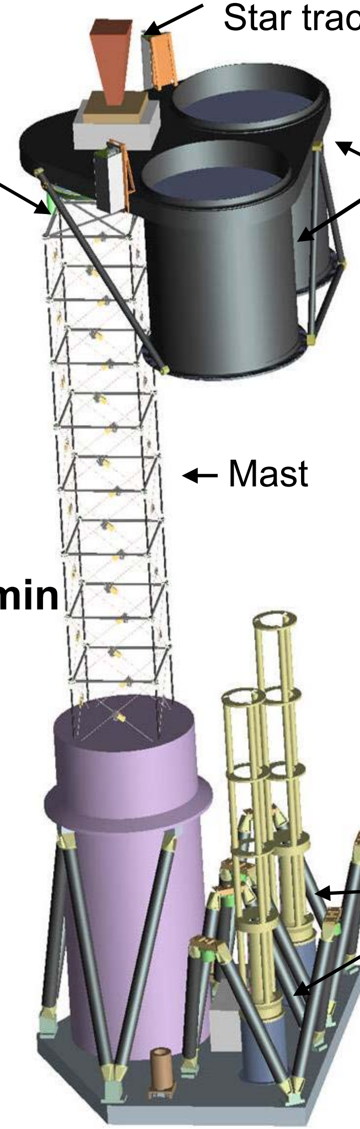
Very bright γ-ray blazar 3C279

Friend of Fermi: Hard X-ray satellite NuSTAR

- Launched in June 2012; led by Caltech
- Two identical co-aligned grazing incidence hard X-ray telescopes:
  - Two multilayer coated segmented glass optics
  - Actively shielded solid state CdZnTe pixel detectors 10 meters away
- Energy bandpass 3 – 80 keV

Point spread function: ~ 1 arc min
Energy resolution: ~ 1 keV

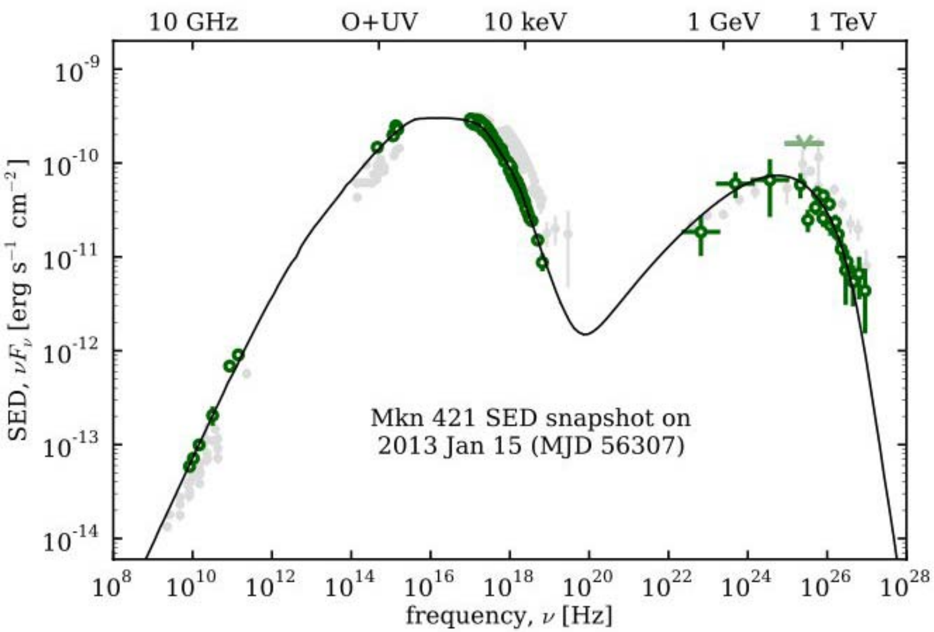


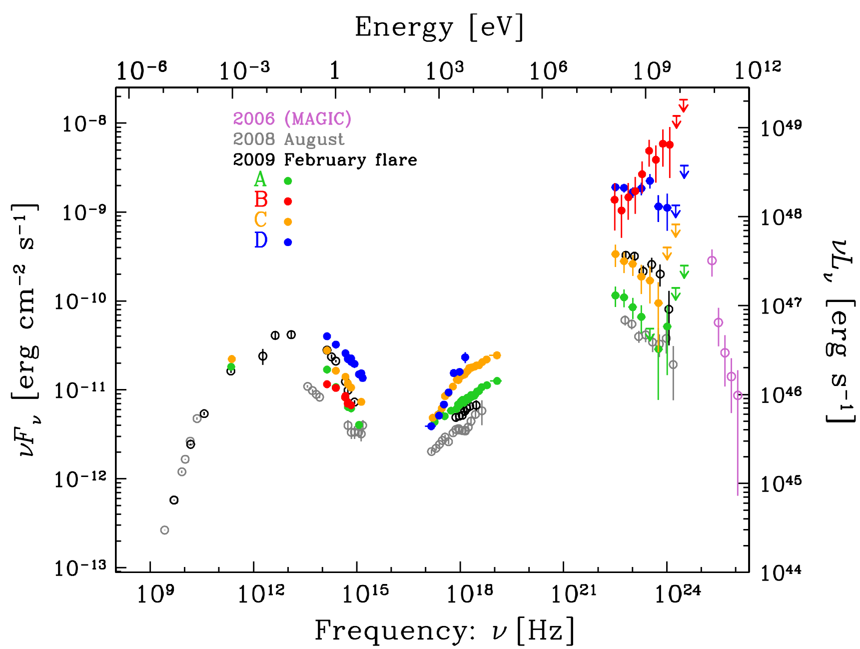


Why are hard X-rays important for blazar studies?

- The hard X-ray band is the intersection of the "tail end" of the synchrotron emission, and the "onset" of the inverse Compton hump
- The "onset" of the inverse Compton peak samples the low-energy particle population in the relativistic plasma - total particle content in the jet (low energy particles are most numerous)

* For the future: X-ray polarization is crucial to verify this picture





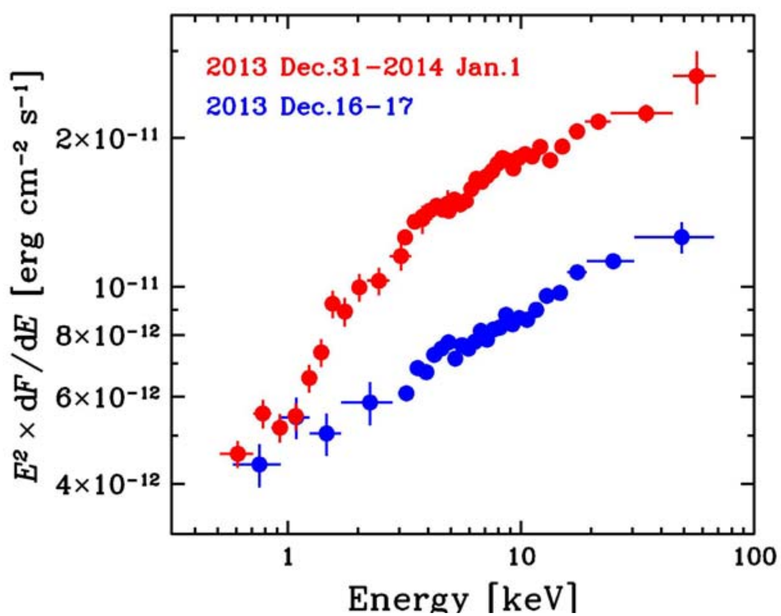
Very bright γ-ray blazar 3C279

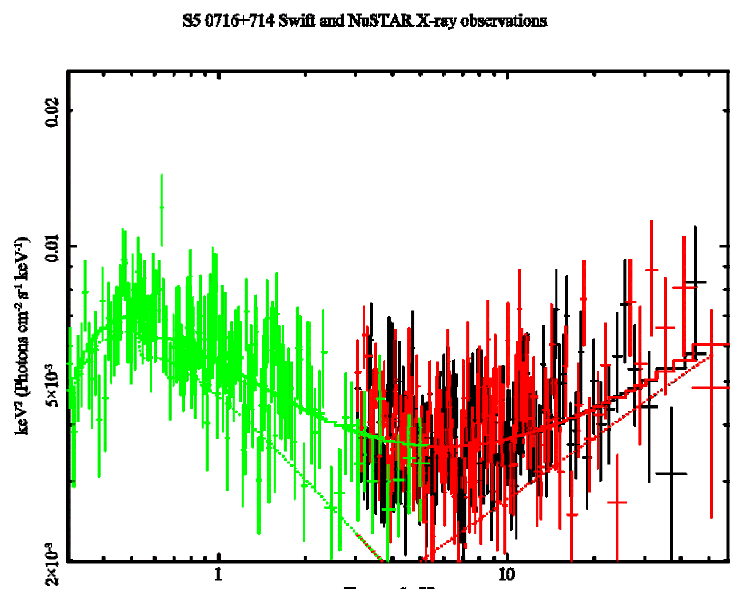
NuSTAR observed ~ 30 blazars

- Blazars were a part of the initial motivation for NuSTAR: to provide a multi-wavelength context to work together with Fermi LAT
- A few were selected for monitoring programs, plus several particularly interesting ones were selected for one or two pointings
- The three famous TeV-emitting objects selected for radio-through-VHE monitoring were Mkn 501, Mkn 421, and PKS 2155-304
- NuSTAR also observed several FSRQs – including high-z objects; implications on formation of black holes in the early Universe
- NuSTAR is poised to observe several more in flaring states

General hard X-ray properties of NuSTAR blazars

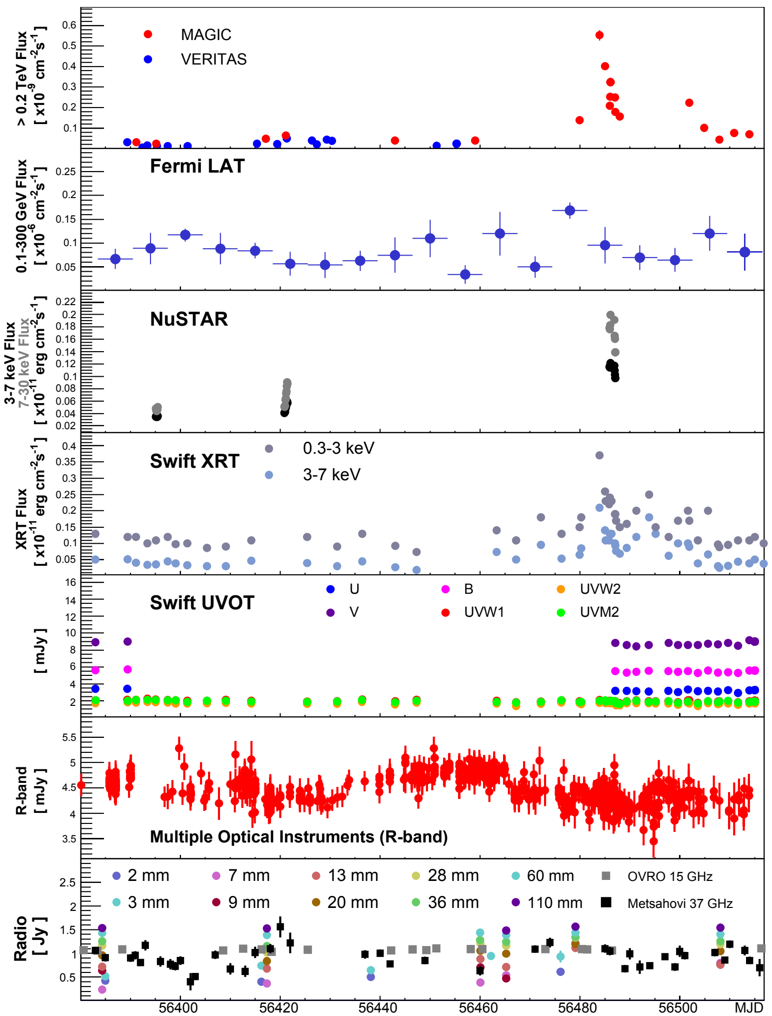
- Very “top-level” summary: best results obtained jointly with Swift, XMM-Newton
- NuSTAR measures 3 – 80 keV spectra of FSRQs to be generally quite hard, they are rising in E x F(E), Γ ~ 1.5, then they often flatten to Γ ~ 2 (example: 3C279 below)
- HBL-type BL Lac objects generally have soft spectra in the NuSTAR band, falling in the E x F(E), Γ ~ 2.5 – 3
- In a few cases, one sees the break in the NuSTAR band: example is S5 0716+714



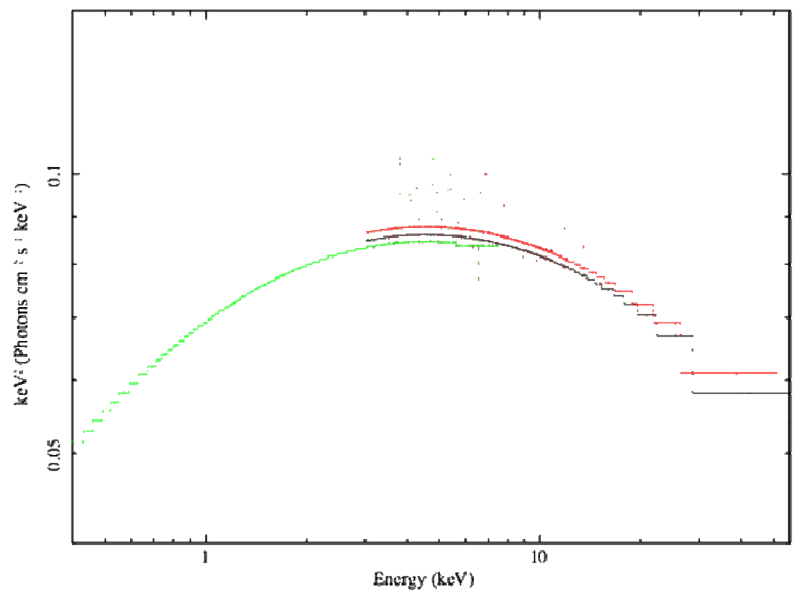


Mkn 501 observed by *NuSTAR*

Campaigns including NuSTAR were conducted generally in the context of MW campaigns…

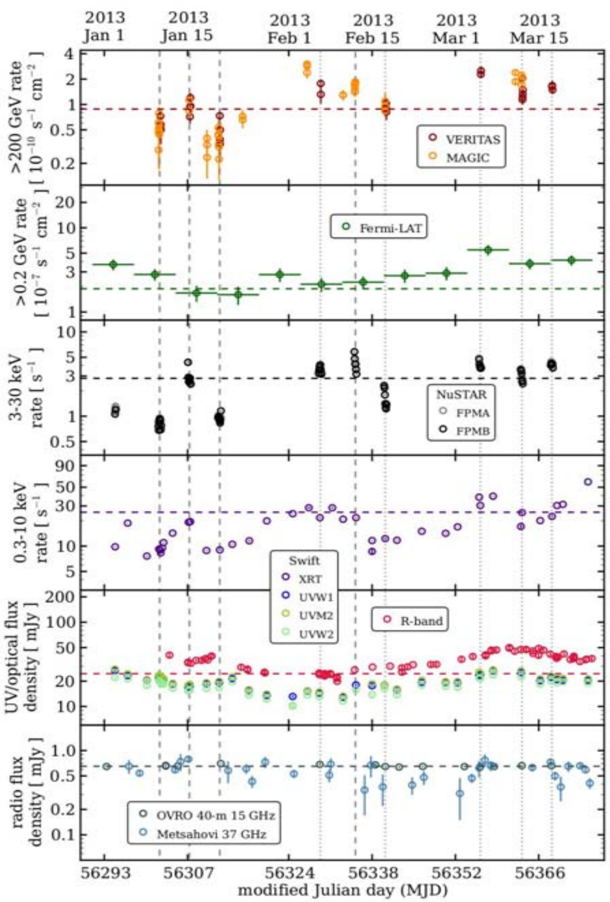


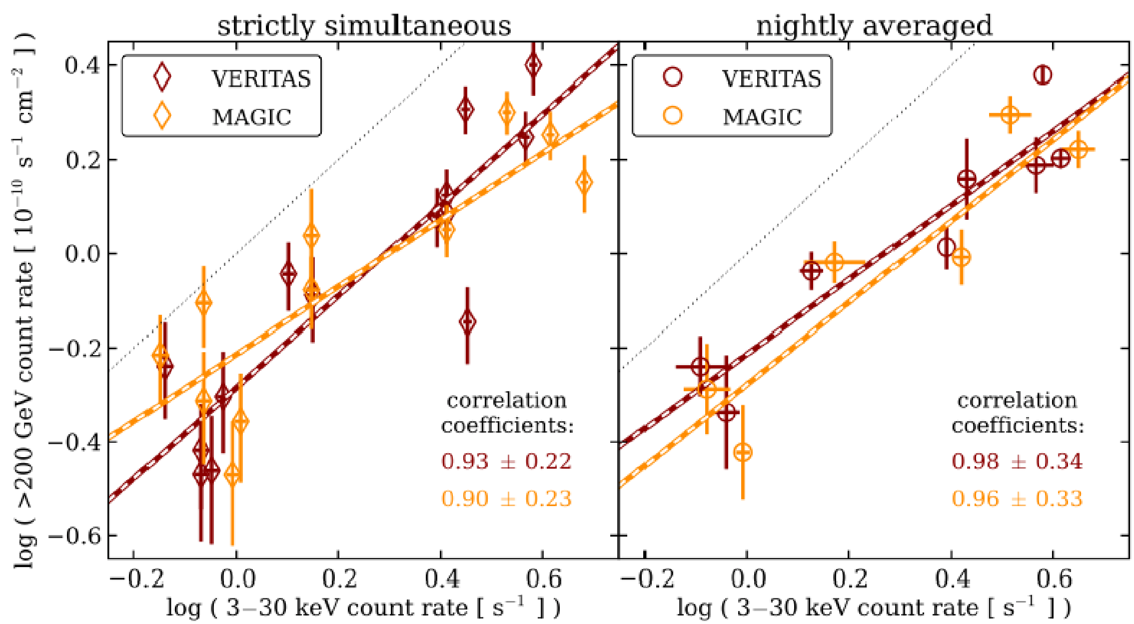
- *NuSTAR*
- MAGIC/VERITAS/HESS
- *Swift, Fermi*
- Optical instruments,
- OVRO/Metzahovi/f-Gamma



Example of Swift and NuSTAR data for Mkn 501 (z = 0.03):
Able to reconstruct full synchrotron peak with *Swift* XRT simultaneous observations: X-ray spectrum not a simple power law

Mkn 421 (NuSTAR Fermi, optical, VHE γ-rays)





- First results published in Balokovic et al. 2015; source beautifully detected up to 60 keV even in a faint state
- Source is softer when fainter, spectrum generally consistent with a broken power law (Γ_{high} ~ 3), not exponential cutoff – radiating particle distribution does not cut off below $E_{electron}$ ~ a few TeV
- Inferences regarding the particle distribution crucial towards studies of EBL
- The X-ray – TeV analysis suggests indicates “linear” correlation – the scattering is likely in the Klein-Nishina regime

PKS 2155-304 and particle content of the jet

* "One object at a time" approach and study a representative case rather than samples

* Well-known and extensively studied blazar, z = 0.117, one of the first BL Lac – type objects detected in X-rays

* Can be very variable: probably most "notorious" aspect of it is the large amplitude, minute-scale variability seen by H.E.S.S. (Benbow et al., Aharonian+ 2007)

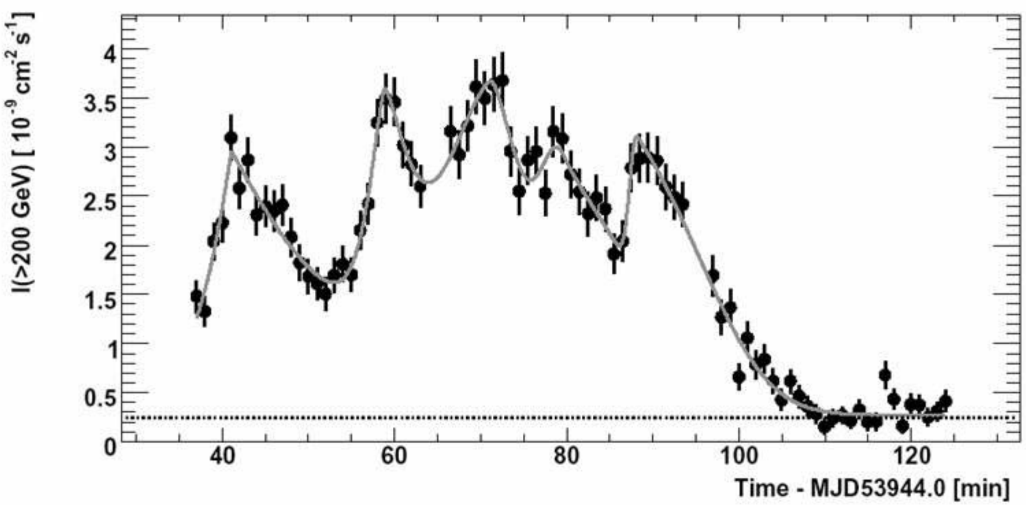


FIG. 1.— The integral flux above 200 GeV observed from PKS 2155−304 on MJD 53944 versus time. The data are binned in 1-minute intervals. The horizontal line represents I(>200 GeV) observed (Aharonian et al. 2006) from the Crab Nebula. The curve is the fit to these data of the superposition of five bursts (see text) and a constant flux.

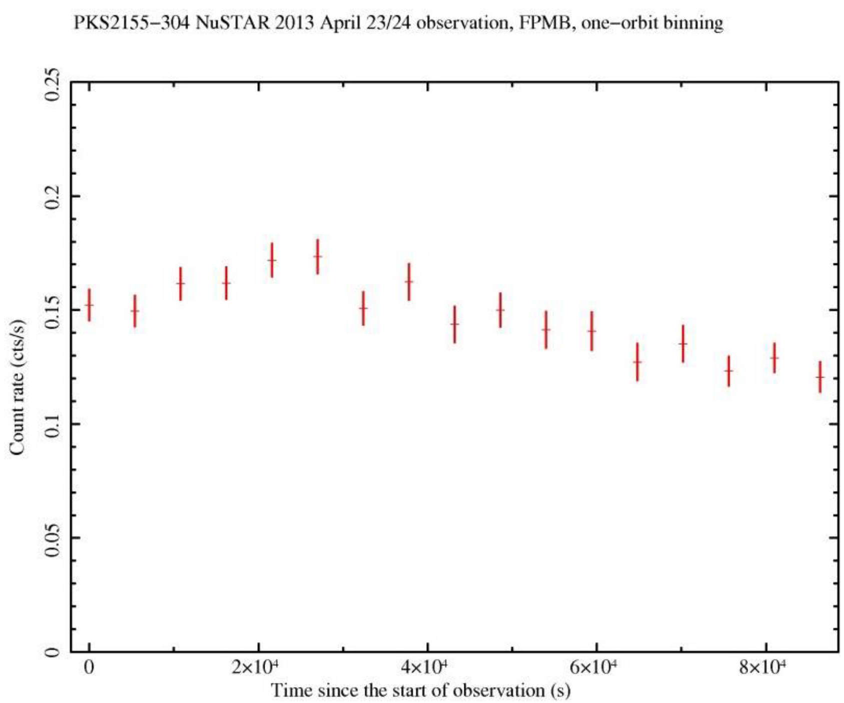
NuSTAR / Fermi observations

*** First NuSTAR pointing was done in April 2013 as a part of the cross-calibration with other X-ray missions -> lots of simultaneous X-ray data!**

*** Object was found in an exceptionally low X-ray state, ~ 10^{-11} erg cm^{-2} s^{-1} (2 – 10 keV), 1/3 of the previously reported "low" state**

*** NuSTAR collected a day's worth of data (~ 40 ks), shows very little variability,**

Fermi analysis straightforward: γ-ray flux during April 2013 is also low, no measurable variability; used +/- 5 days of Fermi data centered on the NuSTAR pointing



NuSTAR and XMM together

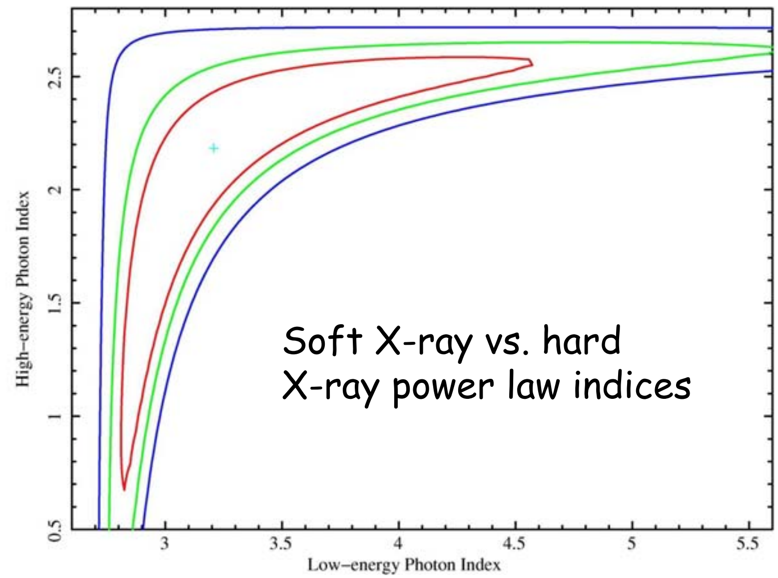
* NuSTAR data alone reveal a “hard tail” - power law indices are 3 & 2

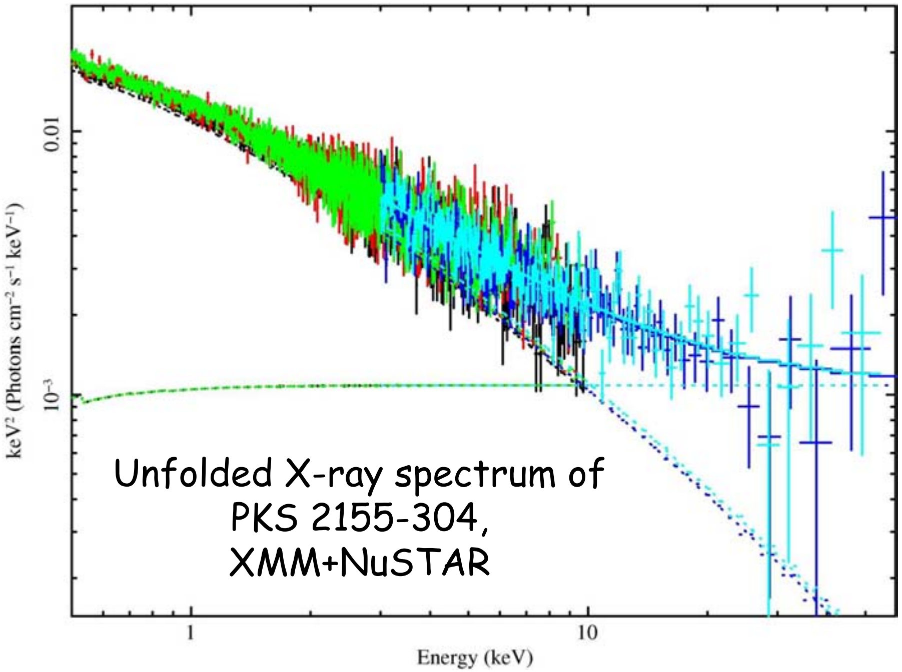
* Adding strictly simultaneous XMM data shows a more complete X-ray picture

- Imposing Galactic column, XMM data alone require a log-parabola model, gradual *steepening* of the spectrum in the 0.5 – 10 keV range

- Joint fit of NuSTAR and XMM implies a log-parabola for the lower E part of the spectrum, + a second, harder power law for the higher E part of the spectrum

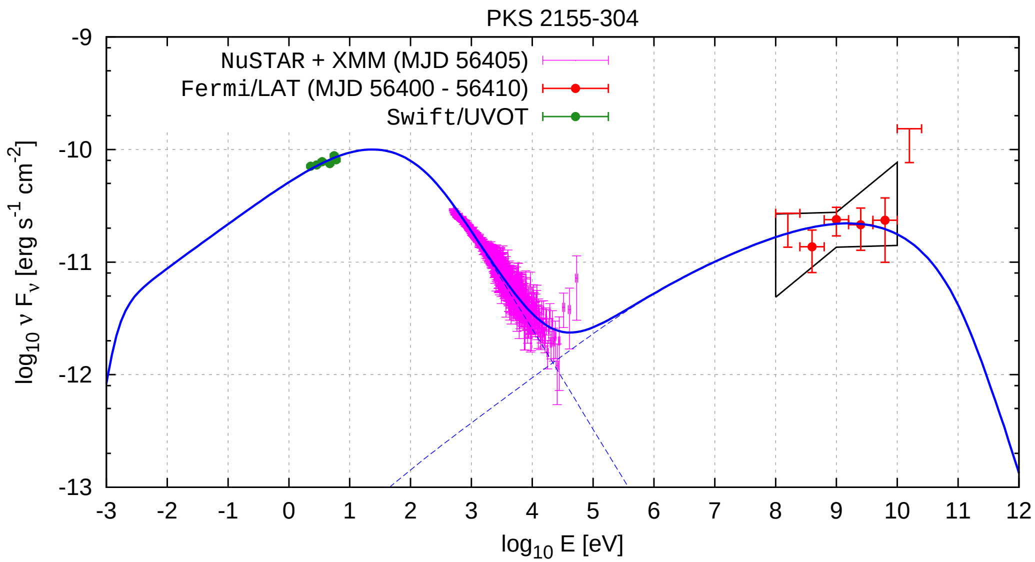
- The 20 – 40 keV flux is ~ 0.8 x 10^{-11} erg/cm^2/s





What does it all mean?
(don't forget about the charge neutrality)

- When we put X-rays together with the Fermi/LAT data, we have a very broad-band picture
- We fit the data with standard synchrotron self-Compton model (Rafal Moderski's "blazar" code, verified via Boettcher / Chiang model)
- The presence of the "hard tail" indicates that the inverse Compton spectrum has to extend to v. low energies!
- But that's where the radiating particles are <u>most numerous</u>!
- Can't be studied in the radio-band synchrotron component (synch. self-absorption), previously unconstrained
- Parameters of the model were calculated using the "standard" $\Gamma_j = 15$, B = 1 G, R = 3 x 10^{15} cm, consistent with all previous modelling
- <u>Important consequence: X-rays definitely should be polarized! (synchrotron process)</u>



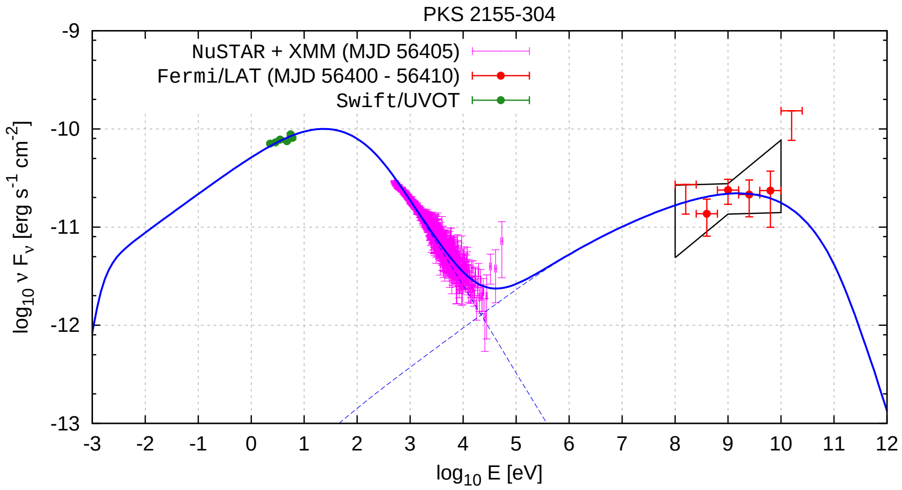
<u>Modelling results:</u>

We find L_B = 6 x 10^{42} erg/s, L_e = 3 x 10^{44} erg/s, L_{rad} = 10^{43} erg/s,

Even without protons, the jet is matter-dominated

Need charge neutrality: assuming one proton per electron yields L_p = 10^{47} erg/s. That's a lot!

Jet in PKS 2155-304 is likely pair-dominated!



Modelling results (repeated):

We find L_B = 6 x 10^{42} erg/s, L_e = 3 x 10^{44} erg/s, L_{rad} = 10^{43} erg/s,

Even without protons, the jet is matter-dominated

Need charge neutrality: assuming one proton per electron yields L_p = 10^{47} erg/s

CONCLUSIONS FOR THIS ANALYSIS:

L_p = 10^{47} erg/s is huge, totally unrealistic, even with a BH mass of 10^9 M_o, the source would need to accrete at L/L_{edd} ~1

This would imply a radiatively efficient accretion which is not the case for PKS 2155 or any HBLs - no thermal disk emission, no emission lines, low-efficiency accretion

HBL – type blazars are supposed to be advection – dominated accretion sources, not accreting at L/L_{edd} ~1!

More likely solution: the jet has a substantial pair content (#positrons / #protons ~ at least 50)

Possibly the first direct indication that HBL blazar jets are pair dominated

Conclusion seems robust to changes in Γ_j, B, … (hard to make a x50 error)

New gamma-ray observations including H.E.S.S.

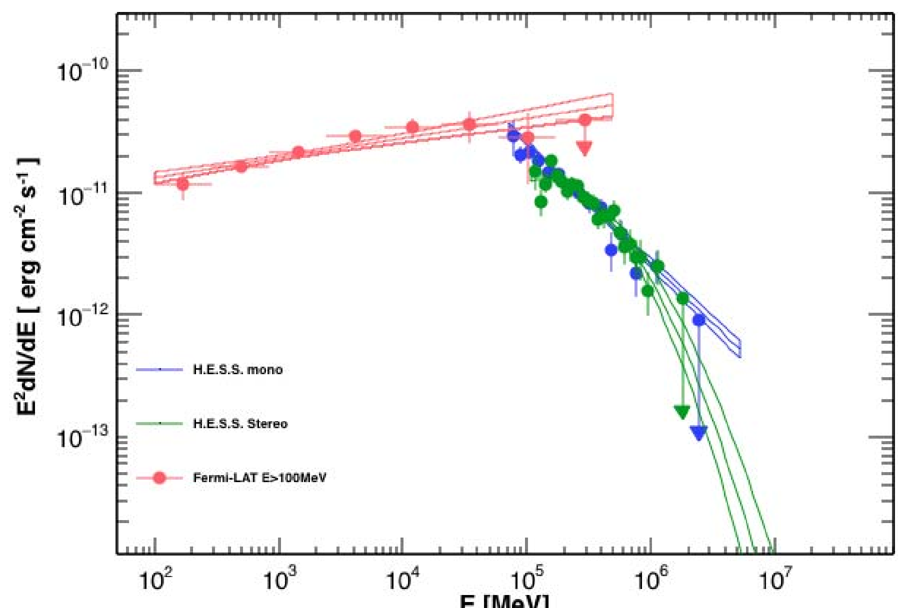
- In 2012 – 2013 H.E.S.S facility was expanded to include the 5th telescope (diameter 28 m!), extending the sensitivity

- PKS 2155 monitoring: June-August 2013

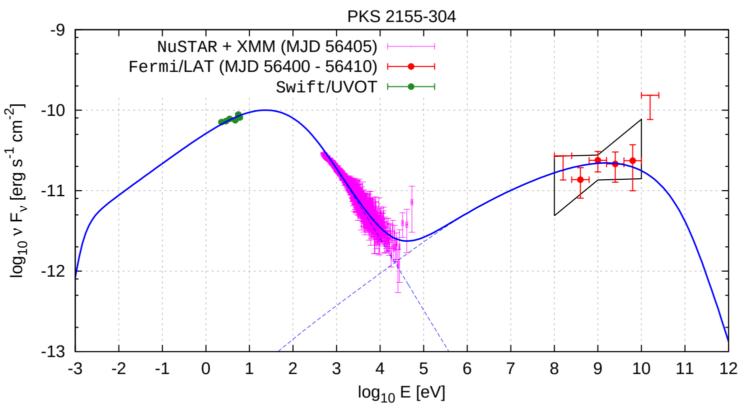
- Fermi shows v. little variability (red points, averaged over 3 months), H.E.S.S. flux a bit more variable

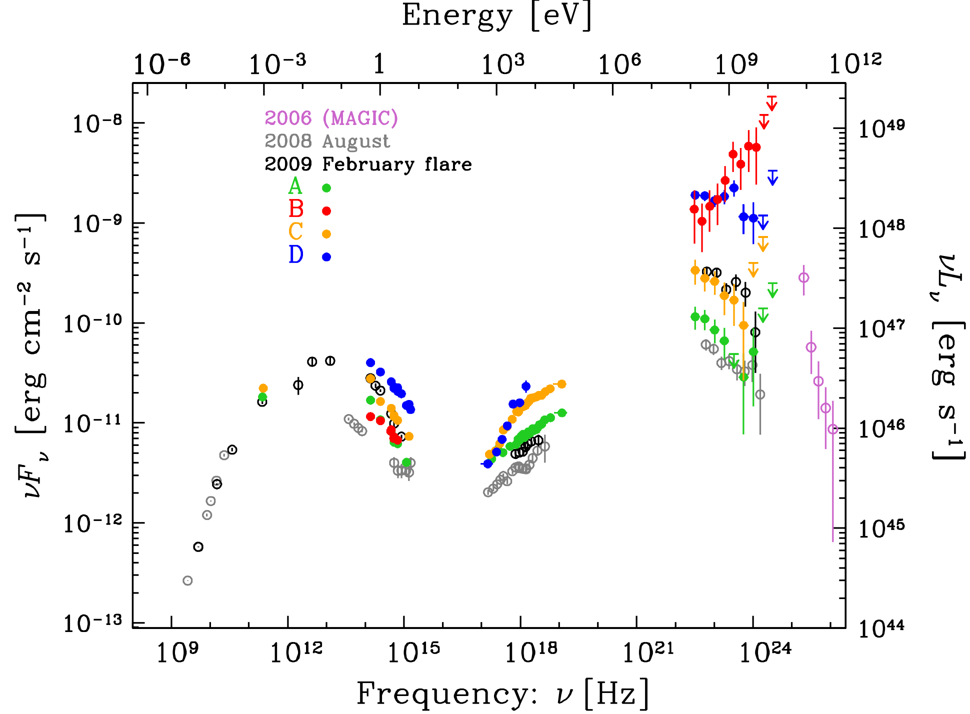
- Object was in a relatively low VHE and X-ray flux state throughout the 6 pointings

- "Hard tail" seems absent in any of the 6 the NuSTAR pointings; the limit on the 20 – 40 keV hard power-law component is ~ 4 times lower than the flux measured 2 months earlier



Other types of blazars?





High-luminosity FSRQ blazar 3C279

<u>OTHER CLASSES OF BLAZARS? HIGH LUMINOSITY TYPES?</u>

* High-luminosity blazars (Flat-Spectrum Radio Quasars) are different!

• Total power provided by accretion can be very large (signatures of luminous, high accretion-rate accretion disk)

• Their jets <u>cannot be entirely devoid of protons</u>:
If the jet was pure pairs, we'd see the "bulk-Compton" feature ("Sikora bump")
- from inverse Compton emission by the cold electrons in the jet (upscattering circum-nuclear AGN radiation) which has not been detected

* Protons (or alternatively, huge Poynting flux) is needed to provide the jet's kinetic energy

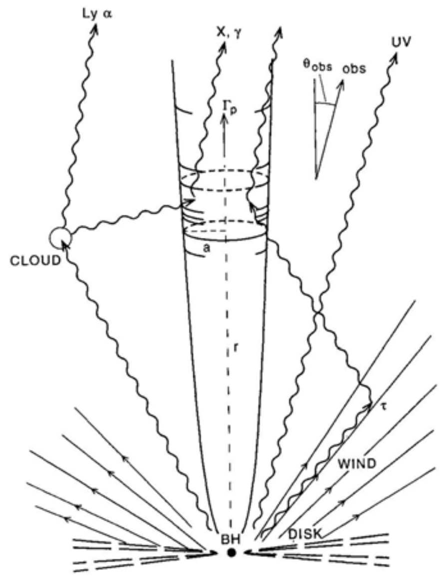


FIG. 2.—Geometry of the source. The radiating region, denoted by short cylinder of dimension a, moves along the jet with pattern Lorentz factor Γ_p. Underlying flow moves with Lorentz factor Γ, which may be different.

Schematic picture of geometry of a blazar jet

Pairs vs. protons in luminous blazars?

<u>Argument goes as follows:</u>

We can estimate the total kinetic power required to be carried by the jet to account for its luminosity for both "pure pair" and "no pair" cases

If we put all this jet kinetic energy into <u>pure pairs</u>, they will Compton-upscatter the Broad Line / accretion disk photon to X-ray energies – this is not seen! - *pure pair jet excluded*

At lease some protons are needed to carry the kinetic power

Some previous papers assumed no pairs -> requiring one proton per electron implied that the jet power was huge (Ghisellini+ 2014 Nature paper)

G+ 2014 invoked tapping the rotation of the black hole (via "Blandford-Znajek process") to power the jet

They argued that pure positron plasma would be slowed down by Compton rocket, produce the "Sikora bump" (spectral feature in the soft X-rays), which is not detected

But, G+ 2014 didn't consider the intermediate cases…

Another tool – "calorimetry" to rescue!

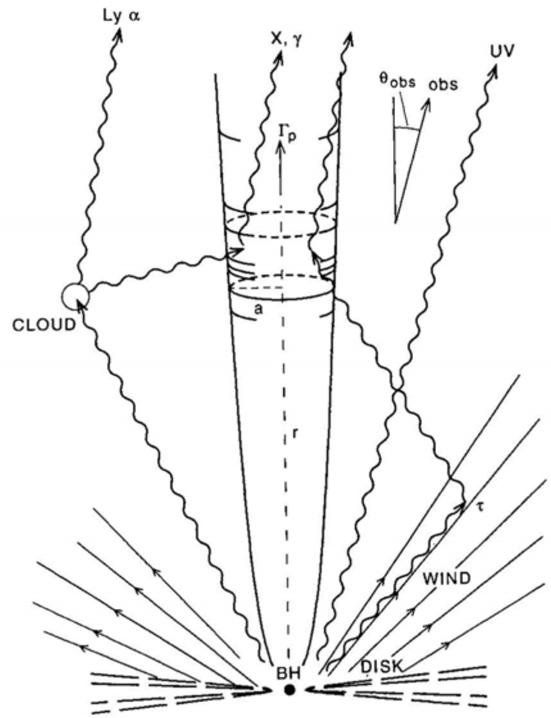
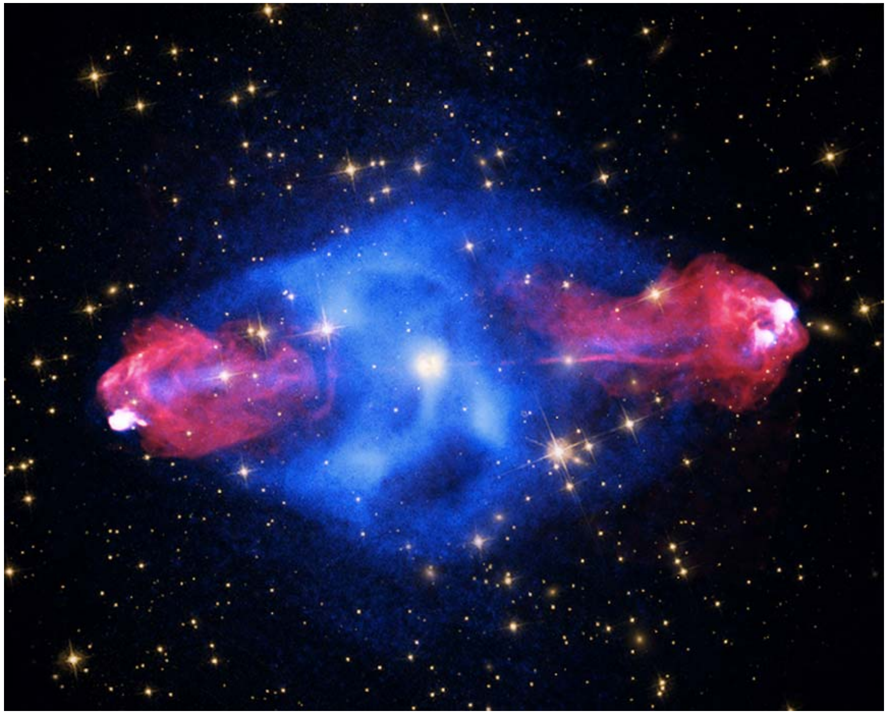


FIG. 2.—Geometry of the source. The radiating region, denoted by short cylinder of dimension a, moves along the jet with pattern Lorentz factor Γ_p. Underlying flow moves with Lorentz factor Γ, which may be different.



Cygnus A radio galaxy:
blazar jet viewed from the side

<u>Hot spots in radio galaxies are a form of a "calorimeter" (beam dump)</u>

One can estimate the total power of the jet from radio lobes
Those emit isotropically – no issues with Doppler boosting
Estimate of the jet power is now more robust – Pjanka, Zdziarski, Sikora (2016) infer that total jet power is lower than inferred by Ghisellini+

Conclusion: invoking Blanford-Znajek in high-luminosity sources is not required - but positron-to-proton ratio needs to be ~ 20

Conclusions

We discussed <u>observations and modelling</u>
of 2 types of jet-dominated AGN (blazars)

BL Lac-type blazar: PKS2155-304; “hard X-ray tail” seen in
NuSTAR - jet must contain appreciable pairs (e+/P > 20)

In high-luminosity, powerful blazars, there is opposite limit:
jet plasma cannot be pure e-/e+ (bulk-Compton
limits); BUT, if pure e/P plasma – jet power
still excessive…
- e+/P ~ 20 OK, consistent with
radio hot spots which provide additional constraints

CHALLENGE TO THEORETICAL EFFORTS: HOW ARE THE PAIRS
PRODUCED IN THE JET?

Future X-ray observations will include measurements of
X-ray polarization

X-ray / γ-ray polarization <u>predictions</u>: depends on the radiation
process

* synchrotron: strong polarization, probably same angle as
optical (X-rays in HBL-type blazars)
* inverse Compton: if seed photons unpolarized – probably
no polarization (γ-rays in FSRQs)
* inverse Compton: if seed photons are polarized – strong
polarization, same angle as synchrotron, same swings
(X-rays in FSRQs)