Parameter estimation of compact binaries from gravitational-wave signals

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University of Birmingham
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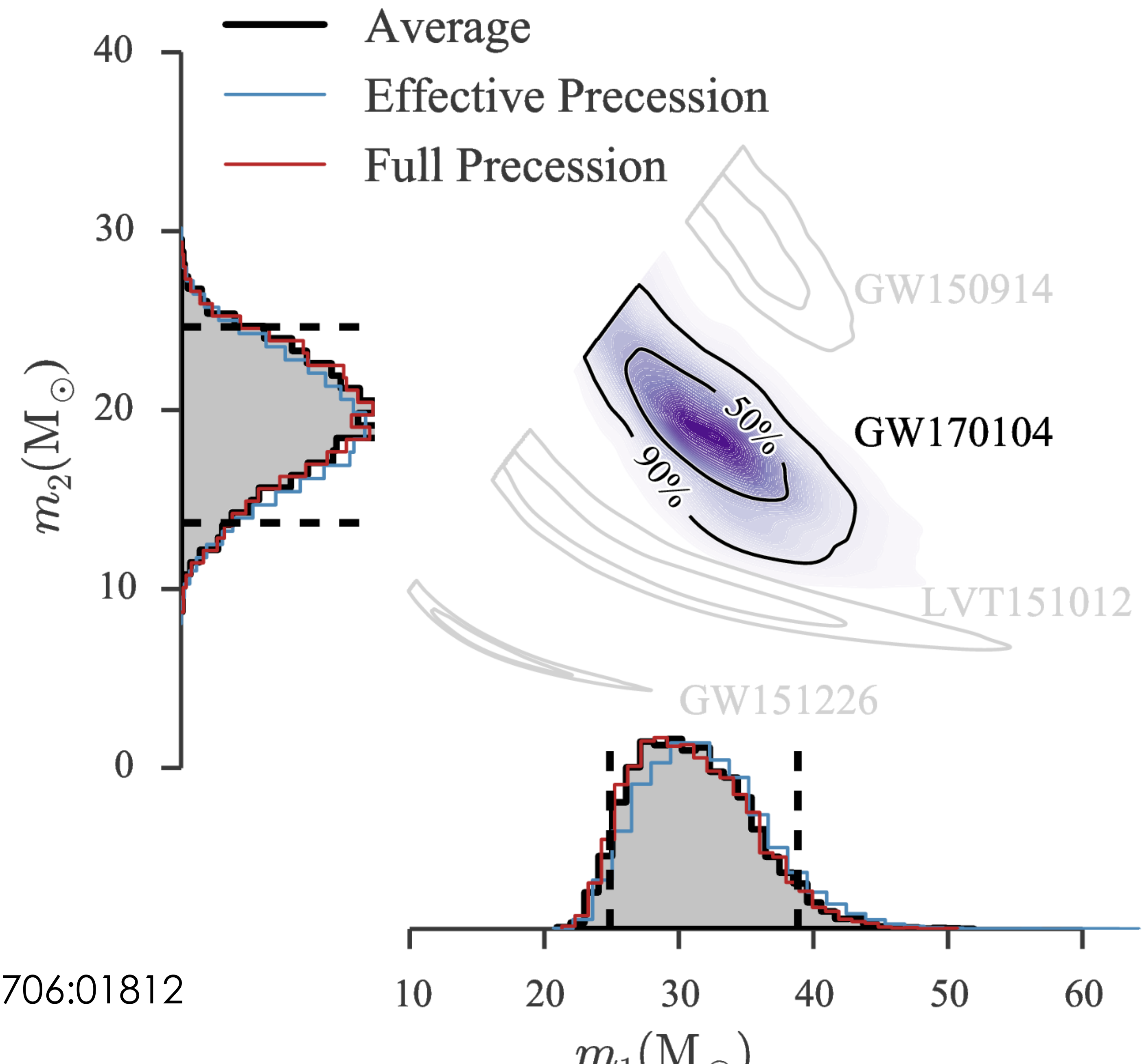
On behalf of the
LIGO Scientific & Virgo Collaborations

DCC G1701119

New Frontiers in Gravitational-wave Astrophysics



LVC
arXiv:1706:01812

How to infer source properties

Our measurements so far

Gravitational-wave observations and formation processes

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Our measurements so far

Gravitational-wave observations and formation processes

Bayes' theorem

$$p(\theta|d) = \frac{p(d|\theta)\,p(\theta)}{p(d)}$$

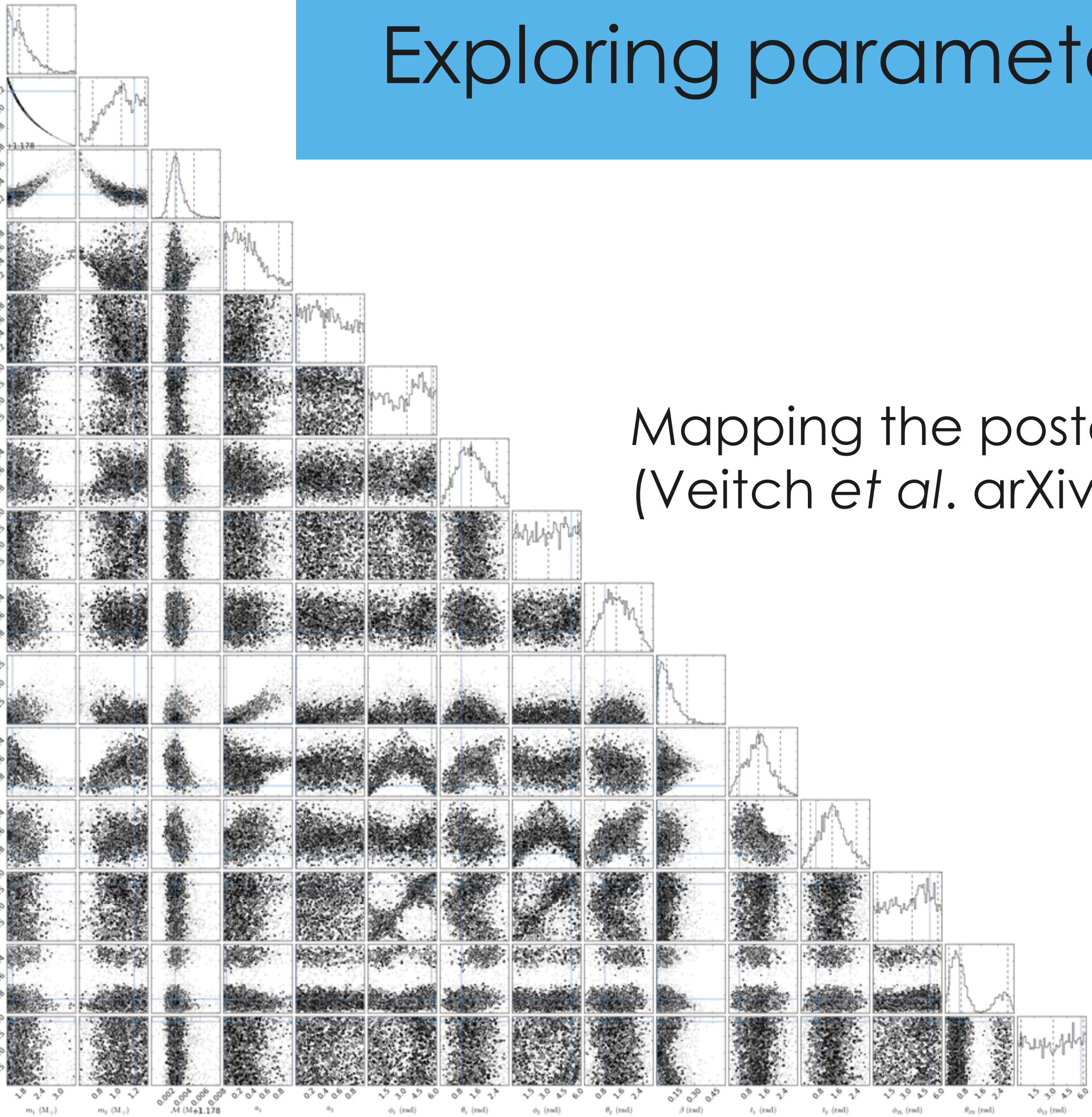
Bayes' theorem

Likelihood Prior

$$p(\theta|d) = \frac{p(d|\theta)\,p(\theta)}{p(d)}$$

Posterior Evidence

Exploring parameter space



Mapping the posterior is difficult (Veitch *et al*. arXiv:1409.7215)

Likelihood

$$p(d|\theta) \propto \exp\left[-\frac{1}{2}\sum_k \langle h_k(\theta) - d_k | h_k(\theta) - d_k \rangle\right]$$

LVC arXiv:1602.03840

Likelihood

Noise-weighting

$$p(d|\theta) \propto \exp\left[-\frac{1}{2}\sum_k \langle h_k(\theta) - d_k | h_k(\theta) - d_k \rangle\right]$$

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$$h_k(\theta) \to h_k(\theta)\left[1 + \delta A_k\right]\exp\left[i\delta\phi_k\right]$$

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Likelihood

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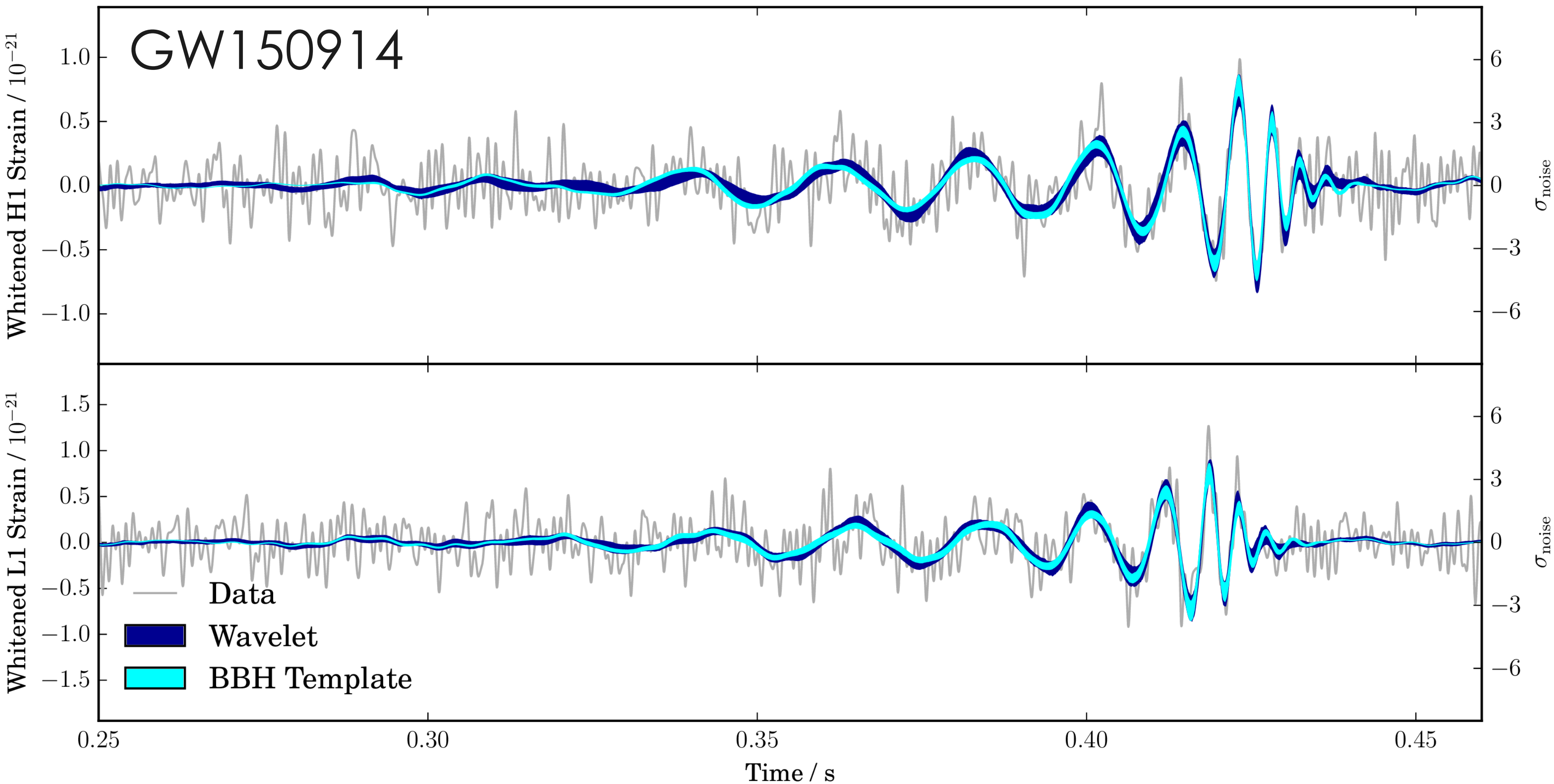
Waveform

LVC arXiv:1602.03840

Waveform



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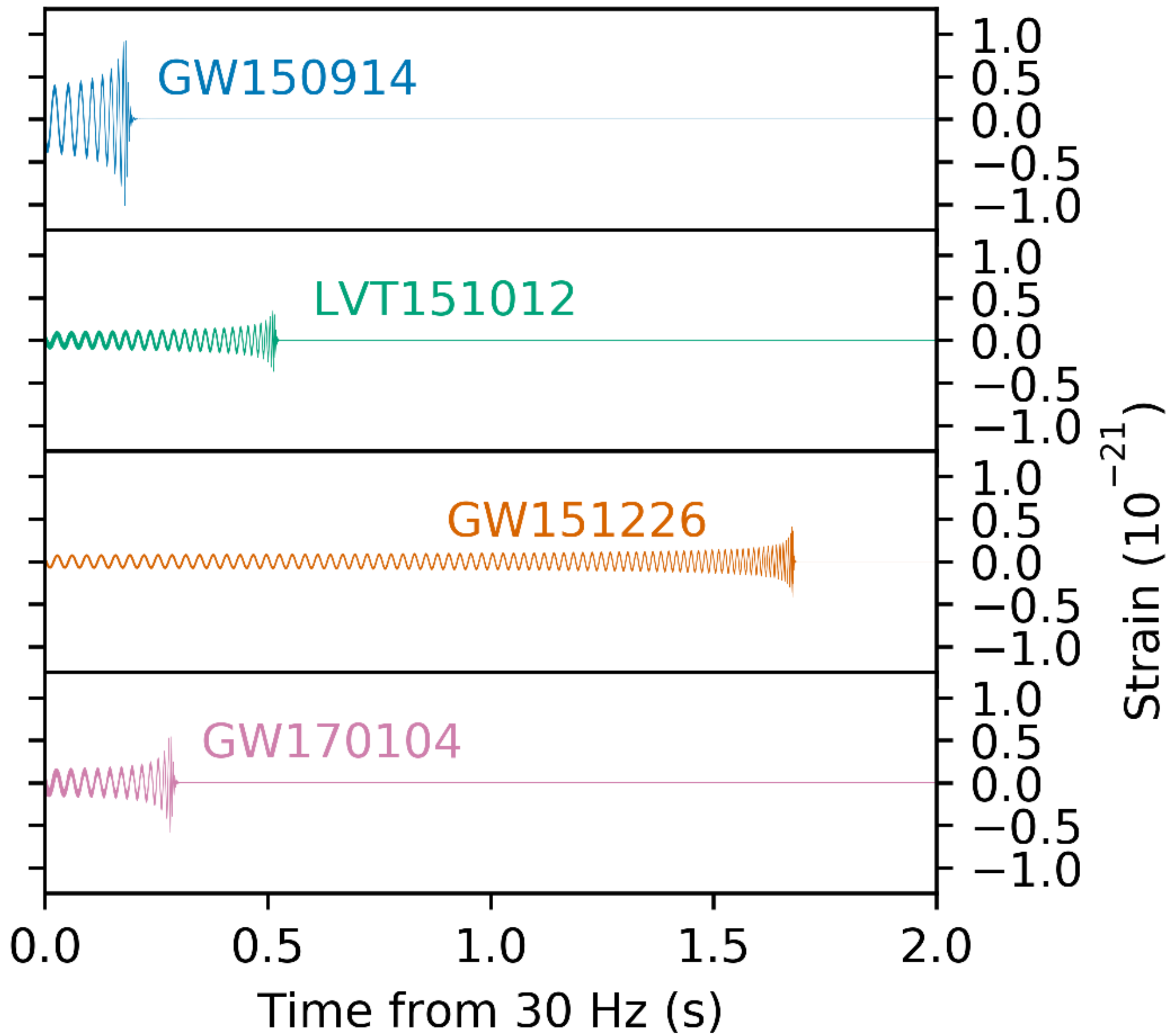
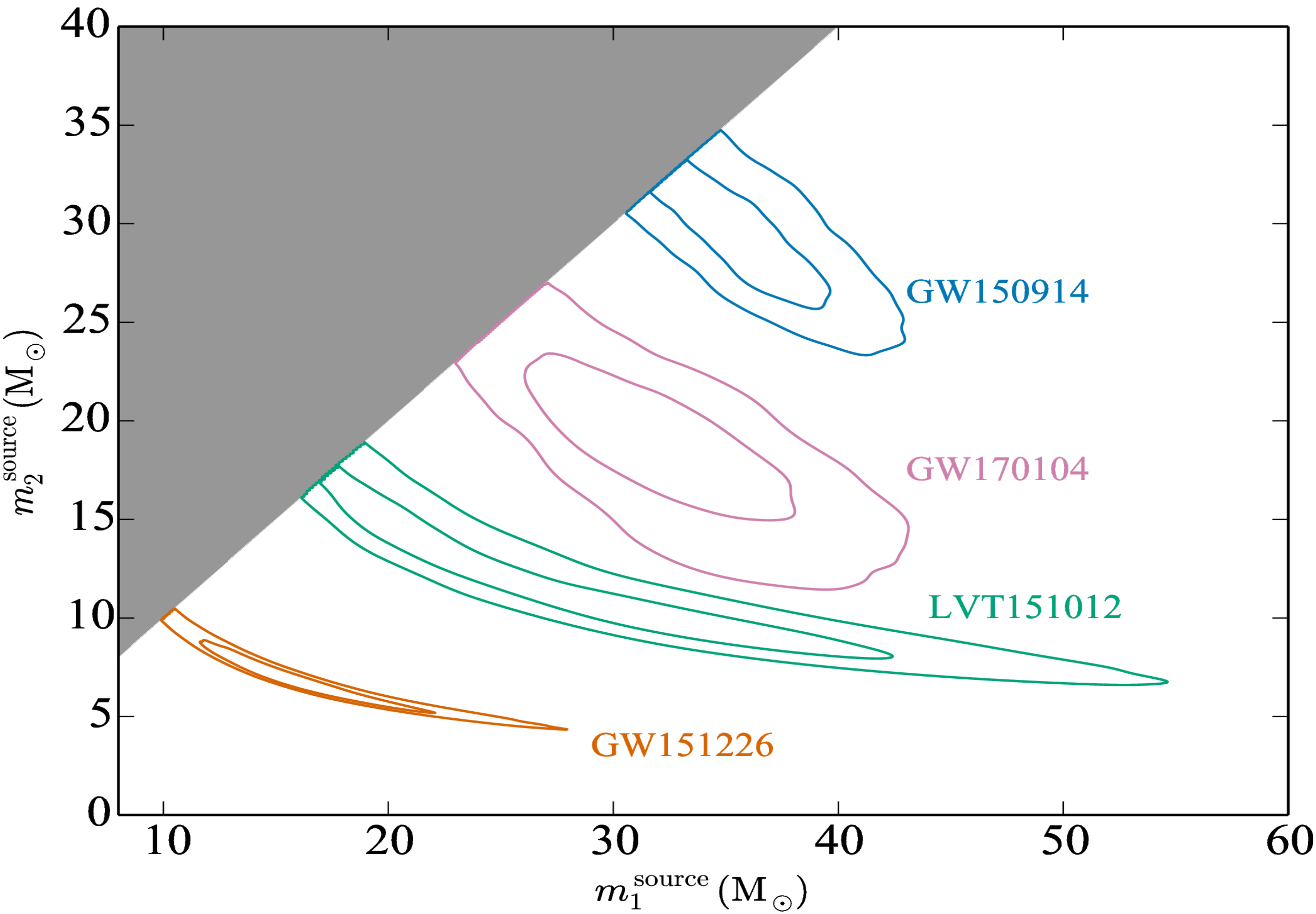
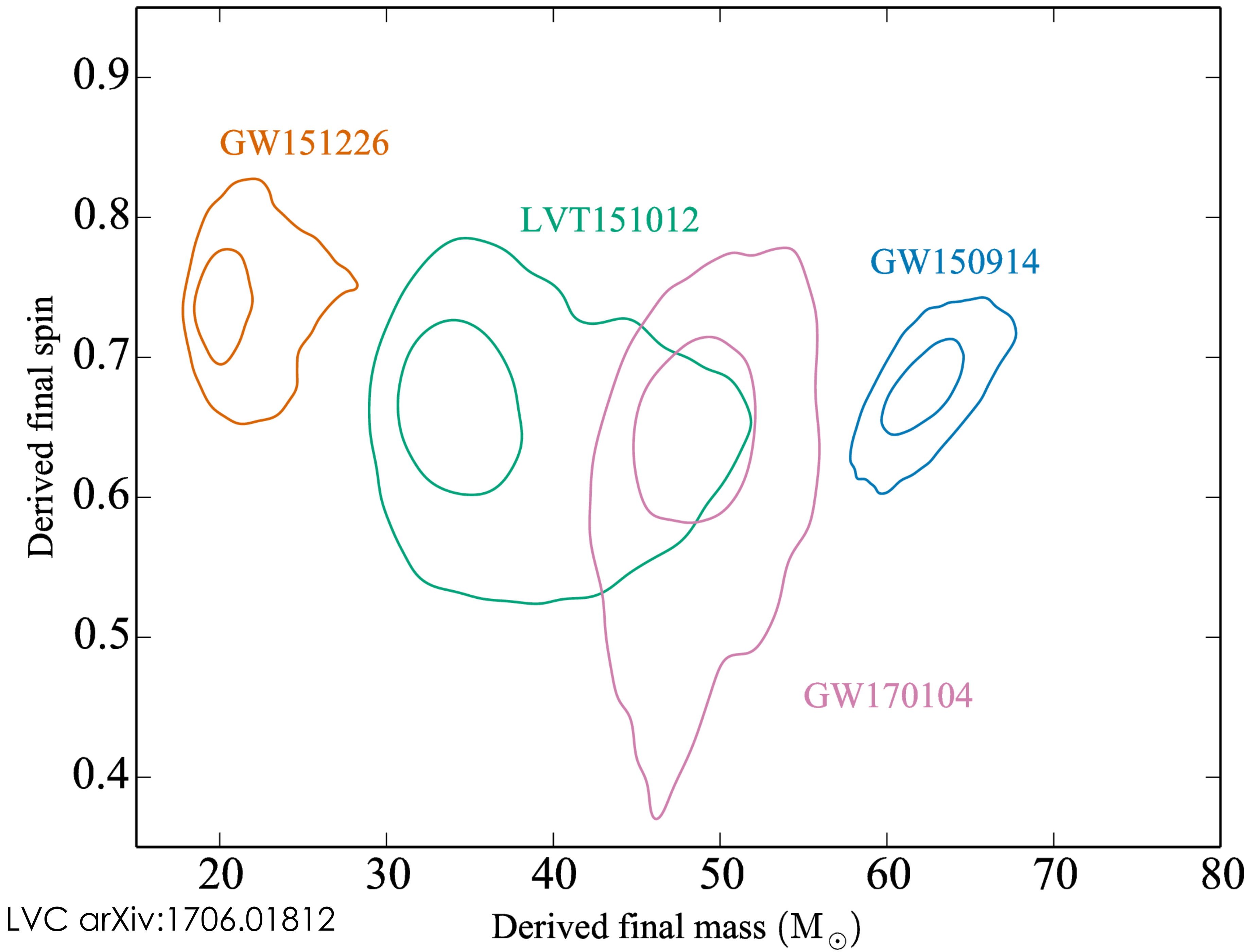
Binary black hole signals



LVC
arXiv:1606.04856,
arXiv:1706.01812



LVC arXiv:1706.01812



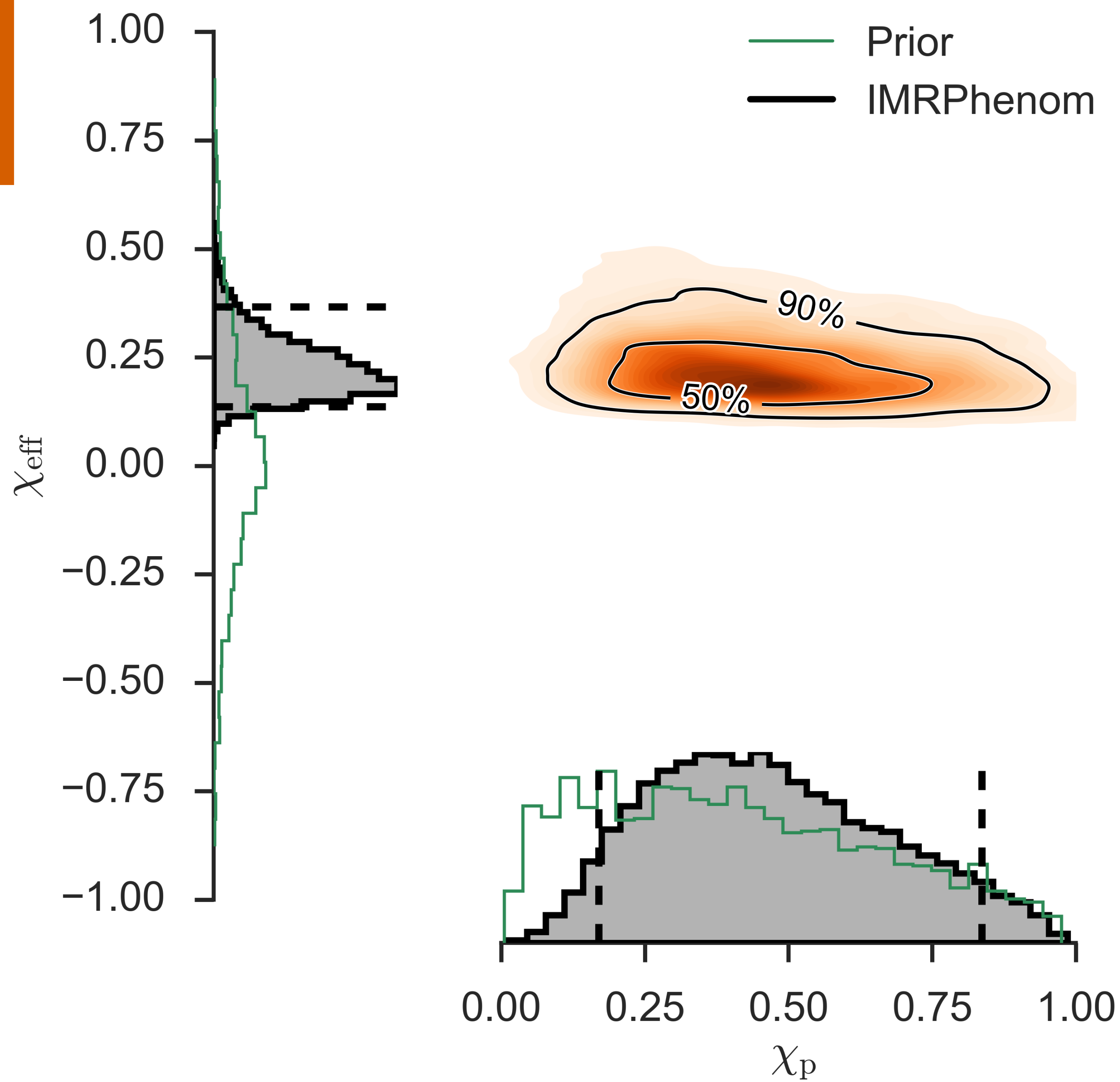
GW151226
LVT151012
GW150914
GW170104
0.9
0.8
0.7
0.6
0.5
0.4
Derived final spin
20
30
40
50
60
70
80
Derived final mass ($M_\odot$)
LVC arXiv:1706.01812

Effective inspiral spin

$$\chi_{\rm eff} = \frac{c}{GM}\left(\frac{\boldsymbol{S}_1}{m_1} + \frac{\boldsymbol{S}_2}{m_2}\right)\cdot\hat{\boldsymbol{L}}$$

Most important combination of spins for evolution of inspiral (arXiv:0909.2867, 1005.3306)

Spin



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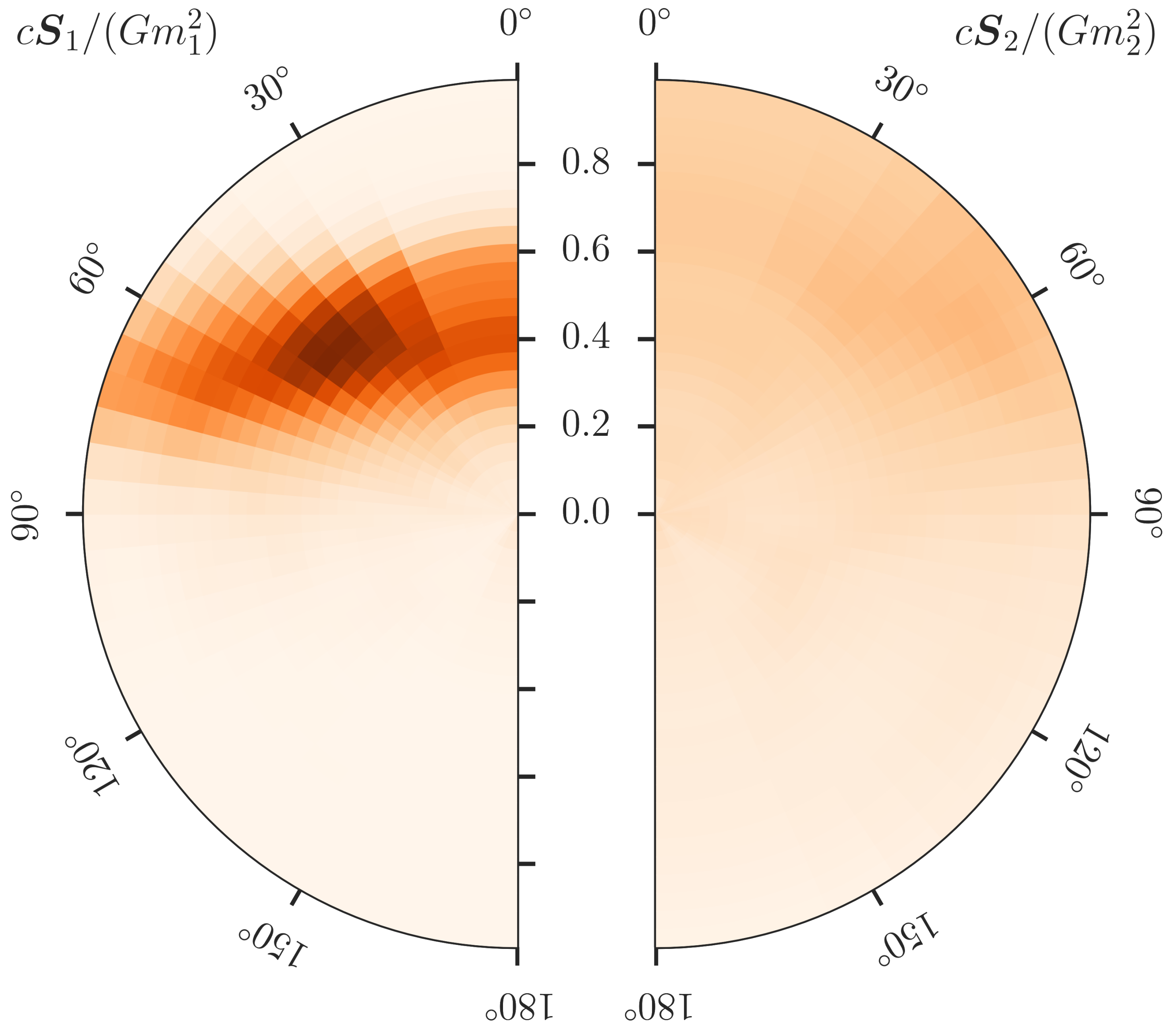
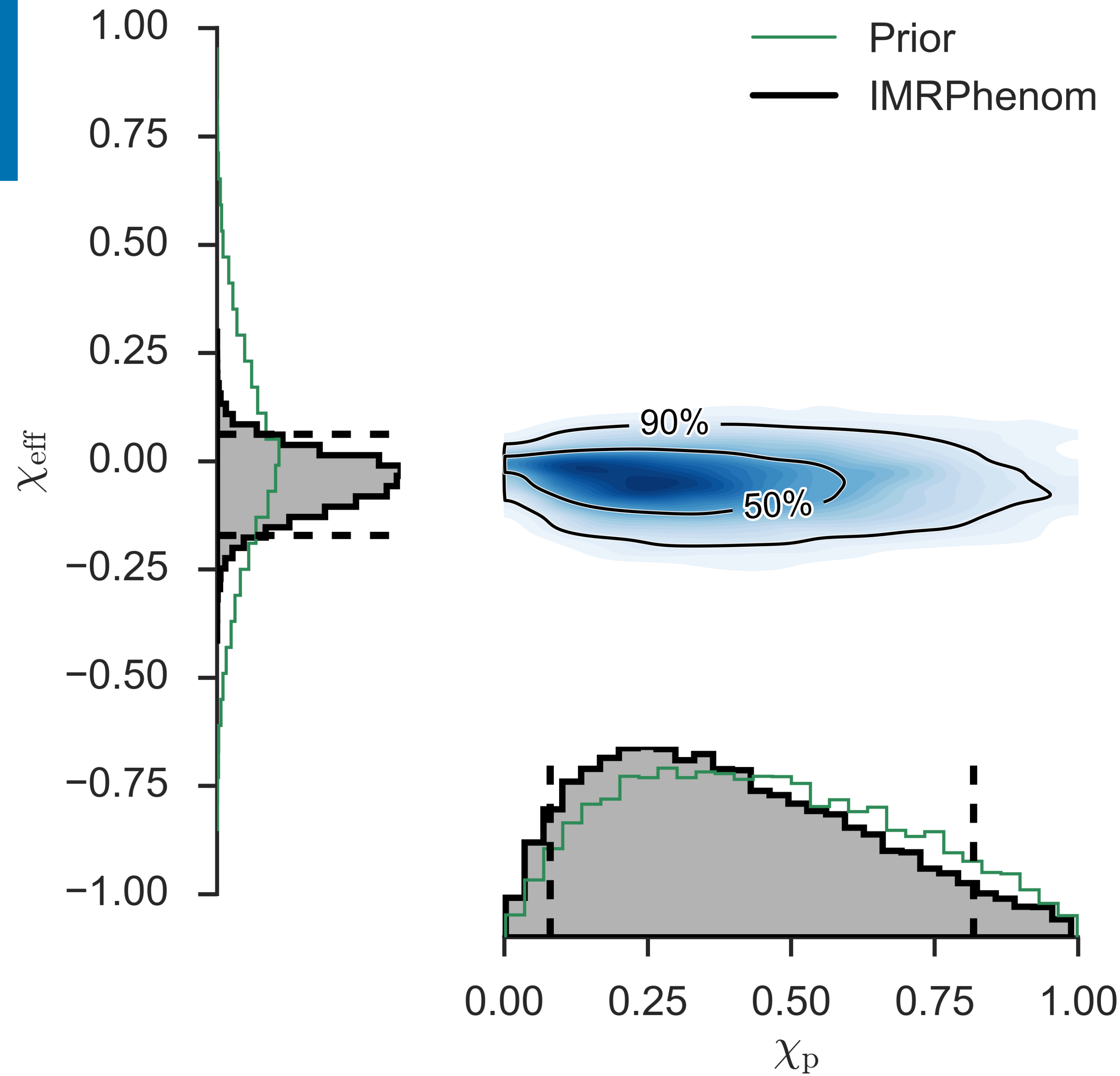
Spin



LVC
arXiv:1606.04855

Spin



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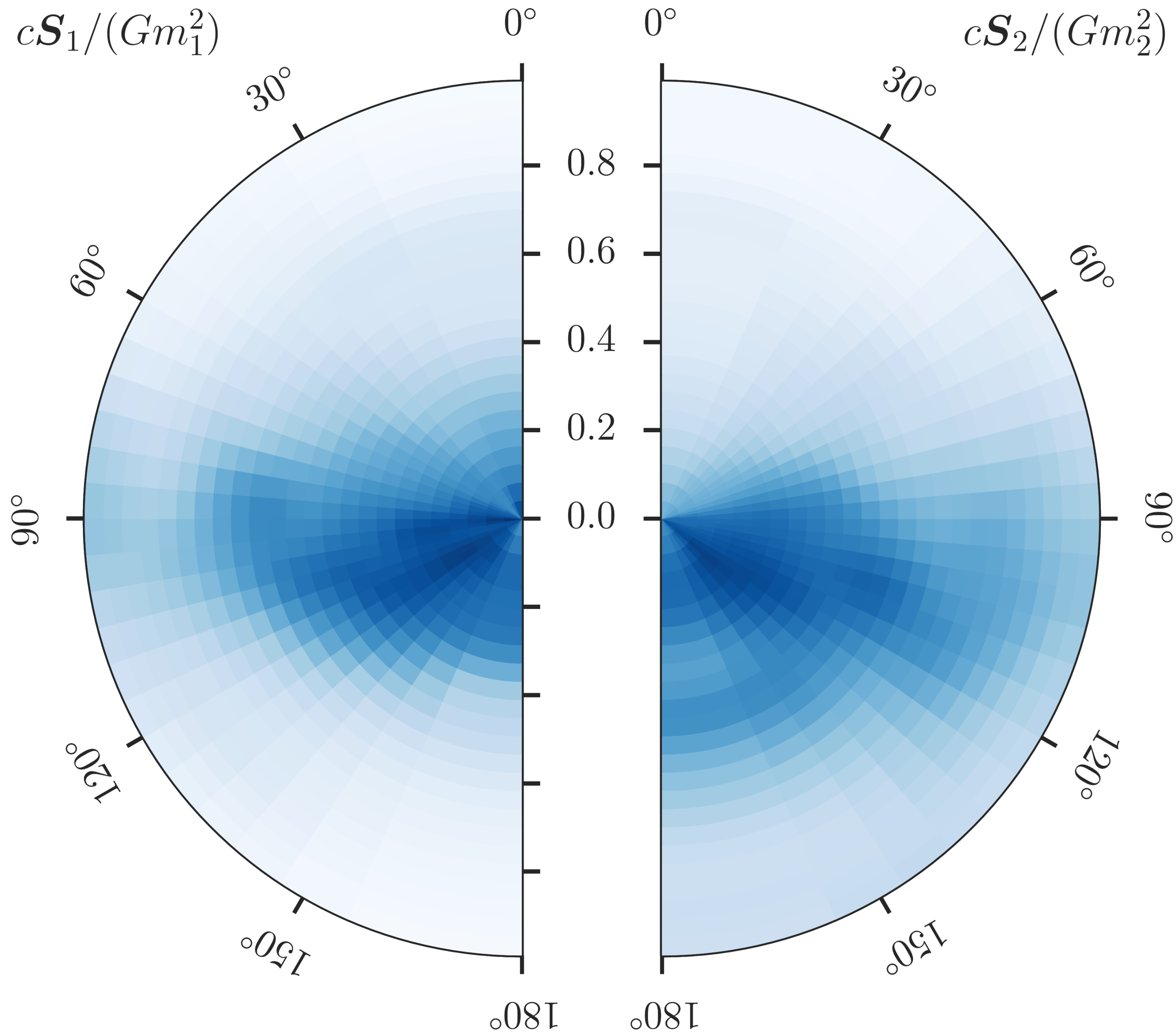
Spin



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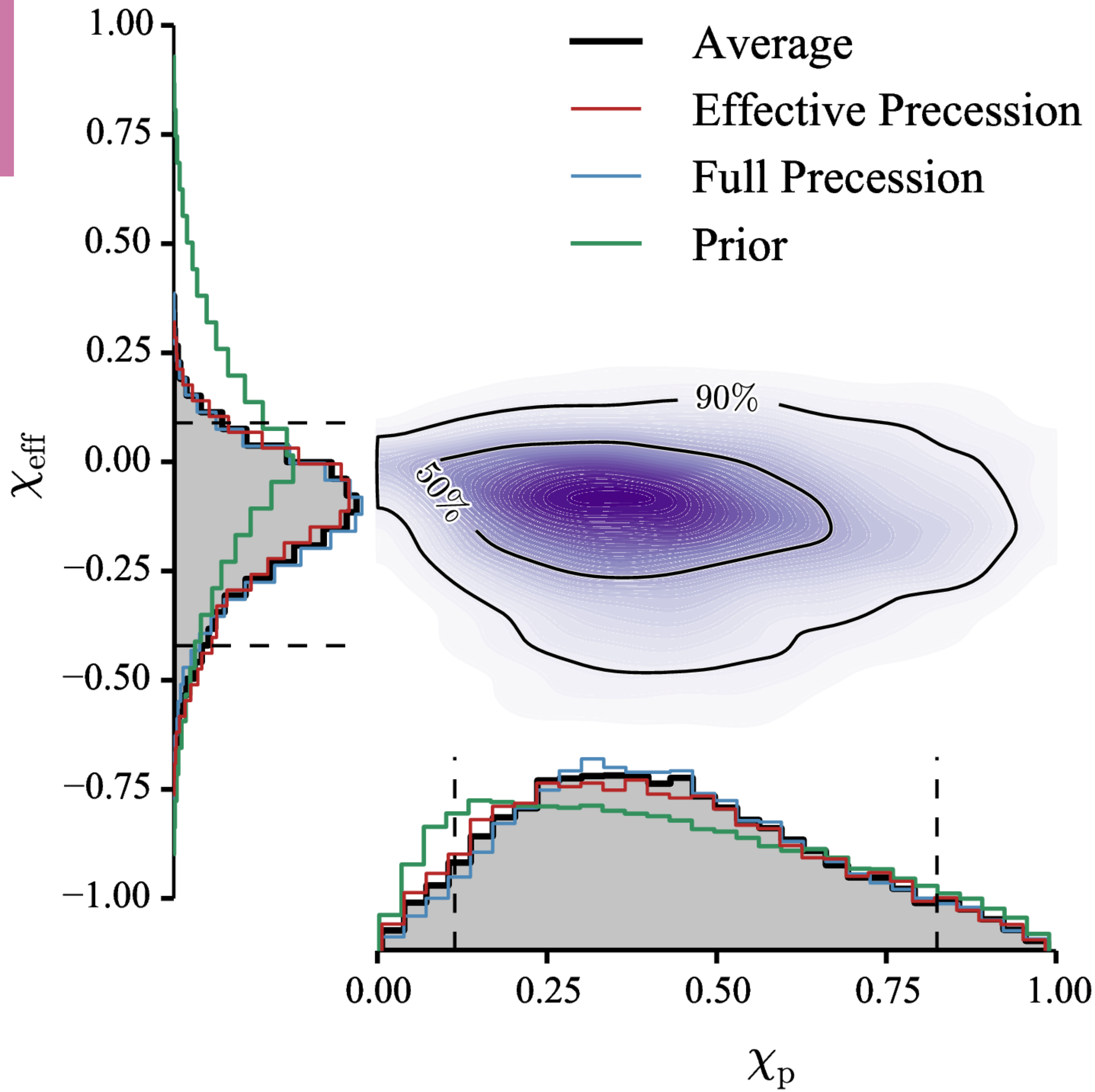

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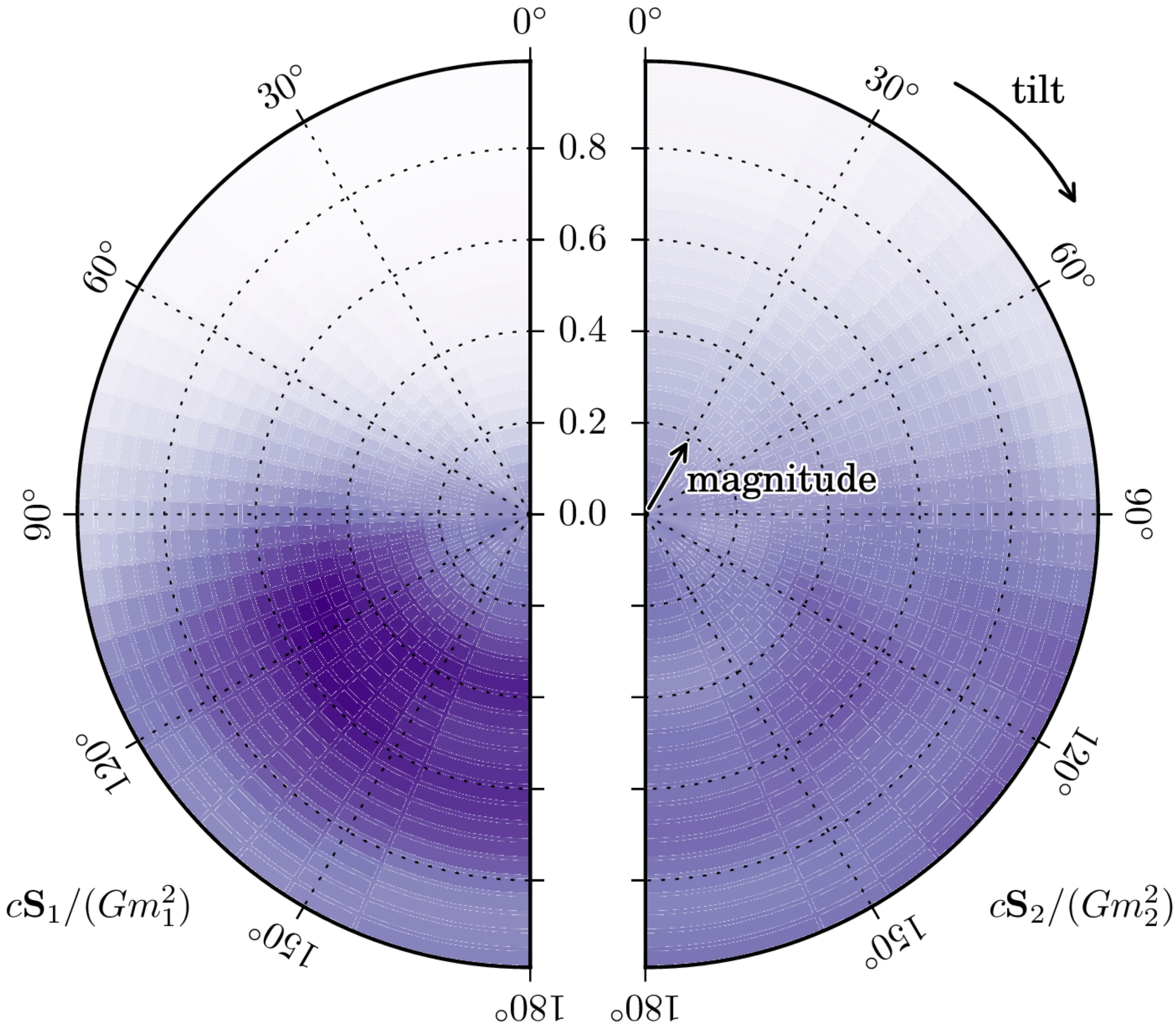

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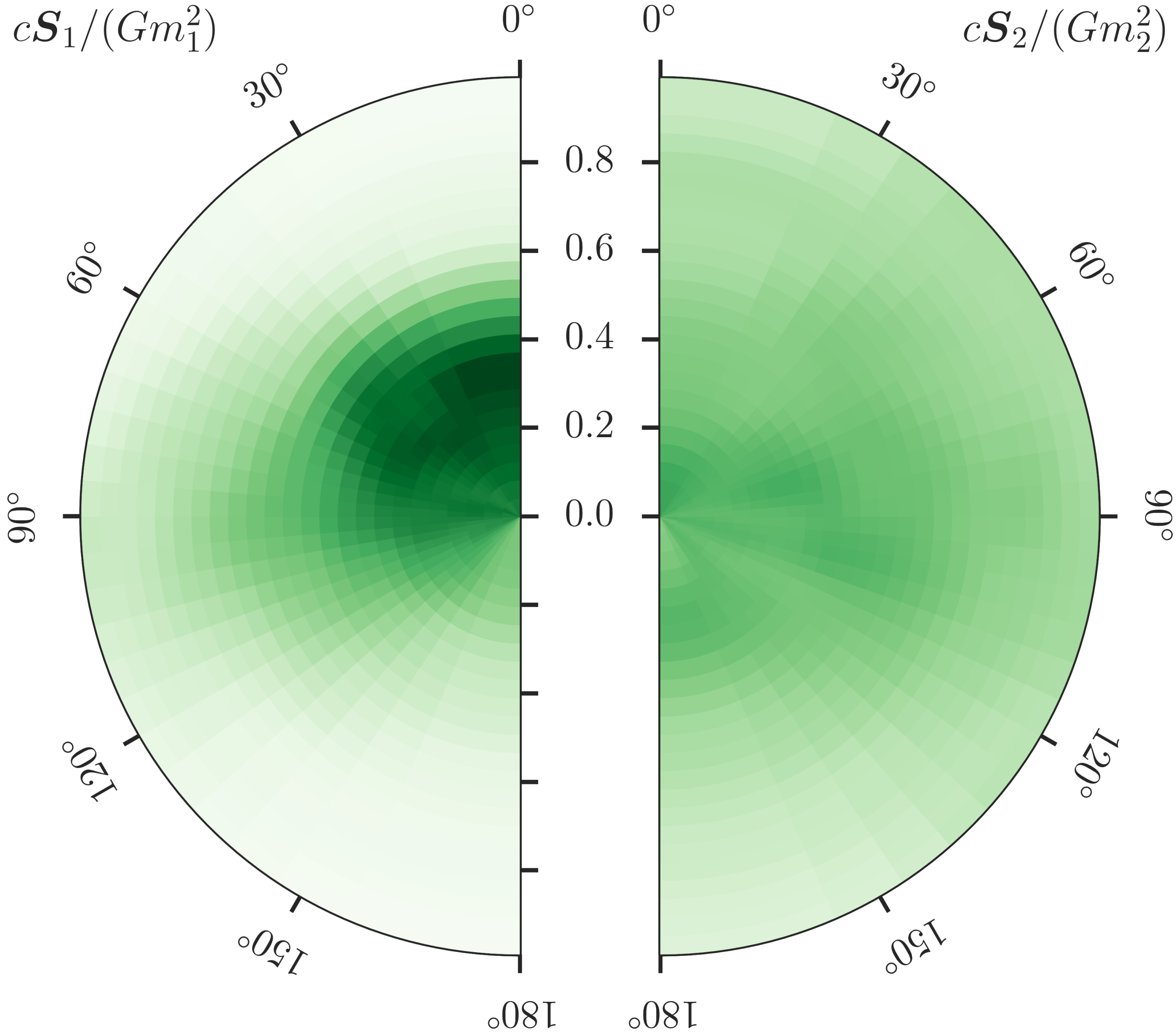

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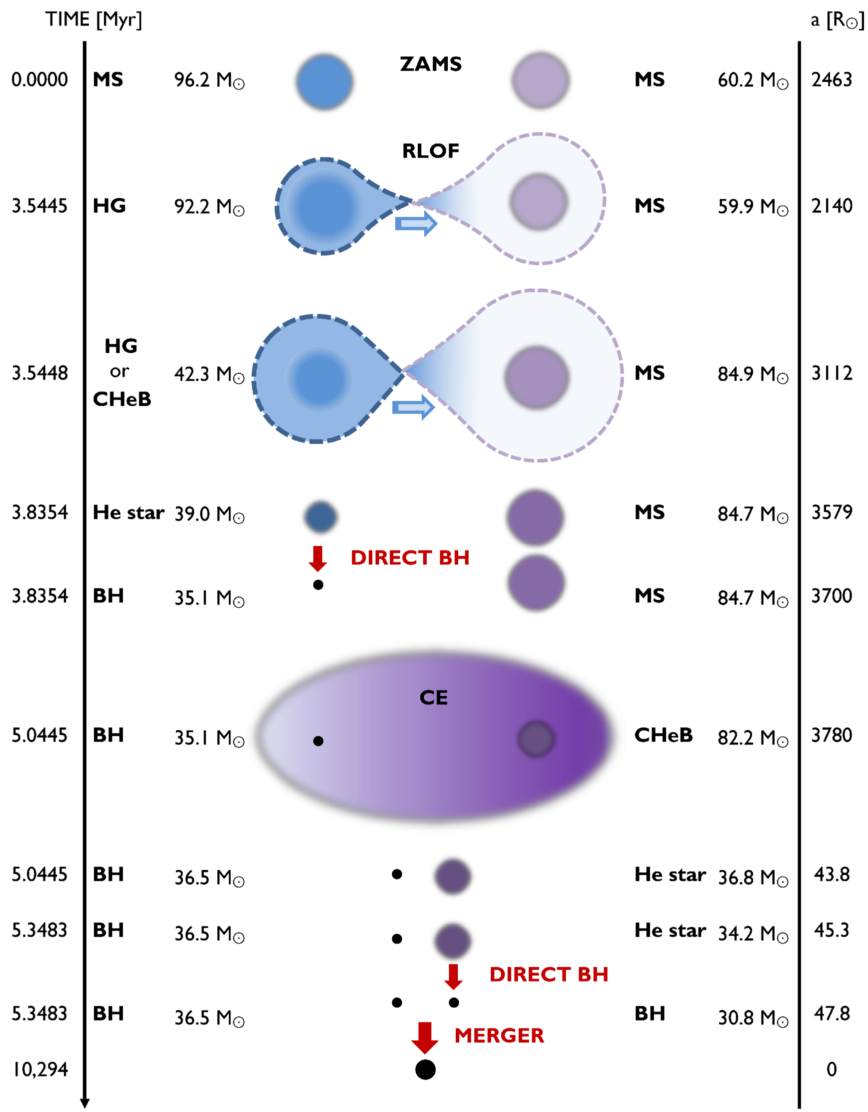
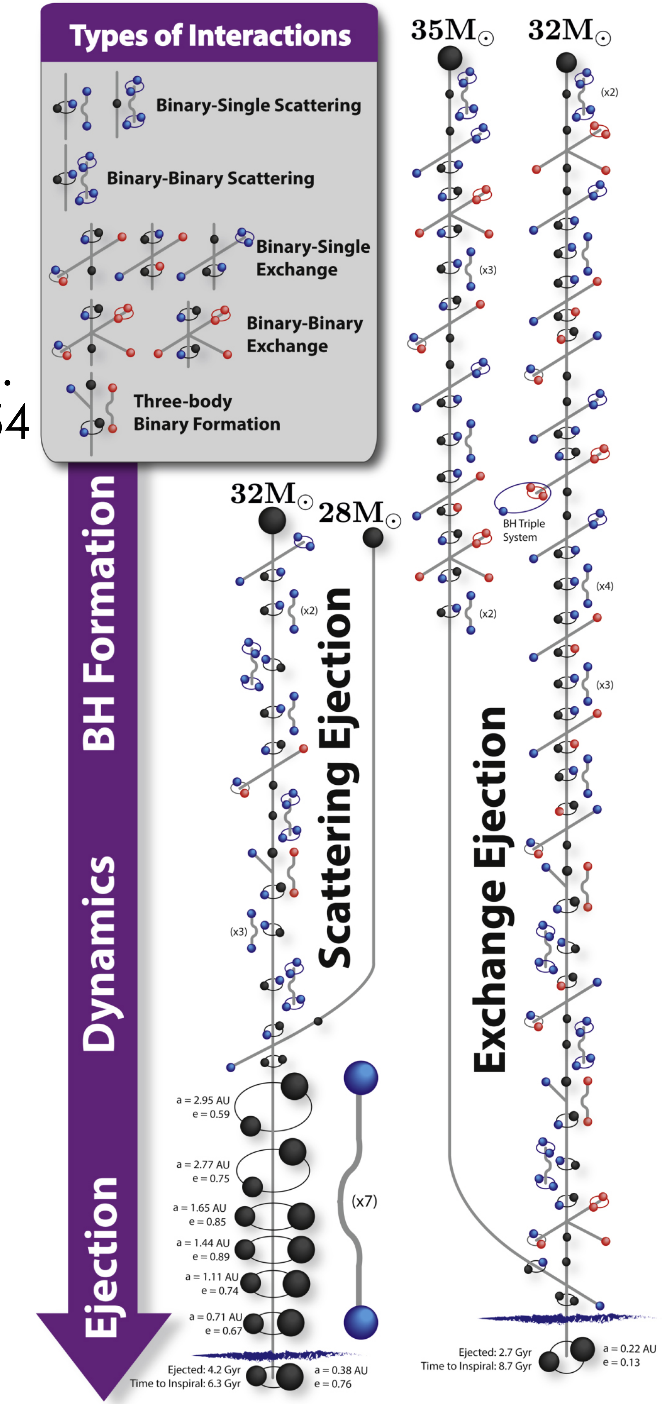
Binary formation

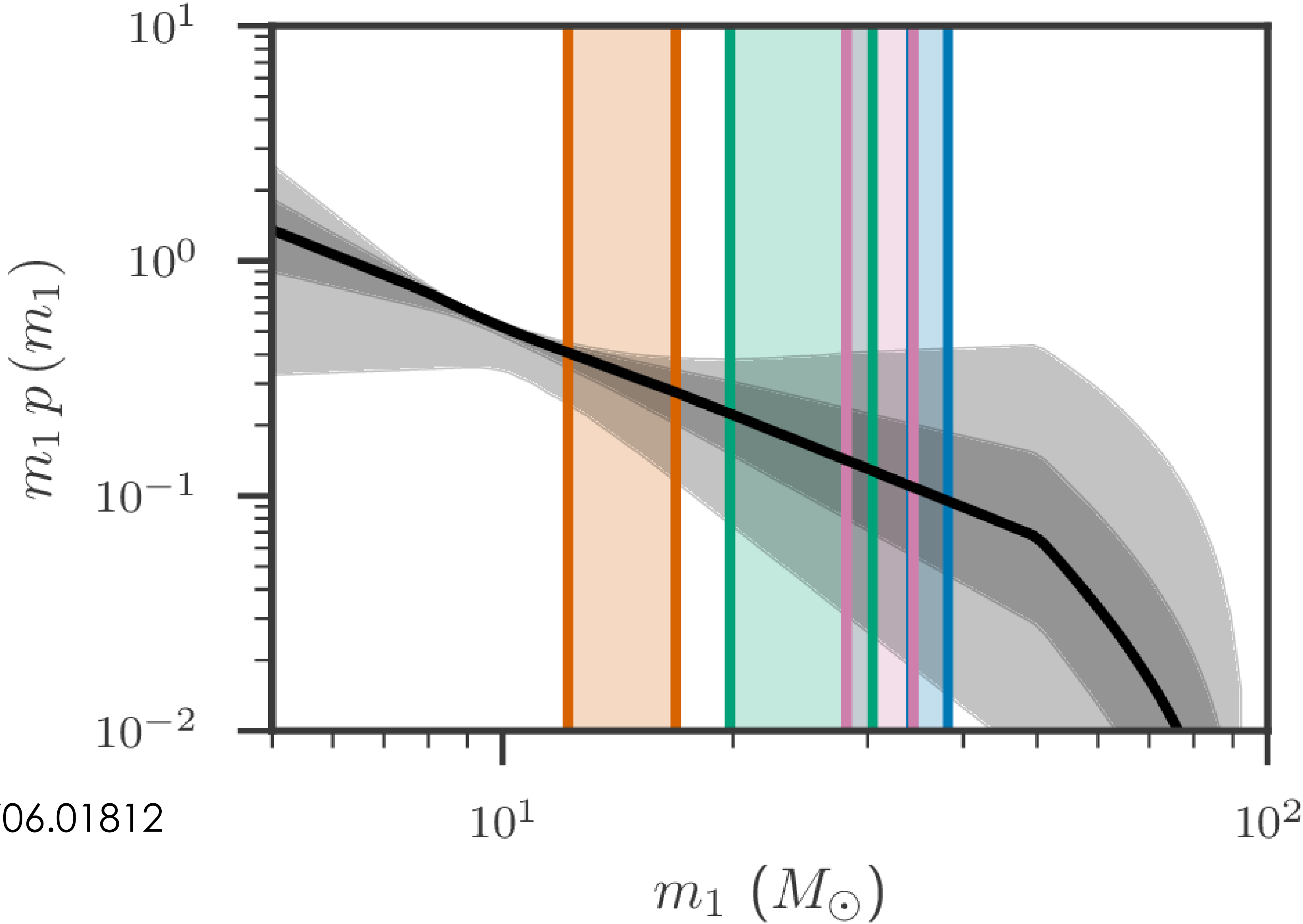


Rodriguez *et al.*
arXiv:1604.04254

Belczynski *et al.*
arXiv:1602.04531



Mass distribution



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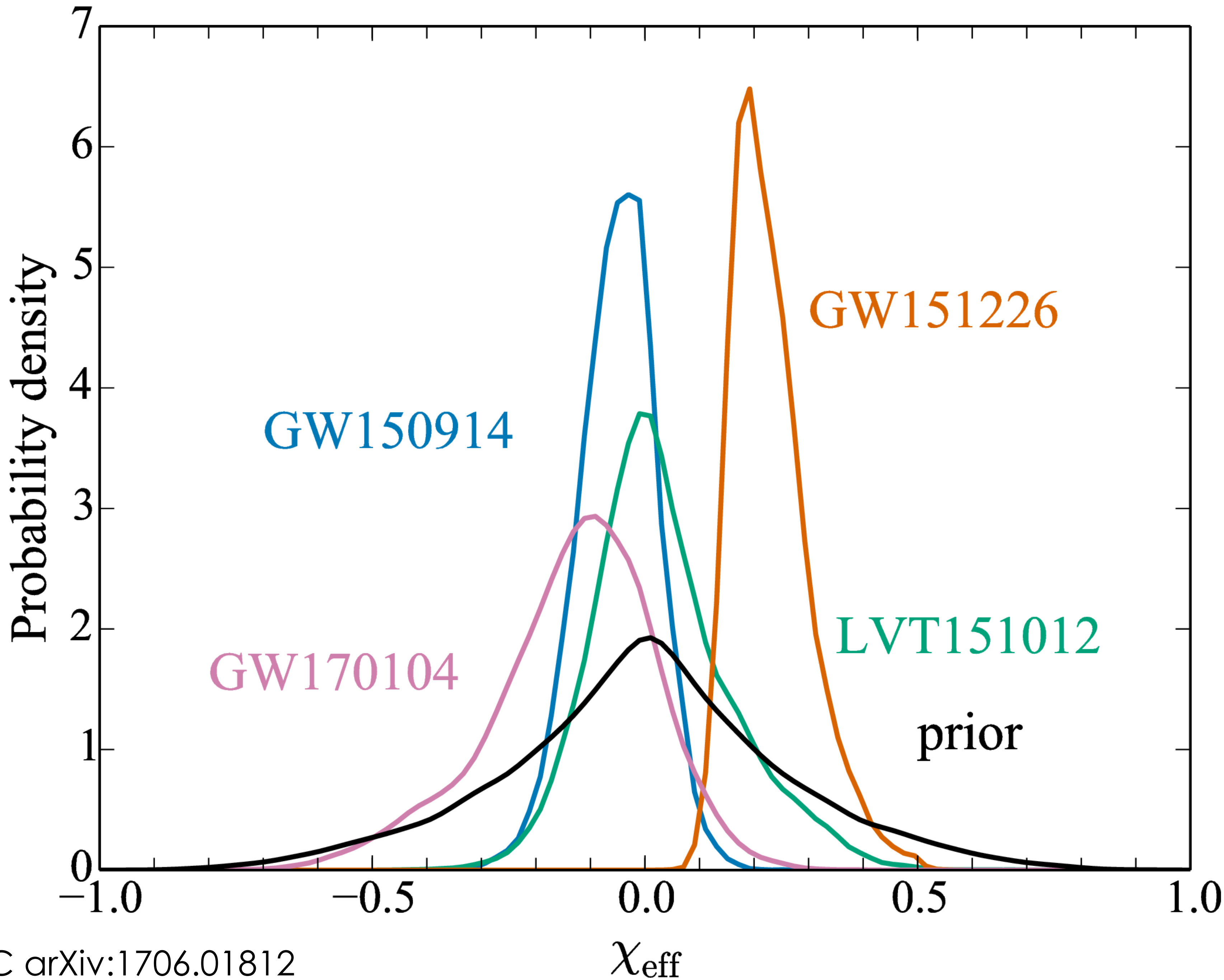

LVC arXiv:1706.01812

Hierarchical analysis

arXiv.org > astro-ph > arXiv:1706.01385

Astrophysics > High Energy Astrophysical Phe

Distinguishing Spin-Aligned and Isotropic Black Hole Populations With Gravitational Waves

Will M. Farr, Simon Stevenson, M. Coleman Miller, Ilya Mandel, Ben Farr, Alberto Vecchio

(Submitted on 5 Jun 2017 (v1), last revised 6 Jun 2017 (this version, v2))

The first direct detections of gravitational waves from merging binary black holes open a unique window into the binary black hole formation environment. One promising environmental signature is the angular distribution of the black hole spins; systems formed through dynamical interactions among already-compact objects are expected to have isotropic spin orientations whereas binaries formed from pairs of stars born together are more likely to have spins preferentially aligned with the binary orbital angular momentum. We consider existing gravitational wave measurements of the binary effective spin, the best-measured combination of spin parameters, in the four likely binary black hole detections GW150914, LVT151012, GW151226, and GW170104. If binary black hole spin magnitudes extend to high values we show that the data exhibit a 2.4σ $(0.015$ odds ratio) preference for an isotropic angular distribution over an aligned one. By considering the effect of 10 additional detections, we show that such an augmented data set would enable in most cases a preference stronger than 5σ $(2.9 \times 10^{-7}$ odds ratio). The existing preference for either an isotropic spin distribution or low spin magnitudes for the observed systems will be confirmed (or overturned) confidently in the near future.

Comments: 32 pages, 9 figures, code and docum

Subjects: **High Energy Astrophysical Phenon**

Report number: LIGO-P1700067

Cite as: **arXiv:1706.01385 [astro-ph.HE]**

(or **arXiv:1706.01385v2 [astro-ph.HE**

Vitale *et al*. arXiv: 1503.04307

Gerosa & Berti arXiv: 1703.06223

Fishbach, Holz & Farr arXiv:1703.06869

Stevenson, **CPLB** & Mandel: 1703.06873

Talbot & Thrane arXiv:1704.08370

Inference uses: waveform templates, noise spectrum, calibration uncertainty & a prior

We have observed a family of binary black holes

Masses and spins give hints to formation mechanisms, but need many detections

- Infer source properties by comparing waveform templates to data
- Some parameter (like chirp mass) measured well, others (like in-plane spins) remain uncertain
- Parameters (masses, spins, redshifts) could give insight into evolution
- Need a hierarchical analysis of a population of ~100 detections

LVC arXiv:1602.03840 arXiv:1602.03846 arXiv:1606.04856 arXiv:1706.01812

Thank you

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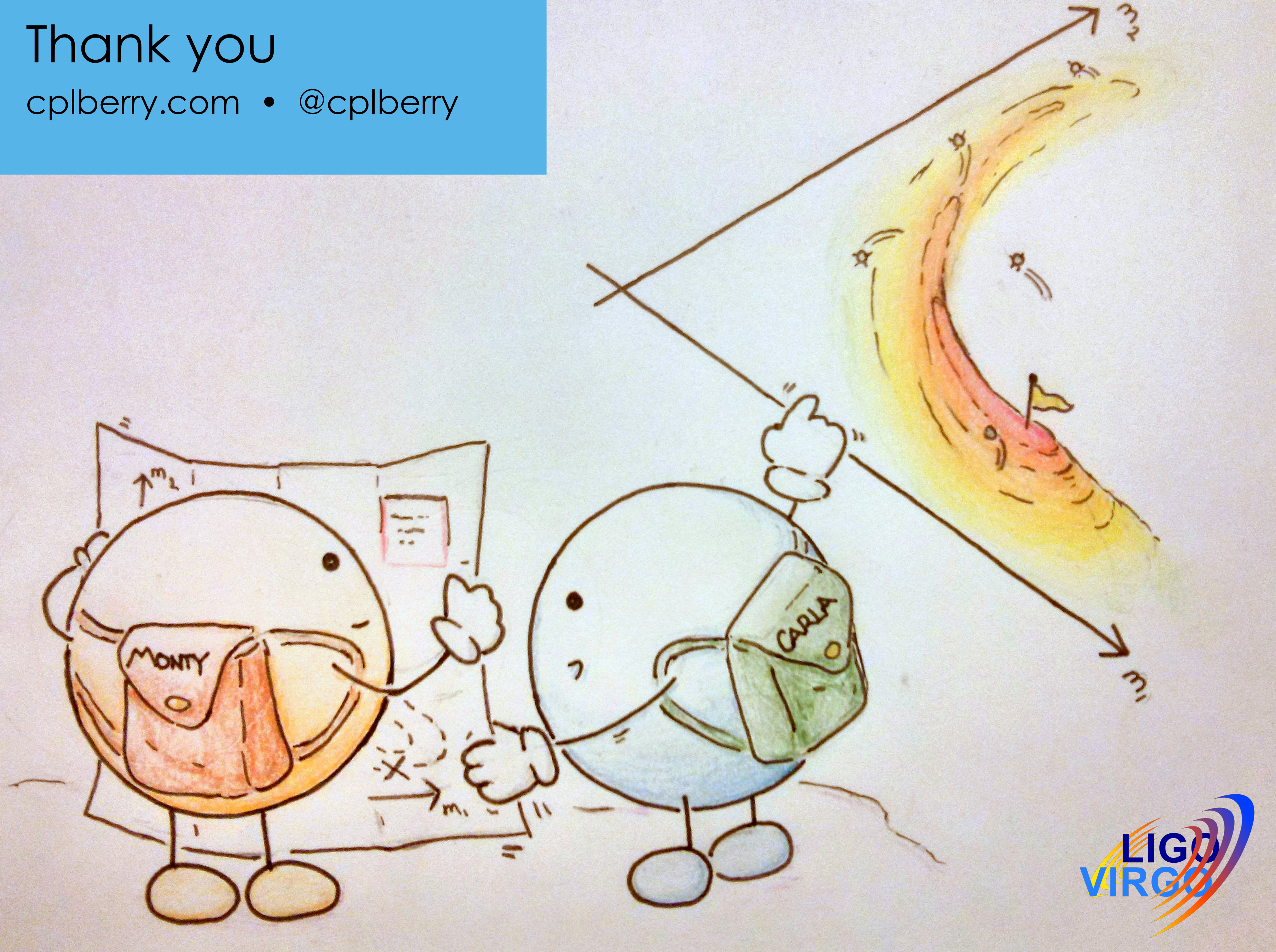


Bayes' theorem

$$p(\theta|d) = \frac{p(d|\theta)\,p(\theta)}{p(d)}$$

Likelihood

Prior

Posterior

Evidence

Bayes' theorem

Likelihood

Prior

$$p(\theta|d,\lambda) = \frac{p(d|\theta,\lambda)\,p(\theta|\lambda)}{p(d|\lambda)}$$

Posterior

Evidence

Bayes' theorem

Evidence

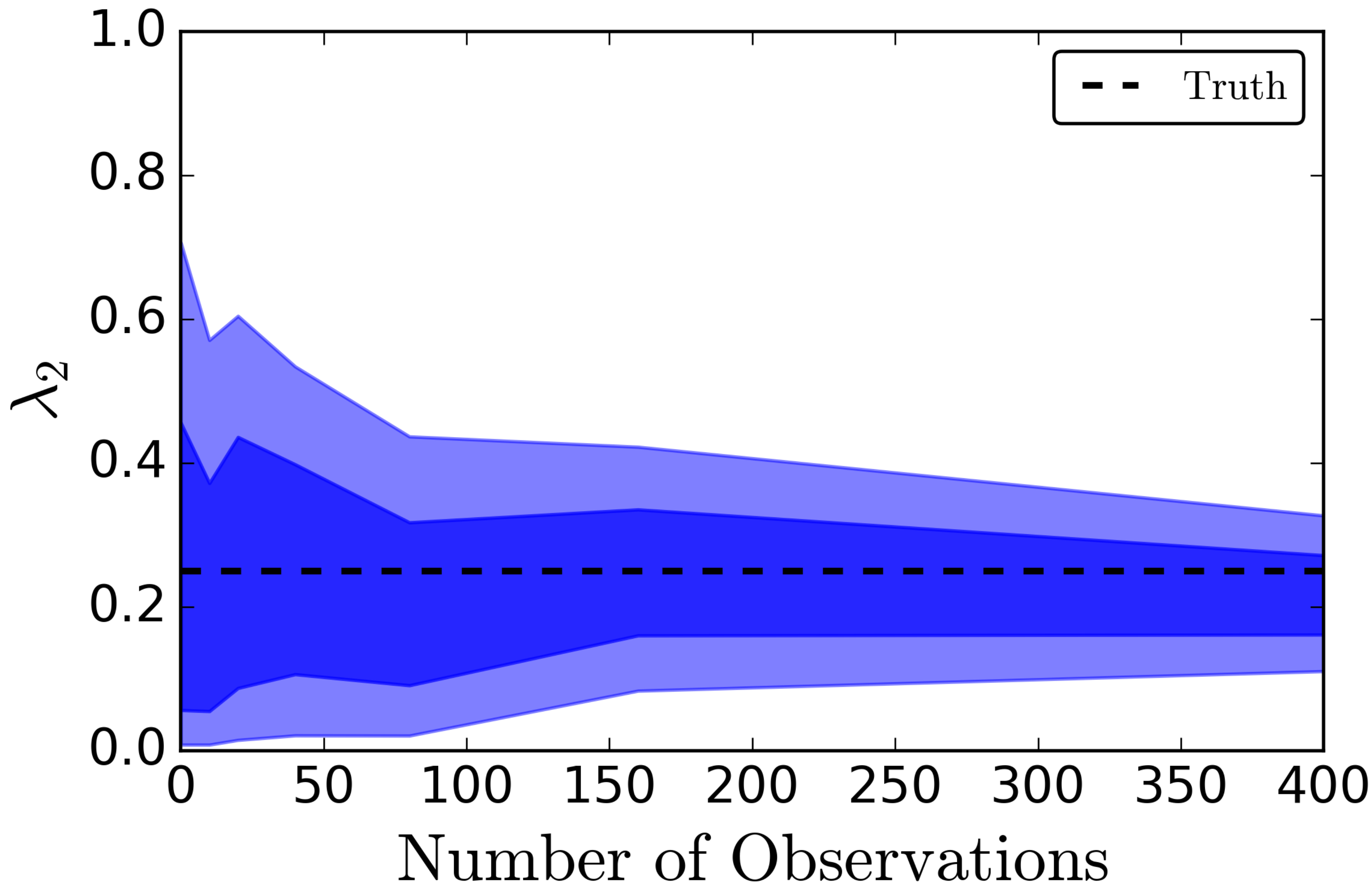
Model prior

$$p(\lambda|\{d\}) = \frac{p(\{d\}|\lambda)\, p(\lambda)}{p(\{d\})}$$

Model posterior

Model inference

Stevenson, CPLB & Mandel
arXiv:1703.06873



Model inference

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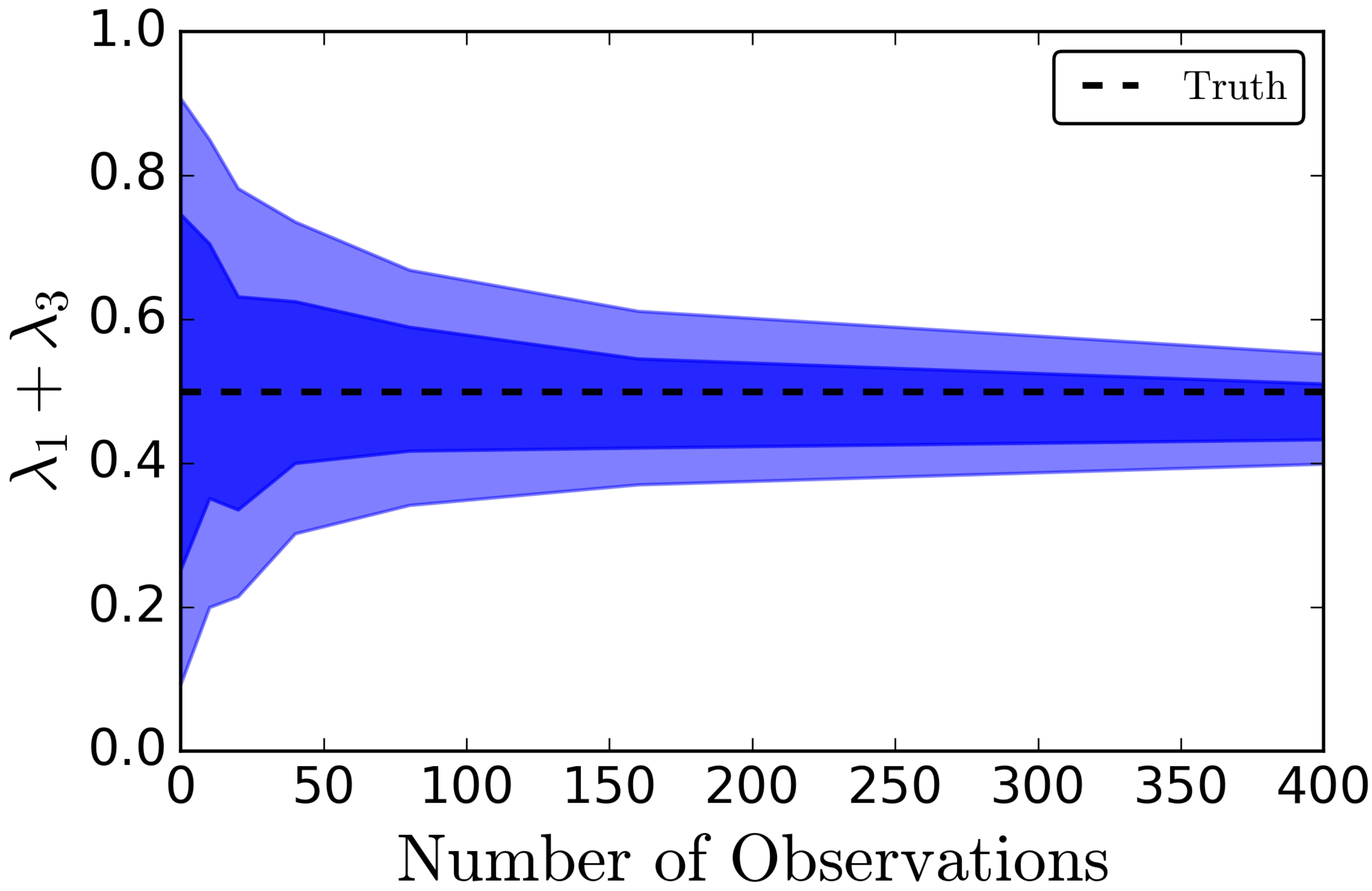
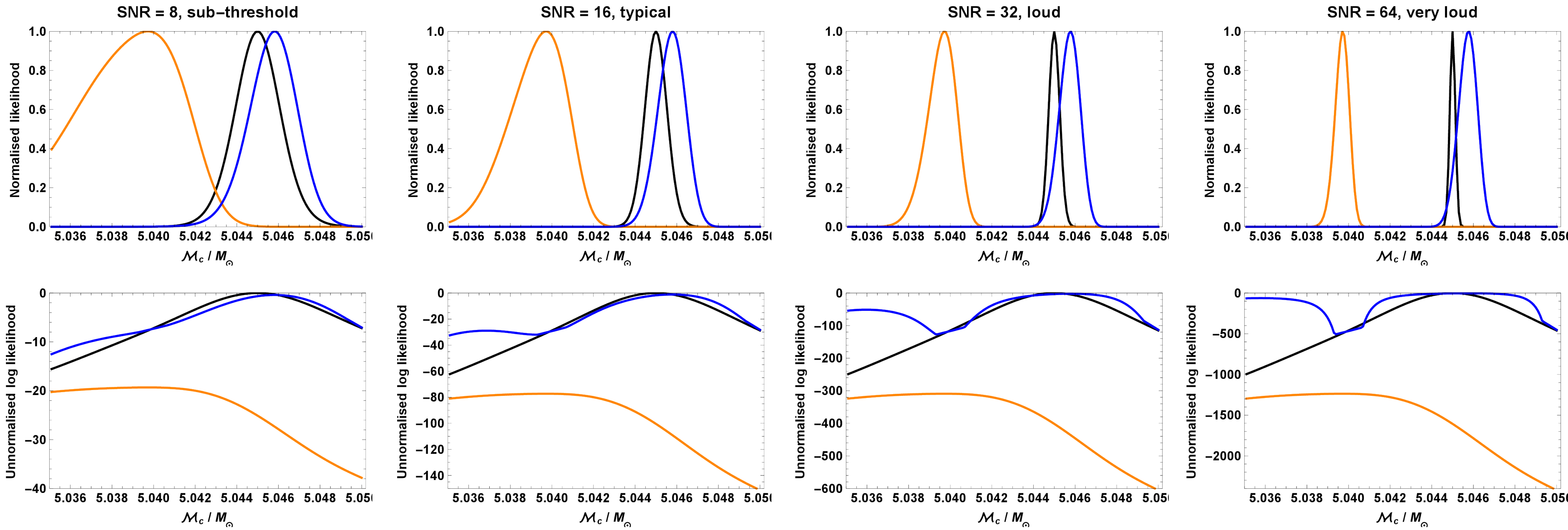


Waveform error

Waveforms introduce theoretical error (arXiv:0707.2982).
Mitigated using Gaussian processes (arXiv:1509.04066).



Moore *et al*. arXiv:1509.04066

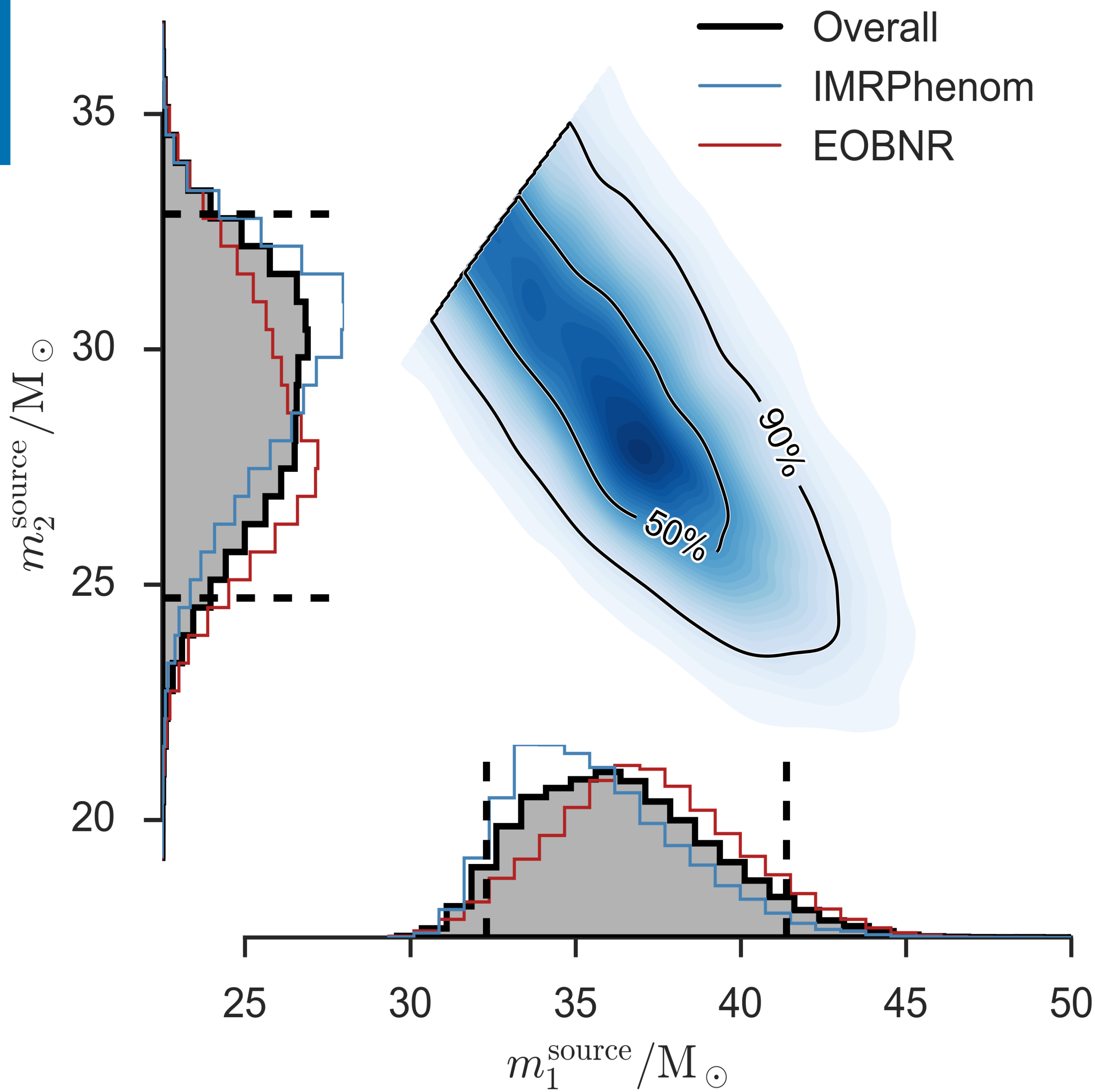
Credit: ButterflyLove1

Chirp mass

$$\mathcal{M}_{\rm c} = \frac{(m_1 m_2)^{3/5}}{(m_1 + m_2)^{1/5}}$$

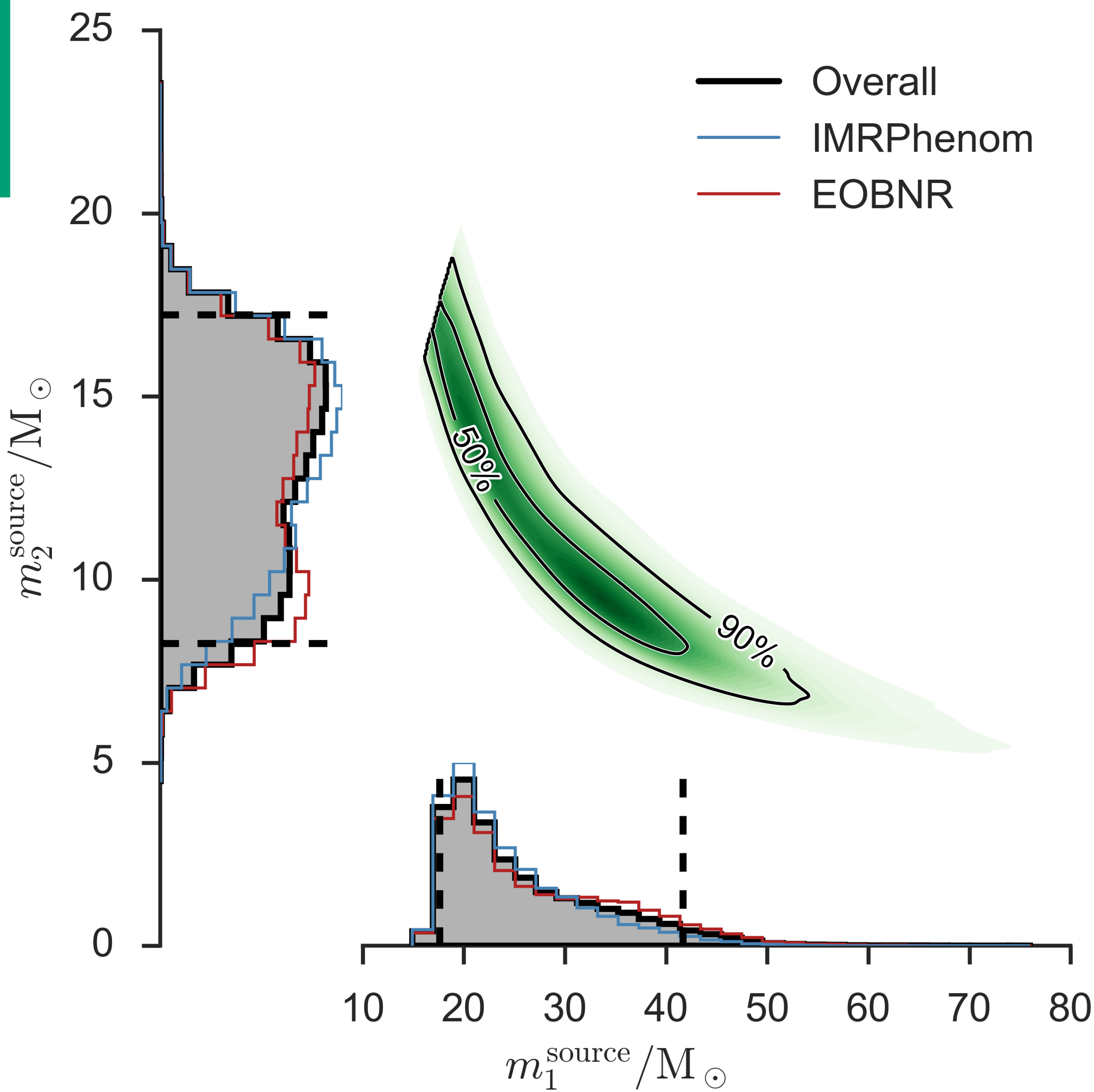
Chirp mass gives leading-order amplitude and phase evolution (Sathyaprakash & Schutz arXiv:0903.0338)

Masses



LVC
arXiv:1606.04856
arXiv:1602.03840

Masses



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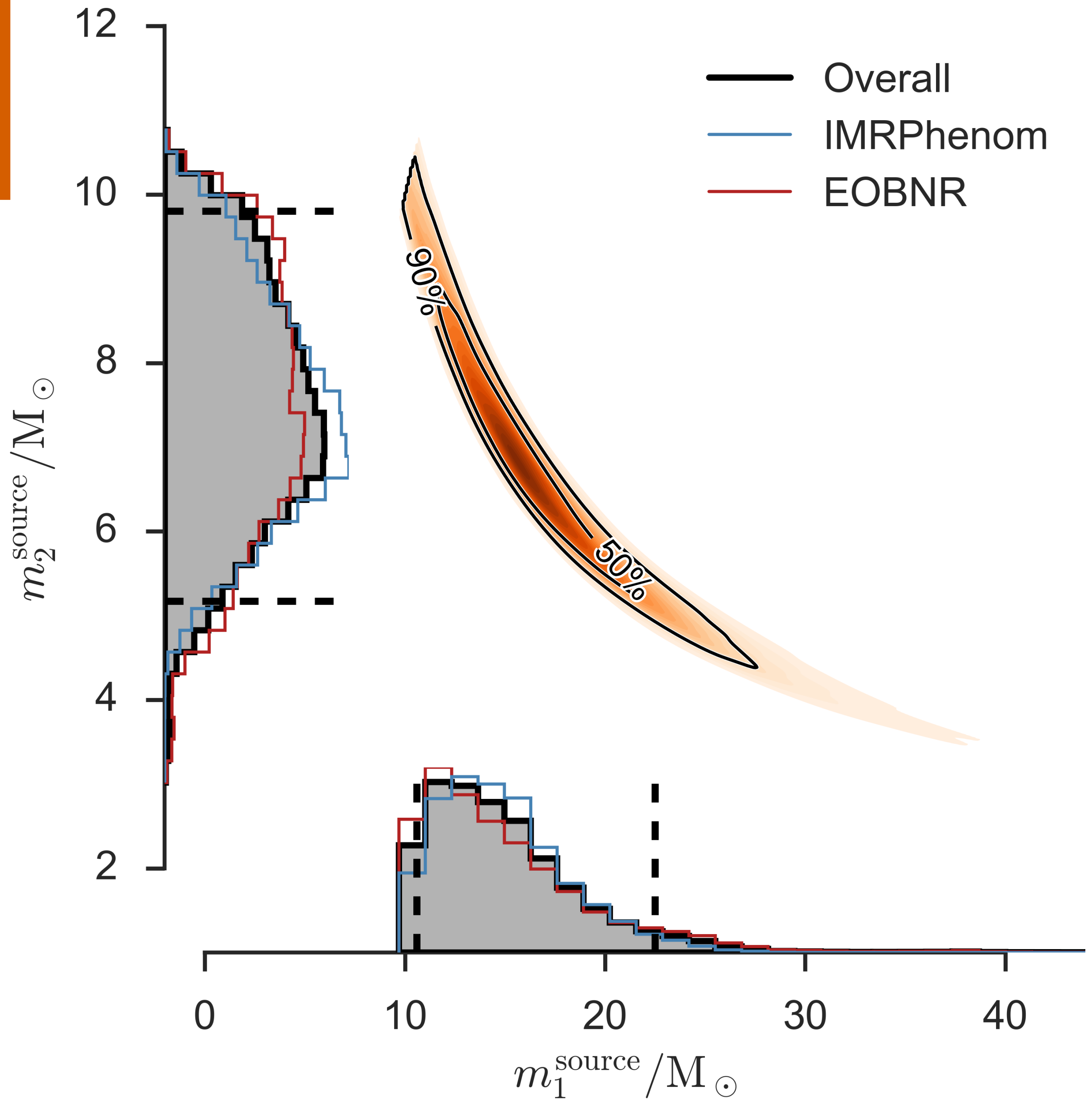
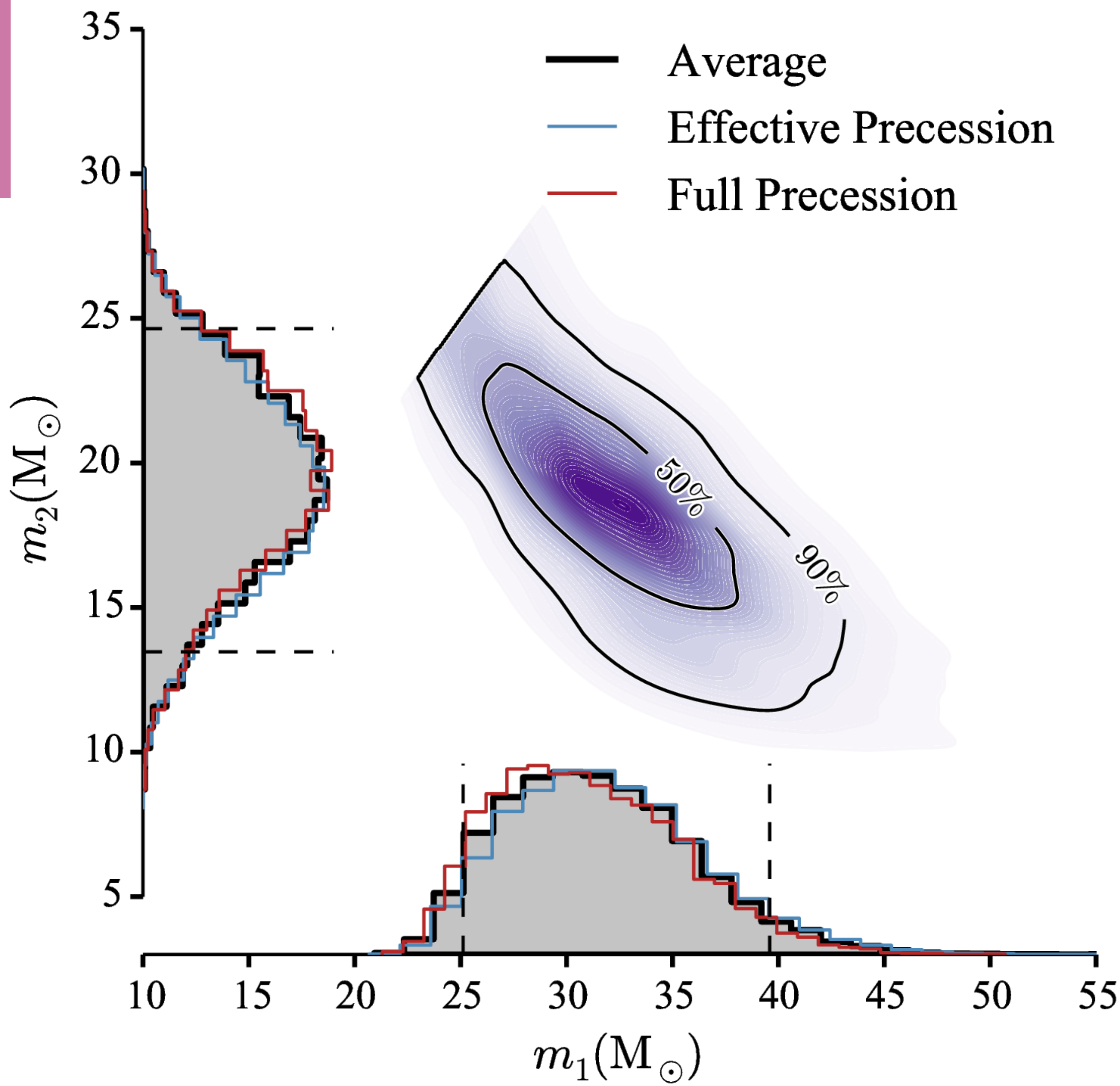


LVC
arXiv:1606.04855

Masses



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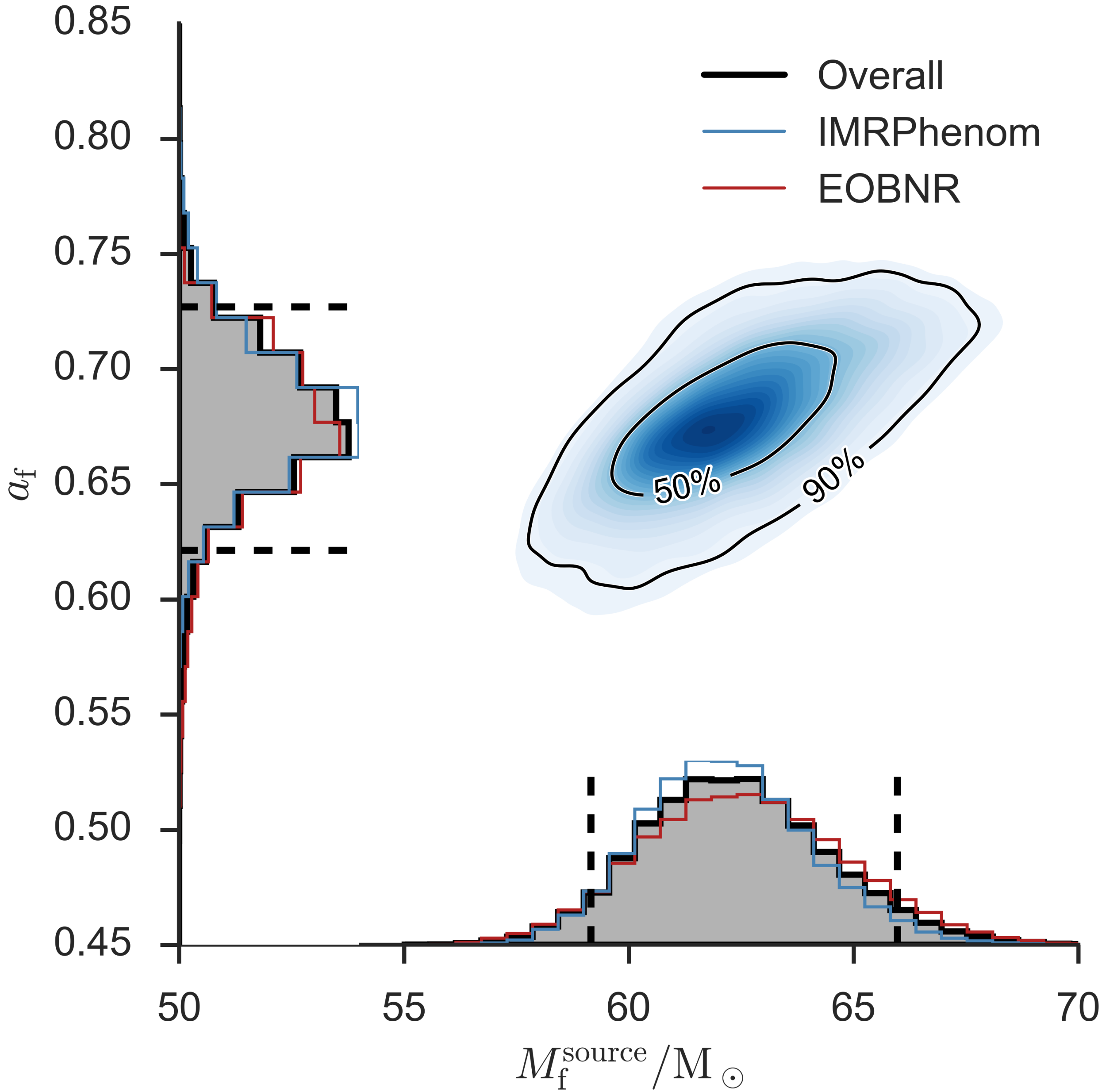
arXiv:1706.01812

Final mass & spin



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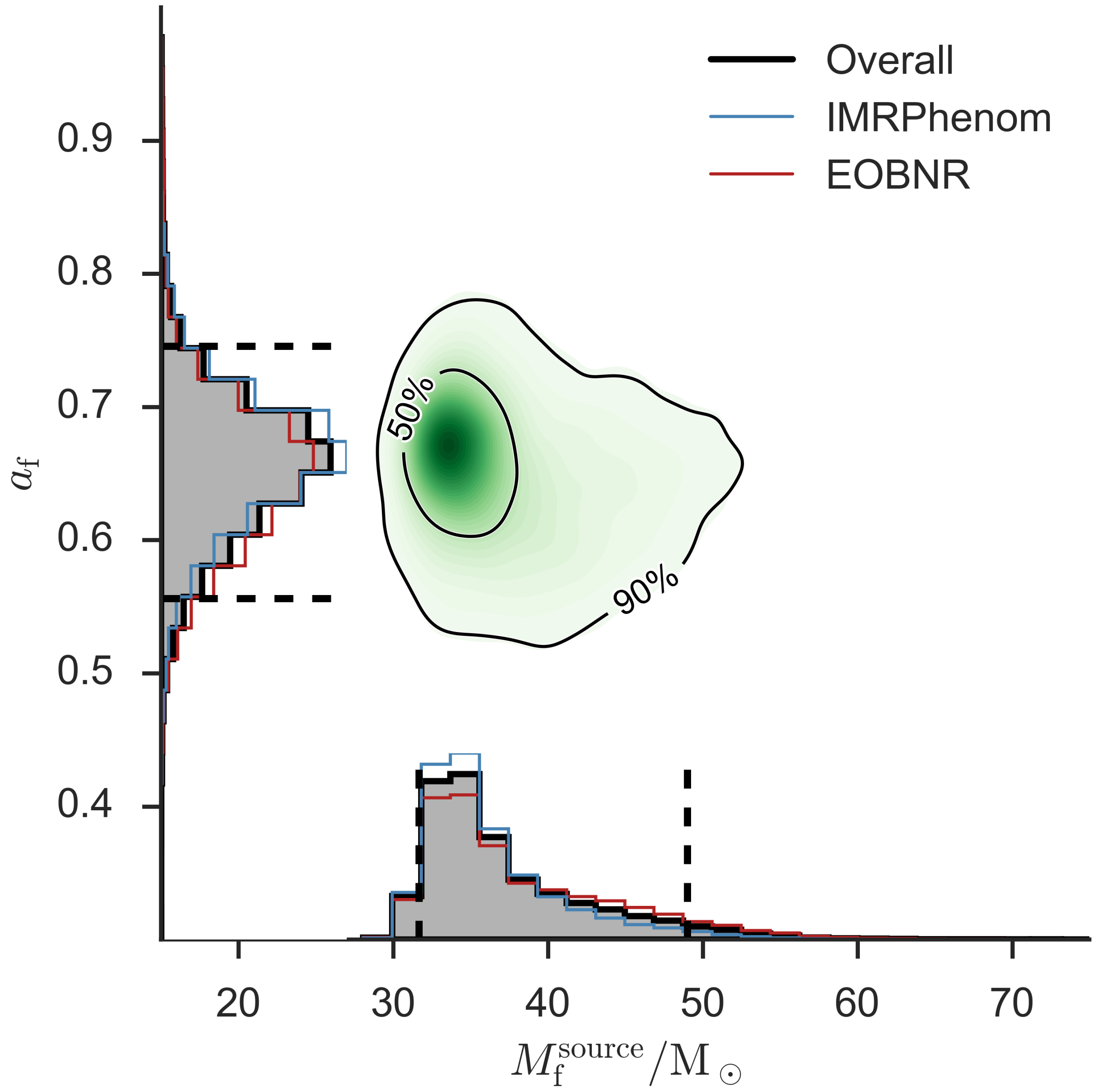
Final mass & spin



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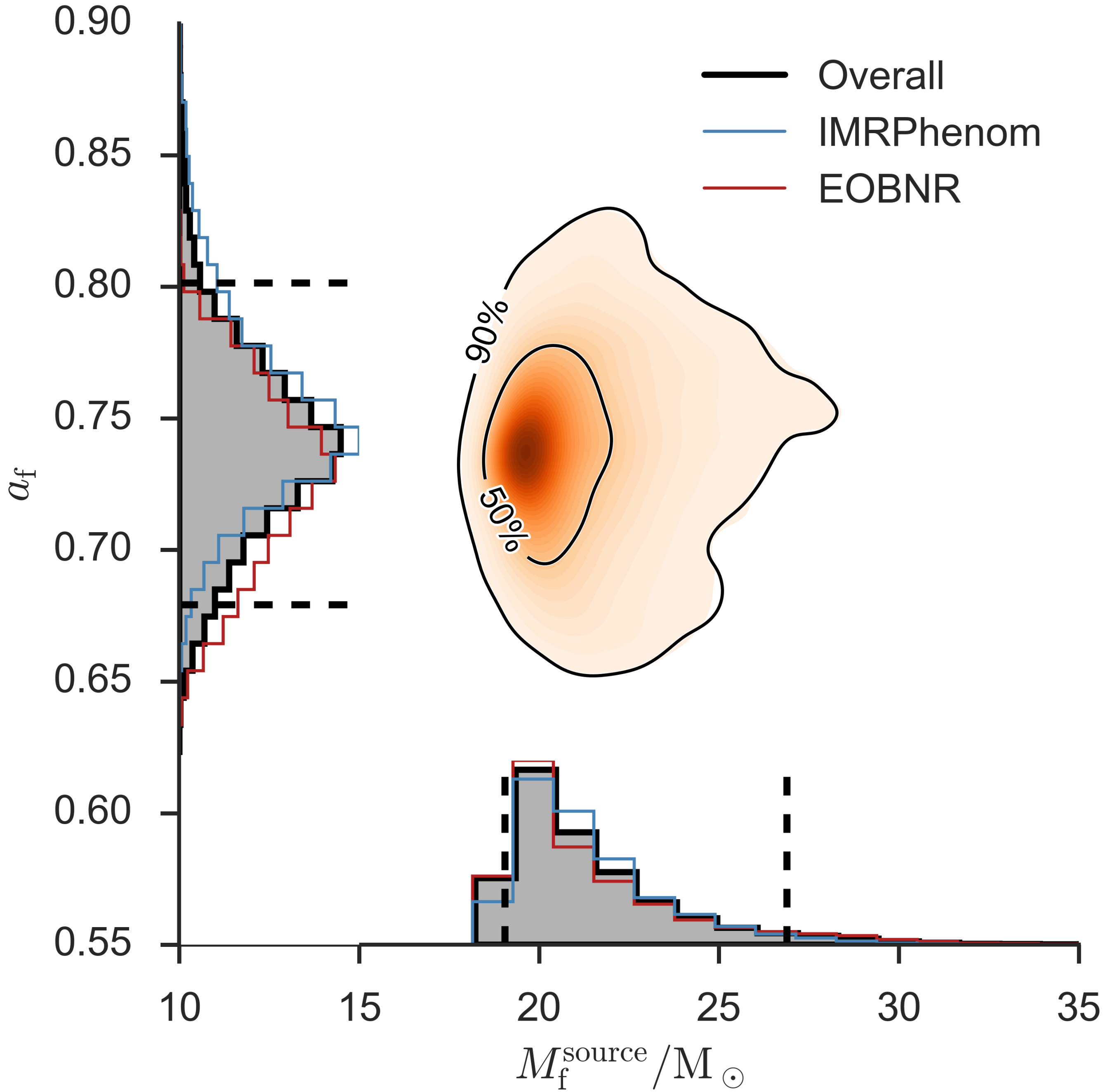

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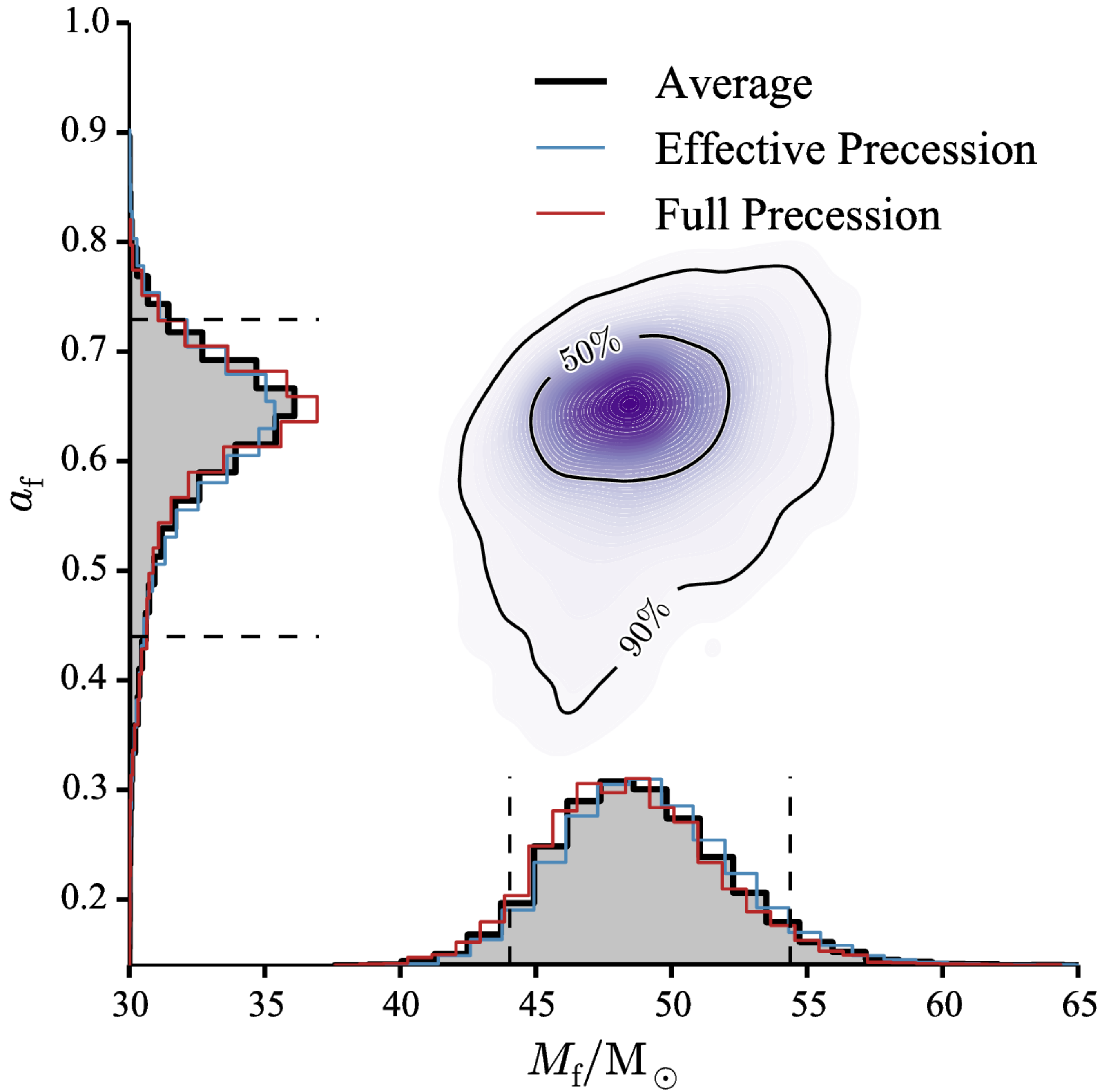

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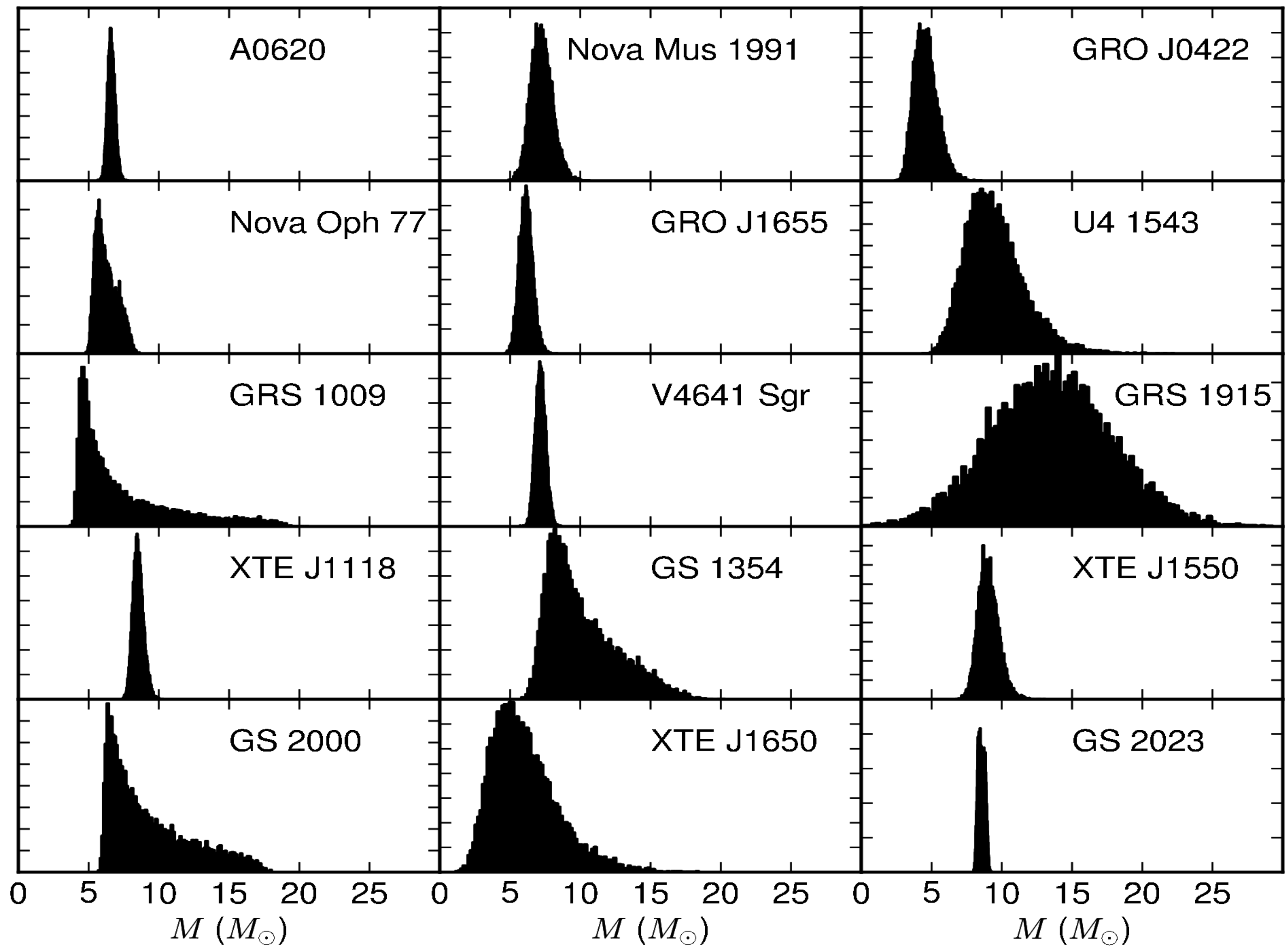
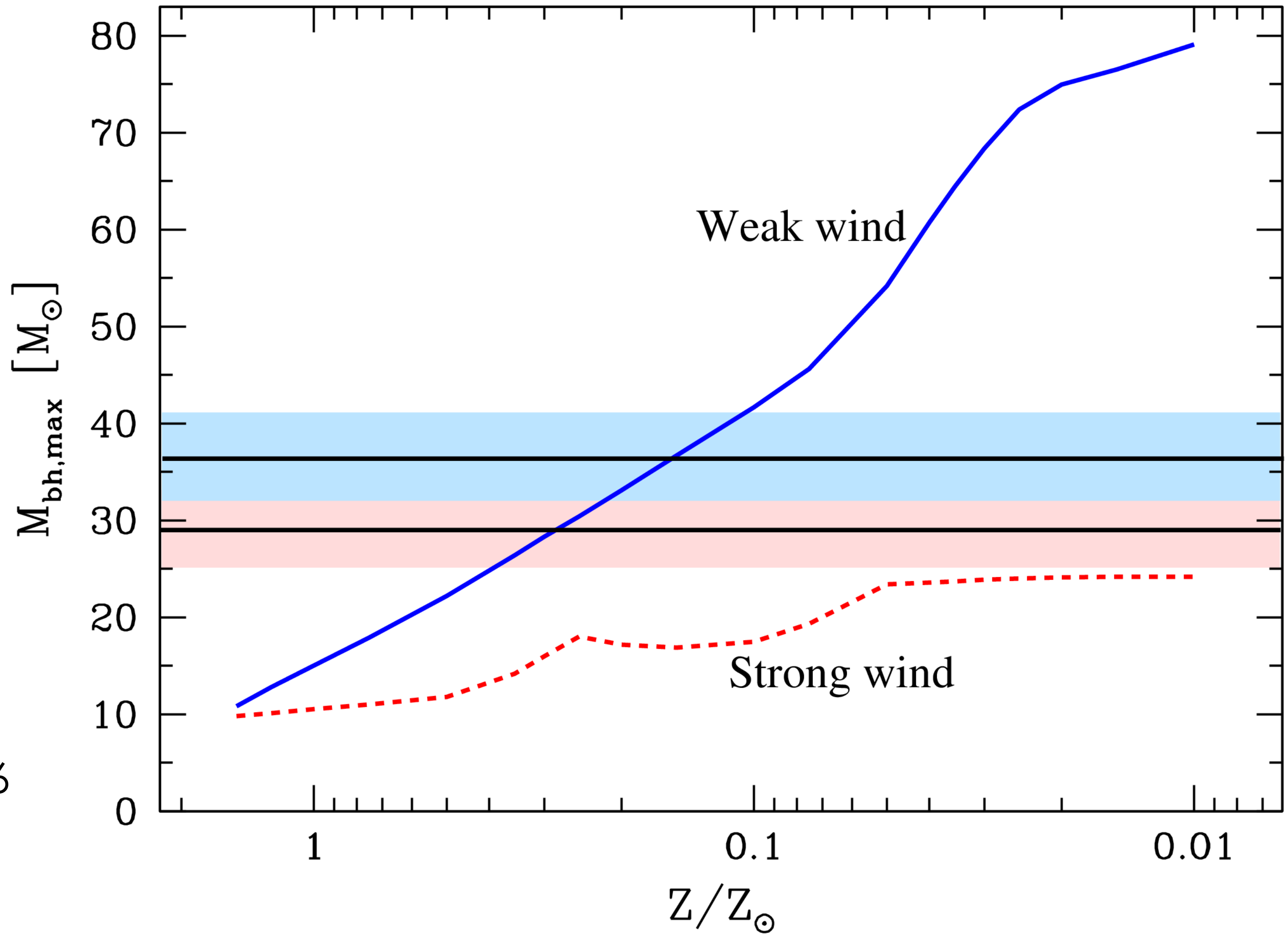
Farr *et al.* arXiv:1011.1459



Metallicity



LVC
arXiv:1602.03846
Belczynski *et al*.
arXiv:0904.2784

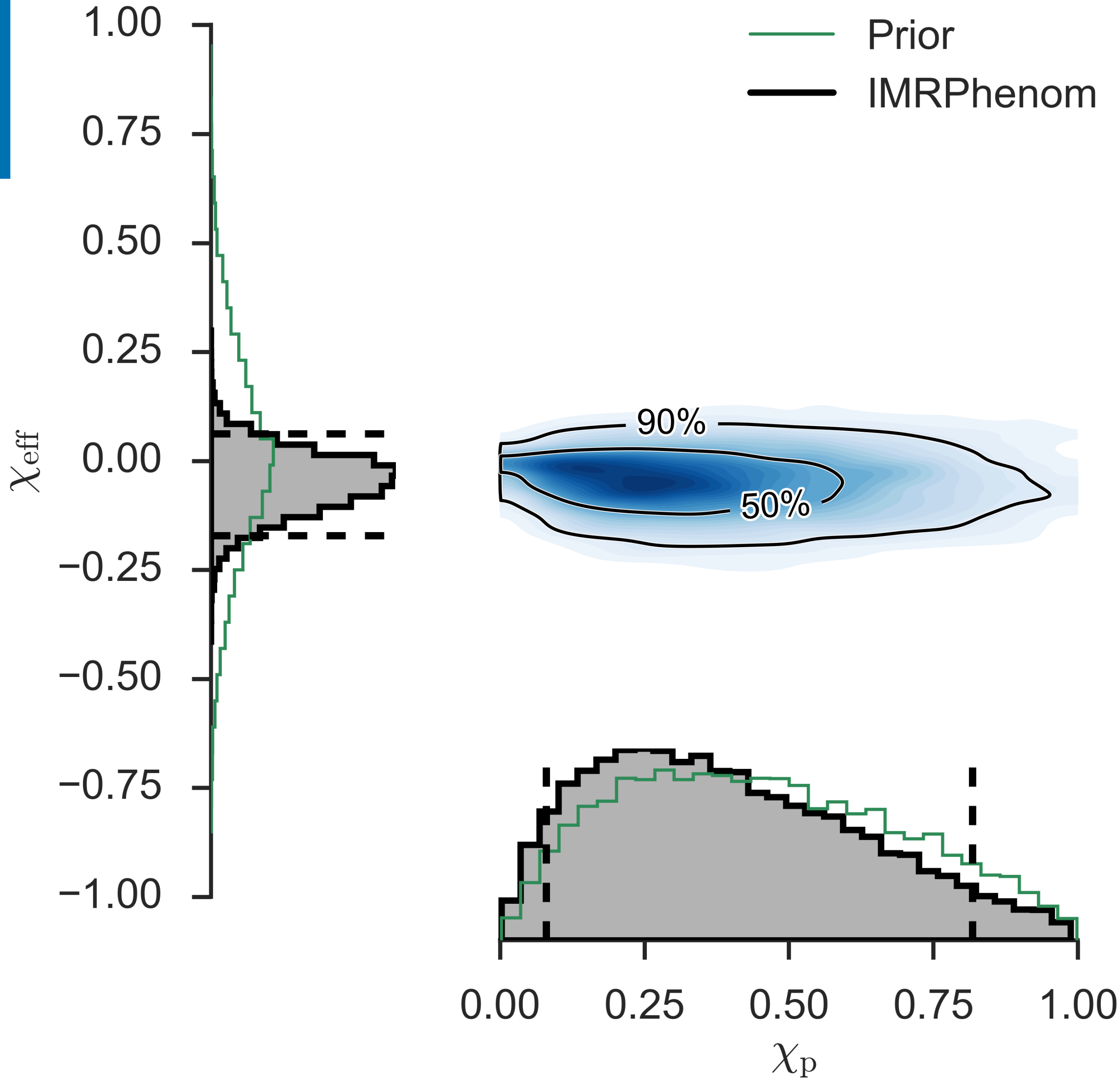
No hair theorem

Black holes have:

1. Mass
2. Spin
3. Electric charge

Image: Matt Groening

Spin



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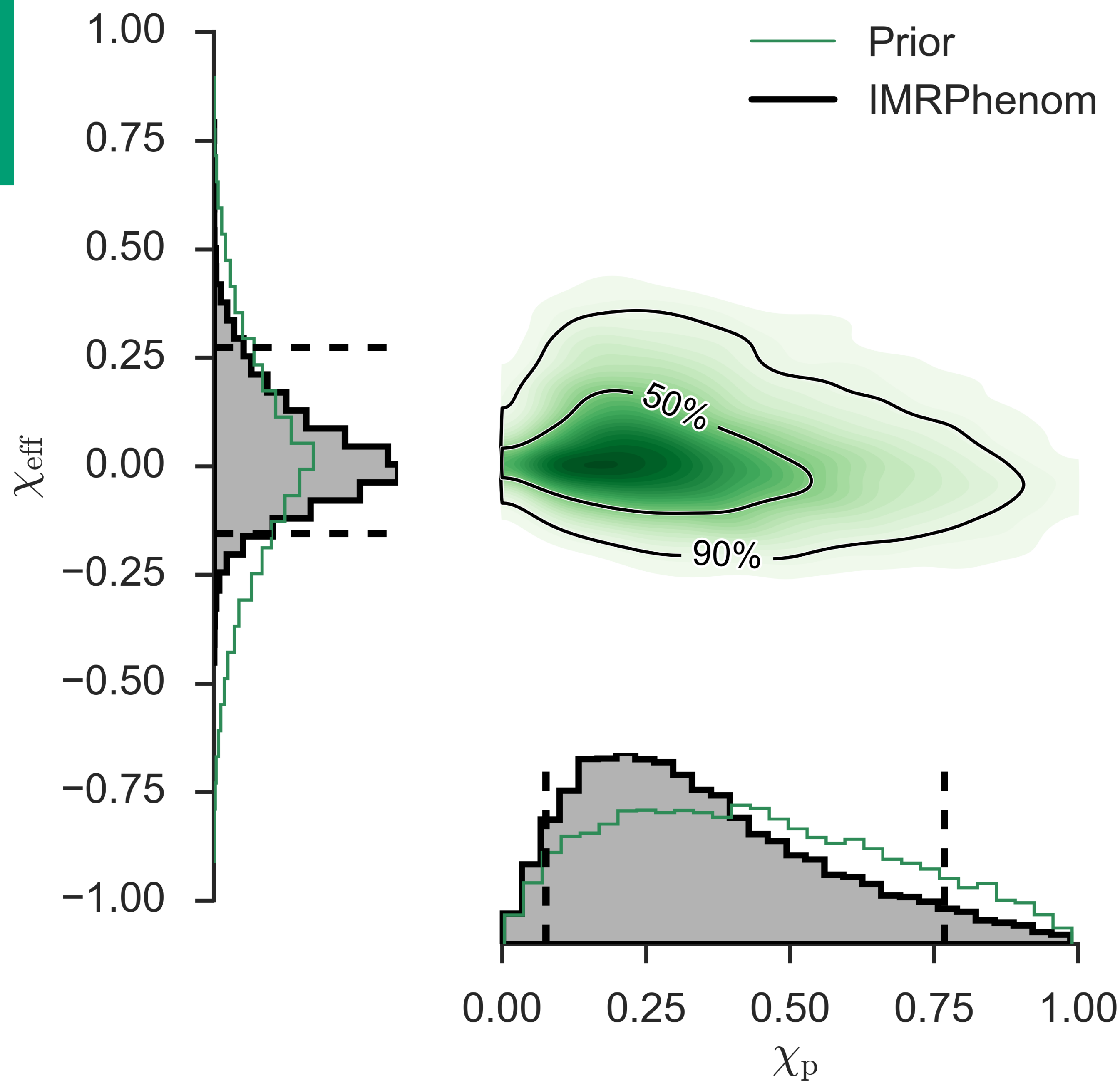
Spin



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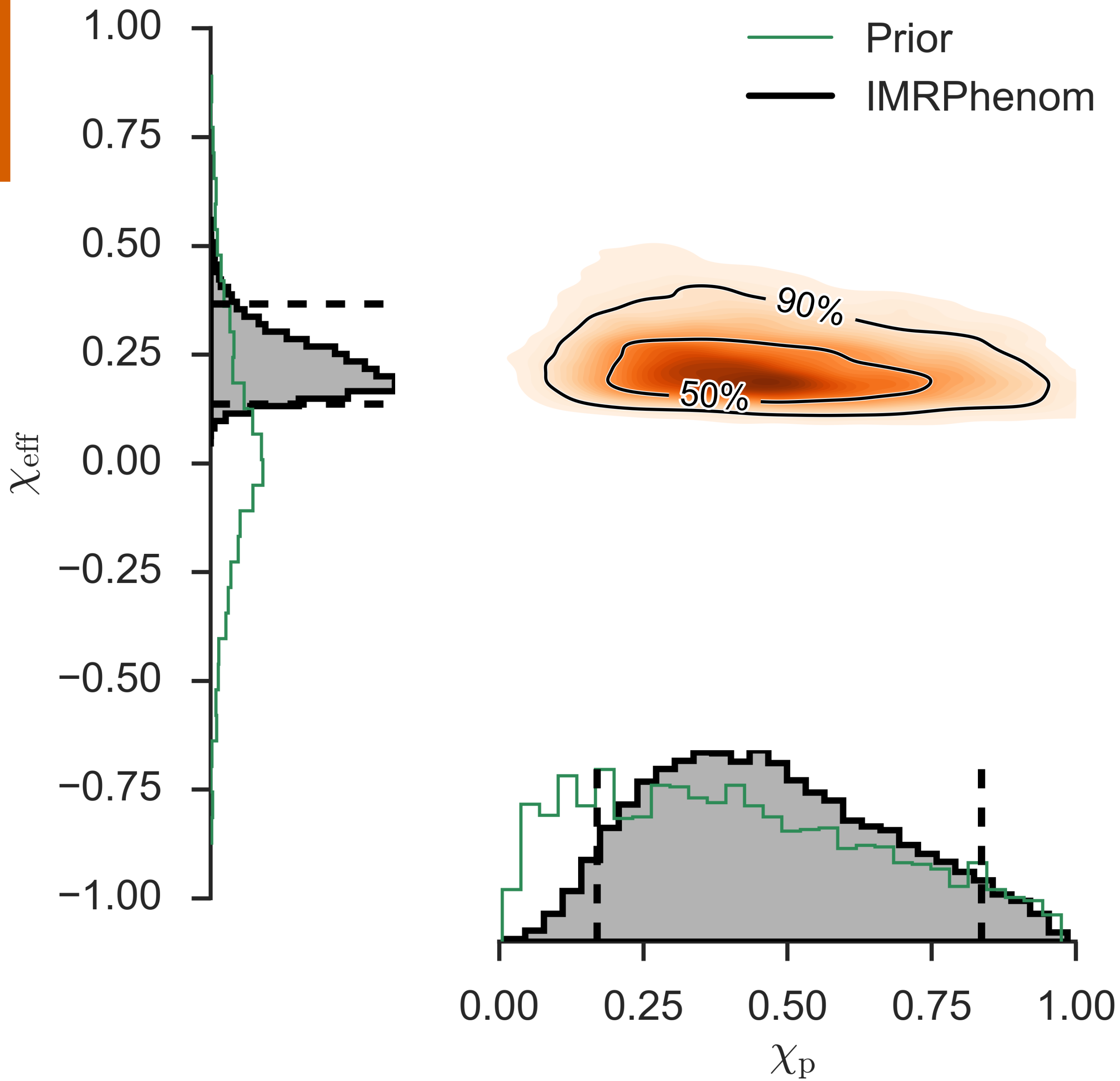

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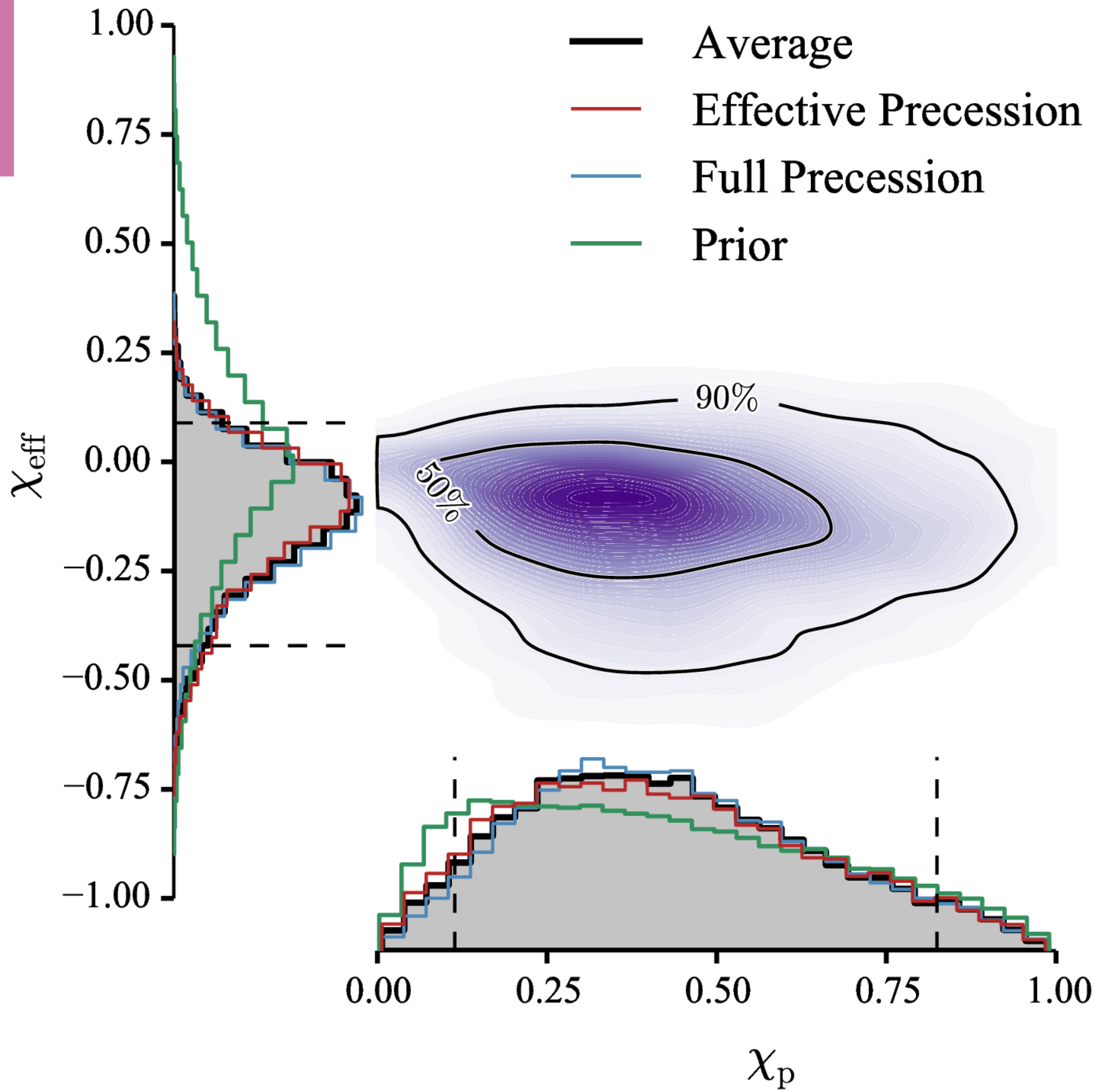

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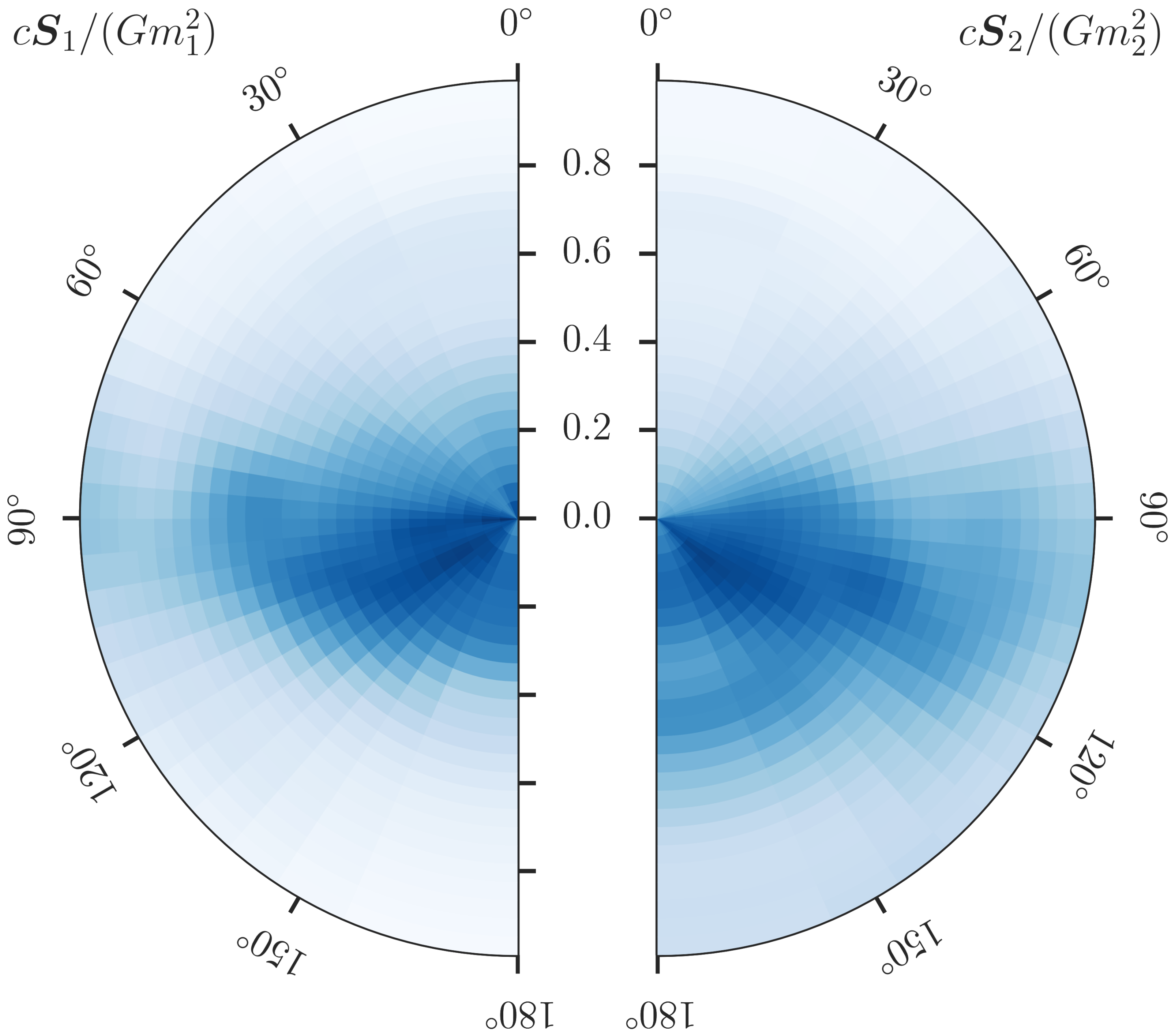

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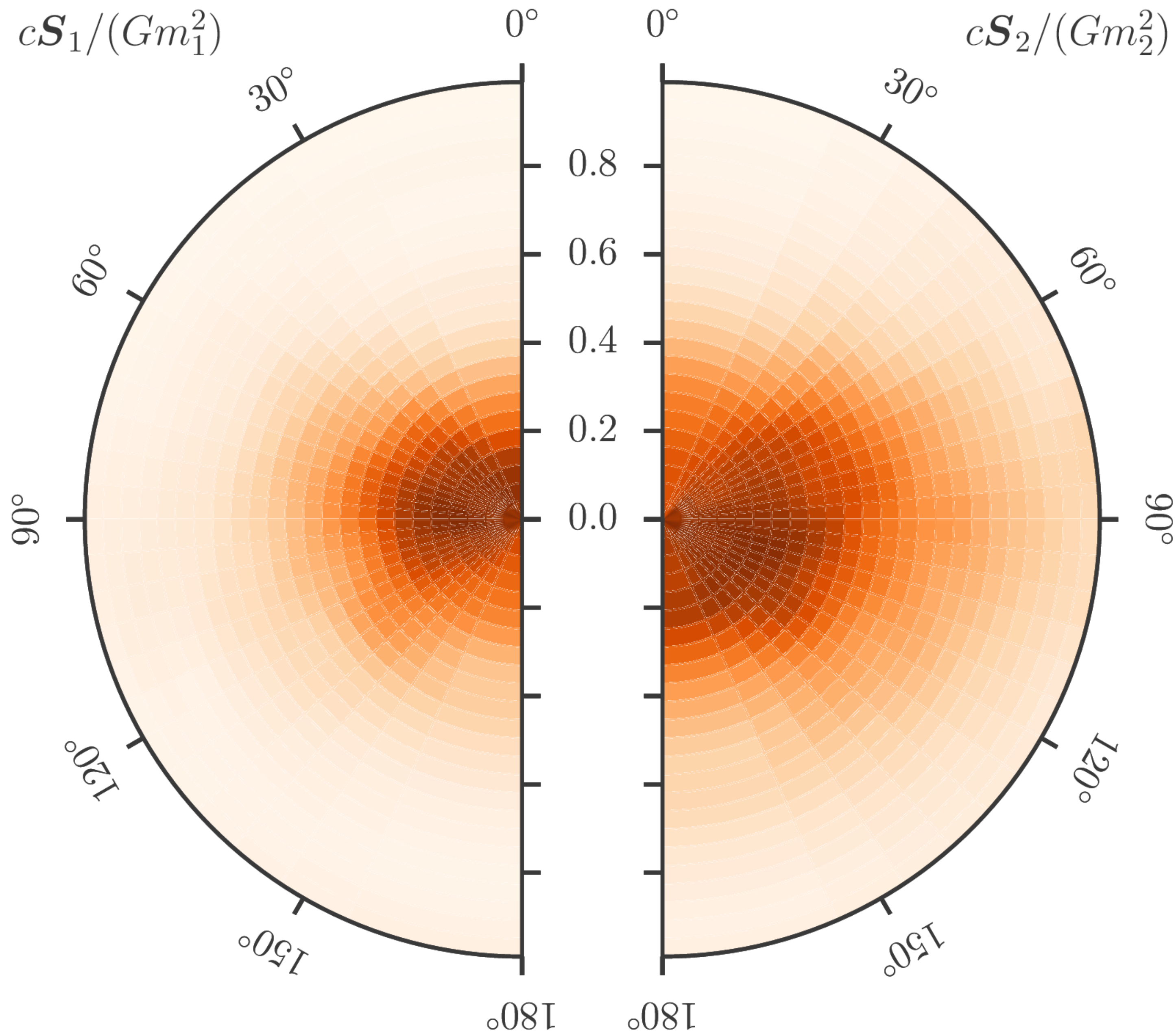
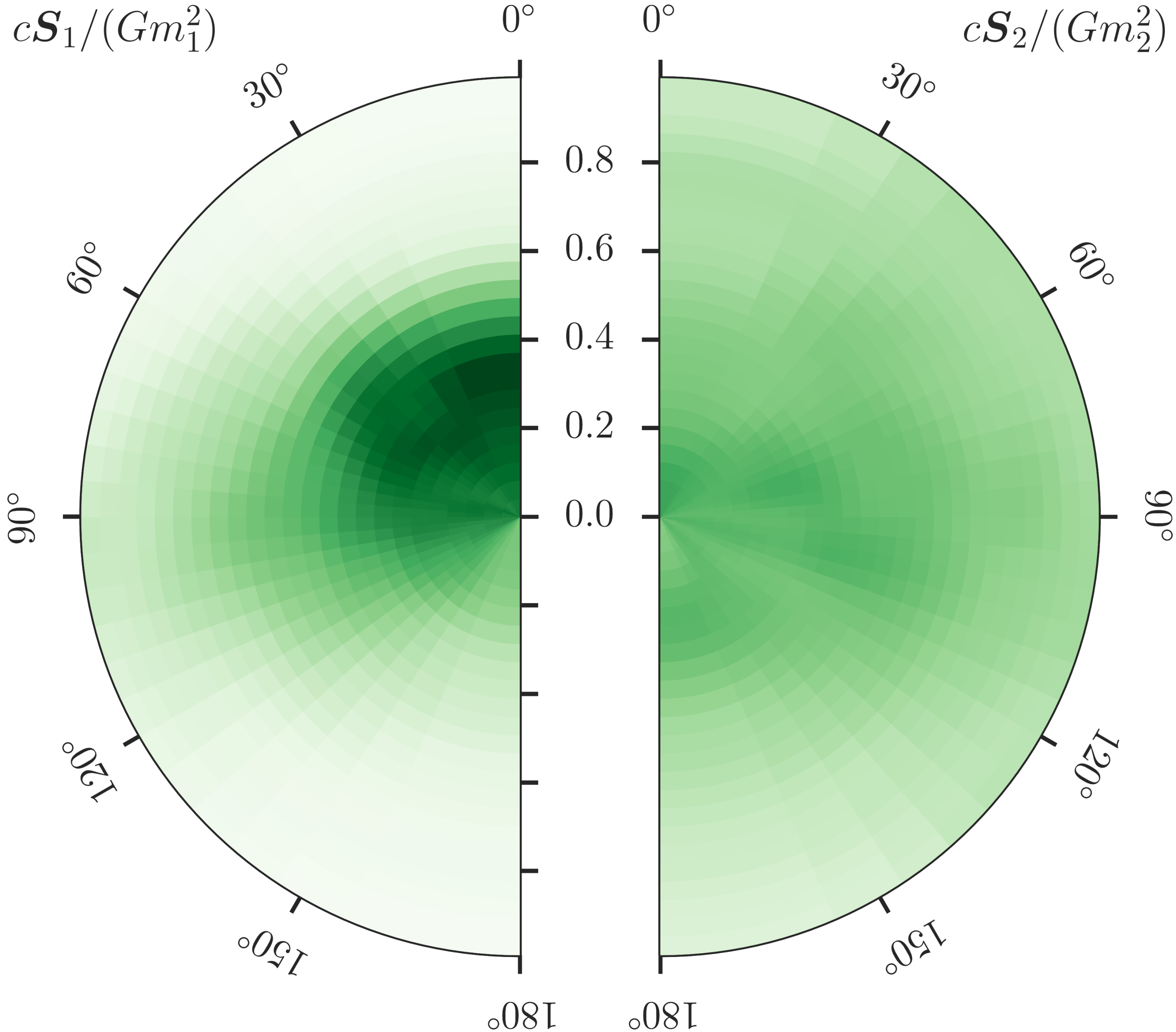


LVC
arXiv:1606.01210

Spin



LVC
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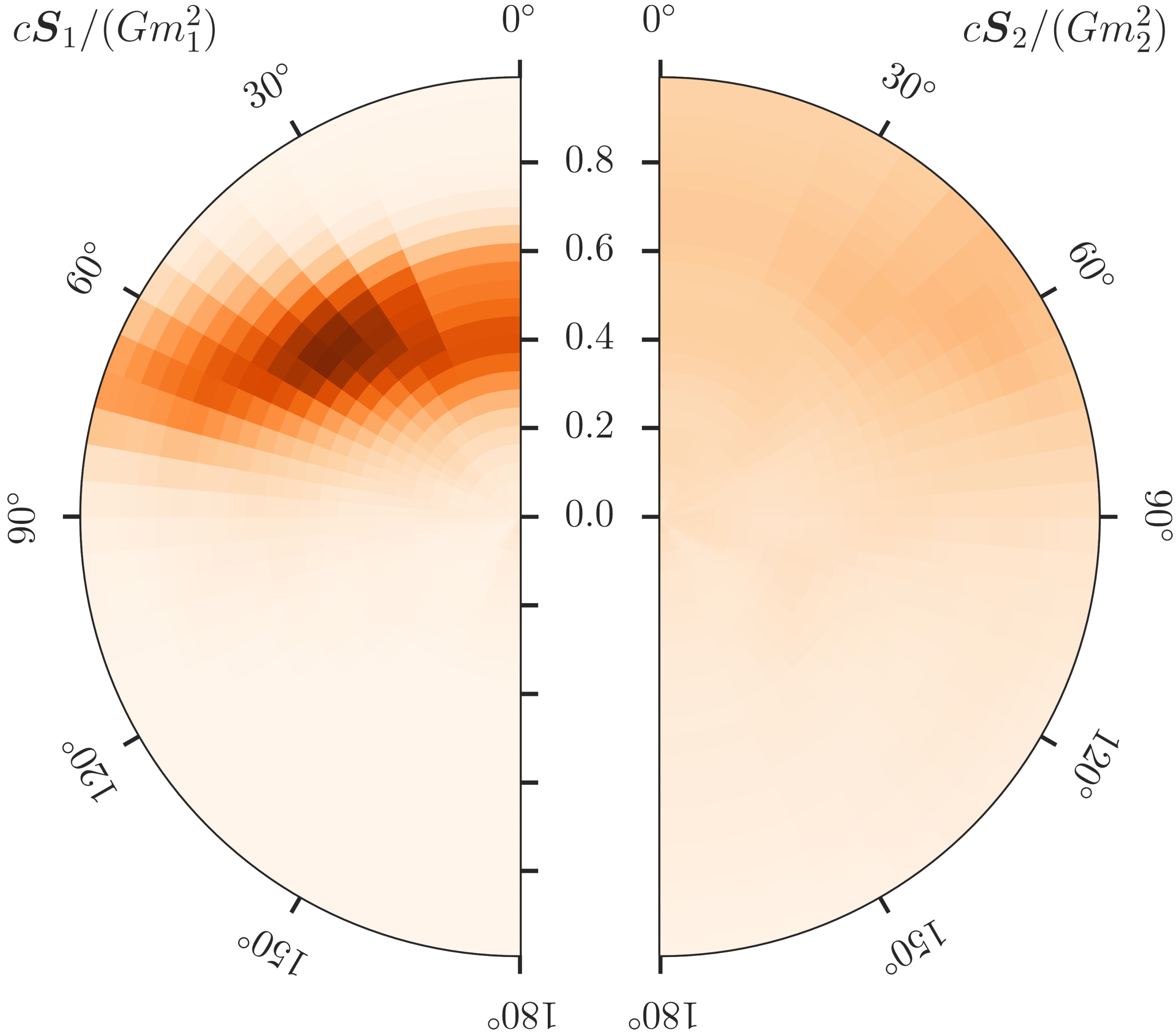
Spin



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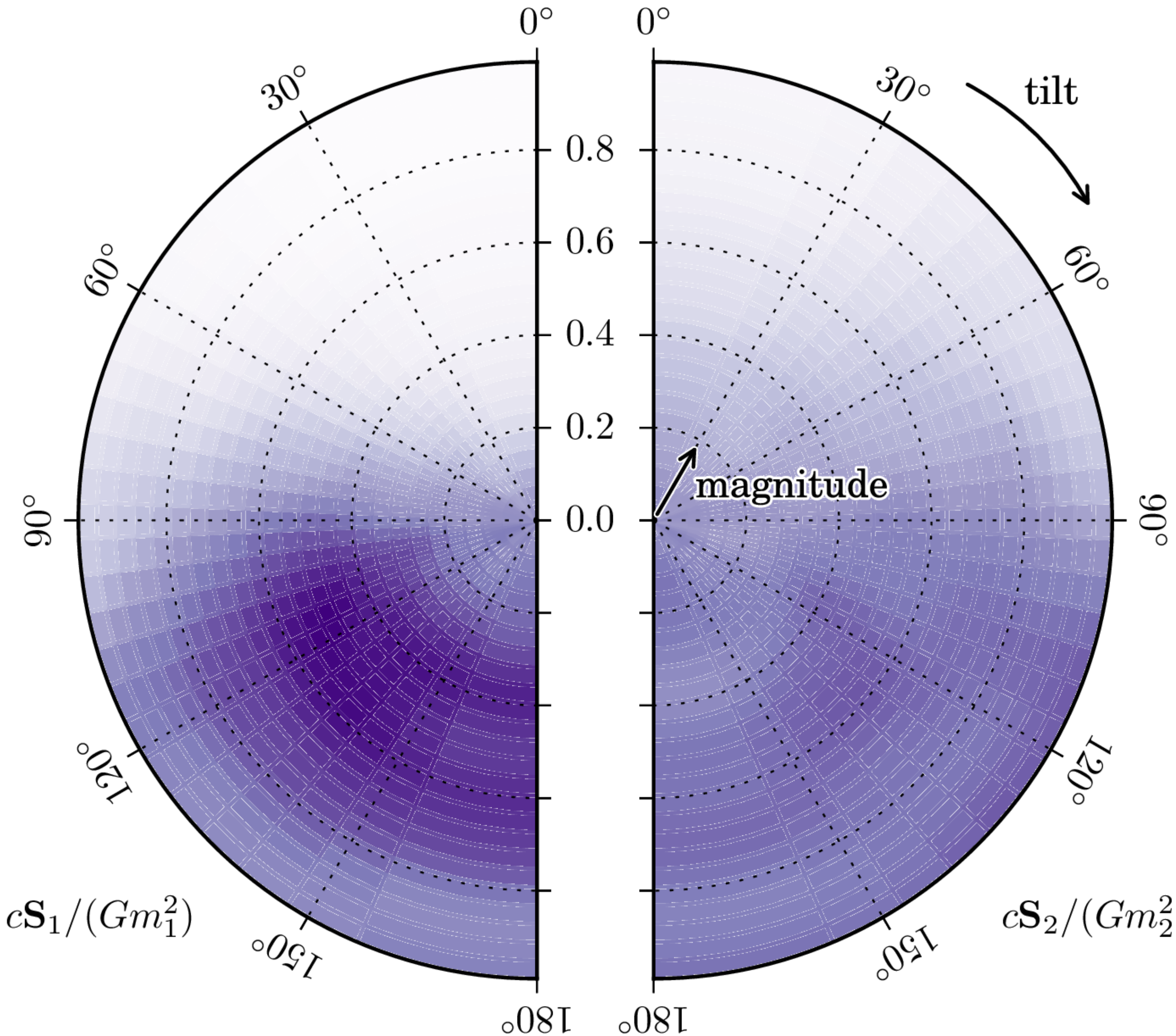

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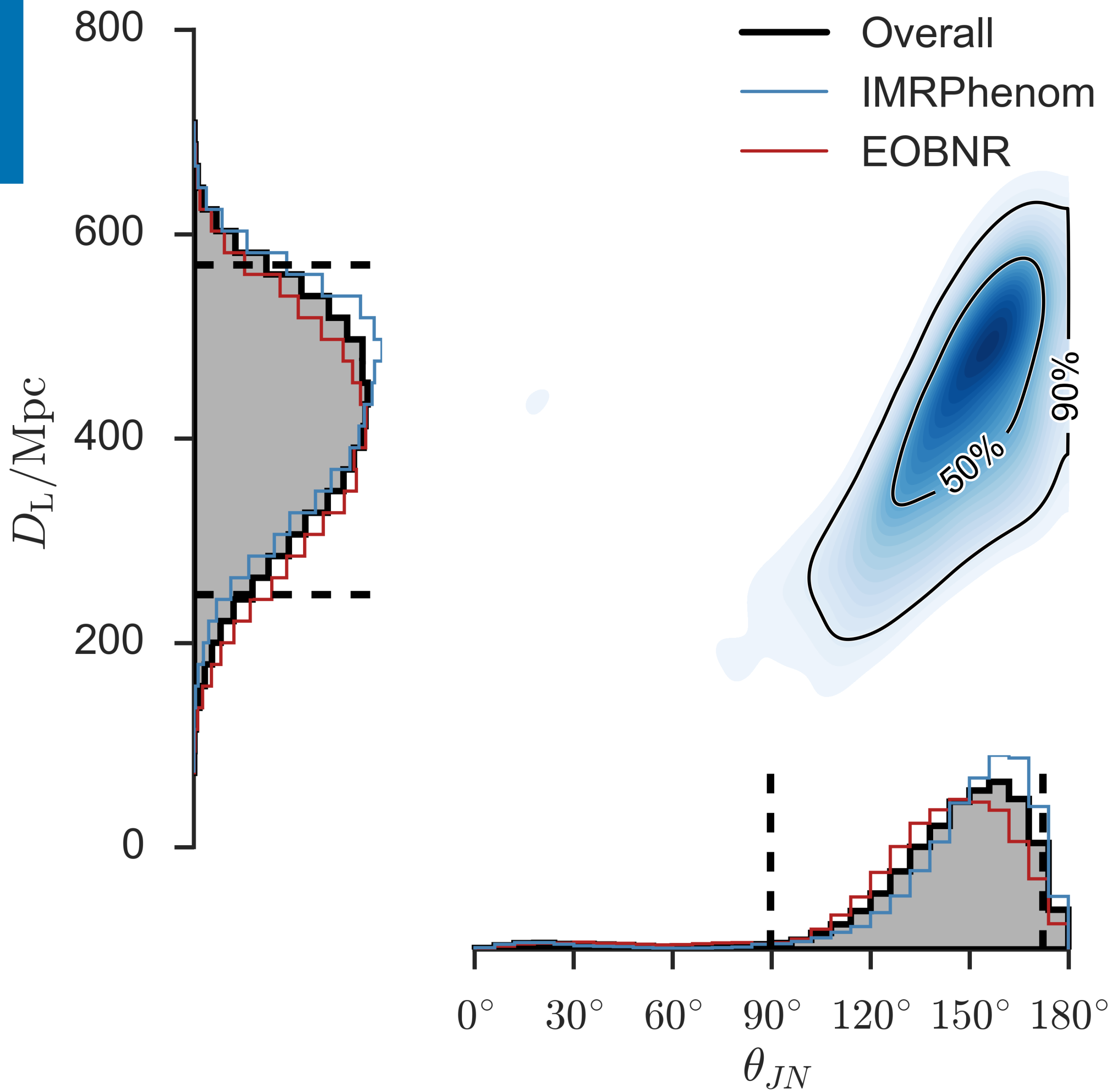
Distance



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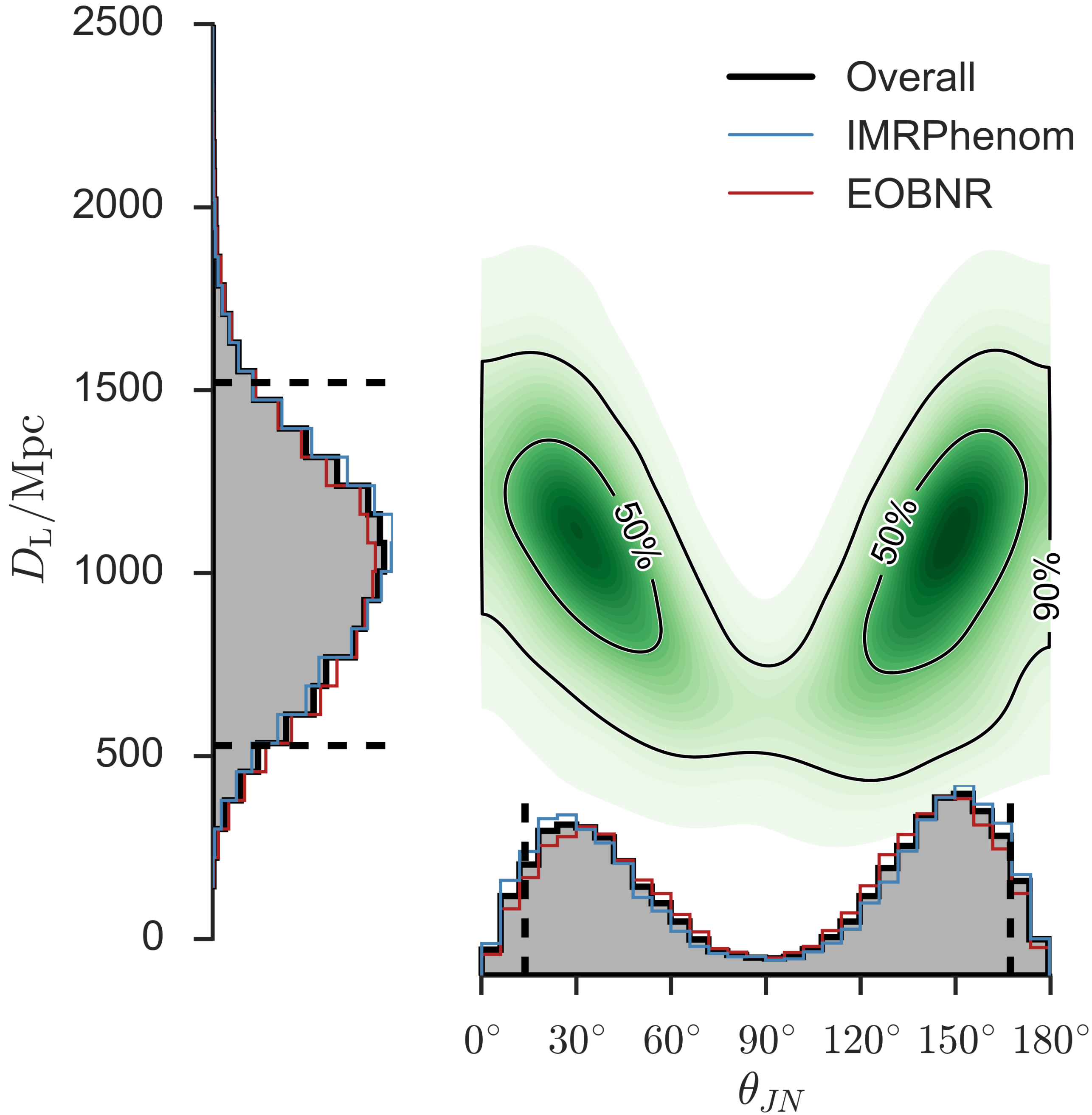

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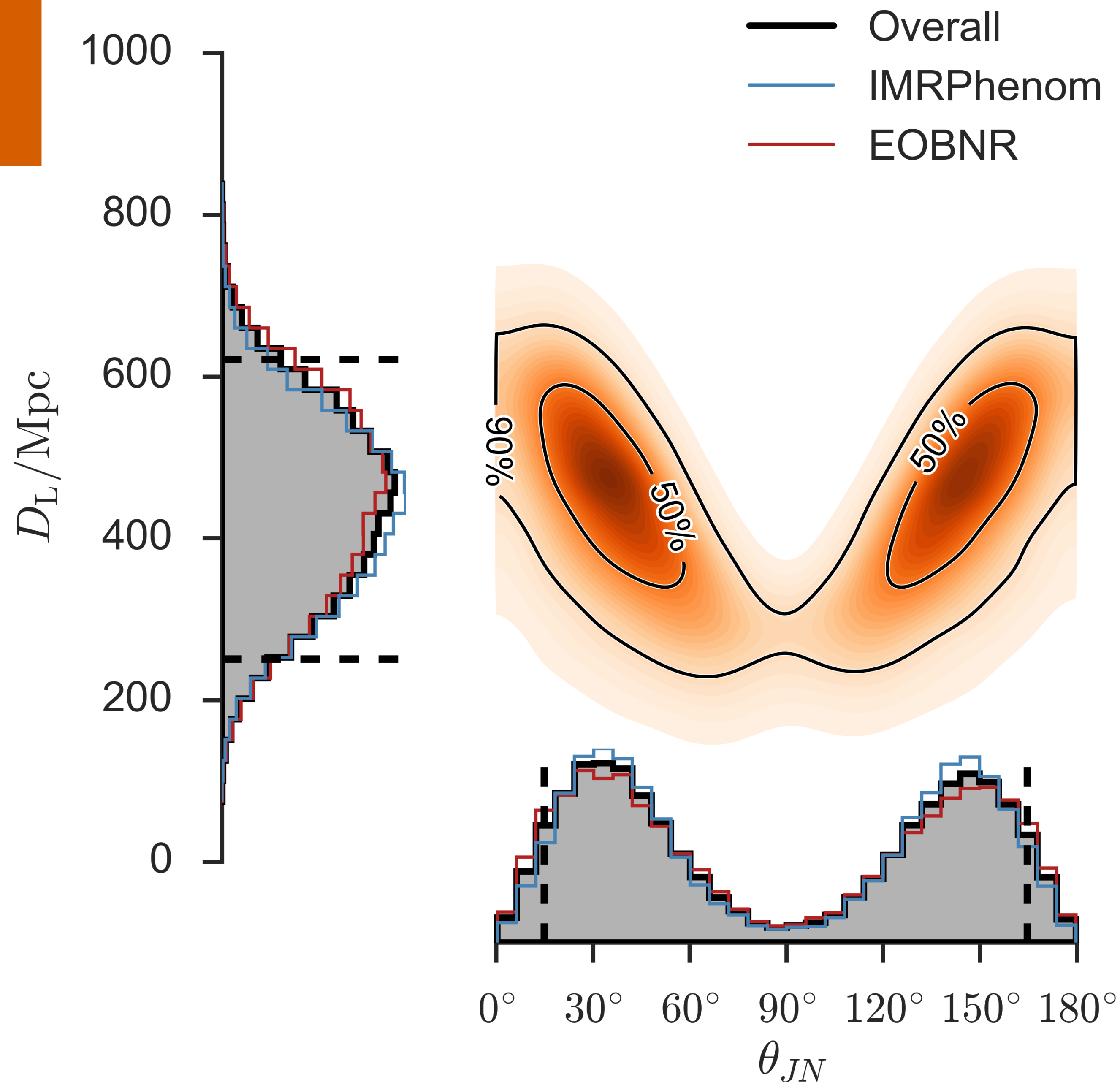

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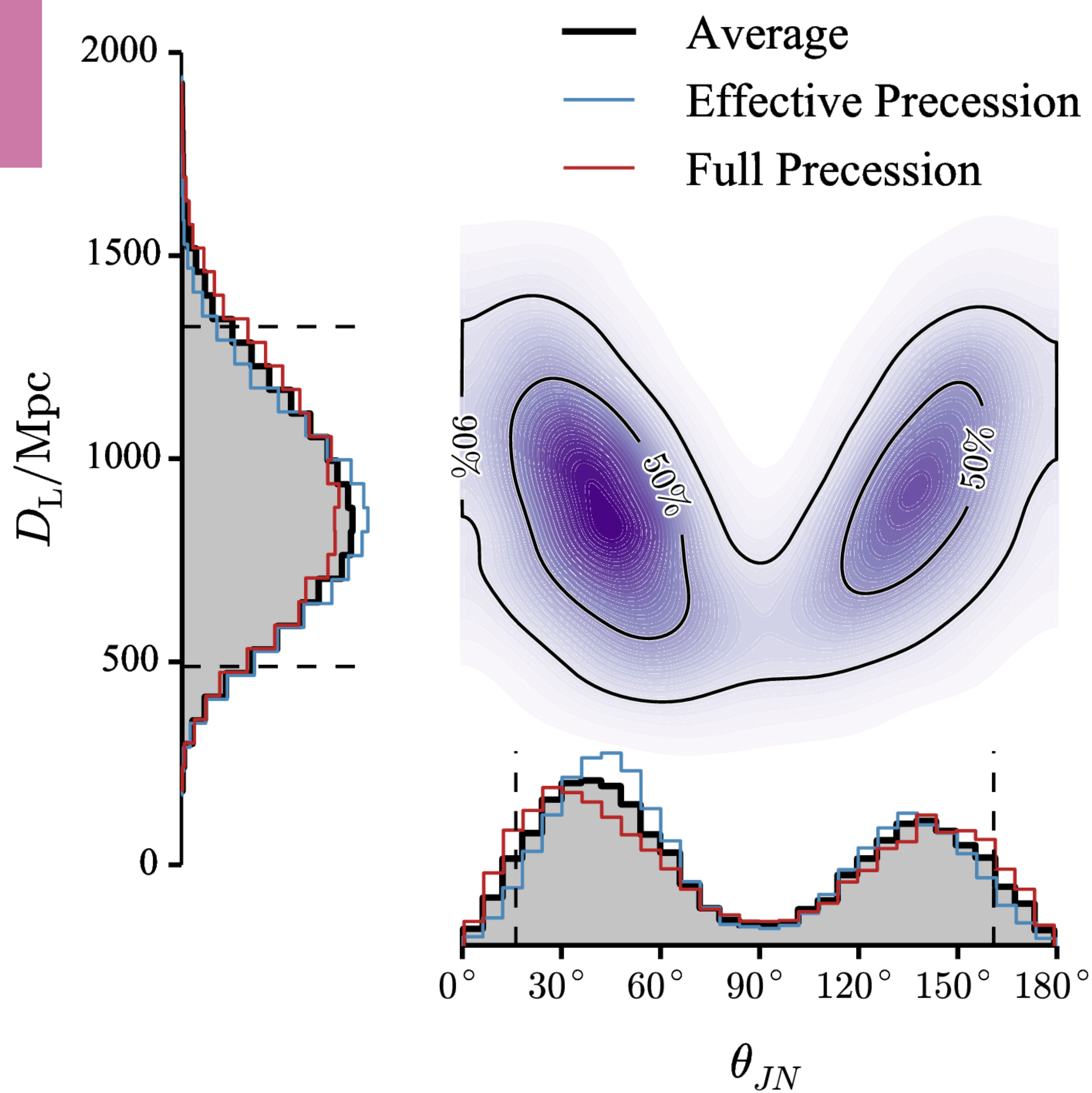

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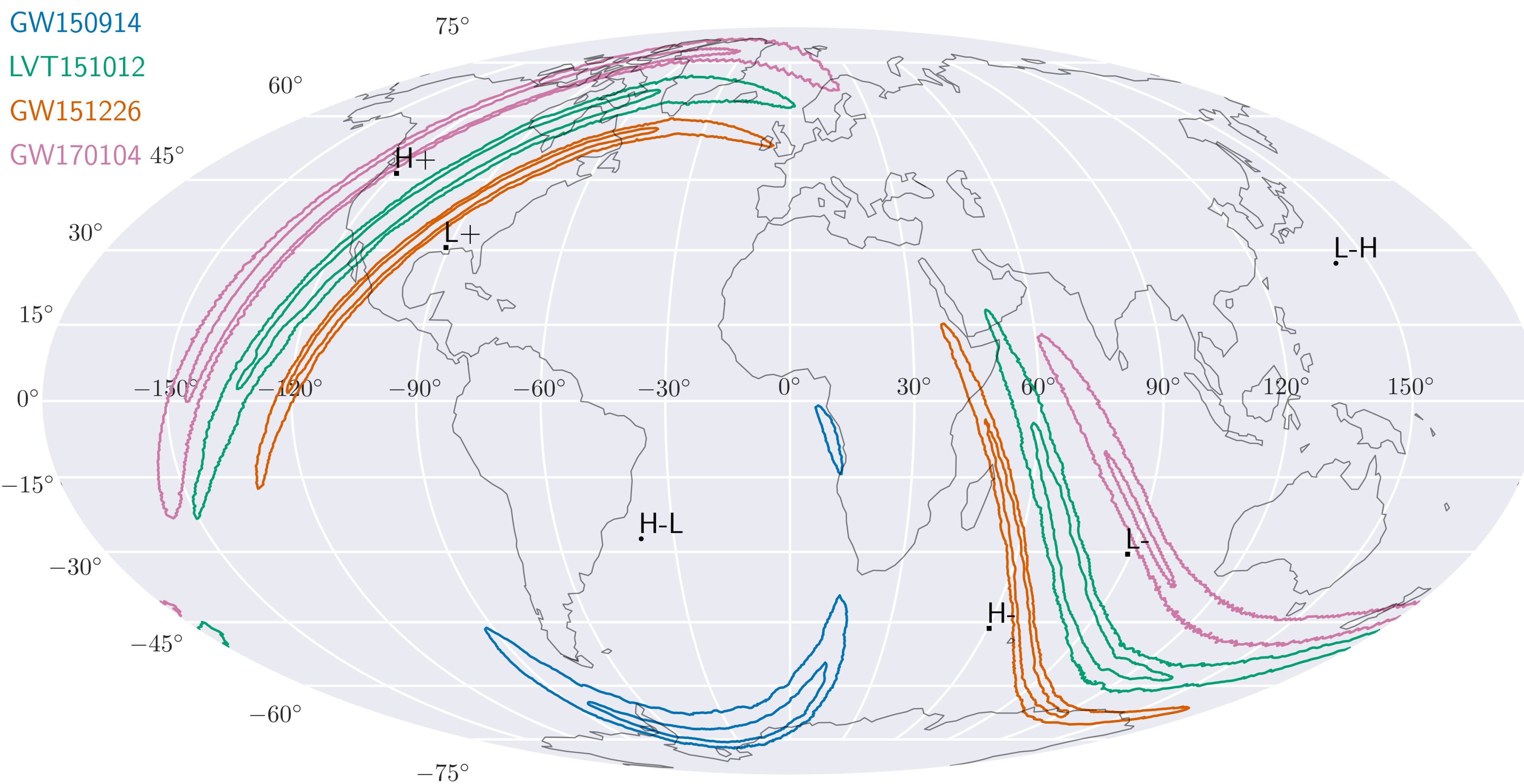

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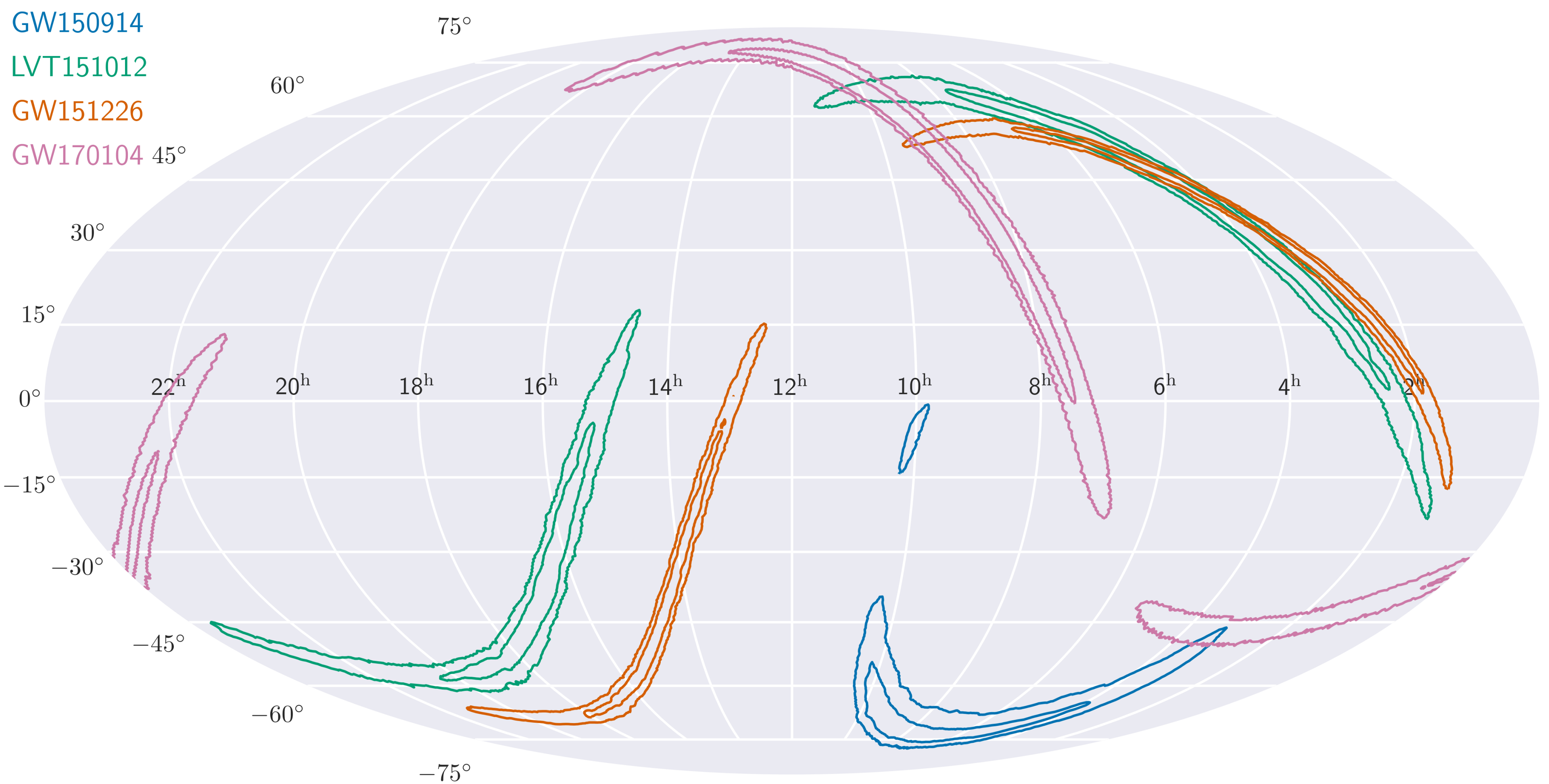

LVC arXiv:1606.04856
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Data release | 2015, HL | 2016, HLV | new 2015, recoloured | Show/hide columns

The First Two Years of Electromagnetic Follow-Up with Advanced LIGO and Virgo

Singer et al. 2014 arXiv:1404.5623

Berry et al. 2015 arXiv:1411.6934

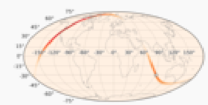
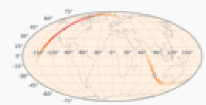
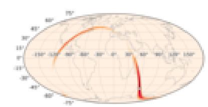
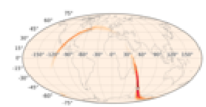
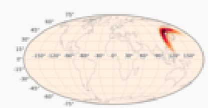
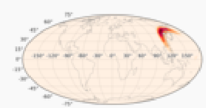
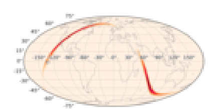
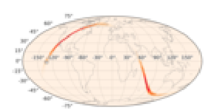
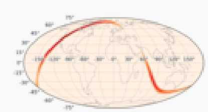
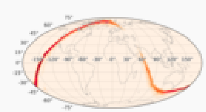
www.ligo.org/scientists/first2years/

asd.gsfc.nasa.gov/Leo.Singer/going-the-distance/

This web page pr
additional online r
related to the pap
Two Years of Ele
Follow-Up with Ad
and Virgo" and th
paper "Parameter
for Binary Neutron
Coalescences wit

Data release | 2015, HL | 2016, HLV | new 2015, recoloured | Show/hide columns

Catalog of simulated events and sky maps for two-detector, HL, 2015 configuration. This is the same configuration as the 2015 tab, except that the simulated detector noise is data from initial LIGO's sixth science run, recoloured (filtered) to have the same PSD as the early Advanced LIGO configuration. See also ASCII tables of simulated signals, detections, and parameter-estimation accuracies in Machine Readable Table format.

event ID	sim ID	network	SNR			BAYESTAR			LALINFERENCE_NEST			sky maps	
			net	H	L	50%	90%	searched	50%	90%	searched	BAYESTAR	LALINFERENCE_NEST
4532	899	HL	13.9	10.1	9.5	180	750	190	170	790	150		
4572	1243	HL	13.2	10.0	8.7	230	830	45	200	920	33		
4618	1768	HL	10.8	8.0	7.3	160	540	220	130	440	280		
4647	1964	HL	12.4	8.6	9.0	260	890	1200	190	780	780		
4711	2704	HL	10.7	8.0	7.1	370	1200	300	450	1600	520		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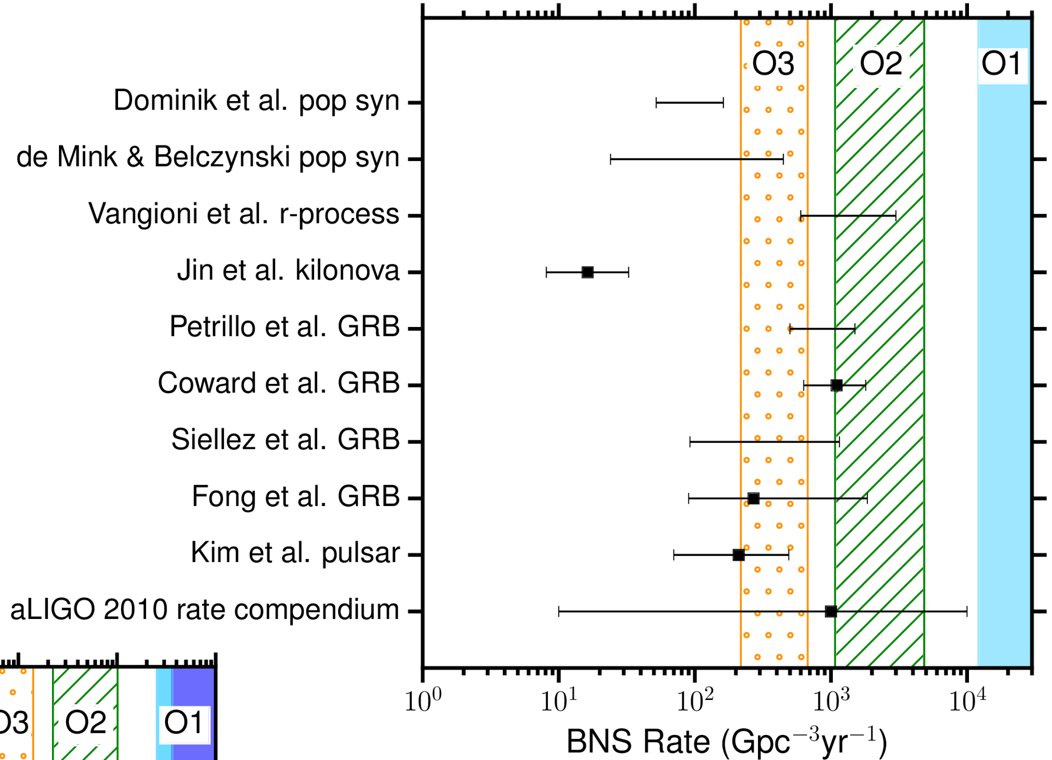

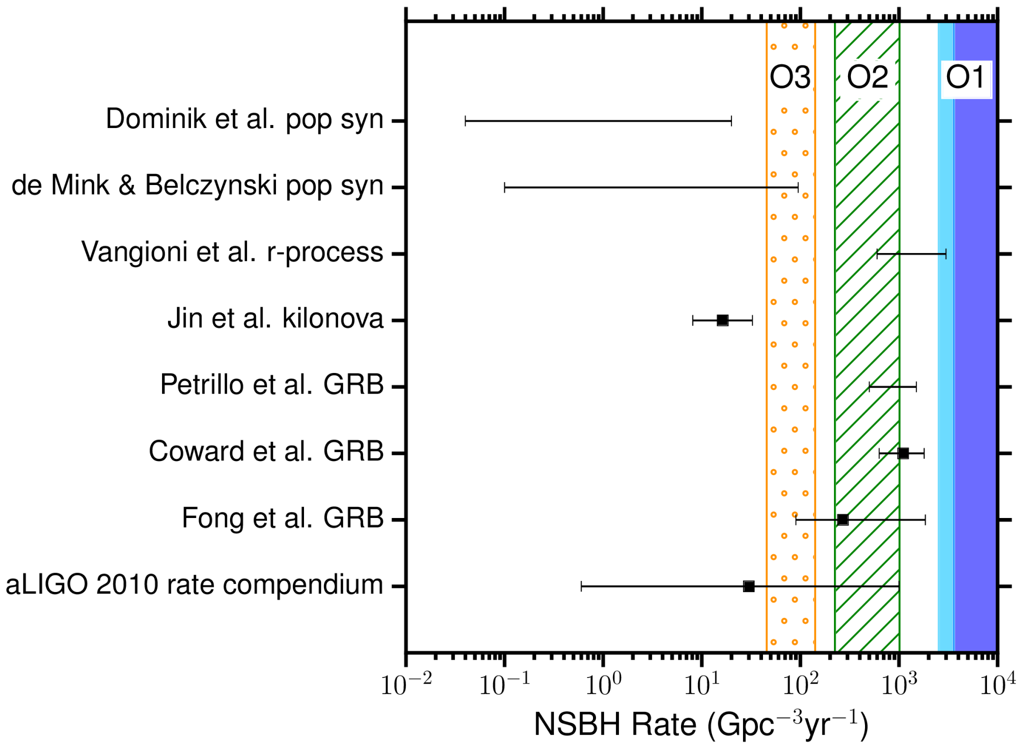

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