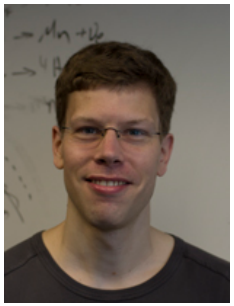
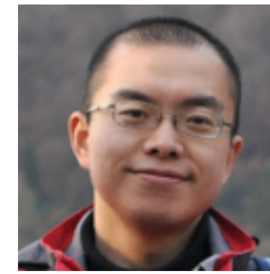
Extragalactic Cosmic Rays above the Iron Knee



Markus Ahlers

Dan Hooper

Felix Aharonian



Ruoyu Liu

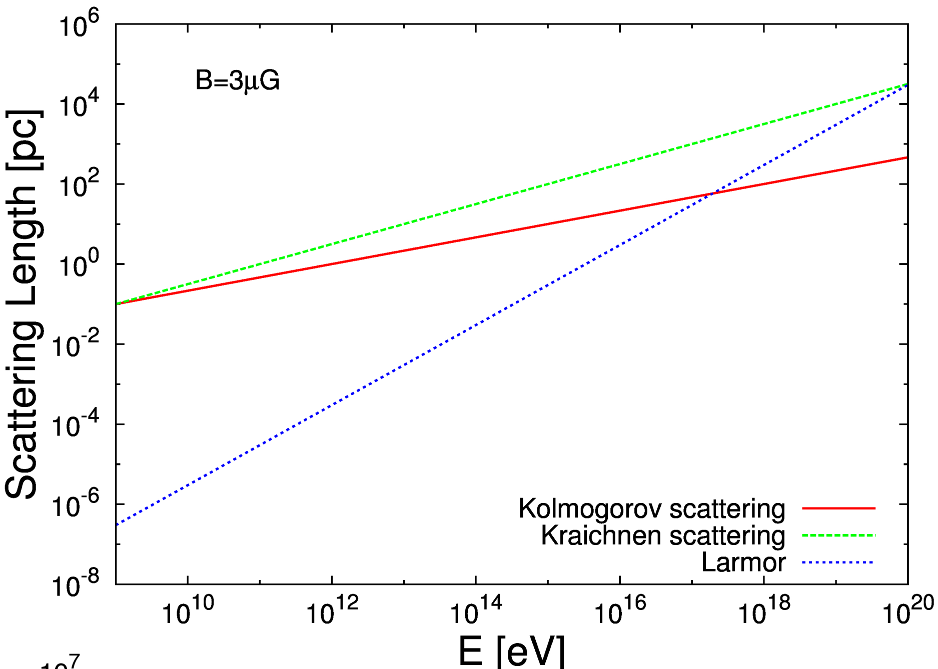
Based on:

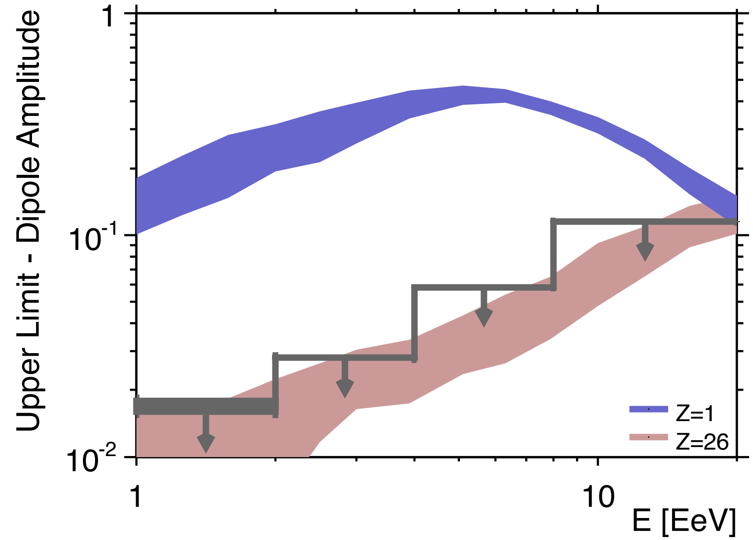
“Indications of Negative Evolution for the Sources of the Highest Energy Cosmic Rays”, **Phys.Rev. D92 (2015) 6, 063011** [astro-ph/**1505.06090]**

“Evidence for a Local "Fog" of Sub-Ankle UHECR”, astro-ph/**1603.03223**

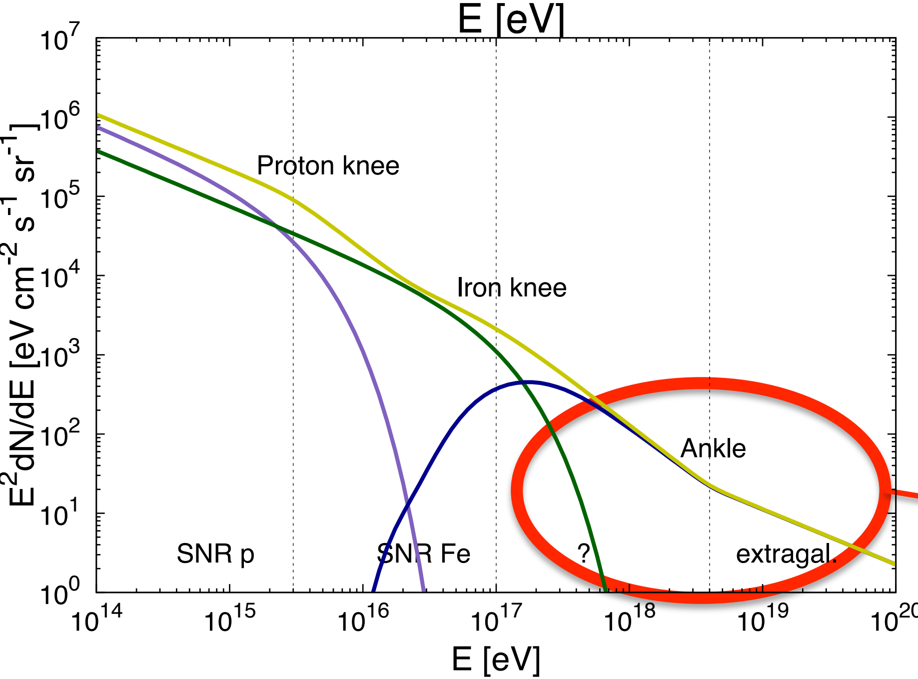
Andrew Taylor

Separate Probes of the Transition Energy



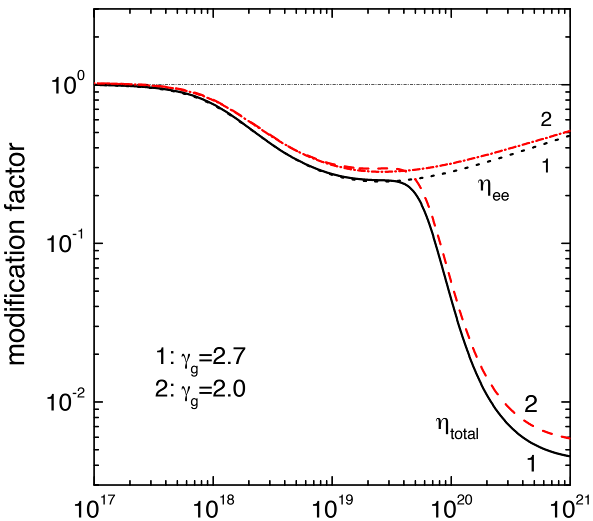


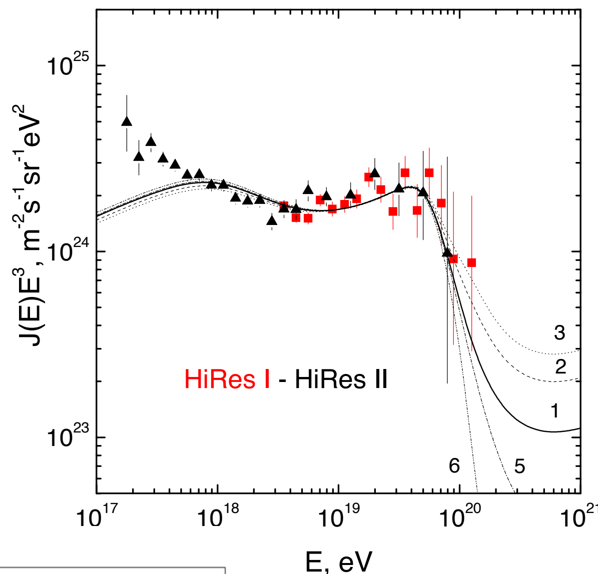
Anisotropy constraint:
Giacinti et al.- astro-ph/**1112.5599**
Pierre Auger Collab.- astro-ph/**1212.3083**



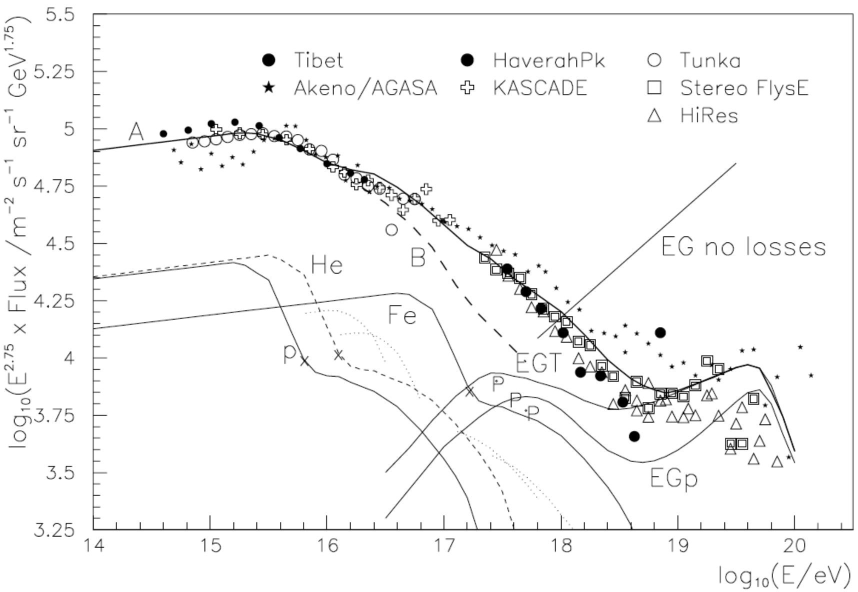
Where does the energy flux go? (or do these CR propagate unimpeded?)
....energy is conserved after all

Historical Debate about the Nature of the Ankle Feature





From Berezinsky et al. (2006)
astro-ph/**0204357**
“Dip Model”



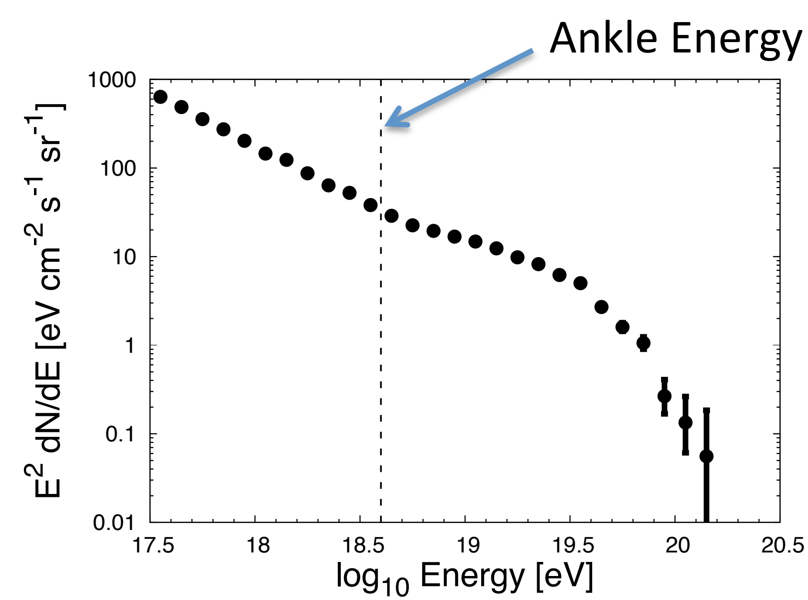
From Hillas et al. (2004)
SKA Conf. Proceedings

What’s changed since then? In which direction do the scales now tip towards?

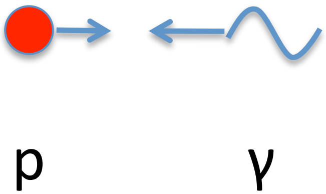
Why Consider UHECR to Understand the Galactic/Extragalactic Transition?

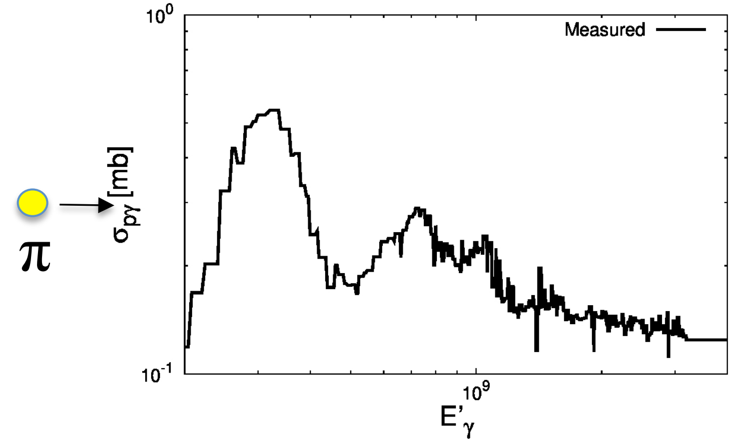
- Since the ankle feature appears at an energy of ~$10^{18.6}$ eV, a new extragalactic source class is presumed to begin to dominate here (in the first instance)

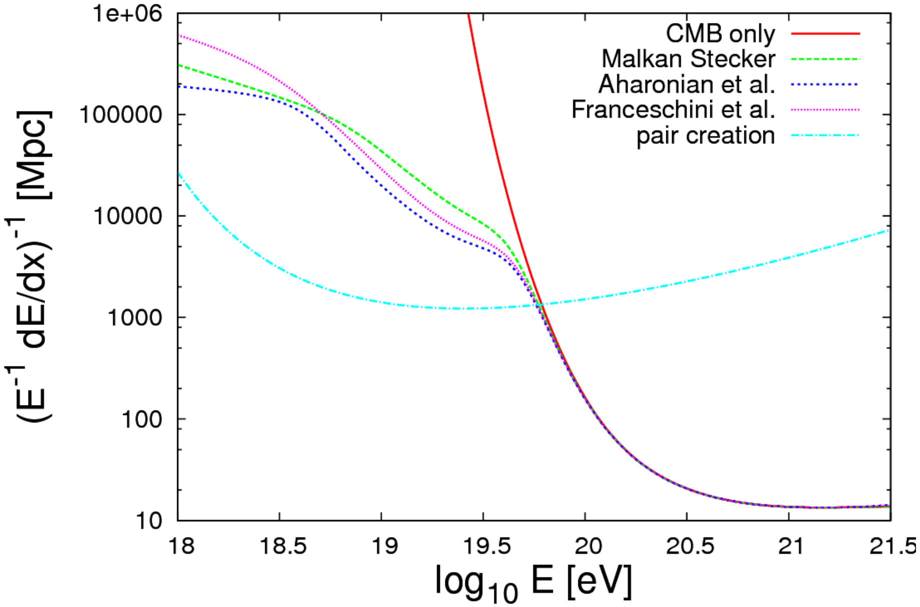
- Information obtained from investigations into the UHECR sources may provide new insights into Galactic-Extragalactic transition energy

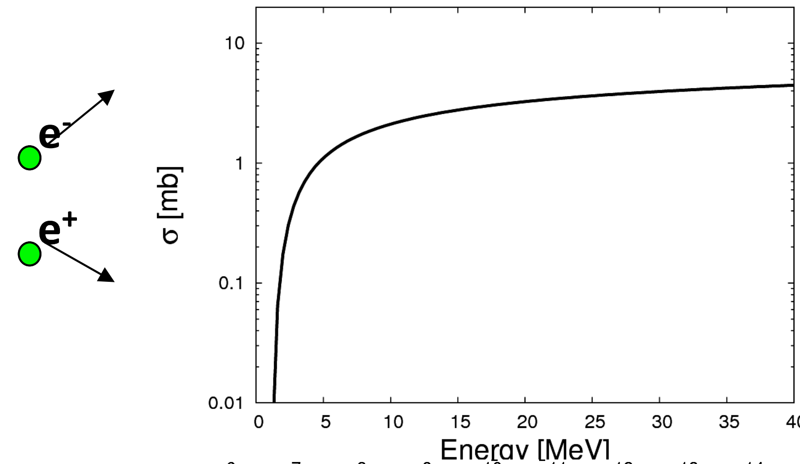


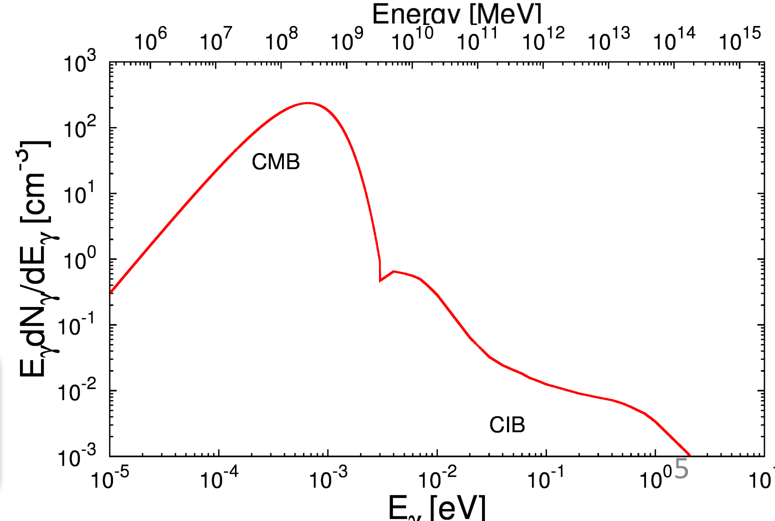
Cosmic Ray Proton Interactions





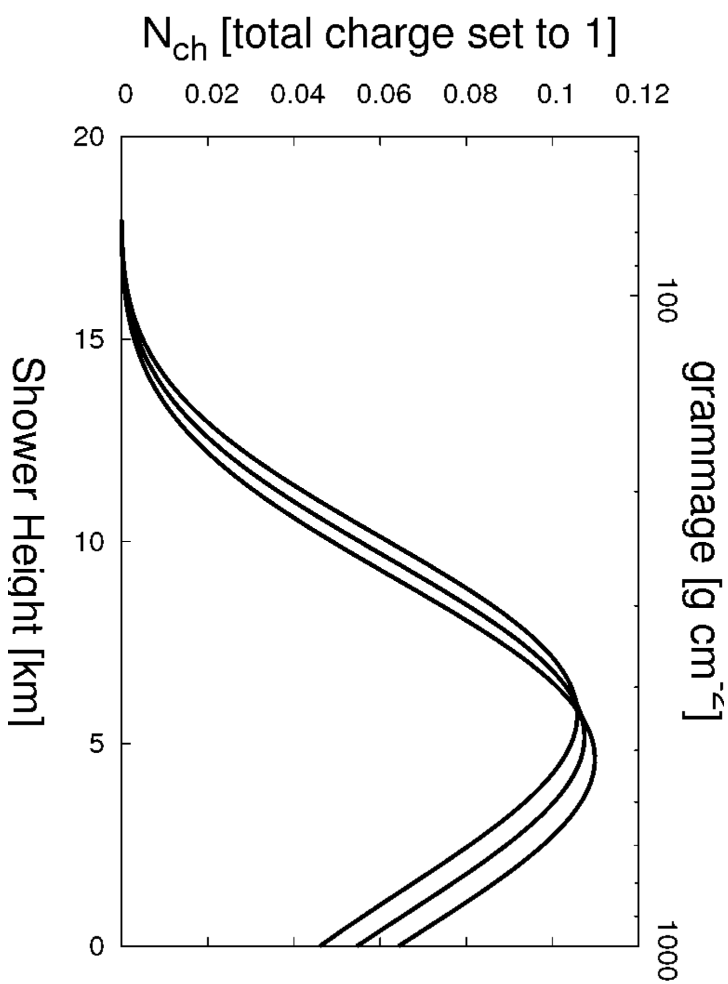


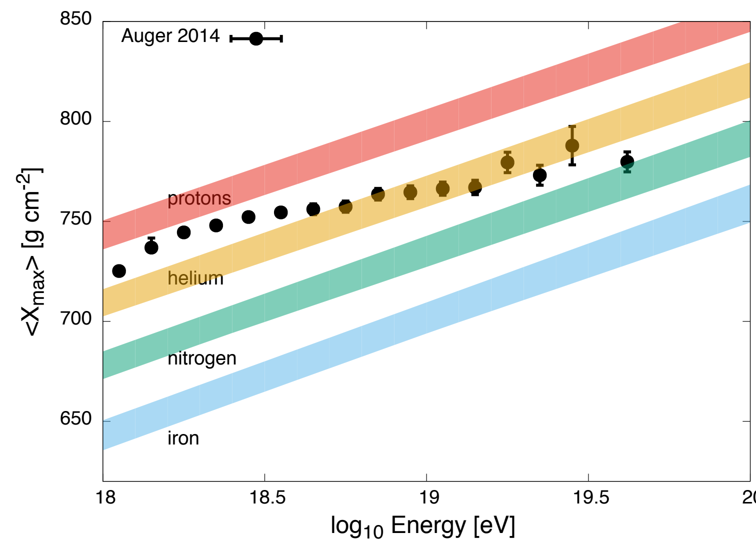


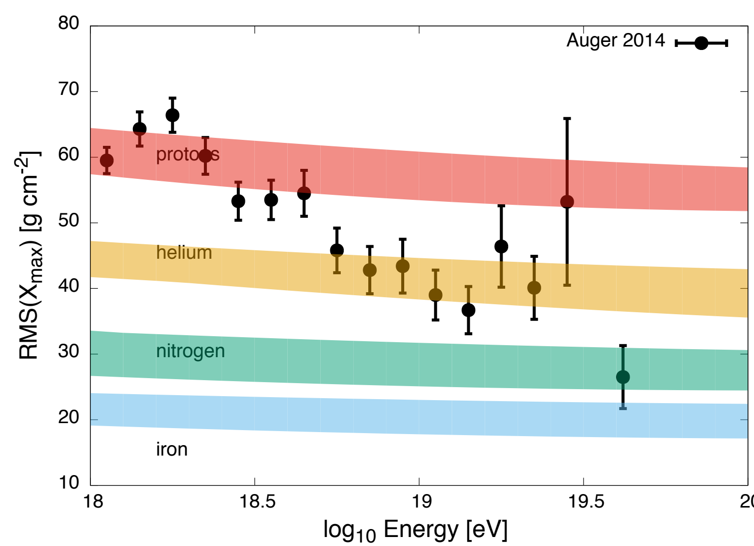


$$R = \frac{2m_p^2}{E_p^2}\int \frac{1}{\epsilon^2}\frac{dN_\gamma}{d\epsilon}\int_0^{4E_p\epsilon/m_p} k_{p\gamma}\epsilon'\sigma_{p\gamma}(E_p,\epsilon')d\epsilon'$$

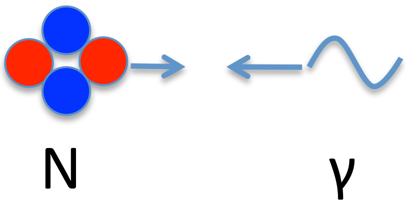
Composition- Consider Nuclei?

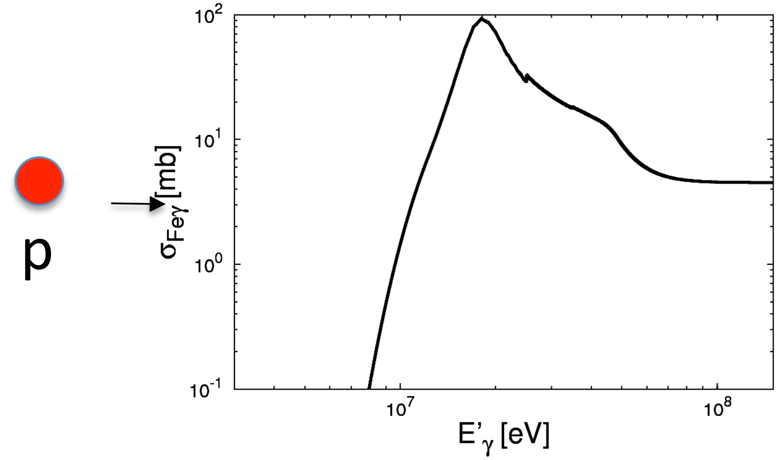


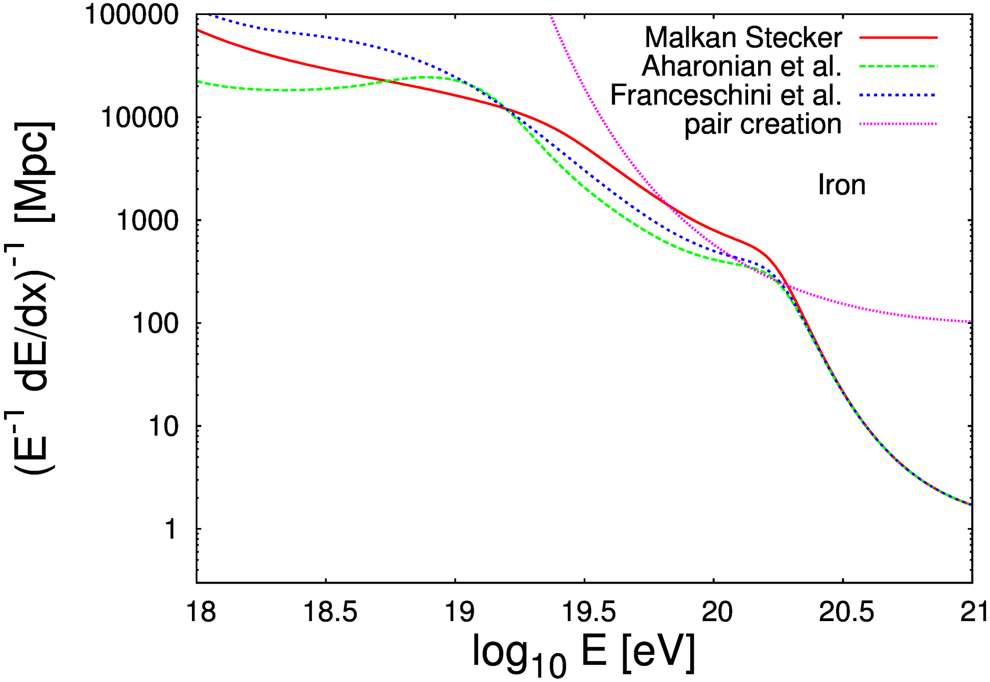


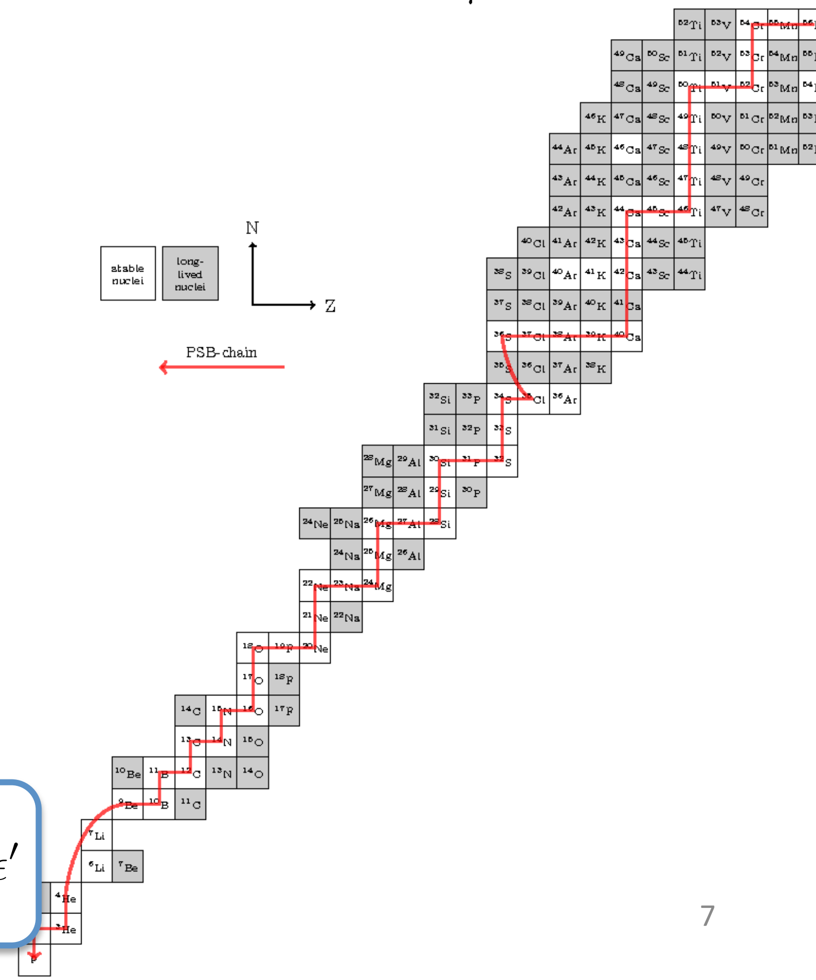


Cosmic Ray Nuclei Interactions







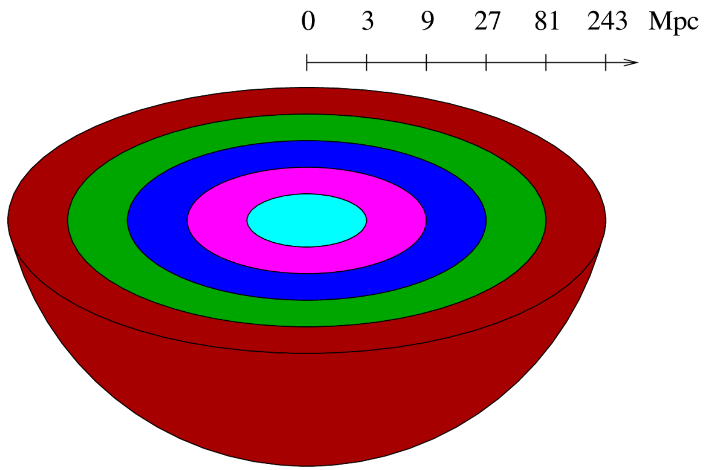


$$R = \frac{2m_N^2}{E_N^2}\int \frac{1}{\epsilon^2}\frac{dN_\gamma}{d\epsilon}\int_0^{4E_N\epsilon/m_N} k_{N\gamma}\epsilon'\sigma_{N\gamma}(E_N,\epsilon')d\epsilon'$$

Assumptions on Source Population

$$\frac{dN}{dV_C} \propto (1+z)^n$$

$$z < z_{max}$$



$$n = -6,\ -3,\ 0,\ 3$$

$$\frac{dN}{dE} \propto E^{-\alpha} \exp[-E/E_{Z,max}]$$

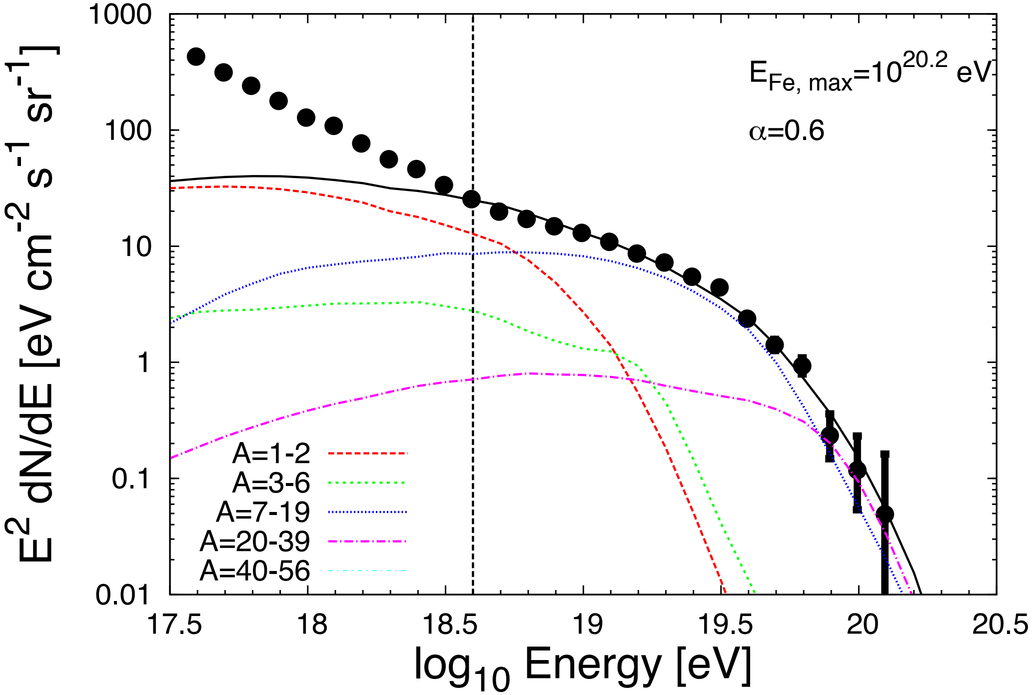
$$E_{Z,max} = (Z/26) \times E_{Fe,max}$$

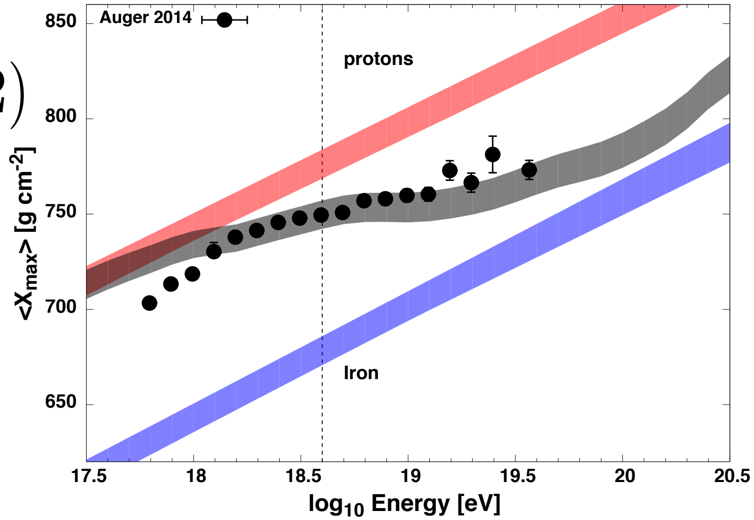
Note- magnetic field horizon effects are neglected in the following. This amounts to assuming: $d_s < (ct_H \lambda_{scat})^{1/2}$
ie. the source distribution may be approximated to be spatially continuous (also note, presence of t_H term comes from temporally continuous assumption)

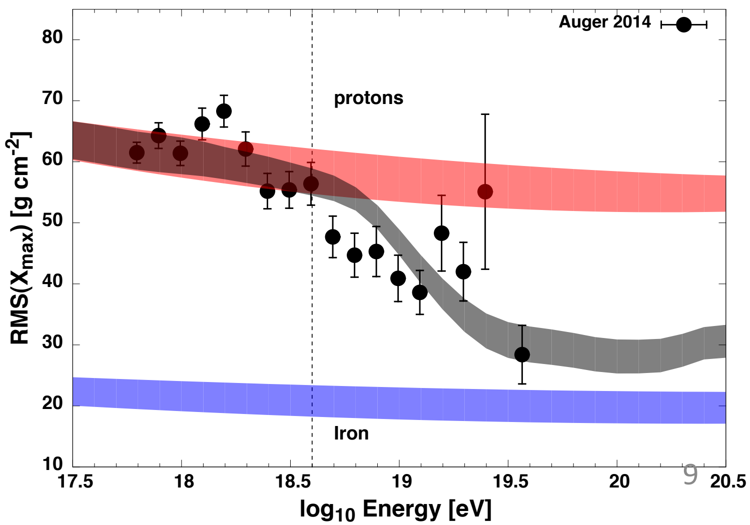
MCMC Likelihood Scan: Spectral + Composition Fits

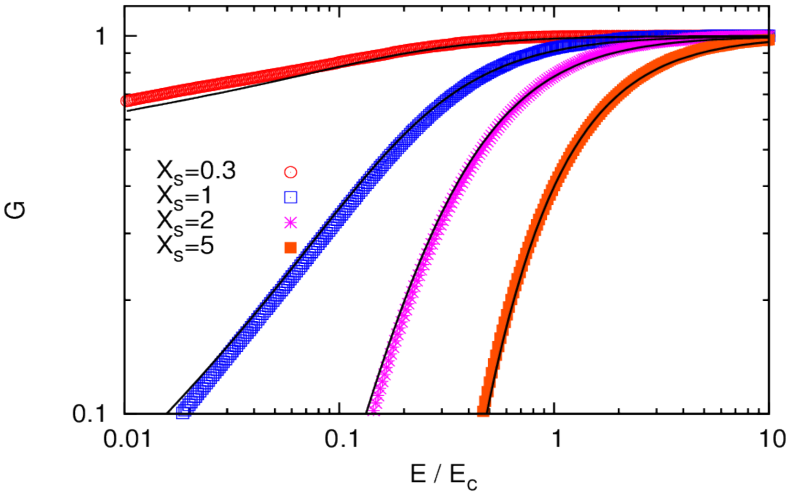
$$\mathbf{L}(\mathbf{f_p}, \mathbf{f_{He}}, \mathbf{f_N}, \mathbf{f_{Si}}, \mathbf{E_{max}}, \alpha) \propto \exp(-\chi^2/2)$$

n=3 evolution result









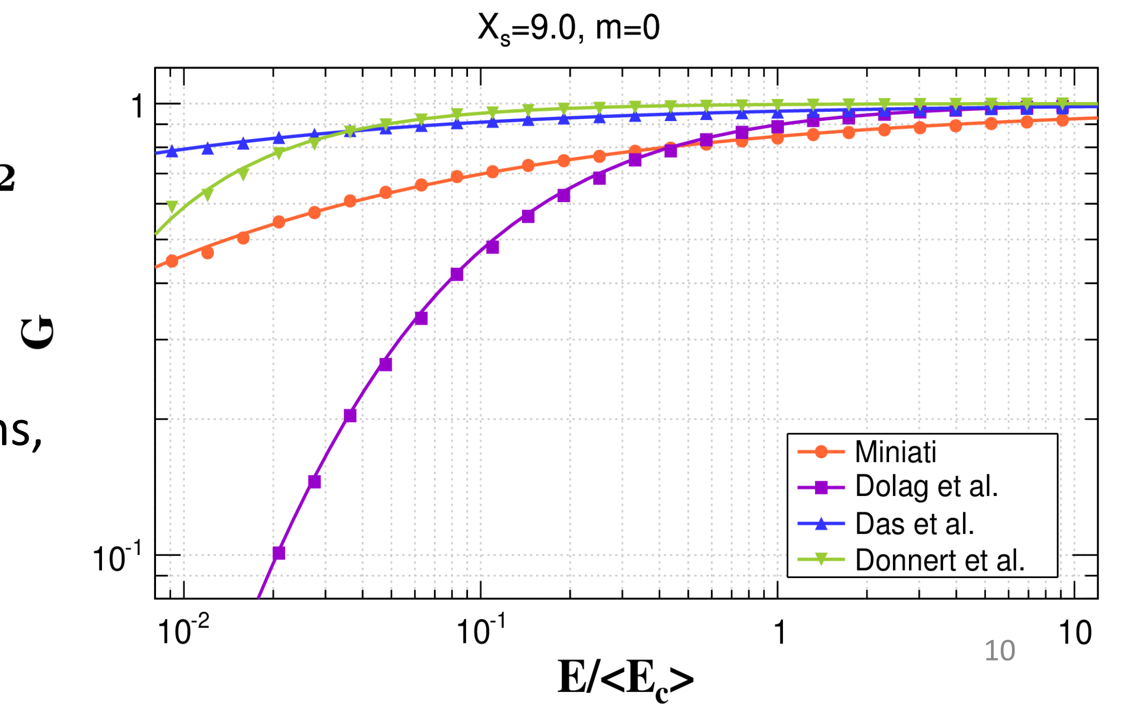
Hard Spectra Problem

Magnetic horizon suppression suggested to resolve "hardness" issue, Mollerach et al. astro-ph/1305.6519

	Miniati	Dolag *et al.*	Das *et al.*	Donnert *et al.*
$\langle B \rangle$ [G]	1.8×10^{-8}	5.5×10^{-11}	1.2×10^{-9}	6.3×10^{-11}
B_{rms} [G]	1.7×10^{-7}	1.5×10^{-8}	5.7×10^{-8}	1.7×10^{-8}

$$X_s = \frac{d_s}{(c t_H l_c)^{1/2}}$$
$$= 0.1 \left(\frac{d_s}{10 \text{ Mpc}}\right)\left(\frac{1 \text{ Mpc}}{l_c}\right)^{1/2}$$

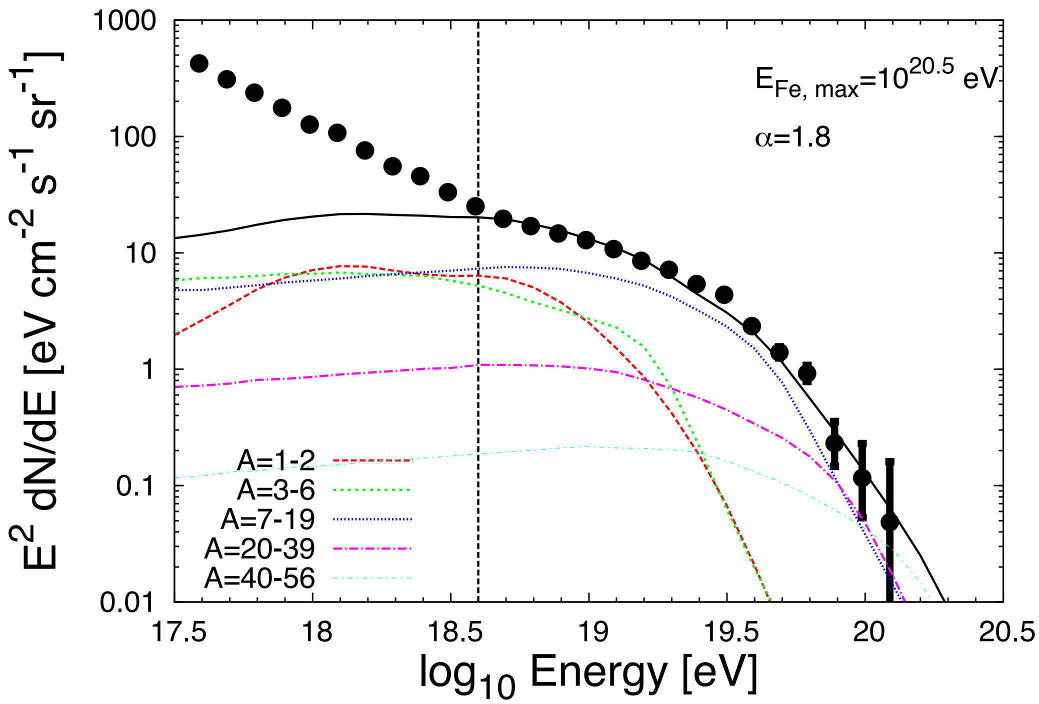
"Realistic" field structures/strengths, however, don't provide sufficient suppression, Alves Batista et al. astro-ph/1407.6150

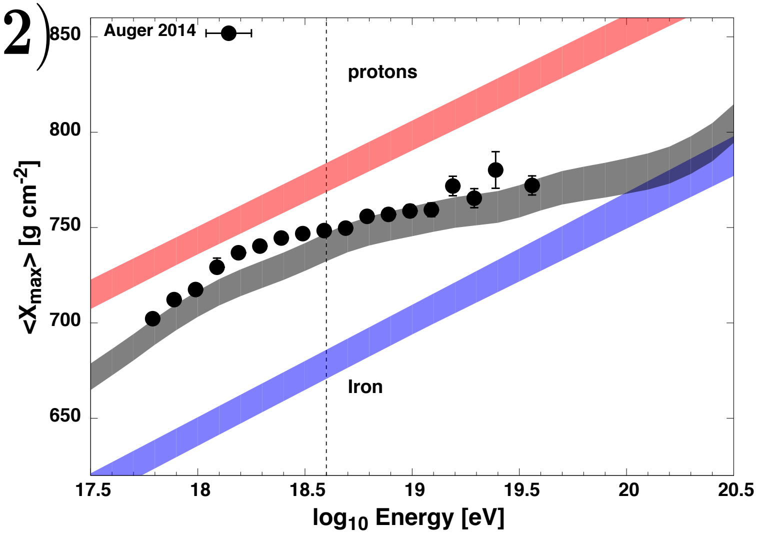


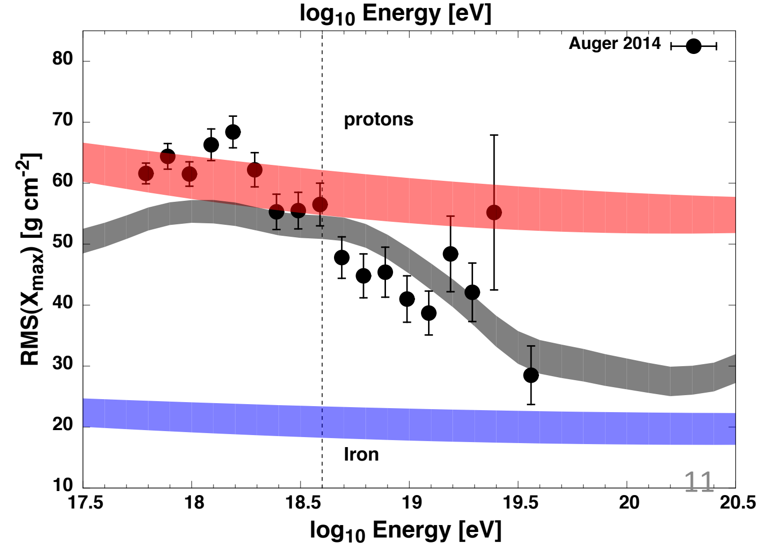
MCMC Likelihood Scan: "Soft" Spectra Solutions

$$\mathbf{L}(\mathbf{f_p}, \mathbf{f_{He}}, \mathbf{f_N}, \mathbf{f_{Si}}, \mathbf{E_{max}}, \alpha) \propto \exp(-\chi^2/2)$$

n=-6 evolution result







MCMC Results Table

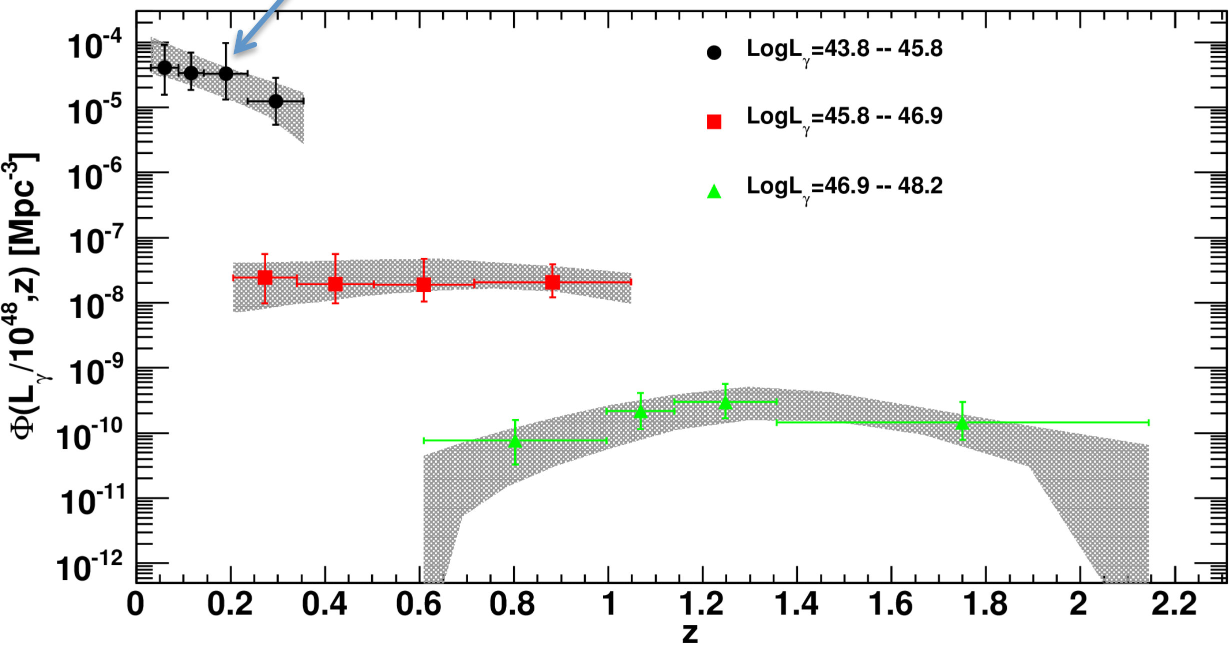
Parameter	$n=-6$		$n=-3$		$n=0$		$n=3$	
	Best-fit Value	Posterior Mean & Standard Deviation	Best-fit Value	Posterior Mean & Standard Deviation	Best-fit Value	Posterior Mean & Standard Deviation	Best-fit Value	Posterior Mean & Standard Deviation
f_p	0.03	0.14 ± 0.12	0.08	0.15 ± 0.13	0.17	0.17 ± 0.16	0.19	0.20 ± 0.16
f_{He}	0.50	0.21 ± 0.17	0.42	0.17 ± 0.16	0.53	0.20 ± 0.17	0.32	0.23 ± 0.20
f_{N}	0.40	0.50 ± 0.18	0.42	0.51 ± 0.19	0.29	0.47 ± 0.19	0.43	0.45 ± 0.21
f_{Si}	0.06	0.11 ± 0.12	0.08	0.12 ± 0.13	0.0	0.11 ± 0.12	0.06	0.078 ± 0.086
f_{Fe}	0.01	0.052 ± 0.039	0.0	0.053 ± 0.042	0.01	0.050 ± 0.038	0.0	0.044 ± 0.034
α	1.8	1.83 ± 0.31	1.6	1.67 ± 0.36	1.1	1.33 ± 0.41	0.6	0.64 ± 0.44
$\log_{10}\left(\frac{E_{\mathrm{Fe,max}}}{\mathrm{eV}}\right)$	20.5	20.55 ± 0.26	20.5	20.52 ± 0.27	20.2	20.38 ± 0.25	20.2	20.16 ± 0.18

Flatter spectra preferred for negative source evolution

Hard spectra preferred for source evolution following that of the SFR

High Spectral Peaked Blazar Evolution

n=-6 evolution result

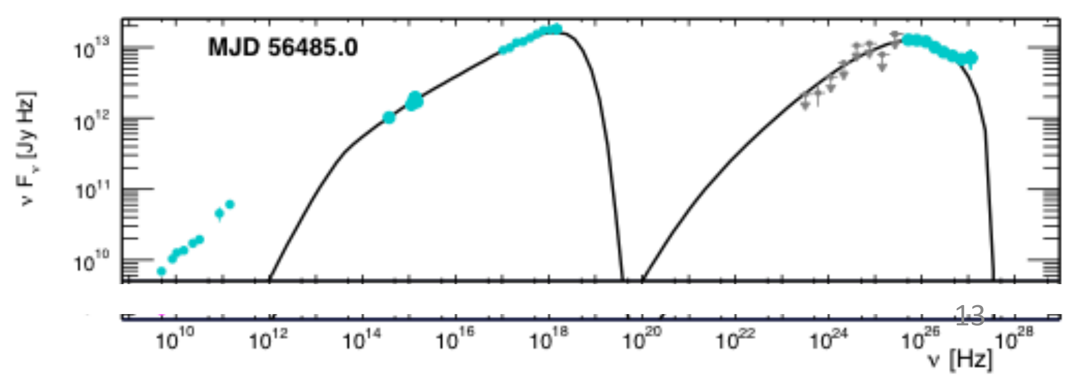


- Reminder:
  Blazar -> BL Lac (FR1) -> HSP

- Supports idea that FSRQ (gas accreting) AGN evolve into BL Lac (gas starved) AGN

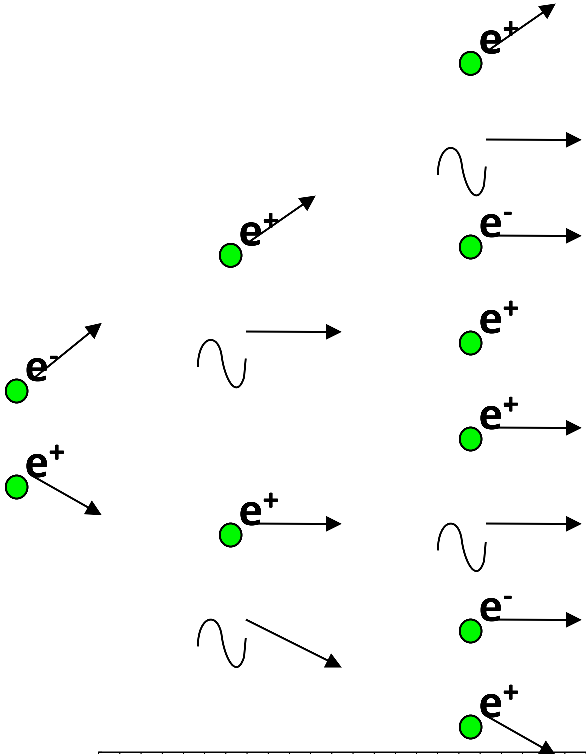
From astro-ph/1310.0006 (Ajello et al. 2014)

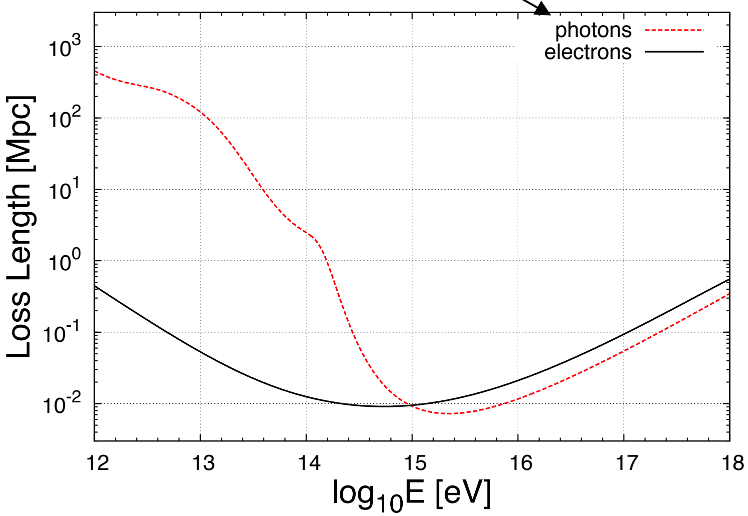
Archetypal HSP example Mrk 501

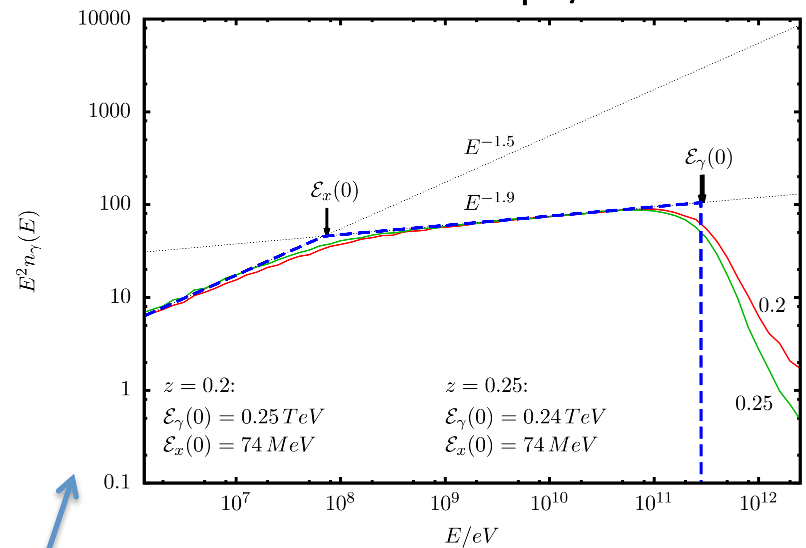


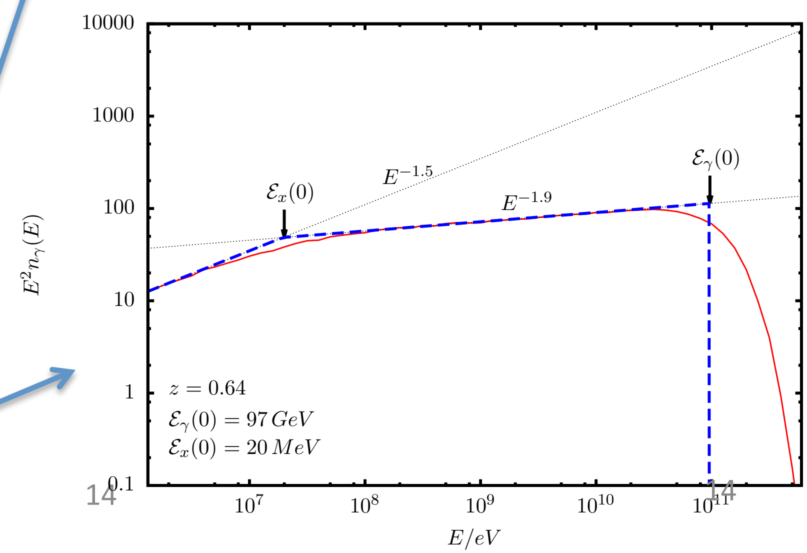
Universality of Cascade Spectra

From Berezinsky et al. astro-ph/**1603.03989**





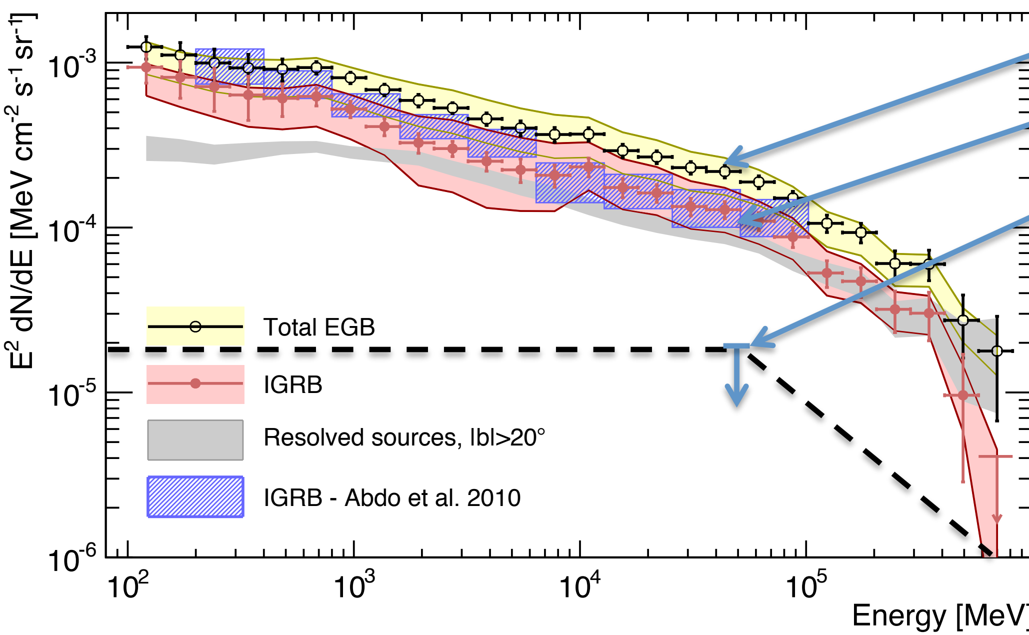




Regardless of where the energy is injected (ie independent of source z), the arriving flux possesses a universal shape

14

The Level of the Gamma-Ray Background



Latitude Cut

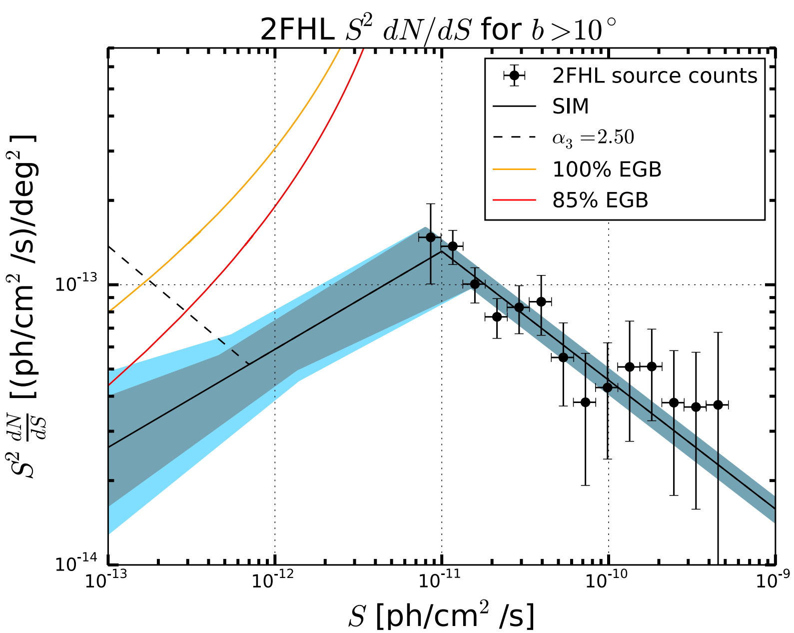
Lat. Cut + Removal of Res. Blazars

Lat. Cut + Removal of Blazars

Using Photon Fluctuation Analysis, the Fermi collaboration pushed a factor of ~10 below the 2FHL sensitivity

$$\frac{\mathbf{dN}}{\mathbf{dS}} \propto \mathbf{S}^{-\alpha}$$

$$\mathbf{I} = \int \mathbf{S}\frac{\mathbf{dN}}{\mathbf{dS}}\mathbf{dS}$$

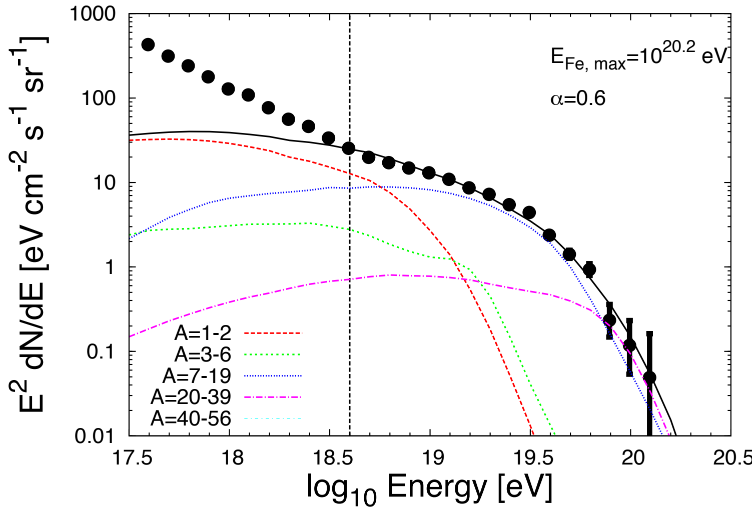


"Our analysis permits us to estimate that point sources, and in particular blazars, explain almost the totality (86^{+16}_{-14} %) of the >50 GeV EGB."

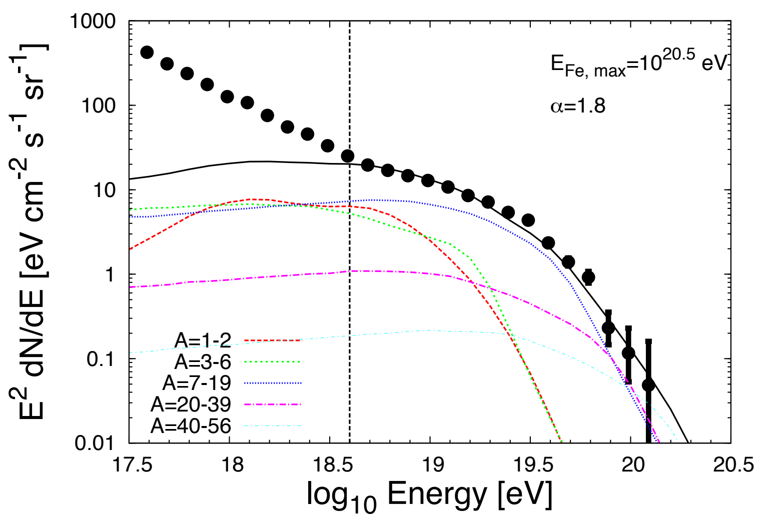
Fermi Collaboration (2015)- astro-ph/1511.00693

Secondary (Guaranteed) Gamma-Ray Fluxes From $>10^{18.6}$eV UHECR Component

n=3 evolution result

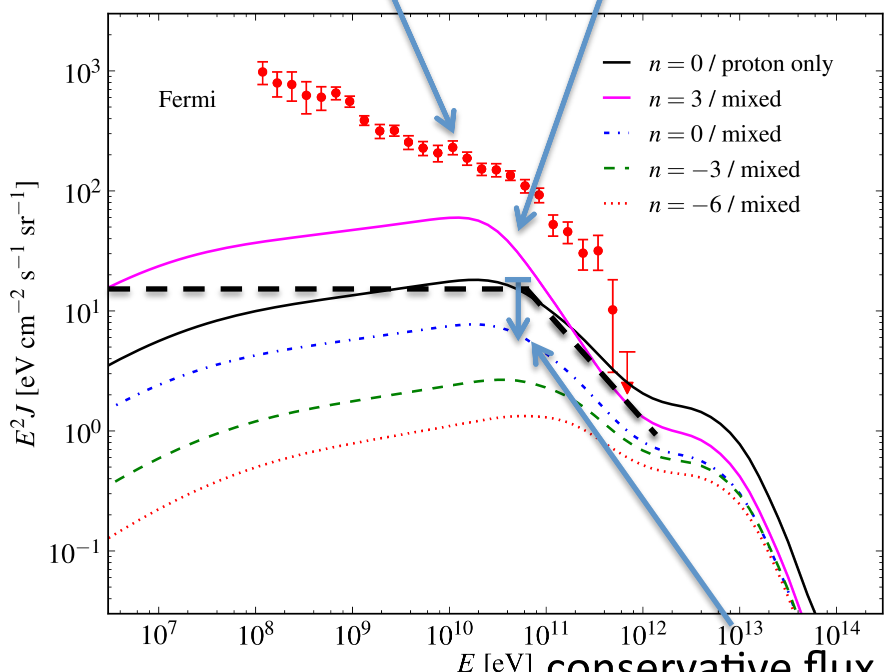


n=-6 evolution result



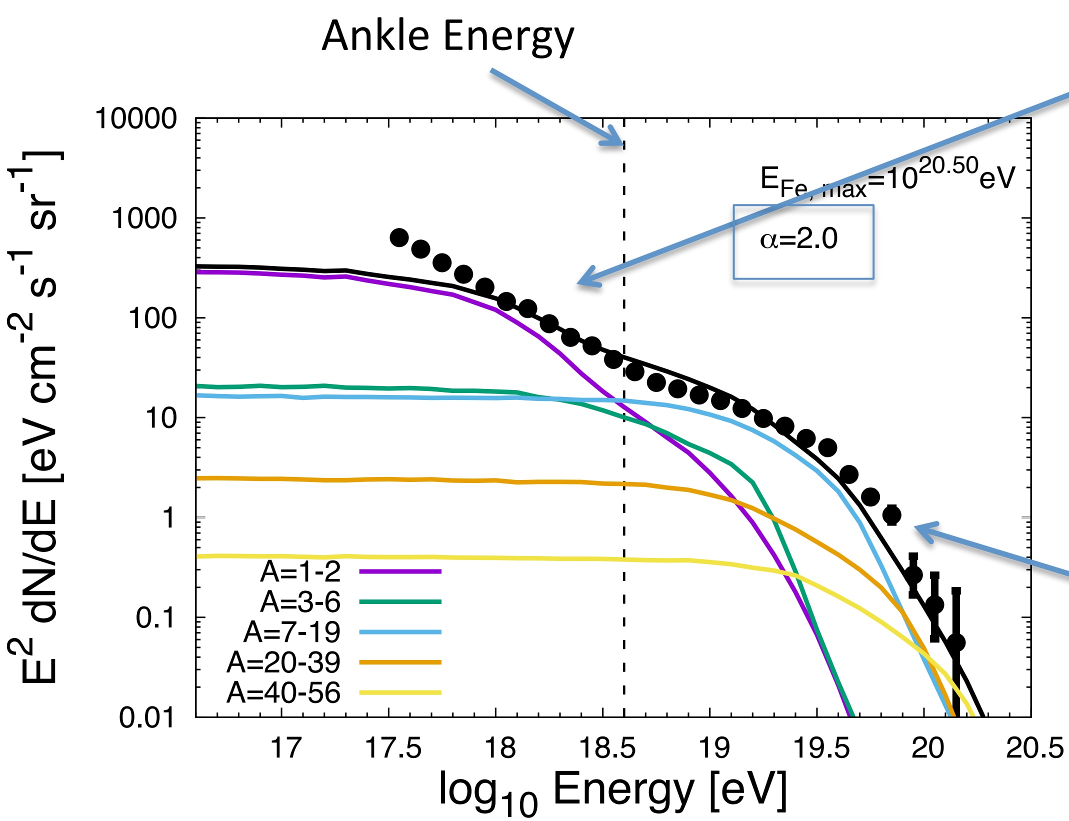
IGRB (EGB with resolved points sources removed)

n=3 to -6 evolution scenarios give rise to between **40%** and **12%** of Fermi limit

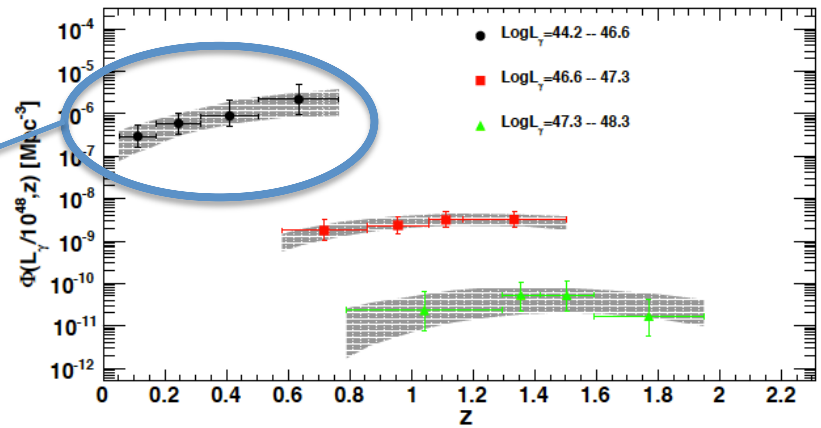


conservative flux upper limit at 50 GeV from astro-ph/1603.03223, Liu et al.

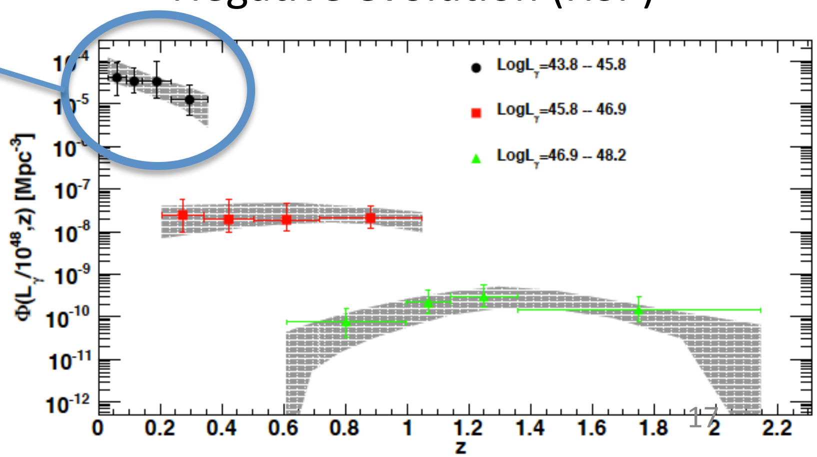
Does a Separate Class of Extragalactic Source Dominate at Sub-Ankle Energies?



Positive evolution (ISP + LSP)



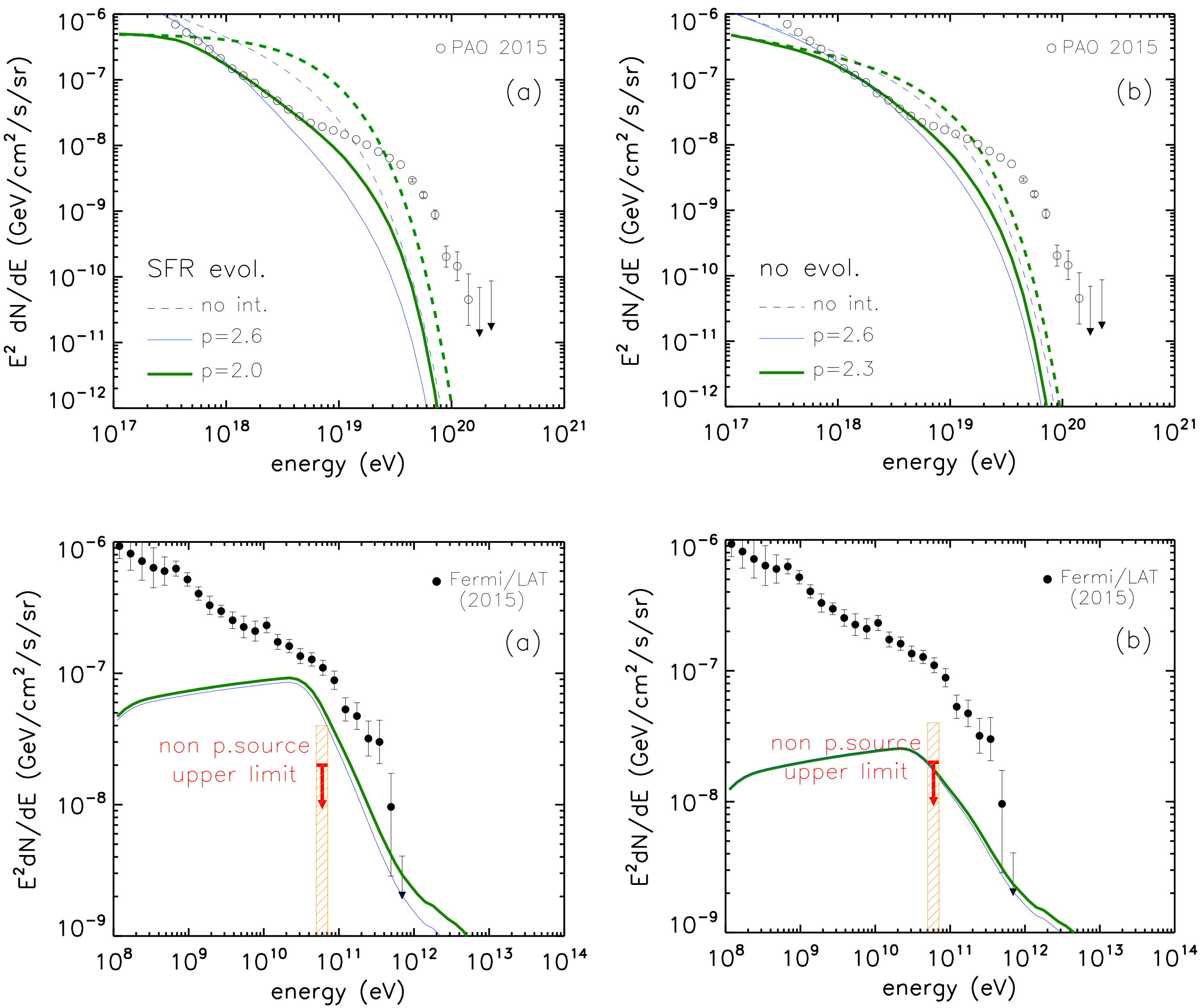
Negative evolution (HSP)



The Origin of Protons Below the Ankle

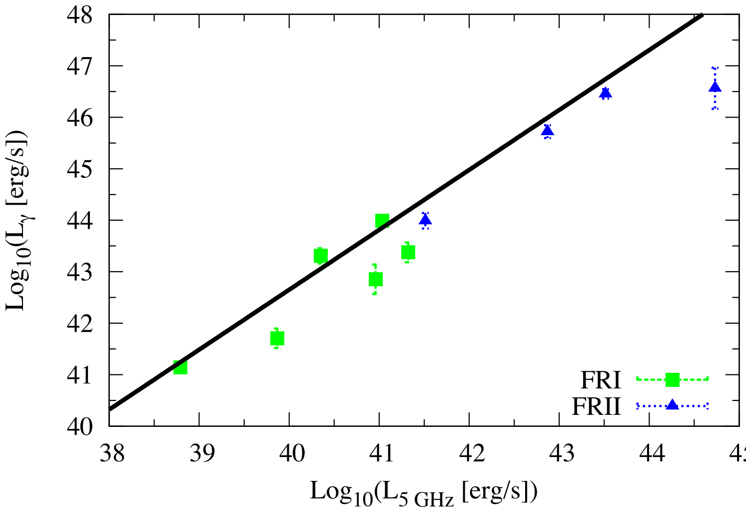
SFR evolution scenario

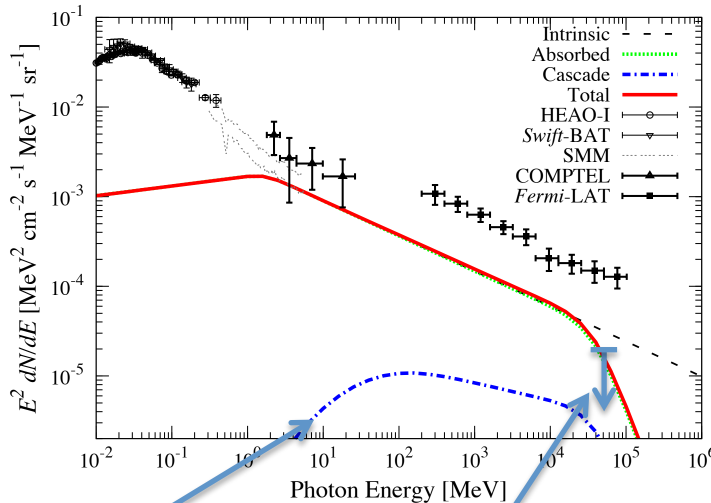
no evolution scenario



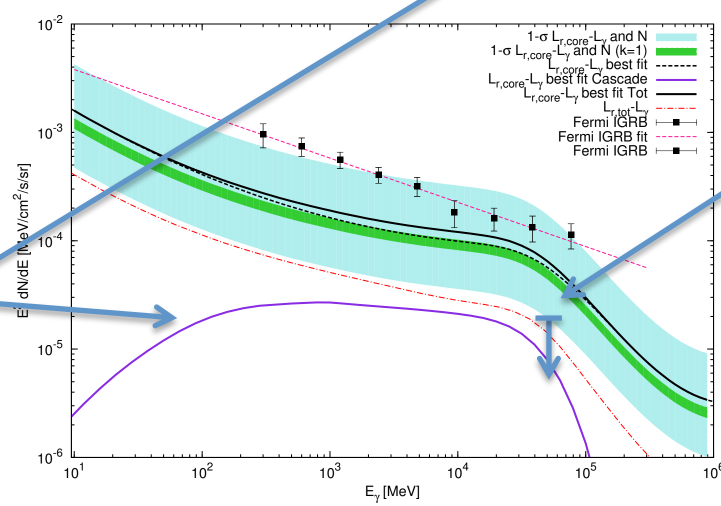
Note- IGRB contribution from cascade losses rather independent of source spectra

.....and Radio Galaxy Contributions Still Not Removed





From astro-ph/1103.3946 (Inoue et al. 2011)



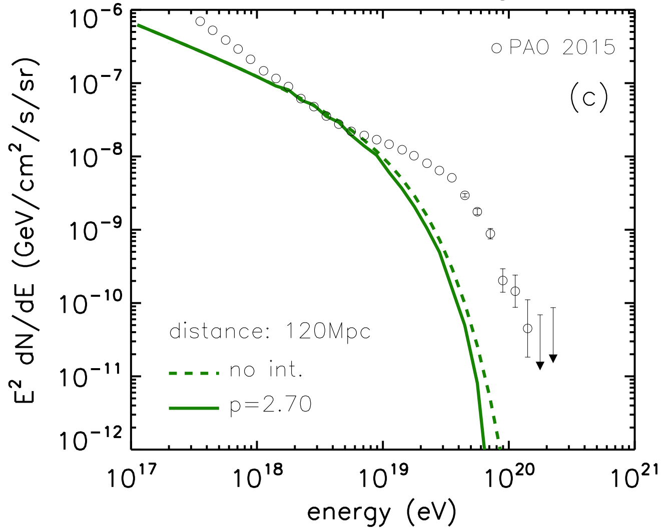
Note level of AGN gamma-ray generated cascades

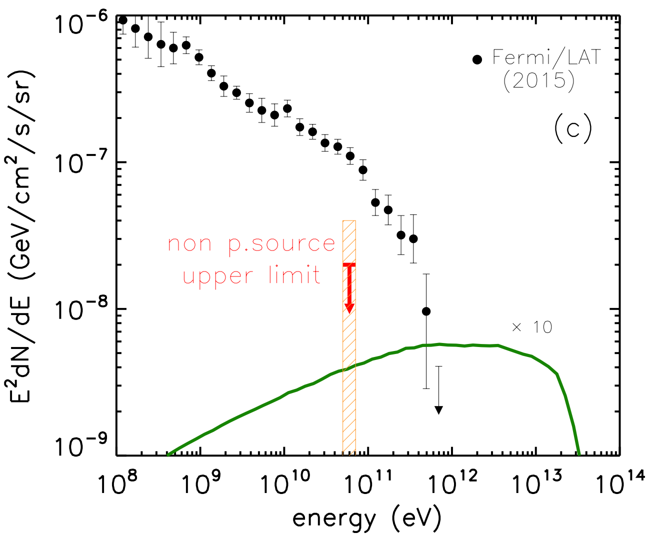
Radio Galaxy contributions are expected to make up a significant fraction of the remaining IGRB.

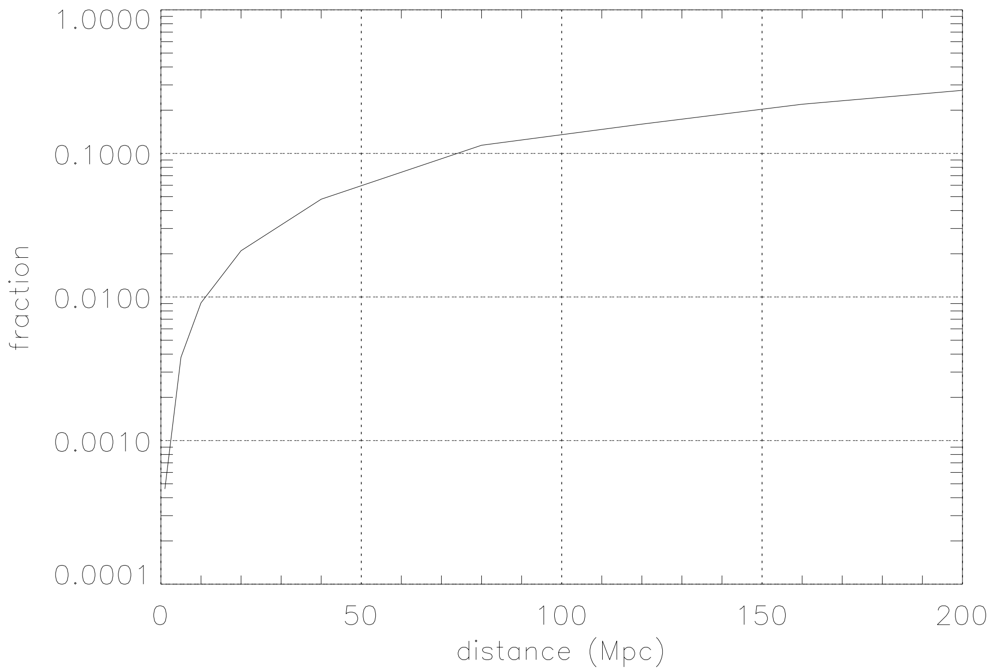
From astro-ph/1304.0908 (Di Mauro et al. 2013)

The Origin of Protons Below the Ankle

Sources at 120 Mpc







If only 1% of EGB comes from sub-ankle UHECR (present limit is 14%), we will be forced to look extremely locally for their sources

An Alternative Interpretation of the Negative Source Evolution Result

At high energies, the negative evolution scenarios help resolve both:

- “hard spectrum”
- “IGRB over-production”

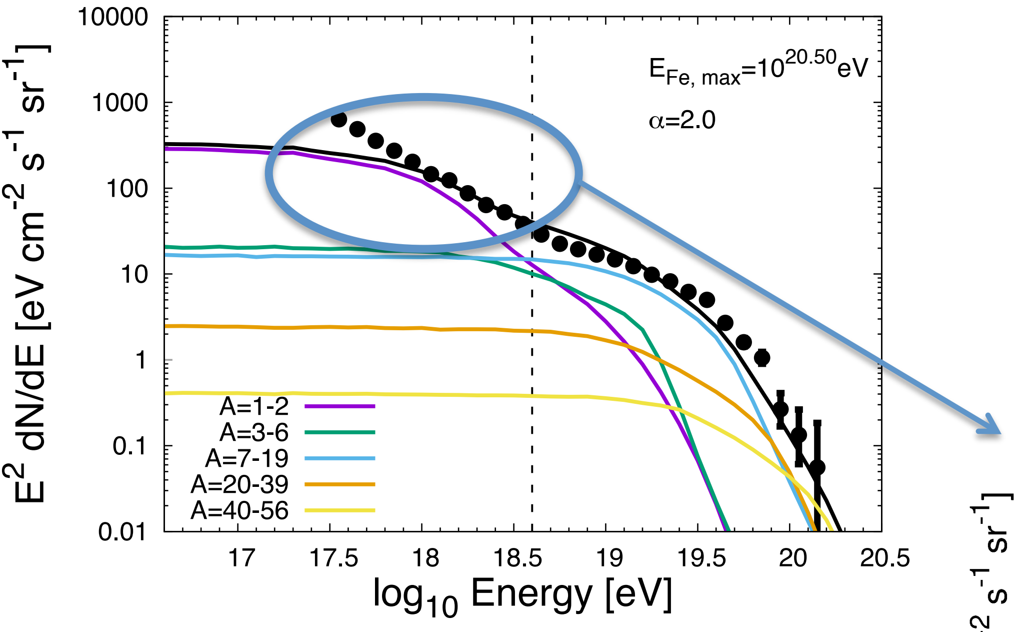
problems.

Alternatively, these scenarios may simply be encapsulating the fact that we’ve a local dominant source and our local value for UHECR is well above the “sea level”!

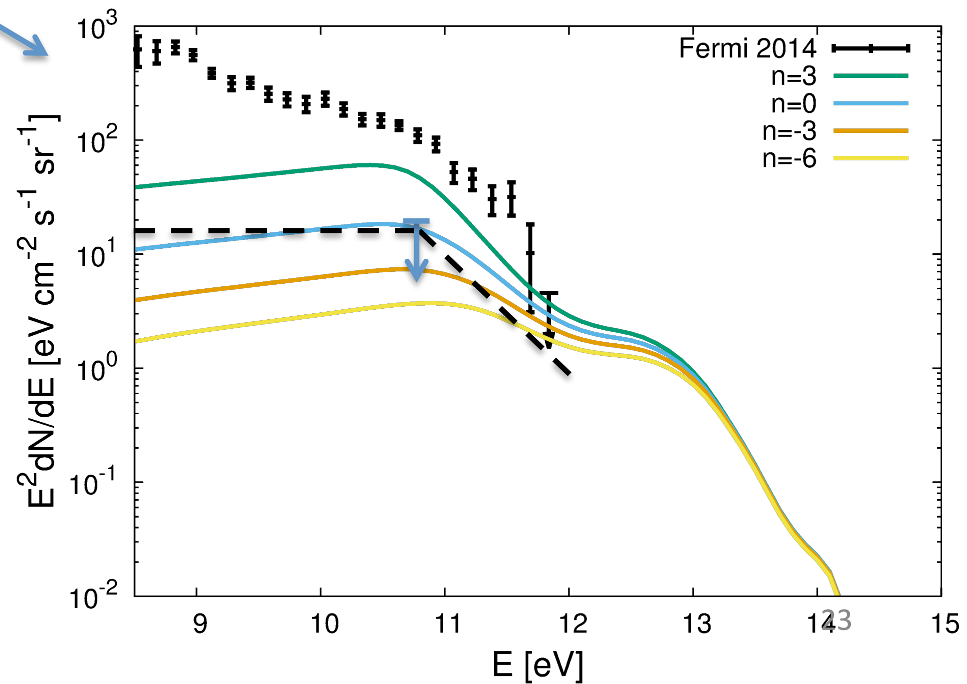
Conclusions

- A negative source evolution allows for an $\mathbf{E^{-2}}$ type spectra to explain CR above the ankle (such an evolution is observed for the HBL blazars)
- A new estimation of the diffuse gamma-ray background limit excludes positive evolution scenarios for these cosmic rays.
- The positive evolution of a separate source class, can account for sub Ankle extragalactic cosmic rays (which again allow an $\mathbf{E^{-2}}$ type spectra for this component)
- New diffuse gamma-ray background limits are challenging for both positive and no-evolution scenarios which account for sub-Ankle extragalactic protons
- These results suggest that UHECR exist in a local fog, with the value locally being well above the "sea level".
- An "understanding" of UHECR sources is possible through an understanding of AGN gamma-ray emission at very high energies!

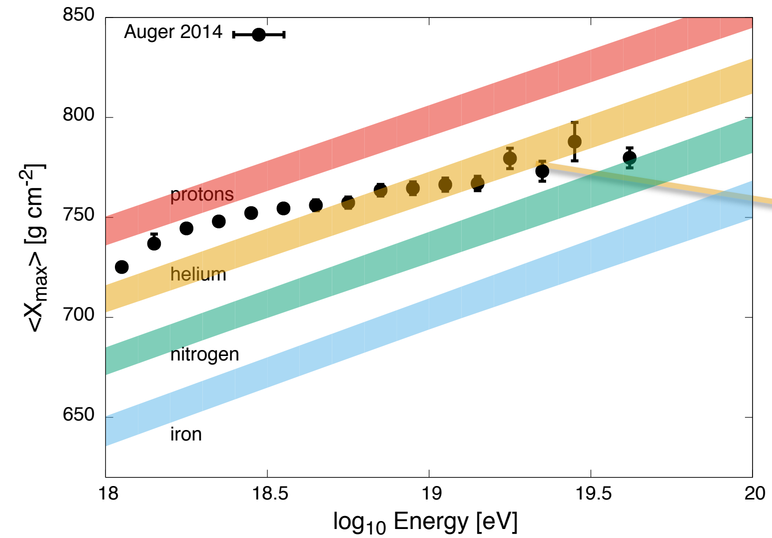
Cascade Contribution from Second Source Population

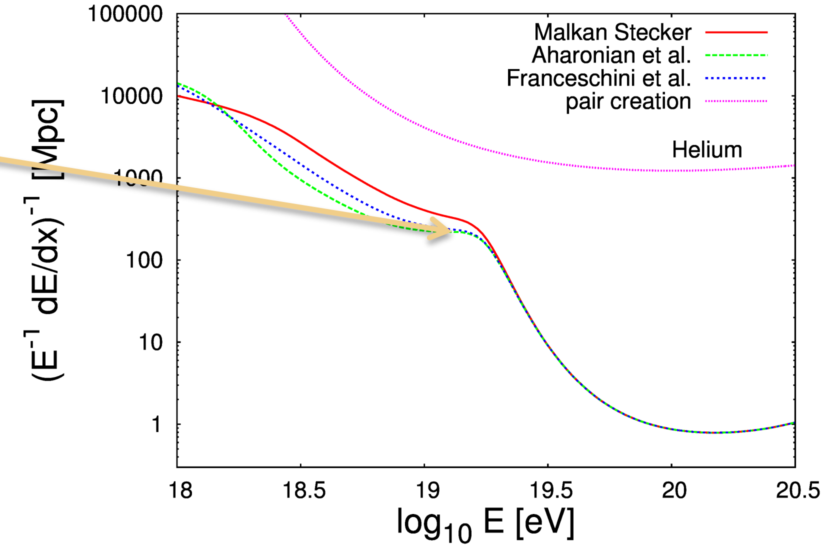


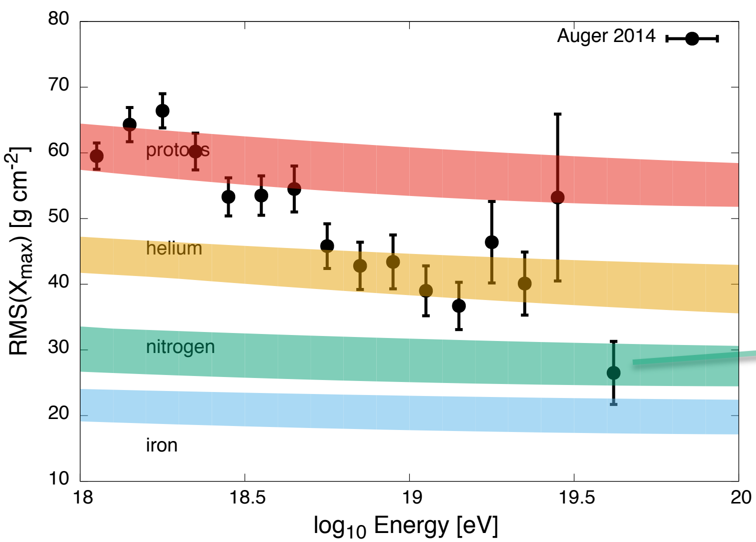
n=3 to -6 evolution scenarios give rise to between **100%** and **40%** of Fermi limit

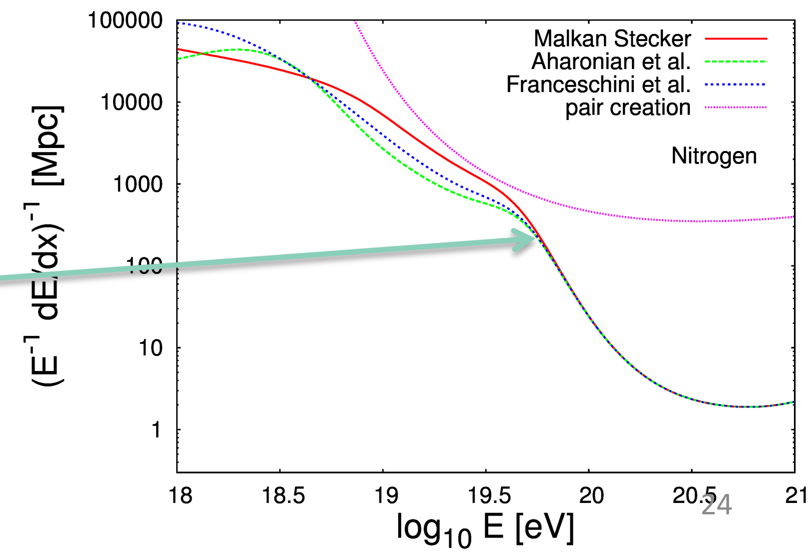


Sources of Cosmic Ray Nuclei Must be Nearby









Big Implications of these (Conservative) New Diffuse Gamma-Ray Limits

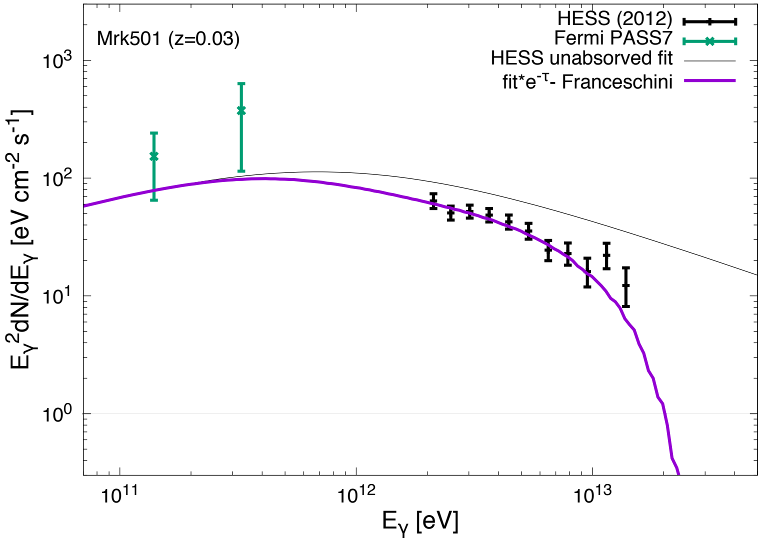
- The positive evolution scenario, favoured by a range of source models, is now disallowed by new limits (a continuation of the excluded scenarios discussed in Gelmini et al. astro-ph/**1107.1672**)

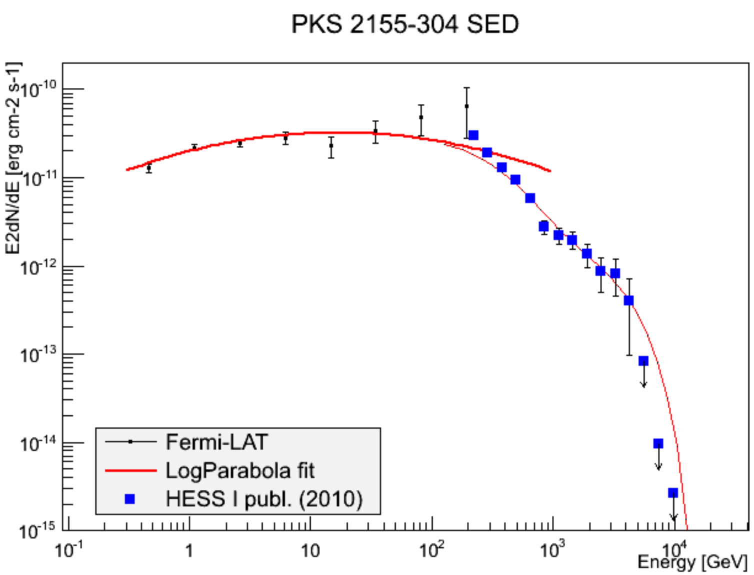
- Even the "no evolution" scenario of the sources, for which only mild spectral softening occurs, is in trouble

- A significant reduction in the cosmogenic neutrino flux is now imposed (bad news for ARA, ANITA, EVA...)

- Strong constraints also placed on sources of PeV neutrinos detected by IceCube (see astro-ph/1511.00688, Bechtol et al.)......

Why Conservative?.....Cascade Contributions from TeV Photons

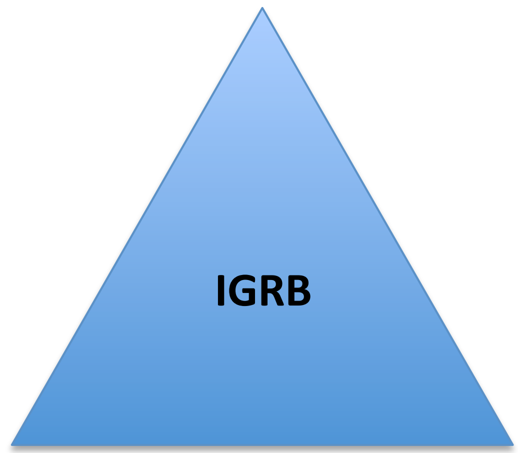




Only takes ~100 such objects to produce 100% of the EGB

The Promise of the IGRB

AGN Gamma-Rays Sources (Blazars + Radio Galaxies)



UHECR Propagation

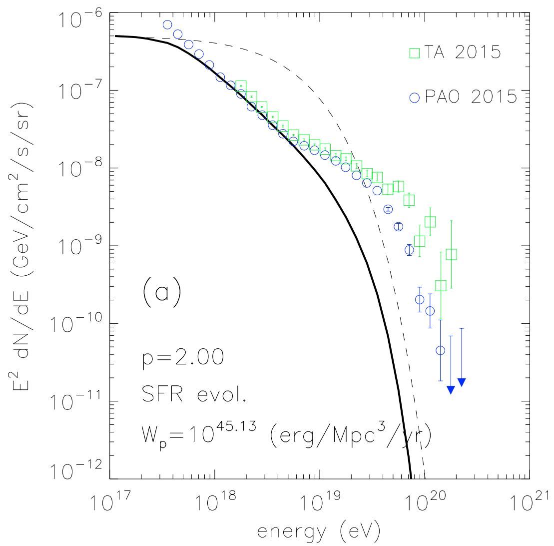
PeV Neutrino Sources ($\pi^{+/-/0}$)

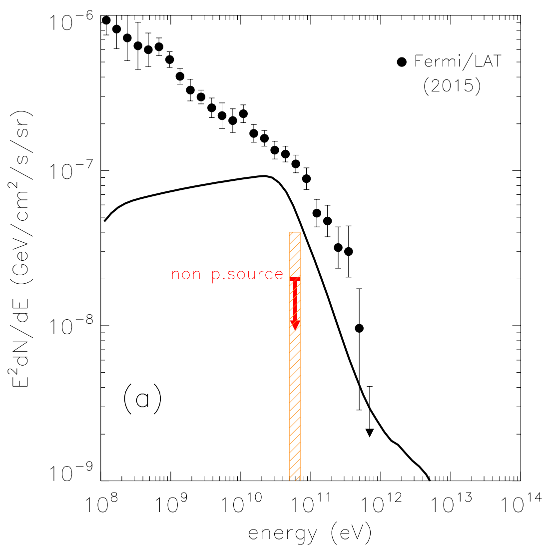
Each of these sectors wants to dominate the diffuse gamma-ray background….understanding this background holds huge potential for understanding these sectors.

Future Directions for IGRB Studies……TeV Bright AGN cascade and radio galaxy contributions

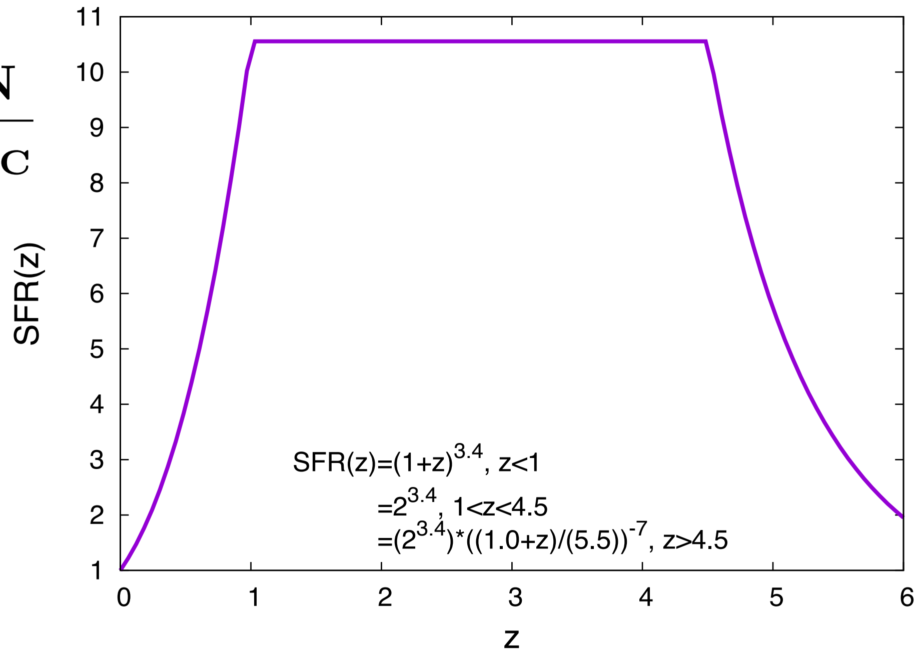
The Origin of Protons Below the Ankle

SFR evolution scenario

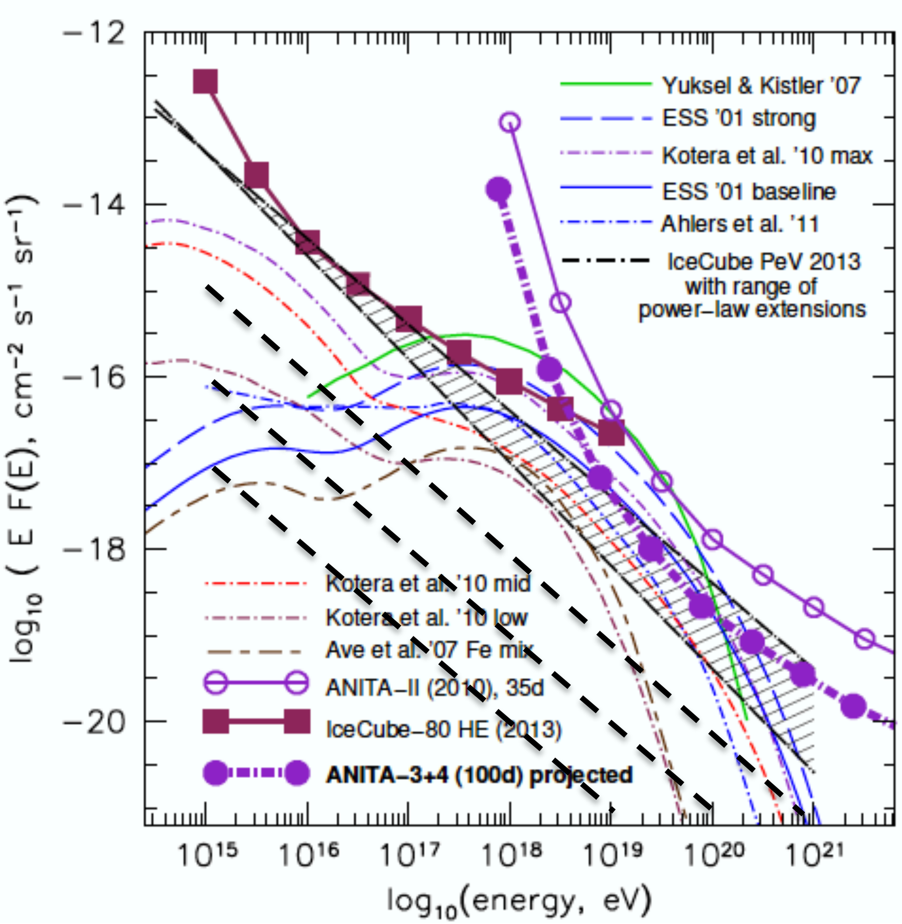


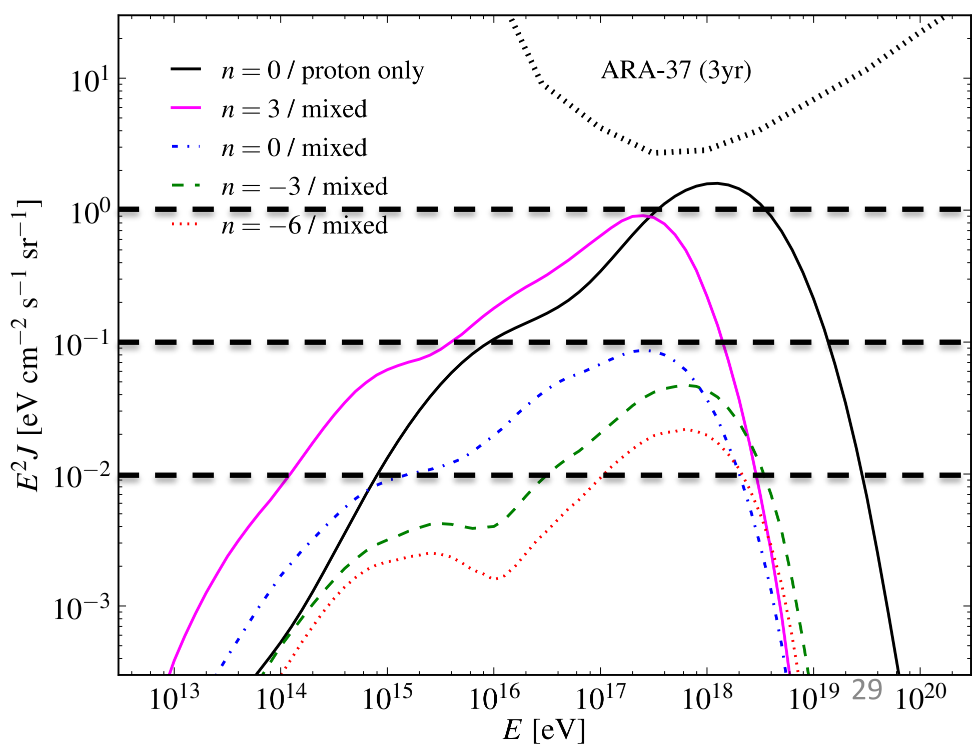


$$\frac{dN}{dV_C}$$



Secondary Neutrino Fluxes

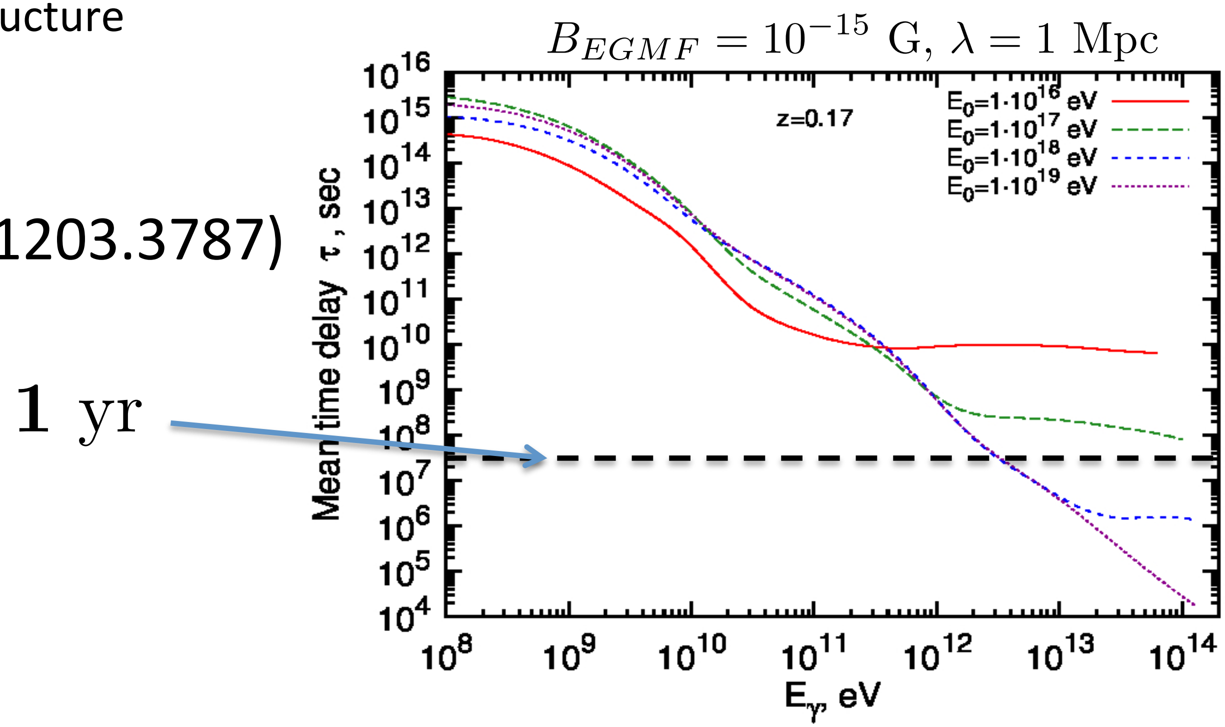




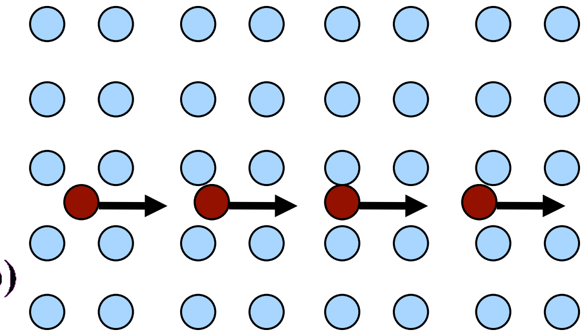
Proton Fed Blazar Emission Model

- Kusenko & Essey have spearheaded the suggestion that some TeV blazars are powered through proton losses in the presence of weak (10^{-15} G) extragalactic magnetic fields
- If this is the case, some subset of the component of resolved/ unresolved blazars should not be removed from the EGB
- However these blazars would not be expected to show short time-scale variability structure

Prosekin et al. (astro-ph/1203.3787)



Plasma Instability Effects (2)

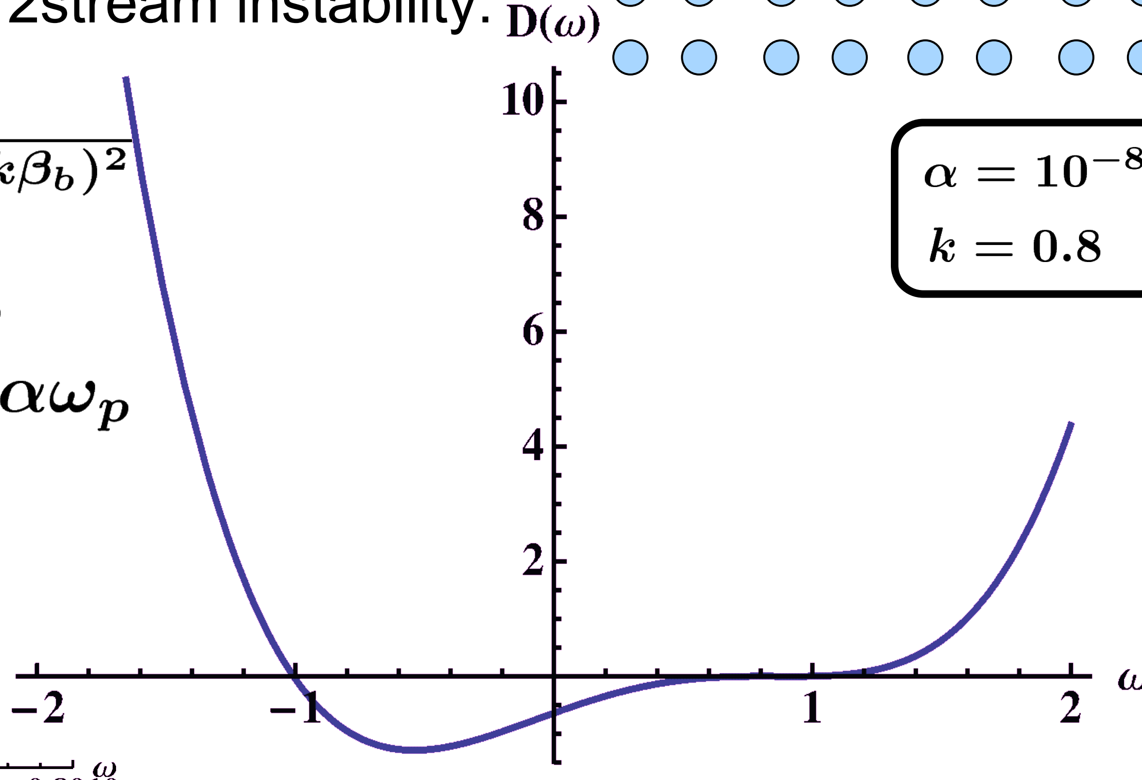


Dielectric equation for cold 2stream instability:

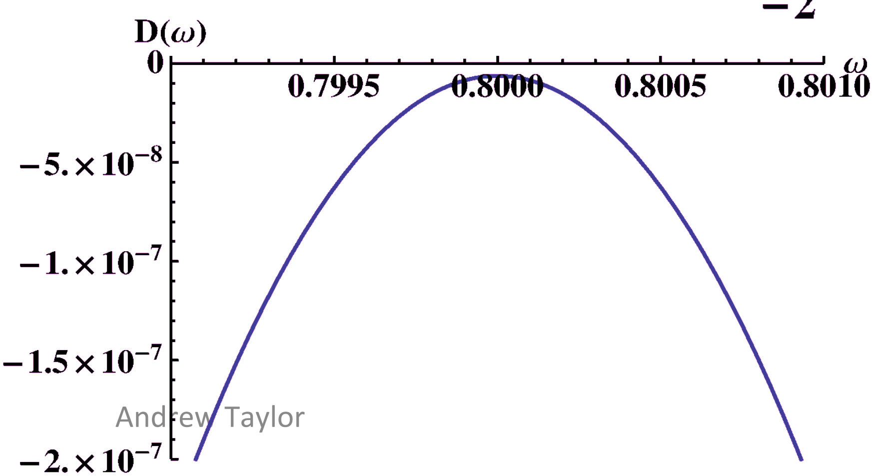
$$D = 1 - \frac{\omega_p^2}{\omega^2} - \frac{\omega_b^2}{\Gamma_b^3(\omega - k\beta_b)^2}$$

$$D = 0: \qquad \omega \approx \pm\omega_p$$

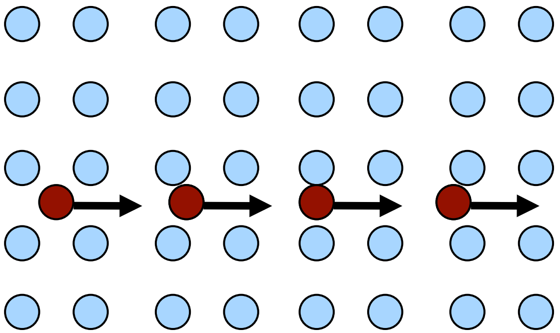
$$\omega \approx k\beta_b \pm \alpha\omega_p$$



$$\alpha = 10^{-8}$$
$$k = 0.8$$

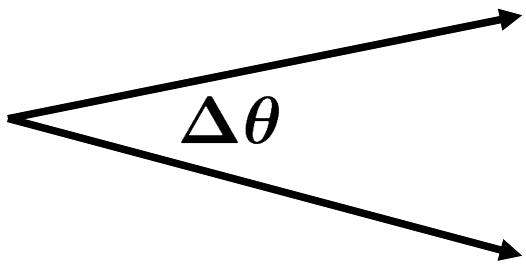


Plasma Cooling?



Growth rate: $$\frac{\delta}{\omega_p} \sim \left(\frac{1}{\Delta\theta}\right)^2 \frac{n_b}{\gamma_b n_p}$$

Is assumed that : $$\Delta\theta \approx \frac{p_\perp}{p_{||}} \approx \frac{1}{\gamma_b}$$



HOWEVER:

- Lorentz factor distribution in CoM frame must be considered fully $$\frac{p_\perp}{p_{||}} \neq \frac{1}{\gamma_b}$$

- Growth time is much larger than plasma frequency- not clear coherence can be maintained over such a long timescale

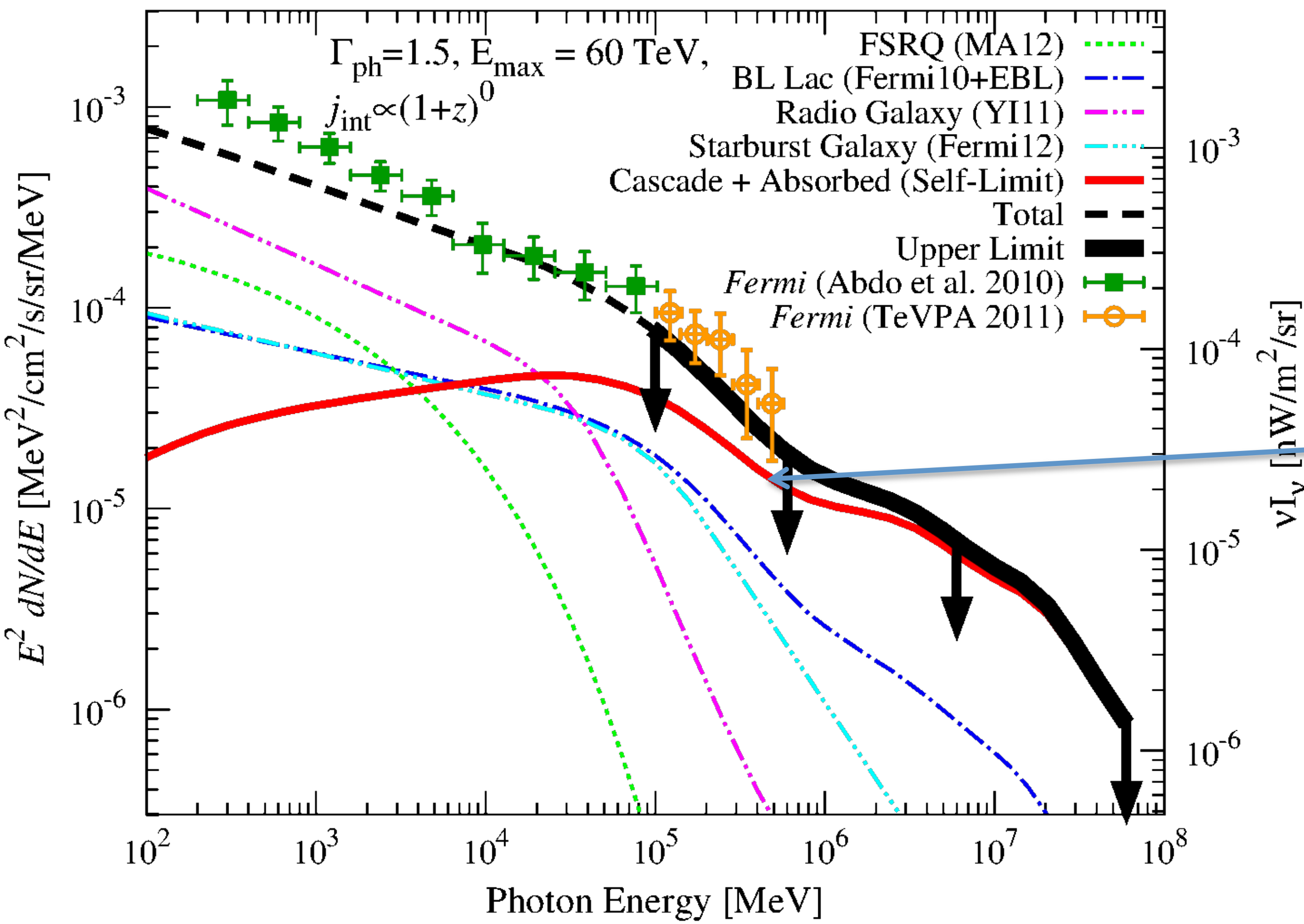
Broderick et al. (2012) astro-ph/1106.5494

Miniati et al. (2013) astro-ph/1208.1761

Sironi et al. (2014) astro-ph/1312.4538

Chang et al. (2014) astro-ph/1410.3797

Cascade Contribution Limit

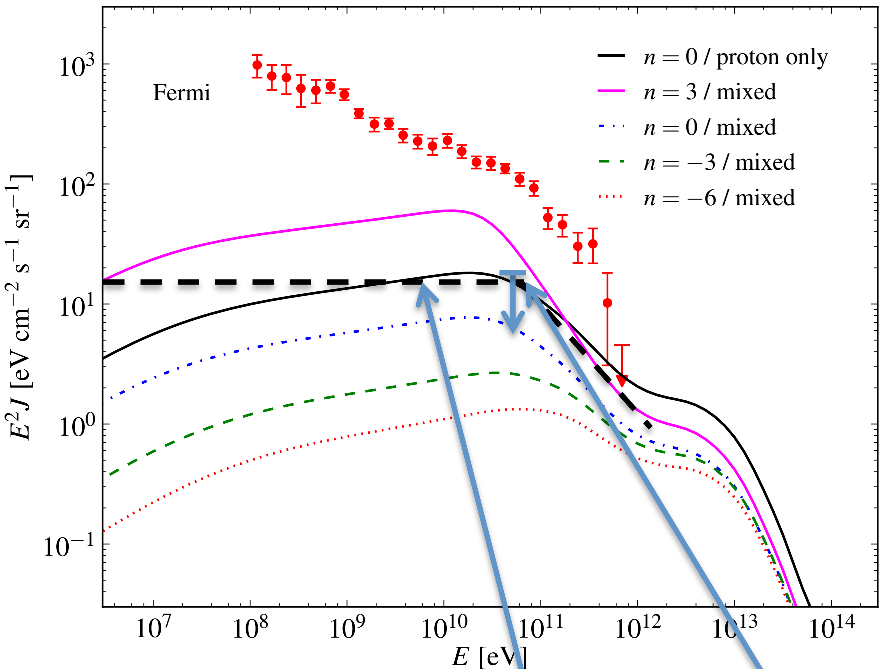


Limit on Cascade Component is around 40% of Fermi limit

Constraints also placed on evolution of contributors to cascade component

From astro-ph/1206.2923 (Inoue et al. 2012)

Revised Cascade Contribution Constraint



nuclei above $10^{18.6}$ eV

The n=3 scenario sits in conflict with this new constraint.

conservative flux upper limit at 50 GeV from astro-ph/1603.03223, Liu et al.

differential cascade limit taken from astro-ph/1511.00688, Bechtol et al.

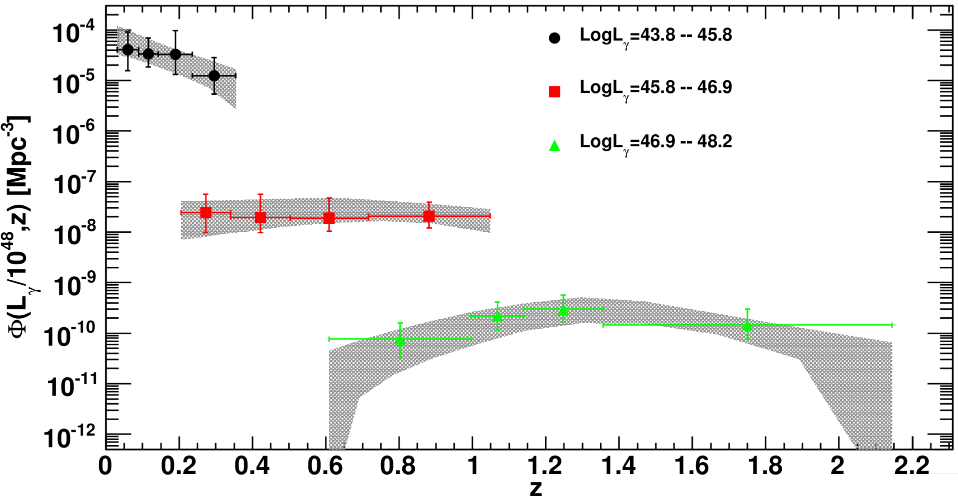
Similar Evolution Observed for Non-Blazar AGN?

Radio Loud AGN are suggested to have positive evolution (n=2) up to z=0.5, followed by negative evolution (n=-4) beyond this.

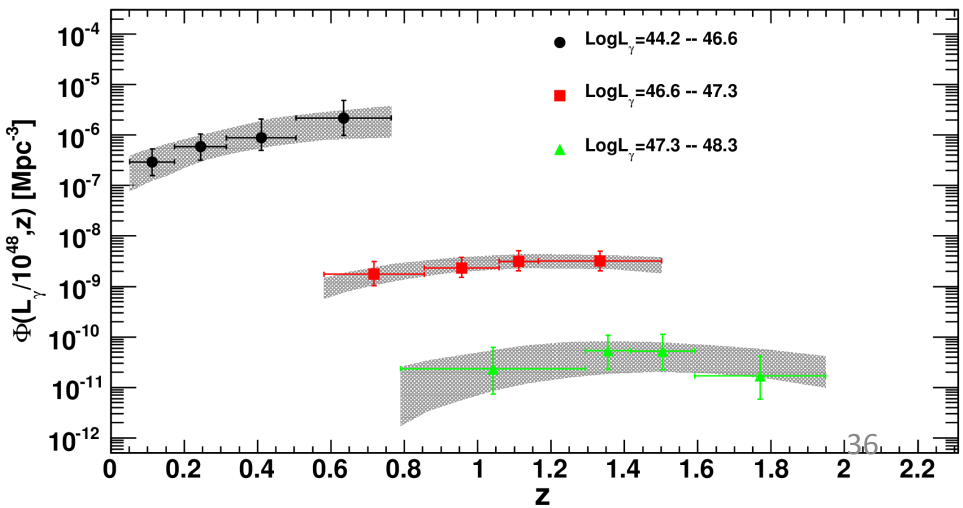
From astro-ph/1506.06554 (Padovani et al. 2015)

What About the Contribution from Other FR1 AGN (LSP + ISP)?

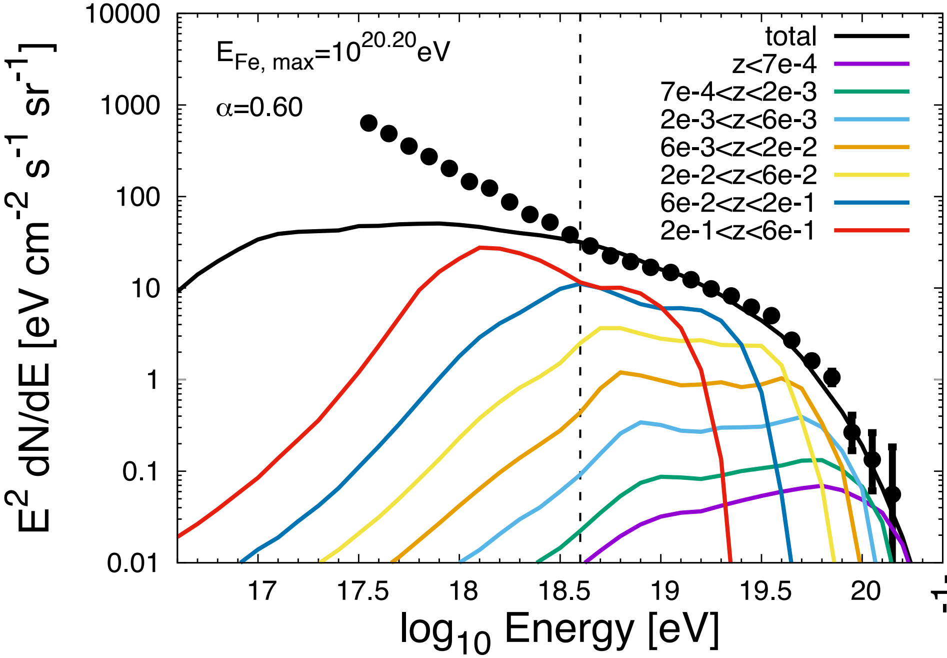
HSP AGN

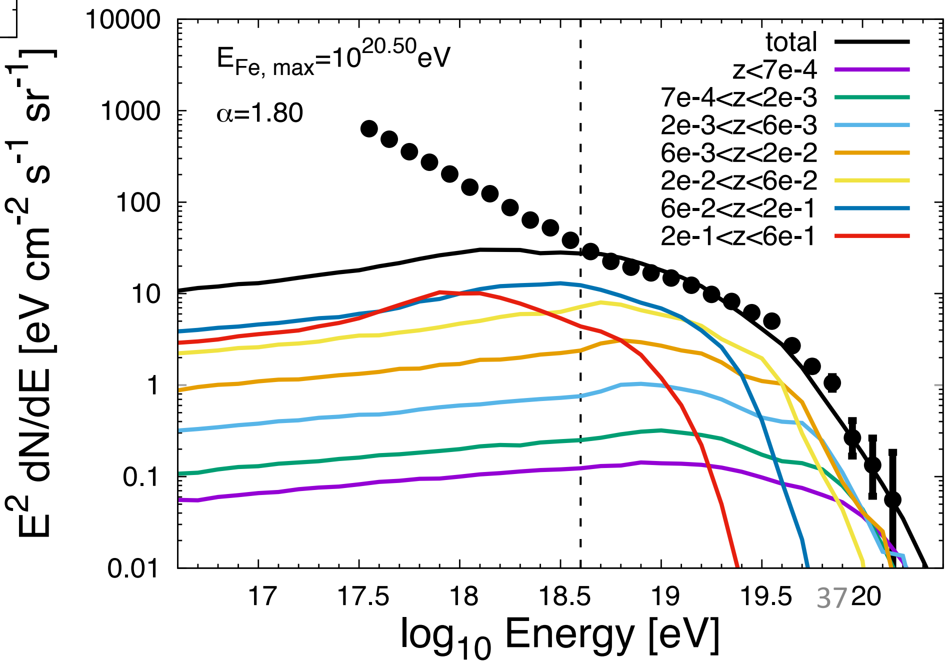


ISP + LSP AGN

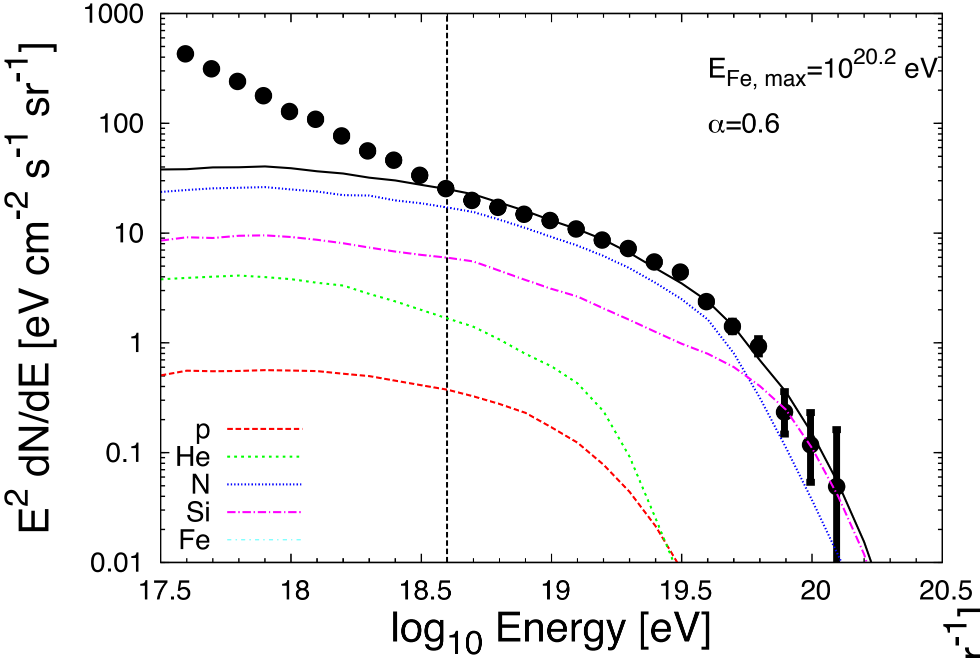


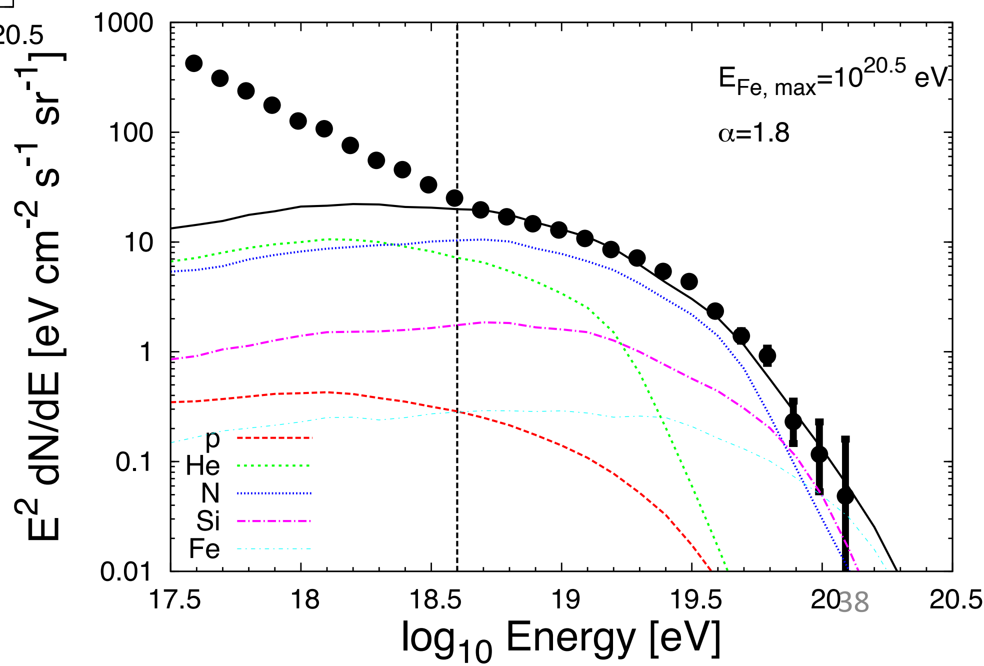
Source Redshifts Contributing to Arriving Flux





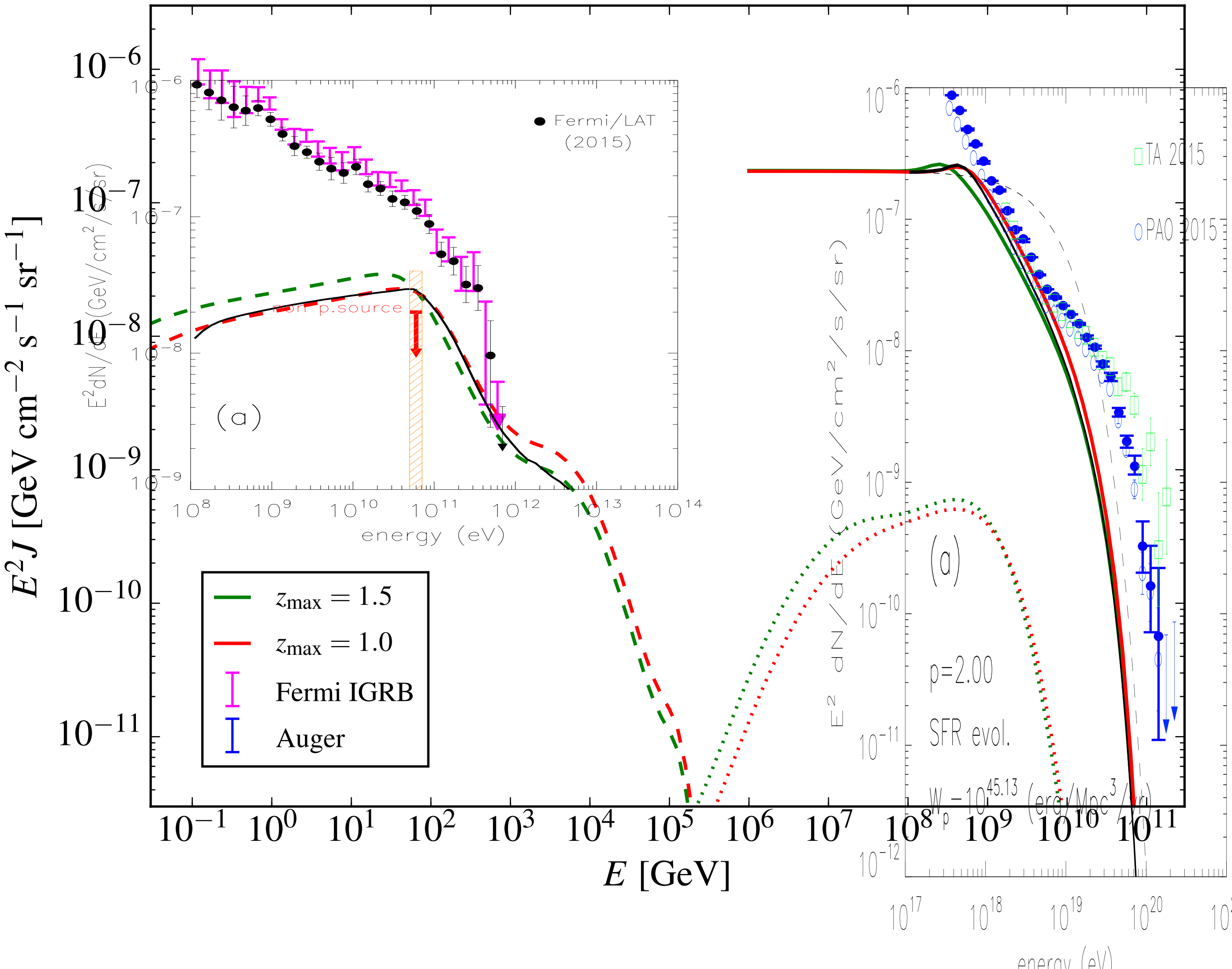
Injection Species Contributing to Arriving Flux





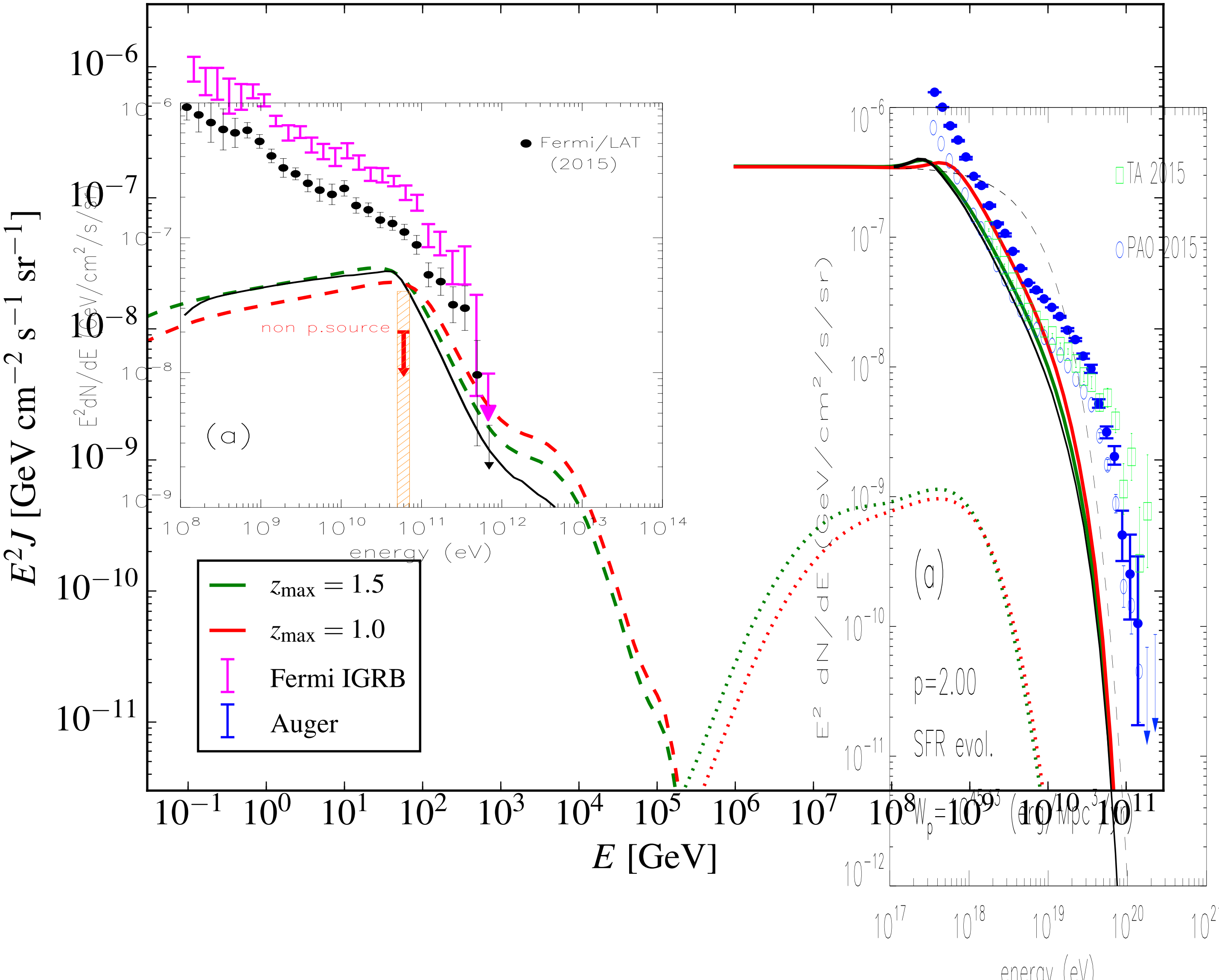
Comparison with Other Kinetic Equation Solver (MA)

$(1+z)^3$ evolution
zmax=1

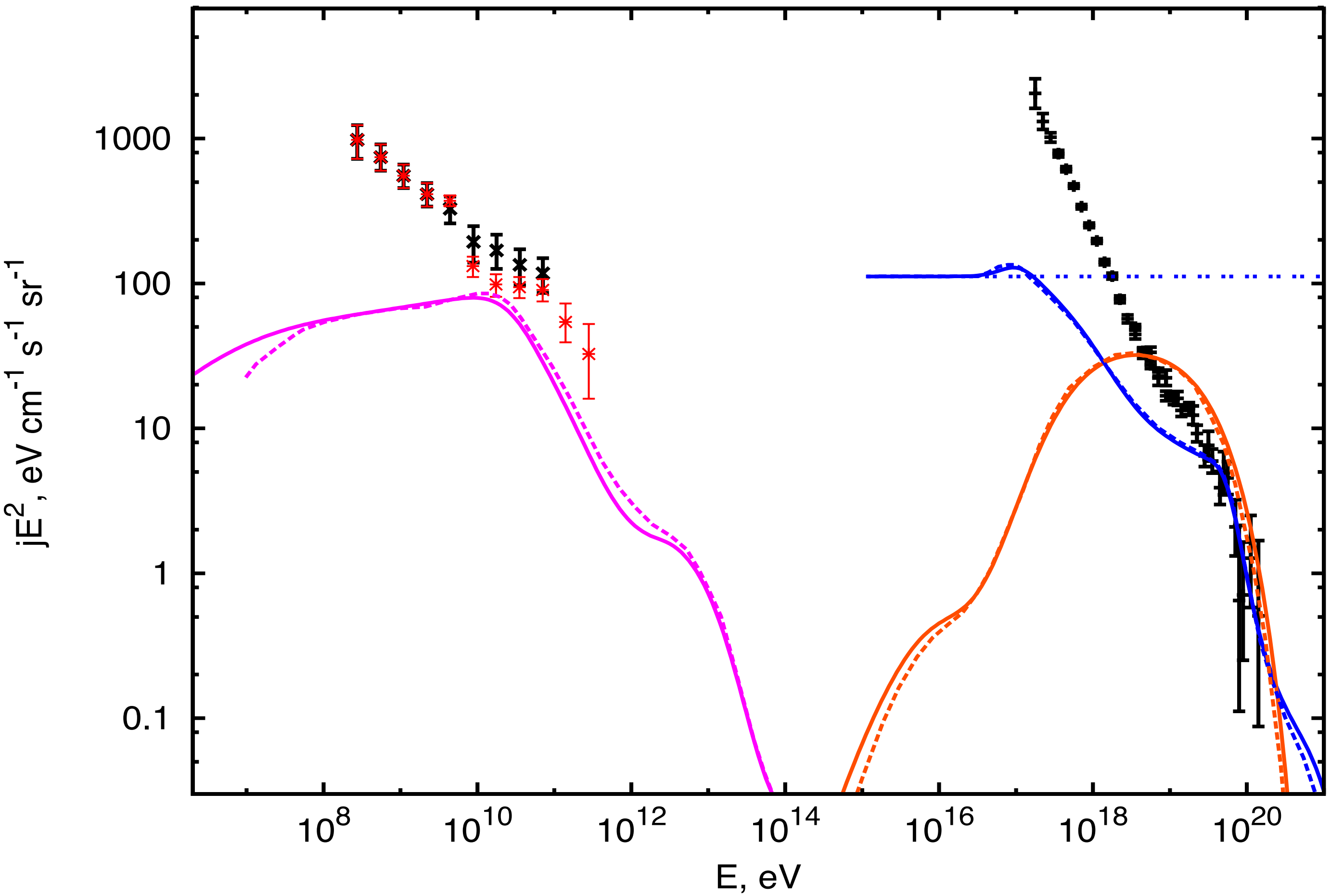


Comparison with Other Kinetic Equation Solver (MA)

$(1+z)^3$ evolution
zmax=1.5



Other Cross-Checks



A comparison is shown between the kinetic equation solver of Markus Ahlers and Oleg Kalashev

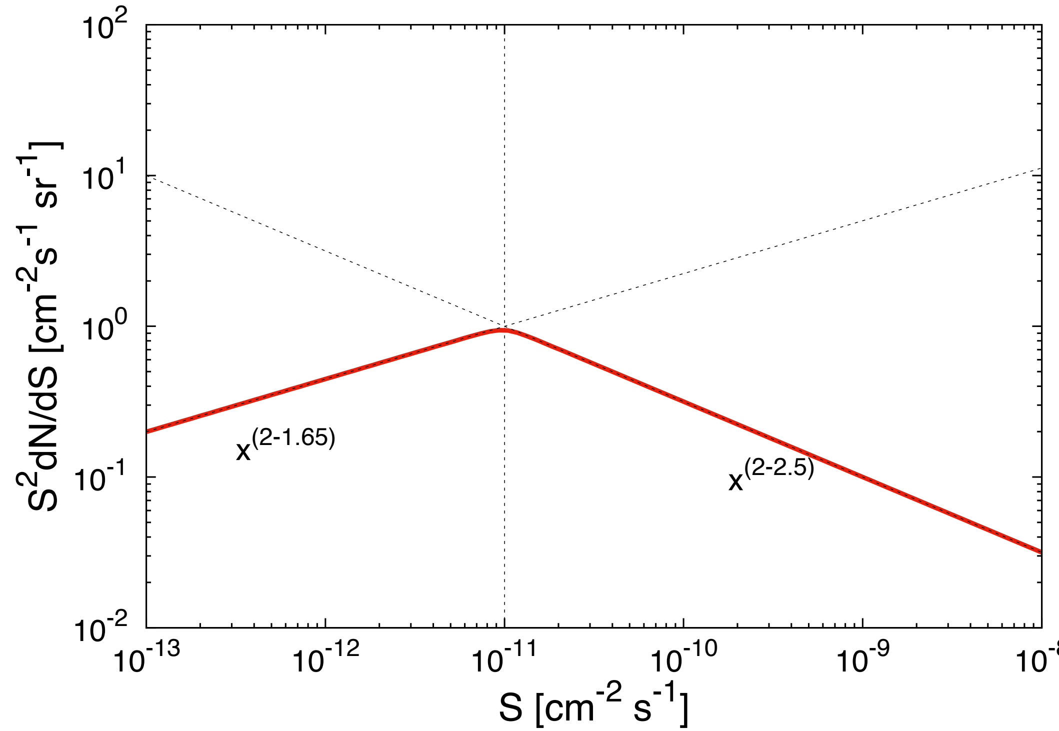
Gelmini et al. astro-ph/**1107.1672**

General Problem for Cascade Contribution?

Fermi Collaboration (2015)- astro-ph/1511.00693

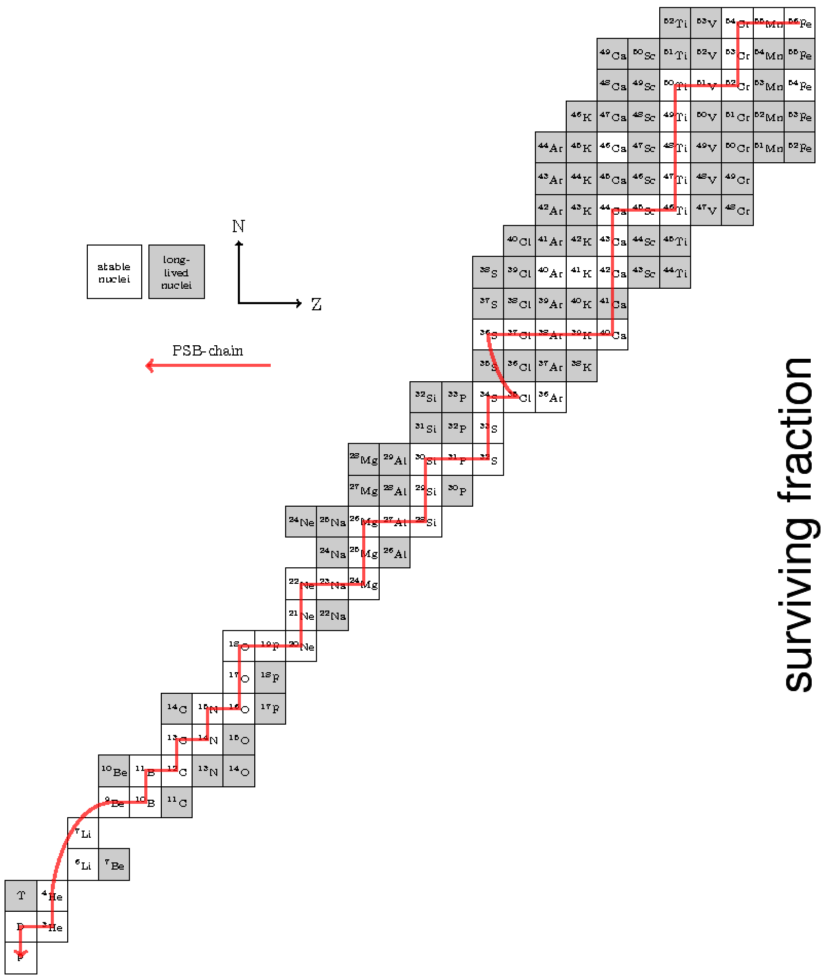
$$\frac{\mathbf{dN}}{\mathbf{dS}} \propto \mathbf{S}^{-\alpha}$$

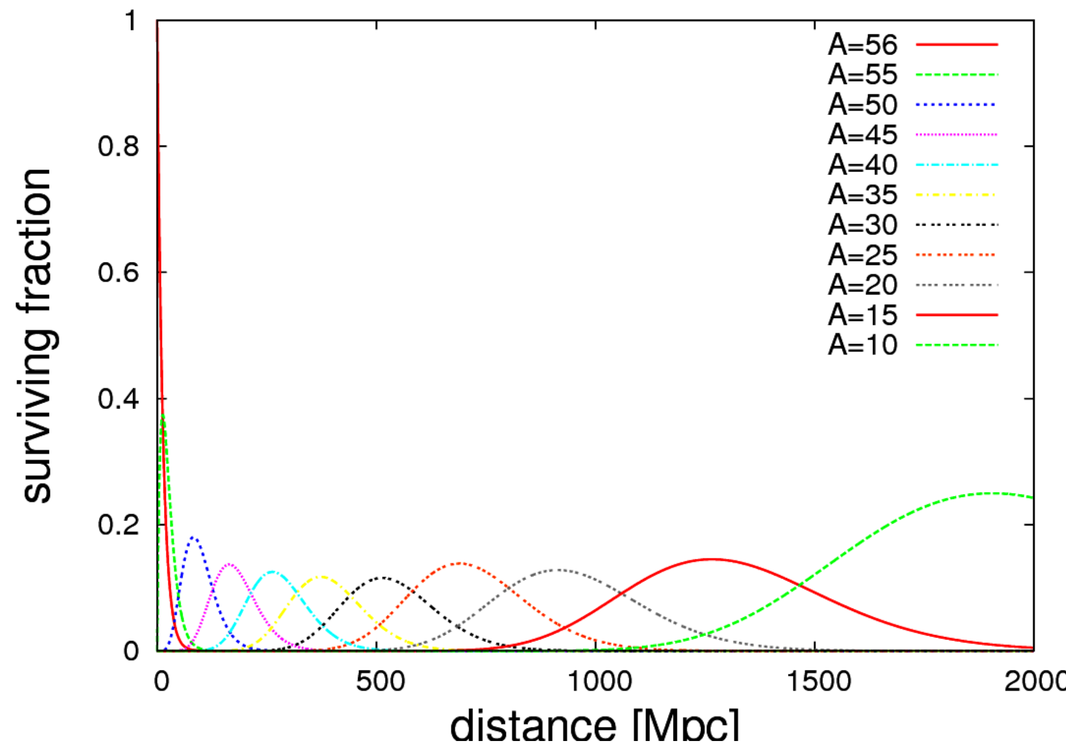
$$\mathbf{I} = \int \mathbf{S} \frac{\mathbf{dN}}{\mathbf{dS}} \mathbf{dS}$$



"Our analysis permits us to estimate that point sources, and in particular blazars, explain almost the totality (86^{+16}_{-14} %) of the >50 GeV EGB."

Nuclei Propagation and Disintegration





From astro-ph/**1107.2055**