

**The beginning of all sciences
is the astonishment that things are
the way they are**

Aristoteles

The Higgs boson as Inflaton

A new view on the SM of particle physics

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Higgs inflation in a Nutshell

You know the SM hierarchy problem?

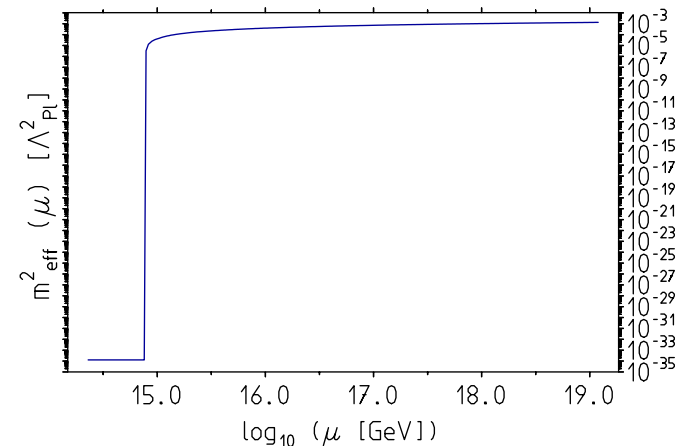
The renormalized Higgs boson mass is small (at EW scale) the bare one is huge due to radiative corrections going with the UV cutoff assumed to be given by the Planck scale $\Lambda_{\text{Pl}} \sim 10^{19}$ GeV.

$$m_{\text{Higgs, bare}}^2 = m_{\text{Higgs, ren}}^2 + \delta m^2$$
$$\delta m^2 = \frac{\Lambda_{\text{Pl}}^2}{(16\pi^2)} C(\mu)$$

● Is this a problem? Is this unnatural?

● It is a prediction of the SM!

□ At low energy we see what we see (what is to be seen): the renormalizable, renormalized SM as it describes close to all we know up to LHC energies.



- What if we go to very very high energies even to the Planck scale?
- Close below Planck scale we start to see the bare theory i.e. a SM with its bare short distance effective parameters, so in particular **a very heavy Higgs boson**, which can be moving at most very slowly, i.e.

① the potential energy

$$V(\phi) = \frac{m^2}{2}\phi^2 + \frac{\lambda}{24}\phi^4 \quad \text{is large}$$

② the kinetic energy

$$\frac{1}{2}\dot{\phi}^2 \quad \text{is small.}$$

The Higgs boson contributes to energy momentum tensor providing

pressure
energy density

$$p = \frac{1}{2}\dot{\phi}^2 - V(\phi)$$

$$\rho = \frac{1}{2}\dot{\phi}^2 + V(\phi)$$

- As we approach the Planck scale (bare theory): slow-roll condition satisfied

$$\frac{1}{2}\dot{\phi}^2 \ll V(\phi) \text{ then } \longrightarrow p \approx -V(\phi) ; \rho \approx +V(\phi) \longrightarrow p = -\rho$$

$\rho = \rho_\Lambda$ **DARK ENERGY!** no other system exhibits such strange equation of state!

- The SM Higgs boson in the early universe provides a huge dark energy!

□ What does the huge DE do? Provides **anti-gravity** inflating the universe!

Friedman equation: $\frac{da}{a} = H(t) dt \rightarrow a(t) = \exp Ht$ exponential growth of the radius $a(t)$ of the universe. $H(t)$ the Hubble constant $H \propto \sqrt{V(\phi)}$. Inflation stops quite quickly as the field decays exponentially. Field equation: $\ddot{\phi} + 3H\dot{\phi} \simeq -V'(\phi)$, for $V(\phi) \approx \frac{m^2}{2} \phi^2$ **harmonic oscillator with friction** \Rightarrow Gaussian inflation (**Planck 2013**)

- the Higgs boson is the inflaton!

- Inflation tunes the total energy density to be that of a **flat space**, which has a particular value $\rho_{\text{crit}} = \mu_{\text{crit}}^4$ with $\mu_{\text{crit}} = 0.00216 \text{ eV}$!

$\rho_{\Lambda} = \mu_{\Lambda}^4$: $\mu_{0,\Lambda} = 0.002 \text{ eV}$ today \rightarrow approaching $\mu_{\infty,\Lambda} = 0.00216 \text{ eV}$ with time

i.e. the large **cosmological constant** gets tamed by inflation to be **part of the critical flat space density**. No cosmological constant problem either?

- Note: inflation is proven to have happened by observation!

Cosmic Microwave Background (CMB) radiation tells it ✓

- Inflation requires the existence of a scalar field,

- ✱ The Higgs field is precisely such a field we need and within the SM it has the properties which promote it to be the inflaton.

Note: the Higgs inflaton is special: almost all properties are known or predictable!

All other inflatons put by hand: all predictions are direct consequences of the respective assumptions

SM Higgs inflation sounds pretty simple but in fact is rather subtle,
because of the high sensitivity to the
SM parameters uncertainties and SM higher order effects

Precondition: a stable Higgs vacuum and a sufficiently large Higgs field at M_{Pl} !

Cosmology, Cosmological Constant and Dark Energy

- Cosmology shaped by Einstein gravity $G_{\mu\nu} = \kappa T_{\mu\nu} +$
 - Weyl's postulate (radiation and matter (galaxies etc) on cosmological scales behave as ideal fluids)
 - Cosmological principle (isotropy of space implying homogeneity)
- ⇒ fix the form of the metric and of the energy-momentum tensor:
- 1. The **metric** (3-spaces of constant curvature $k = \pm 1, 0$)

$$ds^2 = (cdt)^2 - a^2(t) \left(\frac{dr^2}{1-kr^2} + r^2 d\Omega^2 \right)$$

where in the comoving frame $ds = c dt$ with t the **cosmic time**

- 2. The **energy-momentum tensor**

$$T^{\mu\nu} = (\rho(t) + p(t)) (t) u^\mu u^\nu - p(t) g^{\mu\nu} ; \quad u^\mu \doteq \frac{dx^\mu}{ds}$$

Need $\rho(t)$ **energy density** and $p(t)$ **pressure** to get $a(t)$ **radius of the universe**

Einstein [CC $\Lambda = 0$]: curved geometry \leftrightarrow matter [empty space \leftrightarrow flat space]

3. Special form energy-momentum tensor $p(t) = -\rho(t)$??? “Dark Energy” only

$$T^{\mu\nu} = \rho(t) g^{\mu\nu}$$

Peculiar dark energy equation of state: $w = p/\rho = -1$ no known physical system exhibits such strange behavior as anti-gravity!



WHAT IS DARK ENERGY? Well, the simple answer is that we don't know.

First introduced by Einstein as “Cosmological Constant” (CC) as part of the geometry, [where empty space appears curved,] in order to get stationary universe.

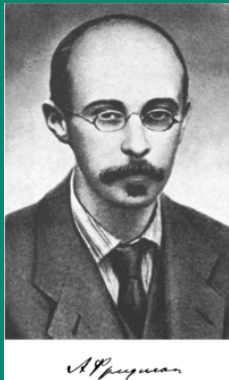
$$G_{\mu\nu} \equiv R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R - \Lambda g_{\mu\nu} = \kappa T_{\mu\nu}$$

Einstein Tensor \Leftrightarrow geometry of space-time

Gravitational interaction strength $\kappa = \frac{8\pi G_N}{3c^2}$

Energy-Momentum Tensor \Leftrightarrow deriving from the Lagrangian of the SM

Cosmological solution: universe as a fluid of galaxies \Rightarrow Friedmann-Equations:



$$\begin{aligned} 3 \frac{\dot{a}^2 + kc^2}{c^2 a^2} - \Lambda &= \kappa \rho \\ - \frac{2 \ddot{a} a + \dot{a}^2 + kc^2}{c^2 a^2} + \Lambda &= \kappa p \end{aligned}$$

$a(t)$ Robertson-Walker radius of the universe
 Λ Cosmological Constant

- universe must be expanding, **Big Bang**, and has finite age t
- Hubble's law [galaxies: $\text{velocity}_{\text{recession}} = H \text{ Distance}$], H Hubble constant
- temperature, energy density, pressure huge at begin, decreasing with time

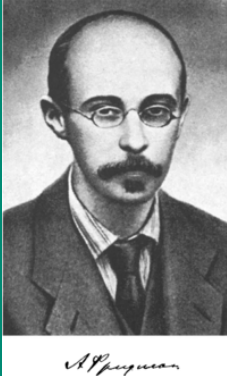
$$G_{\mu\nu} \equiv R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R = \kappa (T_{\mu\nu} + \rho_{\Lambda} g_{\mu\nu}) = \kappa T_{\mu\nu}^{\text{tot}} ; \rho_{\Lambda} = \Lambda/\kappa$$

Einstein Tensor \Leftrightarrow geometry of space-time

Gravitational interaction strength $\kappa = \frac{8\pi G_N}{3c^2}$

Energy-Momentum Tensor \Leftrightarrow deriving from the Lagrangian of the SM

Cosmological solution: universe as a fluid of galaxies \Rightarrow Friedmann-Equations:



$$3 \frac{\dot{a}^2 + kc^2}{c^2 a^2} = \kappa (\rho + \rho_{\Lambda})$$

$$- \frac{2\ddot{a}a + \dot{a}^2 + kc^2}{c^2 a^2} = \kappa (p + p_{\Lambda})$$

$a(t)$ Robertson-Walker radius of the universe
 $p_{\Lambda} = -\rho_{\Lambda}$ **Dark Energy**

- universe must be expanding, **Big Bang**, and has finite age t
- Hubble's law [galaxies: $\text{velocity}_{\text{recession}} = H \text{ Distance}$], H Hubble constant
- temperature, energy density, pressure huge at begin, decreasing with time

Problems of GRT cosmology if dark energy absent:

⇒ **Flatness problem** i.e. why $\Omega \approx 1$ (although unstable) ? CMB $\Omega_{\text{tot}} = 1.02 \pm 0.02$

⇒ **Horizon problem** finite age t of universe, finite speed of light c : $D_{\text{Hor}} = c t$
what we can see at most?

CMB sky much larger [$d_{t_{\text{CMB}}} \simeq 4 \cdot 10^7 \ell_{\text{y}}$] than causally connected patch [$D_{\text{CMB}} \simeq 4 \cdot 10^5 \ell_{\text{y}}$] at t_{CMB} (380 000 yrs), but no such spot shadow seen!

More general: what does it mean **homogeneous** or **isotropic** for causally disconnected parts of the universe? Initial value problem required initial data on space-like plane. Data on space-like plane are causally uncorrelated!

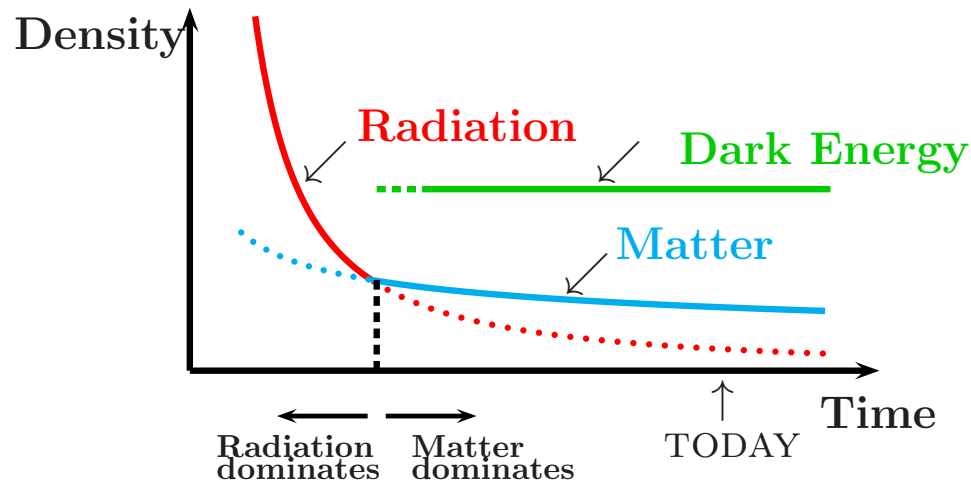
⇒ **Problem of fluctuations** magnitude, various components (dark matter, baryons, photons, neutrinos) related: **same fractional perturbations**
⇒ Planck length ℓ_{Pl} sized quantum fluctuations at Planck time?

As we will see: - $\Omega = 1$ unstable only if not sufficient dark energy!

- dark energy is provided by SM Higgs via $\kappa T_{\mu\nu}$
- no extra cosmological constant $+\Lambda g_{\mu\nu}$ supplementing $G_{\mu\nu}$
- i.e. all is standard GRT + SM (with minimal UV completion)

$$T_{\mu\nu}^{\text{tot}} = T_{\mu\nu}^{\text{SM}}$$

Sketch of the Evolution of the Universe



At Start ➡ a **Light-Flash!**

- high density
- high temperature
- ⇒ radiation dominates!

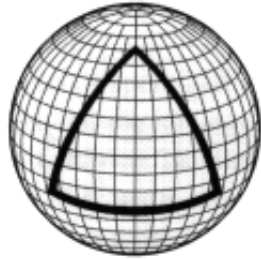
Late Times ➡ **Dark Energy** only
all other stuff dilutes into nothing!
unless universe recontracts [$k = +1$]

Forms of energy:

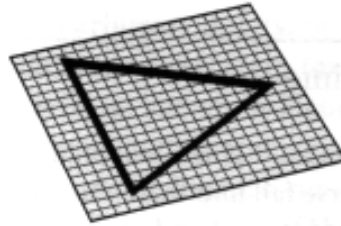
- ☐ **radiation**: photons, highly relativistic particles $p_{\text{rad}} = \rho_{\text{rad}}/3$
- ☐ **normal** and **dark matter** (non-relativistic, dilute) $p_{\text{matter}} \simeq 0$, $\rho_{\text{matter}} > 0$
- ☐ **dark energy** (cosmological constant) $p_{\text{vac}} = -\rho_{\text{vac}} < 0$

Note: Radiation $\rho_{\text{rad}} \propto 1/a(t)^4$, Matter $\rho_{\text{mat}} \propto 1/a(t)^3$, Dark Energy $\rho_{\Lambda} \propto a(t)^0$

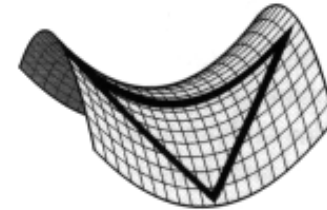
much matter
recontraction



border
case



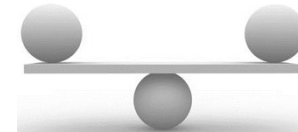
little matter
expanding forever



Curvature: closed $k = 1$ [$\Omega_0 > 1$], flat $k = 0$ [$\Omega_0 = 1$] and open $k = -1$ [$\Omega_0 < 1$]

Interesting fact: flat space geometry \Leftrightarrow specific critical density, “very unstable”

$$\rho_{0,\text{crit}} = \rho_{\text{EdS}} = \frac{3H_0^2}{8\pi G_N} = 1.878 \times 10^{-29} h^2 \text{ gr/cm}^3,$$

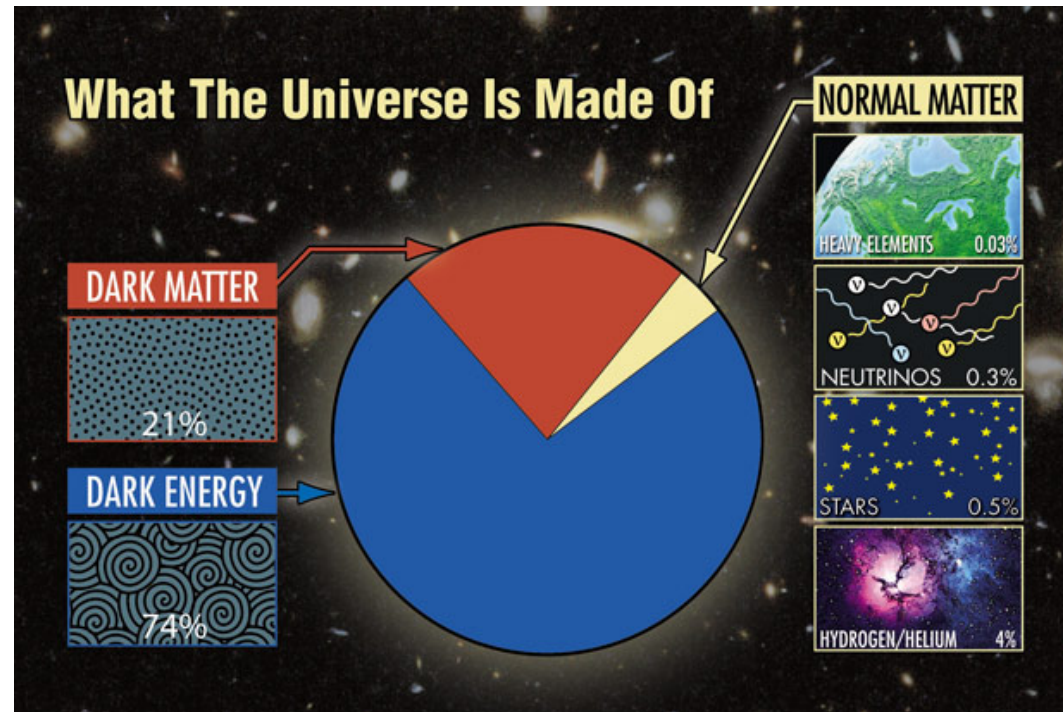


where H_0 is the present Hubble constant, and h its value in units of $100 \text{ km s}^{-1} \text{ Mpc}^{-1}$. Ω expresses the energy density in units of $\rho_{0,\text{crit}}$. Thus the present density ρ_0 is represented by

$$\Omega_0 = \rho_0 / \rho_{0,\text{crit}}$$

Dark energy will turn repulsive state into an attractor!

□ findings from Cosmic Microwave Background (COBE, WMAP, PLANCK)



□ the universe is flat! $\Omega_0 \approx 1$. How to get this for any $k = \pm 1, 0$? \Rightarrow inflation

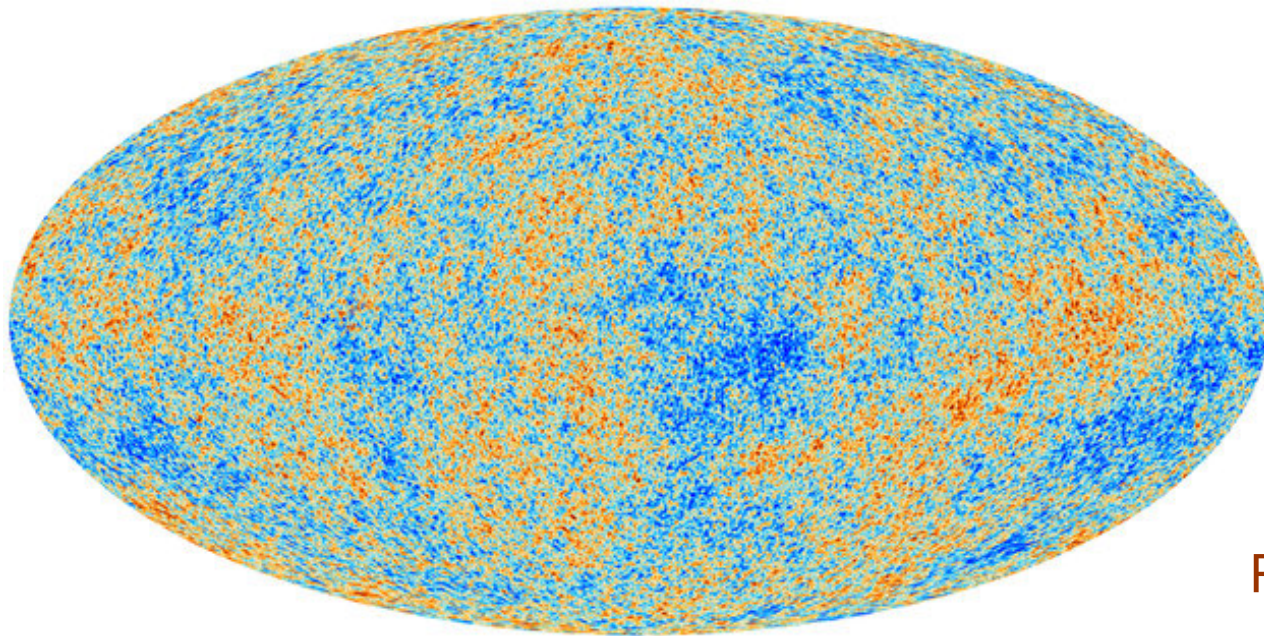
$$\Omega_0 = \Omega_\Lambda + \Omega_{\text{dark matter}} + \Omega_{\text{normal matter}} + \Omega_{\text{radiation}}$$

$$\Omega_\Lambda \simeq 0.74 ; \quad \Omega_{\text{dark matter}} \simeq 0.21 ; \quad \Omega_{\text{normal matter}} \simeq 0.05 ; \quad \Omega_{\text{radiation}} \simeq 0.003$$

The Cosmic Microwave Background

Cosmic black-body radiation of 3 °K Penzias, Wilson 1965, NP 1978

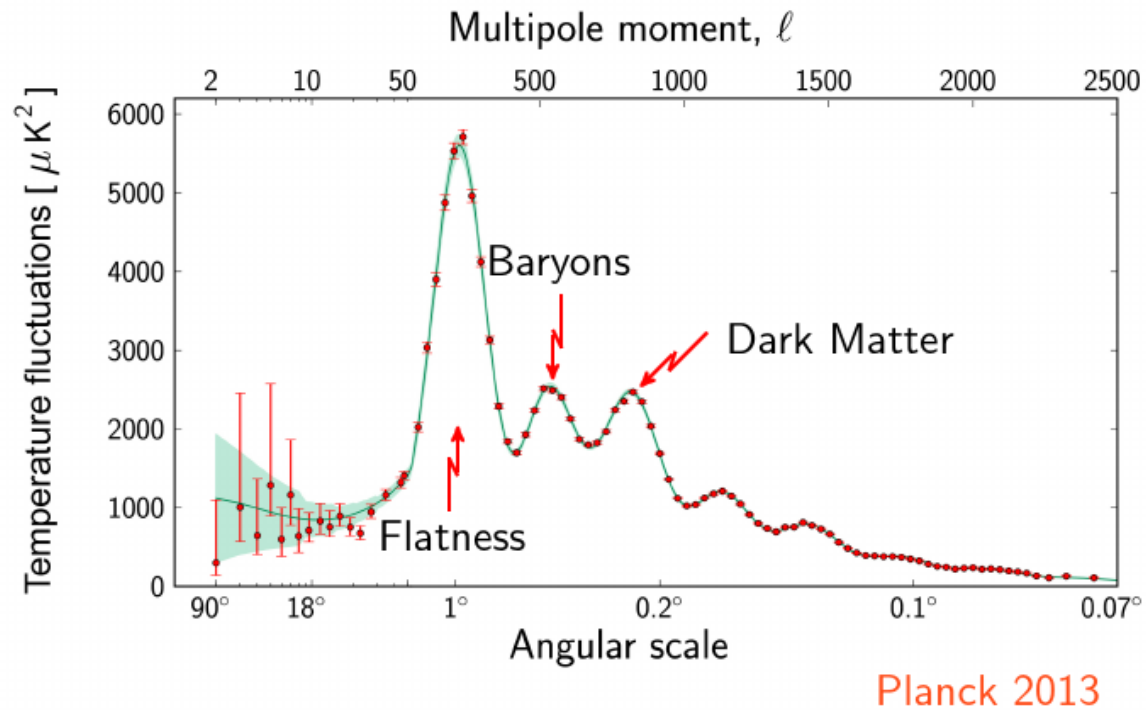
□ The CMB fluctuation pattern: imprinted on the sky when the universe was just **380 000 years** (after B.B.) old. Photons red-shifted by the expansion until they cannot ionize atoms (Hydrogen) any longer (snapshot of surface of last scattering). Smoot, Mather, NP 2006



PLANCK 2013

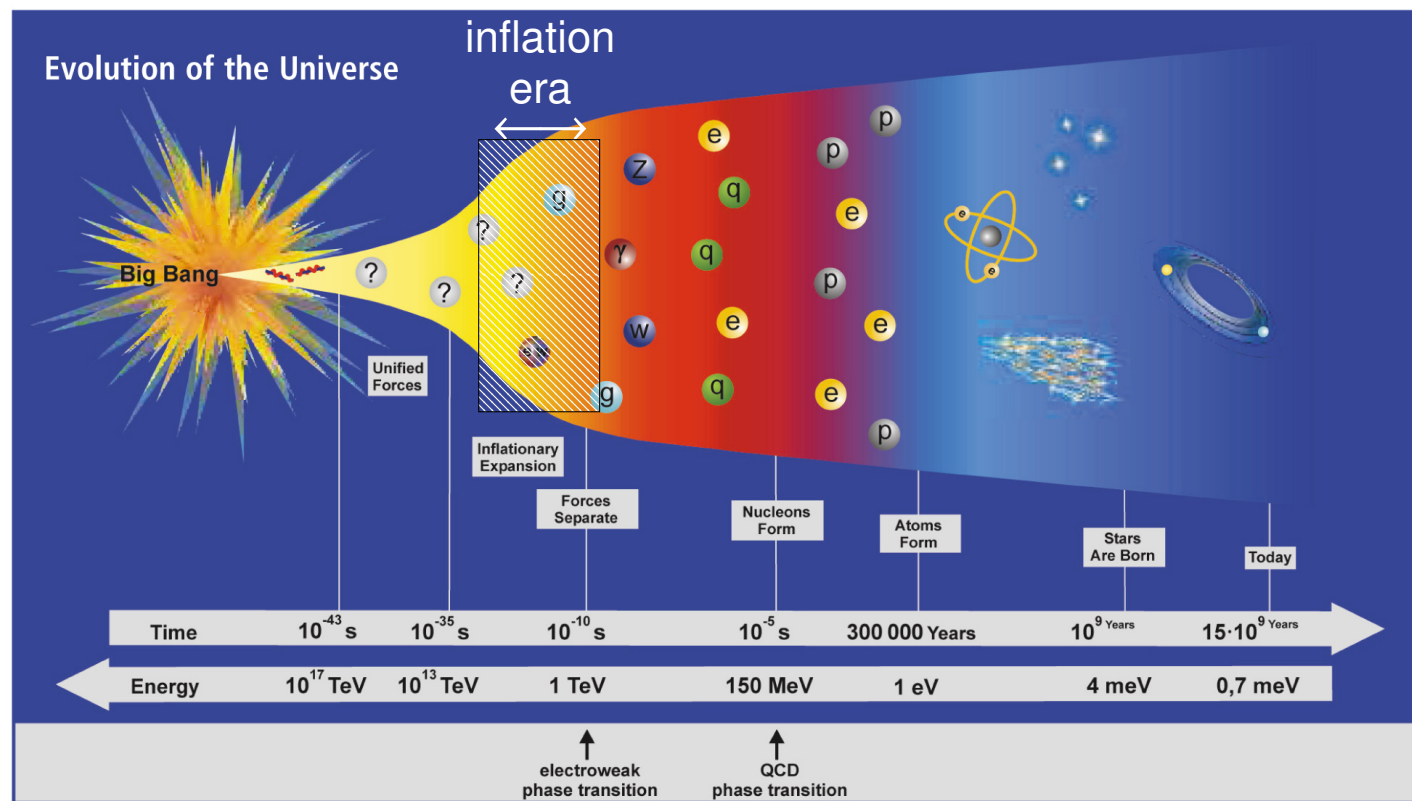
□ The power spectrum of CMB noise: (the acoustic peaks)

CMB shows the Universe is Flat [WMAP, PLANCK]



Inflation

Need inflation! universe must blow up exponentially for a very short period, such that we see it to be flat! [switch on anti-gravity for very short period of time]



Inflation at Work

Flatness, Causality, primordial Fluctuations \Rightarrow Solution: Guth 1980

Inflate the universe

Add an “Inflation term” to the r.h.s of the Friedmann equation, which dominates the very early universe blowing it up such that it looks flat afterwards

Need scalar field $\phi(x) \equiv$ “inflaton” : \Rightarrow inflation term $\frac{8\pi}{3 M_{\text{Pl}}^2} \left(V(\phi) + \frac{1}{2} \dot{\phi}^2 \right)$

Means: switch on strong anti-gravitation for an instant [sounds crazy]

Inflation: $a(t) \propto e^{Ht}$; $H = H(t) \equiv \dot{a}(t)/a(t)$ Hubble “constant”, i.e. $\frac{da}{a} = H(t) dt$

\Rightarrow $N \equiv \ln \frac{a_{\text{end}}}{a_{\text{initial}}} = H (t_e - t_i)$ automatic iff $V(\phi) \gg \frac{1}{2} \dot{\phi}^2$! slow roll!

“flattenization” by inflation: curvature term $k/a^2(t) \sim k \exp(-2Ht) \rightarrow 0$ ($k = 0, \pm 1$ the normalized curvature)

Assume too much matter, universe would re-collapse, inflation creates more space, density drops to critical one, re-collapse just prevented (shifted to $t = \infty$)!

SM Higgs as inflaton?

Energy-momentum tensor of SM $T_{\mu\nu} \hat{=} \Theta_{\mu\nu} = V(\phi) g_{\mu\nu} + \text{derivative terms}$

$$\rho_\phi = \frac{1}{2} \dot{\phi}^2 + V(\phi) \quad ; \quad p_\phi = \frac{1}{2} \dot{\phi}^2 - V(\phi)$$

□ Substitute energy density and pressure into Friedmann and fluid equation

□ Expansion when potential term dominates

$$\ddot{a} > 0 \iff p < -\frac{\rho}{3} \iff \frac{1}{2} \dot{\phi}^2 < V(\phi)$$

Equation of state (provided by the Higgs): $w = \frac{p}{\rho} = \frac{\frac{1}{2} \dot{\phi}^2 - V(\phi)}{\frac{1}{2} \dot{\phi}^2 + V(\phi)}$ is $V(\phi) \gg \frac{1}{2} \dot{\phi}^2$?

● small kinetic energy $\Rightarrow w \rightarrow -1$ is dark energy $p_\phi = -\rho_\phi < 0$!
indeed **Planck (2013)** finds $w = -1.13^{+0.13}_{-0.10}$.

Friedmann equation: $H^2 = \frac{8\pi G_N}{3} \left[V(\phi) + \frac{1}{2} \dot{\phi}^2 \right] \Rightarrow H^2 \simeq \frac{8\pi G_N}{3} V(\phi)$

Field equation: $\ddot{\phi} + 3H\dot{\phi} = -V'(\phi) \Rightarrow 3H\dot{\phi} \simeq -V'(\phi)$, for $V(\phi) \approx \frac{m^2}{2} \phi^2$ **harmonic oscillator with friction** \Rightarrow Gaussian inflation (**Planck 2013**)

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Field equation: $\cancel{\ddot{\phi}} + 3H\dot{\phi} = -V'(\phi) \Rightarrow 3H\dot{\phi} \simeq -V'(\phi)$, for $V(\phi) \approx \frac{m^2}{2} \phi^2$
exponential decay \Rightarrow stops inflation \Rightarrow oscillations set in

$$N \equiv \ln \frac{a(t_{\text{end}})}{a(t_{\text{initial}})} = \int_{t_i}^{t_e} H(t) dt \simeq -\frac{8\pi}{M_{\text{Pl}}^2} \int_{\phi_i}^{\phi_e} \frac{V}{V'} d\phi$$

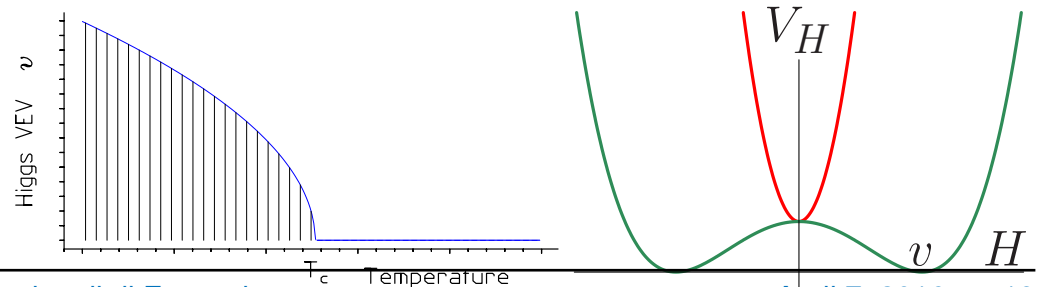
fixed
entirely by
scalar
potential

- need $N \gtrsim 60$, so called ***e*-folds** (CMB causal cone)

Key object of our interest: **the Higgs potential**

$$V = \frac{m^2}{2} H^2 + \frac{\lambda}{24} H^4$$

- Higgs mechanism = spontaneous $H \rightarrow -H$ symmetry breaking!
means: **symmetry at short distance scale, broken at low energies!**
- ❖ when m^2 changes sign and λ stays positive \Rightarrow first order phase transition
- ❖ vacuum jumps from $v = 0$ to $v \neq 0$



Summary part I:

- Inflation is established by observation (Flatness, Primordial Fluctuation etc)
- SM Higgs particle is ideal candidate for the Inflaton and dark energy

Key questions:

- does SM Higgs potential satisfy slow roll condition?
- does the SM provide sufficient amount of inflation?

Key problem:

- renormalized SM Higgs potential established at low energy cannot trigger inflation!

Therefore: standard opinion Higgs cannot be the inflaton

(Guth 1980 originally suggested the Higgs to be the inflaton!)

Standard paradigm:

- renormalizability is fundamental principle, only renormalized SM is physical
- symmetries if broken are broken spontaneously
- the higher the energy the more symmetry (SUSY, GUT, Strings)
- hierarchy problem requires SUSY, extra dimensions, little Higgs, ETC, etc

Emergence Paradigm and UV completion: the LEESM

The SM is a **low energy effective theory** of a **unknown Planck medium** [the “**ether**”], which exhibits the Planck energy as a physical cutoff: i.e. the SM **emerges** from a system shaped by gravitation

$$\Lambda_{\text{Pl}} = (G_N/c\hbar)^{-1/2} \simeq 1.22 \times 10^{19} \text{ GeV}$$

G_N Newton’s gravitational constant, c speed of light, \hbar Planck constant

- SM works up to Planck scale, means that it makes sense to consider the SM as the Planck medium **seen from far away** i.e. the SM is **emergent** at low energies. Expand in $E/\Lambda_{\text{Pl}} \Rightarrow$ see **renormalizable tail** only.
- looking at shorter and shorter distances (higher energies) we can see the bare Planck system as it was evolving from the Big Bang! **Energy scan in time!**
- the tool for accessing early cosmology is the RG solution of SM parameters: we can **calculate the bare parameters from the renormalized ones** determined at low (accelerator) energies.

- In the symmetric phase at very high energy we see the bare system:

the Higgs field is a **collective field** exhibiting an effective mass generated by radiative effects

$$m_{\text{bare}}^2 \approx \delta m^2 \text{ at } M_{\text{Pl}}$$

eliminates fine-tuning problem at all scales!

Many examples in condensed matter systems, **Coleman-Weinberg** mechanism

- “free lunch” in Low Energy Effective SM (LEESM) scenario:

- renormalizability of long range tail automatic!
- so are all ingredients required by renormalizability:
- non-Abelian gauge symmetries, chiral symmetry, anomaly cancellation, fermion families etc
- last but not least the existence of the Higgs boson!

*** all emergent ***
non-renormalizable stuff
heavily suppressed

The low energy expansion at a glance

		dimension	operator	scaling behavior	
hidden world	↑ no data 	·	∞—many irrelevant operators		
		·			
		·			
		$d = 6$	$(\Box\phi)^2, (\bar{\psi}\psi)^2, \dots$	$(E/\Lambda_{\text{Pl}})^2$	
		$d = 5$	$\bar{\psi}\sigma^{\mu\nu}F_{\mu\nu}\psi, \dots$	(E/Λ_{Pl})	
<hr/>					
world as seen	 experimental data ↓	$d = 4$	$(\partial\phi)^2, \phi^4, (F_{\mu\nu})^2, \dots$	$\ln(E/\Lambda_{\text{Pl}})$	tamed by symmetries
		$d = 3$	$\phi^3, \bar{\psi}\psi$	(Λ_{Pl}/E)	
		$d = 2$	$\phi^2, (A_\mu)^2$	$(\Lambda_{\text{Pl}}/E)^2$	
		$d = 1$	ϕ	$(\Lambda_{\text{Pl}}/E)^3$	
	Note: d=6 operators at LHC suppressed by $(E_{\text{LHC}}/\Lambda_{\text{Pl}})^2 \approx 10^{-30}$				

⇒ require chiral symmetry, gauge symmetry, ... ??? self-organized!
 – just looks symmetric as we cannot see the details –

What is natural?

Get everything you need to live (minimal extension), largely what we have.

The higher the energy the higher the symmetry??? Group theory is beautiful mathematics, already Kepler dreamed of the Platonic bodies (regular polyhedra) to become true in celestial mechanics. But finally Kepler's laws came out.

What is more natural: singlets, doublets, triplets, ... or 15-plets?

Actually $U(1) \otimes SU_L(2) \otimes SU_c(3)$ most natural conspiracies!

Hadrons are built from two, three, ... quarks ... pentaquarks already quite exotic (also QCD hadronic spectra follow nice group theory).

The hot and dense medium at the Planck time exhibits all types of fluctuation modes and the simplest combinations conspire to make it to be seen at long distances. The conspiracies in small n -plets are the effective long range symmetries, which protect the masses to freeze out. Chiral fermions, gauge bosons: massless before symmetries are spontaneously broken at low energies.

The Standard Model up to the Planck scale

The Cosmic Bridge

Universe is expanding: began in a very hot and dense state!

At Start a Light-Flash:

BIG BANG!

Light quanta very energetic, all matter totally ionized, all nuclei disintegrated.

Elementary particles only!: $\gamma, e^+, e^-, p, \bar{p}, \dots$

Early cosmology is Particle Physics!

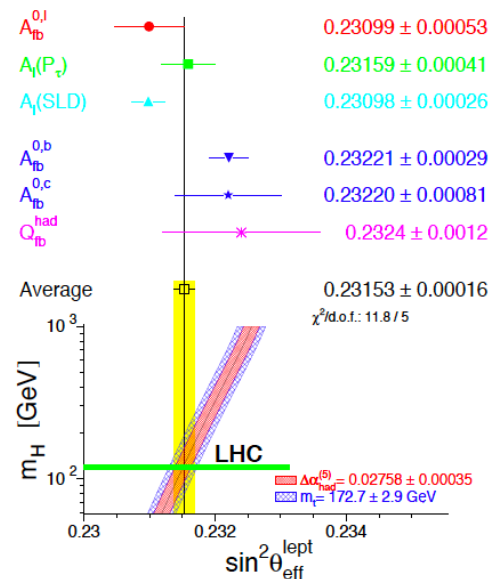
LEP type processes $e^+e^- \leftrightarrow \gamma^* \leftrightarrow X\bar{X}$ new forms of matter

Energy scale \leftrightarrow Temperature \leftrightarrow cosmic Time

$$E = 2 M_X c^2 \Leftrightarrow T = E/k_B \text{ }^\circ\text{K} \Leftrightarrow t = \frac{2.4}{\sqrt{g^*(T)}} \left(\frac{1\text{MeV}}{k_B T} \right)^2 \text{ sec. after B.B.}$$

The Higgs boson discovery – the SM completion

Higgs mass found by ATLAS and CMS agrees perfectly with the indirect bounds



LEP 2005 +++ LHC 2012

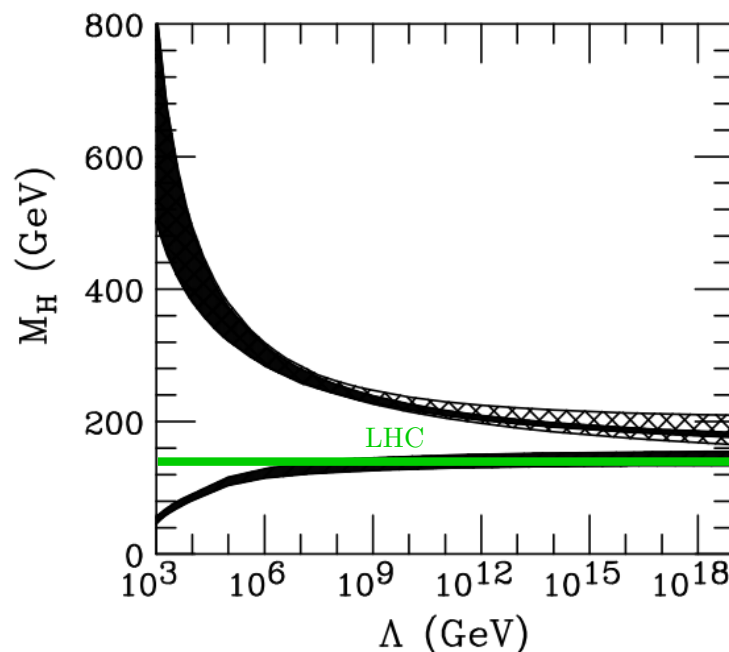
Englert&Higgs Nobel Prize 2013

Higgs mass found in very special mass range $125.9 \pm 0.4 \text{ GeV}$

Higgs boson predicted 1964 by Brout, Englert, Higgs – discovered 2012 at LHC by ATLAS&CMS

Common Folklore: SM hierarchy problem requires a supersymmetric (SUSY) extension of the SM (no quadratic/quartic divergences) **SUSY = infinity killer!**

Do we really need new physics? **Stability bound of Higgs potential** in SM:



$$V = \frac{m^2}{2} H^2 + \frac{\lambda}{24} H^4$$

Riessellmann, Hambye 1996

$$M_H < 180 \text{ GeV}$$

– first 2-loop analysis, knowing M_t –

SM Higgs remains perturbative up to scale Λ_{Pl} if it is light enough (upper bound=avoiding Landau pole) and Higgs potential remains stable ($\lambda > 0$) if Higgs mass is not too light [parameters used: $m_t = 175[150 - 200] \text{ GeV}$; $\alpha_s = 0.118$]

SM – Fermions: 28 per family $\Rightarrow 3 \times 28 = 84$; Gauge-Bosons: $1+3+8=12$; Scalars: 1 Higgs
Photon massless, gluons massless but confined

Before Higgs mechanism (triggering EW phase transition):

SM in symmetric phase : W^\pm, Z and all fermions massless

Higgs “ghosts” ϕ^\pm, ϕ^0 physical, heavy degenerate with the Higgs!

At “low” energy [likely up to 10^{16} GeV]:

$$V = \frac{m^2}{2} H^2 + \frac{\lambda}{24} H^4 ; \quad m^2 = -\mu^2 < 0$$

SM in broken phase : H, W^\pm, Z and all fermions massive [each mass requires separate new interaction via the Higgs: $2+12+1$ decay channels];

3 Higgs “ghosts” ϕ^\pm, ϕ^0 disappear and transmute into longitudinal DOFs of W^\pm, Z

Basic parameters: gauge couplings $g' = g_1, g = g_2, g_3$, top quark Yukawa coupling y_t , Higgs self-coupling λ and Higgs VEV v , besides smaller Yukawas.

Note: $1/(\sqrt{2}v^2) = G_F$ is the Fermi constant! [$v = (\sqrt{2}G_F)^{-1/2}$]

SSB \Rightarrow mass \propto interaction strength \times Higgs VEV v

$$\begin{array}{ll} M_W^2 &= \frac{1}{4} g^2 v^2 ; & M_Z^2 &= \frac{1}{4} (g^2 + g'^2) v^2 ; \\ m_f^2 &= \frac{1}{2} y_f^2 v^2 ; & M_H^2 &= \frac{1}{3} \lambda v^2 \end{array}$$

Effective parameters depend on renormalization scale μ [normalization reference energy!], scale at which ultraviolet (UV) singularities are subtracted

- **running couplings** change substantially with energy and hence as a function of time during evolution of the universe!
- high energy behavior governed by $\overline{\text{MS}}$ Renormalization Group (RG) [$E \gg M_i$]
- key input **matching conditions** between $\overline{\text{MS}}$ and physical parameters !
- running well established for electromagnetic α_{em} and strong coupling α_s :
 α_{em} screening, α_s anti-screening (Asymptotic Freedom)

Asked questions:

- does SM physics extend up to the Planck scale?
- do we need new physics beyond the SM to understand the early universe?
- does the SM collapse if there is no new physics?

“collapse”: Higgs potential gets unstable below the Planck scale; actually several groups claim to have proven vacuum stability break down at 3σ level!

Shaposhnikov et al, Degrassi et al, Maina, Hamada et al, ...

Scenario this talk: Higgs vacuum remains stable up and beyond the Planck scale
⇒ seem to say we do not need new physics affecting the evolution of SM couplings to investigate properties of the early universe. In the focus:

- does Higgs self-coupling stay positive $\lambda > 0$ up to Λ_{Pl} ?
- the key question/problem concerns the size of the top Yukawa coupling y_t
decides about stability of our world! — [$\lambda = 0$ would be essential singularity!]

Will be decided by: ● more precise input parameters
● better established EW matching conditions

Need vacuum stability and Higgs phase transition below M_{Pl} .

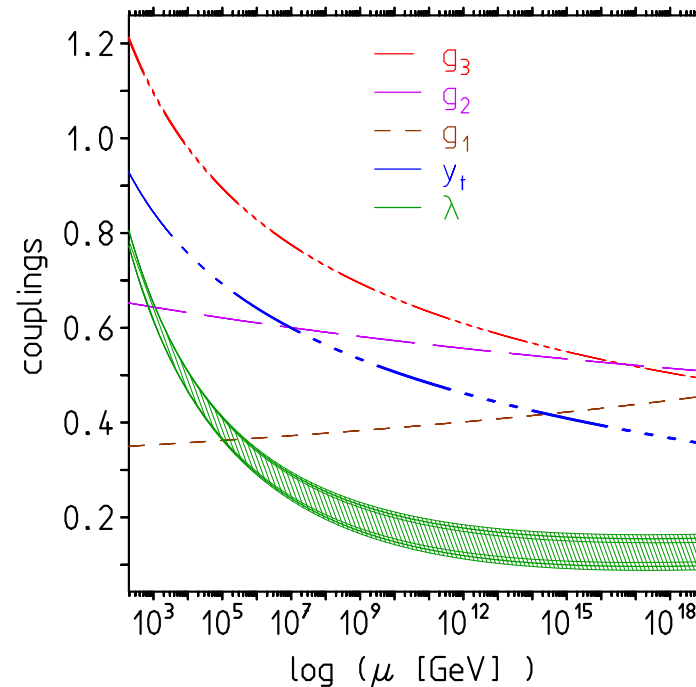
My evaluation of $\overline{\text{MS}}$ parameters revealed Vacuum Stability

Although other evaluations of the matching conditions seem to favor the meta-stability of the electroweak vacuum within the experimental and theoretical uncertainties, one should not exclude the possibility that other experiments and improved matching conditions will be able to establish the absolute stability of Standard Model in the future.

Shaposhnikov et al. [arXiv:1412.3811](#) say about Vacuum Stability

Although the present experimental data are perfectly consistent with the absolute stability of Standard Model within the experimental and theoretical uncertainties, one should not exclude the possibility that other experiments will be able to establish the meta-stability of the electroweak vacuum in the future.

The SM running parameters



The SM dimensionless couplings in the $\overline{\text{MS}}$ scheme as a function of the renormalization scale for $M_H = 124 - 127 \text{ GeV}$.

- perturbation expansion works up to the Planck scale!
no Landau pole or other singularities \Rightarrow Higgs potential remains stable!

- $U(1)_Y$ screening (IR free), $SU(2)_L$, $SU(3)_c$ antiscreening (UV free) [asymptotic freedom (AF)] – g_1, g_2, g_3

Right – as expected (standard wisdom)

- Top Yukawa y_t and Higgs λ : screening if lonely (IR free, like QED)

Wrong!!! – as part of SM, transmutation from IR free to AF

- running top Yukawa – QCD takes over: IR free \Rightarrow UV free
- running Higgs self-coupling – top Yukawa takes over: IR free \Rightarrow UV free
- Higgs coupling decreases up to the zero of β_λ at $\mu_\lambda \sim 3.5 \times 10^{17}$ GeV, where it is small but still positive and then increases up to $\mu = \Lambda_{\text{Pl}}$

The Higgs is special: before the symmetry is broken: all particles massless protected by gauge or chiral symmetry except the four Higgses. Two quantities affected: **Higgs boson mass** and **Higgs vacuum energy**

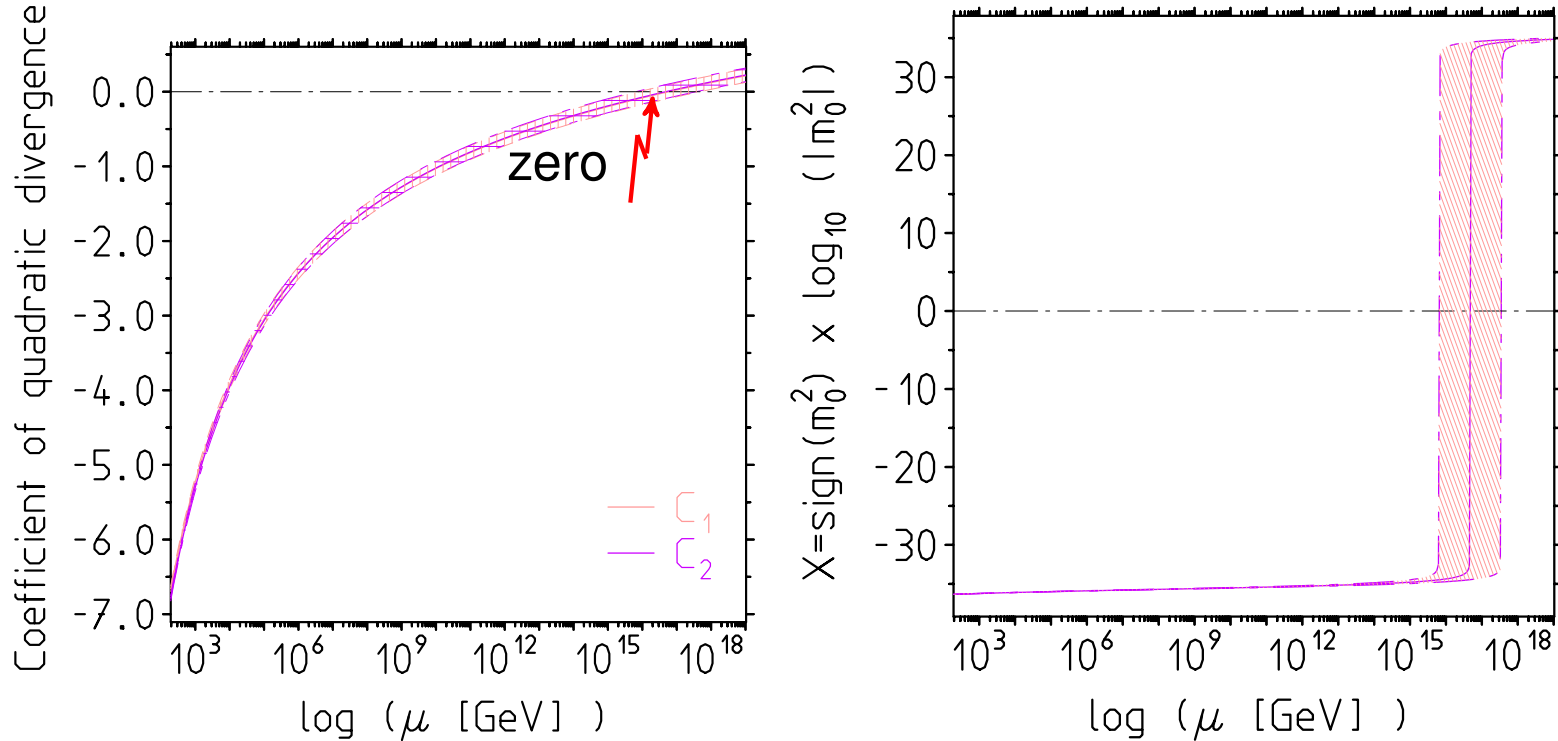
The Role of Quadratic Divergences in the SM

Veltman 1978 [NP 1999] modulo small lighter fermion contributions, one-loop coefficient function C_1 is given by

$$\delta m_H^2 = \frac{\Lambda_{\text{Pl}}^2}{16\pi^2} C_1 ; \quad C_1 = \frac{6}{v^2}(M_H^2 + M_Z^2 + 2M_W^2 - 4M_t^2) = 2\lambda + \frac{3}{2}g'^2 + \frac{9}{2}g^2 - 12y_t^2$$

Key points:

- ➡ C_1 is universal and depends on dimensionless gauge, Yukawa and Higgs self-coupling only, the RGs of which are **unambiguous**.
At two loops $C_2 \approx C_1$ numerically [Hamada et al 2013] stable under RCs!
- ➡ Couplings are running! $C_i = C_i(\mu)$
- ➡ the SM for the given running parameters makes a prediction for the bare effective mass parameter in the Higgs potential:



The Higgs phase transition in the SM [for $M_H = 125.9 \pm 0.4 \text{ GeV}$].

$$m_{\text{bare}}^2 = \text{sign}(m_{\text{bare}}^2) \times 10^X$$

Jump in vacuum energy: wrong sign and 50 orders of magnitude off Λ_{CMB} !!!

$$\Delta V(\phi_0) = -\frac{m_{\text{eff}}^2 v^2}{8} = -\frac{\lambda v^4}{24} \sim -(176.0 \text{ GeV})^4$$

\Rightarrow one version of CC problem

□ in the broken phase $m_{\text{bare}}^2 = \frac{1}{2} m_{H \text{ bare}}^2$, which is calculable!

⇒ the coefficient $C_n(\mu)$ exhibits a zero, for $M_H = 126 \text{ GeV}$ at about $\mu_0 \sim 1.4 \times 10^{16} \text{ GeV}$, not far below $\mu = M_{\text{Planck}}$!!!

⇒ at the zero of the coefficient function the counterterm $\delta m^2 = m_{\text{bare}}^2 - m^2 = 0$ (m the $\overline{\text{MS}}$ mass) vanishes and the bare mass changes sign

⇒ this represents a **phase transition** (PT), which **triggers** the **Higgs mechanism** as well as **cosmic inflation** as $V(\phi) \gg \dot{\phi}^2$ for $\mu > \mu_0$

⇒ at the transition point μ_0 we have $v_{\text{bare}} = v(\mu_0^2)$; $m_{H \text{ bare}} = m_H(\mu_0^2)$,
where $v(\mu^2)$ is the $\overline{\text{MS}}$ renormalized VEV

In any case at the zero of the coefficient function there is a phase transition, which corresponds to a **restoration of the symmetry** in the **early universe**.

Hot universe \Rightarrow finite temperature effects:

□ finite temperature effective potential $V(\phi, T)$:

$$T \neq 0: V(\phi, T) = \frac{1}{2} \left(g_T T^2 - \mu^2 \right) \phi^2 + \frac{\lambda}{24} \phi^4 + \dots$$

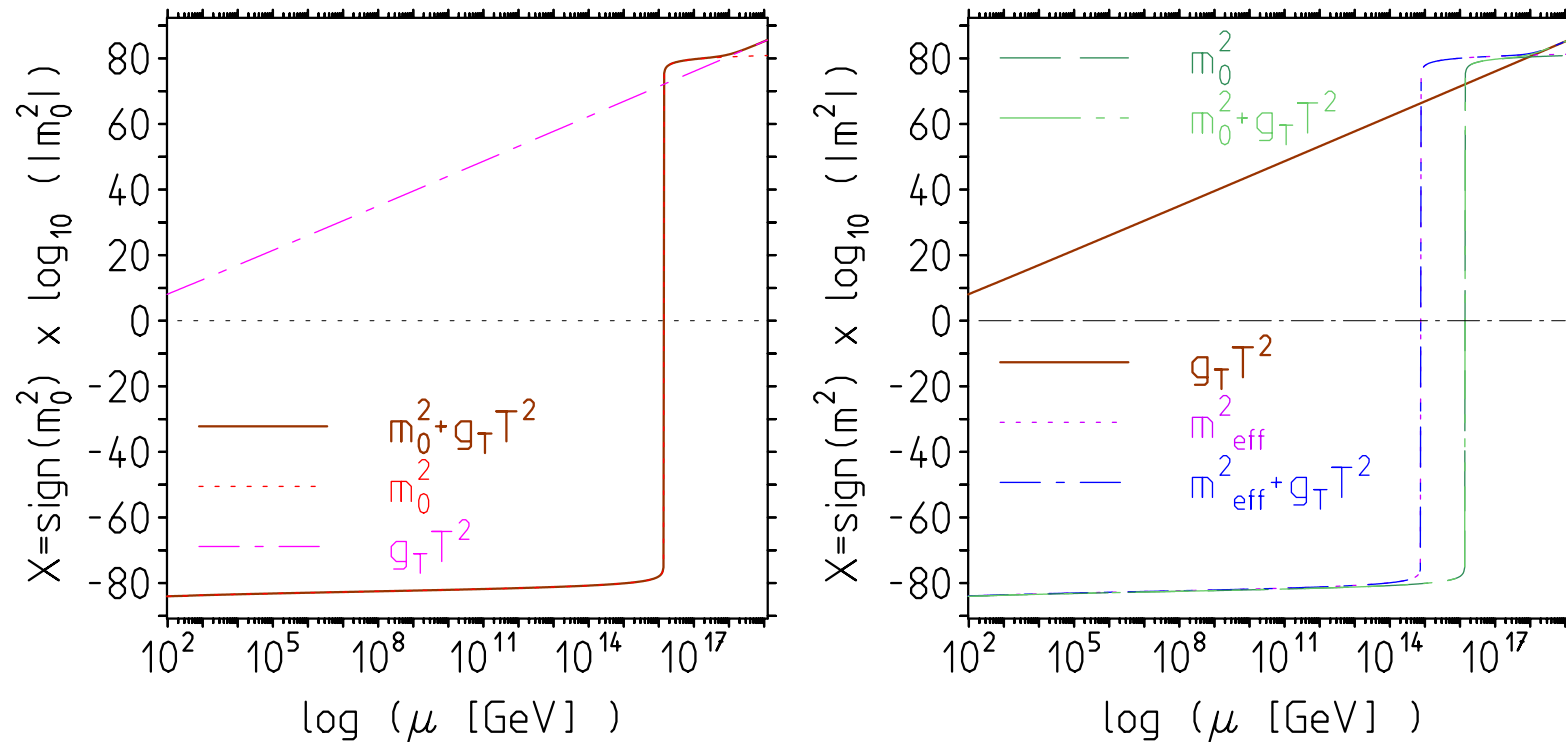
Usual assumption: Higgs is in the broken phase $\mu^2 > 0$ and $\mu \sim v$ at EW scale

EW phase transition is taking place when the universe is cooling down below the critical temperature $T_c = \sqrt{\mu^2/g_T}$.

My scenario: above PT at μ_0 SM in symmetric phase $-\mu^2 \rightarrow m^2 = (m_H^2 + \delta m_H^2)/2$

$$m^2 \sim \delta m^2 \simeq \frac{M_{\text{Pl}}^2}{32\pi^2} C(\mu = M_{\text{Pl}}) \simeq (0.0295 M_{\text{Pl}})^2, \quad \text{or} \quad m^2(M_{\text{Pl}})/M_{\text{Pl}}^2 \approx 0.87 \times 10^{-3}.$$

In fact with our value of μ_0 almost no change of phase transition point by FT effects. True effective mass $m^2 \rightarrow m'^2$ from Wick ordered Lagrangian [$C \rightarrow C + \lambda$].



Effects on the phase transition by finite temperature and vacuum rearrangement

$$\mu_0 \approx 1.4 \times 10^{16} \text{ GeV} \rightarrow \mu'_0 \approx 7.7 \times 10^{14} \text{ GeV},$$

Up to shift in transition temperature PT is triggered by δm^2 and EW PT must be close by at about $\mu_0 \sim 10^{15} \text{ GeV}$ not at EW scale $v \sim 246 \text{ GeV}$!

Important for Baryogenesis!

The Cosmological Constant in the SM

- in symmetric phase $SU(2)$ is a symmetry: $\Phi \rightarrow -U(\omega)\Phi$ and $\Phi^+\Phi$ singlet;

$$\langle 0|\Phi^+\Phi|0\rangle = \frac{1}{2}\langle 0|H^2|0\rangle \equiv \frac{1}{2}\Xi; \quad \Xi = \frac{\Lambda_{\text{Pl}}^2}{16\pi^2}.$$

just Higgs self-loops

$$\langle H^2 \rangle =: \text{loop with one vertex} ; \quad \langle H^4 \rangle = 3 (\langle H^2 \rangle)^2 =: \text{loop with two vertices}$$

⇒ vacuum energy $V(0) = \langle V(\phi) \rangle = \frac{m^2}{2}\Xi + \frac{\lambda}{8}\Xi^2$; mass shift $m'^2 = m^2 + \frac{\lambda}{2}\Xi$

□ for our values of the $\overline{\text{MS}}$ input parameters $m^2 \rightarrow m'^2$

⇒ $\mu_0 \approx 1.4 \times 10^{16} \text{ GeV} \rightarrow \mu'_0 \approx 7.7 \times 10^{14} \text{ GeV},$

- potential of the fluctuation field $\Delta V(\phi)$.

⇒ quasi-constant vacuum density $V(0)$ representing the cosmological constant

⇒ $H \simeq \ell \sqrt{V(0) + \Delta V}$ at M_{Pl} we expect $\phi_0 = \mathcal{O}(M_{\text{Pl}})$ i.e. at start $\Delta V(\phi) \gg V(0)$

□ fluctuation field eq. $3H\dot{\phi} \approx -(m'^2 + \frac{\lambda}{6}\phi^2)\phi$, ϕ decays exponentially, must have been very large in the early phase of inflation

● need $\phi_0 \approx 4.51 M_{\text{Pl}}$, big enough to provide sufficient inflation. Note: this is the only free parameter in SM inflation, the Higgs field is not an observable in the renormalized low energy world (laboratory/accelerator physics).

Decay patterns:

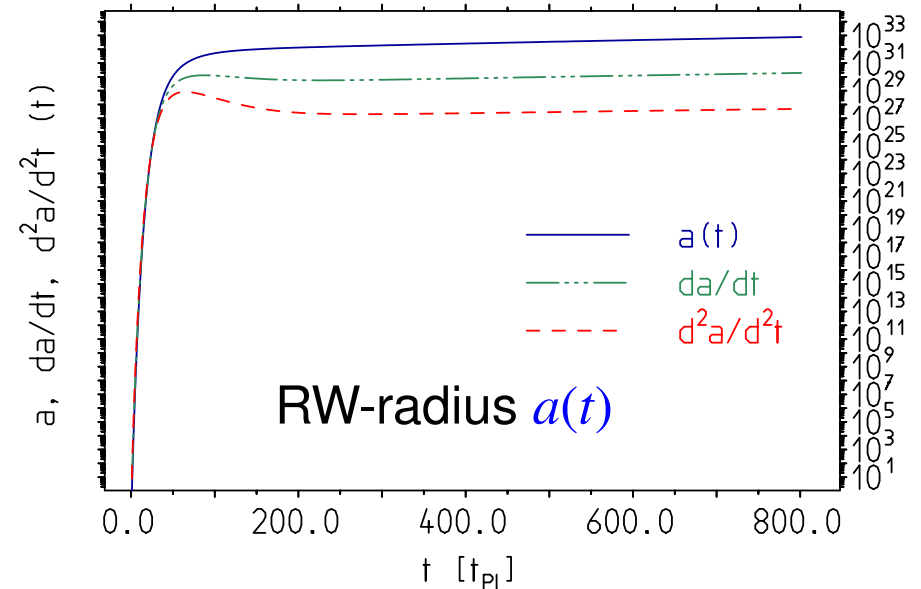
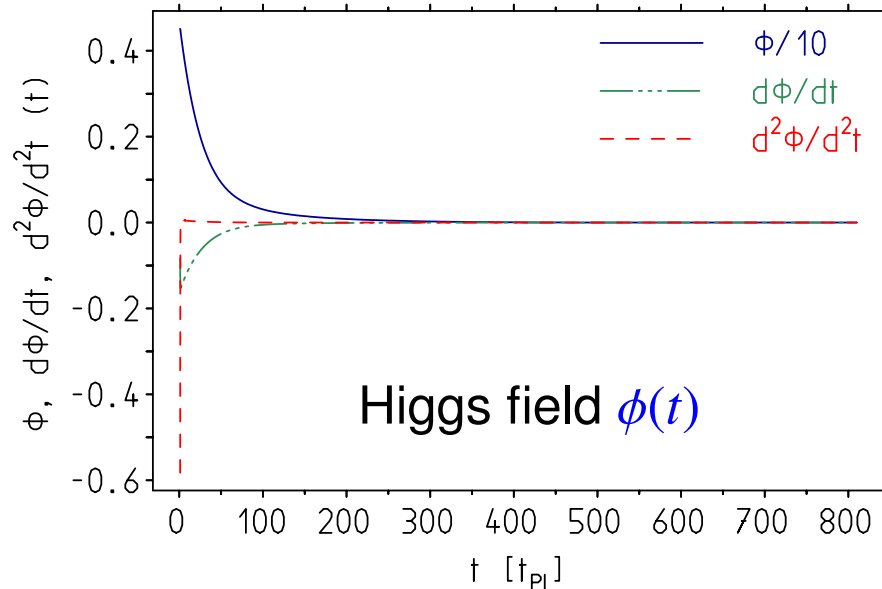
$$\phi(t) = \phi_0 \exp\{-E_0(t - t_0)\}, \quad E_0 \approx \frac{\sqrt{2\lambda}}{3\sqrt{3}\ell}, \quad \approx 4.3 \times 10^{17} \text{ GeV}, \quad V_{\text{int}} \gg V_{\text{mass}}$$

soon mass term dominates, in fact $V(0)$ and V_{mass} are comparable before $V(0)$ dominates and $H \approx \ell \sqrt{V(0)}$ and

$$\phi(t) = \phi_0 \exp\{-E_0(t - t_0)\}, \quad E \approx \frac{m^2}{3\ell \sqrt{V(0)}} \approx 6.6 \times 10^{17} \text{ GeV}, \quad V_{\text{mass}} \gg V_{\text{int}}$$

Note: if no CC ($V(0) \approx 0$) as assumed usually

$$\phi(t) = \phi_0 - X_0(t - t_0), \quad X_0 \approx \frac{\sqrt{2}m}{3\ell} \approx 7.2 \times 10^{35} \text{ GeV}^2, \quad V_{\text{mass}} \gg V_{\text{int}}$$



Note: the Hubble constant in our scenario, in the symmetric phase, during the radiation dominated era is given by (Stefan-Boltzmann law)

$$H = \ell \sqrt{\rho_{\text{rad}}} \simeq 1.66 (k_B T)^2 \sqrt{102.75} M_{\text{Pl}}^{-1}$$

such that at Planck time (SM predicted)

$$H_i \simeq 16.83 M_{\text{Pl}} .$$

i.e. trans-Planckian $\phi_0 \sim 5 M_{\text{Pl}}$ is not unnatural!

Note: inflation stops because of the extremely fast decay of the Higgs field ($t_{\text{end}} \lesssim 100 t_{\text{Pl}}$)

How to get rid of the huge CC?

- $V(0)$ very weakly scale dependent (running couplings): how to get rid of?

Note total energy density as a function of time

$$\rho(t) = \rho_{0,\text{crit}} \left\{ \Omega_{\Lambda} + \Omega_{0,\text{k}} (a_0/a(t))^2 + \Omega_{0,\text{mat}} (a_0/a(t))^3 + \Omega_{0,\text{rad}} (a_0/a(t))^4 \right\}$$

reflects a present-day snapshot. Cosmological constant is constant! Not quite!

- intriguing structure again: the effective CC counterterm has a zero, which again is a point where renormalized and bare quantities are in agreement:

$$\rho_{\Lambda \text{ bare}} = \rho_{\Lambda \text{ ren}} + \frac{M_{\text{Pl}}^4}{(16\pi^2)^2} X(\mu)$$

with $X(\mu) \simeq 2C(\mu) + \lambda(\mu)$ which has a zero close to the zero of $C(\mu)$
when $2C(\mu) = -\lambda(\mu)$, which happens at

$$\mu_{\text{CC}} \approx 3.1 \times 10^{15} \text{ GeV}$$

in between $\mu_0 \approx 1.4 \times 10^{16} \text{ GeV}$ and $\mu'_0 \approx 7.7 \times 10^{14} \text{ GeV}$.

Again we find a matching point between low energy and high energy world:

$$\rho_{\Lambda \text{ bare}} = \rho_{\Lambda \text{ ren}}$$

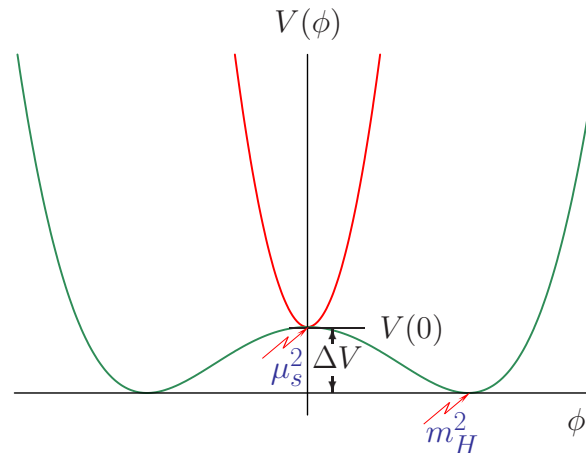
where memory of quartic Planck scale enhancement gets lost!

Has there been a cosmological constant problem?

Crucial point $X = 2C + \lambda = 5\lambda + 3g'^2 + 9g^2 - 24y_t^2$ acquires positive bosonic contribution and negative fermionic ones, with different scale dependence. X can change a lot (pass a zero), while individual couplings are weakly scale dependent $y_t(M_Z)/y_t(M_{\text{Pl}}) \sim 2.7$ biggest, $g_1(M_Z)/g_1(M_{\text{Pl}}) \sim 0.76$ smallest.

□ SM predicts huge CC at M_{Pl} : $\rho_\phi \simeq V(\phi) \sim 2.77 M_{\text{Pl}}^4 \sim (1.57 \times 10^{19} \text{ GeV})^4$
how to tame it?

At Higgs transition: $m'^2(\mu < \mu'_0) < 0$ vacuum rearrangement of Higgs potential



How can it be: $V(0) + \Delta V \sim (0.002 \text{ eV})^4$???

The zero $X(\mu_{CC}) = 0$ provides part of the answer as it makes $\rho_{\Lambda \text{ bare}} = \rho_{\Lambda \text{ ren}}$ to be identified with the observed value?

Seems to be naturally small, since Λ_{Pl}^4 term nullified at matching point.

Note: in principle, like the Higgs mass in the LEESM, also $\rho_{\Lambda \text{ ren}}$ is expected to be a free parameter to be fixed by experiment.

Not quite! there is a big difference: inflation forces $\rho_{\text{tot}}(t) \approx \rho_{0,\text{crit}} = \text{constant}$ after inflation era

$$\Omega_{\text{tot}} = \Omega_{\Lambda} + \Omega_{\text{mat}} + \Omega_{\text{rad}} = \Omega_{\Lambda} + \Omega_{0,\text{k}} (a_0/a(t))^2 + \Omega_{0,\text{mat}} (a_0/a(t))^3 + \Omega_{0,\text{rad}} (a_0/a(t))^4 \approx 1$$

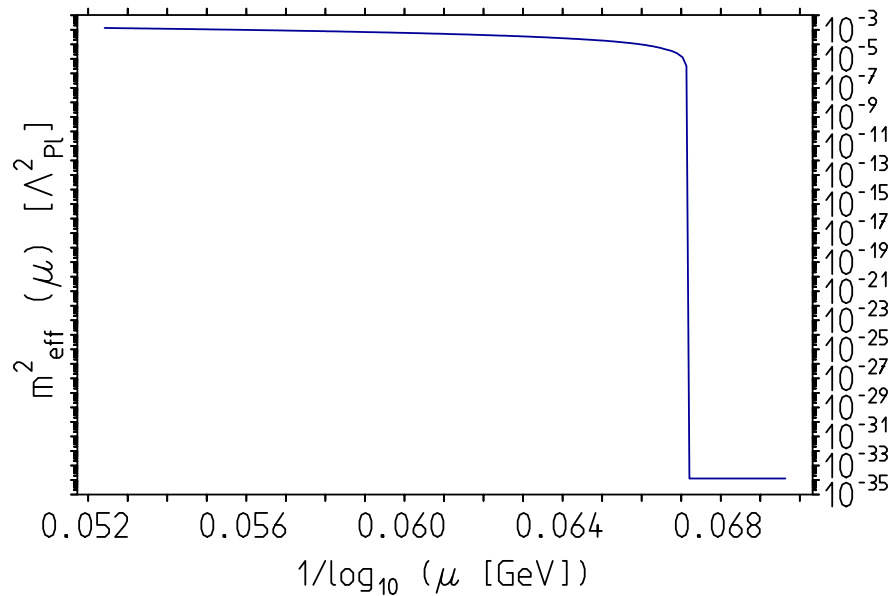
and since $1 > \Omega_{\text{mat}}$, $\Omega_{\text{rad}} > 0$ actually Ω_{Λ} is fixed once we know dark matter, baryonic matter and the radiation density:

$$\Omega_{\Lambda} = 1 - \Omega_{\text{mat}} - \Omega_{\text{rad}}$$

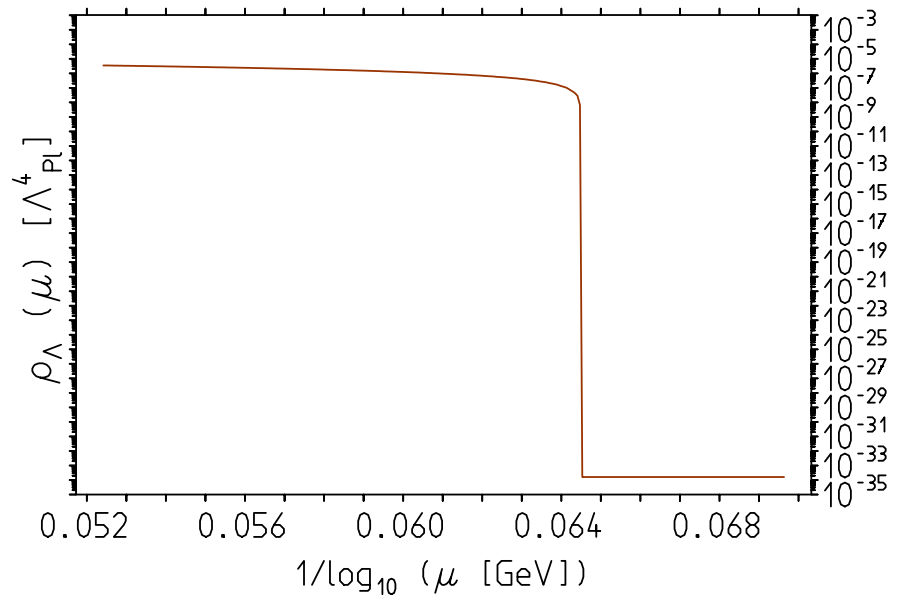
So, where is the miracle to have **CC of the magnitude of the critical density** of a flat universe? Also this then is a prediction of the LEESM!

Note that $\Omega_{\text{tot}} = 1$ requires Ω_{Λ} to be a function of t , up to negligible terms,

$$\Omega_{\Lambda} \rightarrow \Omega_{\Lambda}(t) \approx 1 - (\Omega_{0,\text{dark mat}} + \Omega_{0,\text{baryonic mat}}) (a_0/a(t))^3 \rightarrow 1 ; \quad t \rightarrow \infty$$



effective Higgs mass square



effective dark energy density

in units of Λ_{Pl} , for $\mu < \mu_{\text{CC}}$ we display $\rho_{\Lambda}[\text{GeV}^4] \times 10^{13}$ as predicted by SM

$\rho_{\Lambda} = \mu_{\Lambda}^4$: $\mu_{0,\Lambda} = 0.002 \text{ eV}$ today \rightarrow approaching $\mu_{\infty,\Lambda} = 0.00216 \text{ eV}$ with time

Remark: $\Omega_{\Lambda}(t)$ includes besides the large positive $V(0)$ also negative contributions from vacuum condensates, like $\Delta\Omega_{\text{EW}}$ from the Higgs mechanism and $\Delta\Omega_{\text{QCD}}$ from the chiral phase transition.

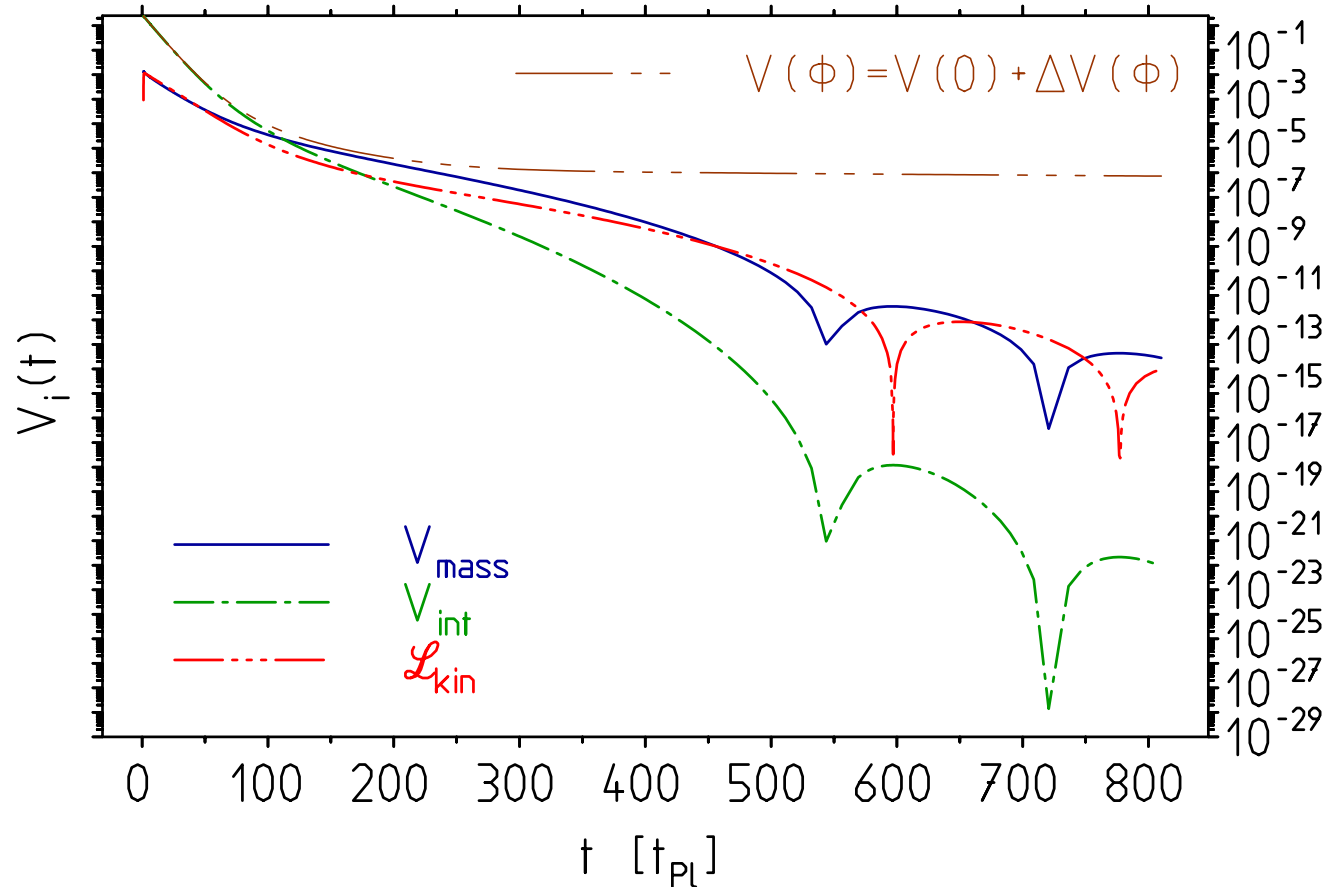
The Higgs Boson is the Inflaton!

- after electroweak PT, at the zeros of quadratic and quartic “divergences”, memory of cutoff lost: renormalized low energy parameters match bare parameters
- in symmetric phase (early universe) bare effective mass and vacuum energy dramatically enhanced by quadratic and quartic cutoff effects
- ⇒ slow-roll inflation condition $\frac{1}{2}\dot{\phi}^2 \ll V(\phi)$ satisfied
- ⇒ Higgs potential provides huge dark energy in early universe which triggers inflation

The SM predicts dark energy and inflation!!!

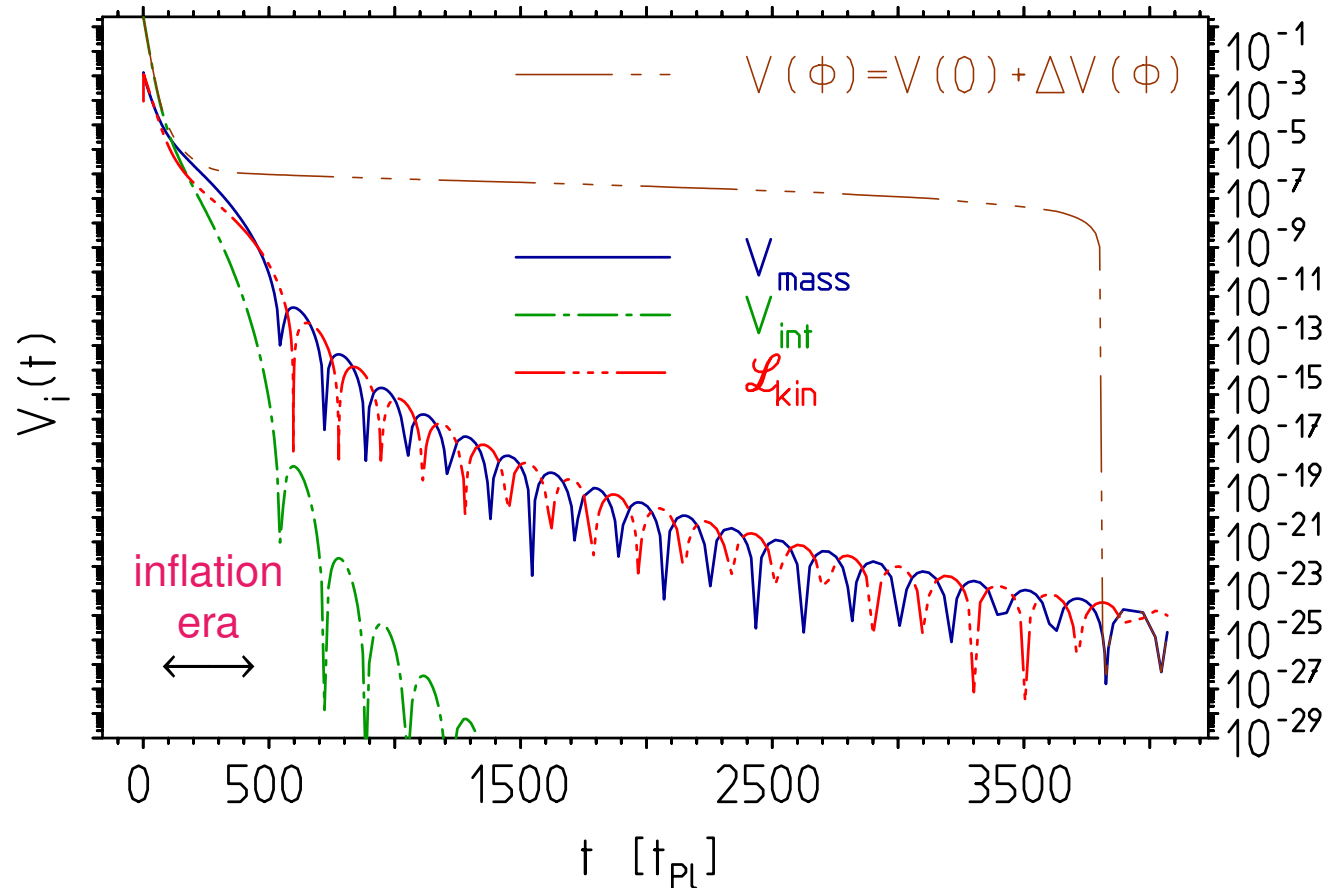
dark energy and inflation are unavoidable consequences of the SM Higgs
(provided new physics does not disturb it substantially)

The evolution of the universe before the EW phase transition:



Inflation epoch ($t \lesssim 450 t_{Pl}$): the mass-, interaction- and kinetic-term of the bare Lagrangian in units of M_{Pl}^4 as a function of time.

The evolution of the universe before the EW phase transition:



Evolution until symmetry breakdown and vanishing of the CC. After inflation quasi-free damped harmonic oscillator behavior (reheating phase).

Comment on Reheating and Baryogenesis

- inflation: exponential growth = exponential cooling
- reheating: pair created heavy states X, \bar{X} in originally hot radiation dominated universe decay into lighter matter states which reheat the universe
- baryogenesis: X particles produce particles of different baryon-number B and/or different lepton-number L . B by SM sphalerons or nearby dim 6 effective interactions

Sacharow condition for baryogenesis:



- small B is natural in LEESM scenario due to the close-by dimension 6 operators Weinberg 1979, Buchmüller, Wyler 1985, Grzadkowski et al 2010
Not really new physics as they are build from SM fields!

□ suppressed by $(E/\Lambda_{\text{Pl}})^2$ in the low energy expansion. At the scale of the EW phase transition the Planck suppression factor is 1.3×10^{-6} .

□ six possible four-fermion operators all $B - L$ conserving!

● \cancel{C} , \cancel{CP} , out of equilibrium

X is the Higgs! – “unknown” X particles now the known very heavy Higgs in symmetric phase of SM: Primordial Planck medium Higgses

All relevant properties known: mass, width, branching fractions, CP violation properties!

Stages: □ $k_B T > m_X \Rightarrow$ thermal equilibrium X production and X decay in balance

□ $H \approx \Gamma_X$ and $k_B T < m_X \Rightarrow X$ -production suppressed, out of equilibrium

- $H \rightarrow t\bar{t}, b\bar{b}, \dots$ predominantly (largest Yukawa couplings)
- CP violating decays: $H^+ \rightarrow t\bar{d}$ [rate $\propto y_t y_d V_{td}$] $H^- \rightarrow b\bar{u}$ [rate $\propto y_b y_u V_{ub}$] and after EW phase transition: $t \rightarrow de^+ \nu$ and $b \rightarrow ue^- \nu_e$ etc.
- Note: before Higgs mechanism bosonic triple couplings like HWW , HZZ are absent (induced by SSB after EW phase transition).
- Preheating absent! Reheating via $\phi \rightarrow f\bar{f}$ while all bosonic decays heavily suppressed (could obstruct reheating)!

Seems we are all descendants of four heavy Higgses via top-bottom stuff!

Baryogenesis most likely a “SM + dim 6 operators” effect!

Unlikely: $B + L$ violating instanton effects $\propto \exp \left[-\frac{8\pi^2}{g^2(\mu)} + \dots \right] \approx e^{-315.8}$ too small.

\Rightarrow observed baryon asymmetry $\eta_B \sim 10^{-10}$ cannot be a SM prediction, requires unknown B violating coupling. But order of magnitude looks to be “explainable”.

Conclusion

- ❑ The LHC made tremendous step forward in SM physics and cosmology: the discovery of the **Higgs boson**, which fills the vacuum of the universe first with **dark energy** and later with the Higgs **boson condensate**, thereby providing mass to quarks leptons and the weak gauge bosons, but **also drives inflation, reheating and all that**
- ❑ “Higgs not just the Higgs”: its mass $M_H = 125.9 \pm 0.4 \text{ GeV}$ has a very **peculiar value**, which opens the narrow window to the Planck world!
- ❑ SM parameter space tailored such that strange exotic phenomena like **inflation** and likely also the continued **accelerated expansion** of the universe are a direct consequence of LEESM physics.
- ➡ ATLAS and CMS results may “revolutionize” particle physics in an unexpected way, namely showing that the SM has higher self-consistency (conspiracy) than expected and previous arguments for the existence of new physics may turn out not to be compelling

⇒ SM as a low energy effective theory of some cutoff system at M_{Pl} consolidated; crucial point $M_{\text{Pl}} \gg \gg \gg \dots$ from what we can see!

⇒ the huge gap $E_{\text{lab}} \ll \ll \ll M_{\text{Pl}}$ lets look particle physics to follow **fundamental** laws (following simple principles, QFT structure)

⇒ change in paradigm:

Natural scenario understands the SM as the “true world” seen from far away

This is anyway what it is!

⇒ Methodological approach known from investigating condensed matter systems. (QFT as long distance phenomenon, critical phenomena)
Wilson 1971, NP 1982; also

⇒ Non-Abelian gauge symmetries as low energy phenomenon
Veltman, Bell, Lewellyn Smith, Cornwall, Levin, Tiktopoulos and others

- ➡ cut-offs in particle physics are important to understand early cosmology, i.e. inflation, reheating, baryogenesis and all that.
- ➡ the LEESM scenario, for the given now known parameters, the SM predicts dark energy and inflation, i.e. they are unavoidable
- ➡ in contrast to “the higher the more symmetric” (SUSY, GUT etc) which have no phenomenological support (only real as imaginations), the LEESM scenario predicts a well established observational fact: dark energy and inflation without the need of any ad hoc assumptions

● So what is “new”?

Take hierarchy problem argument serious, SM should exhibit Higgs mass of Planck scale order (what is true in the symmetric phase), as well as vacuum energy of order Λ_{Pl}^4 , but do not try to eliminate them by imposing supersymmetry or what else, just take the SM regularized by the Planck cutoff as it is.

➡ inflation seems to be strong indication that quadratic and quartic cutoff enhancements are real, as predicted by LatticeSM for instance, i.e.

Power divergences of local QFT are not the problem they are the solution!

● New physics: still must exist

- ① cold dark matter
- ② axions as required by strong CP problem
- ③ singlet neutrino puzzle (Majorana vs Dirac) and likely more \dots , however, NP should not kill huge effects in quadratic and quartic cutoff sensitive terms and it should not deteriorate gross pattern of the running of the SM couplings. As most Yukawa couplings (besides y_t).

Dark Energy: The Biggest Mystery in the Universe

**Unless you accept the SM
supplemented with a physical cutoff!**

Thanks

Thanks for your attention and the kind hospitality!

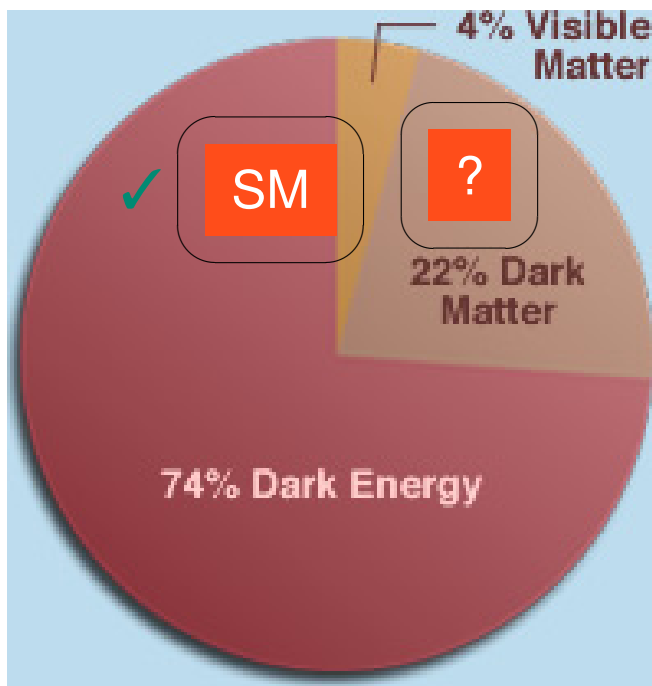


B. Touschek

Backup Slides



Last but not least: today's dark energy = relict Higgs vacuum energy?



WHAT IS DARK ENERGY?

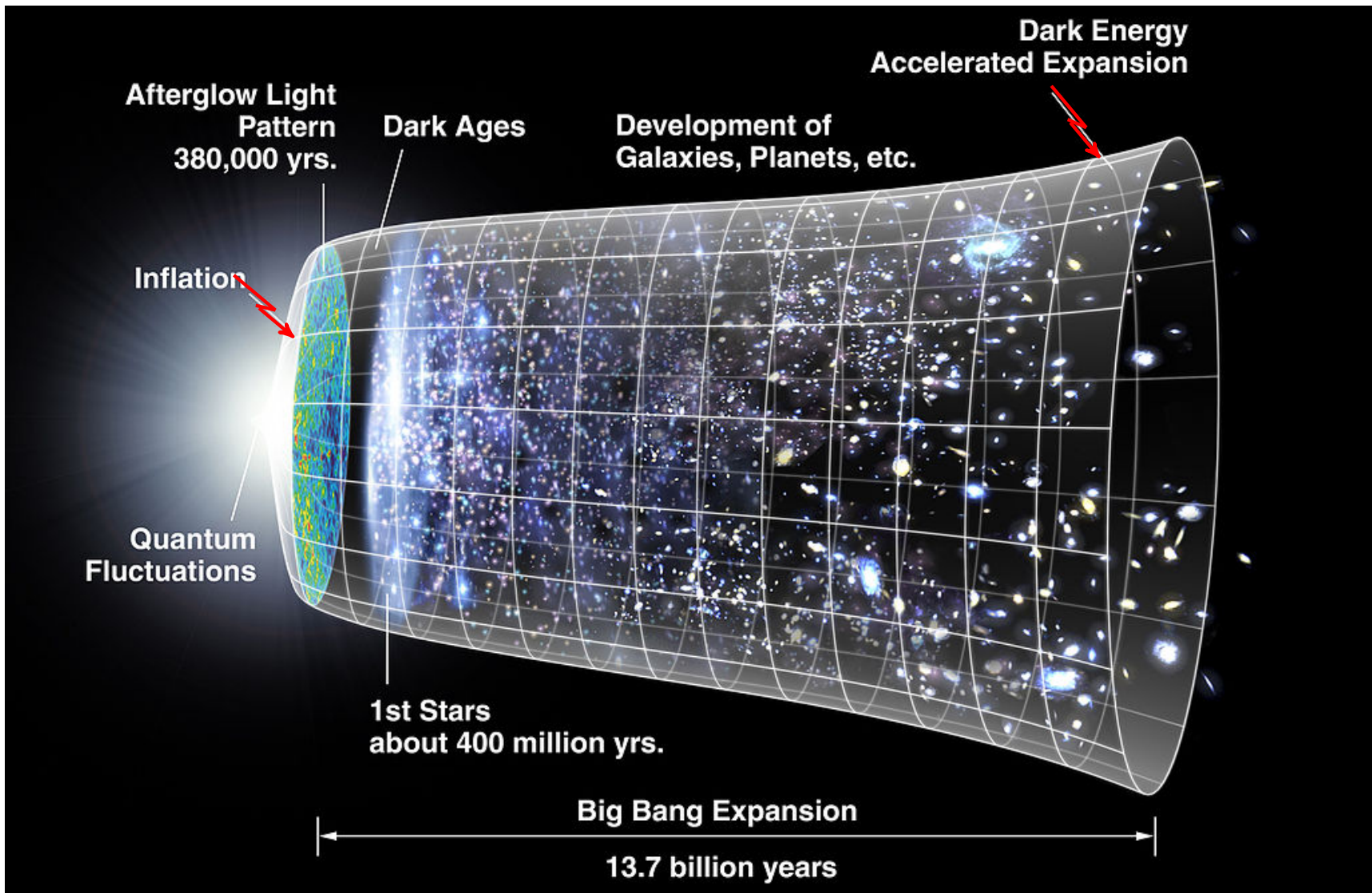
Well, the simple answer is that we don't know.

It seems to contradict many of our understandings about the way the universe works.

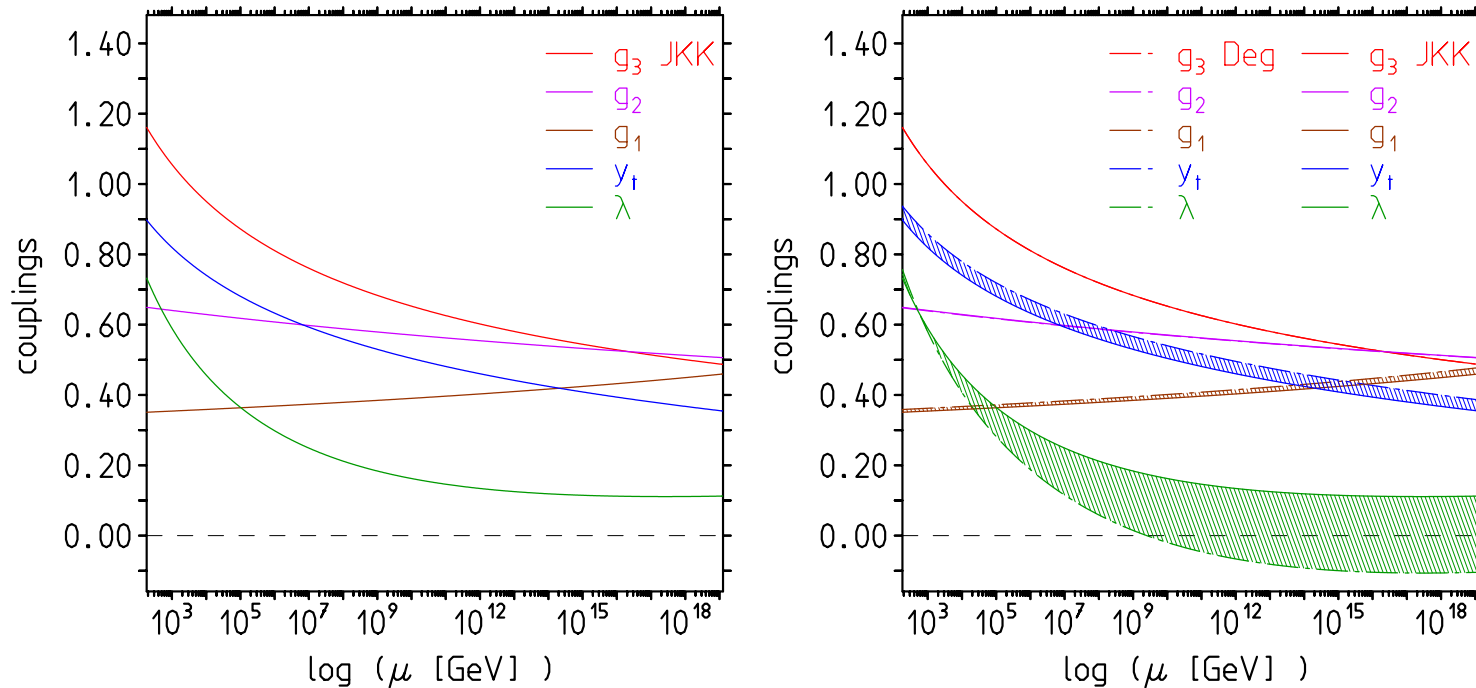
...

Something from Nothing?

It sounds rather strange that we have no firm idea about what makes up 74% of the universe.



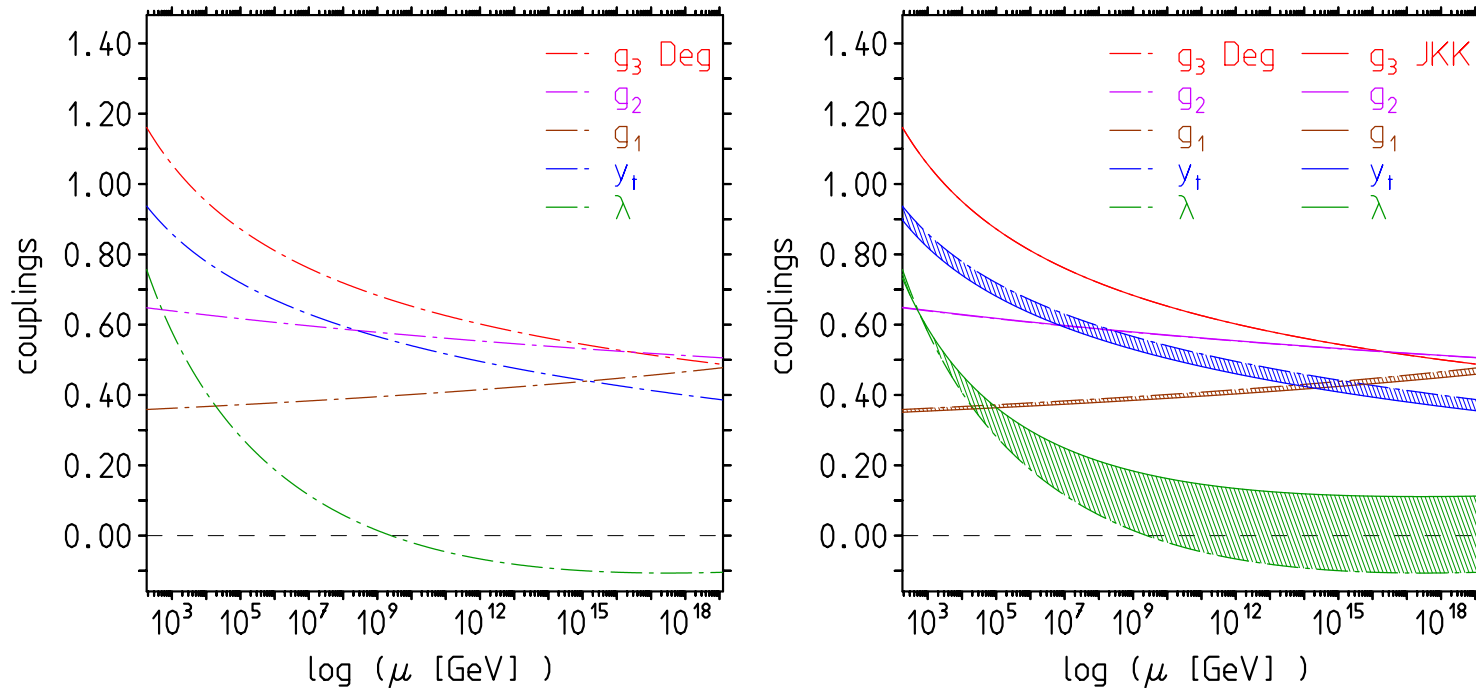
the Higgs at work →



F.J., Kalmykov, Kniehl, On-Shell vs $\overline{\text{MS}}$ parameter matching

- ❖ the big issue is the very delicate **conspiracy between SM couplings**: precision determination of parameters more important than ever \Rightarrow the challenge for LHC and ILC/FCC: precision values for λ , y_t and α_s , and for low energy hadron facilities: more precise hadronic cross sections to reduce hadronic uncertainties in $\alpha(M_Z)$ and $\alpha_2(M_Z)$

New gate to precision cosmology of the early universe!



Shaposnikov et al., Degrassi et al. matching

- ❖ the big issue is the very delicate **conspiracy between SM couplings**: precision determination of parameters more important than ever \Rightarrow the challenge for LHC and ILC/FCC: precision values for λ , y_t and α_s , and for low energy hadron facilities: more precise hadronic cross sections to reduce hadronic uncertainties in $\alpha(M_Z)$ and $\alpha_2(M_Z)$

New gate to precision cosmology of the early universe!

The SM's naturalness problems and fine-tuning “problems”

Issue broached by 't Hooft 1979 as a relationship between macroscopic phenomena which follow from microscopic physics (condensed matter inspired), i.e., bare versus renormalized quantities. Immediately the “hierarchy problem” has been dogmatized as a kind of fundamental principle.

Assume Planck scale $\Lambda_{\text{Pl}} \simeq 1.22 \times 10^{19} \text{ GeV}$ as a UV cutoff regularization:

□ the Higgs mass: [note bare parameters parametrize the true Lagrangian]

$$m_{\text{Higgs, bare}}^2 = m_{\text{Higgs, ren}}^2 + \delta m^2 ; \quad \delta m^2 = \frac{\Lambda_{\text{Pl}}^2}{(16\pi^2)} C(\mu)$$

coefficient typically $C = O(1)$. To keep the renormalized mass at the observed small value $m_{\text{ren}} = O(100 \text{ GeV})$, $\Rightarrow m_{\text{bare}}^2$ has to be tuned to compensate the huge term δm^2 : about 35 digits must be adjusted in order to get the observed value.

Hierarchy Problem!

□ the vacuum energy density $\langle V(\phi) \rangle$:

$$\rho_{\text{vac, bare}} = \rho_{\text{vac, ren}} + \delta\rho ; \quad \delta\rho = \frac{\Lambda_{\text{Pl}}^4}{(16\pi^2)^2} X(\mu)$$

SM predicts huge cosmological constant (CC) at Λ_{Pl} :

$$\rho_{\text{vac, bare}} \simeq V(0) + \Delta V(\phi) \sim 2.77 \Lambda_{\text{Pl}}^4 \sim (1.57 \times 10^{19} \text{ GeV})^4 \text{ vs. } \rho_{\text{vac}} = (0.002 \text{ eV})^4 \text{ today}$$

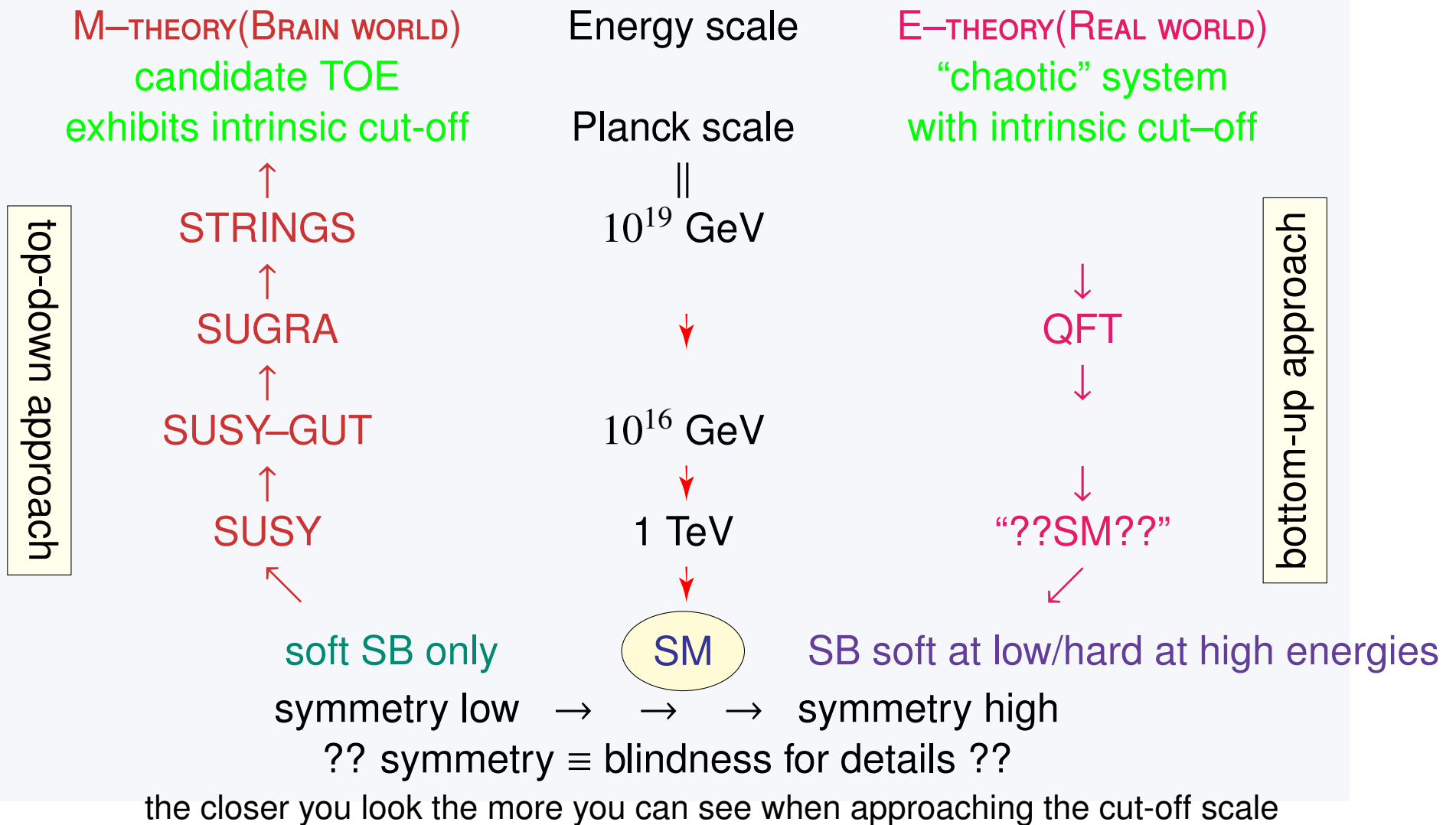
Cosmological Constant Problem!

Note: in symmetric phase the only trouble maker is the Higgs field!

Note: naive arguments do not take into account that quantities compared refer to **very different scales!** $m_{\text{Higgs, bare}}^2$ short distance, $m_{\text{Higgs, ren}}^2$ long distance observables. Also: Λ as a regulator nobody forces you to take it to be Λ_{Pl} .

Need: **UV-completion** of SM: prototype **lattice SM** as true(r) system

Paths to Physics at the Planck Scale



References:

“The Standard model as a low-energy effective theory: what is triggering the Higgs mechanism?,”
Acta Phys. Polon. B **45** (2014) 1167 [arXiv:1304.7813 [hep-ph]].

“The hierarchy problem of the electroweak Standard Model revisited,”
arXiv:1305.6652 [hep-ph] also [arXiv:1503.00809](#) [hep-ph]

“Higgs inflation and the cosmological constant,”
Acta Phys. Polon. B **45** (2014) 1215 [arXiv:1402.3738 [hep-ph]].

Krakow/Durham Lectures:

<http://www-com.physik.hu-berlin.de/~fjeger/SMcosmology.html>

see also: “The Vector Boson and Graviton Propagators in the Presence of Multipole Forces,”
Helv. Phys. Acta **51** (1978) 783.

Summary part II:

- with Higgs discovery: SM essentially complete, Higgs mass $M_H \simeq 126 \text{ GeV}$ very special for Higgs vacuum stability
- SM couplings are energy dependent, all but g' decrease towards M_{Pl} , perturbation theory works well up to Planck scale.
- SM Higgs potential likely remains stable up to M_{Pl} (i.e. $\lambda(\mu) > 0$ for all $\mu < M_{\text{Pl}}$)
- bare parameters are the true parameters at very high energy approaching M_{Pl} , relevant for early universe
- bare parameters are calculable in SM as needed for early cosmology
- cutoff enhanced quantities: effective bare Higgs mass (quadratic $\propto \Lambda_{\text{Pl}}^2$) as well as dark energy (quartic $\propto \Lambda_{\text{Pl}}^4$)
 \Rightarrow provide inflation condition $V(\phi) \gg \frac{1}{2}\dot{\phi}^2$
- SM originally (at very high energies) in symmetric phase, all particles massless except for the four very heavy Higgses
- both the Higgs mass as well as the dark energy exhibit matching points where bare and renormalized values coincide, separates low energy form bare Planck regime responsible for inflation

- need trans-Planckian initial Higgs field $\phi_i = \phi(t_{\text{Pl}}) \sim 5 M_{\text{Pl}}$ in order to get sufficient inflation $N \gtrsim 60$
- trans-Planckian fields do no harm: fast exponential decay of Higgs field
- after inflation in reheating phase: very heavy Higgses mainly decay into top–antitop pairs, which latter (after the EW phase transition) decay into normal baryonic matter
- except for ϕ_i all properties known: inflation and reheating are SM predictions within uncertainties of SM initial parameters and RG evolution approximations
- EW phase transition in this scenario happens at much higher energy than anticipated so far and close by natural Baryon number violating dimension 6 operators likely trigger baryogenesis.
- SM inflation requires very precise input parameters and appropriate higher order corrections (precise knowledge of the SM itself) Presently: $\overline{\text{MS}}$ RG to 3 loops (massless), matching conditions leading 2 loops (need full massive SM calculations, yet incomplete)
- **no SM hierarchy problem**: relation between renormalized and bare Higgs boson mass is a SM prediction, nothing to worry! No CC problem either!

SM inflation vs **added** inflation scenarios

LEESM scenario is easy to rule out:

- ① find any type of New Physics (NP) as motivated by the naive **hierarchy problem** argument. These are most SM extension scenarios (SUSY, Extra Dimensions, Little/st Higgs, ETC and what else), i.e. any physics affecting substantially **the quadratic and quartic “divergences”**.
- ② find any type of “**new heavy states with substantial couplings to SM spectrum**” like 4th family, GUT, “light” heavy (far below M_{Pl}) singlet Majorana neutrino etc, i.e. anything affecting the g', g, g_s, y_t and λ **SM running coupling pattern**.
- ③ confront precise SM inflation predictions with **inflation pattern itself**: large enough N_e , $w \approx -1$, spectral indices n_S , n_T , Gaussianity etc

The LEESM scenario is natural as it **predicts** a bulk of properties, which usually are assumed/ imposed as **basic principles**. All these are **emergent** properties!

Predicted as long range phenomenon:

- ❖ QFT structure,
- ❖ renormalizability and requirements needed for it:
 - ➡ non-Abelian gauge structure,
 - ➡ chiral symmetry,
 - ➡ anomaly cancellation and fermion family structure
 - ➡ the existence of the Higgs particle! (renormalizability)
- ❖ space-time dimensionality $D = 4$, no renormalizable non-trivial QFT in $D > 4$
- ❖ rotation invariance and Lorentz invariance (pseudo-rotations)
- ❖ analyticity, effective unitarity etc

All result are checkable through real calculations (mostly existent).

SM inflation is based on SM predictions, except for the Higgs field value ϕ_0 , which is the only quantity relevant for inflation, which is not related to an observable low energy quantity.

All other inflation scenarios are set up “by hand”: the form of the potential as well as all parameters are **tuned** to reproduce the observed inflation pattern.

Example **Minkowski-Zee-Shaposhnikov et al** so called **non-minimal SM inflation**

- ❶ Change Einstein Gravity by adding $G_{\mu\nu} \rightarrow G_{\mu\nu} + \xi (H^\dagger H) R$ together with renormalized low energy SM $T_{\mu\nu}$ (no relevant operator enhancement)
- ❷ Choose ξ large enough to get sufficient inflation, need $\xi \sim 10^4$, entire inflation pattern essentially depends on ξ only (inflation “by hand”)
- ❸ assume quadratic and quartic SM divergences are absent (argued by dimensional regularization (DR) and $\overline{\text{MS}}$ renormalization)
- ❹ assume SM to be in broken phase at Planck scale, which looks unnatural.
Note: SSB is a low energy phenomenon, which assumes the symmetry to be restored at the short distance scale!)

All but convincing!

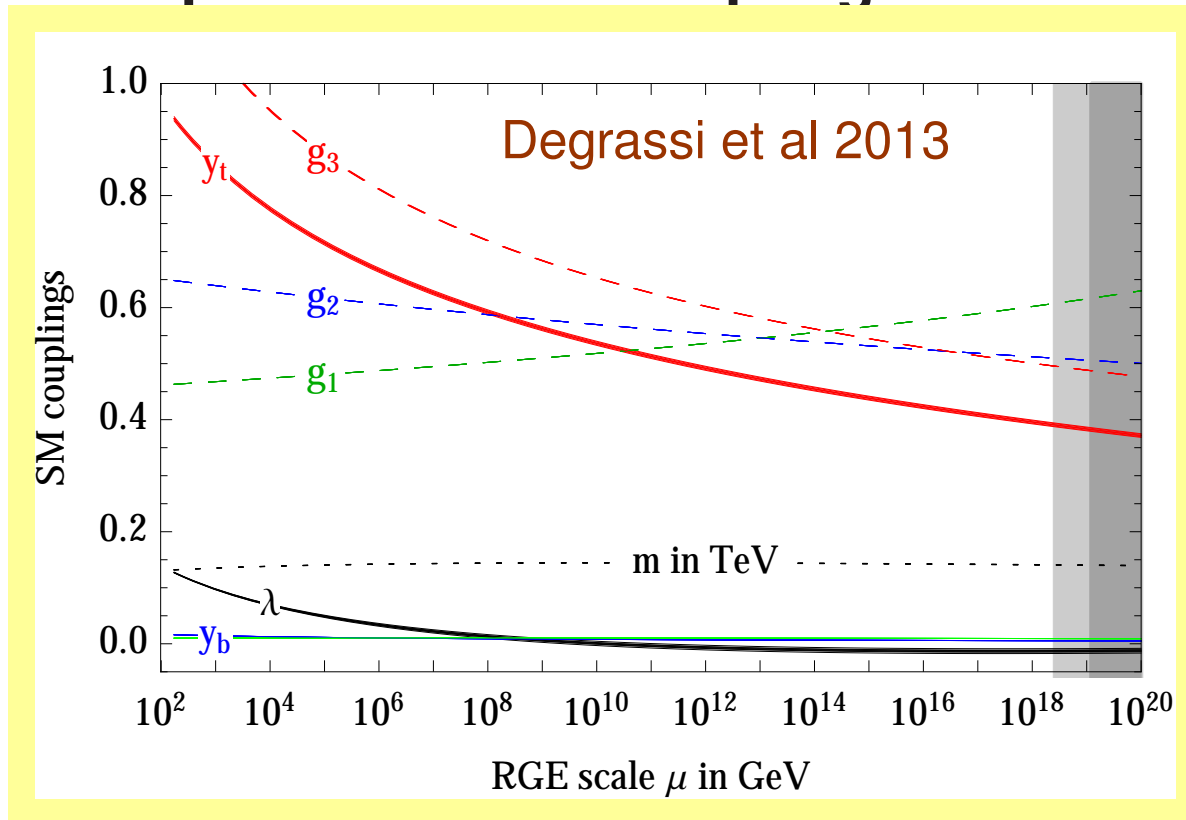
Note: DR and $\overline{\text{MS}}$ renormalization are possible in perturbation theory only. There is no corresponding non-perturbative formulation (simulation on a lattice) or measuring prescription (experimental procedure). It is based on a finite part prescription (singularities nullified by hand), which can only be used to calculate quantities which do not exhibit any singularities at the end. The hierarchy problem cannot be addressed in the $\overline{\text{MS}}$ scheme.

- My scenario: take Einstein Gravity serious (geometry, equivalence principle etc unaffected) together with true SM energy-momentum tensor, i.e. as effective at given scale, beyond $X_{\text{ren}} = X_{\text{bare}}$ matching point: true=bare as relevant near Planck energies.

➡ Additional remarks ⬅

- Test of tricky conspiracy between SM couplings the new challenge
- Very delicate on initial values as we run over 16 orders of magnitude from the EW 250 GeV scale up to the **Planck scale!**
- Running couplings likely have dramatic impact on cosmology! The existence of the world in question?
- LHC and ILC will dramatically improve on Higgs self-coupling λ (Higgs factory) as well as on top Yukawa y_t ($t\bar{t}$ factory)
- for running α_{em} and $\sin^2\Theta_{\text{eff}} \Leftrightarrow g_1$ and g_2 need more information from low energy hadron production facilities, improving QCD predictions and EW radiative corrections! Lattice QCD will play key role for sure.

Comparison of SM coupling evolution



Renormalization of the SM gauge couplings $g_1 = \sqrt{5/3}g_Y$, g_2 , g_3 , of the top, bottom and τ couplings (y_t , y_b , y_τ), of the Higgs quartic coupling λ and of the Higgs mass parameter m . We include two-loop thresholds at the weak scale and three-loop RG equations. The thickness indicates the $\pm 1\sigma$ uncertainties.

Comparison of $\overline{\text{MS}}$ parameters at various scales: Running couplings for $M_H = 126 \text{ GeV}$ and $\mu_0 \simeq 1.4 \times 10^{16} \text{ GeV}$.

coupling \ scale	my findings				Degrassi et al. 2013	
	M_Z	M_t	μ_0	M_{Pl}	M_t	M_{Pl}
g_3	1.2200	1.1644	0.5271	0.4886	1.1644	0.4873
g_2	0.6530	0.6496	0.5249	0.5068	0.6483	0.5057
g_1	0.3497	0.3509	0.4333	0.4589	0.3587	0.4777
y_t	0.9347	0.9002	0.3872	0.3510	0.9399	0.3823
$\sqrt{\lambda}$	0.8983	0.8586	0.3732	0.3749	0.8733	i 0.1131
λ	0.8070	0.7373	0.1393	0.1405	0.7626	- 0.0128

Most groups find just unstable vacuum at about $\mu \sim 10^9 \text{ GeV}$! [not independent, same $\overline{\text{MS}}$ input]

Note: $\lambda = 0$ is an essential singularity and the theory cannot be extended beyond a possible zero of λ : remind $v = \sqrt{6m^2/\lambda}$!!! i.e. $v(\lambda) \rightarrow \infty$ as $\lambda \rightarrow 0$
besides the Higgs mass $m_H = \sqrt{2} m$ all masses $m_i \propto g_i v \rightarrow \infty$ different cosmology

What about the hierarchy problem?

□ In the Higgs phase:

There is no hierarchy problem in the SM!

In the broken phase, characterized by the non-vanishing Higgs field vacuum expectation value (VEV) $v(\mu)$, all the masses are determined by the well known mass-coupling relations

$$\begin{aligned} m_W^2(\mu) &= \frac{1}{4} g^2(\mu) v^2(\mu) ; & m_Z^2(\mu) &= \frac{1}{4} (g^2(\mu) + g'^2(\mu)) v^2(\mu) ; \\ m_f^2(\mu) &= \frac{1}{2} y_f^2(\mu) v^2(\mu) ; & m_H^2(\mu) &= \frac{1}{3} \lambda(\mu) v^2(\mu) . \end{aligned}$$

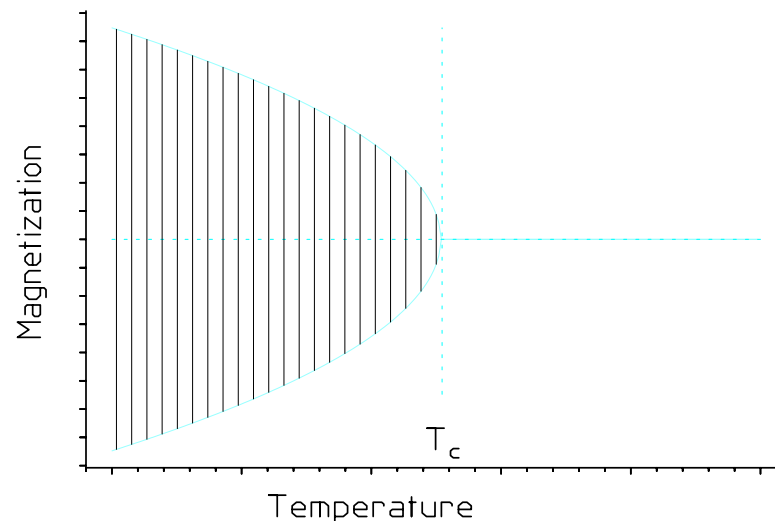
□ Higgs mass cannot be much heavier than the other heavier particles!

□ Extreme point of view: all particles have masses $O(M_{\text{Pl}})$ i.e. $v = O(M_{\text{Pl}})$.
This would mean the symmetry is **not** recovered at the high scale,

notion of SSB obsolete! Of course this makes no sense.

- Higgs VEV v is an **order parameter** resulting from long range collective behavior,
can be as small as we like.

Prototype: magnetization in a ferromagnetic spin system



$M = M(T)$ and actually $M(T) \equiv 0$ for $T > T_c$ furthermore $M(T) \rightarrow 0$ as $T \rightarrow T_c$

● $v/M_{\text{Pl}} \ll 1$ just means we are close to a 2nd order phase transition point.

□ In the symmetric phase at very high energy we see the bare system:

the Higgs field is a **collective field** exhibiting an effective mass generated by radiative effects

$$m_{\text{bare}}^2 \approx \delta m^2 \text{ at } M_{\text{Pl}}$$

eliminates fine-tuning problem at all scales!

Many example in condensed matter systems.

In my view the hierarchy problem is a pseudo problem!

★ What rules the β -functions:

Naively:

□ $U(1)_Y$ screening (IR free), $SU(2)_L$, $SU(3)_c$ antiscreening (UV free) [asymptotic freedom (AF)]

Right – as expected

□ Yukawa and Higgs: screening (IR free, like QED)

Wrong!!! – transmutation from IR free to AF

At the Z boson mass scale: $g_1 \simeq 0.350$, $g_2 \simeq 0.653$, $g_3 \simeq 1.220$, $y_t \simeq 0.935$ and $\lambda \simeq 0.796$

Leading (one-loop) β -functions at $\mu = M_Z$: [$c = \frac{1}{16\pi^2}$]

❖ gauge couplings:

$$\beta_1 = \frac{41}{6} g_1^3 c \simeq 0.00185 ; \quad \beta_2 = -\frac{19}{6} g_2^2 c \simeq -0.00558 ; \quad \beta_3 = -7 g_3^3 c \simeq -0.08045 ,$$

❖ top Yukawa coupling:

$$\begin{aligned}\beta_{y_t} &= \left(\frac{9}{2} y_t^3 - \frac{17}{12} g_1^2 y_t - \frac{9}{4} g_2^2 y_t - 8 g_3^2 y_t \right) c \\ &\simeq 0.02328 - 0.00103 - 0.00568 - 0.07046 \\ &\simeq -0.05389\end{aligned}$$

not only depends on y_t , but also on mixed terms with the gauge couplings g' , g and g_3 which have a negative sign.

In fact the QCD correction is the leading contribution and determines the behavior. Notice the critical balance between the dominant strong and the top Yukawa couplings: QCD dominance requires $g_3 > \frac{3}{4} y_t$ in the gaugeless limit.

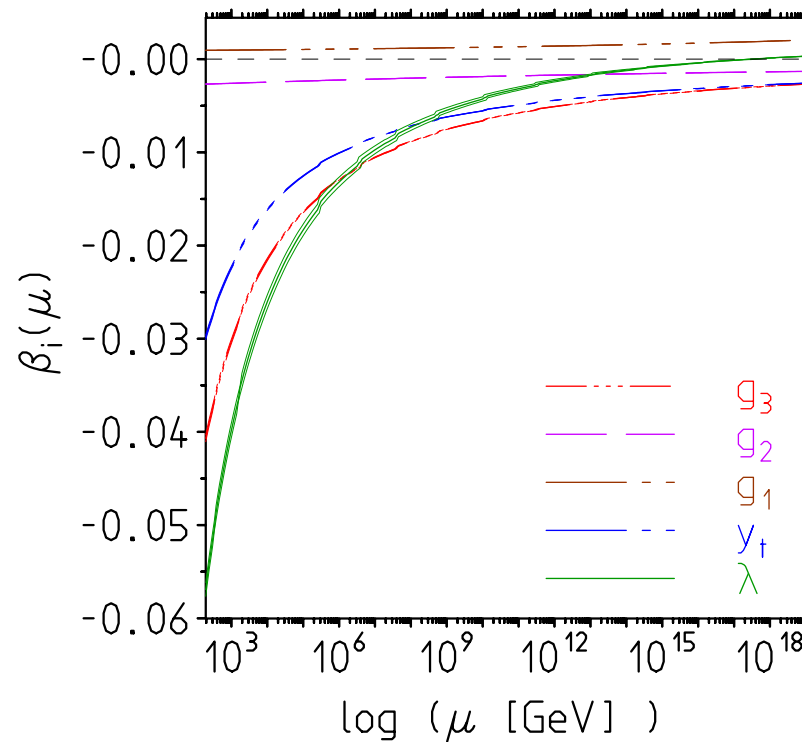
❖ the Higgs self-coupling

$$\begin{aligned}\beta_\lambda &= \left(4 \lambda^2 - 3 g_1^2 \lambda - 9 \lambda g_2^2 + 12 y_t^2 \lambda + \frac{9}{4} g_1^4 + \frac{9}{2} g_1^2 g_2^2 + \frac{27}{4} g_2^4 - 36 y_t^4 \right) c \\ &\simeq 0.01606 - 0.00185 - 0.01935 + 0.05287 + 0.00021 + 0.00149 + 0.00777 - 0.17407 \\ &\simeq -0.11687\end{aligned}$$

dominated by y_t contribution and not by λ coupling itself. At leading order it is not subject to QCD corrections. Here, the y_t dominance condition reads $\lambda < \frac{3(\sqrt{5}-1)}{2} y_t^2$ in the gaugeless limit.

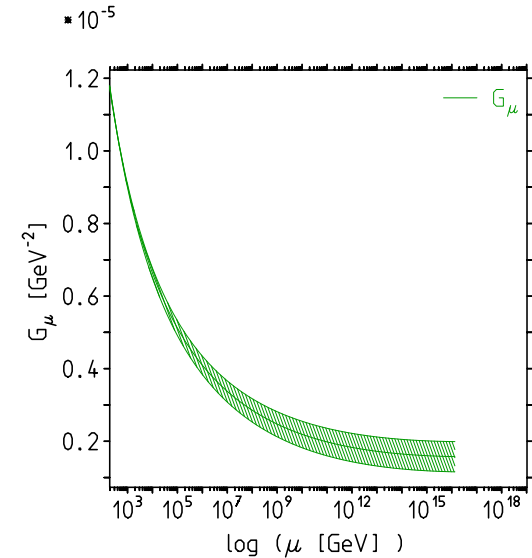
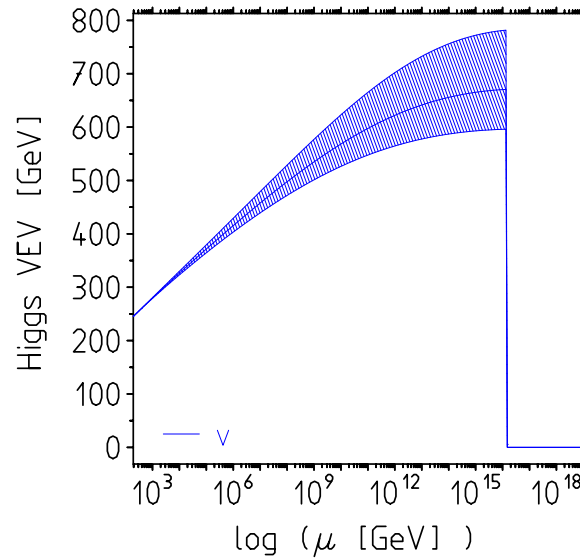
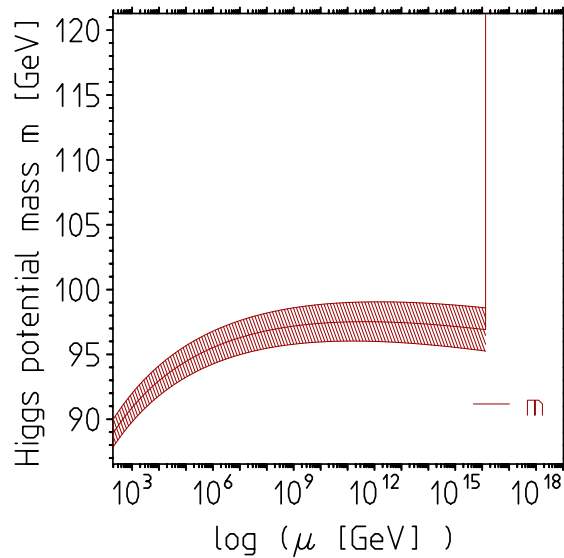
□ running top Yukawa QCD takes over: IR free \Rightarrow UV free

□ running Higgs self-coupling top Yukawa takes over: IR free \Rightarrow UV free



Including all known RG coefficients (EW up incl 3-loop, QCD up incl 4-loop)

- ➡ except from β_λ , which exhibits a zero at about $\mu_\lambda \sim 10^{17} \text{ GeV}$, all other β -functions do not exhibit a zero in the range from $\mu = M_Z$ to $\mu = M_{\text{Pl}}$.
- ➡ so apart from the $U(1)_Y$ coupling g_1 , which increases only moderately, all other couplings decrease and perturbation theory is in good condition.
- ➡ at $\mu = M_{\text{Pl}}$ gauge couplings are all close to $g_i \sim 0.5$, while $y_t \sim 0.35$ and $\sqrt{\lambda} \sim 0.36$.
- effective masses moderately increase (largest for m_Z by factor 2.8): scale like $m(\kappa)/\kappa$ as $\kappa = \mu'/\mu \rightarrow \infty$,
i.e. mass effect get irrelevant as expected at high energies.



Non-zero dimensional $\overline{\text{MS}}$ running parameters: m , $v = \sqrt{6/\lambda} m$ and $G_F = 1/(\sqrt{2} v^2)$. Error bands include SM parameter uncertainties and a Higgs mass range $125.5 \pm 1.5 \text{ GeV}$ which essentially determines the widths of the bands. Note that v increases by a factor about 2.5 before it jumps to zero at the transition point.