

Dark Matter models beyond the WIMP paradigm

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JGU Mainz

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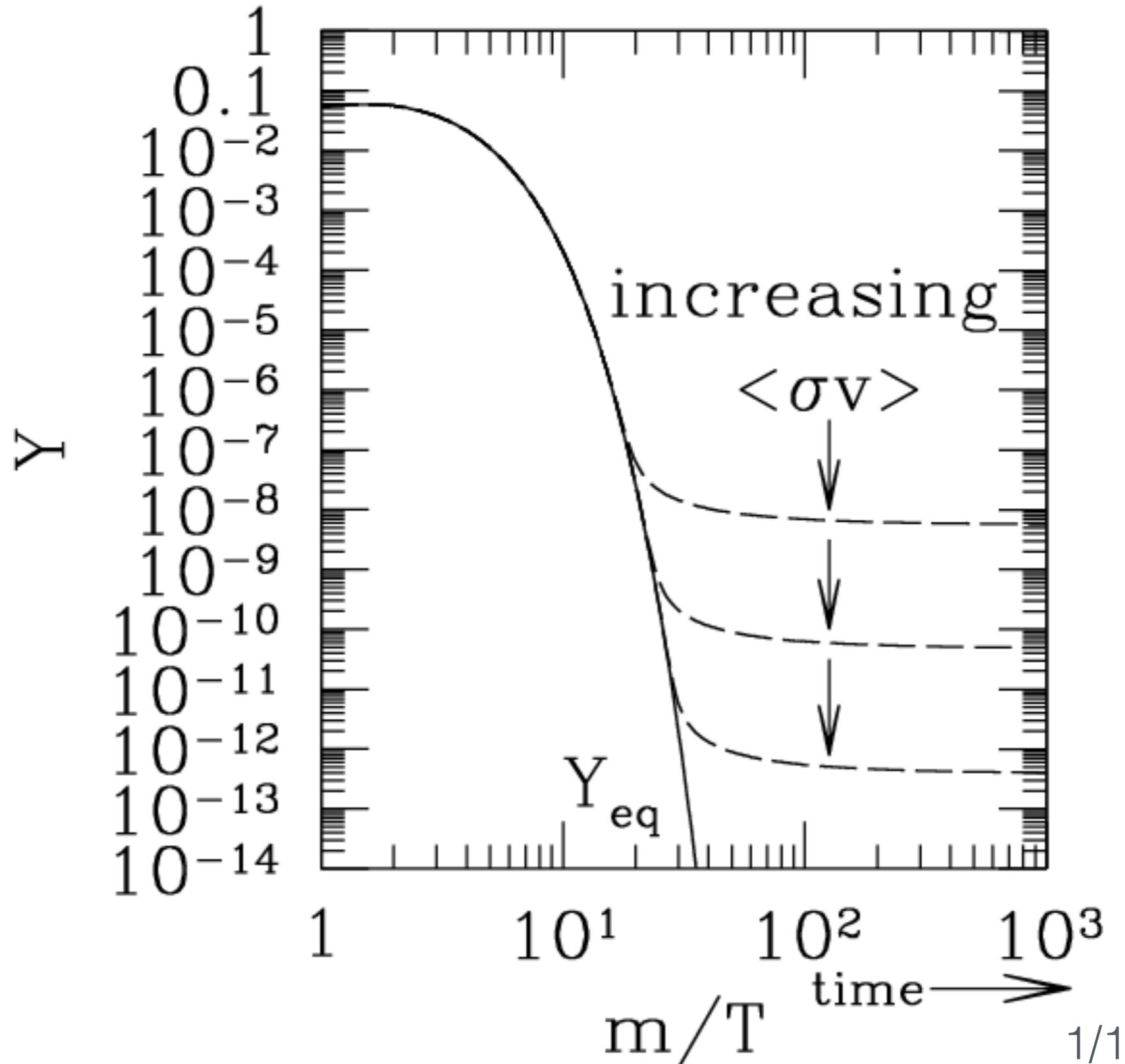
- The thermal WIMP paradigm
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- Vev Flip Flop: Dark Matter Decay between Weak Scale Phase Transitions

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- The thermal WIMP paradigm

1009.3690 — Gondolo & Gelmini

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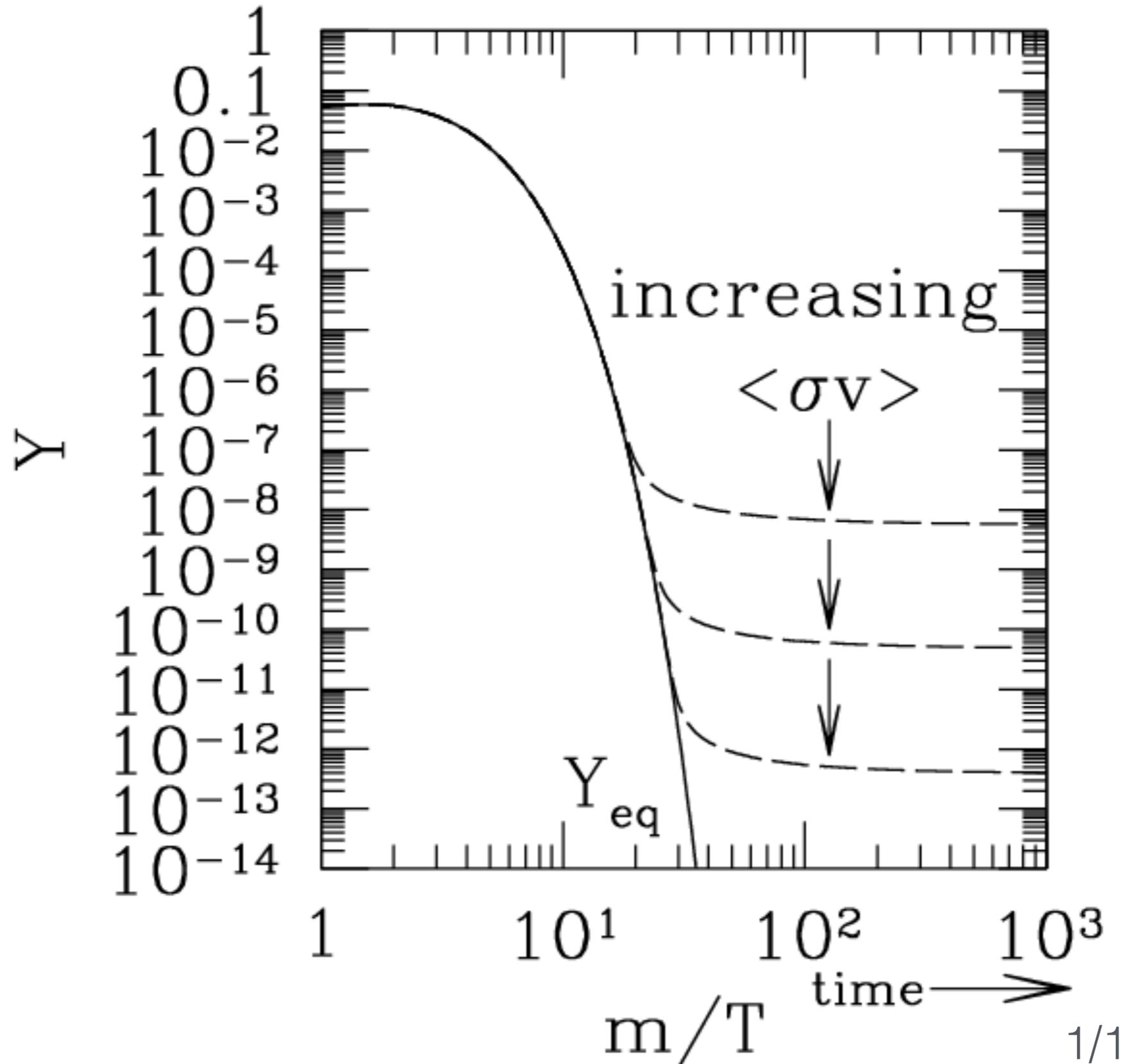


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- 2) Expansion rate greater than annihilation rate leads to freeze-out.

$$H \gtrsim \Gamma$$



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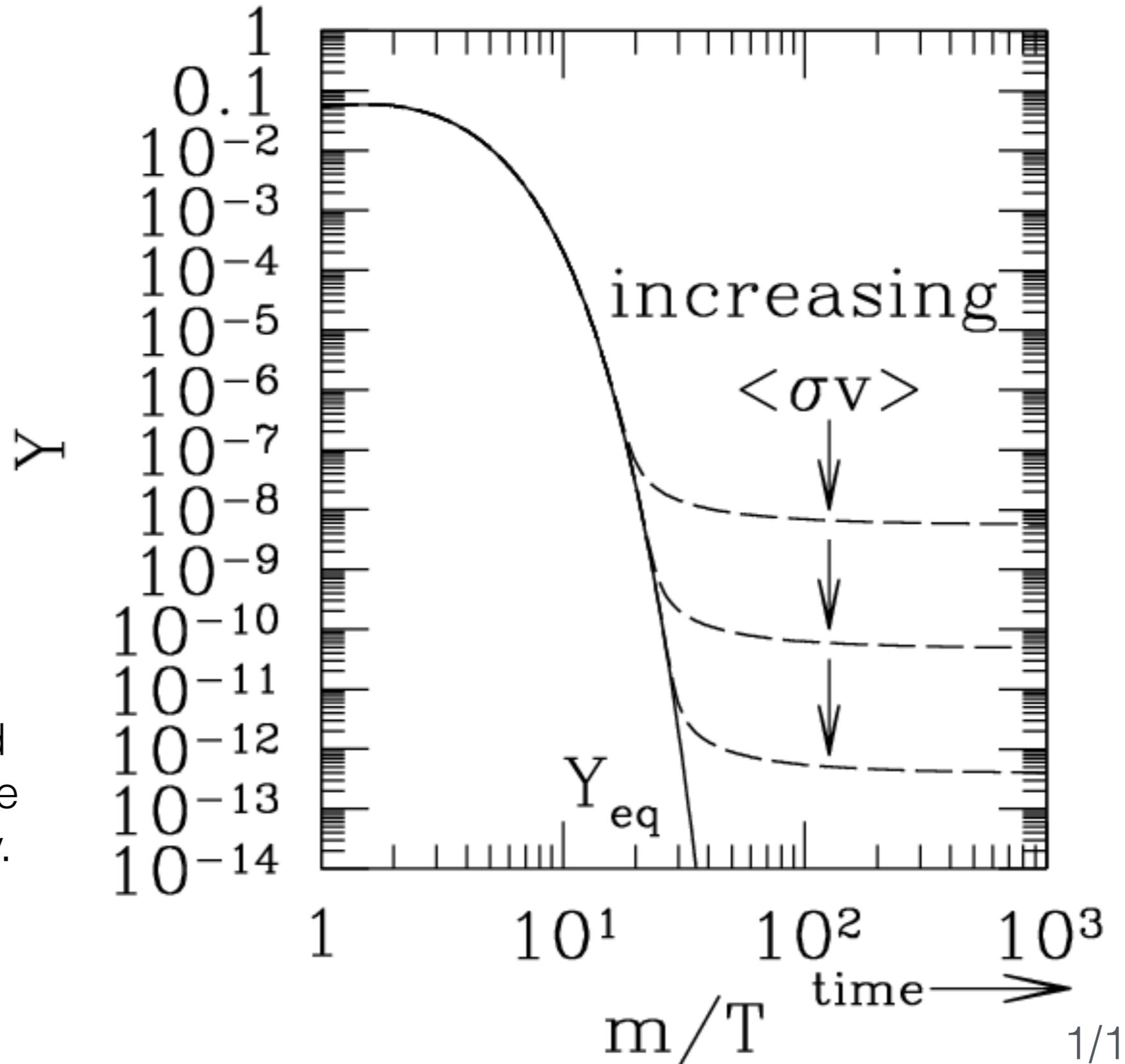
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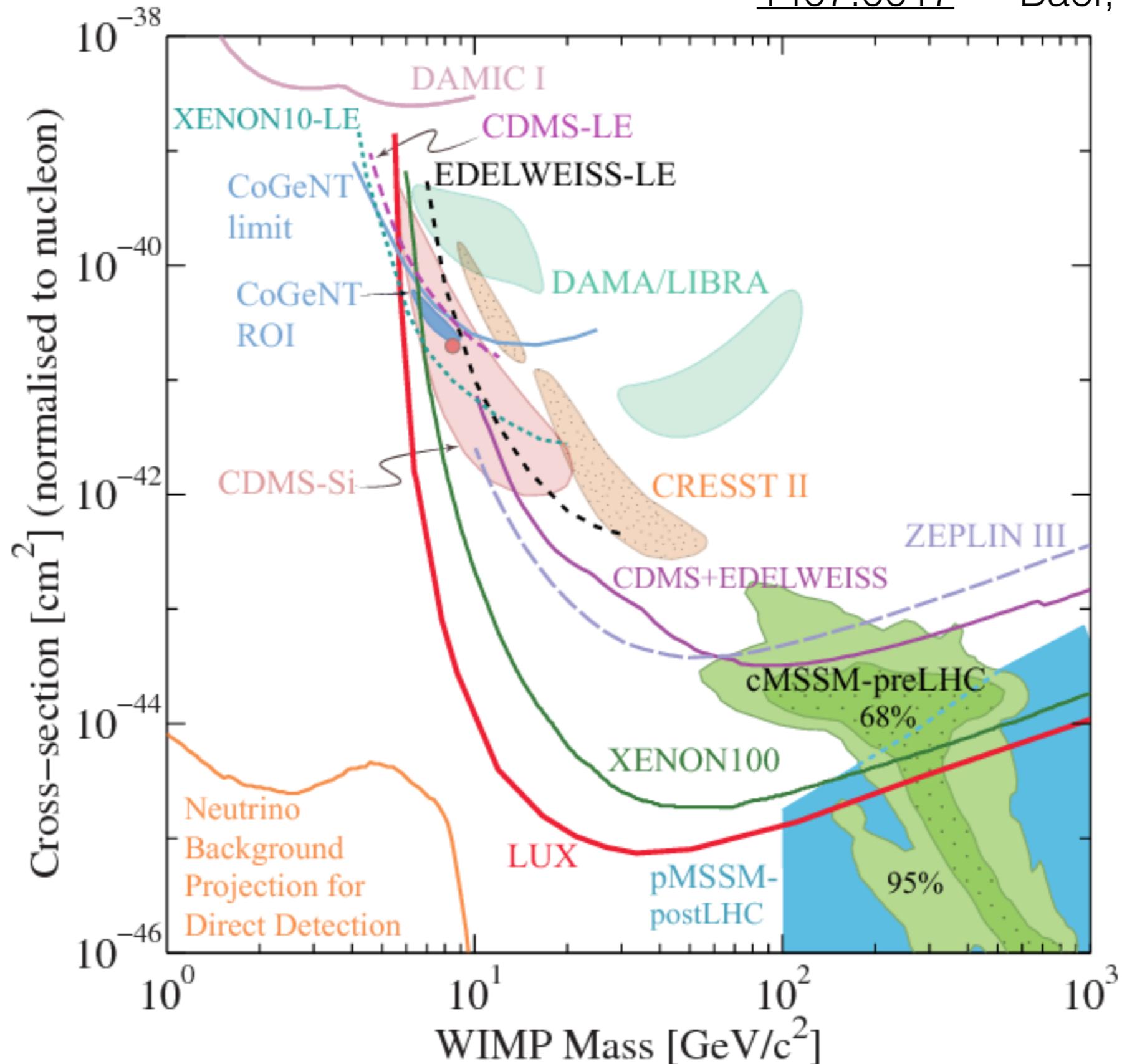
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3) Weak scale masses and order one couplings give ~ observed relic density.



- The thermal WIMP paradigm

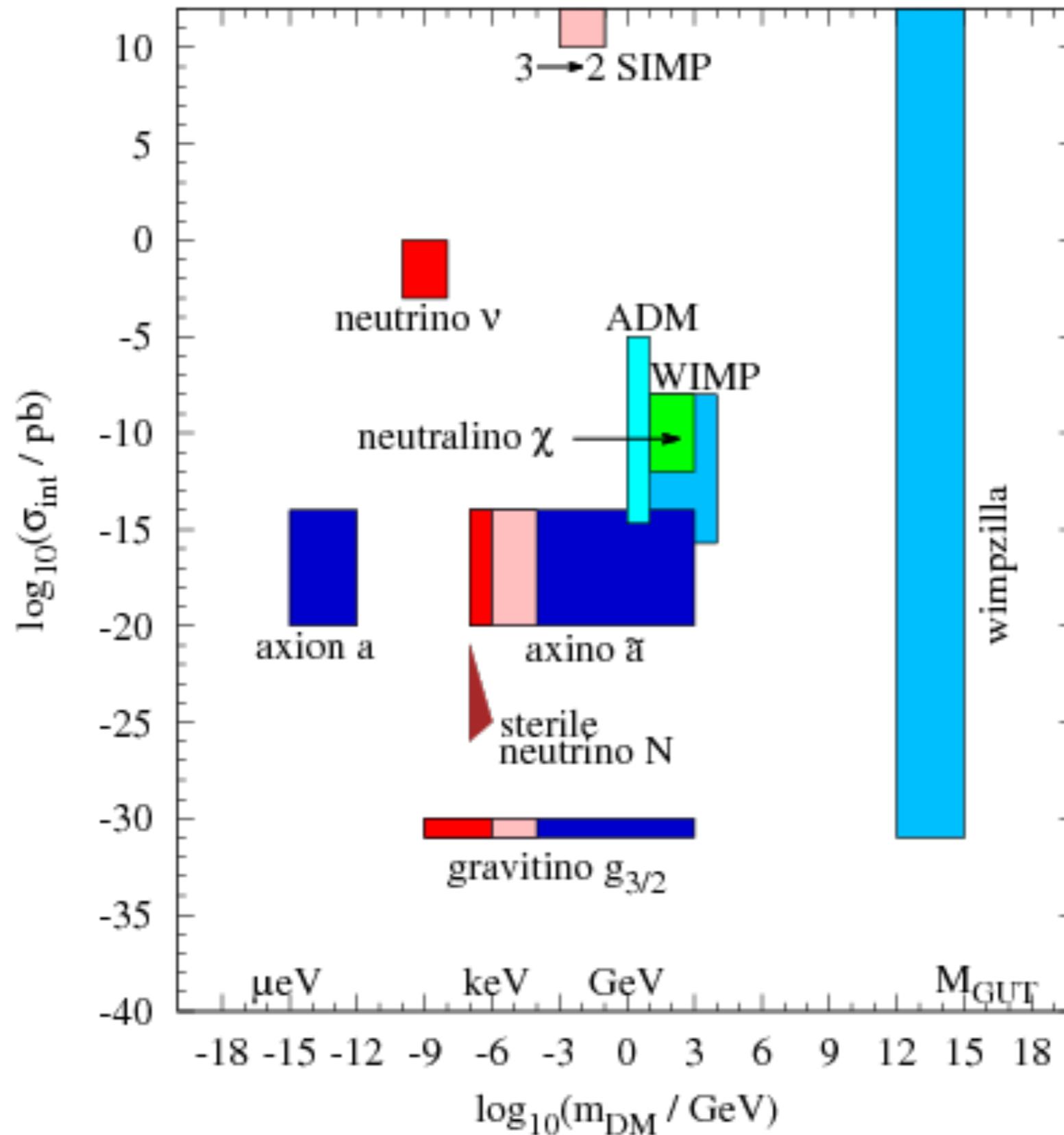
1407.0017 — Baer, Howard et al.



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Based on [1608.07578](#), MJB & Joachim Kopp

- Vev Flip Flop: Model

Field	Spin	SM	\mathbb{Z}_3	mass scale
χ	$\frac{1}{2}$	$(1, 1, 0)$	$\chi \rightarrow e^{2\pi i/3} \chi$	TeV
S_3	0	$(1, 3, 0)$	$S_3 \rightarrow e^{2\pi i/3} S_3$	100 GeV
Ψ_3	$\frac{1}{2}$	$(1, 3, 0)$	$\Psi_3 \rightarrow e^{-2\pi i/3} \Psi_3$	TeV
Ψ'_3	$\frac{1}{2}$	$(1, 3, 0)$	$\Psi'_3 \rightarrow e^{-2\pi i/3} \Psi'_3$	TeV

$$\begin{aligned} \mathcal{L}_{\text{Yuk}} = & y_\chi S_3^\dagger \bar{\chi} \Psi_3 + y'_\chi S_3^\dagger \bar{\chi} \Psi'_3 \\ & + y_\Psi \epsilon^{ijk} S_3^i \overline{\Psi_3^j} (\Psi_3'^k)^c + h.c. \end{aligned}$$

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$$y_\chi \sim y'_\chi \sim 10^{-7} \qquad y_\Psi \sim 1$$

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$$\begin{aligned}
 V(H, S_3) = & -\mu^2 H^\dagger H + \lambda(H^\dagger H)^2 \\
 & -\mu_S^2 S_3^\dagger S_3 + \lambda_S(S_3^\dagger S_3)^2 + \lambda_3(S_3^\dagger T^a S_3)^\dagger S_3^\dagger T^a S_3 \\
 & + \alpha H^\dagger H S_3^\dagger S_3 + \beta H^\dagger \tau^a H S_3^\dagger T^a S_3
 \end{aligned}$$

- Vev Flip Flop: Effective Potential

$$\begin{aligned} V^{\text{eff}} &= V^{\text{tree}} + V^{\text{1-loop}} \\ &= V^{\text{tree}} + \sum_i \frac{n_i T}{2} \sum_{n=-\infty}^{+\infty} \int \frac{d^3 \vec{k}}{(2\pi)^3} \log \left[\vec{k}^2 + \omega_n^2 + m_i^2(h, S) \right] \end{aligned}$$

$$V^{\text{eff}} = V^{\text{tree}} + V^{\text{CW}} + V^{\text{T}} + V^{\text{daisy}}$$

- Vev Flip Flop: Effective Potential

$$V^{\text{eff}} = V^{\text{tree}} + \underline{V^{\text{CW}}} + V^T + V^{\text{daisy}}$$

$$V^{\text{CW}}(h, S) = \sum_i \frac{n_i}{64\pi^2} m_i^4(h, S) \left[\log \frac{m_i^2(h, S)}{\Lambda^2} - \frac{3}{2} \right]$$

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$$m_A^2(h, S) = -\mu_S^2 + \frac{1}{2}\alpha h^2 + \lambda_S S^2$$

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$$m_\gamma^2(h, S) = 0$$

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$$i \in \{A, G^0, W^\pm, Z, \gamma, t, h - S, S^+ - S^- - G^+\}$$

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Coleman & Weinberg, 1973

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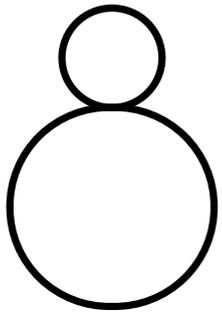
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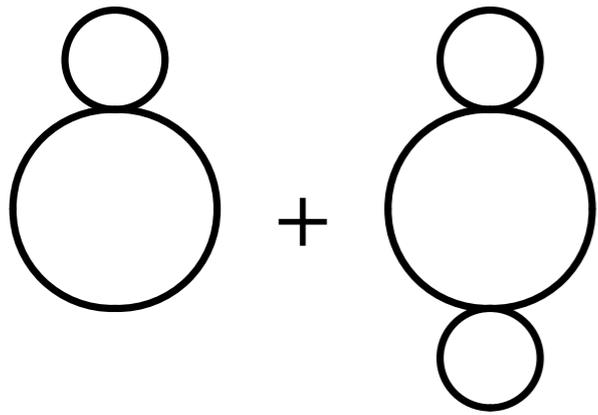
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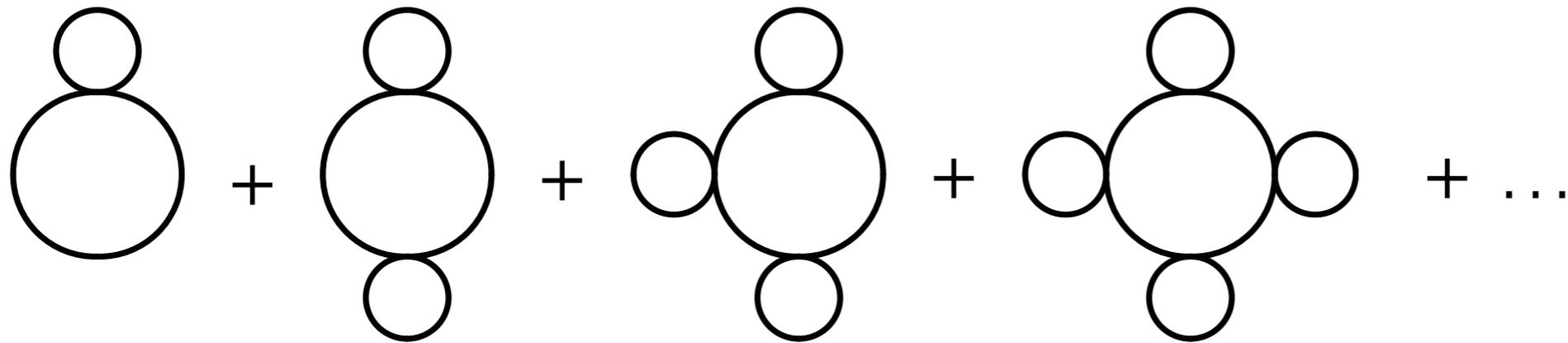
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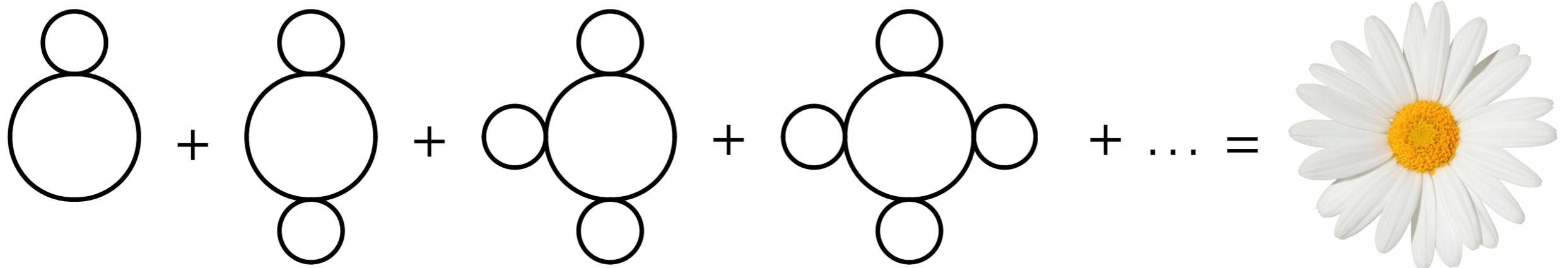
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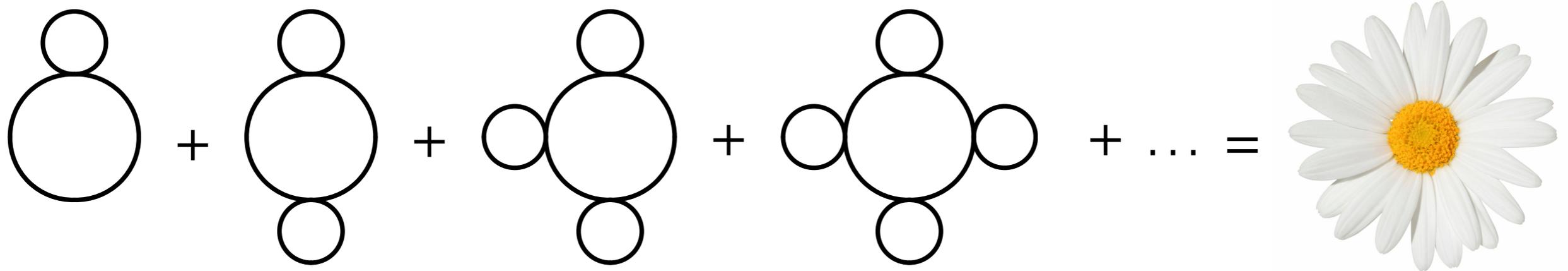
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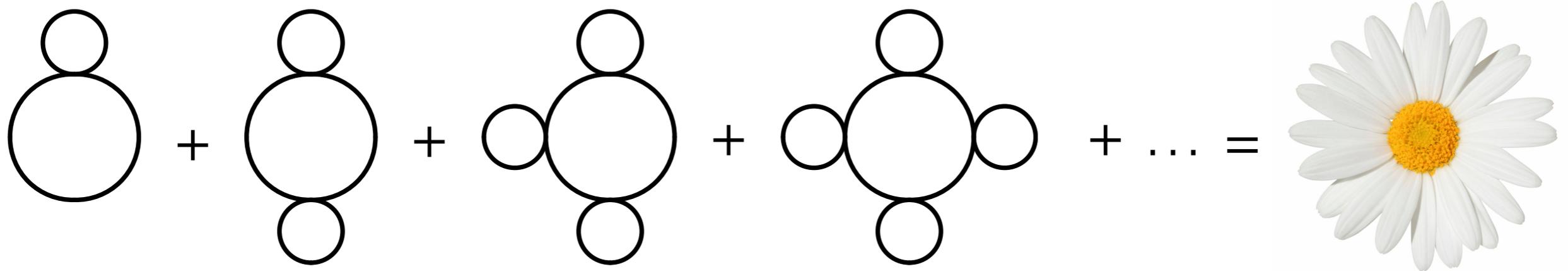
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$$V^{\text{daisy}} = -\frac{T}{12\pi} \sum_i n_i \left([m_i^2(h, S) + \Pi_i(T)]^{\frac{3}{2}} - [m_i^2(h, S)]^{\frac{3}{2}} \right)$$

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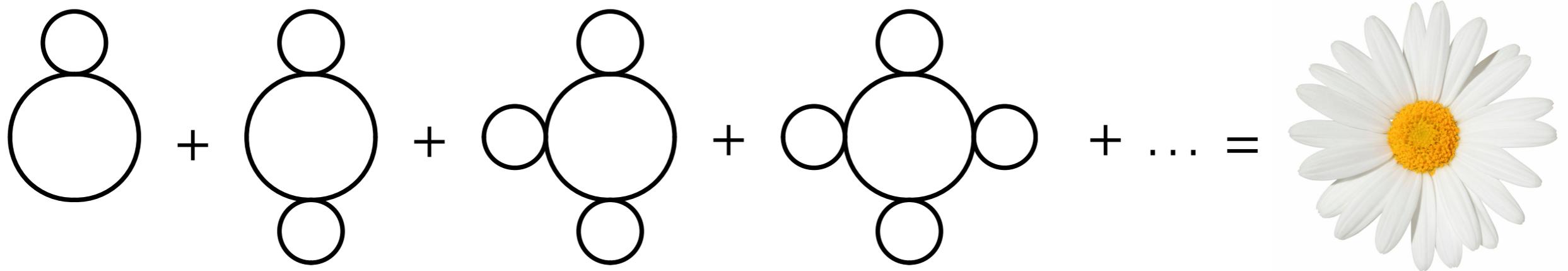


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$$\Pi_{S, A, S_R^+, S_I^+, S_R^-, S_I^-} = \frac{T^2}{6} (3g^2 + \alpha + 2\lambda_3 + 4\lambda_S)$$

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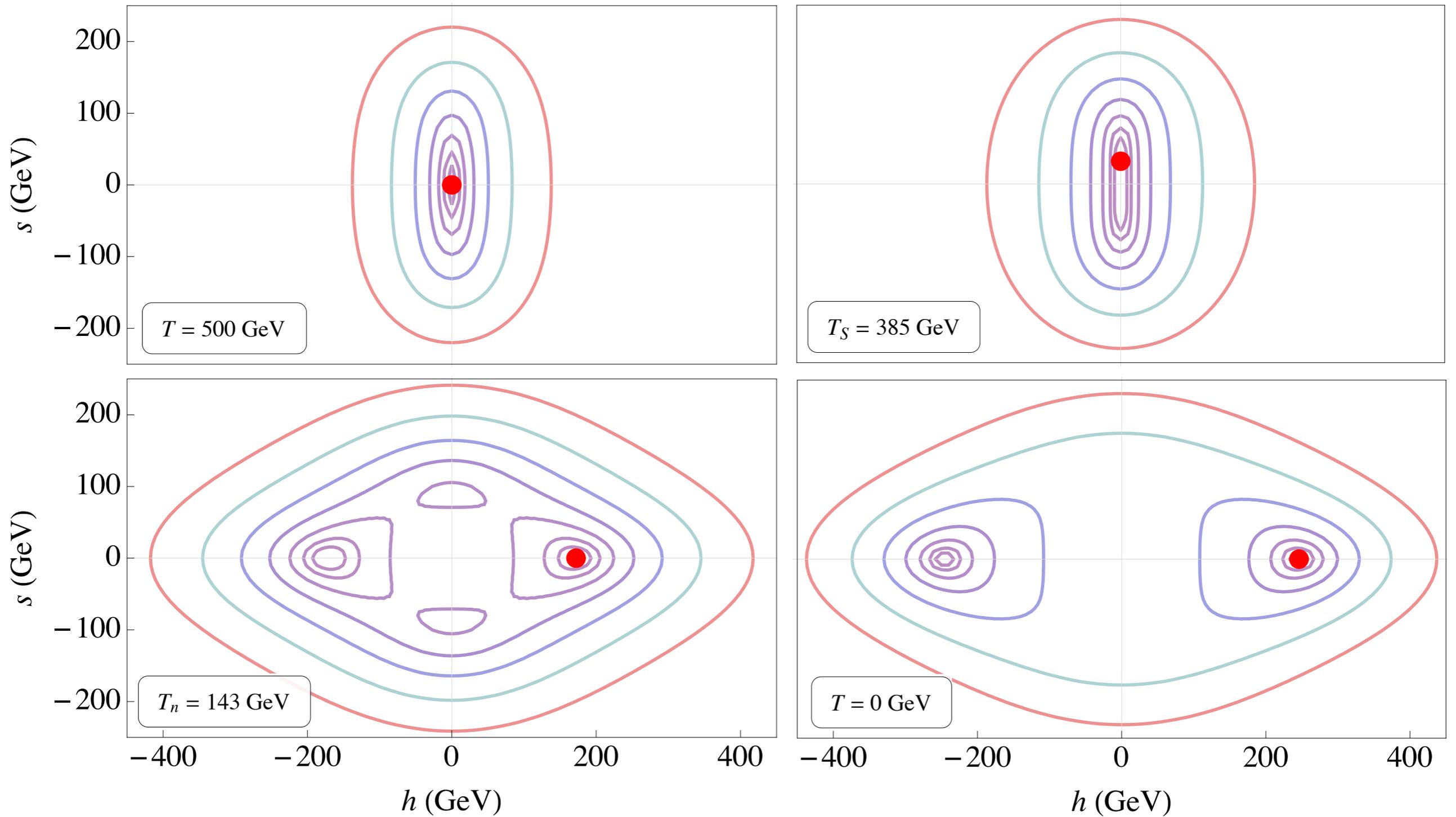
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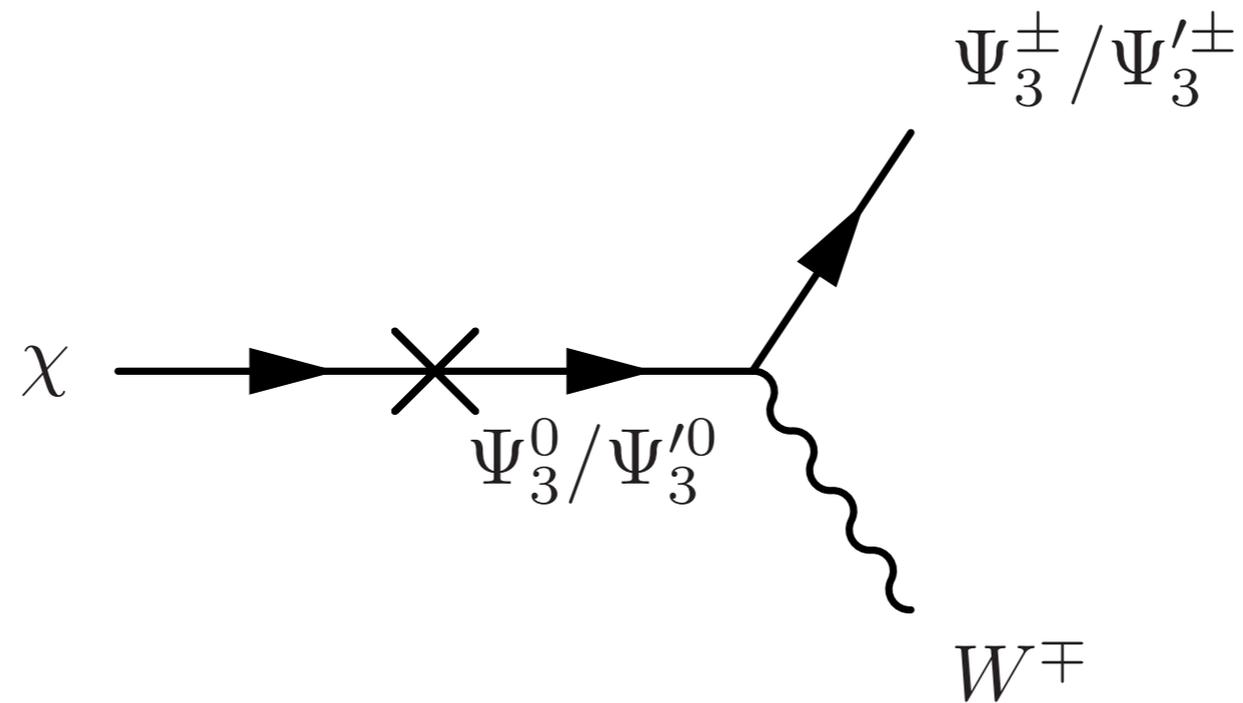
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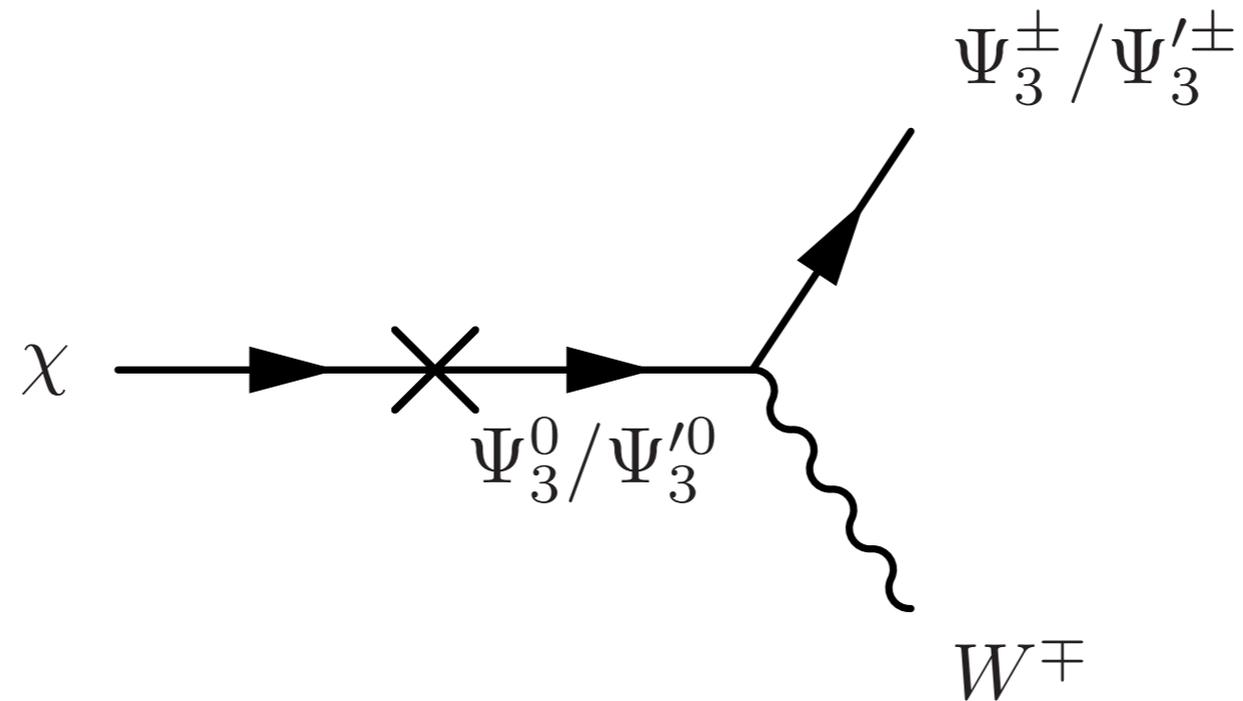


- Vev Flip-Flop: Dark Matter Decay

$$\langle S \rangle \neq 0 \implies \mathcal{L} \supset y_\chi \langle S \rangle \bar{\chi} \Psi_3 + y'_\chi \langle S \rangle \bar{\chi} \Psi'_3$$



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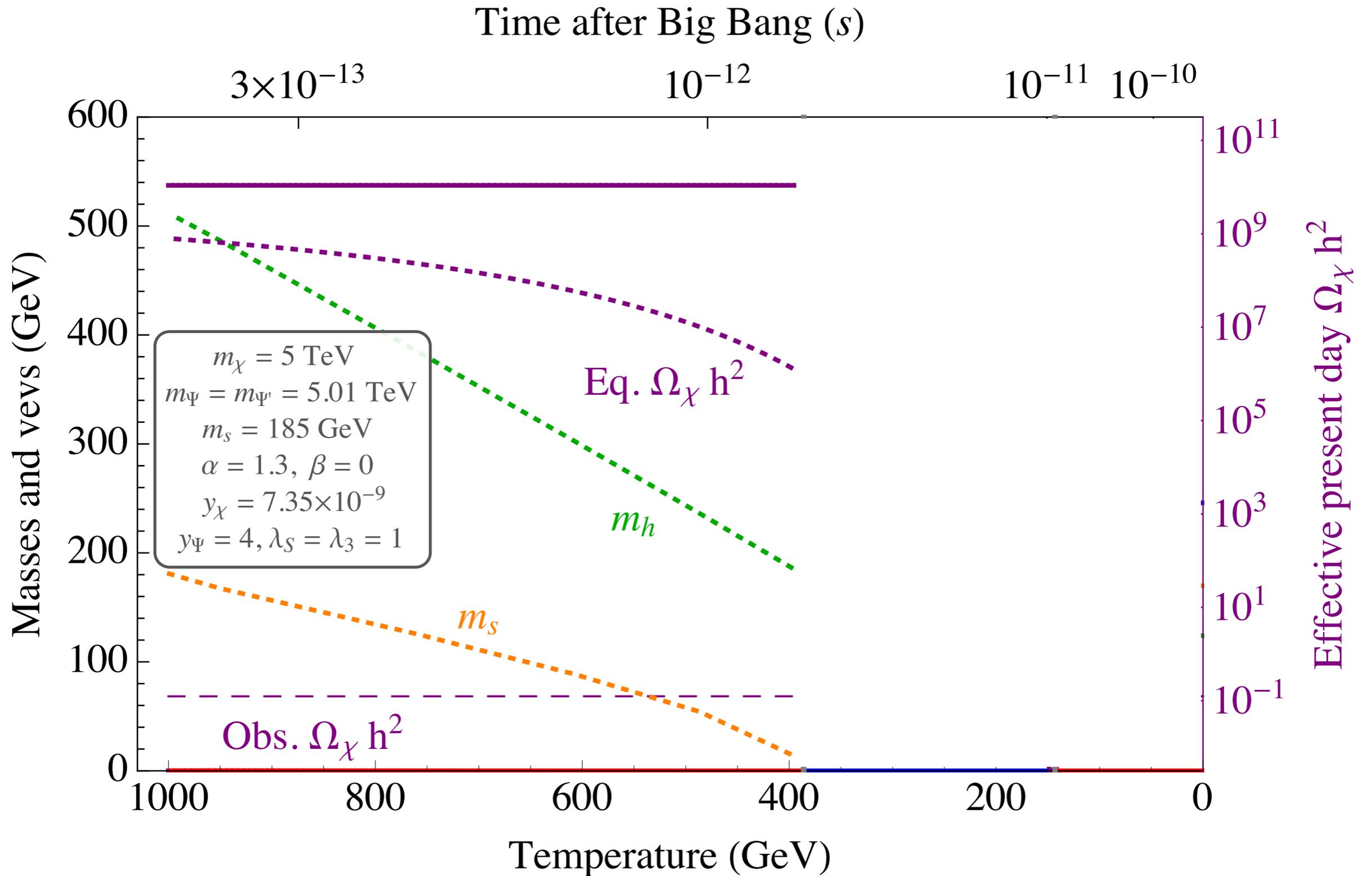


$$\dot{n}_\chi^j + 3Hn_\chi^j = -\frac{\Gamma}{\gamma^j} (n_\chi^j - n_\chi^{j,\text{eq}})$$

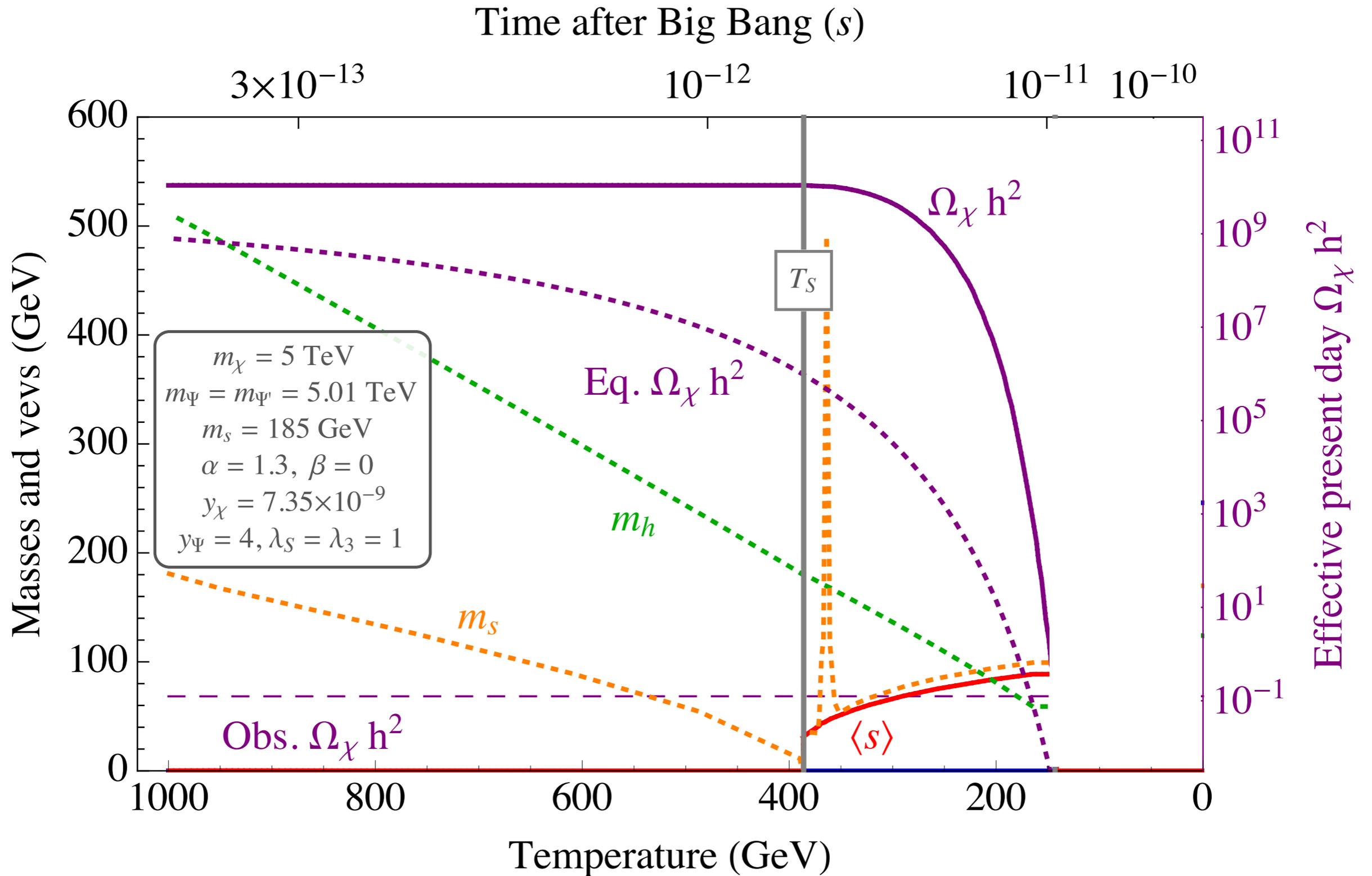
$$H^2 = \left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G_N}{3} (\rho_{\text{SM}} + \rho_\chi)$$

$$\dot{\rho}_{\text{SM}} = 4\rho_{\text{SM}}\frac{\dot{T}}{T} = m_\chi\Gamma \sum_j (n_\chi^j - n_\chi^{j,\text{eq}}) - 4H\rho_{\text{SM}}$$

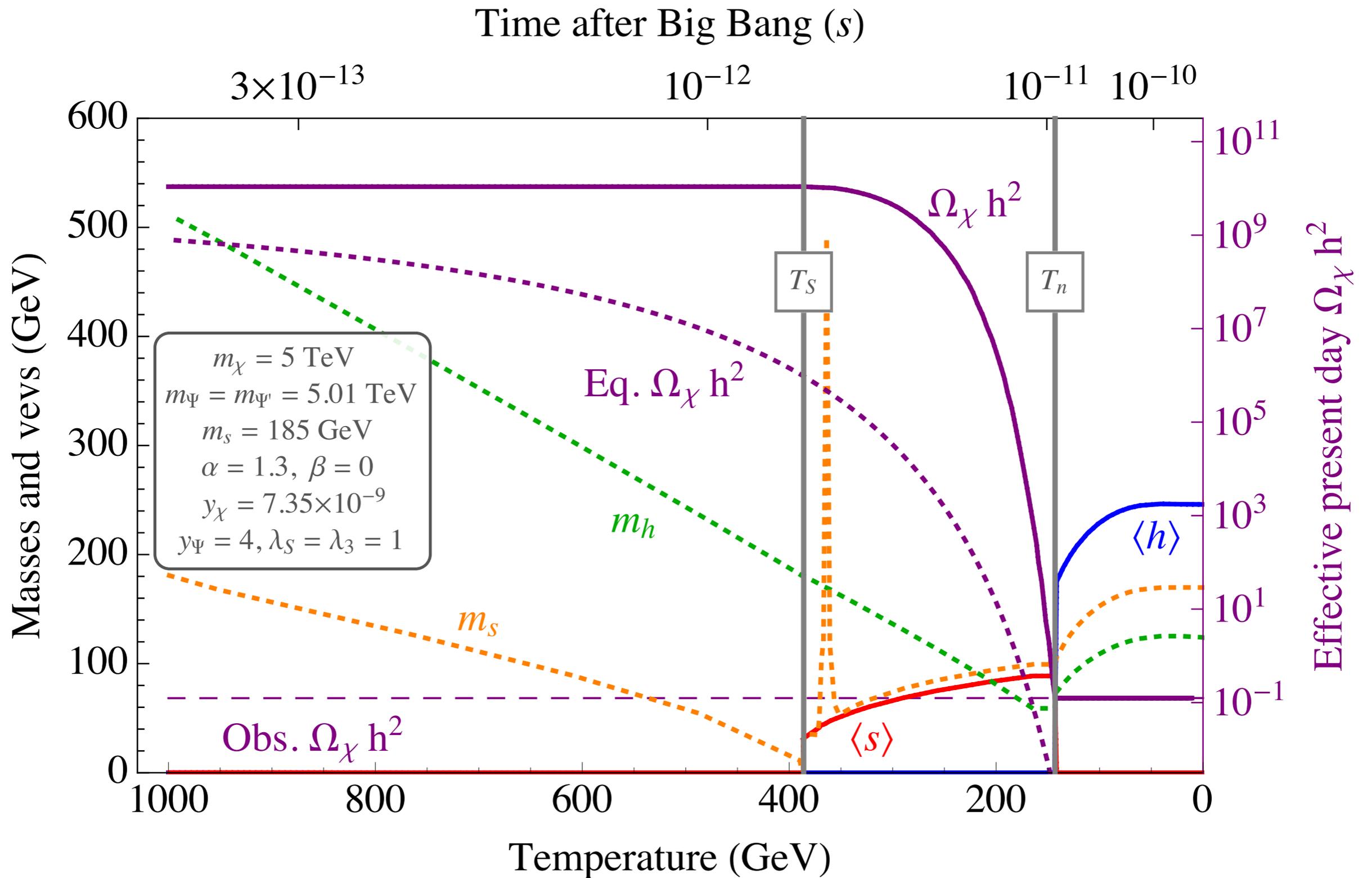
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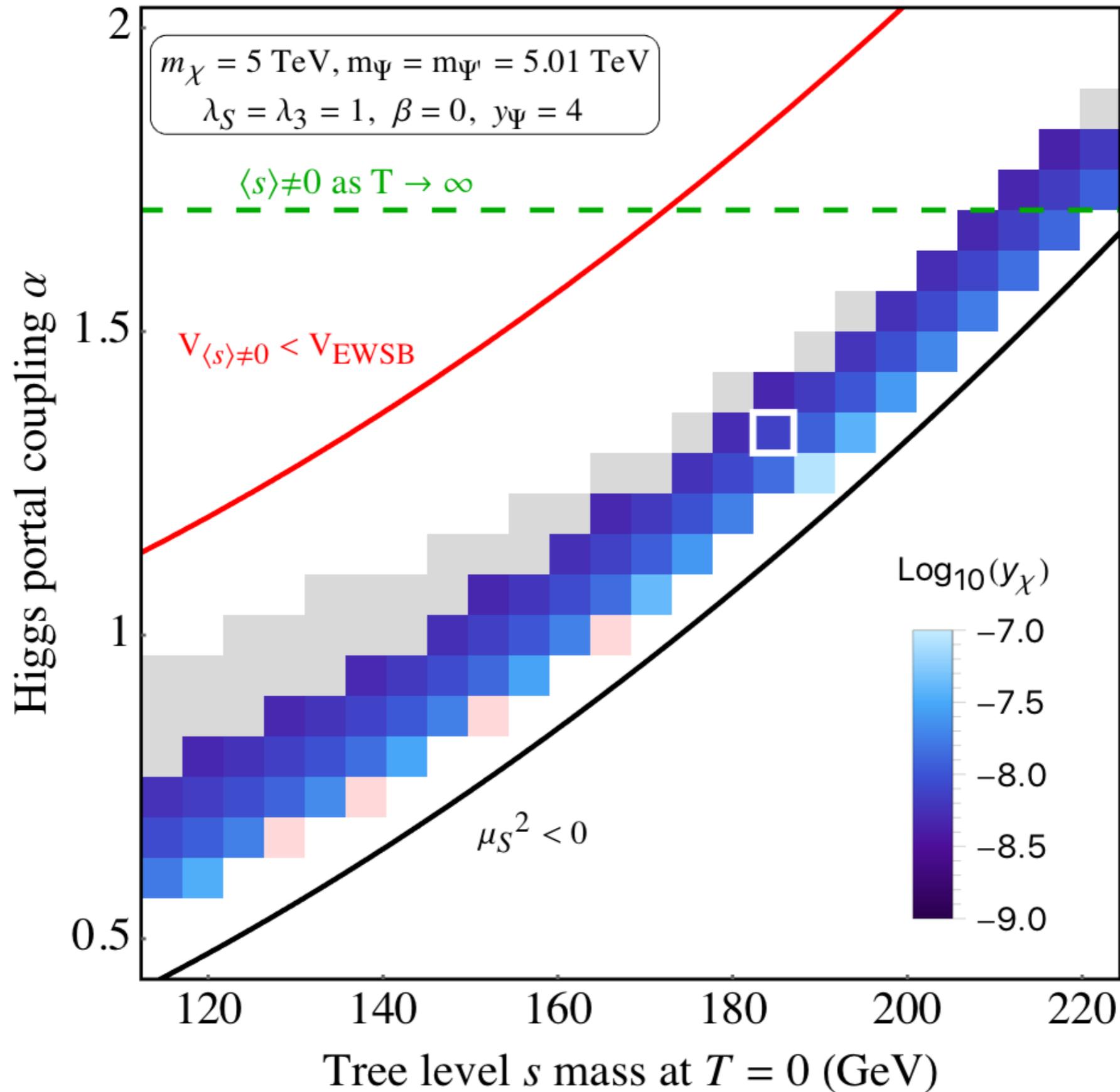
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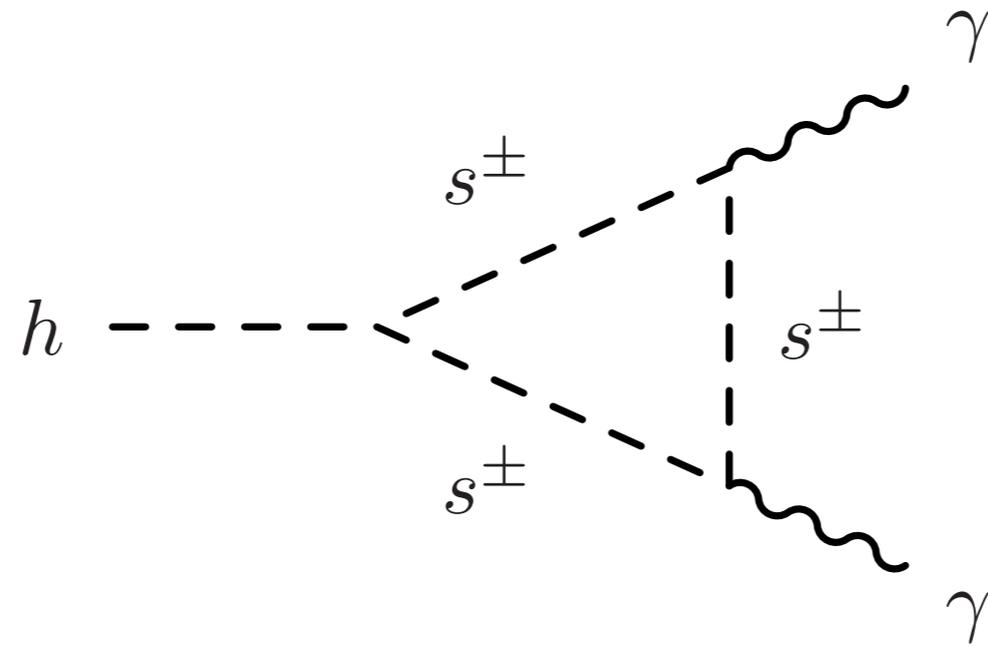
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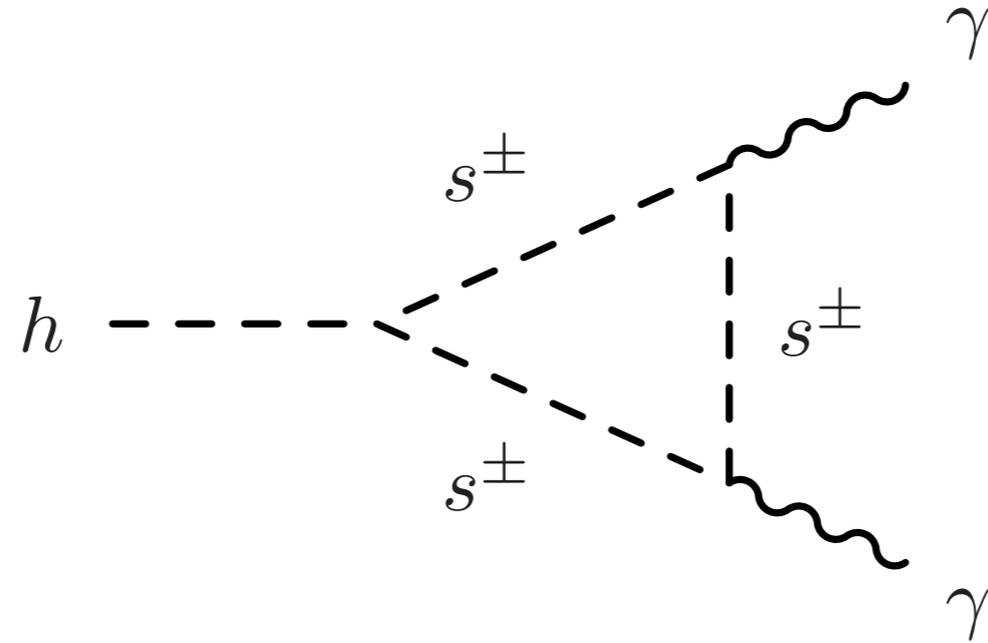
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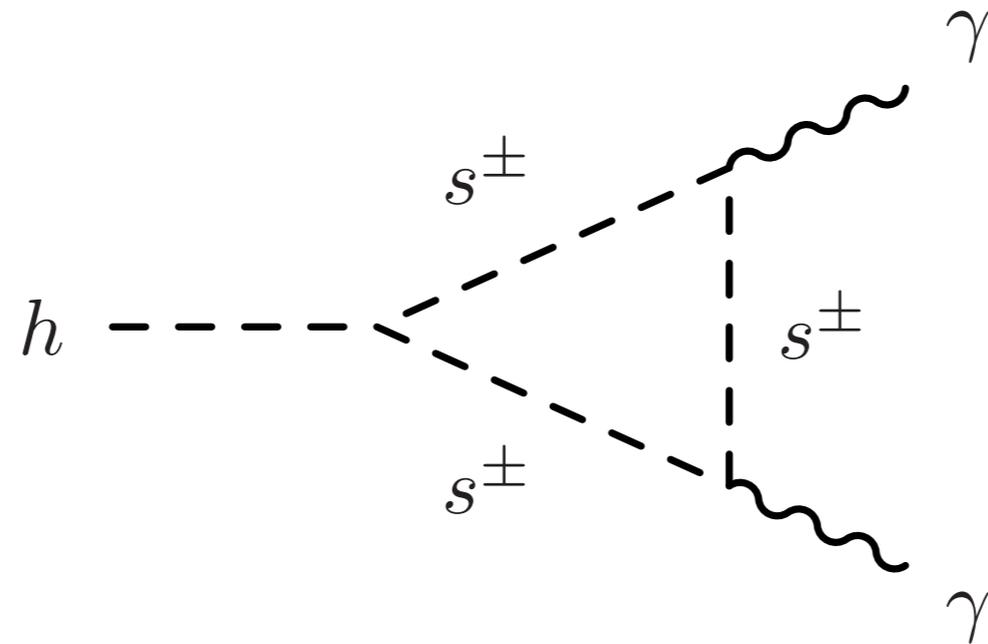


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$$|F^{\gamma\gamma}|^2 = (|F_W^{\gamma\gamma}|^2 - |F_f^{\gamma\gamma}|^2 - |F_{S_3}^{\gamma\gamma}|^2)^2$$

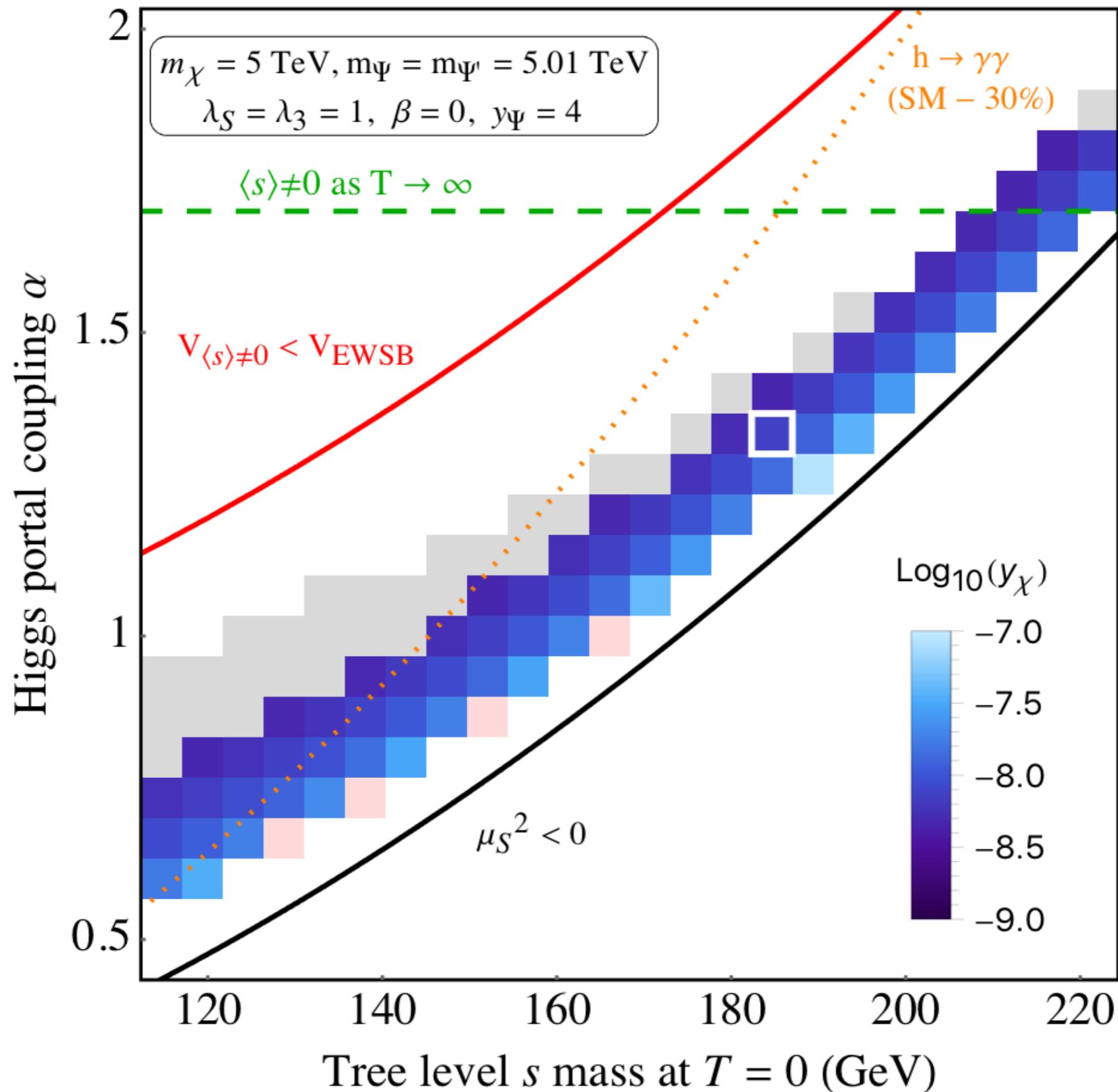
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$$|F^{\gamma\gamma}|^2 = (|F_W^{\gamma\gamma}|^2 - |F_f^{\gamma\gamma}|^2 - |F_{S_3}^{\gamma\gamma}|^2)^2$$

$$F_{S_3}^{\gamma\gamma} = \frac{\alpha_{\text{EW}}}{\pi} \alpha \left[1 - \frac{4m_{s^\pm}^2}{m_h^2} f \left(\frac{4m_{s^\pm}^2}{m_h^2} \right) \right]$$

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- Outlook:
 - More realisations,
 - LHC phenomenology,
 - gravitational wave signals,
 - baryogenesis,...