Fermi bubbles

Dmitry Malyshev
Erlangen Center for Astroparticle Physics

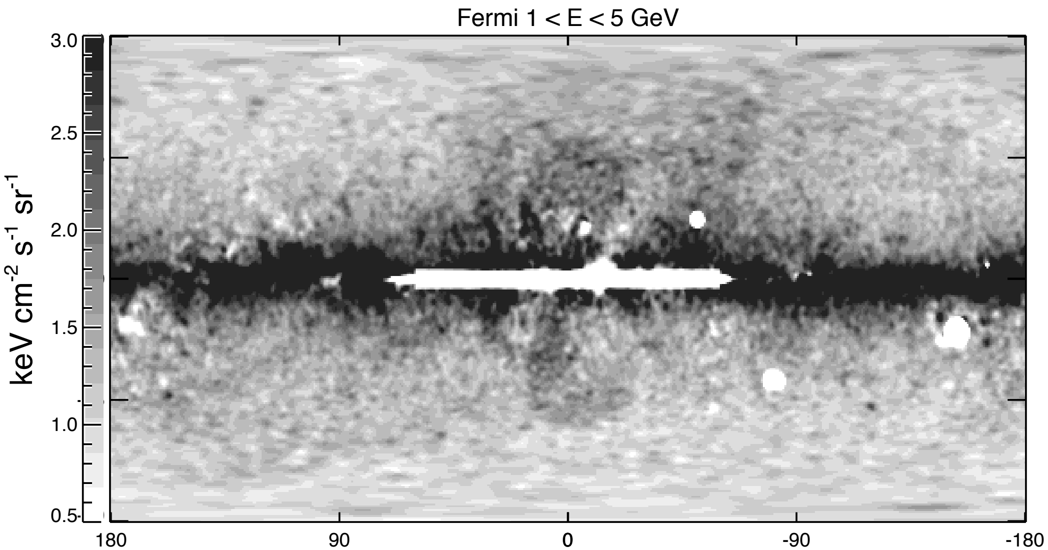
Anna Franckowiak,
Vahe' Petrosian

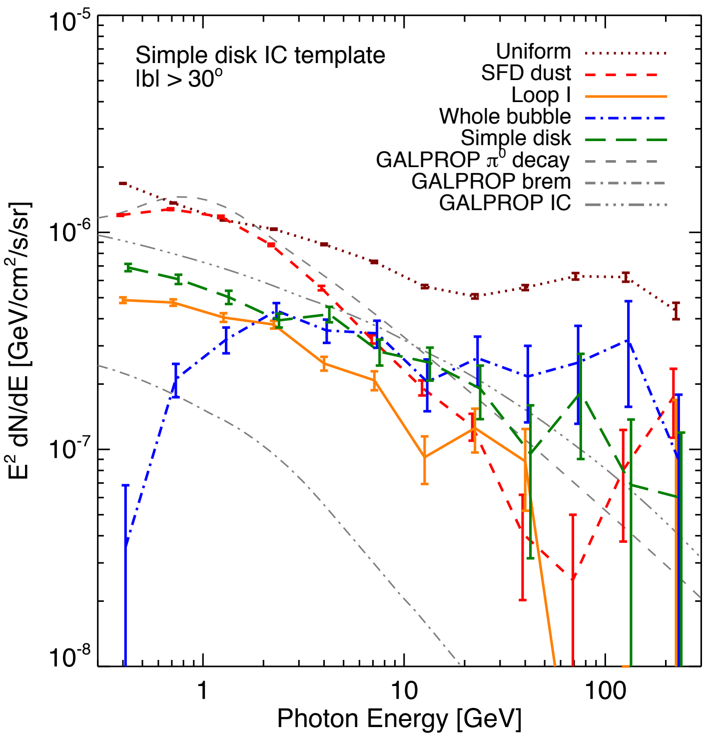
on behalf of the Fermi LAT collaboration

RICAP, Rome
June 22, 2016

Discovery of the Fermi bubbles

- **Su, Slatyer, Finkbeiner, May 2010**
  - **E^{-2} spectrum up to 100 GeV**
  - **have narrow edges**
  - **stretch up to 55° above and below the Galactic center**





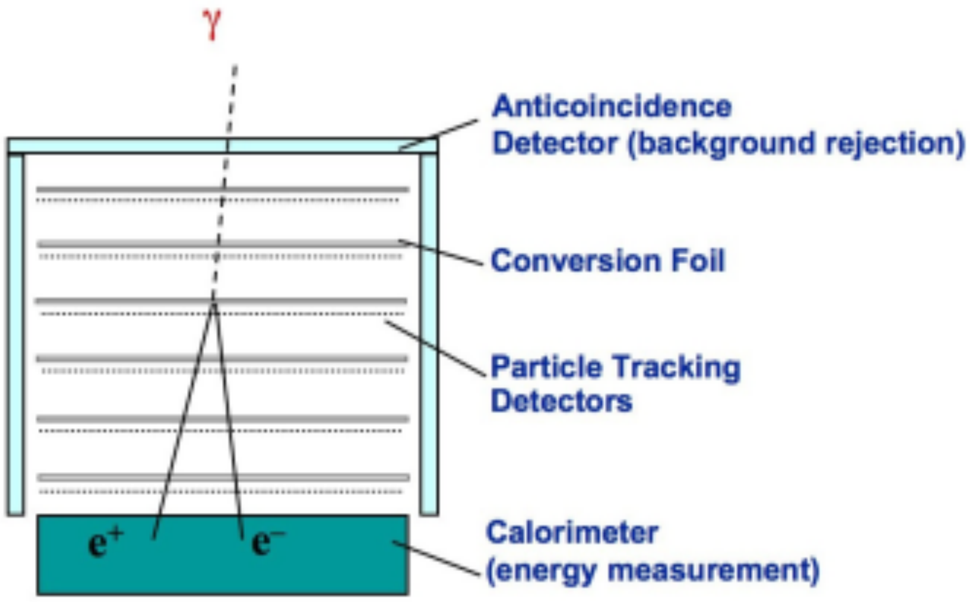
Su, Slatyer, Finkbeiner, ApJ 724 (2010)

Fermi bubbles – an elephant in gamma-ray sky

- **Fermi bubbles' solid angle is about 1 sr**
  - **This is comparable to an elephant at 3 m**

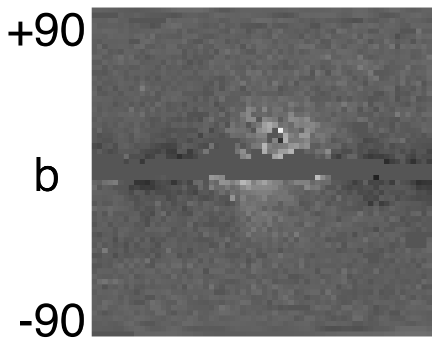
Fermi-LAT

- Fermi Large Area Telescope – gamma ray space telescope
- Launched on June 11, 2008
  - 20 MeV to more than 1 TeV
  - 2.4 sr field of view
  - Better than 1° resolution above 1 GeV
  - Covers the sky in two orbits (3 hours)

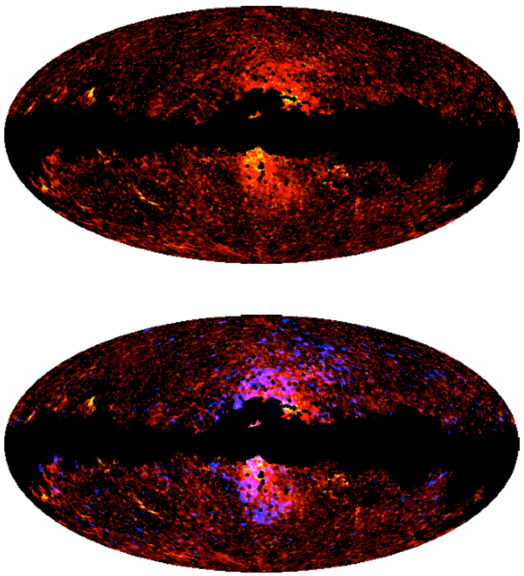


Haze

- **Microwave haze**

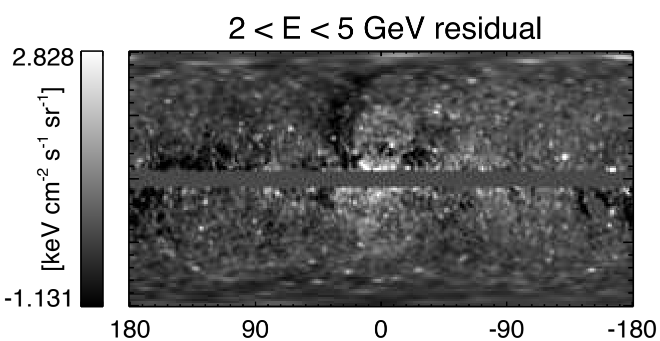


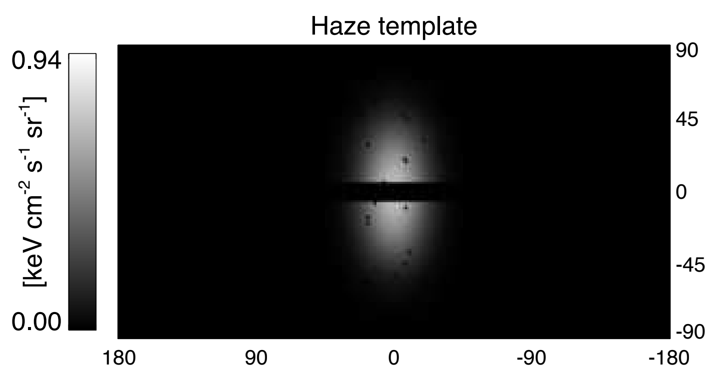
Finkbeiner, ApJ 614 (2004)



Planck Collaboration A&A 554 (2013)

- **Gamma-ray haze**

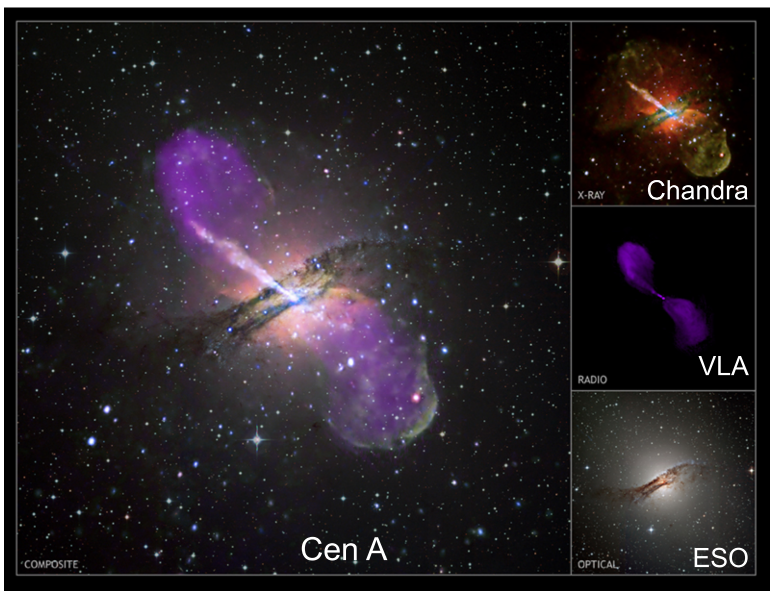




Dobler et al, ApJ 717 (2010)

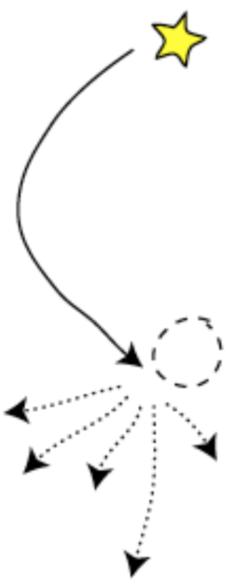
Fermi bubbles origin

- Emission mechanisms
  - Leptonic (inverse Compton)
  - Hadronic
- Origin
  - AGN-like activity (~ leptonic)
  - Star formation or star-burst (~ hadronic)

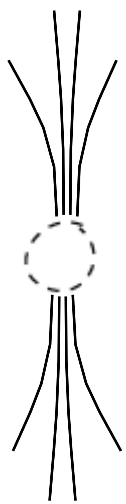


Leptonic Model

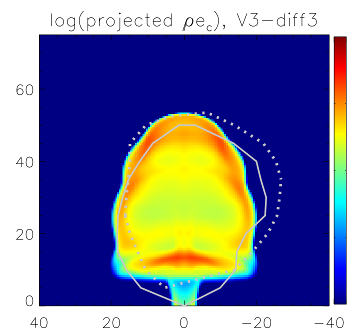
- **Electrons accelerated in the jet**
- **Gamma rays by inverse Compton scattering on radiation fields**
- **Microwave haze by synchrotron of same population of electrons**



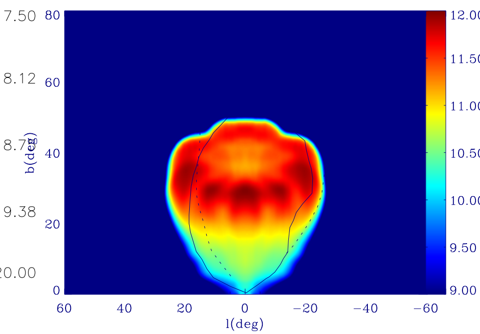
disruption of stars or molecular clouds by central black hole



AGN-like jet transports particles to high latitudes



Guo & Mathews
ApJ 756 (2012)



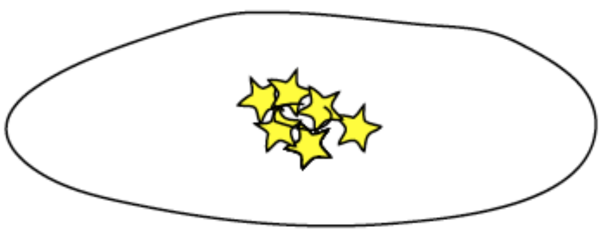
Yang et al.
ApJ 761 (2012)

Jets interact with interstellar medium to form bubbles

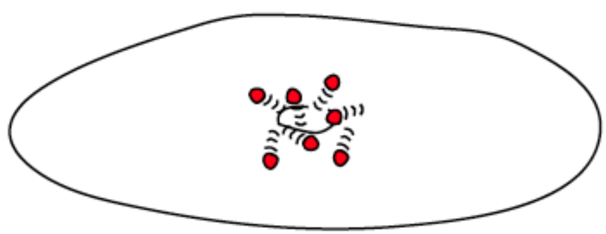
Illustrations by P. Mertsch

Hadronic Model

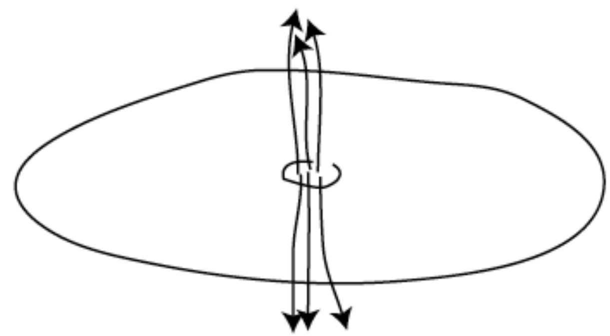
- **Cosmic rays accelerated by Supernovae shells**
- **Gamma rays by π^0 on thermal gas (density ~ 0.01 cm^{-3})**
- **Secondary e^+e^- produce synchrotron radiation**



increased star formation rate close to GC



acceleration of CR protons and nuclei in SNRs



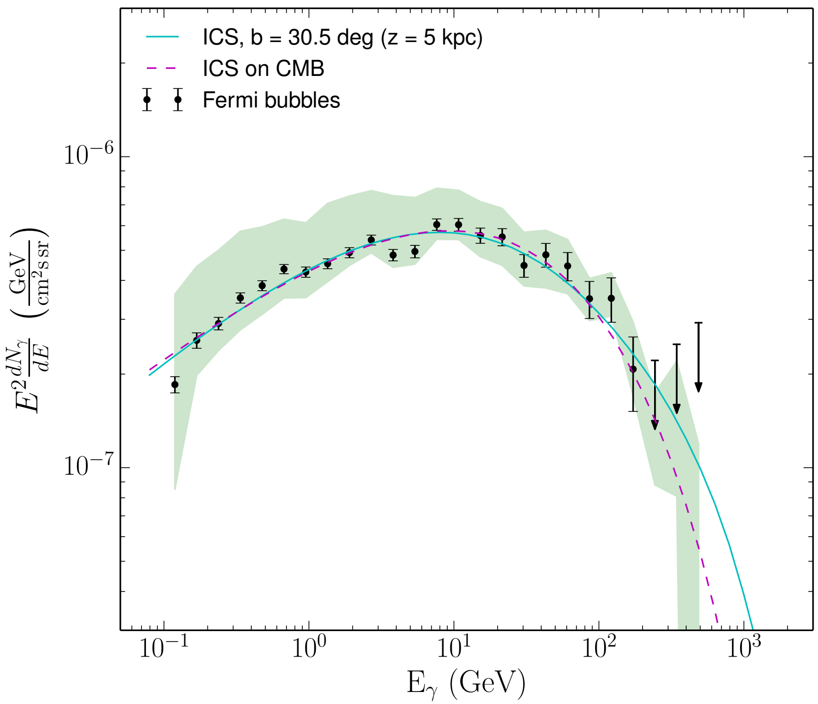
wind convects CRs away from disk

Aharonian & Crocker, PRL, 106 (2011)

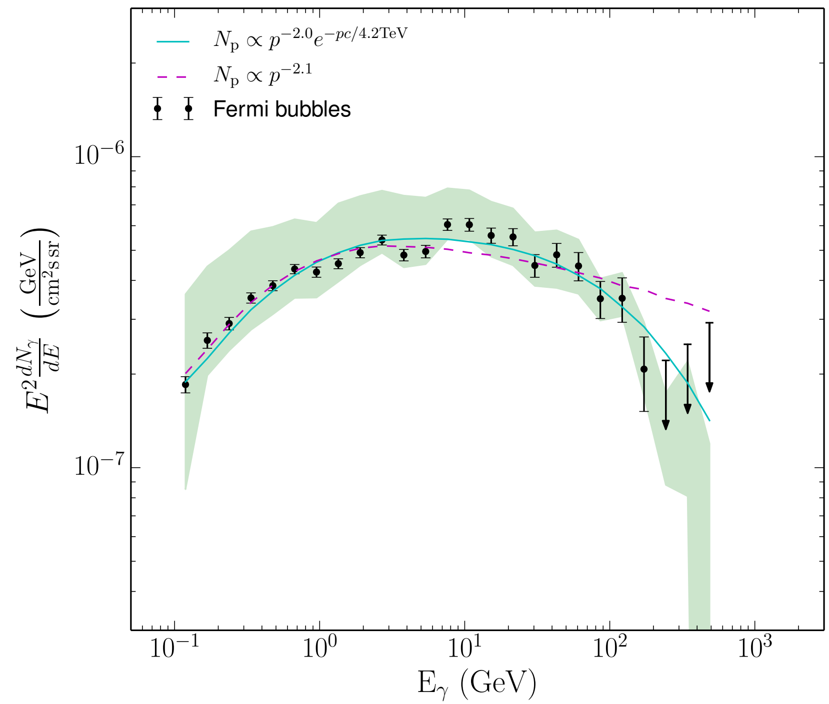
Illustrations by P. Mertsch

Gamma-ray spectrum

Leptonic model



Hadronic model + secondary IC

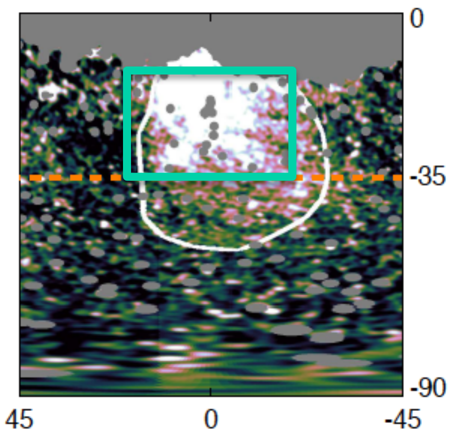


Ackermann et al (Fermi LAT), ApJ 793 (2014)

- **Both leptonic and hadrnoic models fit the spectrum**

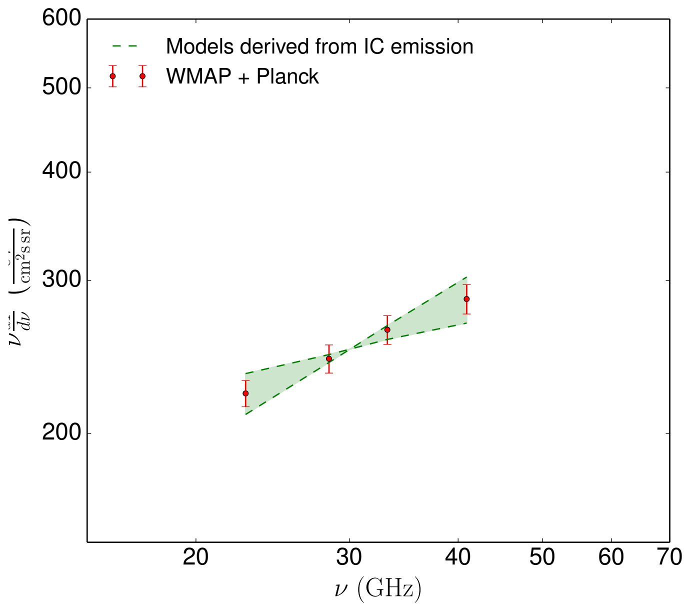
Microwave haze

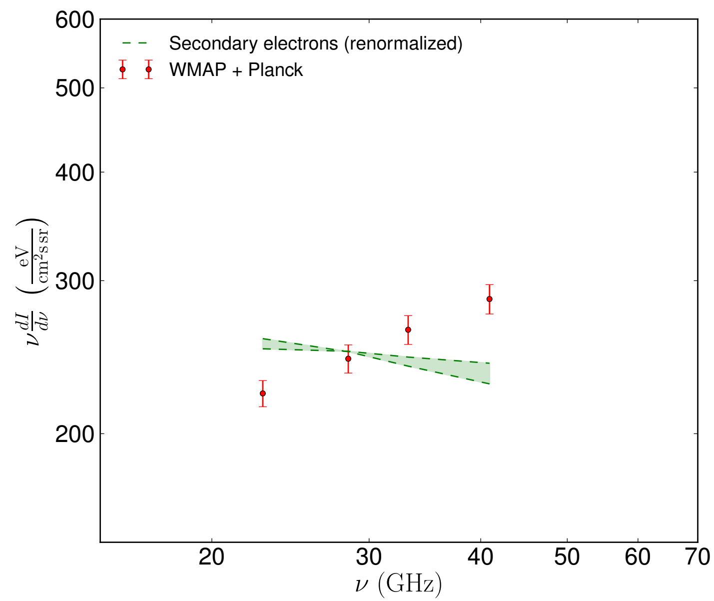
Leptonic | Hadronic (secondary leptons)



|l| < 25° and
-35° < b < -10°

Planck Collaboration
A&A 554 (2013)



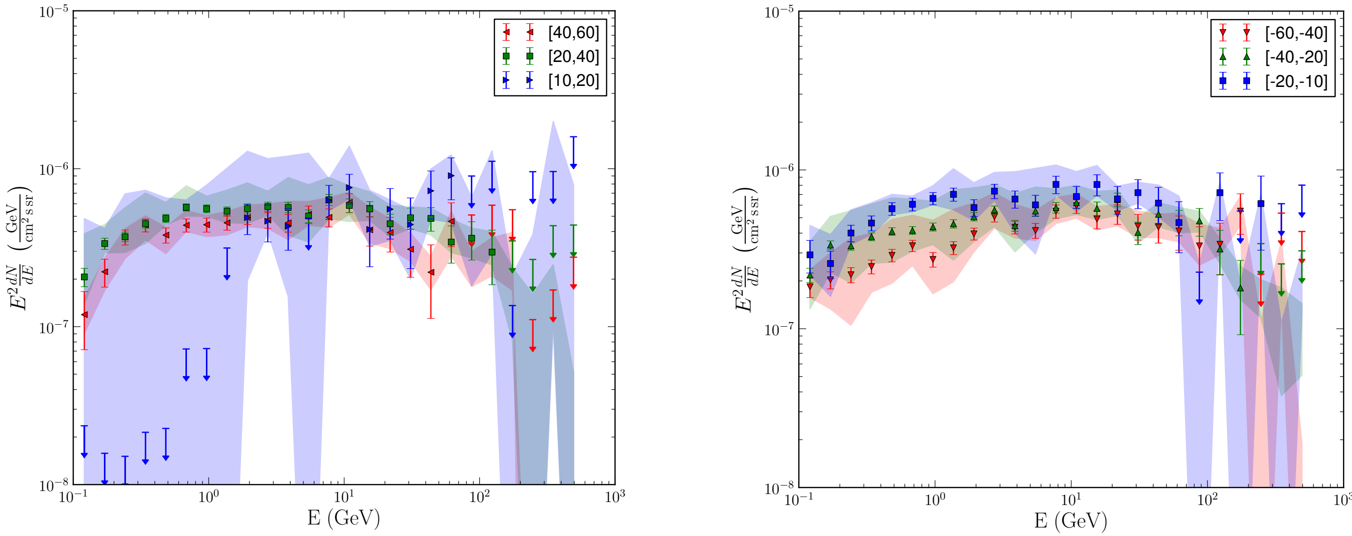


Ackermann et al (Fermi LAT), ApJ 793 (2014)

- **Synchrotron emission from secondary leptons in hadronic models cannot explain the microwave haze**

Spectrum in stripes

- At latitudes |b| > 10°, the spectrum is uniform



Ackermann et al (Fermi LAT), ApJ 793 (2014)

- Natural in hadronic models
- In leptonic models the velocity should be > 10000 km/s to avoid e^+e^- cooling before they reach z ~ 10 kpc distance
  - stochastic reacceleration: **Mertsch & Sarkar PRL 107 (2011)**

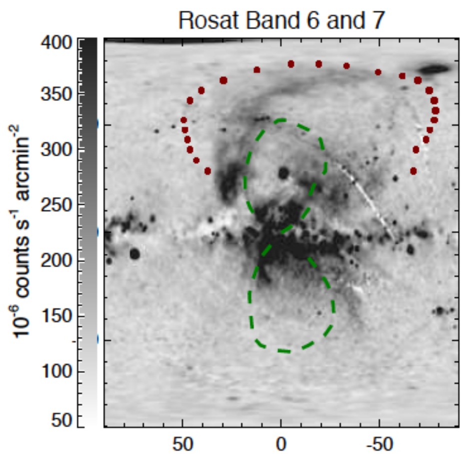
Bubbles boundary

- Narrow boundary
  - Natural in AGN models – result of expansion
  - In star-formation models, one needs a mechanism that keeps CR from escaping, e.g., magnetic draping
- Absence of a shock
  - Natural in star-formation / hadronic models
  - In leptonic models one needs to (re)accelerate electrons

More puzzles

- **X-rays**

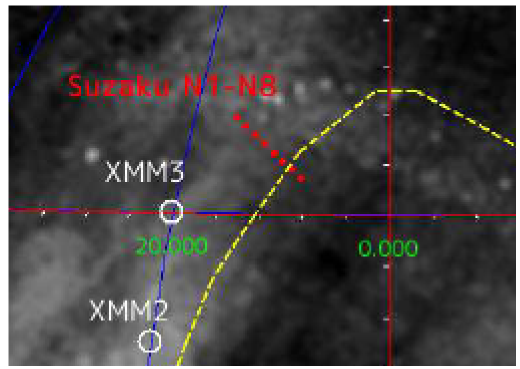
ROSAT



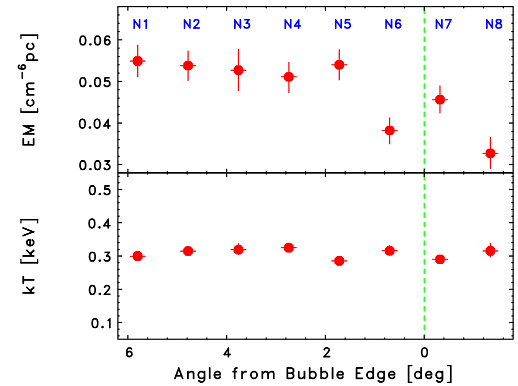
Su, Slatyer, Finkbeiner
ApJ 724 (2010)

Suzaku

Pointings



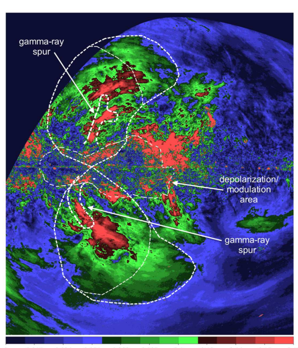
Emission measure



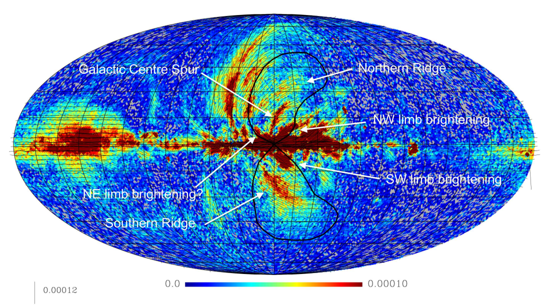
Kataoka et al, ApJ 779 (2013)

- **Polarization**

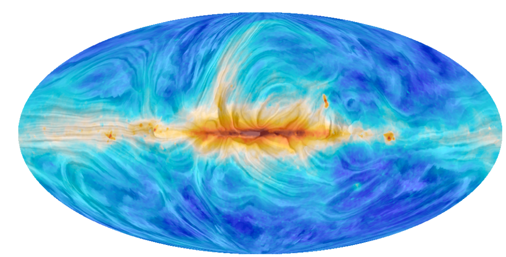
S-PASS, 2.3 GHz



WMAP, 23 GHz
polarization



Planck, 30 GHz
polarization



Adam et al (Planck), arXiv:1502.01582

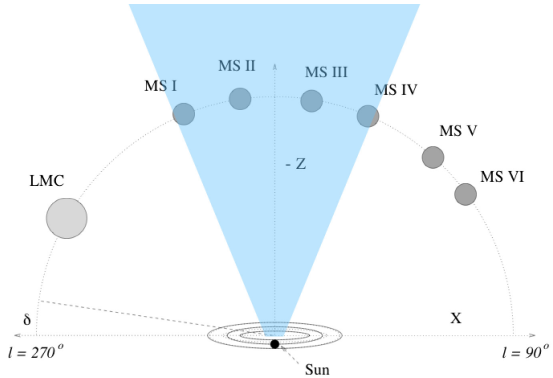
Carretti et al, Nature 493 (2013)

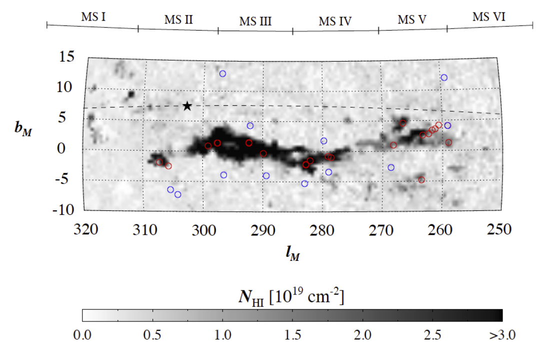
Scoreboard

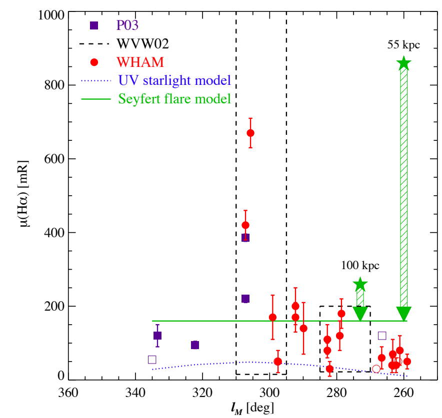
	Leptonic	Hadronic
Energy spectrum	✔	with secondary IC
WMAP / Planck haze	✔	extra component
Isotropic emission	reacceleration	✔
Narrow boundary	✔	magnetic draping
No visible shock	?	✔

AGN or starburst

- Often happen together
- Evidence for an AGN-like activity 0.5 – 5 Myr ago
  - Magellanic stream ionization

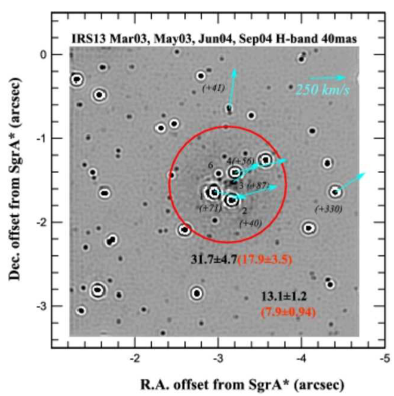


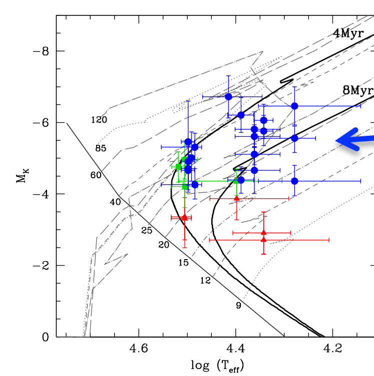




Bland-Hawthorn et al, ApJ 778 (2013)

- Young (~ 6 Myr) stellar population near the GC (~ 10^4 M_{sun})



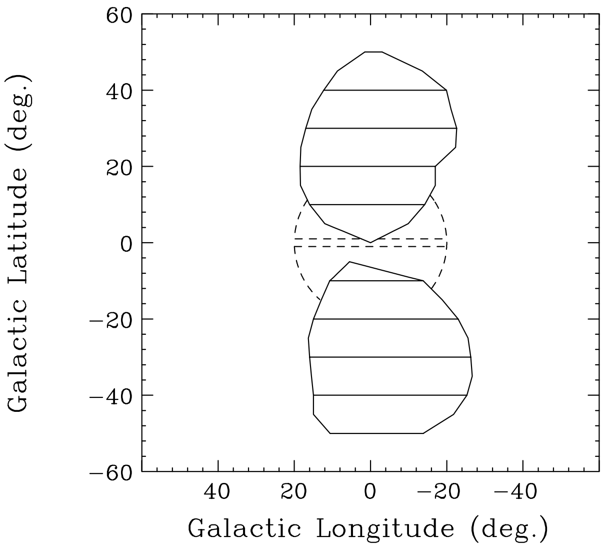


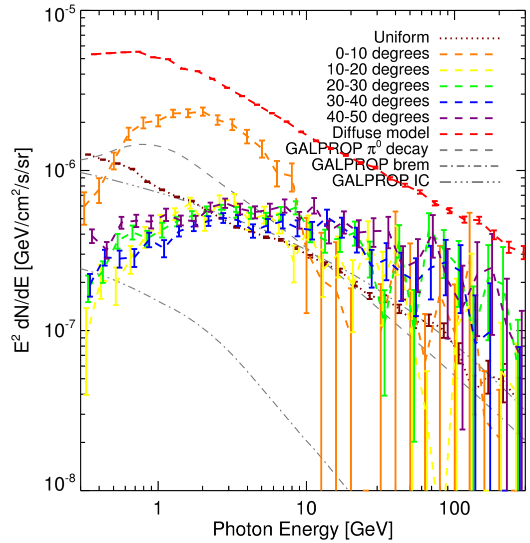
OB supergiants
(~ few Myr lifetime)

Paumard et al,
ApJ 643 (2006)

Galactic center and the Fermi bubbles

- Fermi bubbles spectrum for |b| < 10°





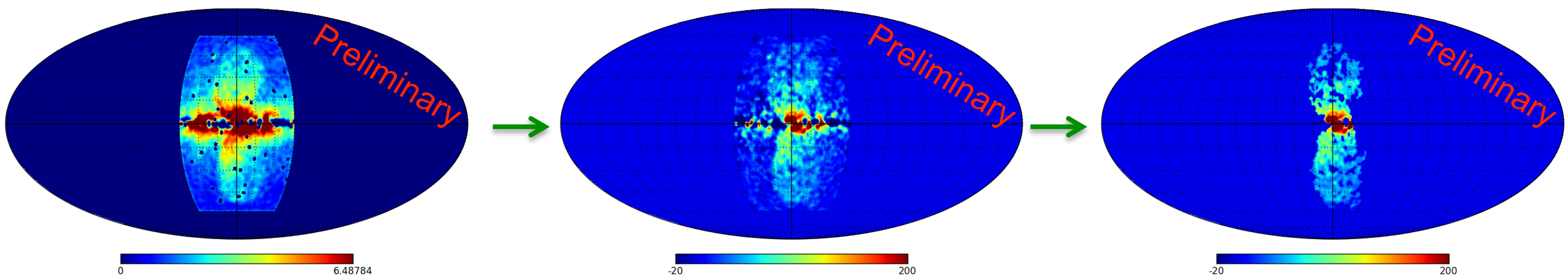
Hooper & Slatyer
Phys.Dark Univ. 2 (2013)

- Is it a part of the Fermi bubbles or a separate component?
- Options
  - Only bubbles
  - No bubbles
  - Both the bubbles and a new component

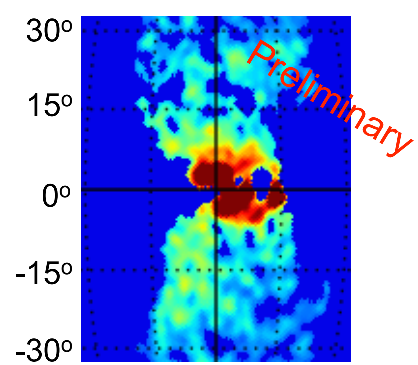
Fermi bubbles at low latitudes

- Assume that the bubbles have the same spectrum near the GC as at high latitudes ~ E^{-2} between 1 and 10 GeV
- Subtract π^0 component and PS from data and represent the residual using two components:
  - Bubble-like ~ E^{-2}
  - Other components (IC, ISO, Loop I etc.) ~ $E^{-2.4}$

Data – gas – PS | ~E^{-2} component | Bubbles template



- Fermi bubbles template near the GC:
  - Larger intensity
  - Displaced to the right from the GC



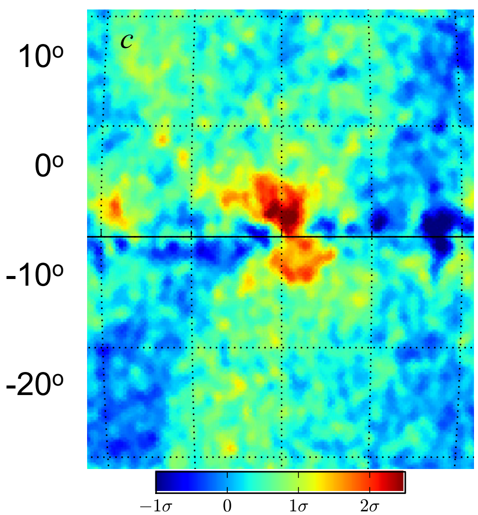
Fermi bubbles near the GC

- **Center of the Fermi bubbles intersection with the Galactic plane:**

  ~ 1° – 2° or about 100 – 300 pc to the right of the GC?

Fermi LAT Pass 7 diffuse model

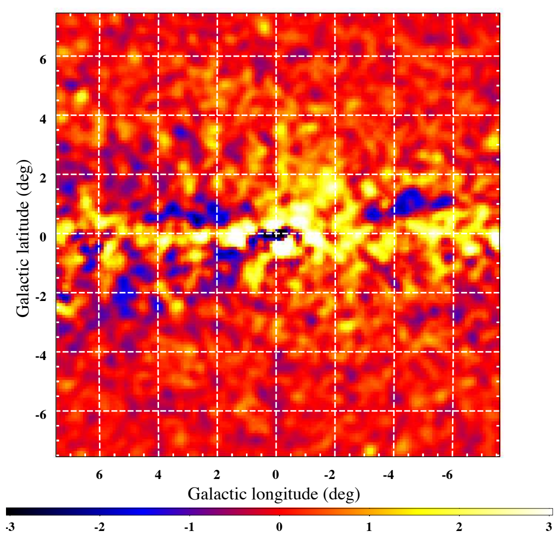
Fermi bubbles



Acero et al (Fermi LAT) ApJS 223 (2016)

Fermi LAT analysis of the GC excess?

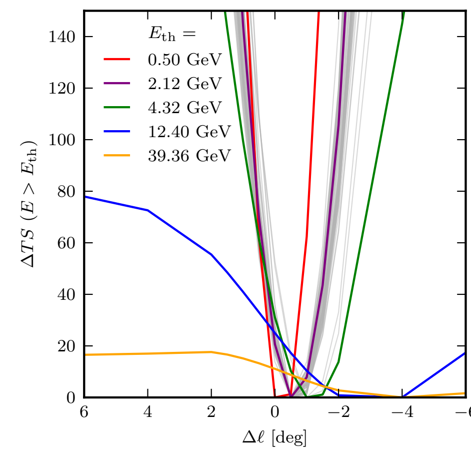
Residual 1.6 – 10 GeV



Ajello et al (Fermi LAT) ApJ 819 (2016)

Calore et al GC excess analysis?

Displacement of the GC excess:



Calore et al, JCAP 1503 (2015)

Future

- eROSITA
  - Search for cavity in hot gas plasma due to CR pressure inside the Fermi bubbles
- HESS, MAGIC, VERITAS, CTA, HAWC
  - Fermi bubbles near the GC seem to be brighter
  - Possible to see with Cherenkov telescopes?
- IceCube, KM3net
  - Search for neutrinos from the Fermi bubbles
- More analysis of existing data
  - Fermi LAT (Pass 8 data)
  - Planck polarization

Conclusions

- Fermi bubbles are a unique feature on gamma-ray sky
  - Relatively bright in gamma rays
  - No clear counterpart at high latitudes in X-rays or radio
- Possible origin and emission mechanisms
  - AGN-like activity of Sgr A* (IC gamma rays)
  - Enhanced star formation near the GC (π^0 gamma rays)
    - Both scenarios have advantages and disadvantages
- Tentative characterization at low latitudes:
  - Enhanced intensity near the Galactic plane
  - Displaced to the right (negative longitudes) from the GC
- Origin of the Fermi bubbles is an exciting question
  - Should learn more soon using new data from
    - eROSITA, CTA, IceCube and KM3net